Neutrino oscillations

Pablo del Amo Sanchez
GraSPA
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Overview

Non historical approach: minimal effort

- Atmospheric neutrinos: SK
- The saga of Solar neutrinos
- Closing the trilogy: reactor neutrinos
- Teaser: LBNO/LBNE

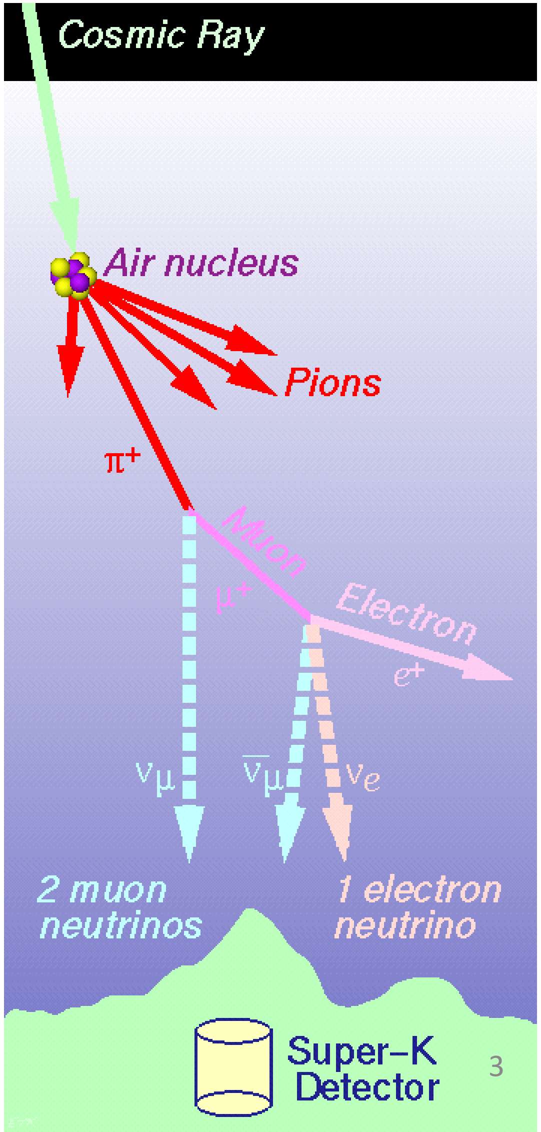
Atmospheric neutrinos

- Cosmic rays collisions in upper atmosphere (15 km)
- Plenty of pions from hadronic interactions
- $\pi^+ \rightarrow \mu^+\nu_\mu$ and $\mu^+ \rightarrow e^+\nu_e\bar{\nu}_\mu$

so

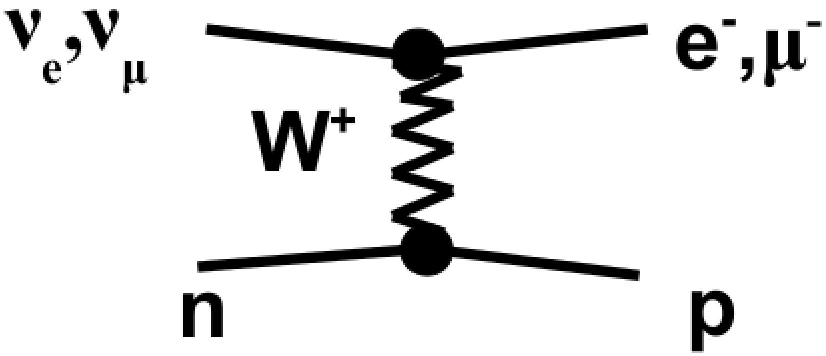
$$\nu_\mu : \nu_e = 2 : 1$$

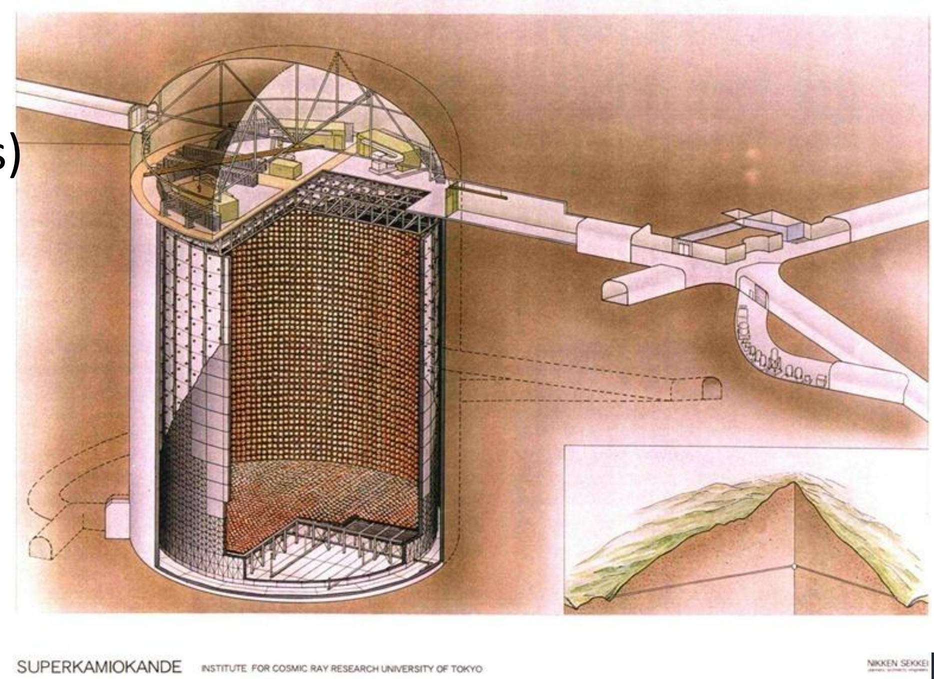
(known better than 3% below 5 GeV)



Water Cerenkov detectors

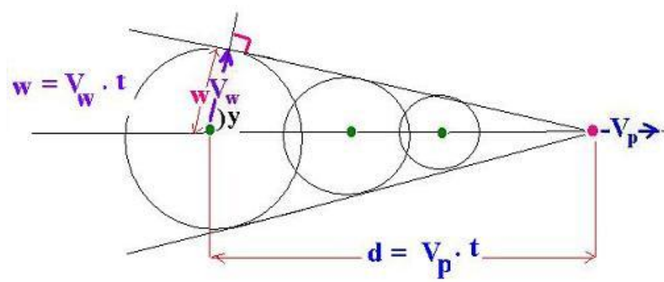
- Huge underground water tanks surrounded by photomultiplier tubes (PMTs)





- Interacting particles produce light, light gives electrical signal in PMTs

Cerenkov effect: particles faster than speed of light in medium radiate light (e.g. blueish light in nuclear reactors)



- Ex: (Super-)KamiokaNDE et SNO

Water Cerenkov detectors

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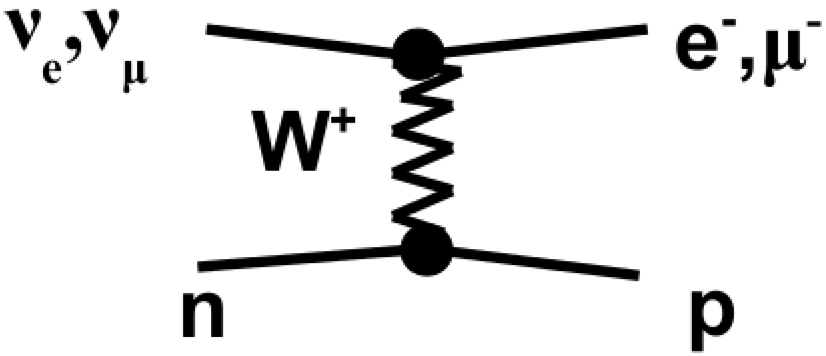

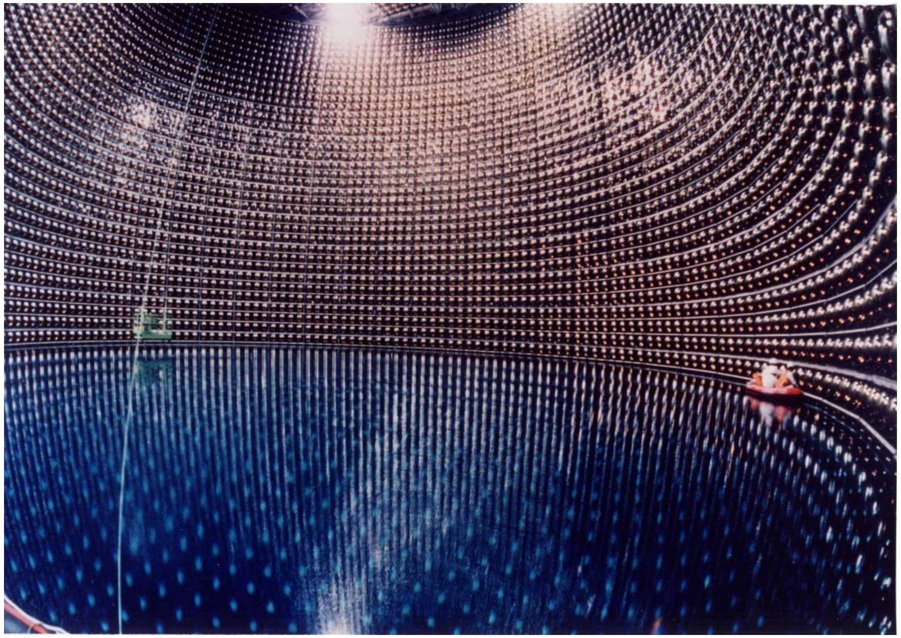

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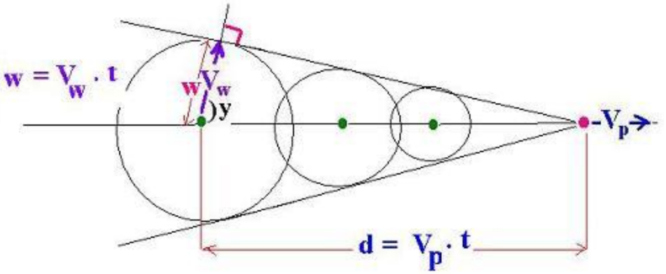
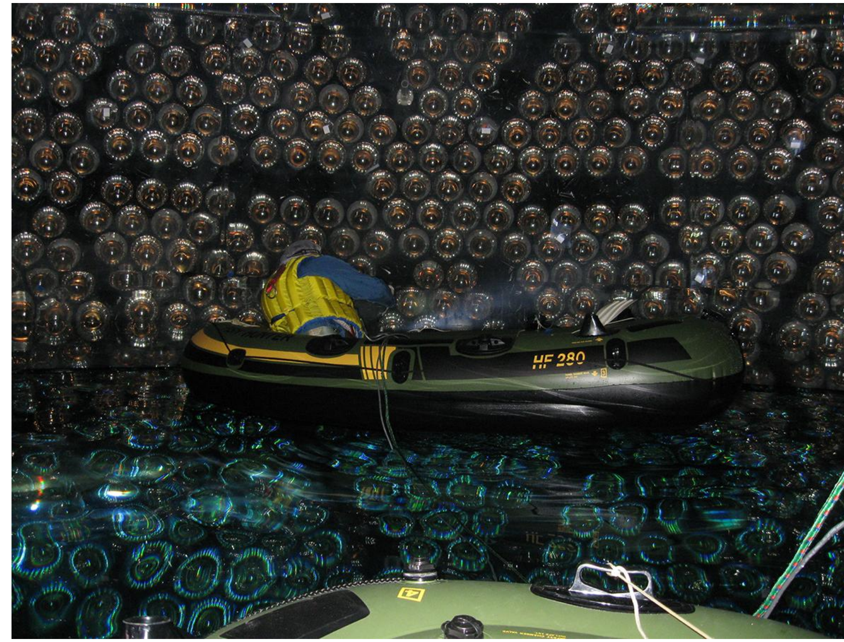
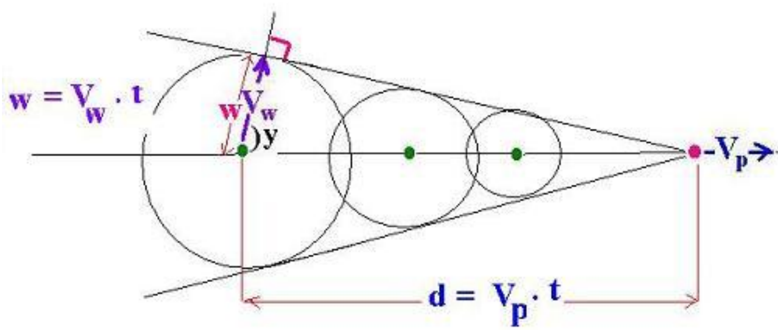


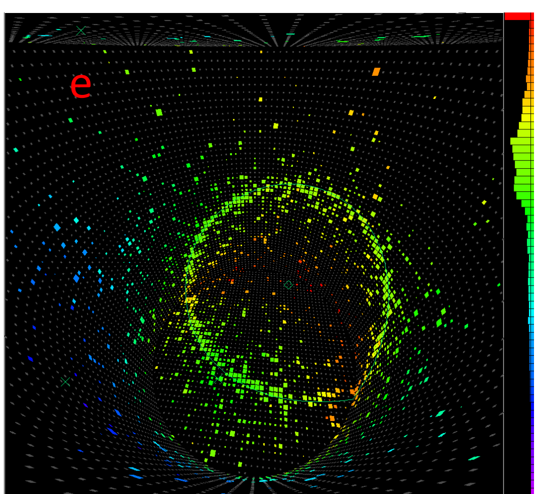
- Ex: (Super-)KamiokaNDE et SNO

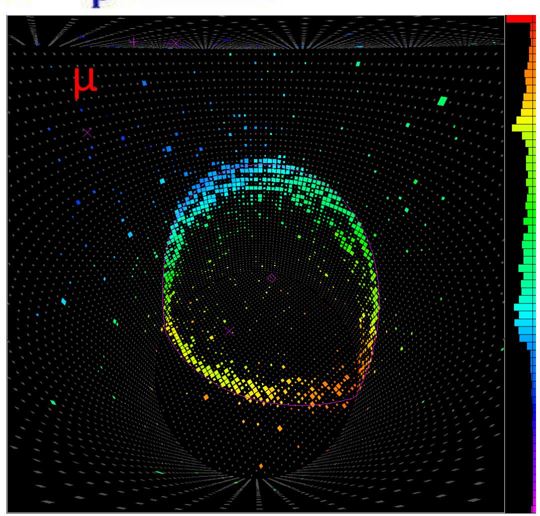
Water Cerenkov detectors

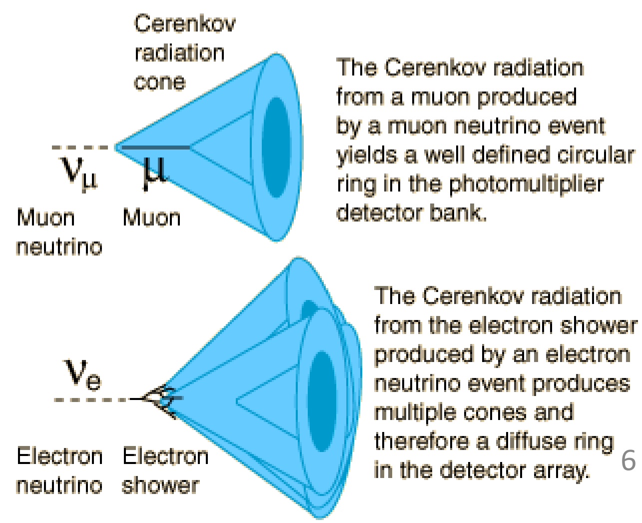
- SNO et (Super-)KamiokaNDE
- Directionality from Cerenkov cone
- Energy from total collected light
- Distinction between electrons and muons





