Dark Radiation

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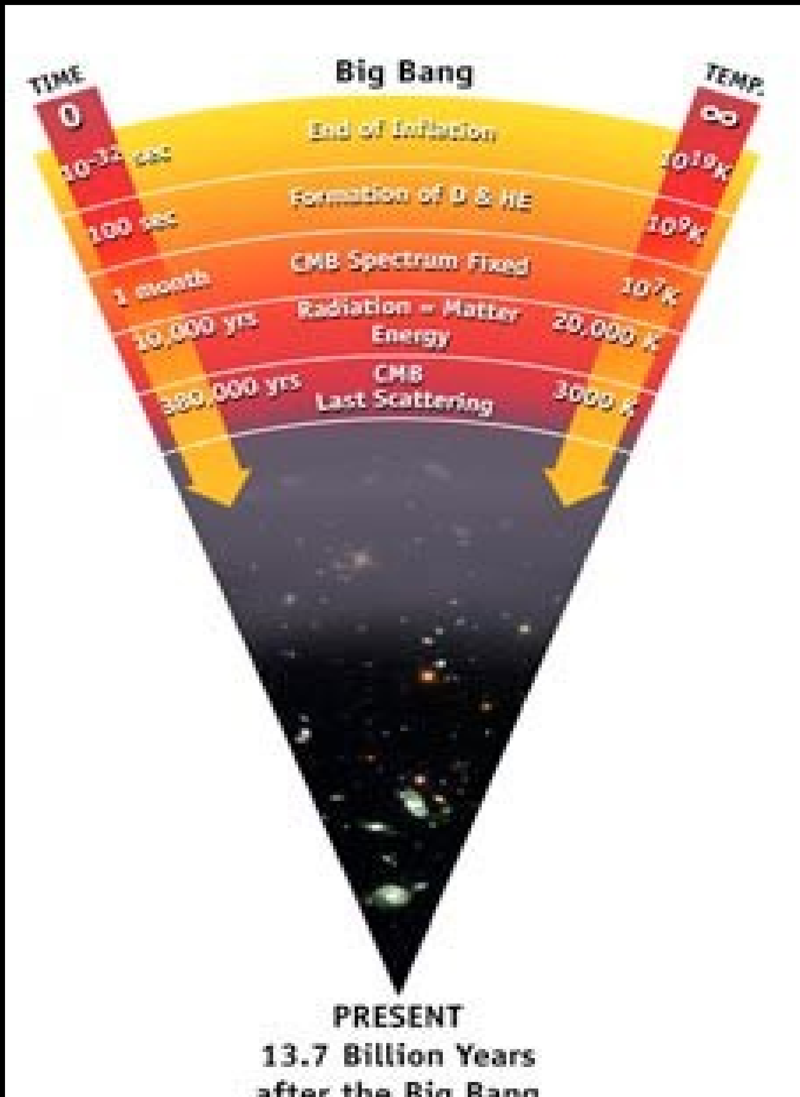
Outline:

- Introduction to cosmology: the CMB angular power spectrum and datasets used;
- The Cosmic Neutrino Background (CNB) and its clustering parameters;
- Constraints on new physics damping tail parameters: the Neutrino effective Number (N_{eff}) and the Lensing Amplitude (A_L);
- Some Dark radiation models
- Indipendent ways to constrain N_{eff}: the neutrino isocurvature mode and the bispectrum signal
- Conclusions

Introduction to cosmology



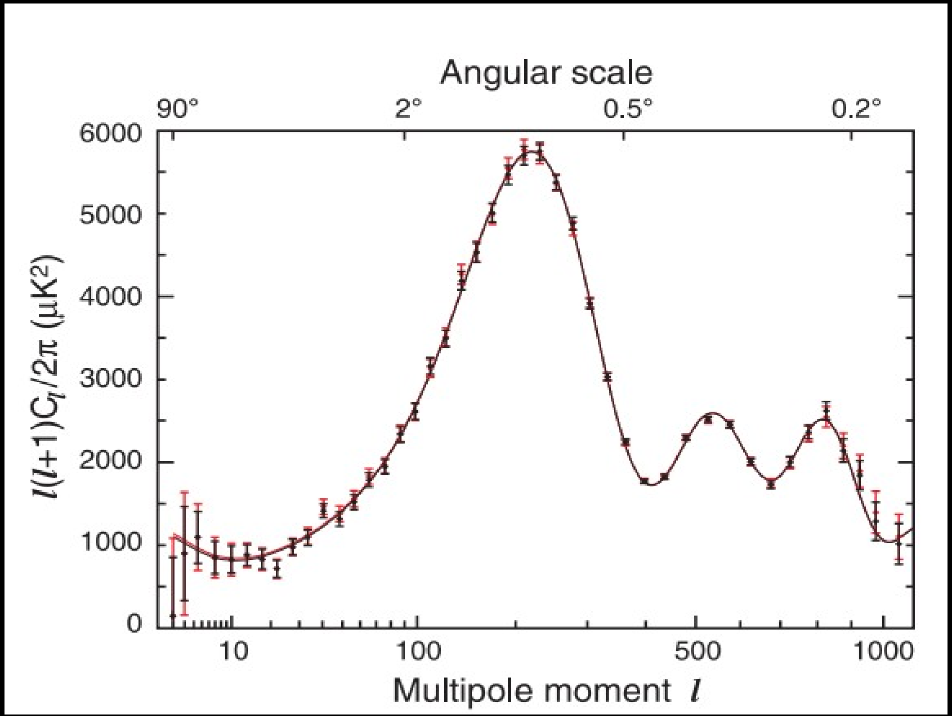
The Universe originates from a hot Big Bang. The primordial plasma in thermodynamic equilibrium cools with the expansion of the Universe. It passes through the phase of decoupling, in which the Universe becomes trasparent to the motion of photons, and the phase of recombination, where electrons and protons combine into hydrogen atoms.

The Cosmic Microwave Background (CMB) is the radiation coming from the recombination, emitted about 13 billion years ago, just 400,000 years after the Big Bang. The CMB provides an unexcelled probe of the early Universe and today it is a black body a temperature T=2.726K.

Introduction to cosmology

The main tool of research in cosmology is the angular power spectrum of CMB temperature anisotropies.

$$\left\langle \frac{\Delta T}{T}\left(\vec{\gamma}_1\right)\frac{\Delta T}{T}\left(\vec{\gamma}_2\right)\right\rangle = \frac{1}{2\pi}\sum_{\ell}(2\ell+1)C_\ell P_\ell\left(\vec{\gamma}_1\cdot\vec{\gamma}_2\right)$$



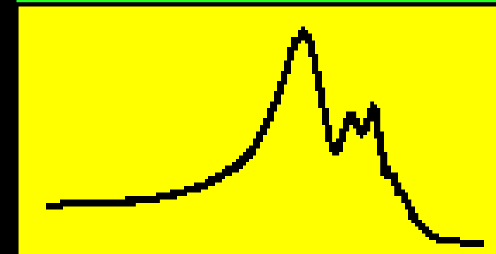
Hinshaw et al., arXiv: 1212.5226

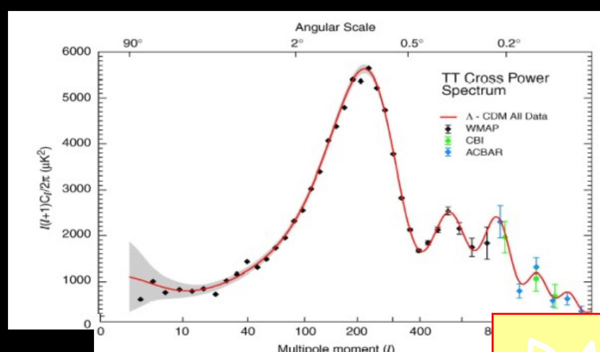
- Dots: cosmological data
- Solid lines: theoretical models

Introduction to cosmology

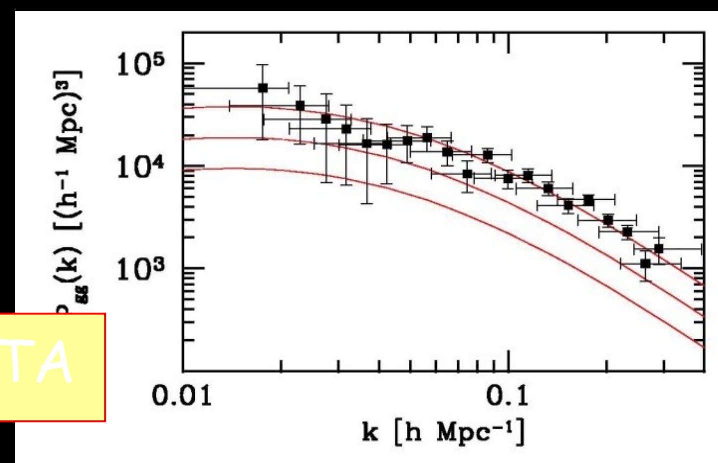
Cosmological parameters:
$(\Omega_b h^2, \Omega_m h^2, h, n_s, \tau, \Sigma m_\nu)$

Theoretical model





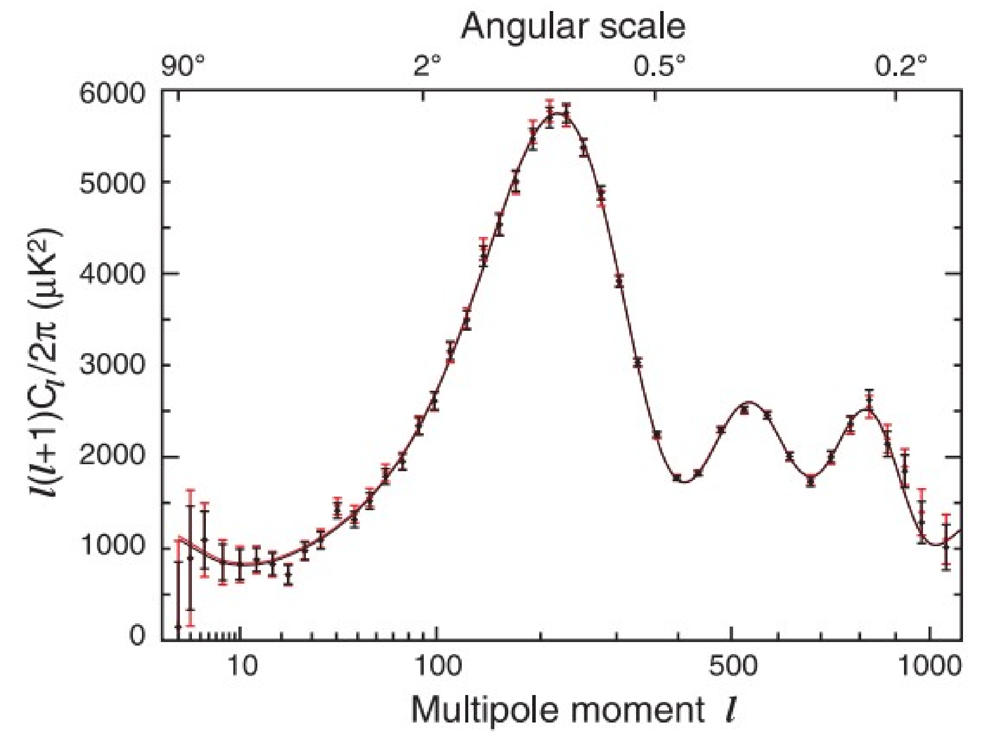
DATA



PARAMETER CONSTRAINTS

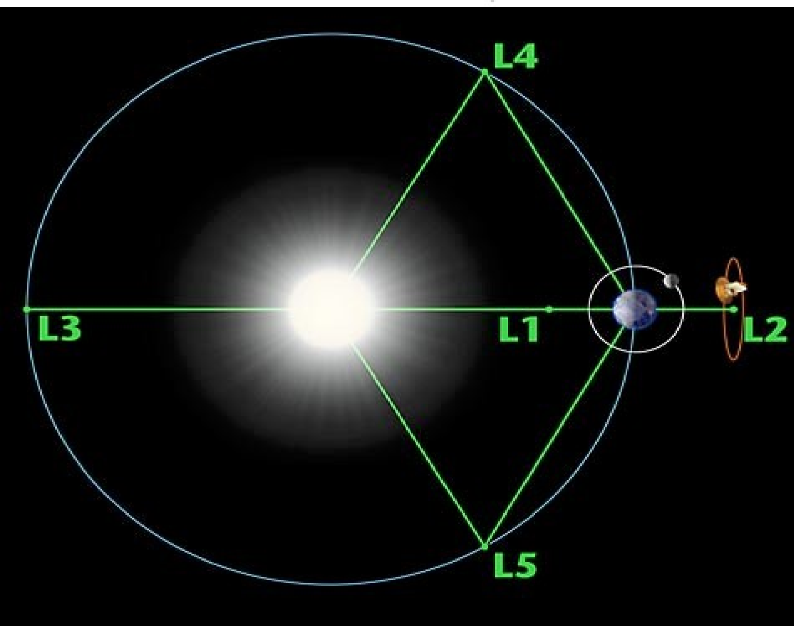
Cosmological data

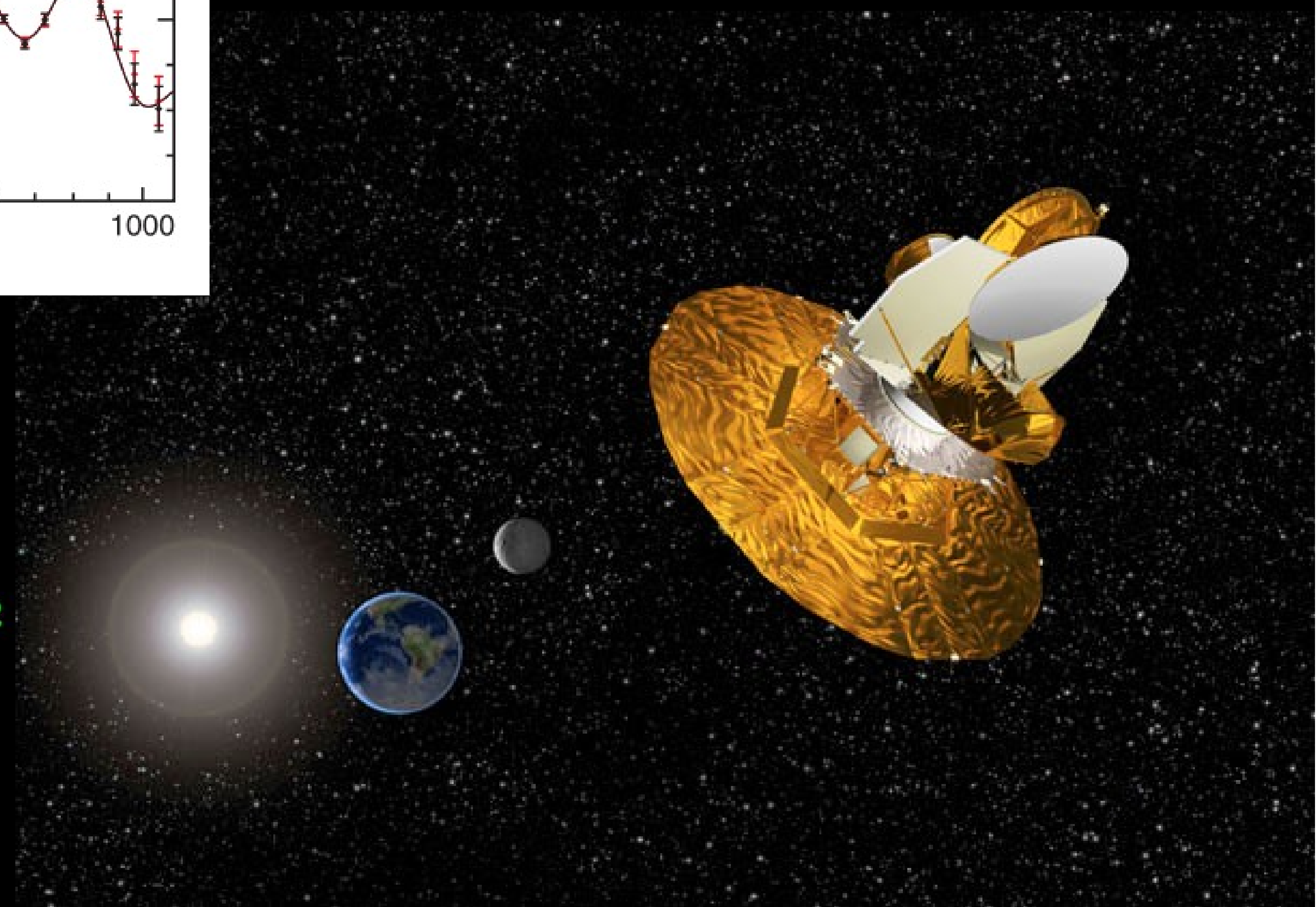
Hinshaw et al., arXiv: 1212.5226



WMAP satellite experiment:

- Angular resolution 0.3°;
- 1.4 and 1.6 meters telescopes;
- Orbit around L2;
- Frequency range of 22Ghz to 90Ghz.