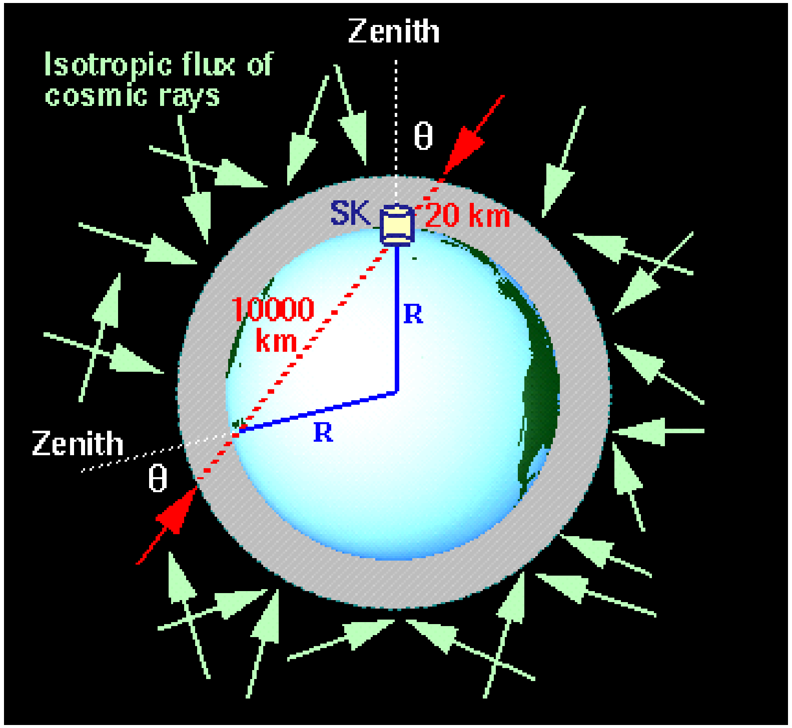


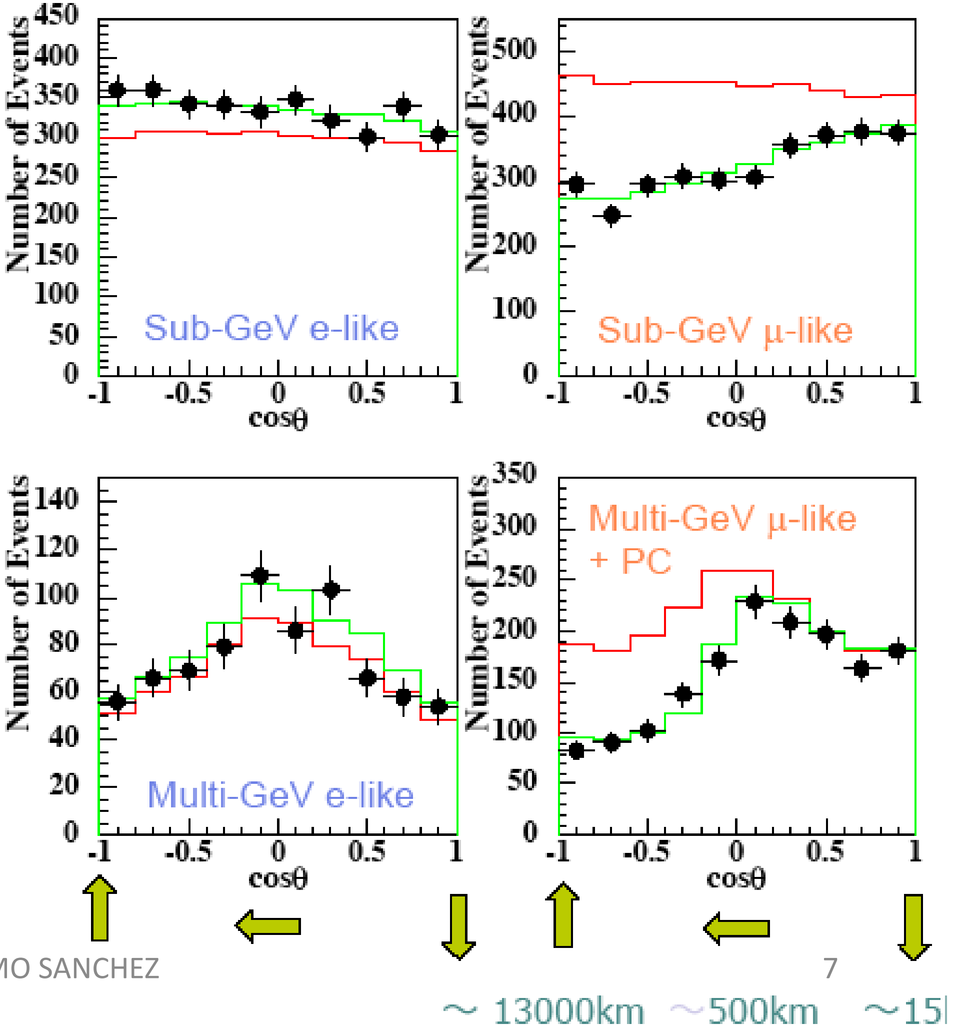




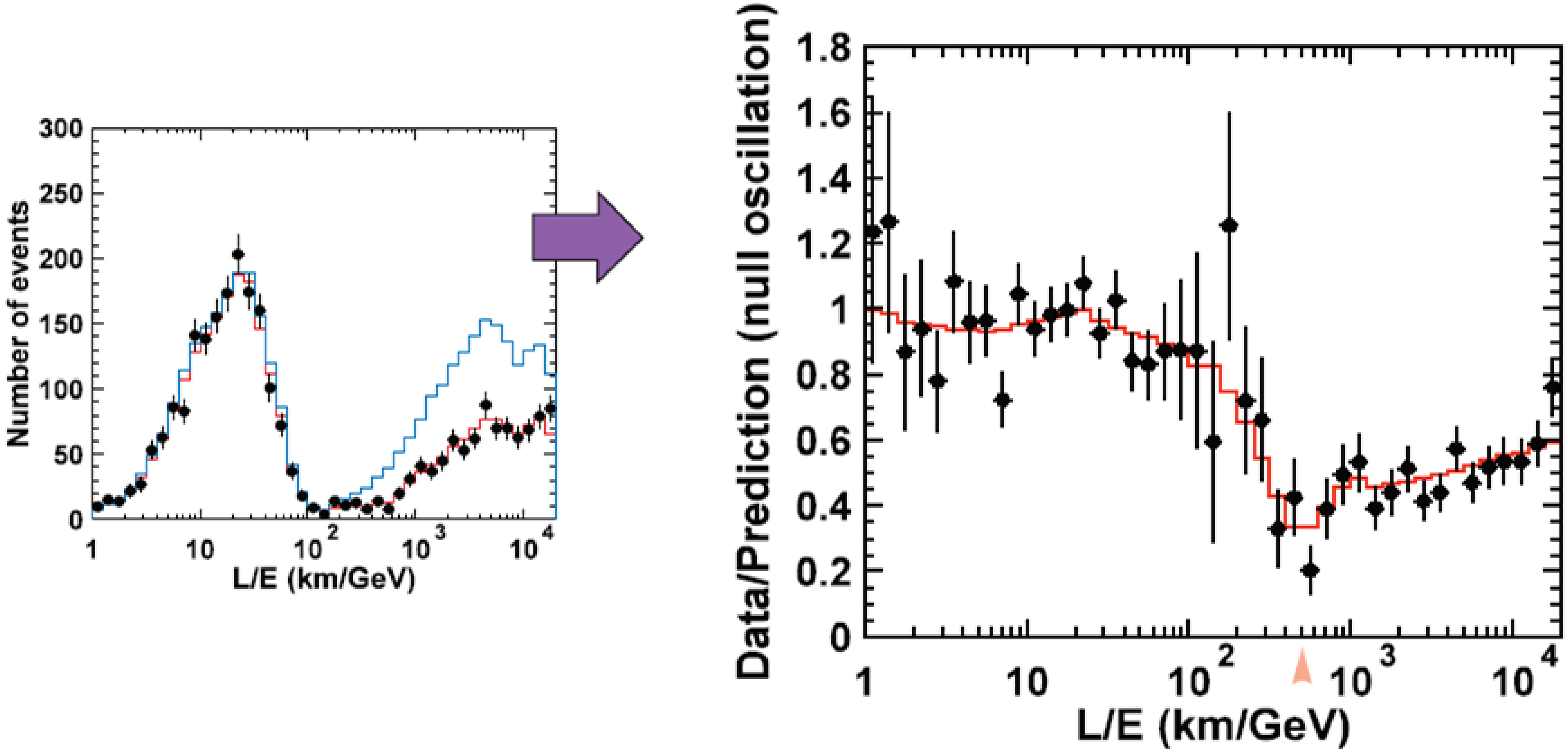
Super-KamiokaNDE

- 1000m deep, 50000 tons of water, 11000 PMTs
- Observed expected number of downgoing ν_μ, deficit in upgoing
- No excess in ν_e, so $\nu_\mu \rightarrow \nu_\tau$?



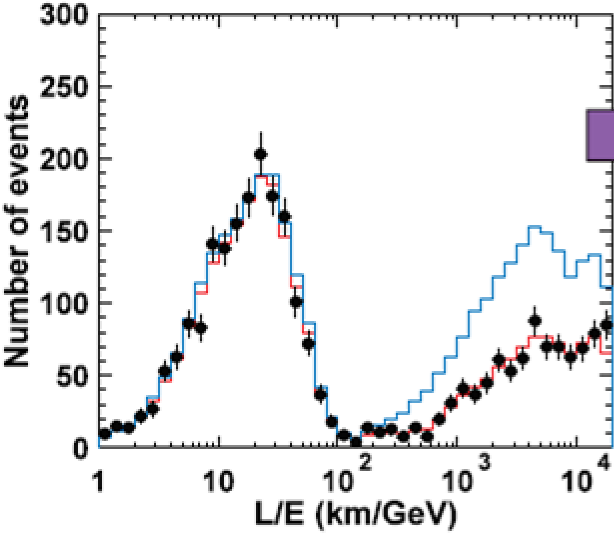


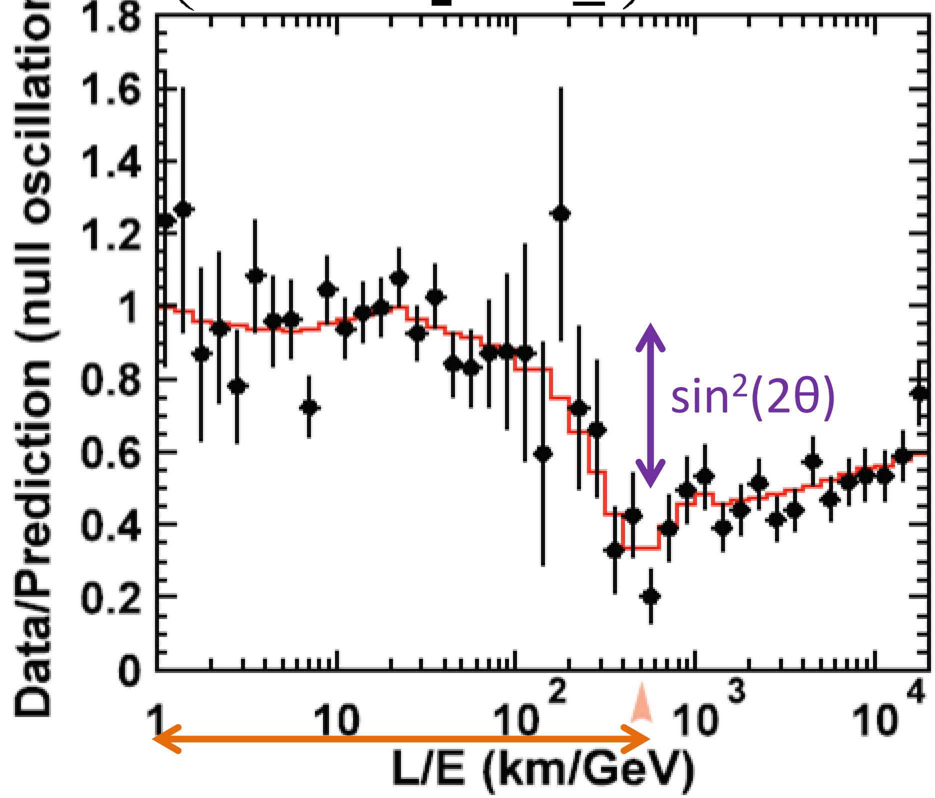
Atmospheric neutrinos disappear?



Atmospheric neutrinos oscillate!

$$P(\nu_\alpha \to \nu_\beta) = \sin^2(2\theta)\sin^2\left(1.27\frac{\Delta m^2 L[km]}{E[GeV]}\right)$$





L/E ~500 km/GeV ↔ Δm^2 ~ 2.3 x10^{-3} eV^2

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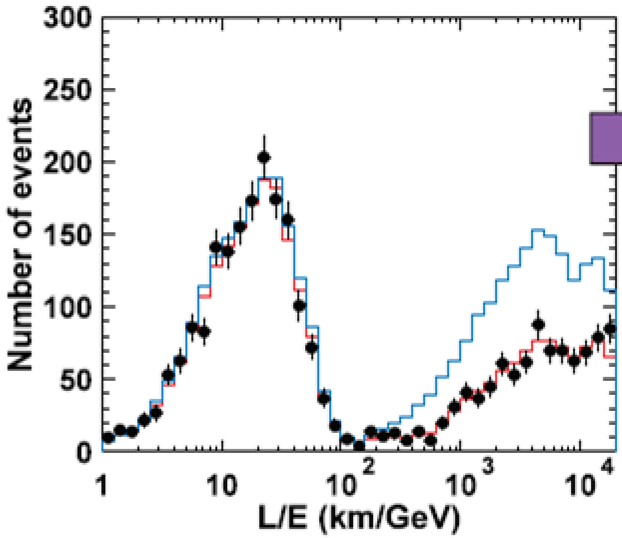
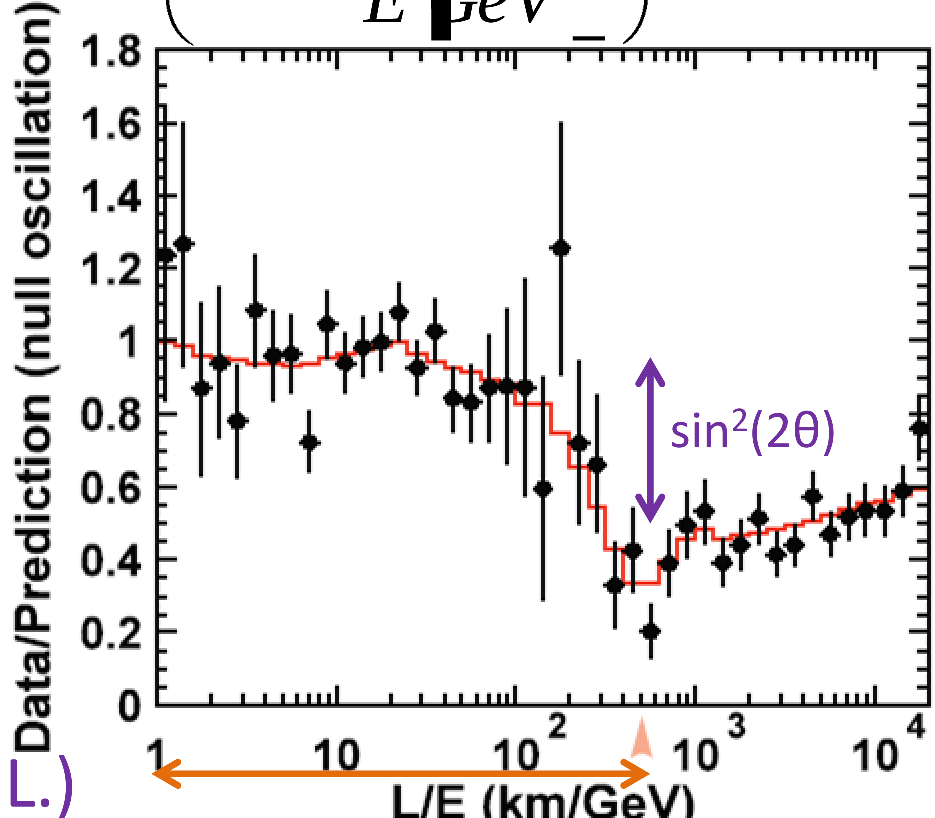




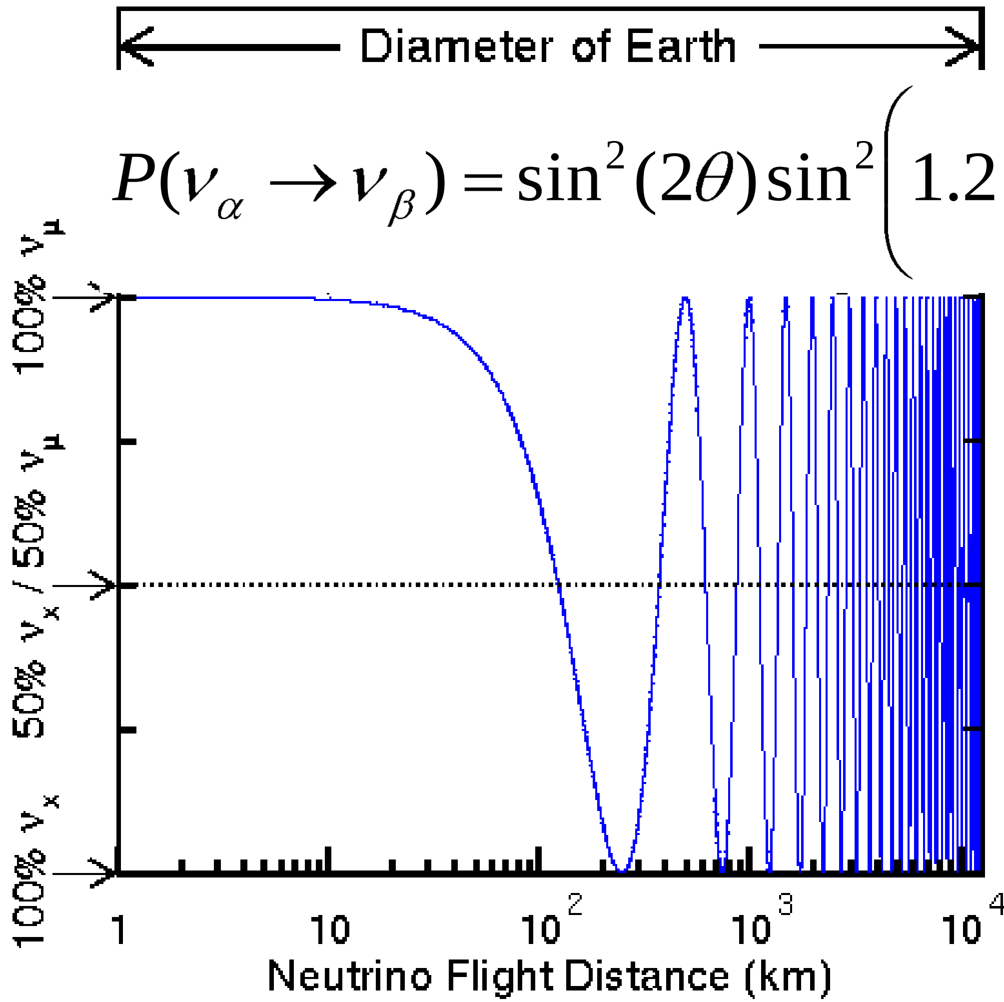
$\sin^2(2\theta) = 1.00$ (>0.93 90% C.L.)