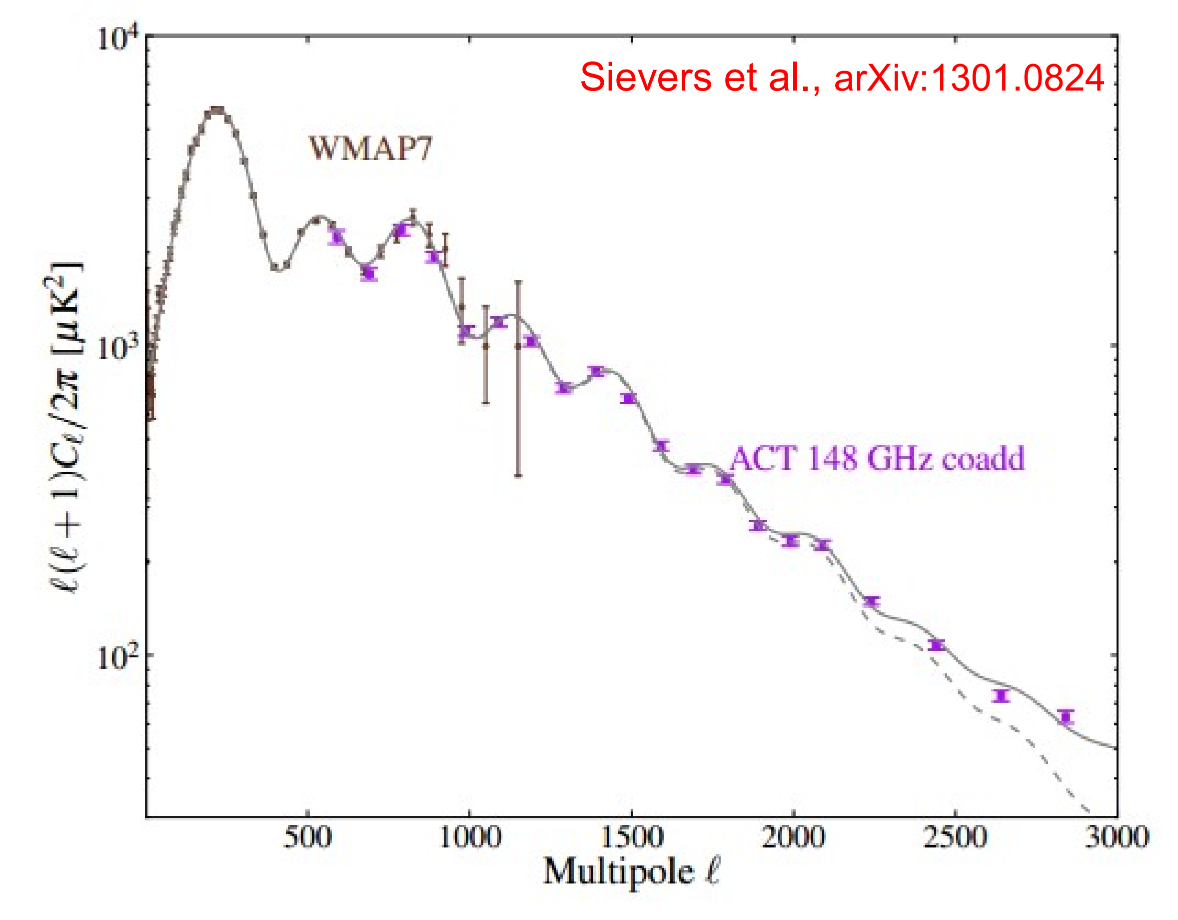


Cosmological data

Ground based observations have a better angular resolution and they measure the CMB spectrum at smaller angular scales:

The Atacama Cosmology Telescope (ACT):

- Six-meters telescope on Cerro Toco in the Atacama Desert in the north of Chile.



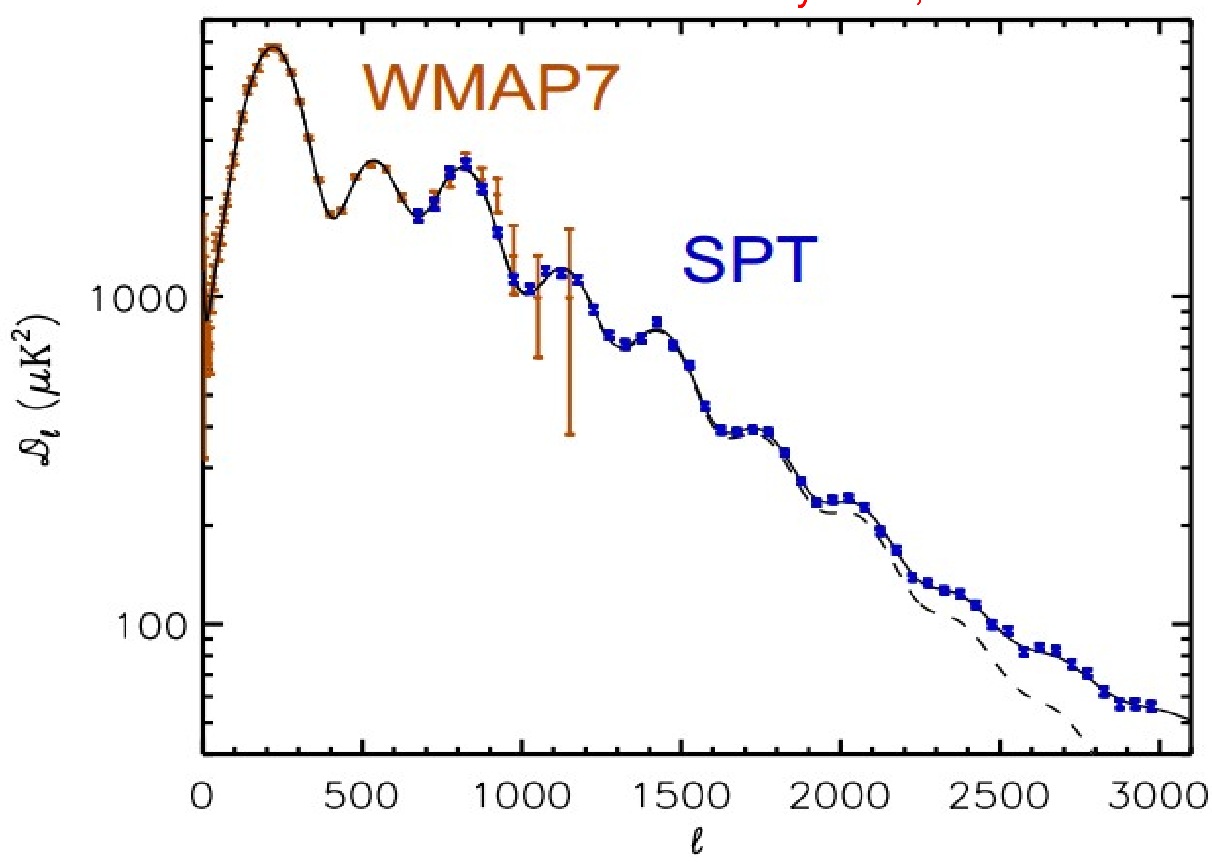
Cosmological data

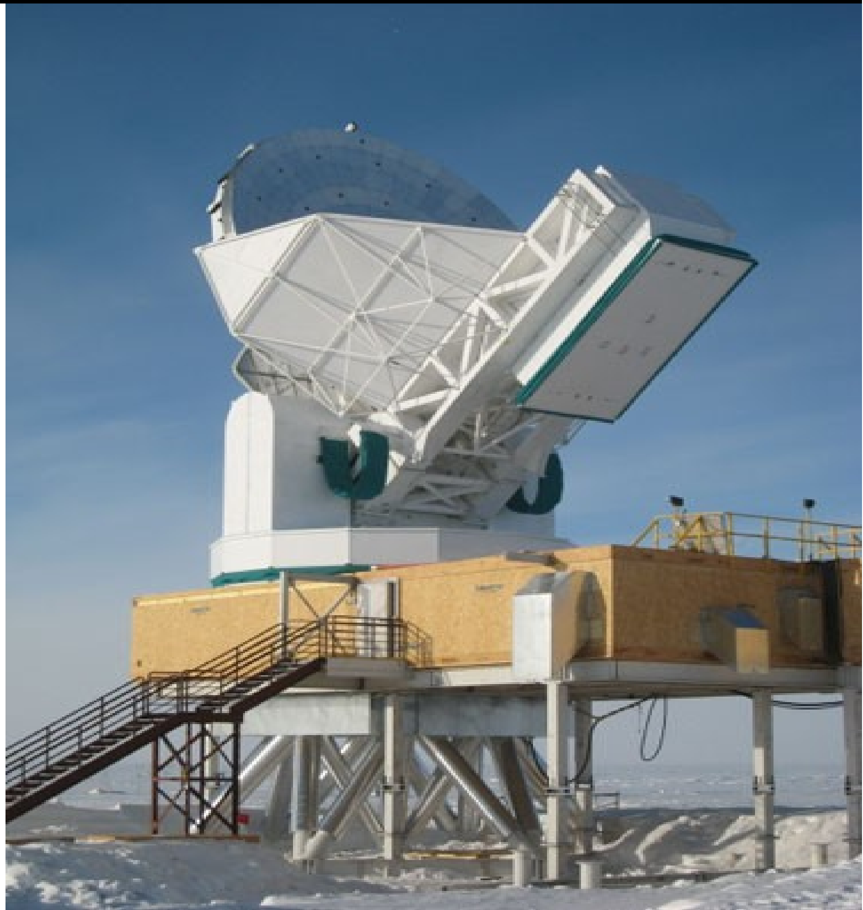
Ground based observations have a better angular resolution and they measure the CMB spectrum at smaller angular scales:

The South Pole Telescope (SPT):

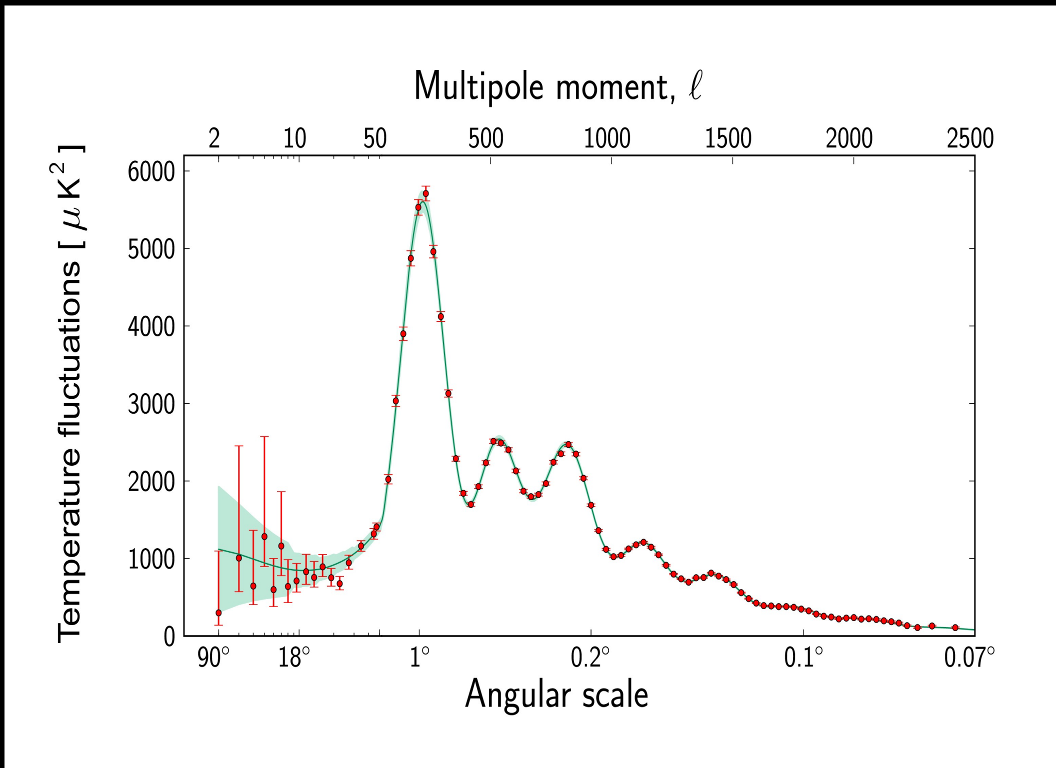
- Ten meters diameter telescope in the Amundsen-Scott South Pole Station, Antarctica.

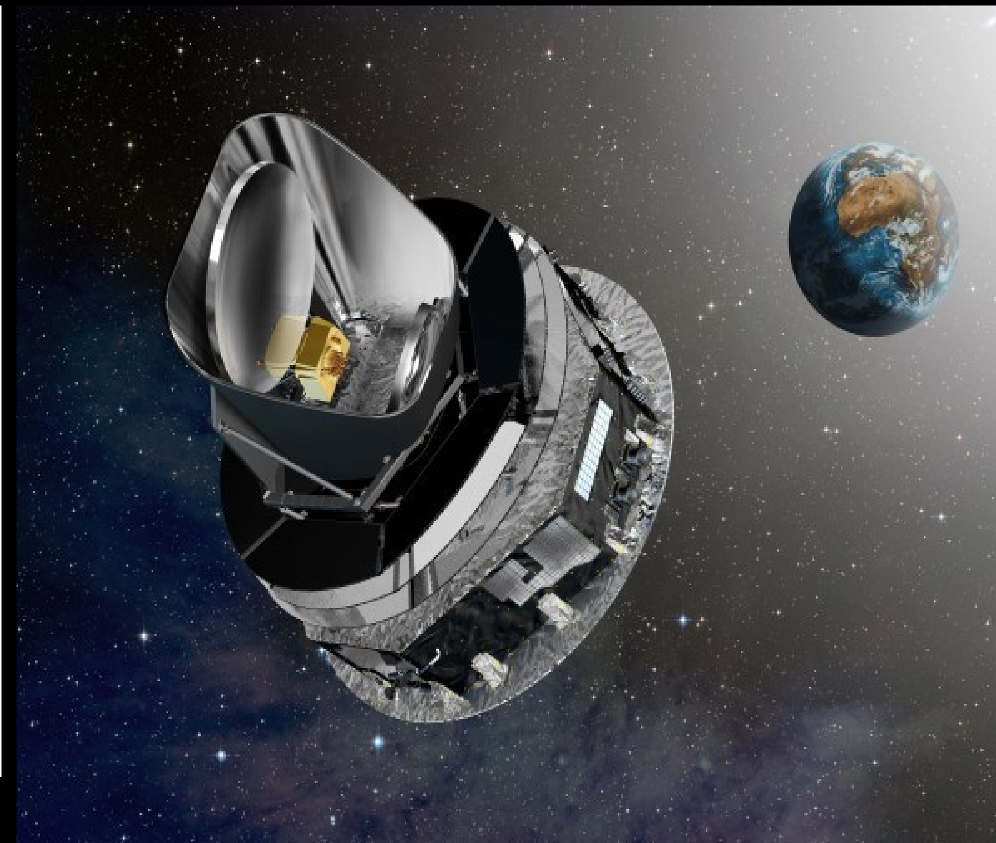
Story et al., arXiv:1210.7231





Cosmological data



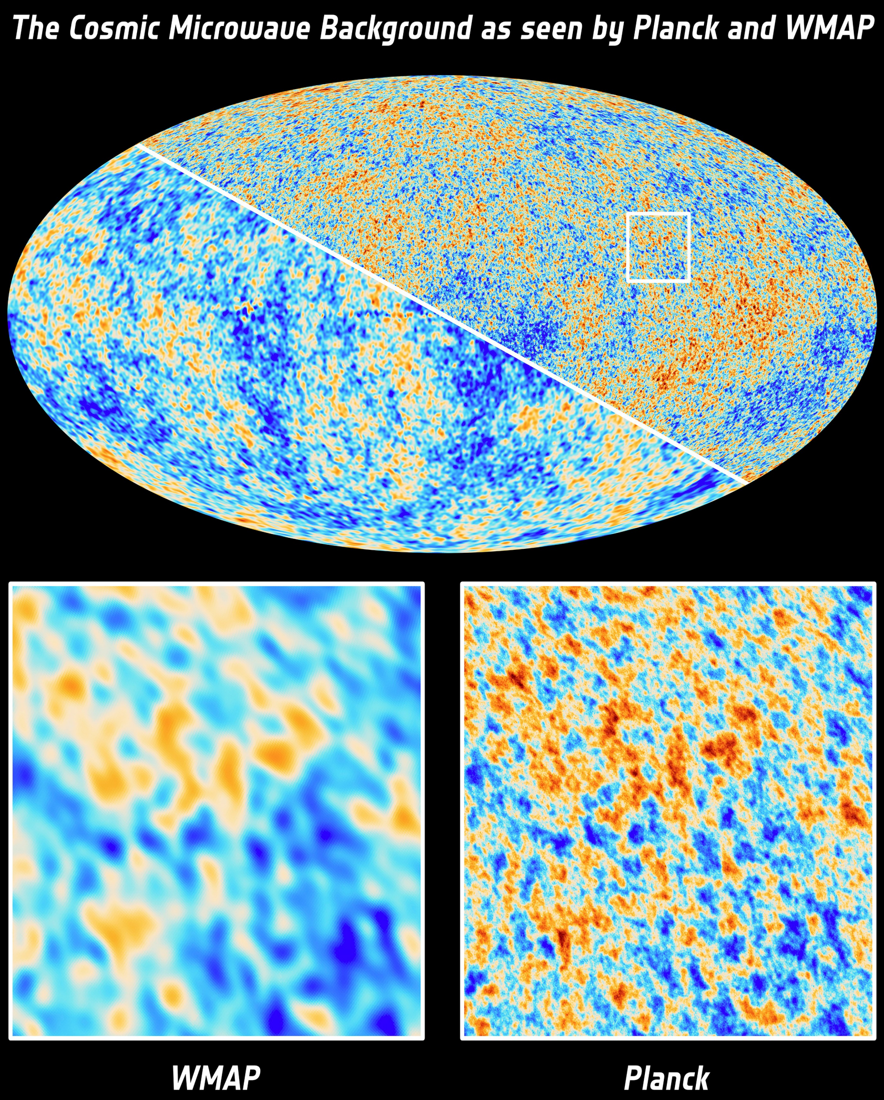


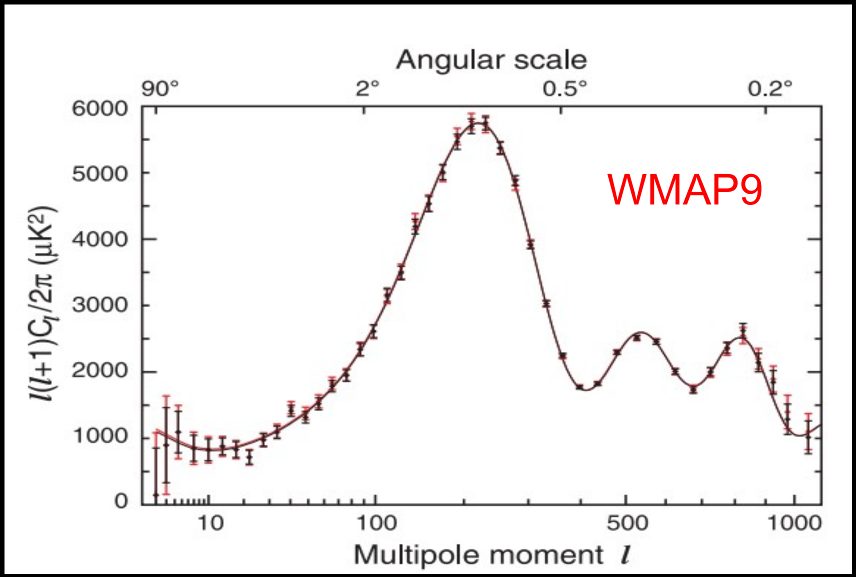
Planck satellite experiment:

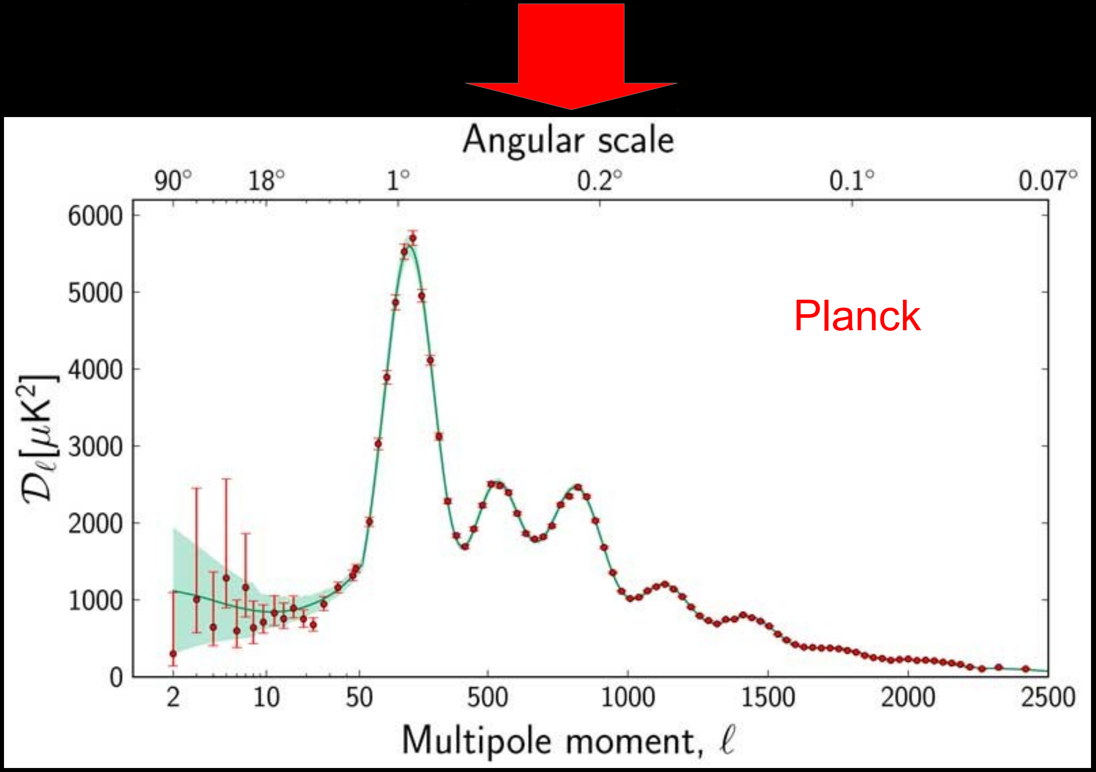
- Frequency range of 30GHz to 857GHz;
- Orbit around L2;
- Composed by 2 instruments:
  - LFI $\rightarrow$ 1.5 meters telescope; array of 22 differential receivers that measure the signal from the sky comparing with a black body at 4.5K.
  - HFI $\rightarrow$ array of 52 bolometers cooled to 0.1K.

Cosmological data

The recent Planck satellite data have improved the WMAP 9 years measurements:







The Cosmic Neutrino Background

When the rate of the weak interaction reactions, which keep neutrinos in equilibrium with the primordial plasma, becomes smaller than the expansion rate of the Universe, neutrinos decouple:

$$T_{dec} \approx 1MeV$$

After neutrinos decoupling, photons are heated by electrons-positrons annihilation. When also photons decouple, the ratio between the temperatures of photons and neutrinos will be fixed, despite the temperature decreases with the expansion of the Universe. We expect today a Cosmic Neutrino Background (CNB) at a temperature:

$$T_\nu = \left(\frac{4}{11}\right)^{1/3} T_\gamma \approx 1.945K \rightarrow \; kT_\nu \approx 1.68 \cdot 10^{-4} eV$$

With a number density of:

$$n_f = \frac{3}{4}\frac{\varsigma(3)}{\pi^2} g_f T_f^3 \rightarrow n_{\nu_k, \bar{\nu}_k} \approx 0.1827 \cdot T_\nu^3 \approx 112 cm^{-3}$$

N_{eff}

The relativistic neutrinos contribute to the present energy density of the Universe:

$$\rho_{rad} = \rho_\gamma + \rho_\nu = g_\gamma \left(\frac{\pi^2}{30}\right) T_\gamma^4 + g_\nu \left(\frac{\pi^2}{30}\right)\left(\frac{7}{8}\right) T_\nu^4$$

$$\rho_{rad} = \left(1 + \left(\frac{7}{8}\right)\left(\frac{4}{11}\right)^{\frac{4}{3}} \left(\frac{g_\nu}{g_\gamma}\right)\right) \rho_\gamma$$

We can introduce the effective number of relativistic degrees of freeedom:

$$\rho_{rad} = \left[1 + \frac{7}{8}\left(\frac{4}{11}\right)^{4/3} N_{\rm eff}\right] \rho_\gamma$$

The expected value is $N_{eff} = 3.046$, if we assume standard electroweak interactions and three active massless neutrinos. The 0.046 takes into account effects for the non-instantaneous neutrino decoupling and neutrino flavour oscillations. (Mangano et al. 2005)

Neff

Measuring $\Delta N_{eff} \equiv N_{eff} - 3.046$ we can constrain the dark radiation.

Increasing N_{eff} essentially increases the expansion rate H at recombination.

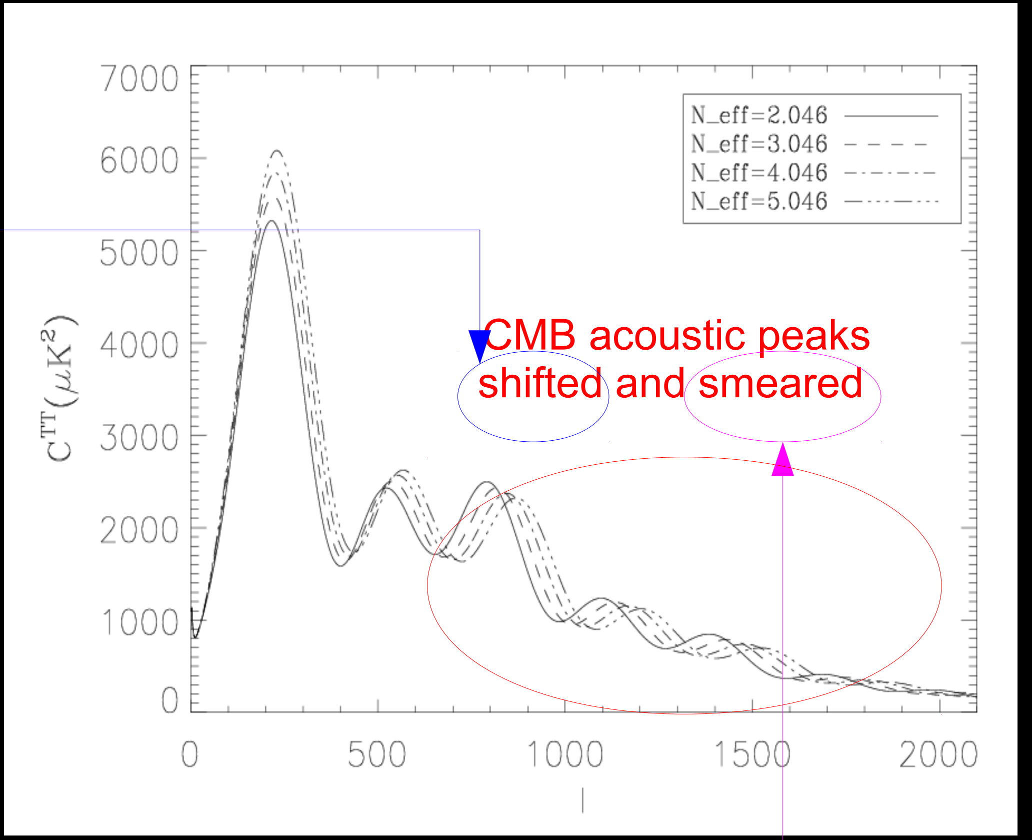
$$H^2 = \left(\frac{\dot{a}}{a}\right)^2 = H_0^2\left(\frac{\Omega_r}{a^4} + \frac{\Omega_m}{a^3} + \frac{\Omega_k}{a^2} + \Omega_\Lambda\right)$$

So decrease the sound horizon at recombination,

$$r_s = \int_0^{t_*} c_s\, dt/a = \int_0^{a_*} \frac{c_s\, da}{a^2 H}.$$

and the diffusion distance (damping scale):

$$r_d^2 = (2\pi)^2 \int_0^{a_*} \frac{da}{a^3 \sigma_T n_e H}\left[\frac{R^2 + \frac{16}{15}(1+R)}{6(1+R^2)}\right]$$



Probing the Neutrino Number with CMB data

Once the angular size of the sound horizon θ_s is fixed, we are fixing the angular scales of the acoustic peaks.

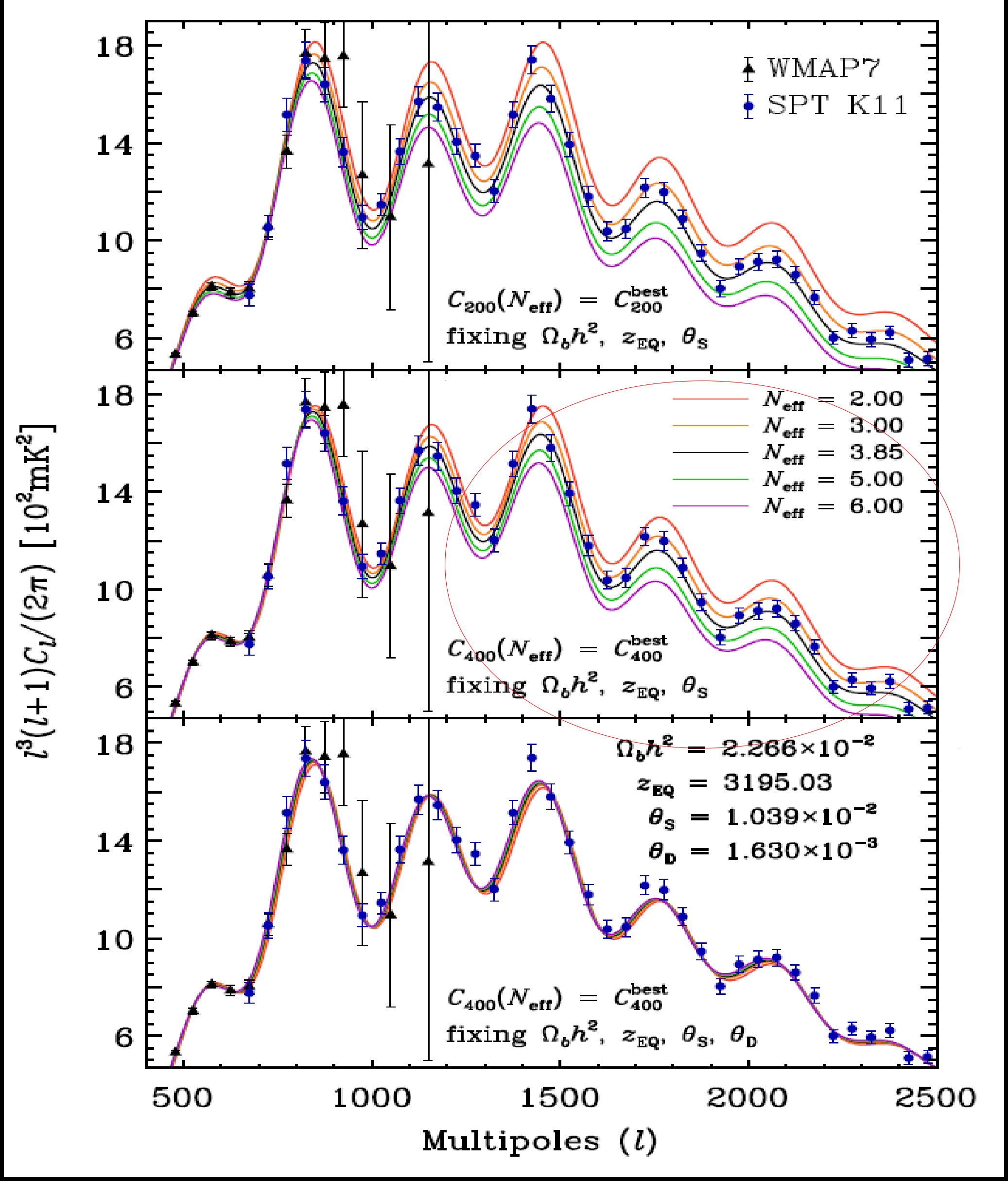
$$\theta_s = \frac{r_s}{D_A}$$

When we increase N_{eff}, we are increasing the angular scale of the diffusion length θ_d

$$\theta_d = \frac{r_d}{D_A}$$

and the result is an increasing of the damping in the small angular scale anisotropy.

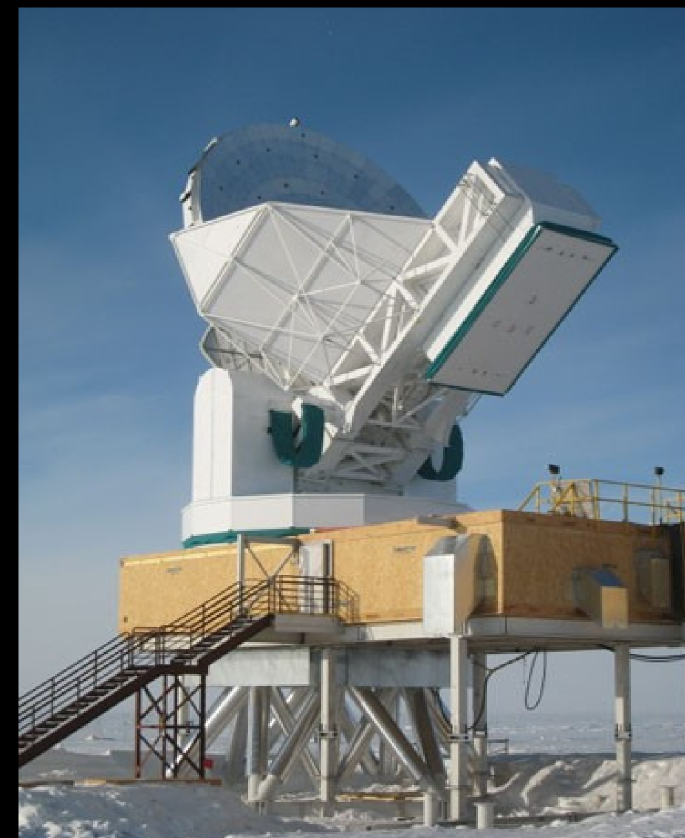
$$\theta_d = \frac{r_d}{r_s}\theta_s \simeq \frac{\sqrt{H}^{-1}}{H^{-1}} = \sqrt{H}$$



Hou et al., arXiv:1104.2333v2

CMB constraints on N_{eff} before Planck (December 2012)

South Pole Telescope (+WMAP7)



(Z. Hou et al. 2012)

$$N_{\mathrm{eff}} = 3.62 \pm 0.48$$

Atacama Cosmology Telescope (+WMAP7)

(J. L. Sievers et al. 2013)

$$N_{\mathrm{eff}} = 2.78 \pm 0.55$$

The lensing amplitude

This tension between SPT and ACT was not related only to N_{eff}..

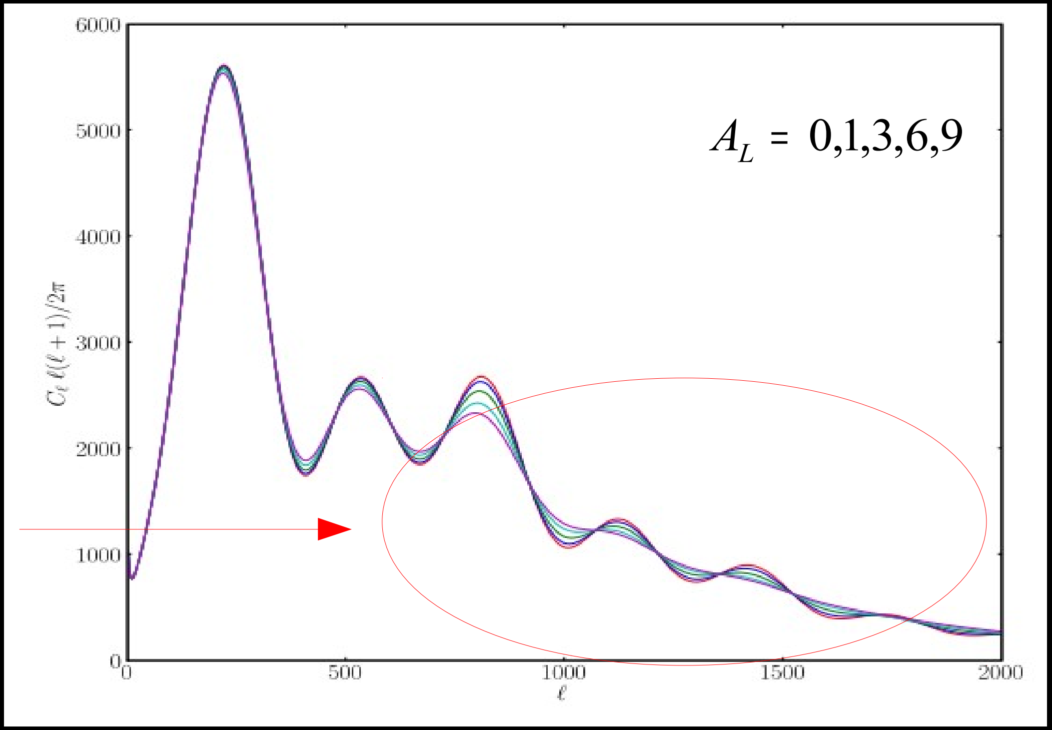
The lensing amplitude A_L parameterizes the rescaling of the lensing potential $\phi(n)$, then the power spectrum of the lensing field:

$$C_\ell^{\phi\phi} \rightarrow A_L C_\ell^{\phi\phi}$$

The gravitational lensing deflects the photon path by a quantity defined by the gradient of the lensing potential $\phi(n)$, integrated along the line of sight n, remapping the temperature field.

Its effect on the power spectrum is the smoothing of the acoustic peaks, increasing A_L.