$\Delta m^2 = (2.50 \pm 0.27) \times 10^{-3}\ eV^2$

L/E ~500 km/GeV ↔ $\Delta m^2 \sim 2.3 \times 10^{-3}\ eV^2$

But why don’t we see this?

$$P(\nu_\alpha \to \nu_\beta) = \sin^2(2\theta)\sin^2\left(1.27\frac{\Delta m^2 L\,[km]}{E\,[GeV]}\right)$$

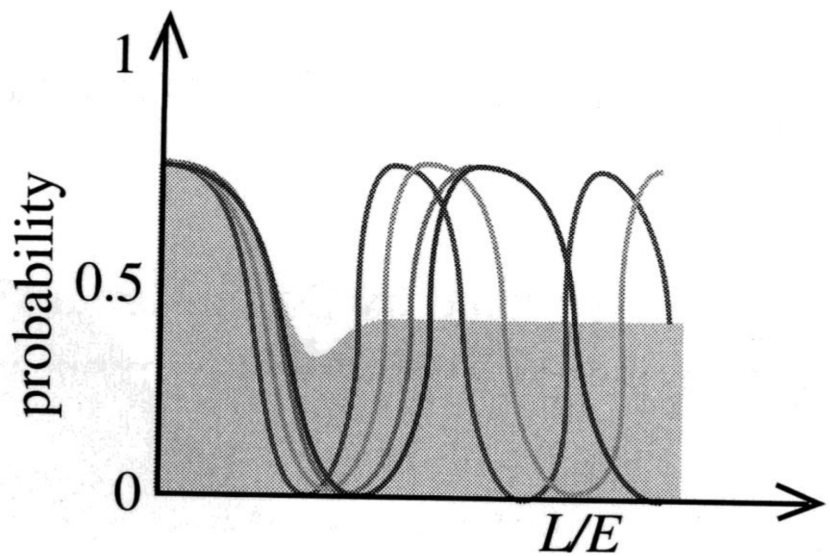


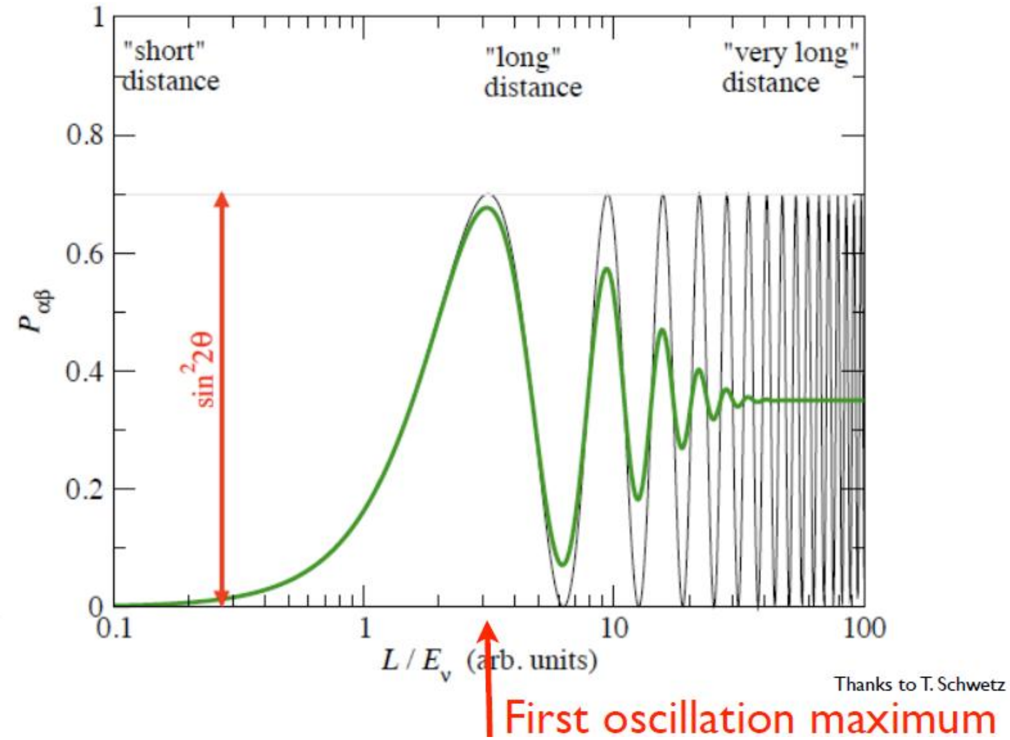
Because…

- Two effects:

  Neutrinos not monochromatic → different oscillation lengths

  Experimental resolution: if too close, maxima and minima blurred

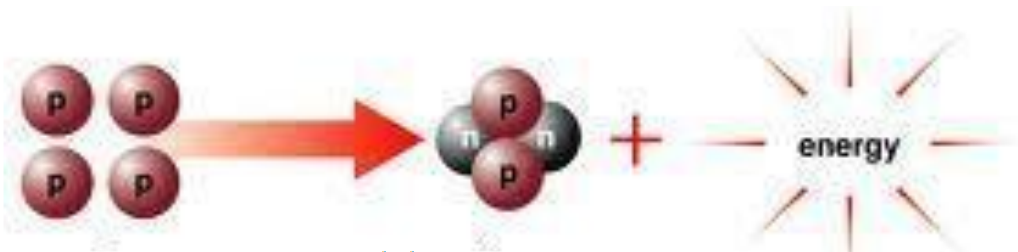




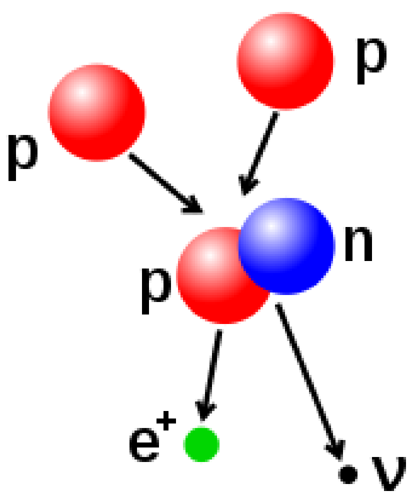
The solar neutrino saga

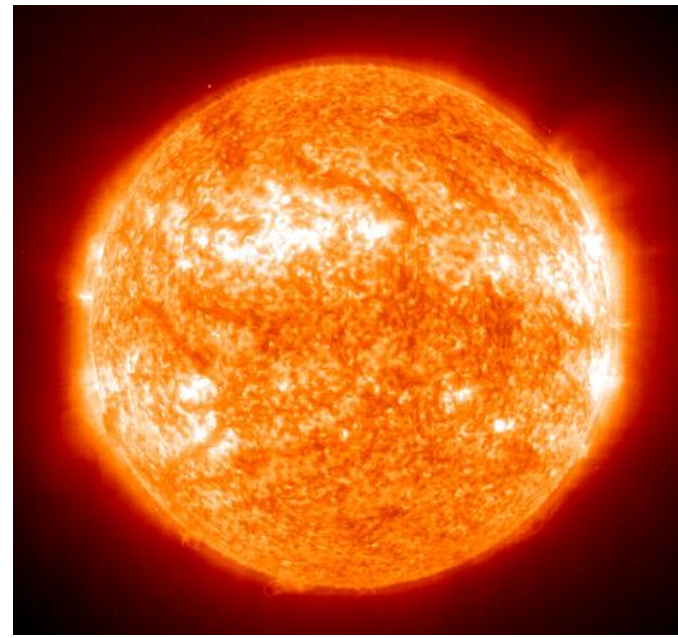
Neutrinos from the Sun

- Hydrogen fusion in the Sun requires inverse beta decay:



$$4p + 2e^- \rightarrow \text{He}^{++} + 2\nu_e + 26,7\ \text{MeV}$$



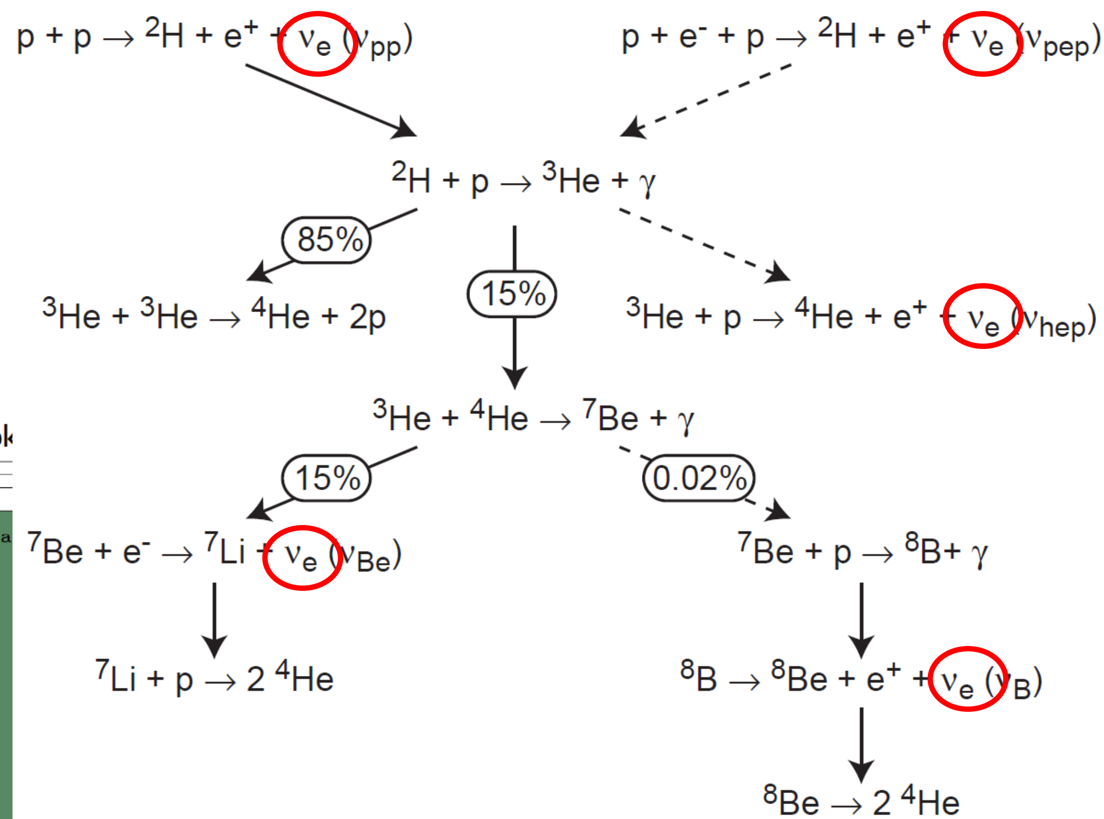


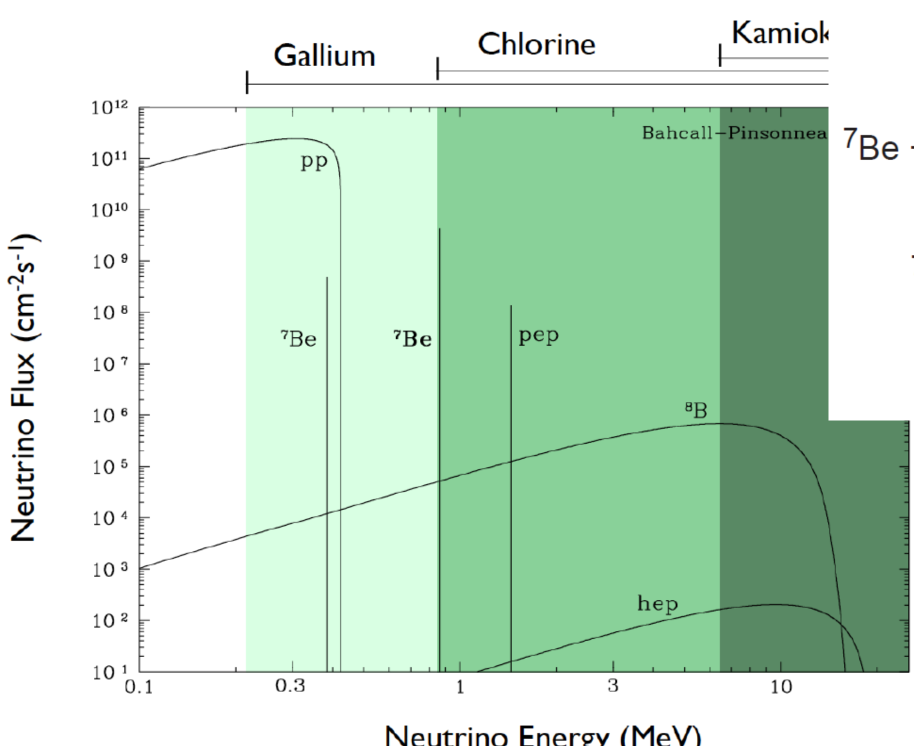
Solar constant = 1361 J/s m^2

$\phi_{\nu e}^{sun}$ = 6.4 x 10^{14} ν_e/s m^2

Neutrinos from the Sun

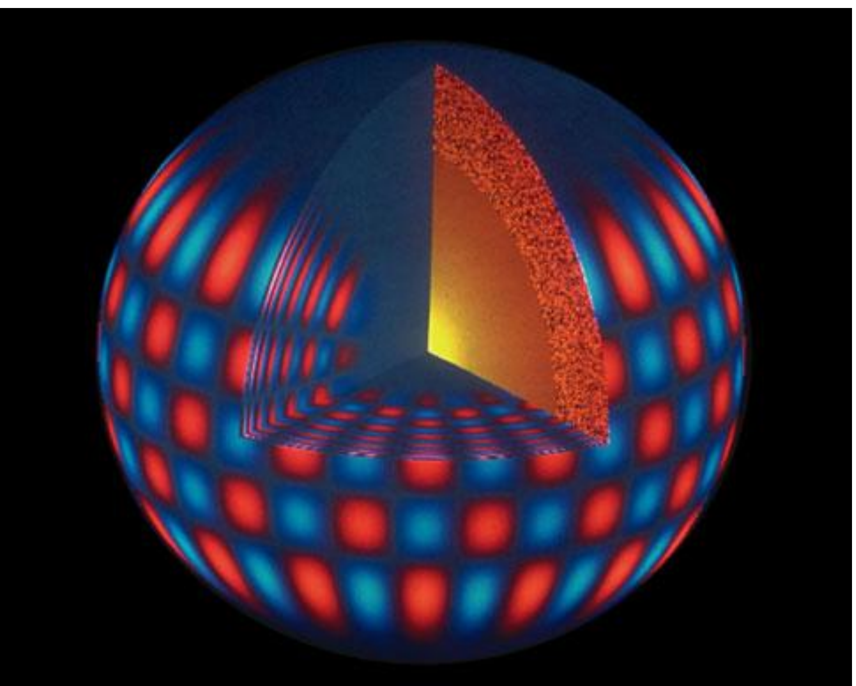
- Neutrino flux from the Sun accurately predicted (Bahcall et al)

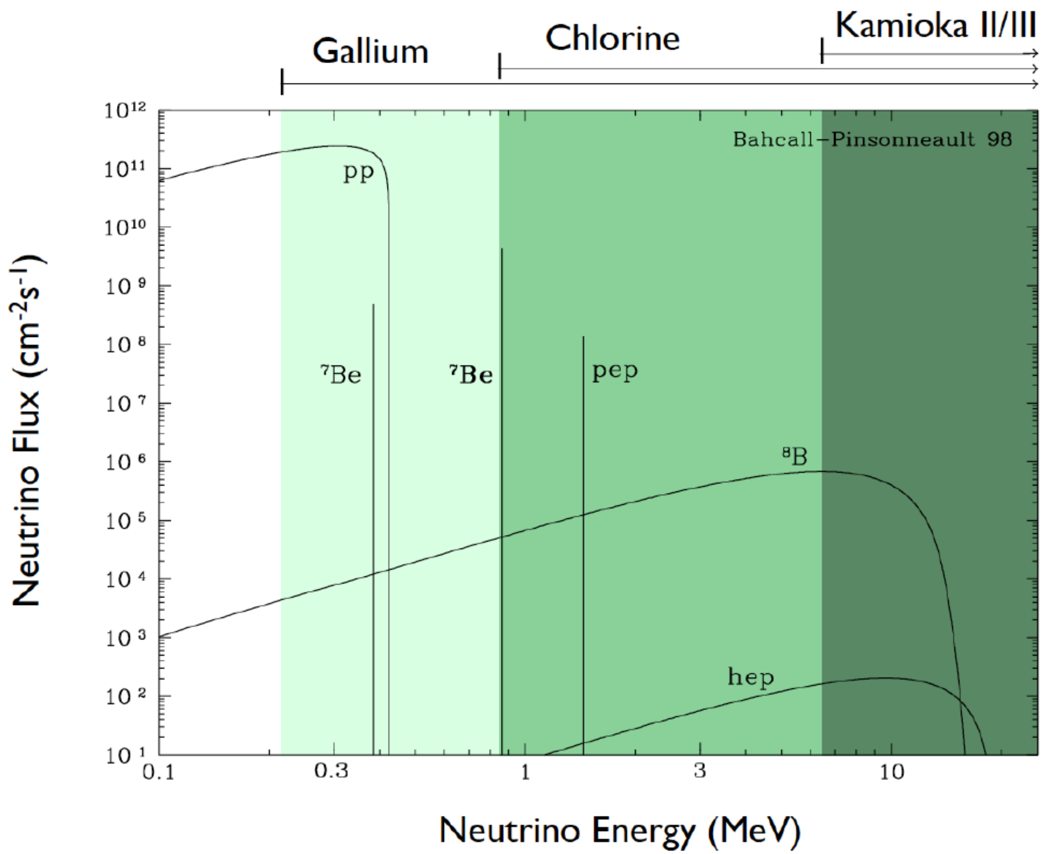




Neutrinos from the Sun

- Neutrino flux from the Sun accurately predicted (Bahcall et al)
- Model in good agreement with results from helioseismology





Homestake experiment

Late 1960s: Ray Davis set to test ν_e flux predictions in underground mine (under 1500m of rock)

Experiment run for 30 years (till 1994):

- observed 2.56 ± 0.23 SNU
- expected 8.2 ± 1.8 SNU

} ~30%

1 Solar Neutrino Unit = 10^{-36} interactions/s atom

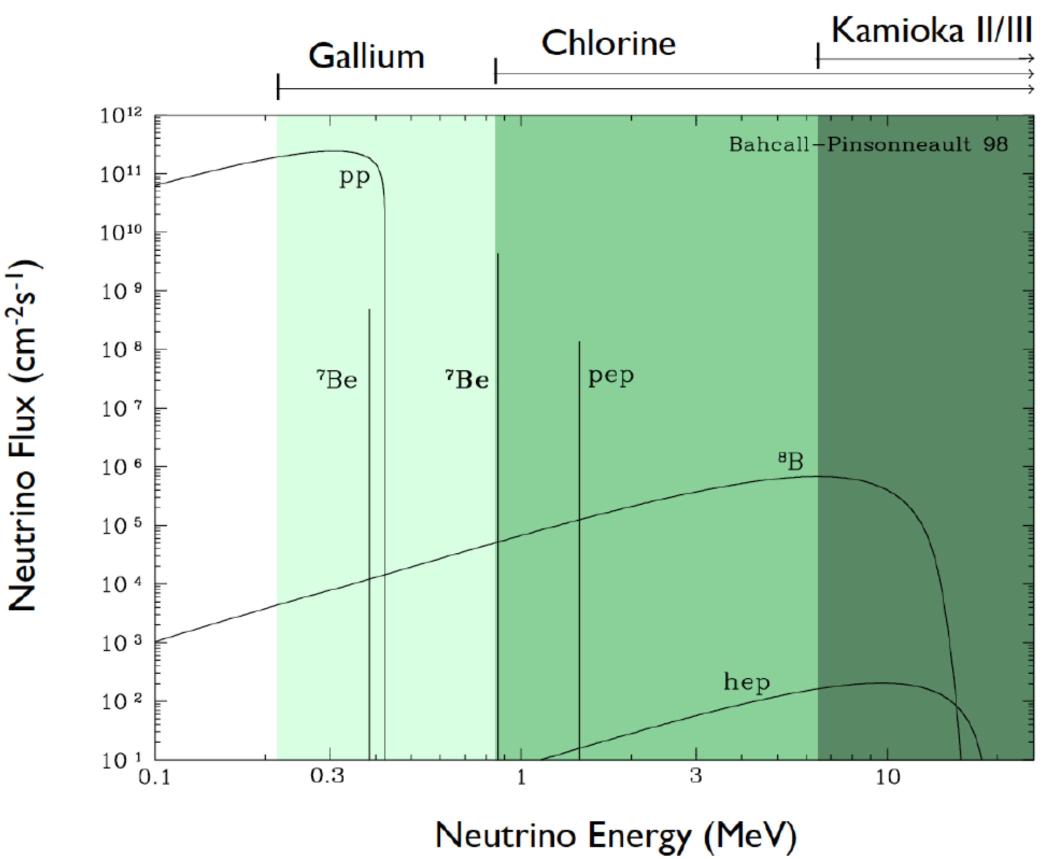
$\nu_e + n \rightarrow p + e^-$

- **Homestake:** $\nu_e + {}^{37}Cl \rightarrow {}^{37}Ar + e^-$
  - Located in Lead, SD
  - 615 tons of C_2Cl_4 (Cleaning fluid)
  - Extraction method:
    - Pump in He that displaces Ar
    - Collect Ar in charcoal traps
    - Count Ar using radioactive decay
  - Never Calibrated with source

Raymond Davis, Nobel Prize 2002

Problems?

- Problems with experiment? With ν_e flux predictions?
- Test other parts of the ν_e spectrum with different experimental techniques

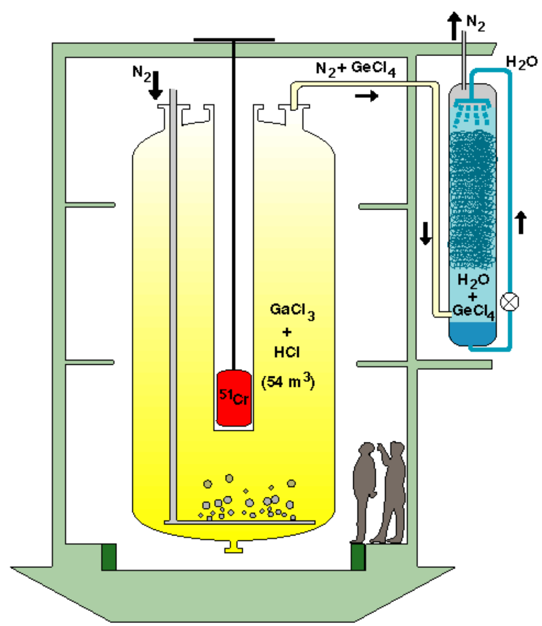


Gallex: $\nu_e + {}^{71}Ga \rightarrow {}^{71}Ge + e^-$

Observed 68.1 ± 3.75 SNU

Expected 127 ± 12 SNU

~50%



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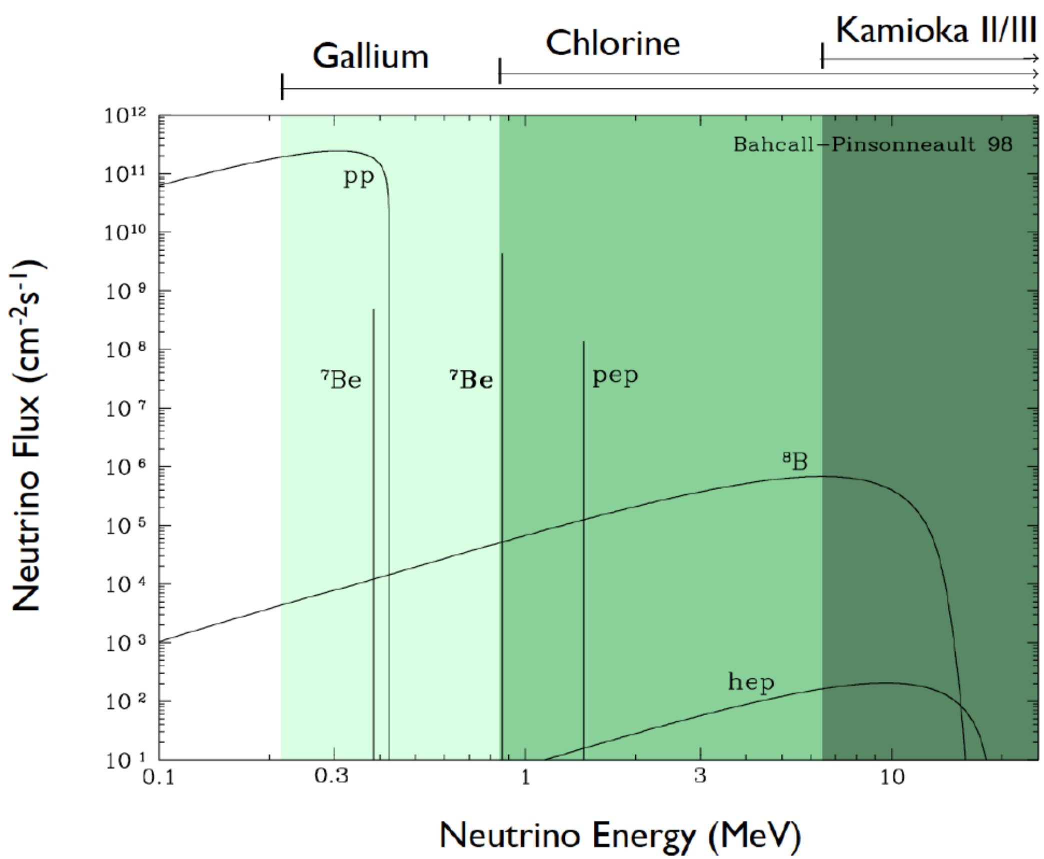
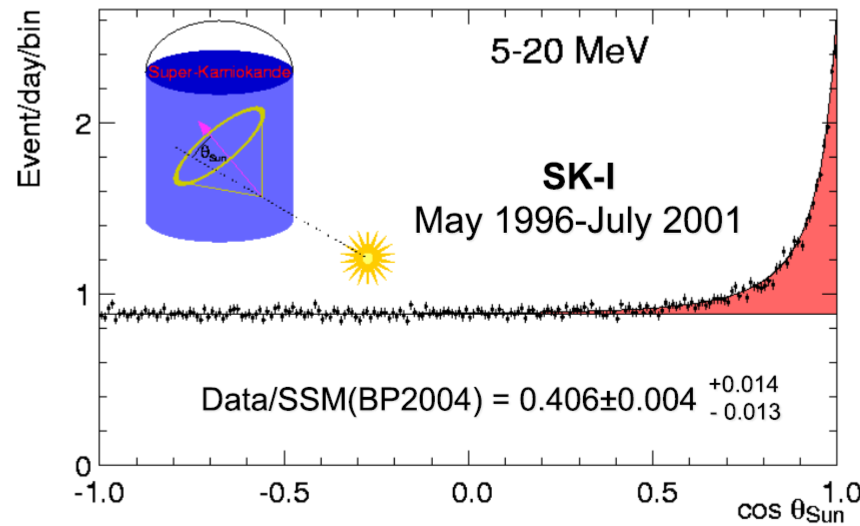
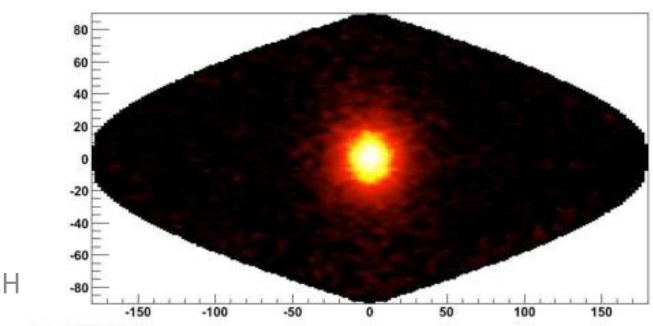


KamiokaNDE: $\nu_e + e^- \rightarrow \nu_e + e^-$

Observed ~40% of expectation





Problems?

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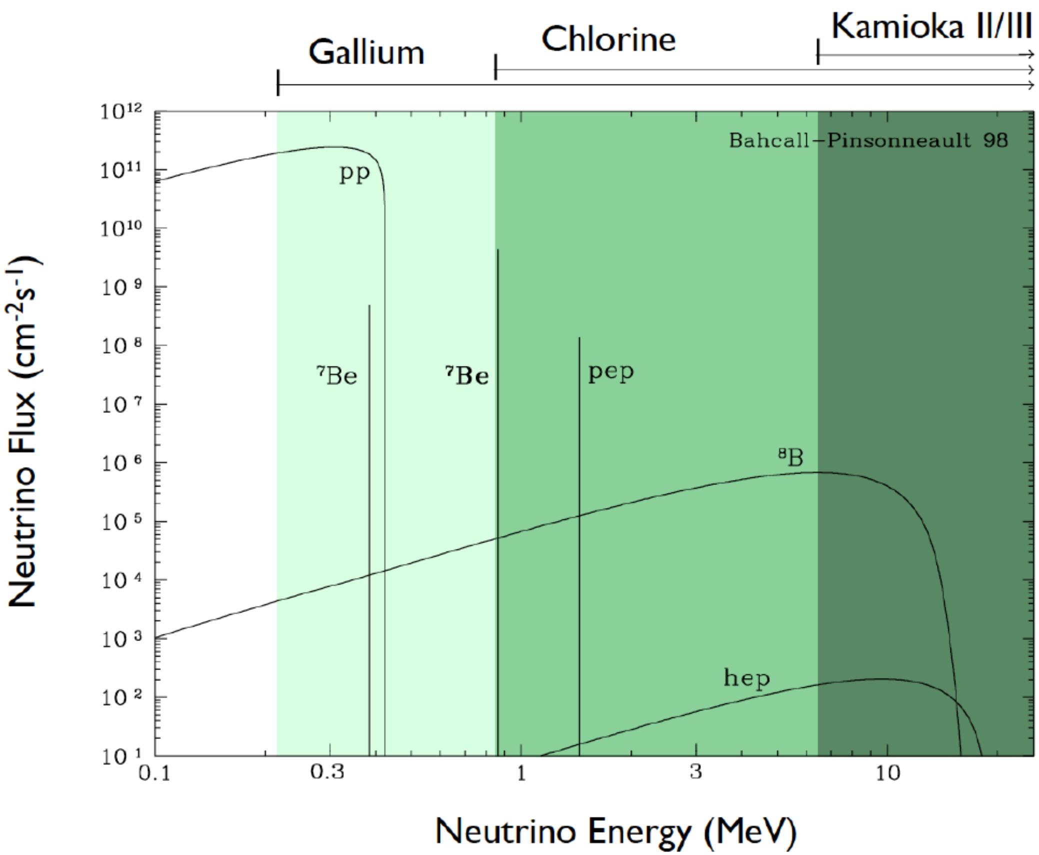


Experiment type	Observed/Expected
Chlorine	~30%
Gallium	~60%
KamiokaNDE	~40%

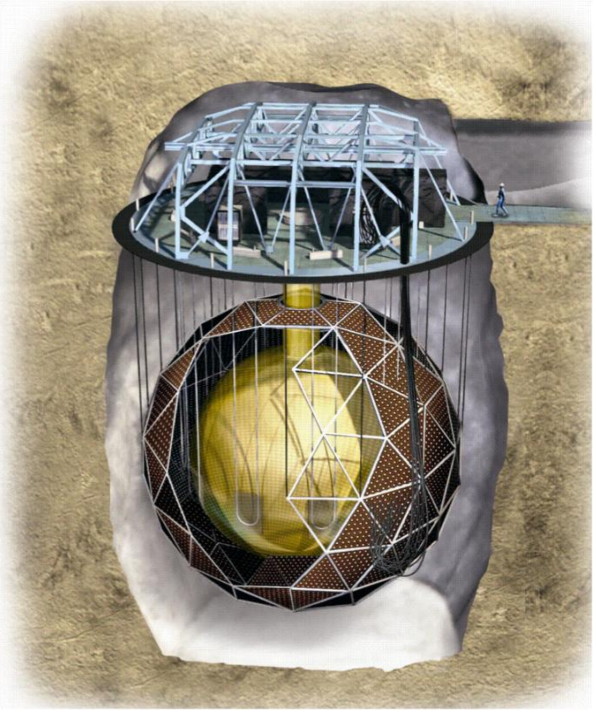
Perhaps neutrinos are oscillating after all, as suggested by Pontecorvo et al?

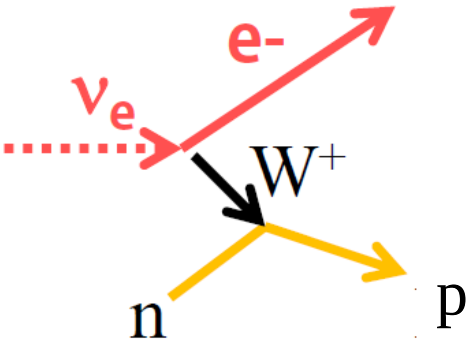
These experiments only sensitive to ν_e try and detect ν_μ and ν_τ too! → SNO

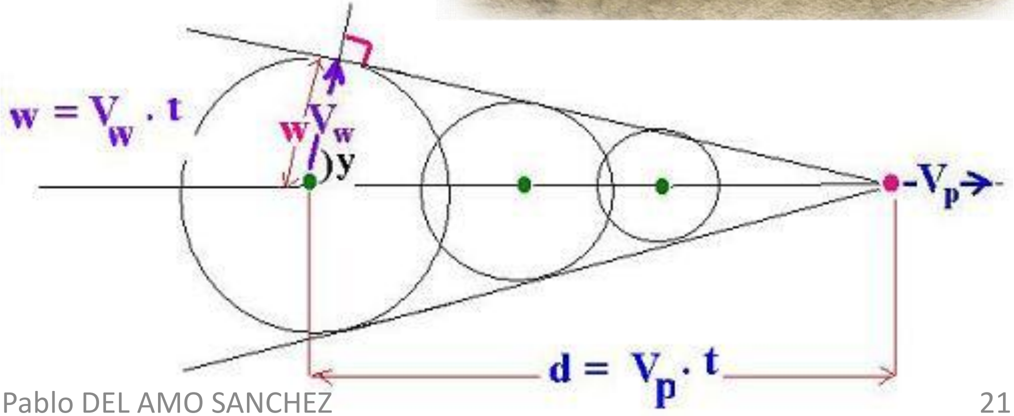
Sudbury Neutrino Observatory (SNO)

- 2000 m deep (Sudbury, Ontario)
- Cosmics veto
- 1000 tons of Heavy water (D_20),
  shielded by 7000 tons light water (H_20)
  seen by 9500 photomultiplier tubes (PMTs)
- So-called Water Cerenkov detector

Particles faster than speed of light in medium radiate light (e.g. blueish light in nuclear reactors)







SNO

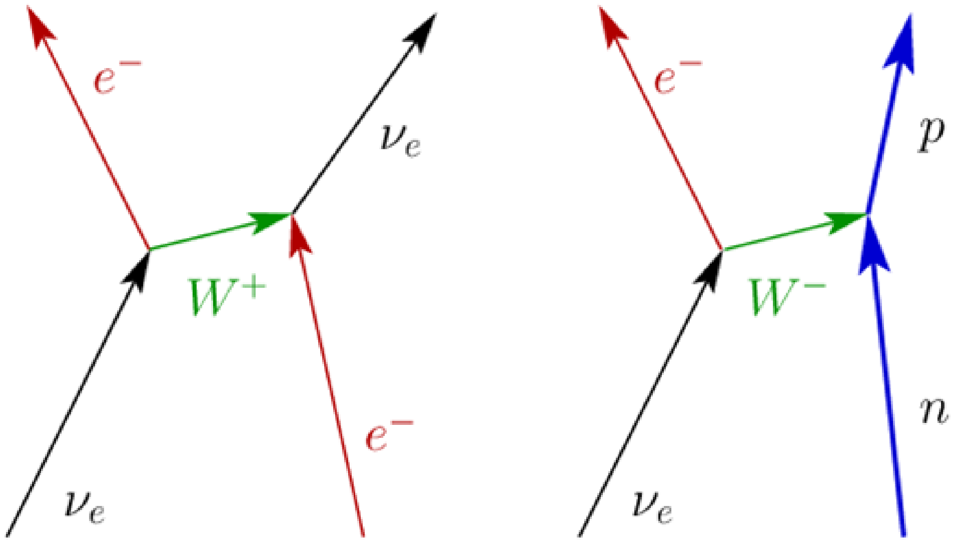
- SNO measures well ν_e flux:

CC : $\nu_e + d \rightarrow p + p + e^-$

- Good measurement of the ν_e spectrum.
- Some directional information.
- Only sensitive to ν_e.

ES : $\nu_e + e^- \rightarrow \nu_e + e^-$

- Strong directional sensitivity.
- Low statistics.



Charged current

- Cannot see ν_μ / ν_τ flux in this way: neutrinos from Sun not energetic enough to produce heavy μ or τ particles in interactions

SNO

- But it measures the total $\nu_e+\nu_\mu+\nu_\tau$ flux by means of Neutral Current interactions!

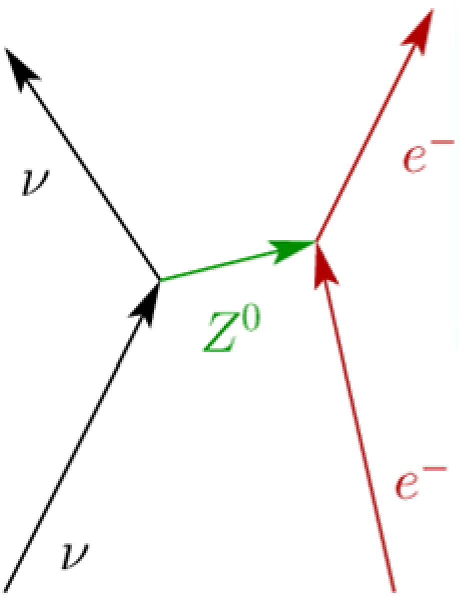
NC : $\nu_x + d \rightarrow n + p + \nu_x$

- Measures total ^{8}B flux from the Sun.
- Equal cross-section to all (active) neutrino flavours.