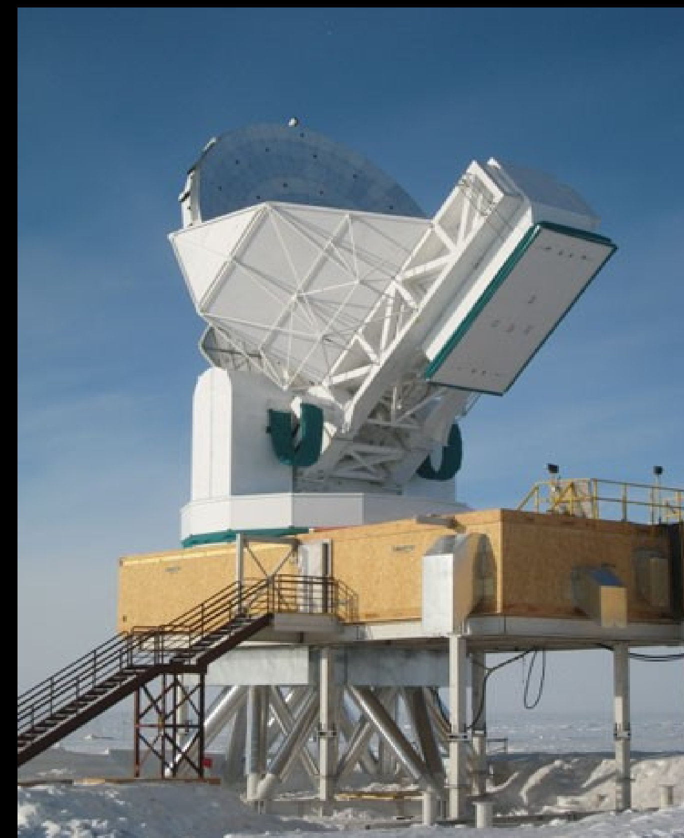
If A_L =1 then the theory is correct, else we have a new physics/systematics?



(E. Calabrese et al. 2008)

CMB constraints on A_L before Planck (December 2012)

South Pole Telescope (+WMAP7)



Atacama Cosmology Telescope (+WMAP7)

(J. L. Sievers et al. 2013)

(K. T. Story et al. 2012)

$$A_{\mathrm{L}} = 1.70 \pm 0.38$$

$$A_L = 0.86^{+0.15}_{-0.13}$$

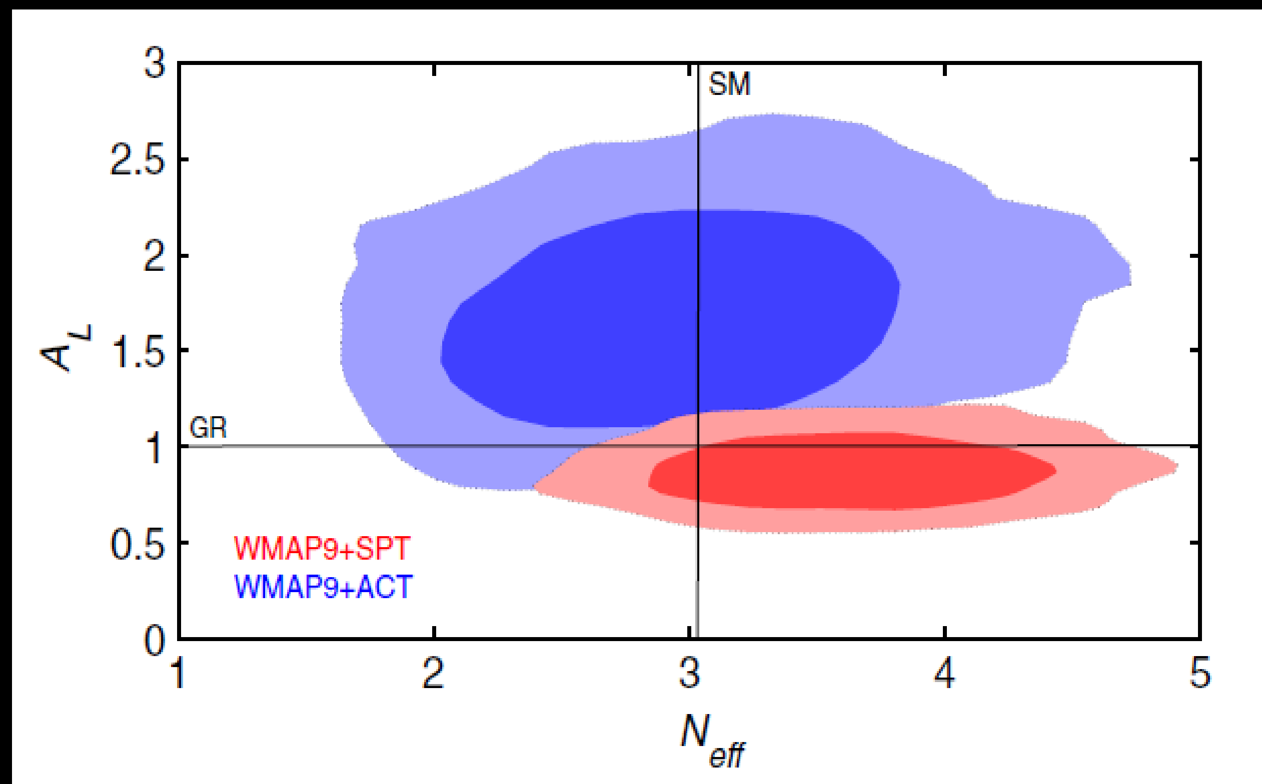
Pre-Planck constraints

Both N_{eff} and A_L affect the damping tail.

Tension between ACT and SPT is more clear in the A_L vs N_{eff} plane.

Parameter	WMAP9 + SPT
$N_{\rm eff}$	3.72 ± 0.46
$A_{\rm L}$	0.85 ± 0.13

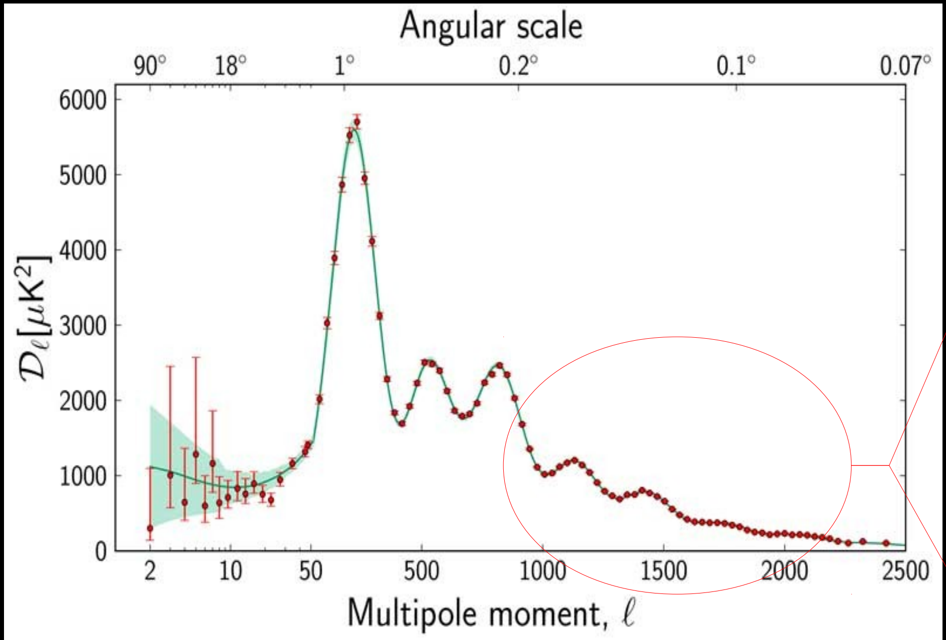
Parameter	WMAP9 + ACT
$N_{\rm eff}$	3.00 ± 0.61
$A_{\rm L}$	1.70 ± 0.37

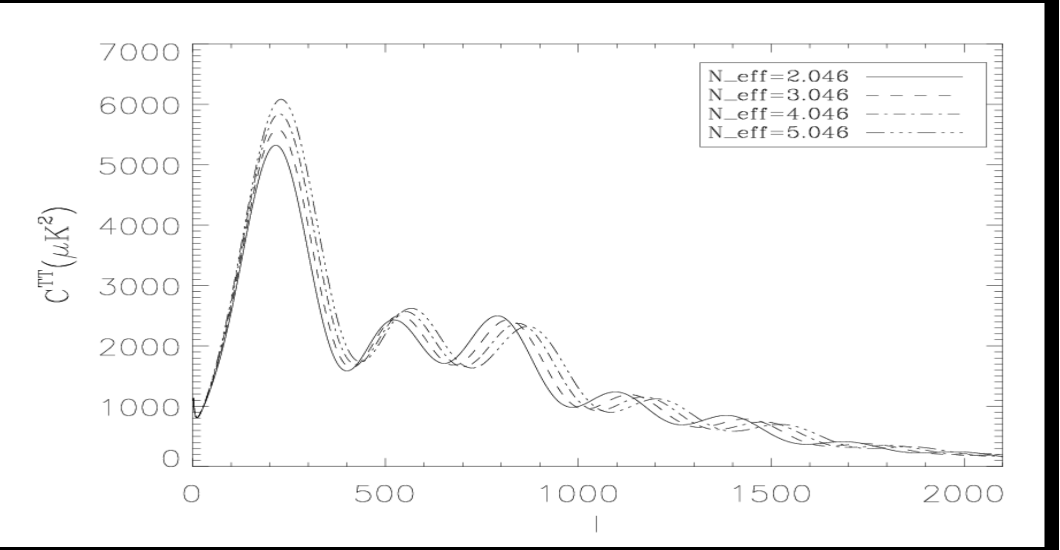


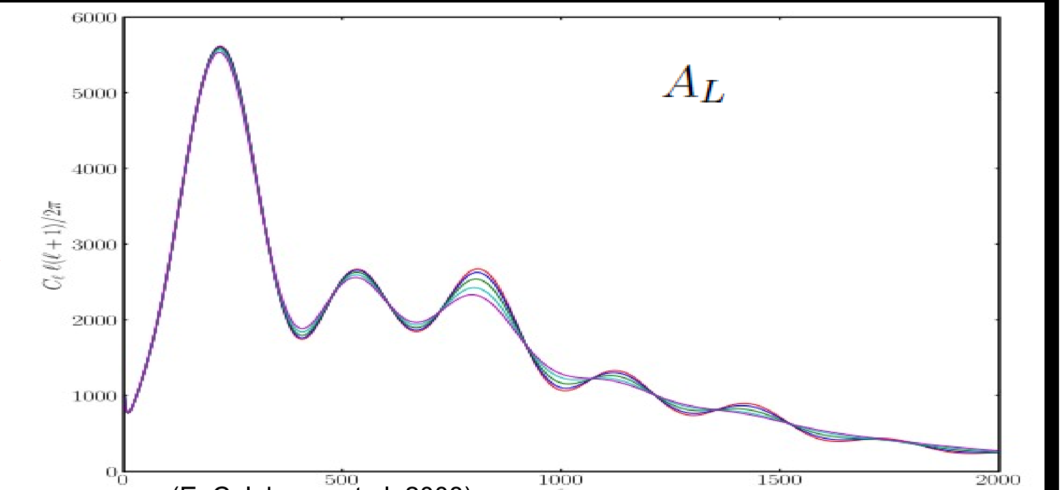
E. Di Valentino et al, Phys. Rev D, 88, 023501, 2013

The damping tail

The Planck satellite detected with high precision the anisotropy damping tail, allowing to better constrain these two parameters of new physics.







(E. Calabrese et al. 2008)

What about Planck?

Planck+WP 2013 result does not solve the issue!

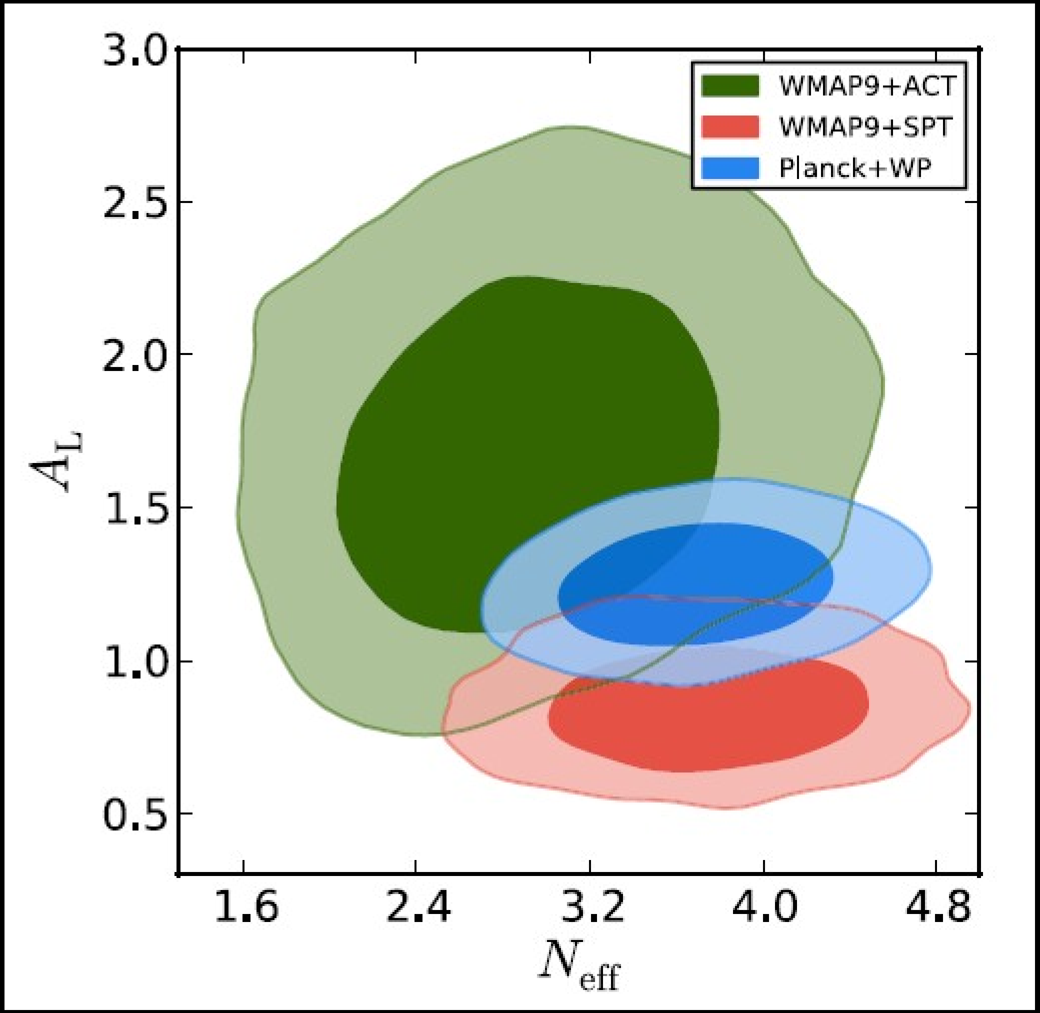
N_{eff} from Planck is in better agreement with SPT !

A_L from Planck is in better agreement with ACT !

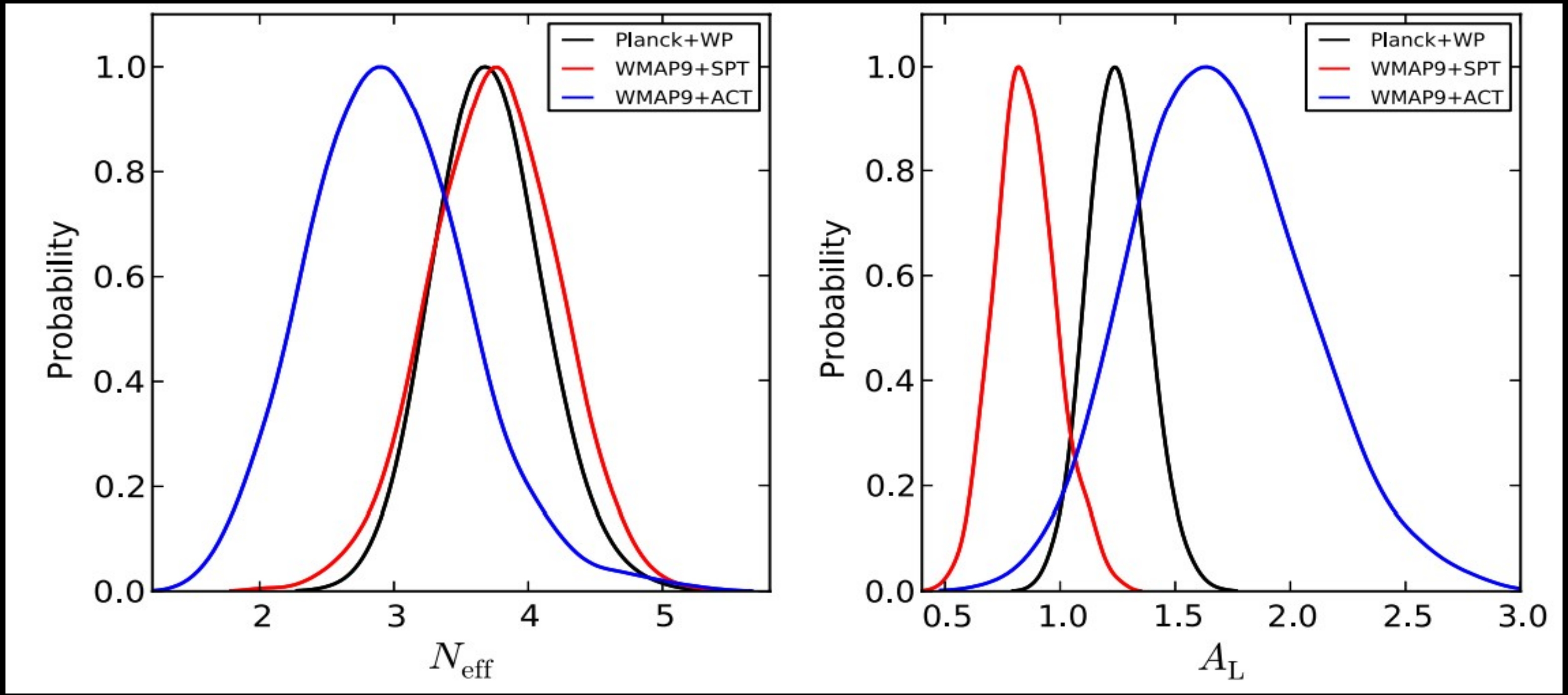
$$N_{\rm eff} = 3.71 \pm 0.40$$

$$A_{\rm L} = 1.25 \pm 0.13$$

(68%; Planck+WP)