Signature event of SNO

3 neutron detection methods:

$n + d \rightarrow t + \gamma + 6.26\,\text{MeV}$	**SNO-I**
$n + {}^{35}\text{Cl} \rightarrow {}^{36}\text{Cl} + \gamma + 8.6\,\text{MeV}$	**SNO-II**
$n + {}^{3}\text{He} \rightarrow p + t + 0.76\,\text{MeV}$	**SNO-III**

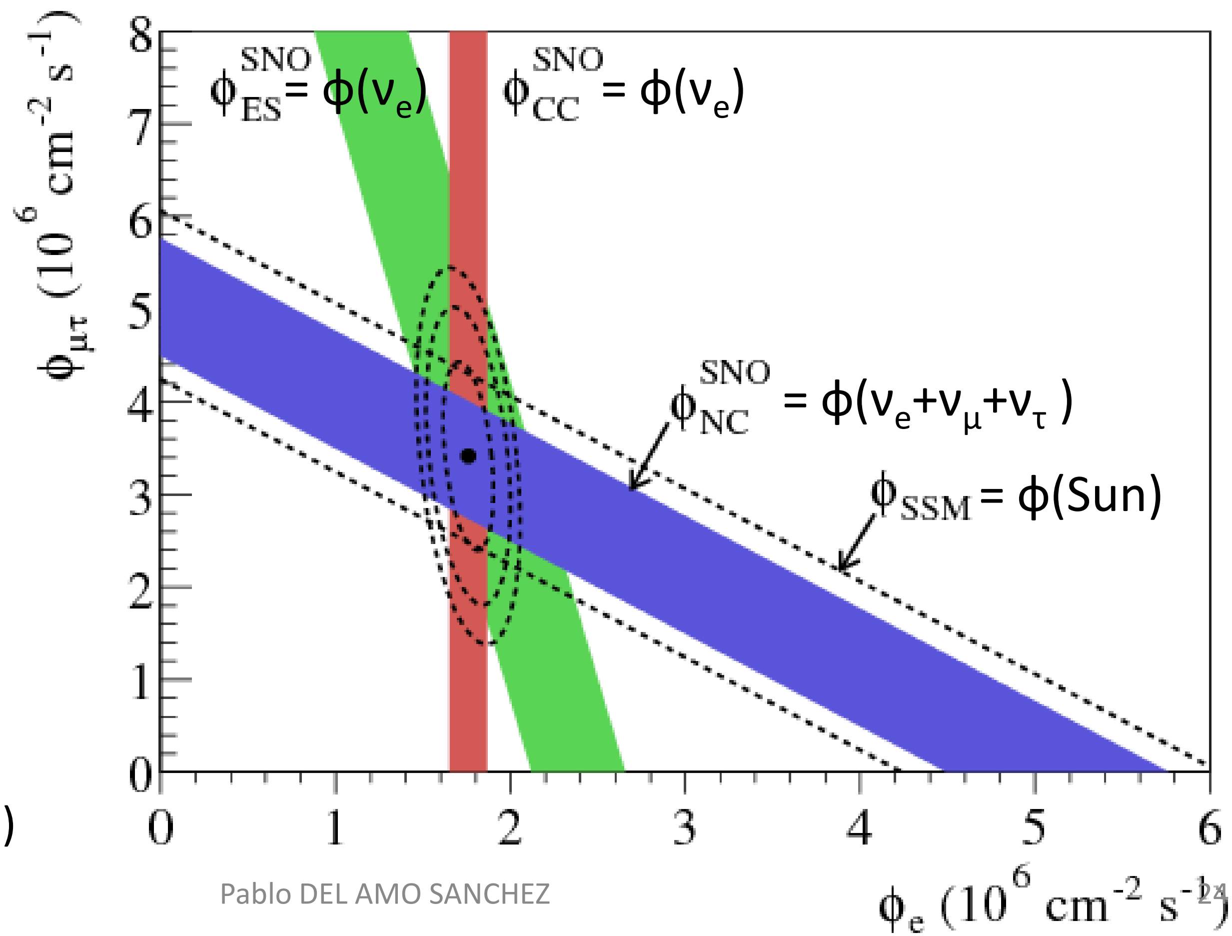


Neutral current

Solar neutrinos oscillate!

Less ν_e than predicted but total $\nu_e+\nu_\mu+\nu_\tau$ correct!

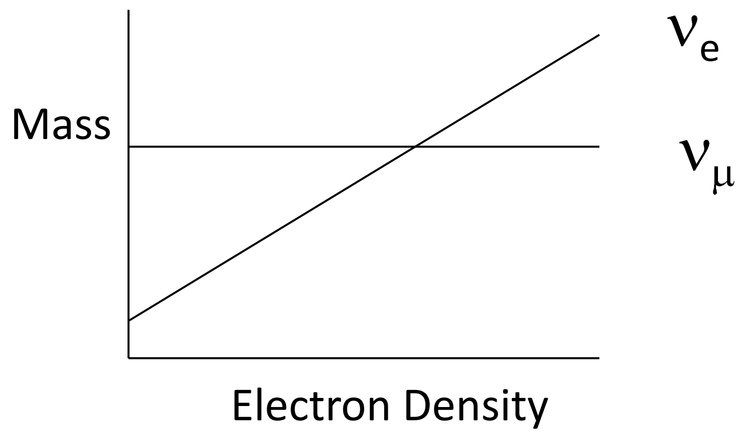
Bruno Pontecorvo (1957)



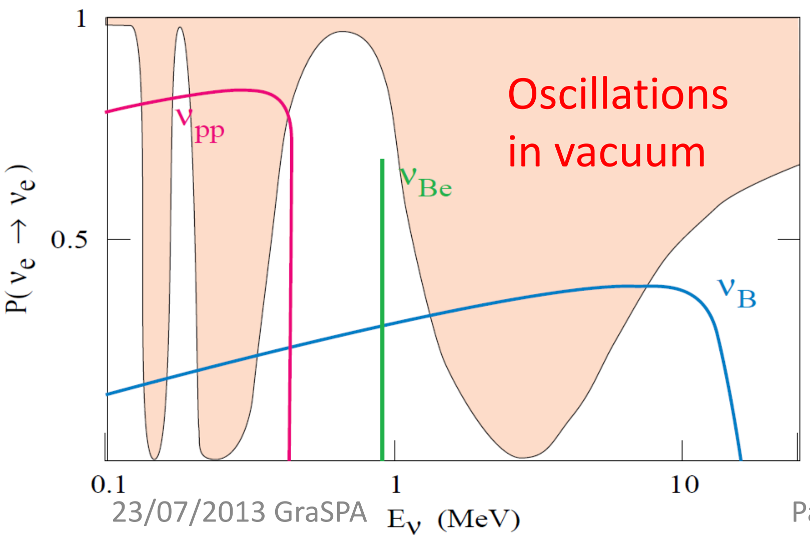
Matter effects are important!

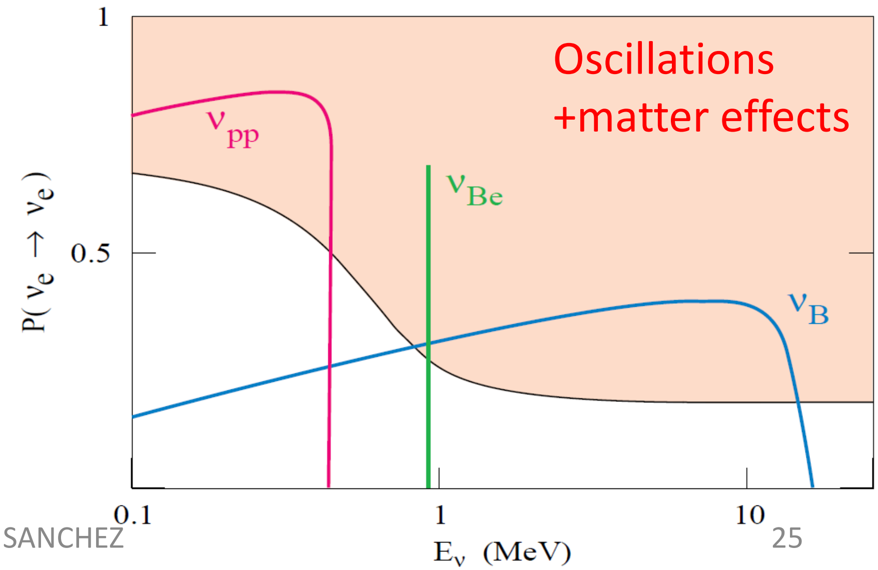
- High electron density in Sun → matter effects!
- ν_e get heavier, ν_μ & ν_τ unaffected.

  Resonance effects may enhance oscillation



$$P(\nu_e \rightarrow \nu_e) = 1 - \sin^2(2\theta)\sin^2\left(1.27\frac{\Delta m^2 L[km]}{E[GeV]}\right)$$



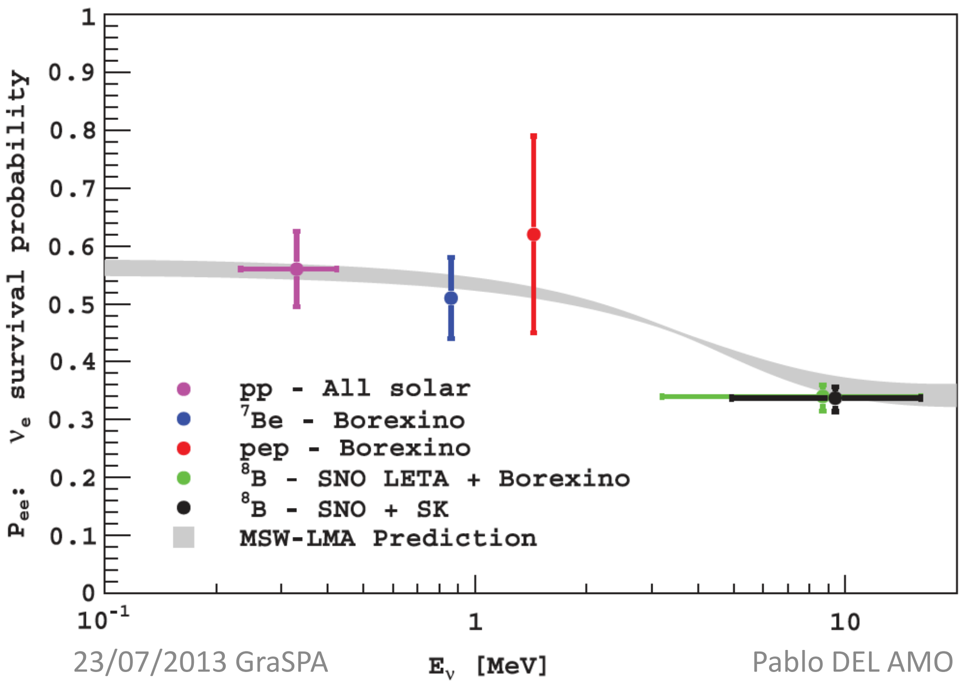


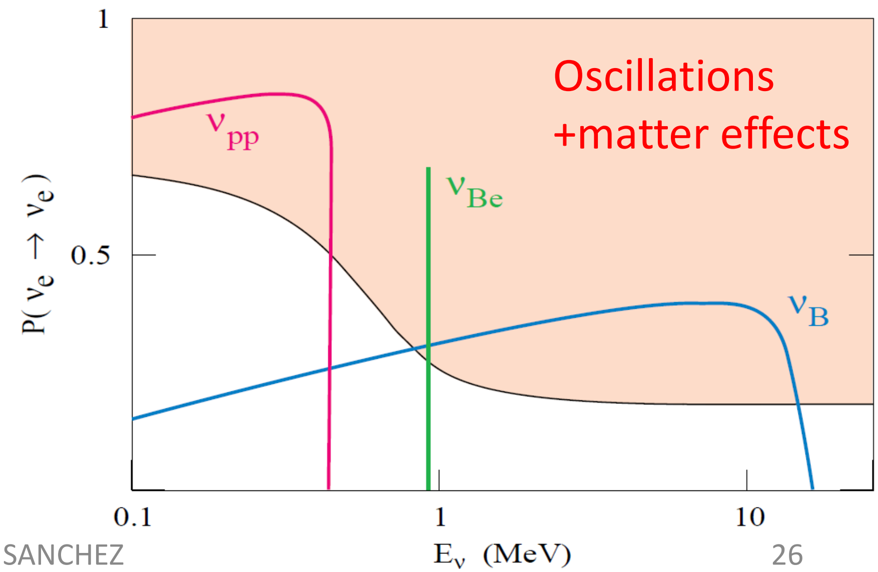
Matter effects are important!

- Found oscillation parameters for solar neutrinos:

$$P(\nu_e \rightarrow \nu_e) = 1 - \sin^2(2\theta)\sin^2\left(1.27\frac{\Delta m^2 L\,[km]}{E\,[GeV]}\right)$$

$\sin^2(2\theta) = 0.857 \pm 0.024$ $\Delta m^2 = (7.5 \pm 0.20) \times 10^{-5} eV^2$





Closing the trilogy:
reactor neutrino experiments

Reactor neutrinos

- Nuclear reactors, source of abundant antineutrinos! $\bar{\nu}_e$

Fission products are neutron rich

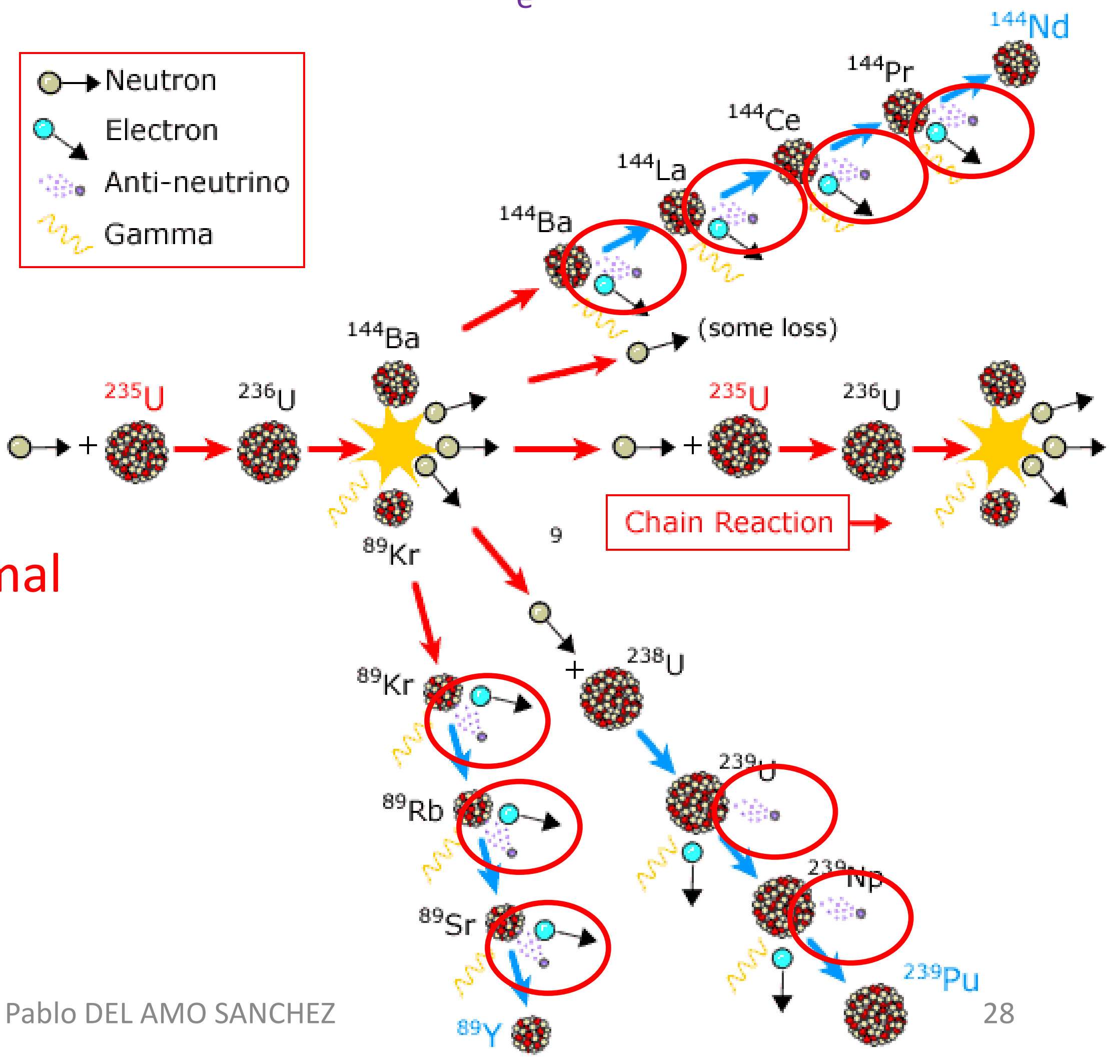
Too many neutrons to be stable

→ plenty of beta decays!

- ~ 6 $\bar{\nu}_e$ /fission

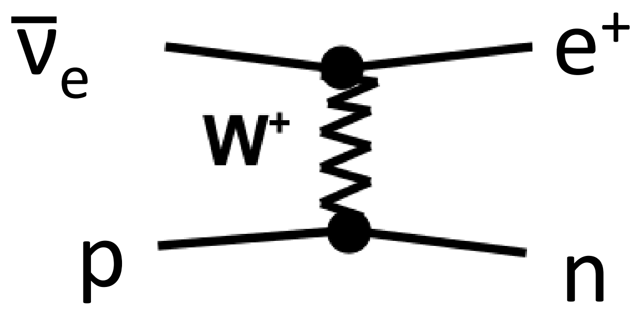
  ~ 200 MeV/fission

$2 \times 10^{20}\ \bar{\nu}_e / GW_{thermal}$



Liquid scintillator detectors

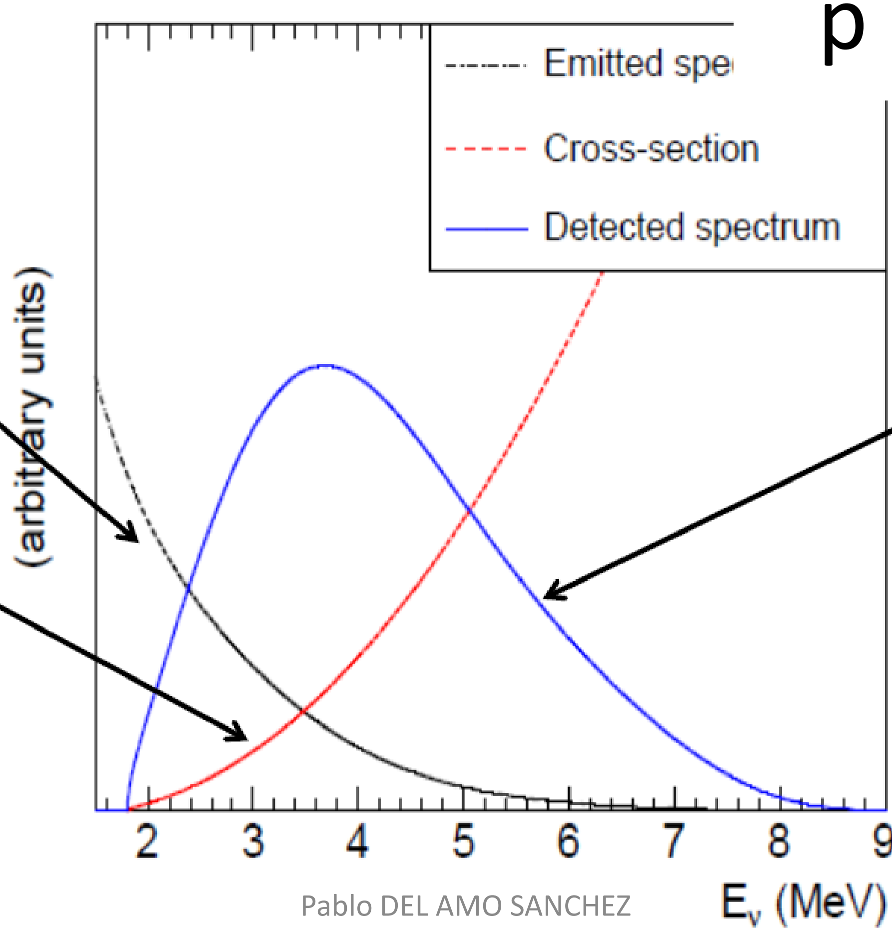
- Detect reactor $\overline{\nu}_e$ through inverse beta decay