N. Said, E. Di Valentino, M. Gerbino, Phys. Rev D, 88, 023513, 2013



Parameter	Planck + WP	WMAP9 + SPT	WMAP9 + ACT
$\Omega_b h^2$	0.02306 ± 0.00051	0.02264 ± 0.00051	0.02295 ± 0.00052
$\Omega_c h^2$	0.1239 ± 0.0054	0.1232 ± 0.0080	0.112 ± 0.011
θ	1.04124 ± 0.00077	1.0415 ± 0.0012	1.0410 ± 0.0025
τ	0.095 ± 0.015	0.088 ± 0.014	0.090 ± 0.015
n_s	0.996 ± 0.018	0.982 ± 0.018	0.975 ± 0.019
$\log[10^{10}A_s]$	3.111 ± 0.034	3.169 ± 0.048	3.083 ± 0.044
N_{eff}	3.71 ± 0.40	3.72 ± 0.46	3.00 ± 0.61
A_{L}	1.25 ± 0.13	0.85 ± 0.13	1.70 ± 0.37
Ω_Λ	0.736 ± 0.022	0.736 ± 0.023	0.731 ± 0.025
t_0[Gyr]	13.08 ± 0.38	13.14 ± 0.43	13.74 ± 0.57
Ω_m	0.264 ± 0.022	0.264 ± 0.023	0.269 ± 0.025
H_0[km/s/Mpc]	74.9 ± 3.7	74.6 ± 3.7	70.9 ± 3.9

E. Di Valentino et al, Phys. Rev D, 88, 023501, 2013

N. Said, E. Di Valentino, M. Gerbino, Phys. Rev D, 88, 023513, 2013

Current Constraints on Dark Radiation from CMB data

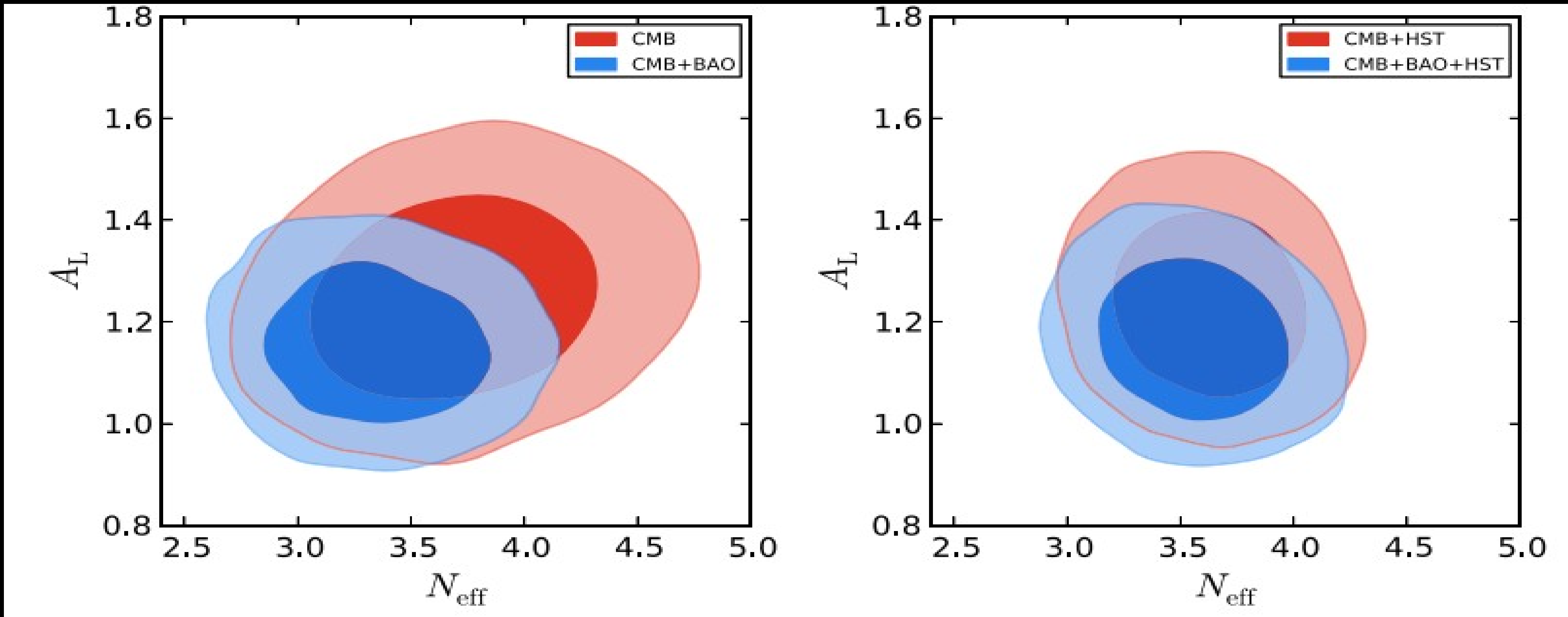
Parameter	WMAP7 + SPT [11]	WMAP9 + SPT [3]	WMAP7 + ACT [10]	WMAP9 + ACT [3]	Planck + WP [2]	PLANCK + WP
N_{eff}	3.62 ± 0.48	3.72 ± 0.46	2.78 ± 0.55	3.00 ± 0.61	3.51 ± 0.39	3.71 ± 0.40
A_{L}	1.00	0.85 ± 0.13	1.00	1.70 ± 0.37	1.00	1.25 ± 0.13
Parameter	WMAP7 + SPT [12]	WMAP9 + SPT [3]	WMAP7 + ACT [10]	WMAP9 + ACT [3]	Planck + WP [2]	PLANCK + WP
N_{eff}	3.046	3.72 ± 0.46	3.046	3.00 ± 0.61	3.046	3.71 ± 0.40
A_{L}	$0.86^{+0.15}_{-0.13}$	0.85 ± 0.13	1.70 ± 0.38	1.70 ± 0.37	$1.22^{+0.11}_{-0.13}$	1.25 ± 0.13

Allowing variations in A_L increases the mean value for N_{eff} by 5-10%.

There is a small, but not negligible correlation between the two parameters.

N. Said, E. Di Valentino, M. Gerbino, Phys. Rev D, 88, 023513, 2013

What about including HST or BAO data?



Results:

- HST brings N_{eff}>3.046 at more than 95% c.l.
- BAO brings N_{eff}=3.046 in between 68% c.l.
- BAO+HST gives N_{eff}>3.046 at about 95% c.l.

(HST wins over BAO)

We consider recent HST measurements (Riess et al. 2011) of H_0 = (73.8 ± 2.4) km/s/Mpc;

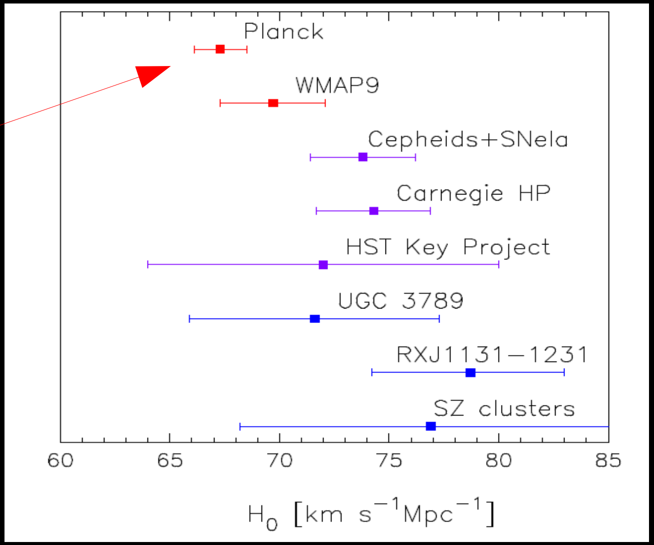
and for BAO surveys we include:

- SDSS-DR7 at redshift z=0.35
- SDSS-DR9 at z=0.57
- WiggleZ at z=0.44, 0.60, and 0.73.

N. Said, E. Di Valentino, M. Gerbino, Phys. Rev D, 88, 023513, 2013

What about H_0?

Moreover, analyzing the Planck data combined with the WMAP9 polarization data, only varying both N_{eff} and A_L we solve the tension to the value of the Hubble constant:



([Planck Collaboration], arXiv:1303.5076)

$$H_0 = (67.3 \pm 1.2)\ \mathrm{km\,s^{-1}\,Mpc^{-1}} \quad (68\%;\ \textit{Planck}+\mathrm{WP}+\mathrm{highL}).$$

$$H_0 = (70.0 \pm 2.2)\ \mathrm{km\,s^{-1}\,Mpc^{-1}} \quad (68\%;\ \textit{WMAP}\text{-}9)$$

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Ω_Λ	0.736 ± 0.022
t_0[Gyr]	13.08 ± 0.38
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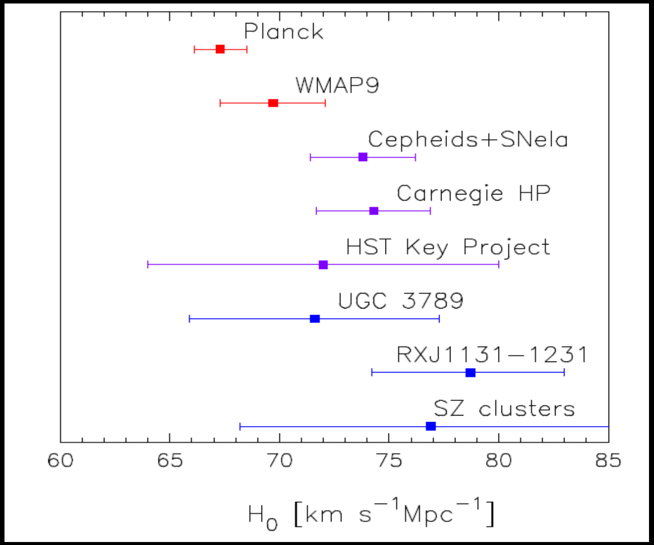
N. Said, E. Di Valentino, M. Gerbino, Phys. Rev D, 88, 023513, 2013



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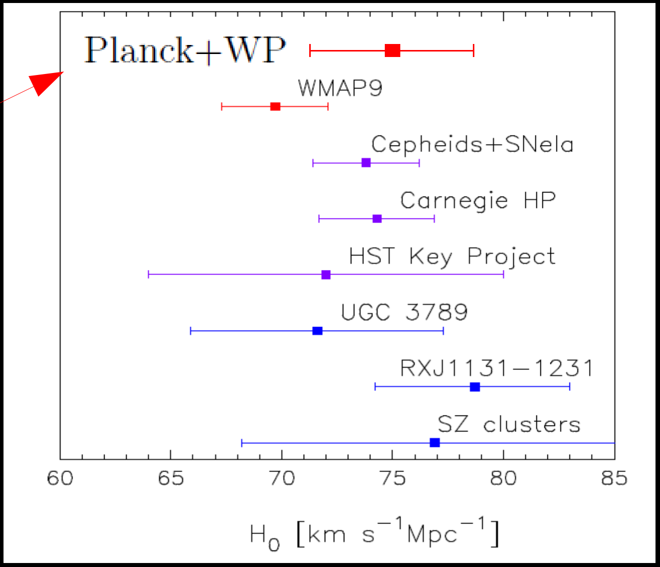


A test for the CNB: Perturbations

The CNB can be further checked by considering two additional parameters: the sound speed in the CNB rest frame c^2_{eff} and the neutrino viscosity c^2_{vis}.

Modelling the CNB as a Generalized Dark Matter (GDM) component (W. Hu 1998), the background evolution is given by the conservation equation:

$$\frac{\dot{\rho}_g}{\rho_g} = -3(1+w_g)\frac{\dot{a}}{a}$$

where $w_g = p_g/\rho_g$ is the equation of state.

The perturbations evolution has only two degrees of freedom. We introduce c^2_{eff} that describes pressure fluctuations respect to density perturbations:

$$w_g \Gamma_g = (c^2_{\rm eff} - c^2_g)\delta_g^{(\rm rest)}$$

and c^2_{vis} that parameterizes the relationship between the anisotropic stress and the metric shear:

$$w_g\left(\dot{\pi}_g + 3\frac{\dot{a}}{a}\pi_g\right) = 4c^2_{\rm vis}(kv_g - \dot{H}_T)$$

A test for the CNB: Perturbations

A combination of these three parameters specifies the clustering properties of GDM, helping to determine its nature from the observational constraints.

We have:

- a scalar-field dark matter $\longrightarrow$ $\left(w_g, c_{eff}^2, c_{vis}^2\right) = (w_g, 1, 0)$
- CDM $\longrightarrow$ $\left(w_g, c_{eff}^2, c_{vis}^2\right) = (0, 0, 0)$
- Radiation $\longrightarrow$ $\left(w_g, c_{eff}^2, c_{vis}^2\right) = \left(\frac{1}{3}, \frac{1}{3}, \frac{1}{3}\right)$

So, for standard neutrinos we expect $\longrightarrow$ $c_{eff}^2 = c_{vis}^2 = \frac{1}{3}$

A test for the CNB: Perturbations

Indeed, we consider the set of Boltzmann equations that describes perturbations in massless neutrino (Archidiacono et al. 2011), to set constraints on the clustering parameters:

$$\dot{\delta}_\nu = \frac{\dot{a}}{a}\left(1-3c_{\rm eff}^2\right)\left(\delta_\nu + 3\frac{\dot{a}}{a}\frac{q_\nu}{k}\right) - k\left(q_\nu + \frac{2}{3k}\dot{h}\right)$$

$$\dot{q}_\nu = k\,c_{\rm eff}^2\left(\delta_\nu + 3\frac{\dot{a}}{a}\frac{q_\nu}{k}\right) - \frac{\dot{a}}{a}q_\nu - \frac{2}{3}k\pi_\nu$$

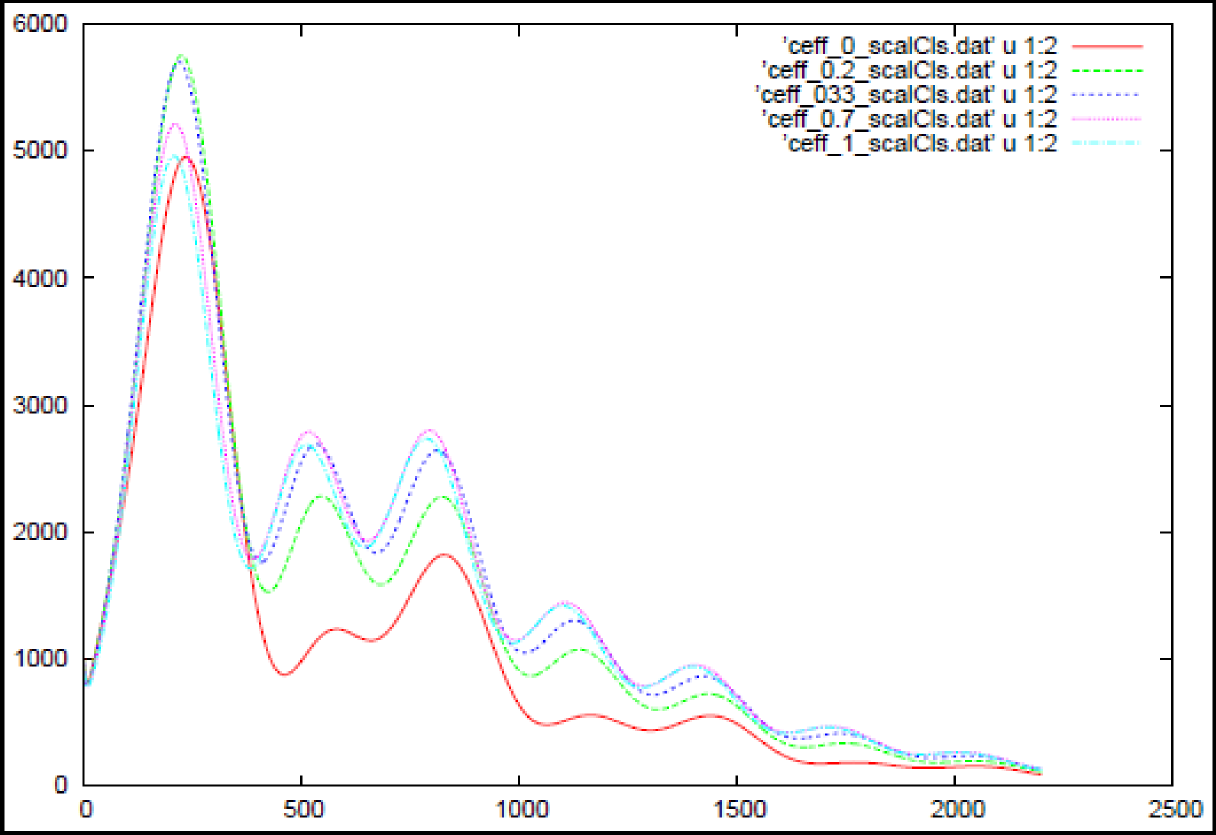
$$\dot{\pi}_\nu = 3\,c_{\rm vis}^2\left(\frac{2}{5}q_\nu + \frac{8}{15}\sigma\right) - \frac{3}{5}kF_{\nu,3}$$

$$\frac{2l+1}{k}\dot{F}_{\nu,l} - lF_{\nu,l-1} = -\left(l+1\right)F_{\nu,l+1}\quad l\geq 3$$

The effect on the CMB:

The increasing of the c^2_{eff} affects the CMB power spectrum with:

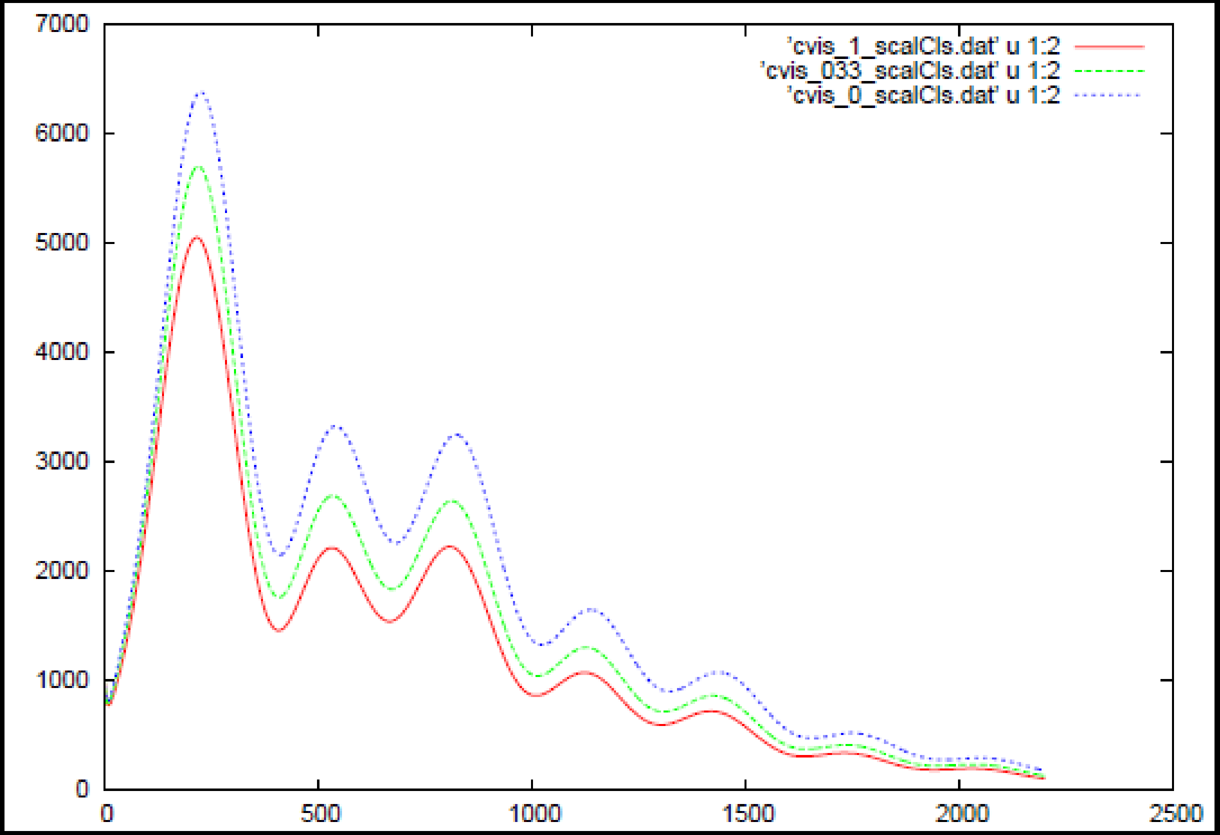
- The increasing of the even peaks amplitude;
- A slight shifting of the peaks towards smaller multipoles.