Exponential decrease of emitted spectrum

β−inverse detection process

$$\overline{\nu}_e + p \rightarrow e^+ + n$$

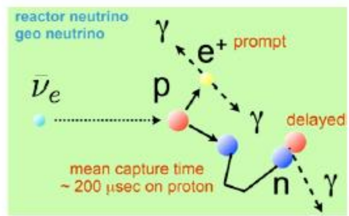


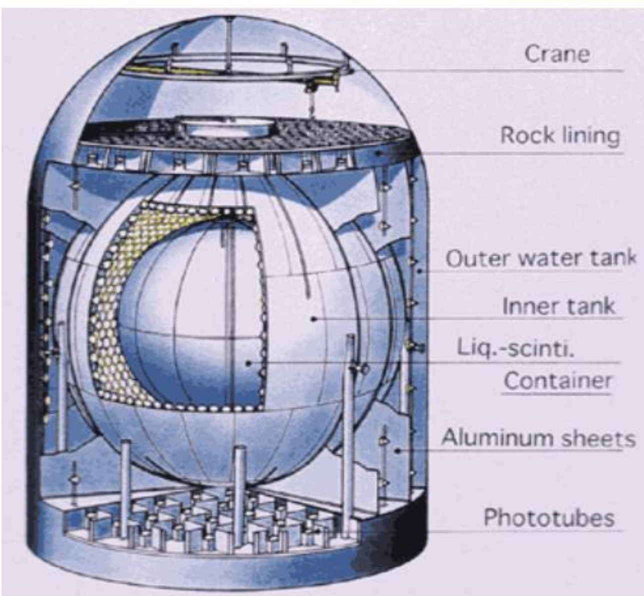
Detected Spectrum

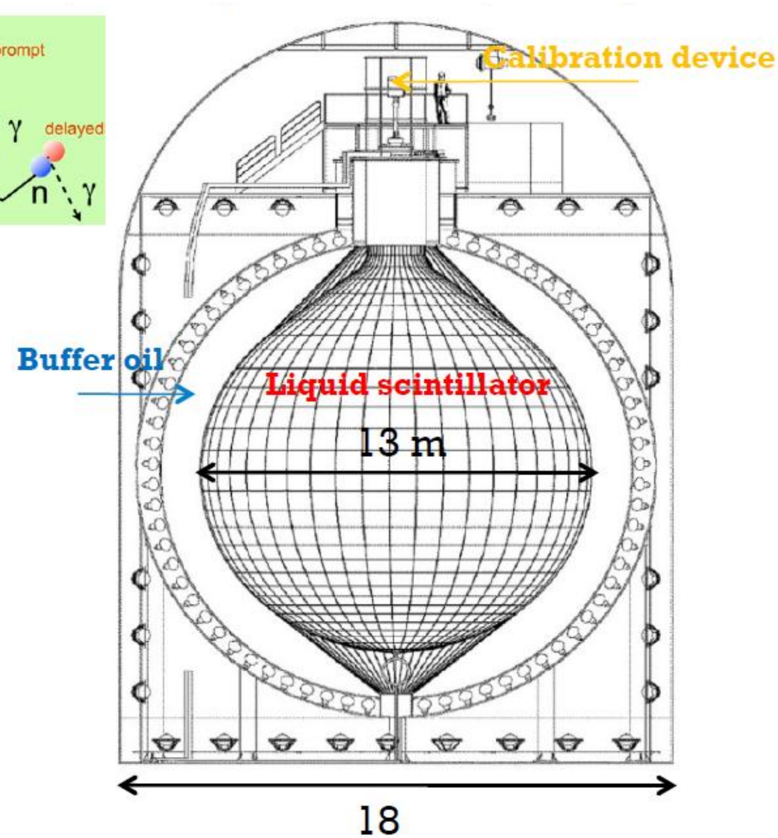
Relevant E range [1.8 – 8] MeV

Liquid scintillator detectors

- KamLAND: Kamioka Liquid scintillator AntiNeutrino Detector
  - **1000 ton liquid scintillator:**
  - Spherical plastic balloon
  - 1325 17" + 554 20" PMTs
  - Inverse β decay detection

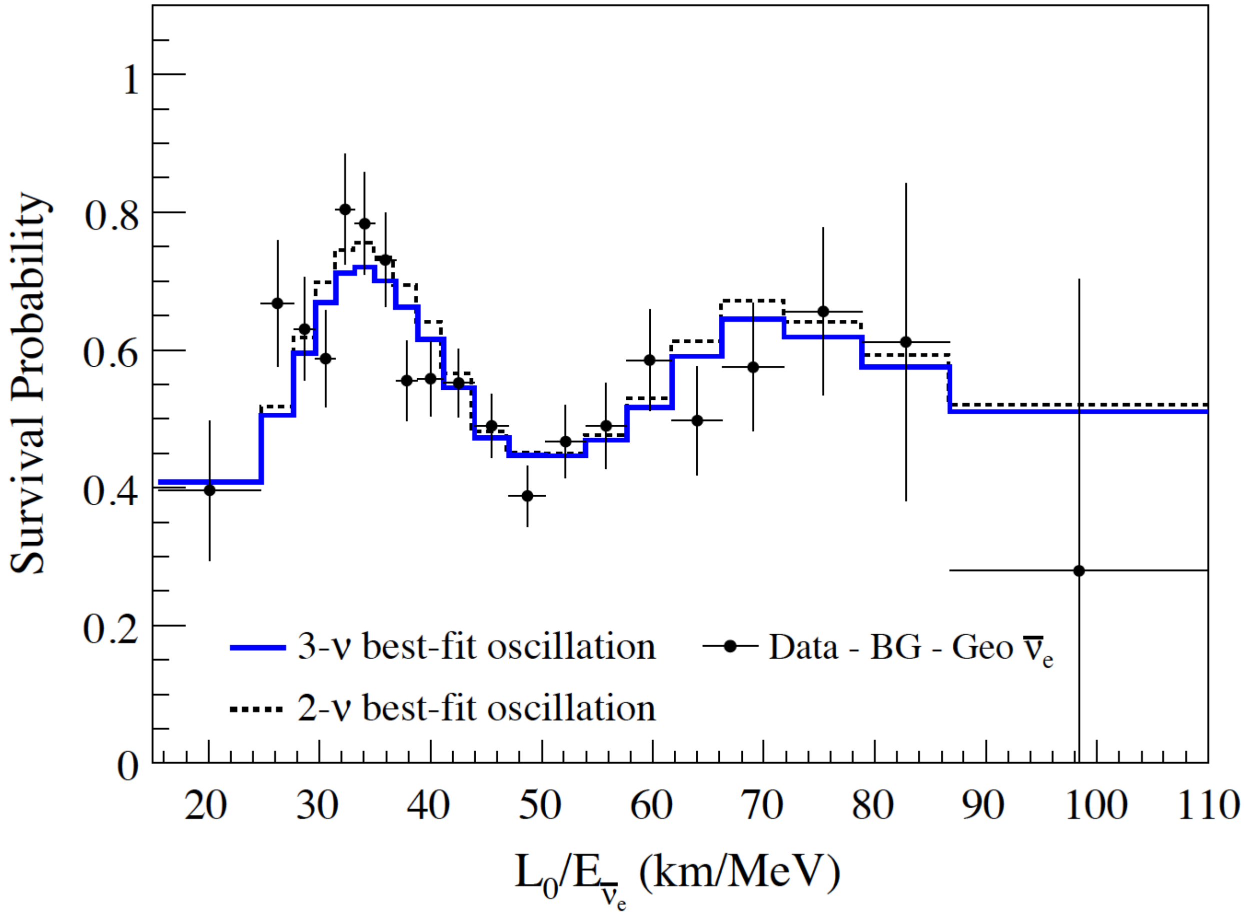






Reactor neutrinos oscillate!

- Confirm solar neutrino oscillations



What have we learnt so far?

- **Neutrinos oscillate!**

  ν_e, ν_μ, ν_τ different from ν_1, ν_2, ν_3

- Two different oscillation frequencies:

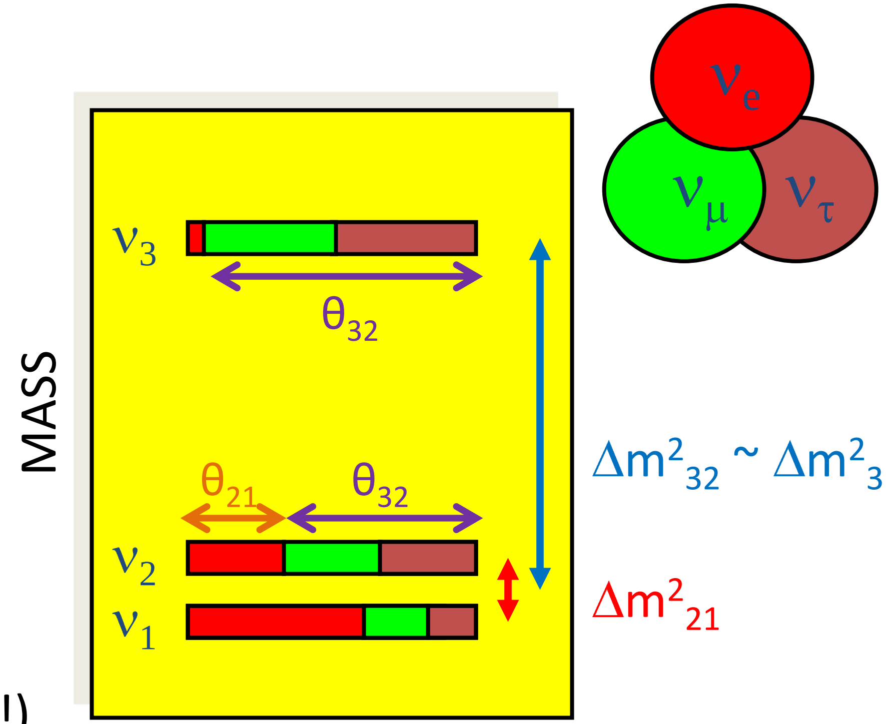
  fast: atmospheric, $\Delta m^2_{32} \sim \Delta m^2_{31}$

  slow: solar, Δm^2_{21} atm ~ 20 x solar

- Neutrinos mix a lot! (Mixing angles large!)

  atmospheric, maximal $\theta_{32} = 45° \pm 6°$

  solar, large $\theta_{21} = 34° \pm 1°$



What have we learnt so far?

- Neutrinos oscillate!

  ν_e, ν_μ, ν_τ different from ν_1, ν_2, ν_3

- Two different oscillation frequencies:

  fast: atmospheric, $\Delta m^2_{32} \sim \Delta m^2_{31}$

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- Neutrinos mix a lot! (Mixing angles large!)

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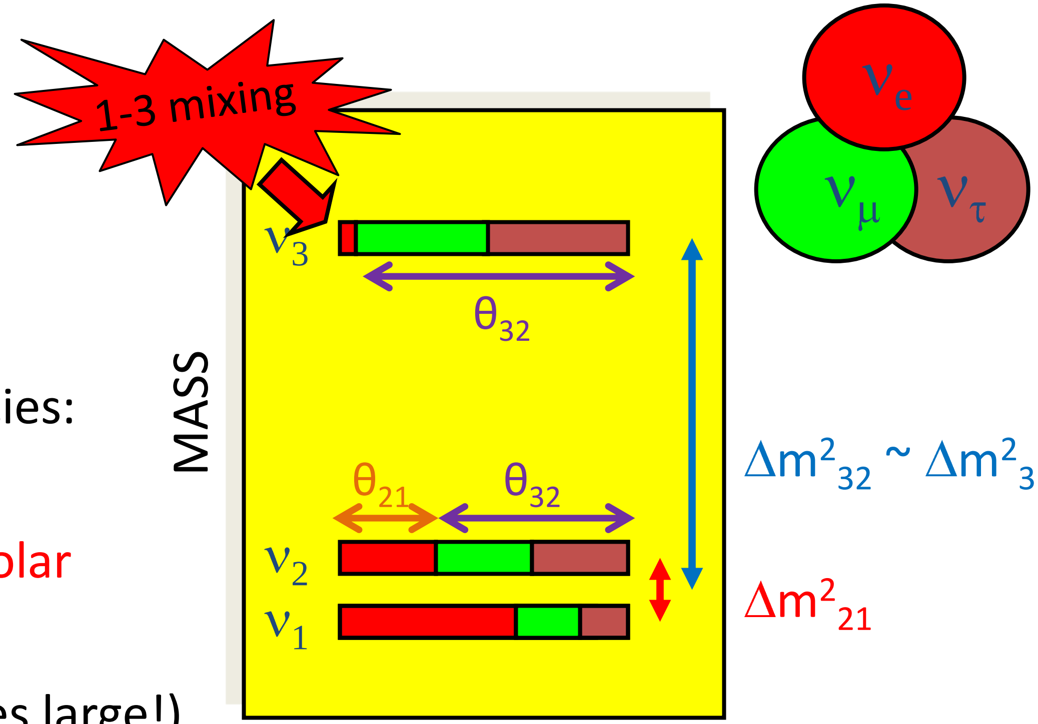
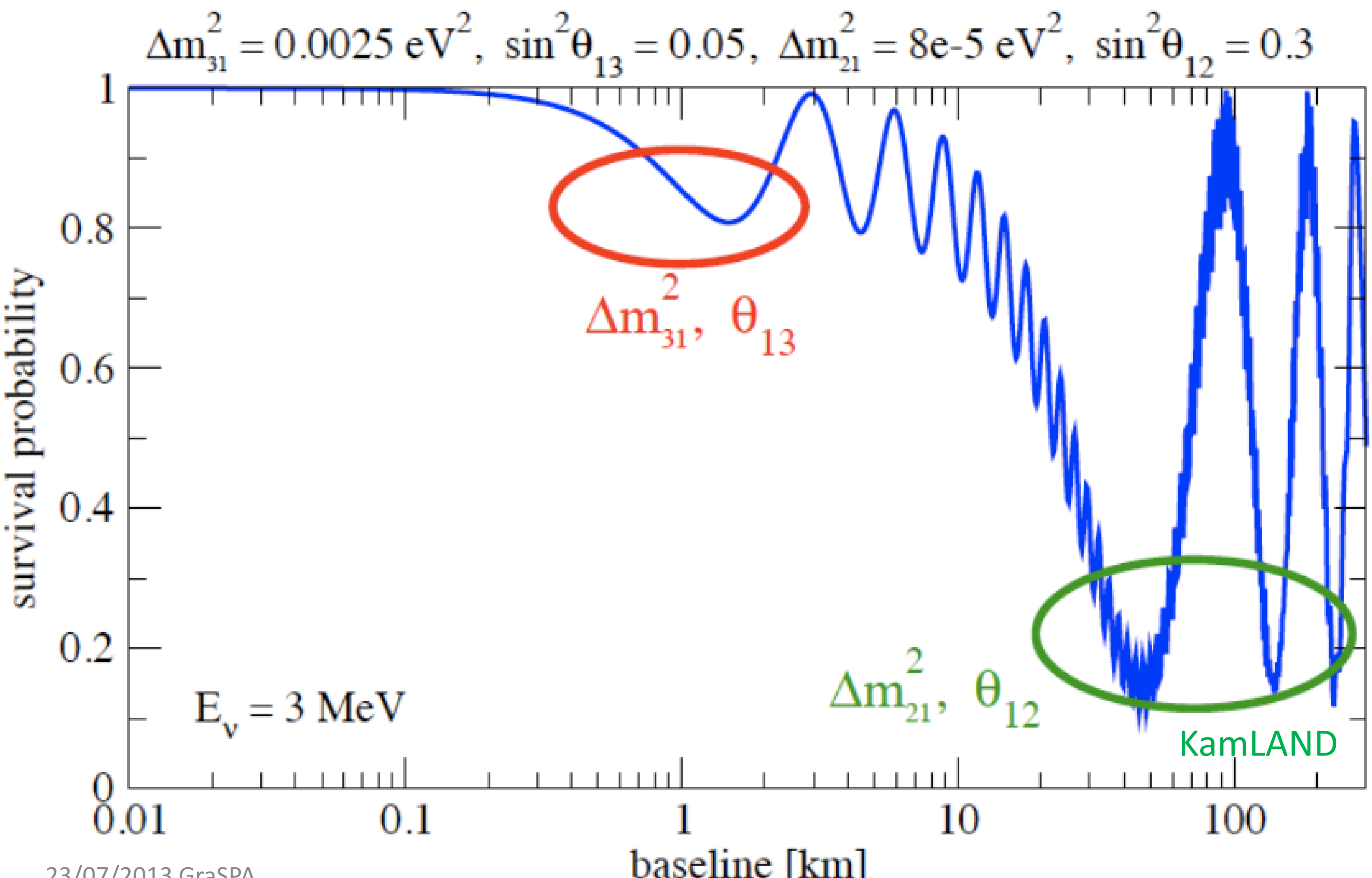
- What is the amount of ν_e in ν_3 (θ_{13})?



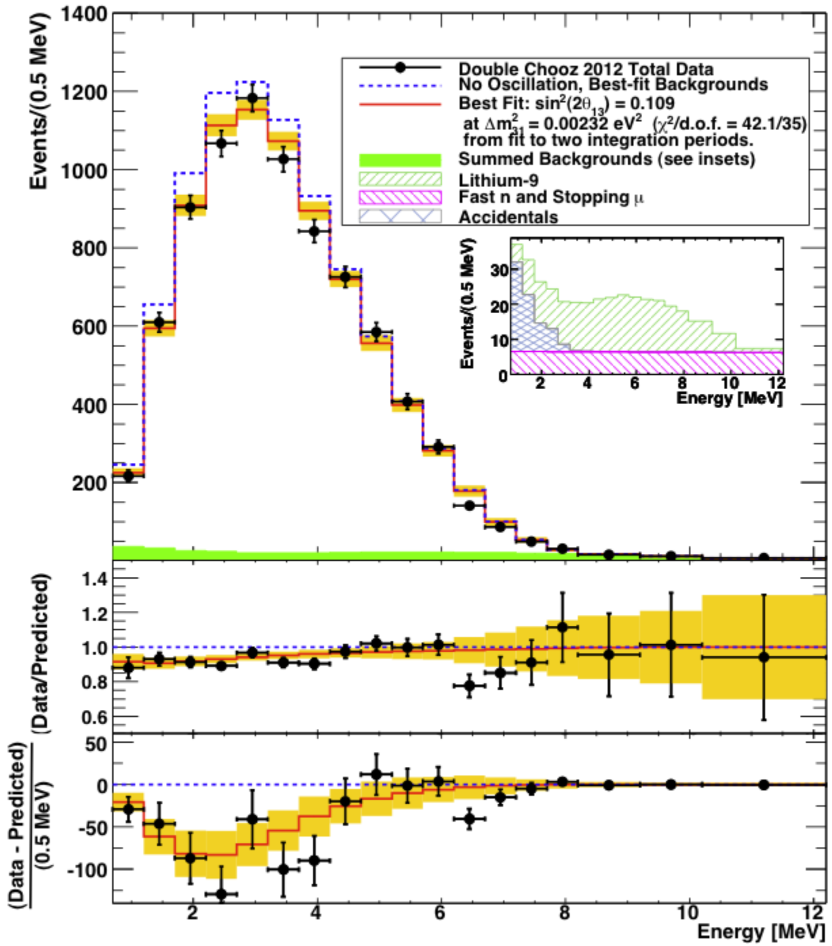
Amount of ν_e in faster oscillations (θ_{13})



Amount of ν_e in fast oscillations (θ_{13})

Oscillation probability depends on energy → search for energy-dependent depletion

- Double Chooz: liquid scintillator detector, 1 km away from reactors

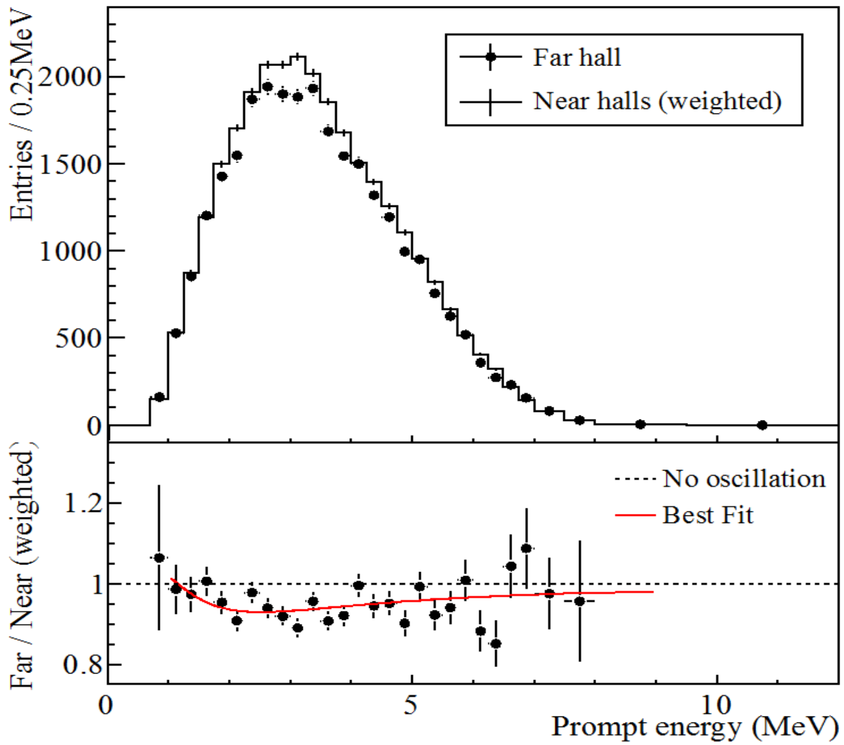


23/07/2013 GraS

Amount of ν_e in fast oscillations (θ_{13})

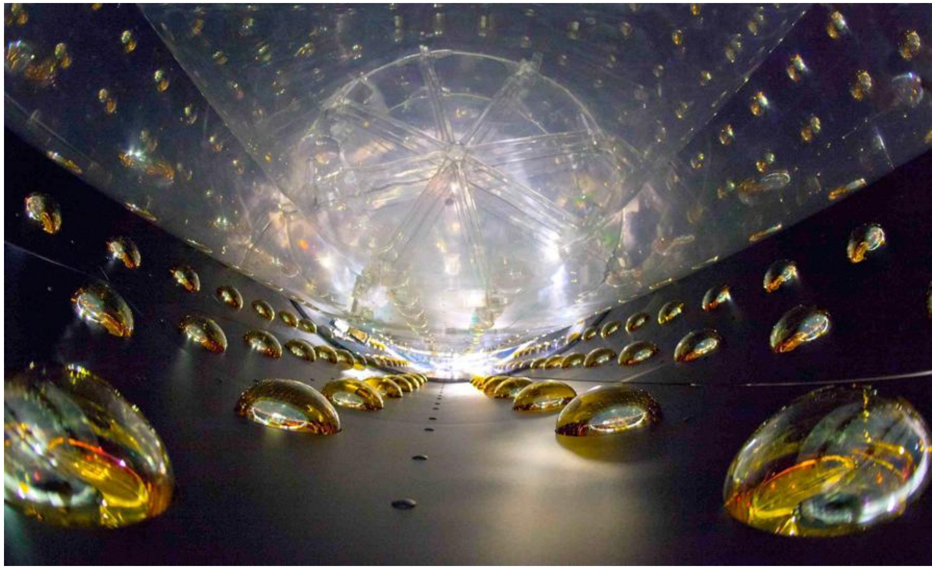
Oscillation probability depends on energy → search for energy-dependent depletion

- Daya Bay: very similar detector to Double Chooz and Reno, all 1-2 km away from reactors



$\sin^2(2\theta_{13}) = 0.089 \pm 0.012$

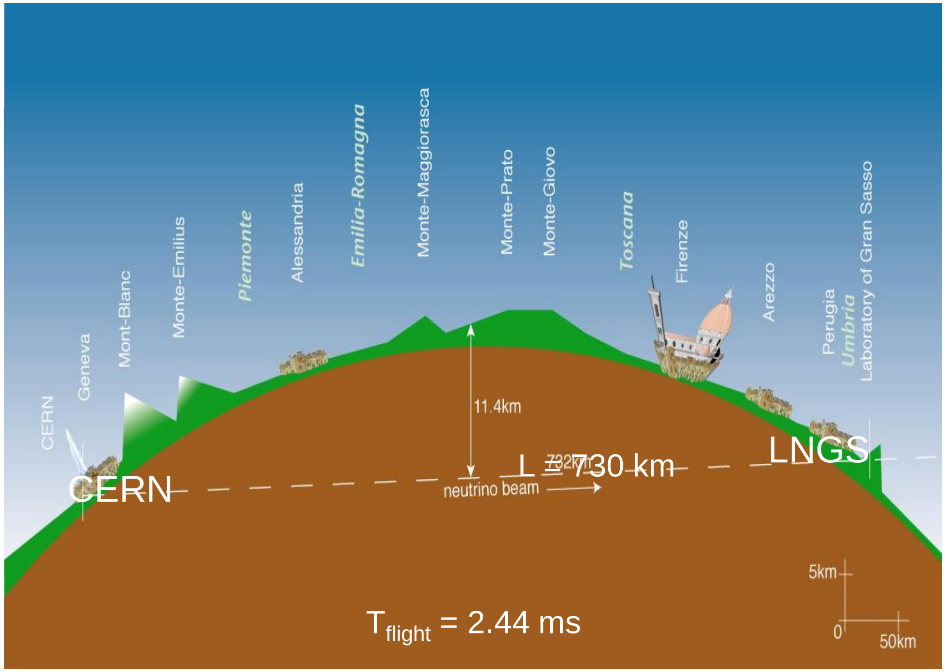
$\theta_{13} = 9.1° \pm 0.6°$

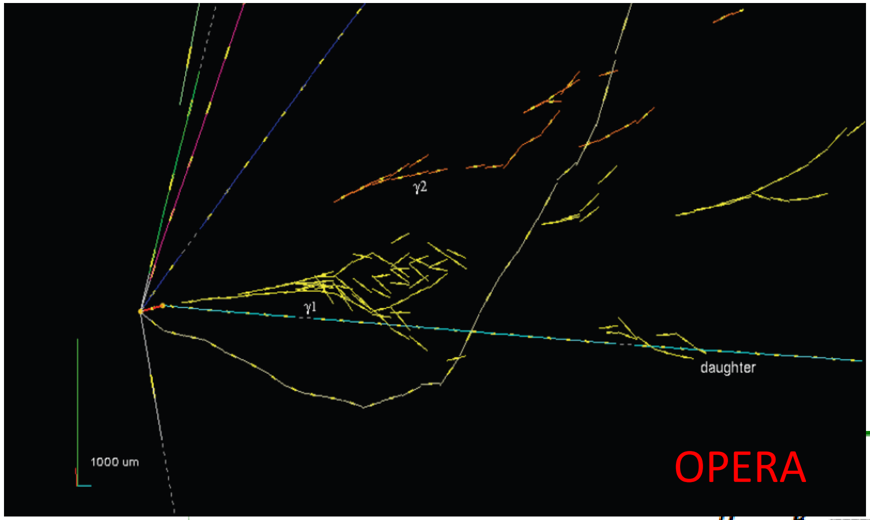


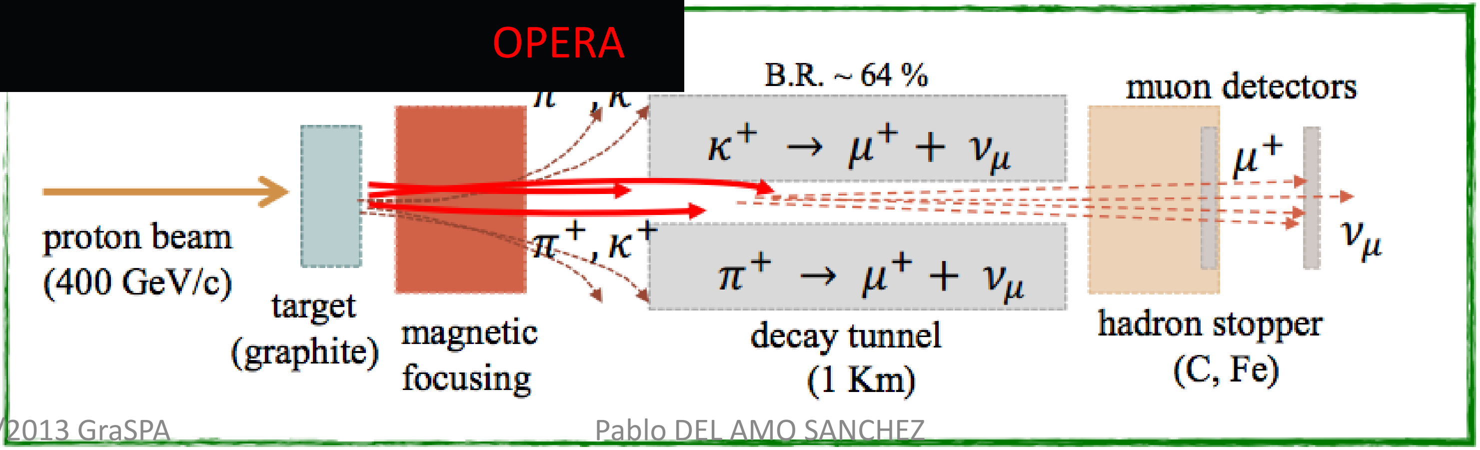
Accelerator experiments

- Can also produce neutrino beams:
- Results in excellent agreement with

other neutrino sources:







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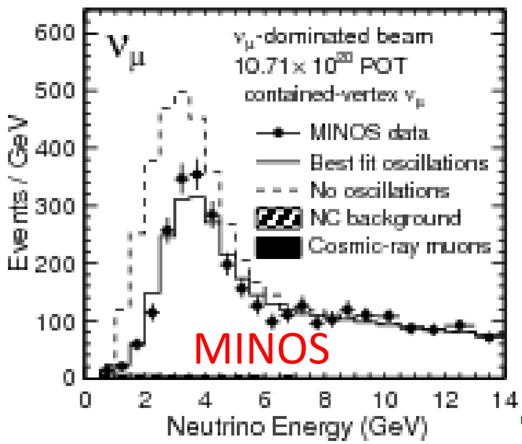
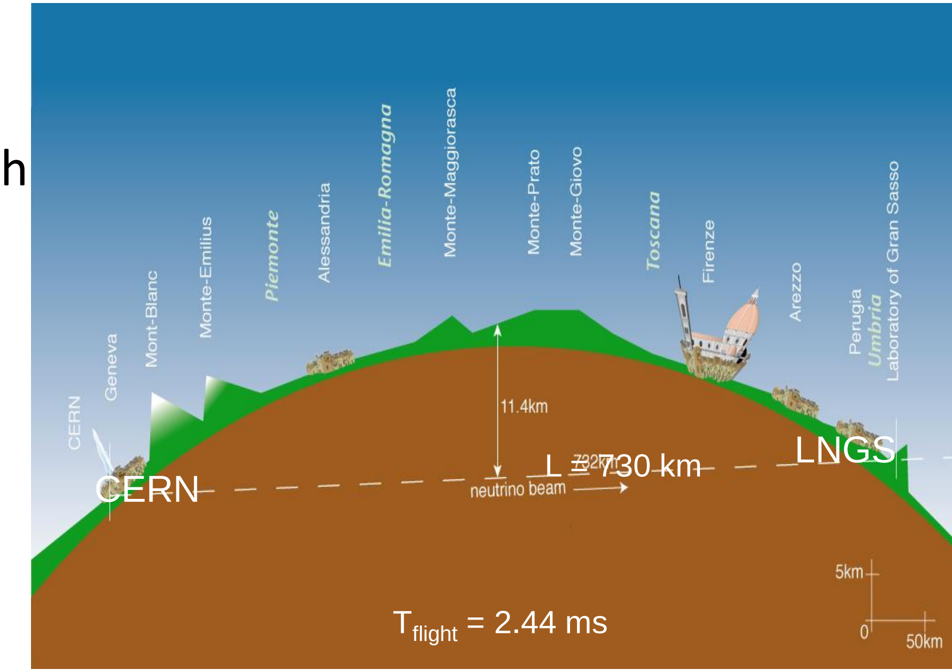
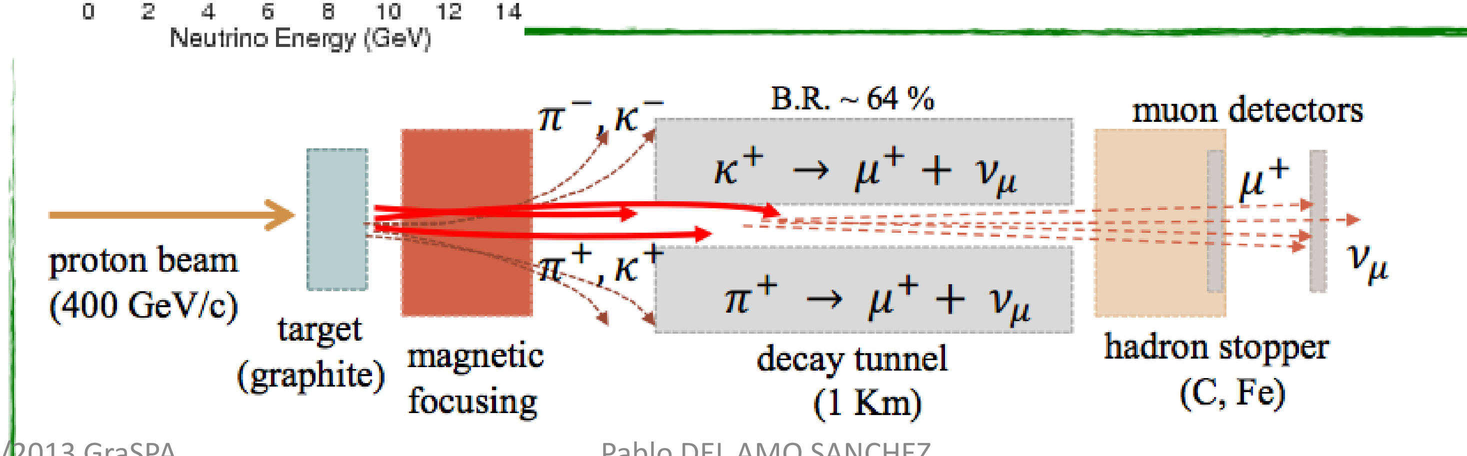
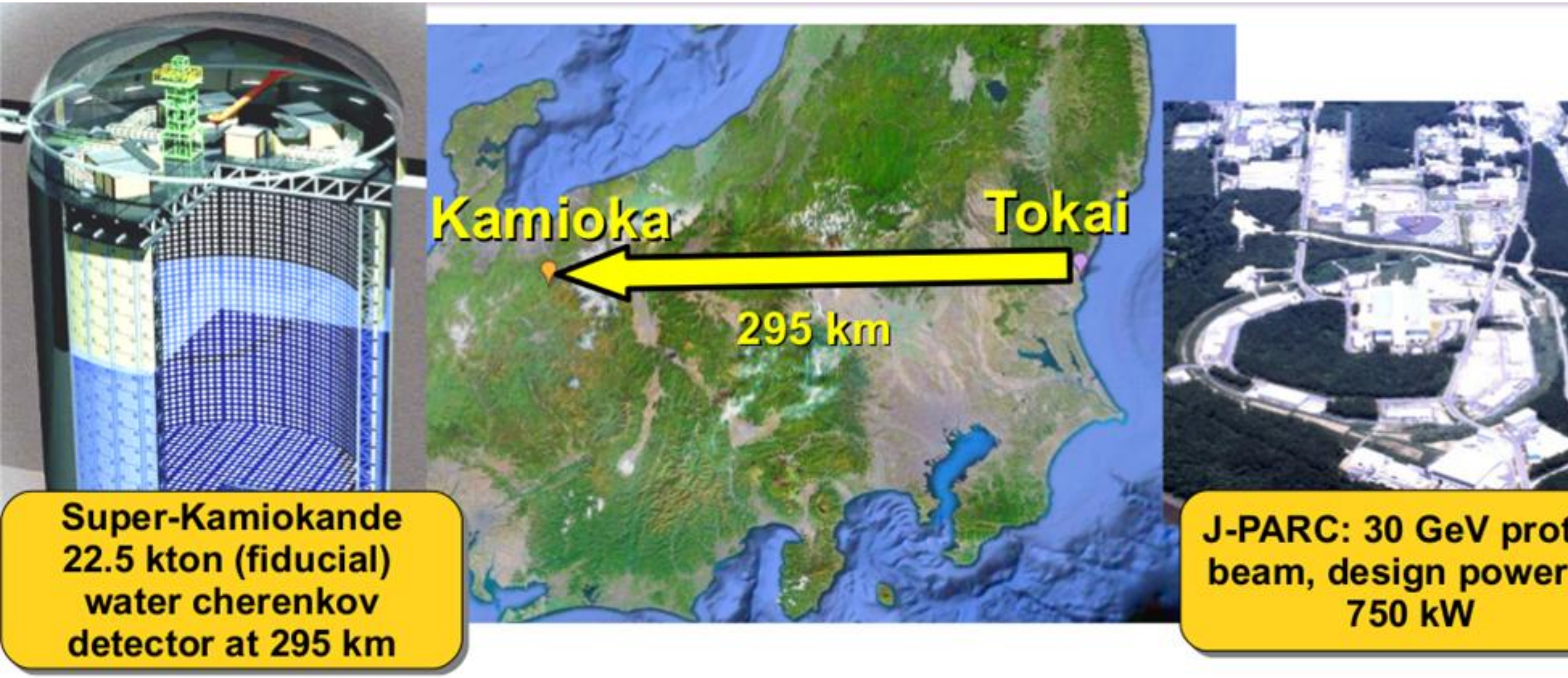






Recent results: ν_e appearance

- T2K observes 28 ν_e events, 4.6 background events expected
- Appearance of different flavour at 7.5 σ