The effect on the CMB:

The increasing of the c^2_{vis} affects the CMB power spectrum with:

- The decreasing of the peaks amplitude.



CNB Perturbations: Results

Assuming N_{eff}=3.046, Planck+WP suggests a higher value of the viscosity parameter c^2_{vis}, in tension with the standard value at about 1.5 standard deviations, and a lower value of the sound speed c^2_{eff}, ruling out the standard value at more than 95% c.l.:

Parameter	ΛCDM	$+c^2_{\rm vis} + c^2_{\rm eff}$
$100\,\Omega_b h^2$	2.206 ± 0.028	2.118 ± 0.047
$\Omega_c h^2$	0.1199 ± 0.0027	0.1157 ± 0.0038
$100\,\theta$	1.0413 ± 0.0006	1.0412 ± 0.0014
$\log[10^{10} A_S]$	3.089 ± 0.025	3.173 ± 0.052
τ	0.090 ± 0.013	0.089 ± 0.013
n_S	0.9606 ± 0.0073	0.998 ± 0.018
$A_{\rm L}$	$\equiv 1$	$\equiv 1$
$c^2_{\rm vis}$	$\equiv 0.33$	0.60 ± 0.18
$c^2_{\rm eff}$	$\equiv 0.33$	0.304 ± 0.013
H_0 [(a)]	67.3 ± 1.2	68.0 ± 1.3

M. Gerbino, E. Di Valentino, N. Said, PRD in press 2013

CNB Perturbations: Results

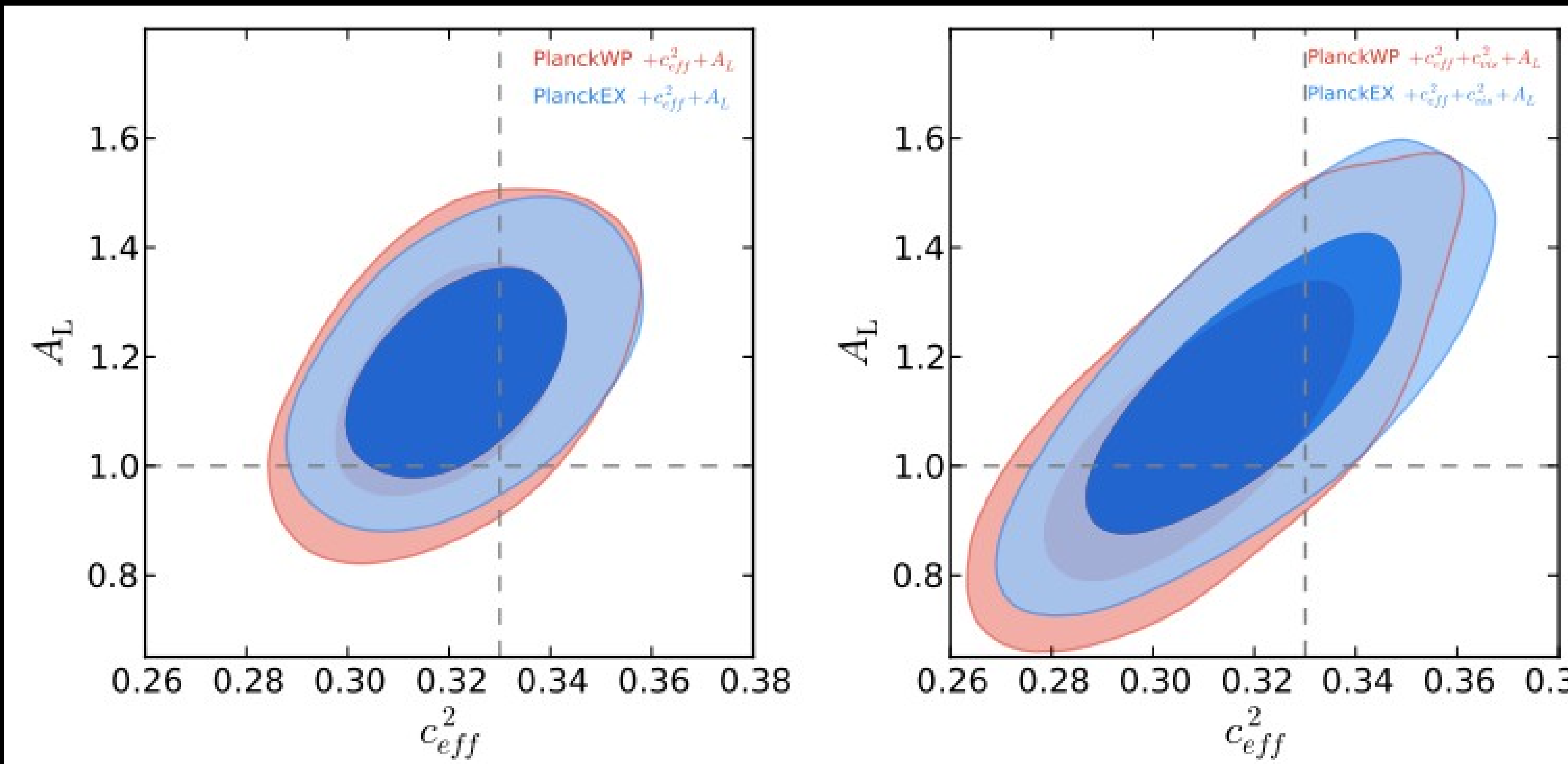
We find a correlation between the neutrino parameters and the lensing amplitude:

Parameter	ΛCDM	$+c^2_{\rm eff} + c^2_{\rm vis} + A_{\rm L}$
$100\,\Omega_b h^2$	2.206 ± 0.028	2.162 ± 0.095
$\Omega_c h^2$	0.1199 ± 0.0027	0.1159 ± 0.0036
$100\,\theta$	1.0413 ± 0.0006	1.0420 ± 0.0020
$\log[10^{10} A_S]$	3.089 ± 0.025	3.141 ± 0.078
τ	0.090 ± 0.013	0.089 ± 0.014
n_S	0.9606 ± 0.0073	0.989 ± 0.023
$A_{\rm L}$	$\equiv 1$	1.08 ± 0.18
$c^2_{\rm vis}$	$\equiv 0.33$	0.51 ± 0.22
$c^2_{\rm eff}$	$\equiv 0.33$	0.311 ± 0.019
H_0 $^{(a)}$	67.3 ± 1.2	68.6 ± 1.7

M. Gerbino, E. Di Valentino, N. Said, PRD in press 2013

CNB Perturbations: Results

We find a correlation between the neutrino parameters and the lensing amplitude:

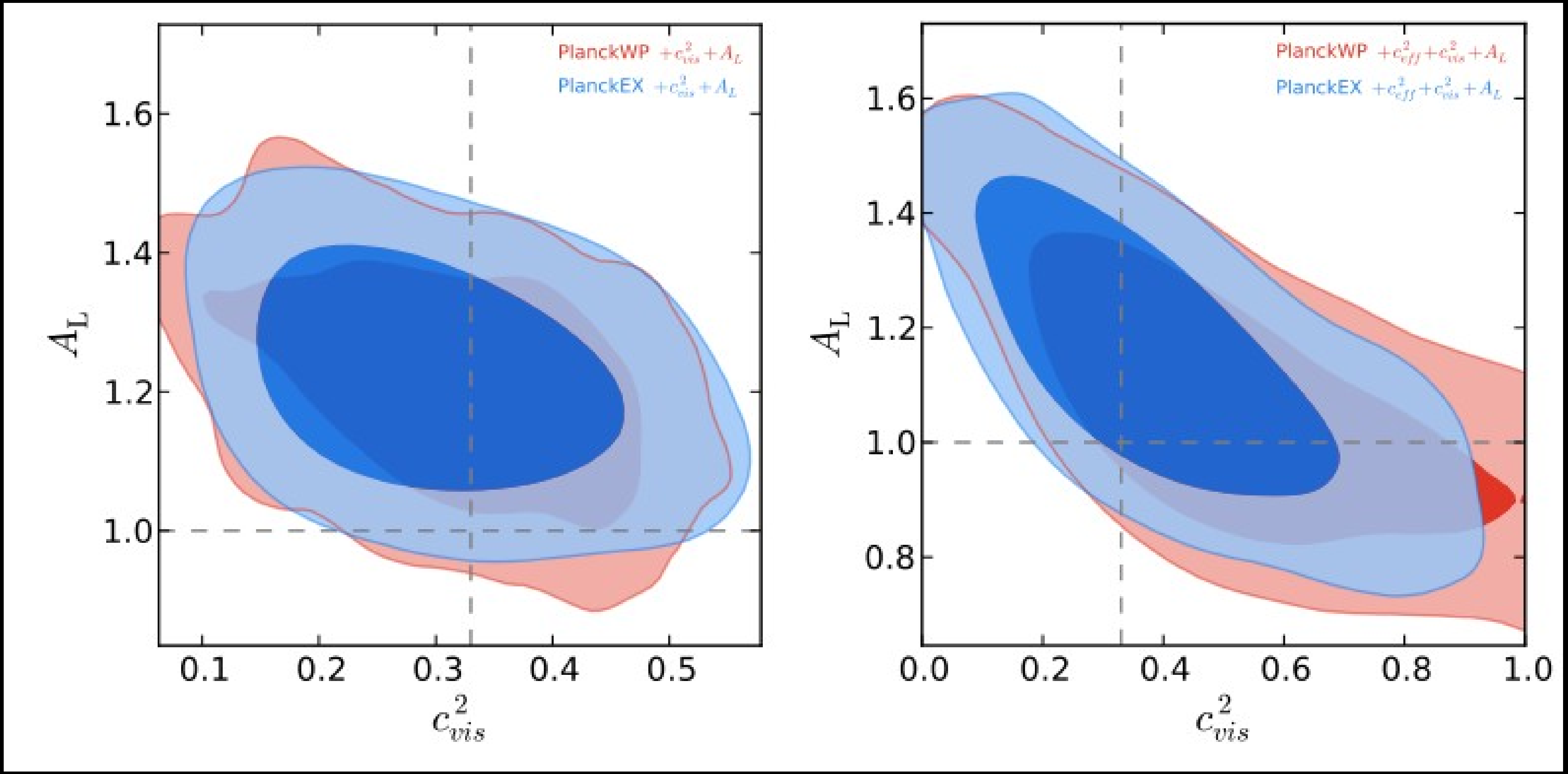


A_L=1 can be in more agreement with data with a lower c^2_{eff}.

M. Gerbino, E. Di Valentino, N. Said, PRD in press 2013

CNB Perturbations: Results

We find a correlation between the neutrino parameters and the lensing amplitude:

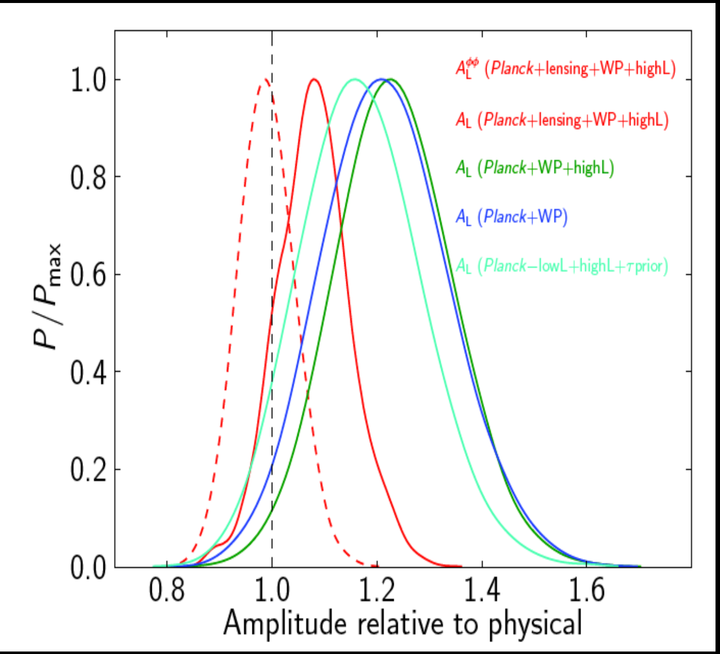


A_L=1 can be in more agreement with data with a higher c^2_{vis}.

M. Gerbino, E. Di Valentino, N. Said, PRD in press 2013

CNB Perturbations: Results

If we allow the CNB clustering parameters to vary, the anomalous large value of A_L measured by Planck disappears..



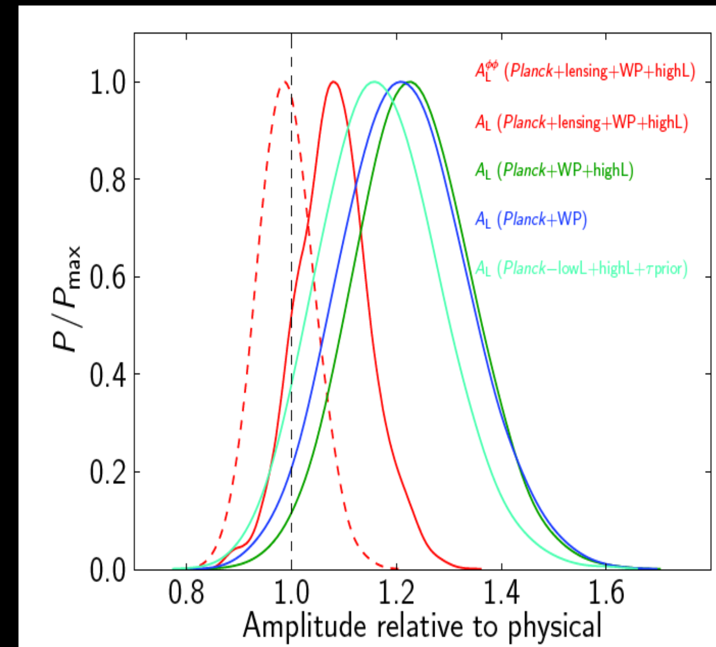
$$A_L^{\phi\phi} = 0.99 \pm 0.05 \quad (68\%; \textit{Planck}+\text{lensing}+\text{WP}+\text{highL})$$

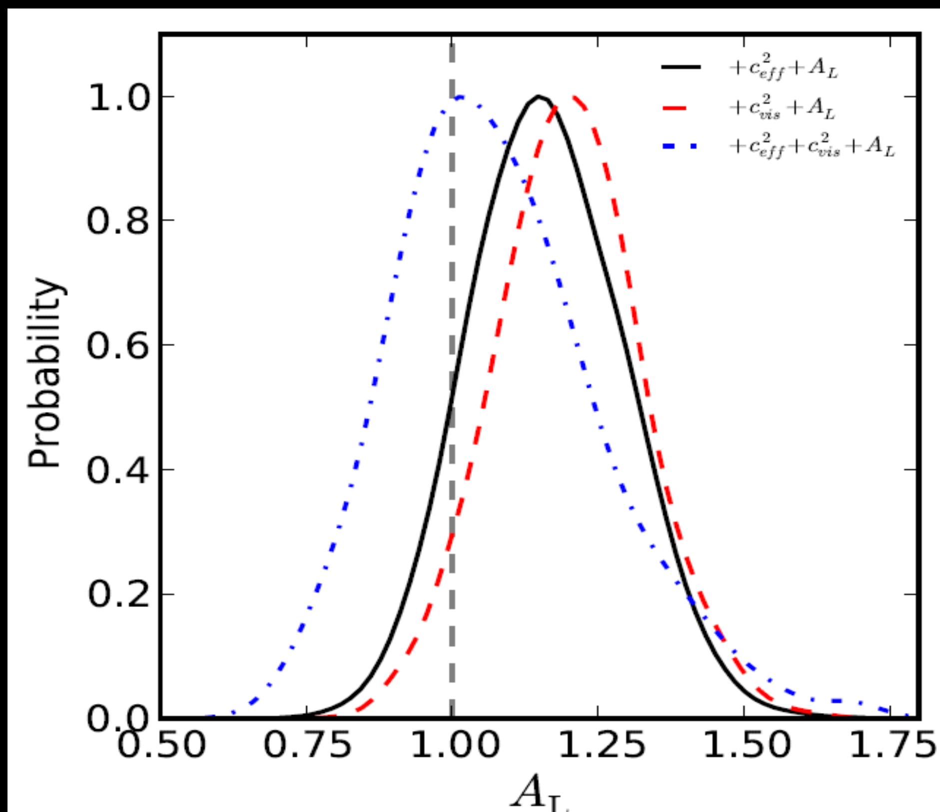
$$A_L = 1.23 \pm 0.11 \quad (68\%; \textit{Planck}+\text{WP}+\text{highL})$$

M. Gerbino, E. Di Valentino, N. Said, PRD in press 2013

CNB Perturbations: Results

If we allow the CNB clustering parameters to vary, the anomalous large value of AL measured by Planck disappears..





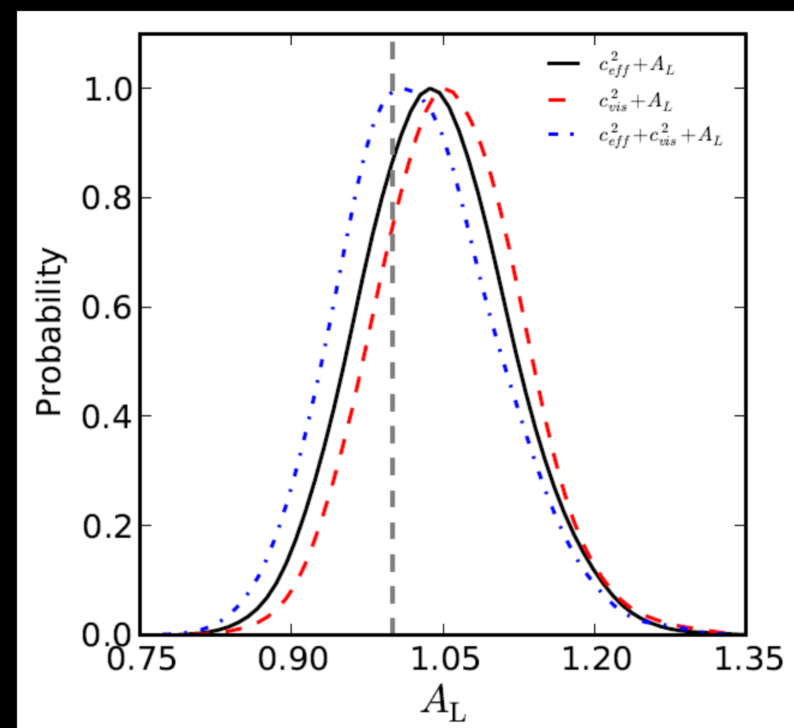
$$A_{\mathrm{L}} = 1.08 \pm 0.18 \text{ at } 68\% \text{ c.l.}$$

M. Gerbino, E. Di Valentino, N. Said, PRD in press 2013

CNB Perturbations: Results

And the same result is obtained adding the lensing likelihood to the Planck+WP dataset:

Parameter	$+c^2_{\rm vis} + c^2_{\rm eff}$	$+c^2_{\rm eff} + c^2_{\rm vis} + A_{\rm L}$
$100\,\Omega_b h^2$	2.155 ± 0.049	2.166 ± 0.063
$\Omega_c h^2$	0.1151 ± 0.0034	0.1151 ± 0.0037
$100\,\theta$	1.0414 ± 0.0015	1.0415 ± 0.0017
$\log[10^{10} A_S]$	3.146 ± 0.054	3.137 ± 0.064
τ	0.090 ± 0.013	0.088 ± 0.013
n_S	0.990 ± 0.019	0.989 ± 0.021
$A_{\rm L}$	$\equiv 1$	1.025 ± 0.076
$c^2_{\rm vis}$	0.52 ± 0.18	0.50 ± 0.19
$c^2_{\rm eff}$	0.312 ± 0.013	0.314 ± 0.015
H_0 $^{(a)}$	68.5 ± 1.1	68.8 ± 1.5

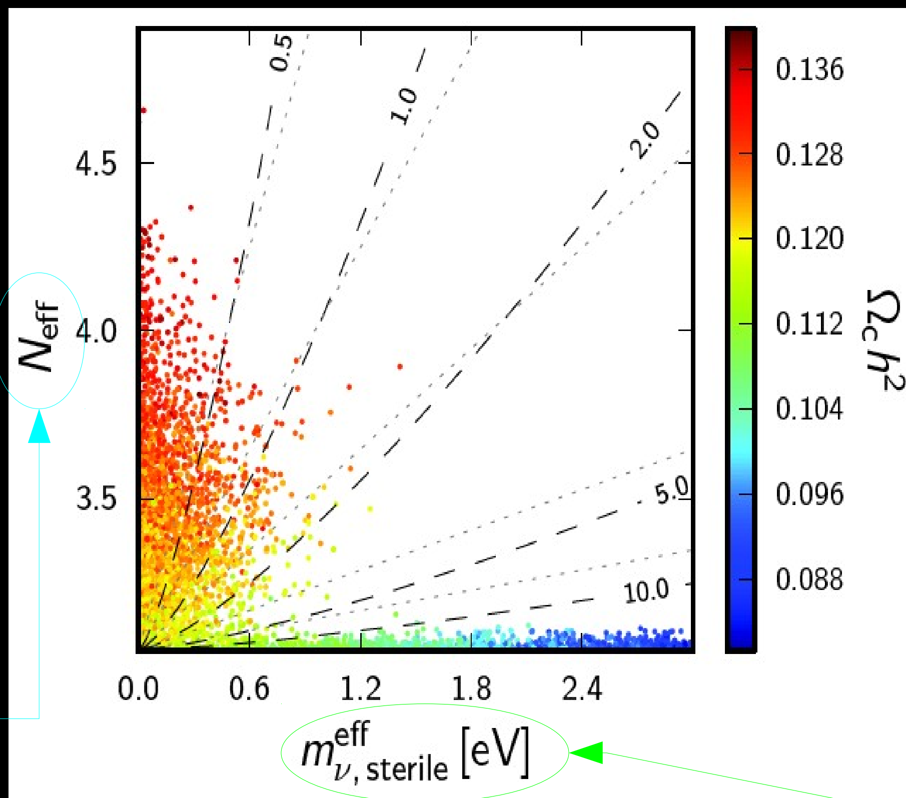


M. Gerbino, E. Di Valentino, N. Said, PRD in press 2013

The effective sterile neutrino mass

We can constrain simultaneously the effective sterile neutrino mass and N_{eff}. Their relationship is strongly model dependent, but fixed the model we can infer the physical mass of the particle.

Contribution of the sterile neutrino when it is massless.



Thermally distributed

$$m^{eff}_{\nu,\,sterile} = (T_s/T_\nu)^3 m^{thermal}_{sterile} = (\Delta N_{eff})^{3/4} m^{thermal}_{sterile}$$

Distributed proportionally to active neutrinos (Dodelson-Widrow model)

$$m^{eff}_{\nu,\,sterile} = \chi_s m^{DW}_{sterile}$$

$$\Delta N_{eff} = \chi_s$$

Contribution of the sterile neutrino when it is massive.

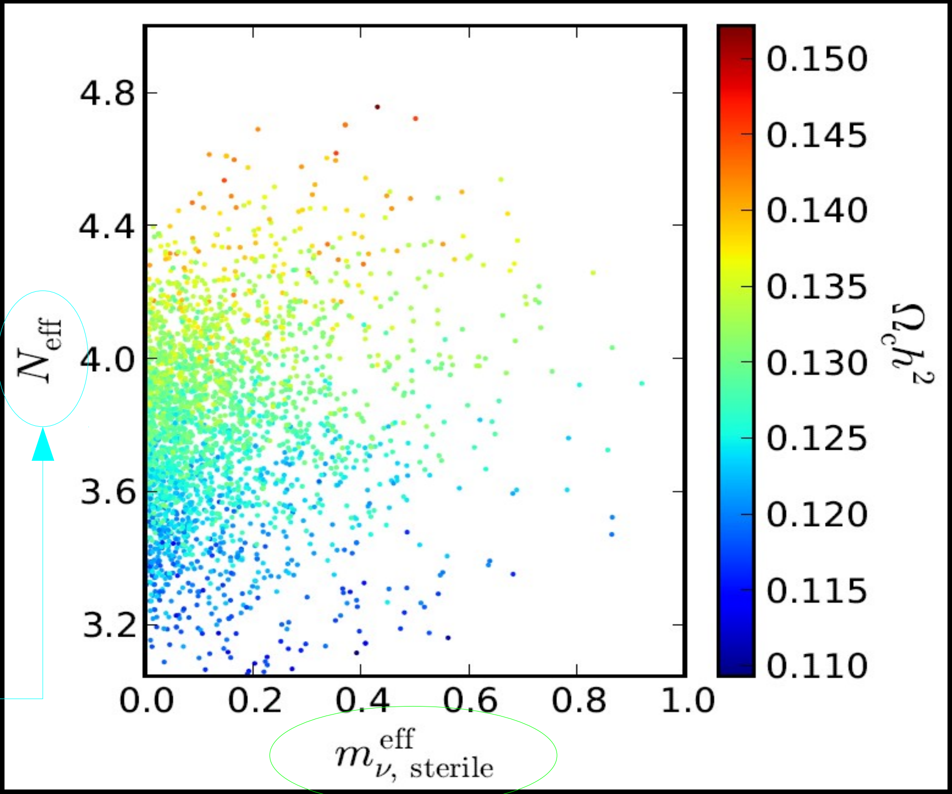
$$\left.\begin{aligned} N_{eff} &< 3.91 \\ m^{eff}_{\nu,\,sterile} &< 0.59\,\mathrm{eV} \end{aligned}\right\} \quad (95\%;\ \text{CMB for } m^{thermal}_{sterile} < 10\,\mathrm{eV})$$

(Planck collaboration 2013)

The effective sterile neutrino mass

Adding the HST measurements we obtain stronger bounds on these parameters, and the standard value N_{eff}=3.046 is excluded at more than 2 sigma.

Contribution of the sterile neutrino when it is massive.



Contribution of the sterile neutrino when it is massive.

$$3.25 < N_{\rm eff} < 4.37$$

$$m^{\rm eff}_{\nu,\rm sterile} < 0.44\ \rm eV$$

(95%; Planck+HST)

Dark radiation models

When we combine the Planck Satellite data with HST measurements we have at 95% c.l.:

$$N_{\rm eff} = 3.83 \pm 0.54$$

And when low multipole polarization measurements from the WMAP 9 years data release and high multipole CMB data from both ACT and SPT are added in the analysis, we find a 95% c.l.:

$$N_{\rm eff} = 3.62^{+0.50}_{-0.48}$$

These bounds indicate the presence of an extra dark radiation $\Delta N_{\rm eff} \equiv N_{\rm eff} - 3.046$ at the $\sim 2.4\sigma$ confidence level and can be exploited to set limits on any model containing extra dark radiation species.

Light Sterile Neutrino Model

The simplest way to explain the extra dark radiation is to include extra sterile neutrino species.

We focus here on the so-called (3+1) neutrino mass model.

If the sterile neutrino never reaches a complete thermalization, then its abundance is much lower than the thermal one and depends dramatically on the flavour mixing processes operating at the decoupling period.

We consider small mixing both between the active and heavy sectors and between the sterile and light neutrino sectors.

The sterile neutrino contributes to the energy density of the Universe (A. Melchiorri et al. 2008):

$$\Omega_s h^2 \simeq 7 \times 10^{-5} \left(\frac{\Delta m^2_{41}}{eV^2}\right) \sum_a \frac{g_a}{\sqrt{C_a}} \left(\frac{U_{a4}}{10^{-2}}\right)^2$$

and to ΔN_{eff}:

$$\Delta N_{\rm eff} = \frac{\Omega_s h^2}{\frac{7}{8}\left(\frac{4}{11}\right)^{\frac{4}{3}} \Omega_\gamma h^2}$$

(3+1) neutrino mass model

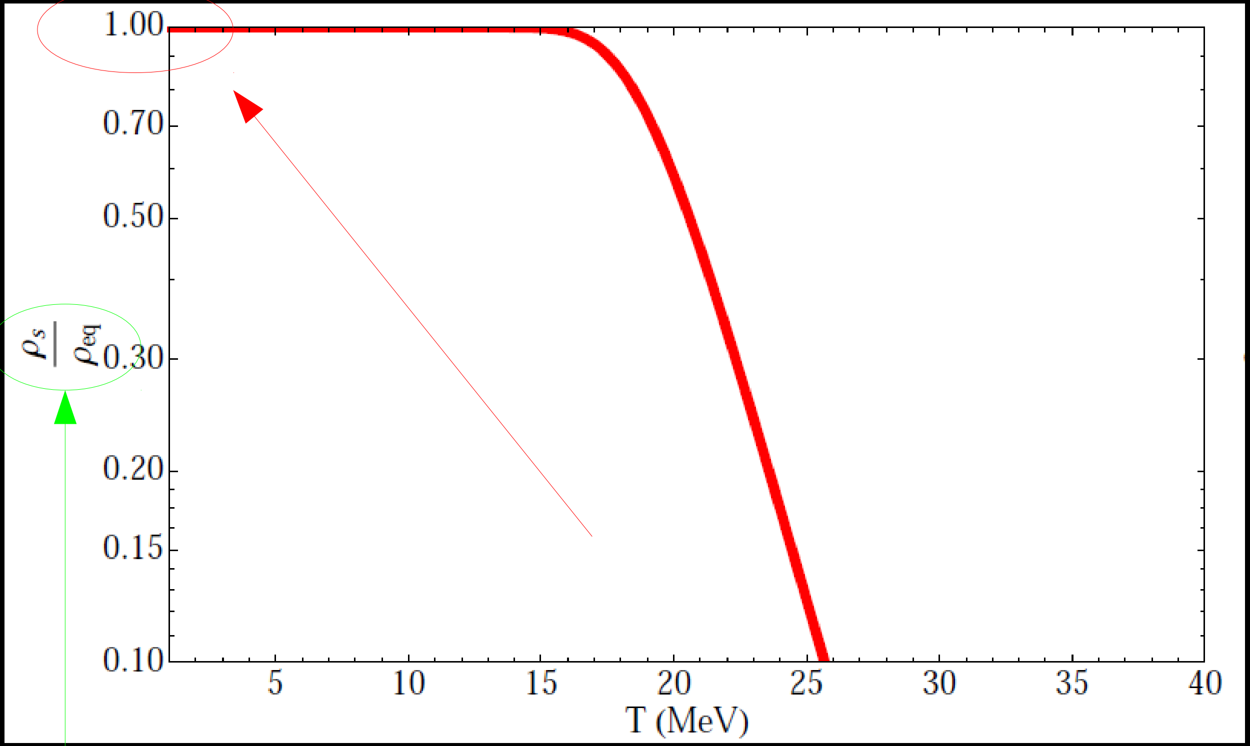
We computed the sterile neutrino abundance at its decoupling from T=100 MeV up to T = 1 MeV.

$$\left(\frac{\partial\rho}{\partial T}\right) = -\frac{1}{HT}\left(i\left[H_{\rm m} + V_{\rm eff}, \rho\right] - \{\Gamma, (\rho - \rho_{\rm eq})\}\right)$$

We have assumed:

- for the active neutrino mixing parameters, the global fit from M. C. Gonzales-Garcia et al. (2012);
- for the sterile neutrino mixing parameters the global fit from neutrino oscillation data (J. M. Conrad et al. 2013): Ue4 = 0.14, Uμ4 = 0.17 and $\Delta m^2 14$ = 0.93 eV^2;
- a normal hierarchy scheme with m1=0;
- the sum of the active neutrino masses equal to 0.056.

We normalized the sterile neutrino abundance to the equilibrium distribution.



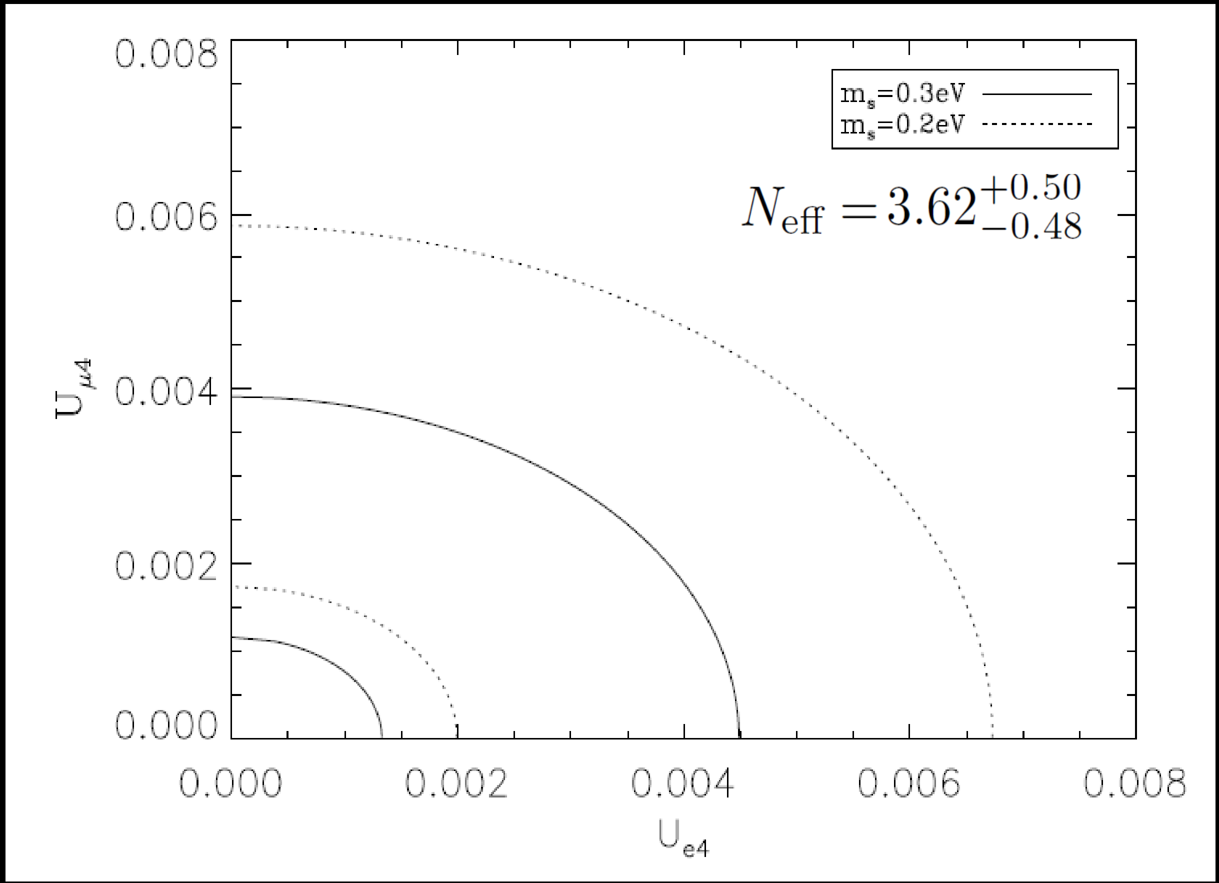
(E. Di Valentino, A. Melchiorri and O. Mena, arXiv:1304.5981)

(3+1) neutrino mass model

We can set constraints on the sterile neutrino mixing parameters from the recent ΔN_{eff} measurements, for a given value of the sterile neutrino mass m <~0.3 eV, setting $U_{\tau 4}=0$.

We find that the relatively large values of the sterile neutrino mixing parameters preferred by <u>short baseline oscillation data in (3 + 1) model</u> (J. M. Conrad et al. 2013) (Δm^2_{14} = 0.15 eV^2, U_{e4} = 0.39 and $U_{\mu 4}$ = 0.39) <u>are excluded</u> for 0.2 eV < m_s < 0.3 eV.

For lower sterile neutrino masses m_s < 0.1 eV we have higher mixing parameters, but this mass is highly disfavored by oscillation analyses.



(E. Di Valentino, A. Melchiorri and O. Mena, arXiv:1304.5981)

Extended dark sector models

These contain a dark sector, with relativistic degrees of freedom, that eventually decouple from the standard model sector contributing to N_{eff}.

We consider the so-called asymmetric dark matter scenario, in which the extra degrees of freedom are produced by the annihilations of the thermal symmetric dark matter component. The dark sector contains both light (g_ℓ) and heavy (g_h) degrees of freedom at the temperature of decoupling T_d from the standard model .

For higher T_d > MeV, the contribution to N_{eff} will be (M. Blennow et al. 2012):

$$\Delta N_{\rm eff} = \frac{13.56}{g_{\star S}(T_D)^{\frac{4}{3}}} \frac{(g_\ell + g_h)^{\frac{4}{3}}}{g_\ell^{\frac{1}{3}}}$$

While for lower temperatures (T_d < MeV), if the dark sector couples to neutrinos, we have:

$$N_{\rm eff} = (3 + \frac{4}{7}\frac{(g_h + g_\ell)^{\frac{4}{3}}}{g_\ell^{\frac{1}{3}}})(\frac{3\times\frac{7}{4} + g_H + g_h + g_\ell}{3\times\frac{7}{4} + g_h + g_\ell})^{\frac{4}{3}}$$

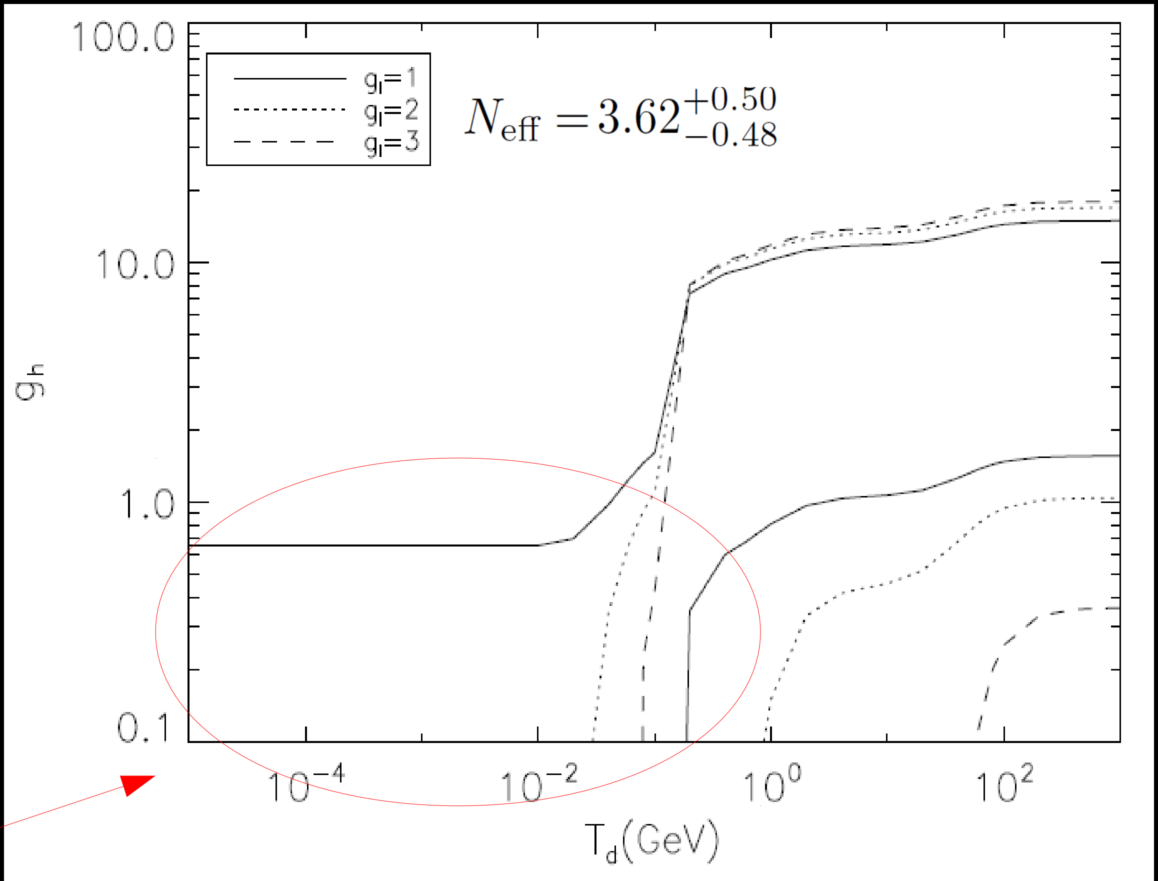
where g_H the number of degrees of freedom that become non relativistic between BBN and the dark sector decoupling period.

Asymmetric dark matter scenario

We can use the measured value of N_{eff} to find the required g_h that heat the dark radiation plasma as a function of T_d, for a fixed value of g_ℓ, for $g_H = 0$.

For T_d > MeV, in order to have a larger value of N_{eff}, the standard model relativistic degrees of freedom will be heated, requiring therefore heating in the dark sector, then the larger values of g_h.

On the other hand, at lower decoupling temperatures, the required g_h decrease as N_{eff} does, so the extra heavy degrees of freedom are disfavoured.



(E. Di Valentino, A. Melchiorri and O. Mena, arXiv:1304.5981)

Neutrino isocurvature

Another indipendent way to investigate the value of N_{eff}, is to consider the neutrino isocurvature modes.

The presence of neutrino isocurvature could bias the value of N_{eff}, and implies a non-zero chemical potential in the neutrino background distribution.

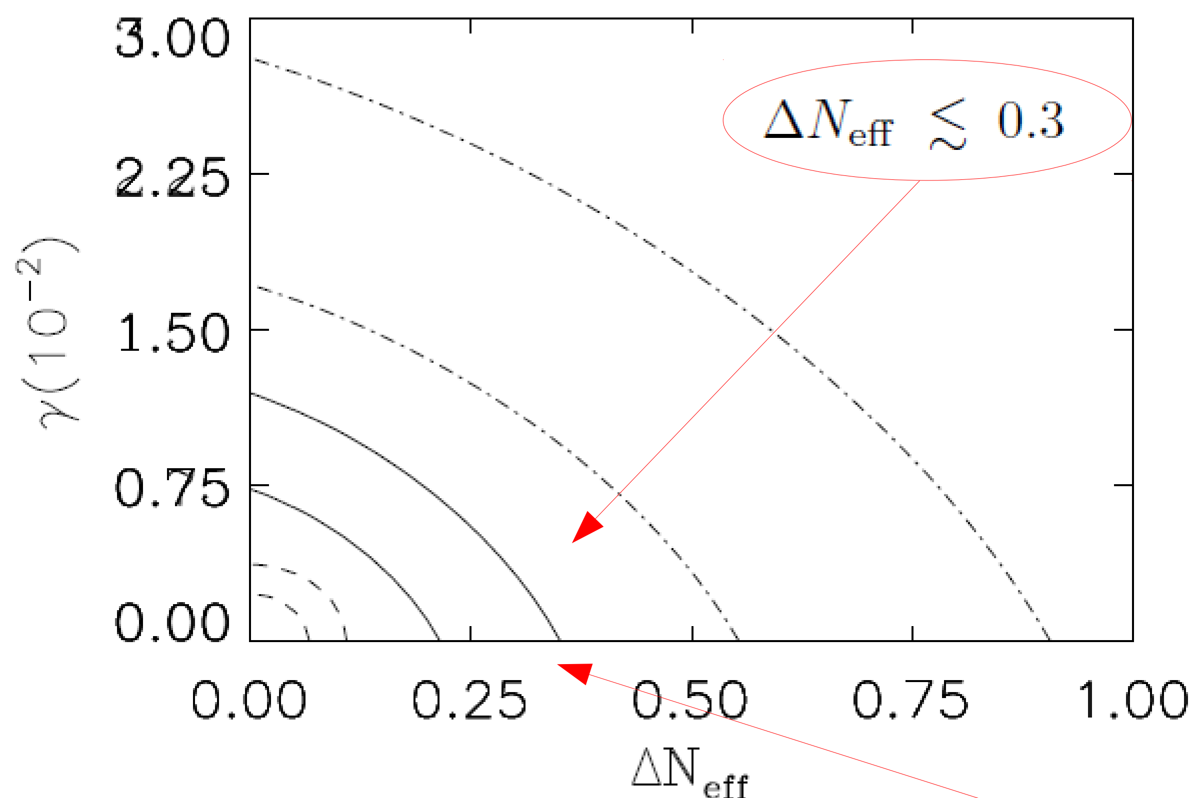
We have forecasted in E. Di Valentino, et al., Phys. Rev. D85 (2012) 043511 the possibility for Planck satellite to constrain simultaneosly the amplitude of neutrino isocurvature perturbations and the extra energy density, parametrized by N_{eff}, associated to the neutrino chemical potential in the curvaton scenario. After the decay of the curvaton field, residual isocurvature perturbations can be imprinted in the neutrino component of the cosmological fluid.

We expect the release of the Planck high frequency polarization likelihood to update this work.

Neutrino isocurvature

The total CMB power spectrum can be parameterized in terms of the adiabatic, neutrino isocurvature density and totally anticorrelated spectra:

$$\boxed{C_\ell = (1-\gamma^2)C_\ell^{ad} + \gamma^2 C_\ell^{nid} - 2\gamma\sqrt{(1-\gamma^2)}C_\ell^{corr}}$$



68% and 95% c.l. likelihood contours for Planck (solid line), SPIDER (dot-dashed line) CMBPol (dashed line).

E. Di Valentino, et al., Phys.Rev. D85 (2012) 043511

The CMB lensing bispectrum

Another indipendent way to investigate the value of N_{eff} is to consider the Bispectrum signal.

The CMB Bispectrum includes a lot of cosmological informations, thanks to the cross-correlation between the Integrated Sachs-Wolfe effect (ISW) and the gravitational Weack Lensing (WL).

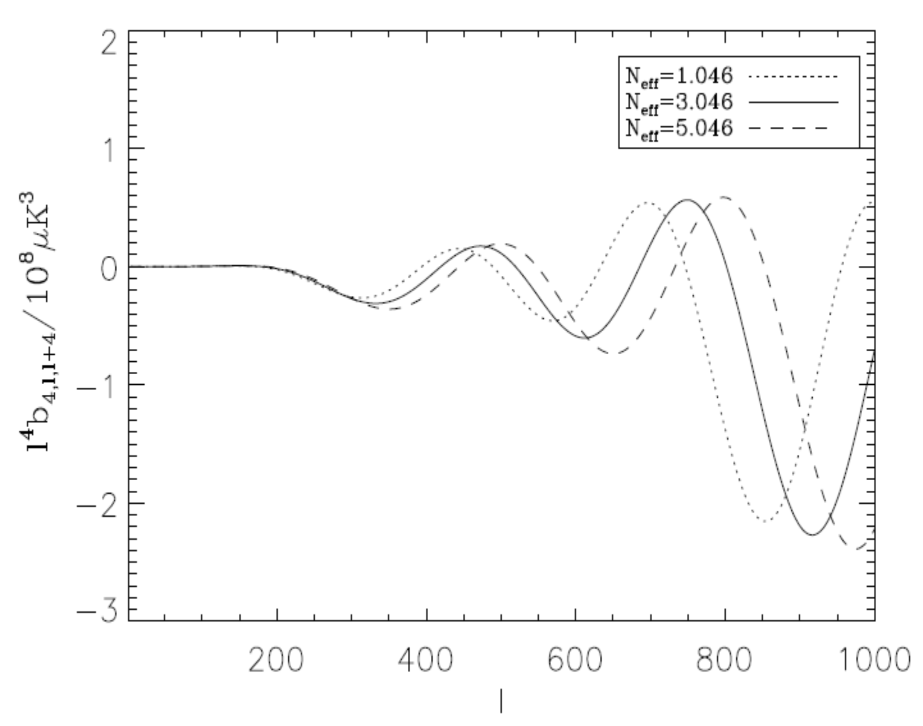
- ISW: Photons change their frequency crossing a hole gravitational potential;
- WL: Deflection path photons due to the presence of large scale structures.

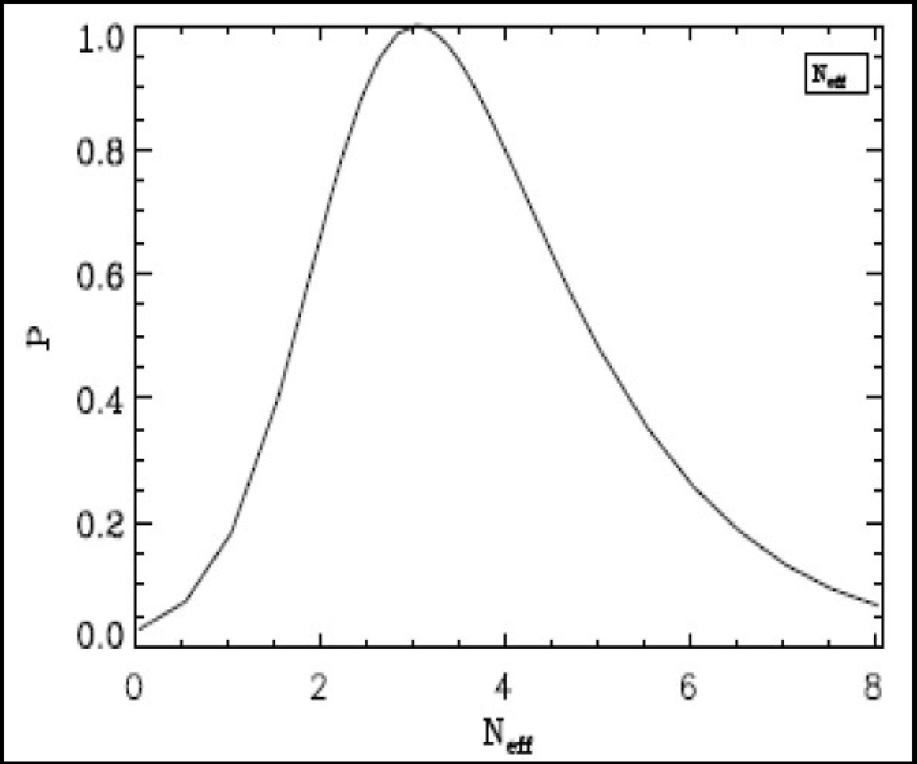
We have forecasted the possibility for the Planck satellite to constrain, using the bispectrum signal, N_{eff} (E. Di Valentino, et al., Phys. Rev. D 87, 103523, 2013), and some modified gravity models (E. Di Valentino, et al., Phys. Rev. D86 (2012) 063517), and we expect the release of the Planck polarization data to update these works.

The CMB lensing bispectrum

The bispectrum is sensitive to N_{eff}, and we found that the ability of the Planck satellite to constrain it will be:

$$2.0 < N_{eff} < 4.6 \text{ at } 68\% \text{ c.l.}$$





E. Di Valentino, et al., Phys. Rev. D 87, 103523, 2013

Non-Gaussianity

The estimators of local non-Gaussianity f_{NL}, that may be subtracted for an unbiased detection of primordial non-Gaussianity, change significantly if we consider the presence of a non-standard neutrino background, as reported in the table.

Model	σ_{fnl}	σ_{lens}	correlation	bias on f_{NL}	σ_{fnl}^{marge}
$N_{eff}^{rel} = 3.046$ $\sum m_\nu = 0$	4.33	0.18	0.24	9.7	4.47
$N_{eff}^{rel} = 0.046$ $\sum m_\nu = 0$	4.40	0.16	0.24	12.5	4.54
$N_{eff}^{rel} = 5.046$ $\sum m_\nu = 0$	4.30	0.19	0.25	9.3	4.44
$N_{eff}^{rel} = 0.046$ $N_{eff}^{mass} = 3$ $\sum m_\nu = 1eV$	4.17	0.22	0.23	7.5	4.29
$N_{eff}^{rel} = 0.046$ $N_{eff}^{mass} = 4$ $\sum m_\nu = 2eV$	4.13	0.24	0.24	7.1	4.26

E. Di Valentino, et al., Phys. Rev. D 87, 103523, 2013

Conclusions:

- The Planck experiment doesn't solve the tension between the ACT and SPT data:

Parameter	Planck + WP	CMB + HST	CMB + BAO	CMB + BAO + HST
$N_{\rm eff}$	3.71 ± 0.40	3.63 ± 0.27	3.35 ± 0.31	3.56 ± 0.27
$A_{\rm L}$	1.25 ± 0.13	1.24 ± 0.12	1.16 ± 0.10	1.17 ± 0.10

- The large anomalous value of the A_L can be explained with the CNB clustering parameters.

Parameter	$+c^2_{\rm eff} + c^2_{\rm vis} + A_{\rm L}$
$A_{\rm L}$	1.08 ± 0.18
$c^2_{\rm vis}$	0.51 ± 0.22
$c^2_{\rm eff}$	0.311 ± 0.019

- Adding the HST data we obtain an evidence at more than 2 sigma for the dark radiation, and we can costrain a couple of models.
- We can use neutrino isocurvature modes and the bispectrum signal to constrain indipendently the value of $N_{\rm eff}$.

The next Planck data release, expected around summer 2014, will be decisive in confirming or ruling out the dark radiation.

Thank you!

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