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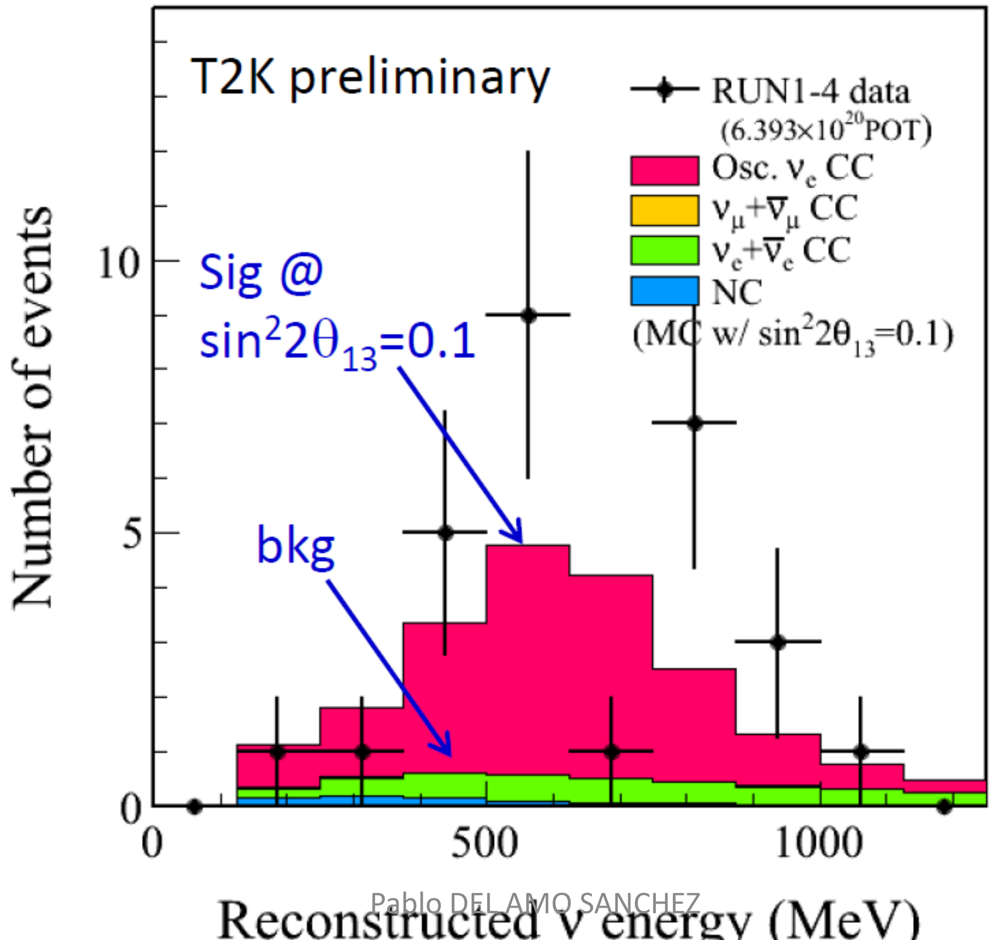


Neutrino mixing matrix

3 angles and 1 CP phase: $\theta_{12}, \theta_{13}, \theta_{23}, \delta$ + 2 phases

$$U_{PMNS} = \begin{pmatrix} 1 & 0 & 0 \\ 0 & c_{23} & s_{23} \\ 0 & -s_{23} & c_{23} \end{pmatrix} \begin{pmatrix} c_{13} & 0 & s_{13} \cdot e^{-i\delta} \\ 0 & 1 & 0 \\ -s_{13} \cdot e^{i\delta} & 0 & c_{13} \end{pmatrix} \begin{pmatrix} c_{12} & s_{12} & 0 \\ -s_{12} & c_{12} & 0 \\ 0 & 0 & 1 \end{pmatrix} \begin{pmatrix} 1 & 0 & 0 \\ 0 & e^{i\alpha_1} & 0 \\ 0 & 0 & e^{i\alpha} \end{pmatrix}$$

atmospheric ν

Reactor

solar ν

Dirac

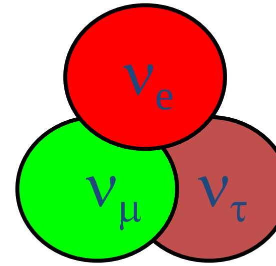
Majorana

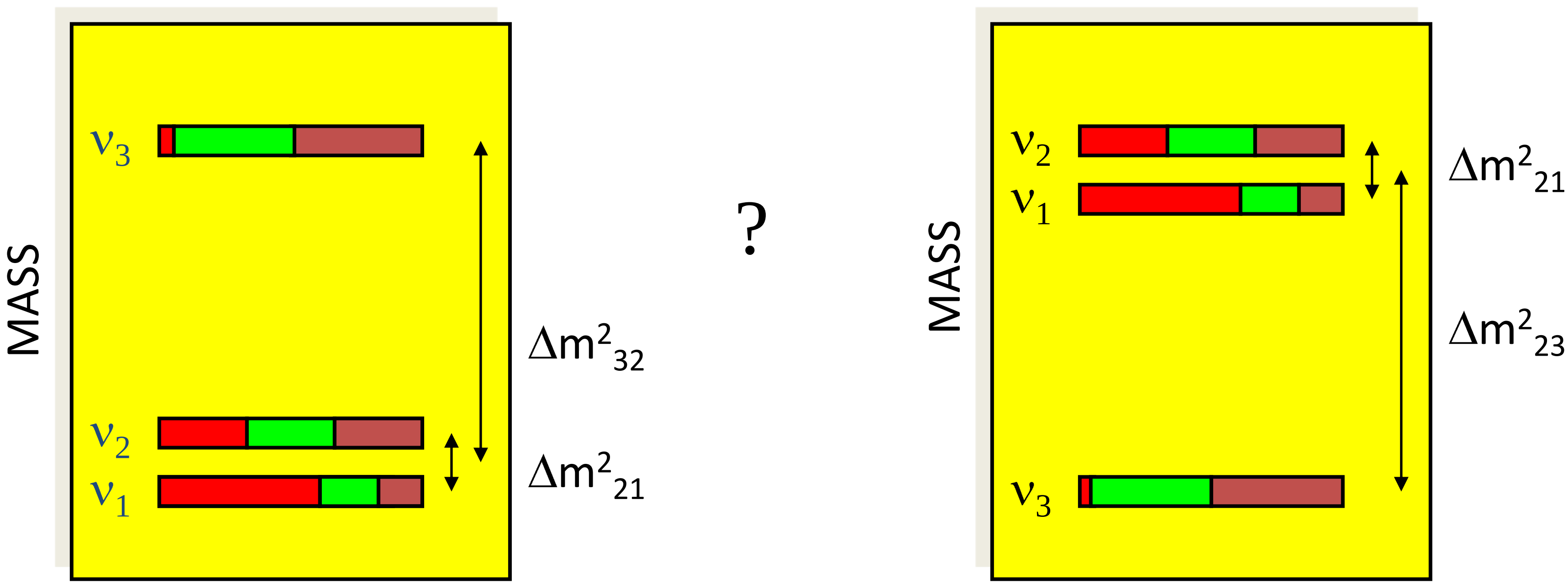
$$c_{ij} = \cos\theta_{ij}$$

$$s_{ij} = \sin\theta_{ij}$$

δ, matter-antimatter asymmetry in neutrinos?

Mass hierarchy?





Normal mass hierarchy

Inverted mass hierarchy

Which mass state is the lightest?

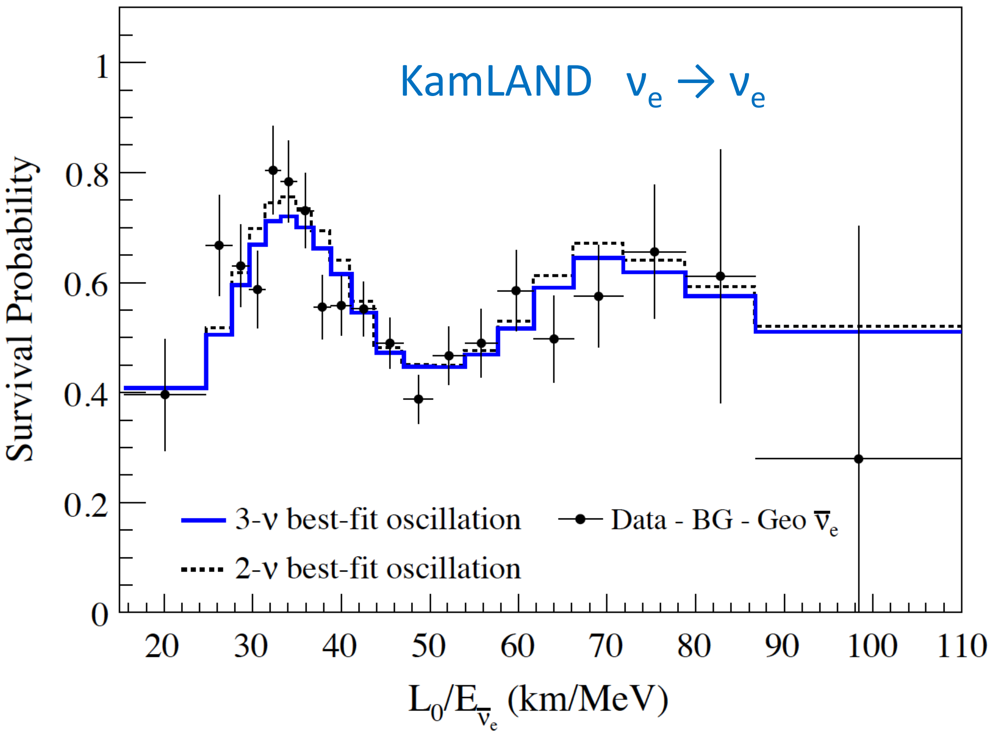
Future long baseline projects…

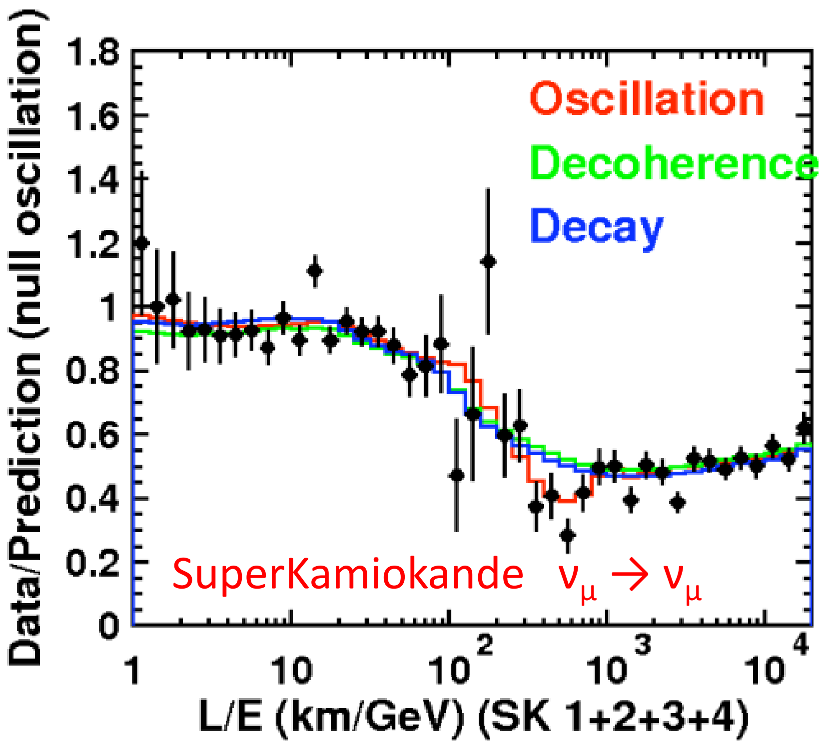
2100 km from RAL, 1500 km from DESY, and 1160 km from Protvino.

Conclusions

- Neutrinos oscillate! Masses ≠ 0

ν_e, ν_μ, ν_τ different from ν_1, ν_2, ν_3





Conclusions

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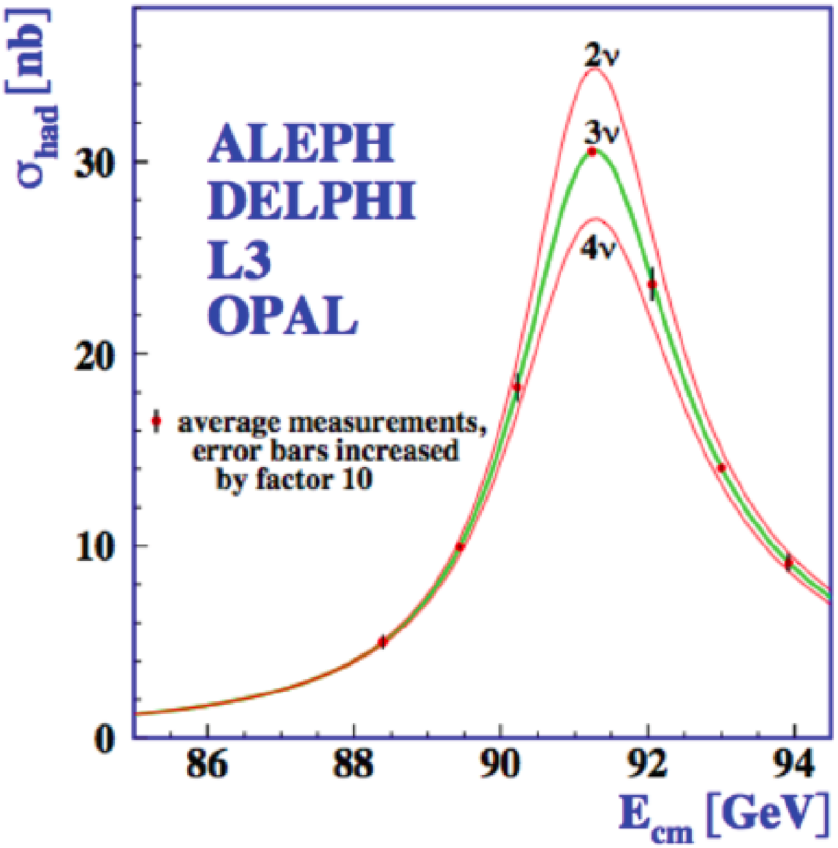
  reactor, not so small $\theta_{13} = 9.1° \pm 0.6°$

- For the future: matter-antimatter asymmetry in neutrinos?

  which is the lightest mass state?

BACK UP SLIDES

How many neutrinos are there?



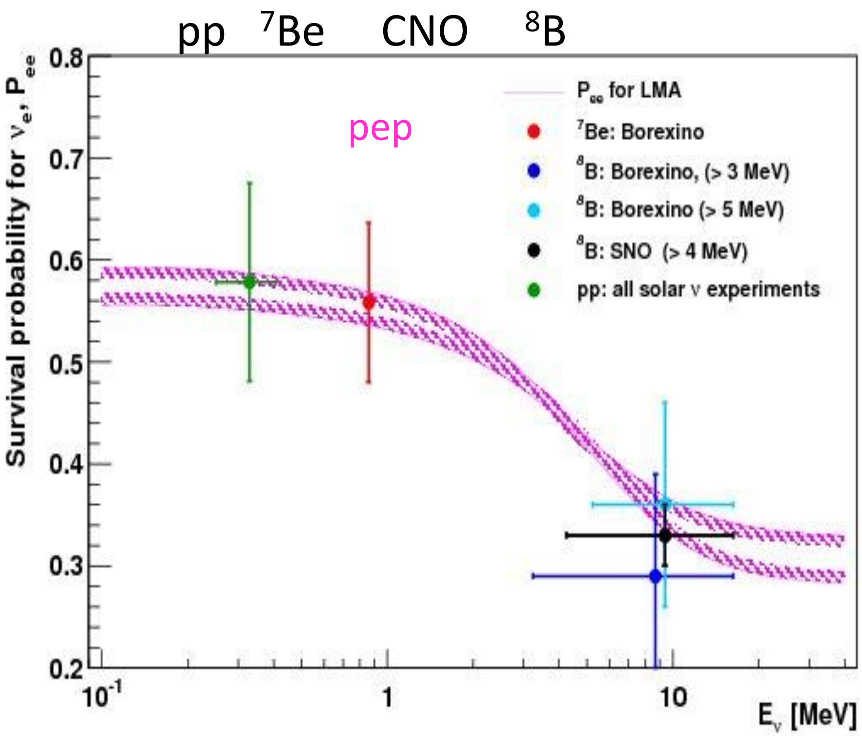
$$\Gamma_{\text{inv}} = \Gamma_Z - \Gamma_{\text{had}} - 3\Gamma_l$$

$$\Gamma_{\text{inv}} = N_\nu \cdot \Gamma_\nu$$

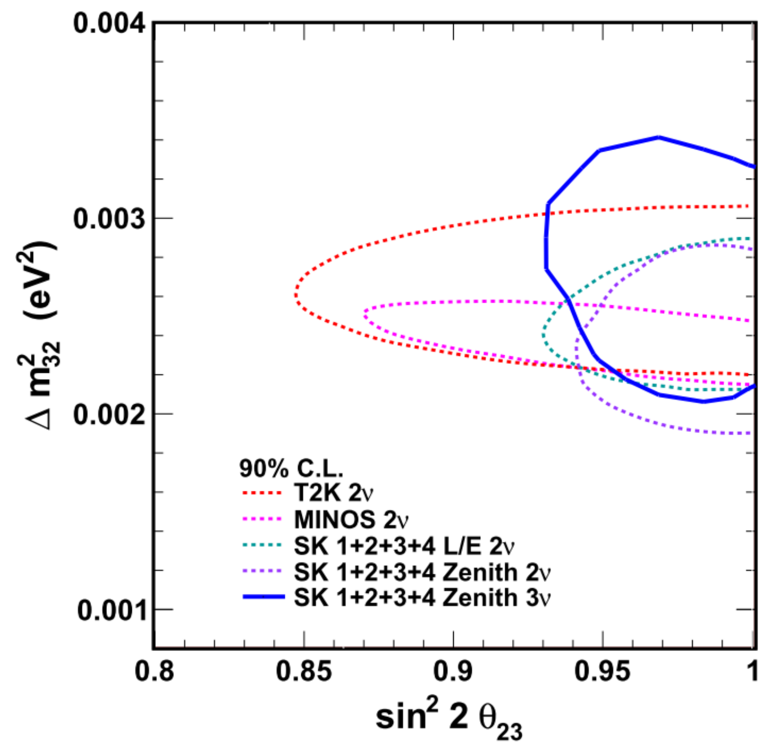
PDG K. Nakamura et el., JPG 37, 075021 (2010)

Number N = 2.984 ± 0.008
(Standard Model fits to LEP data)

Number N = 2.92 ± 0.05 (S=1.2)
(Direct measurement of invisible Z width)



pp
^{7}Be
CNO
^{8}B
pep
Survival probability for ν_e, P_{ee}
0.8
0.7
0.6
0.5
0.4
0.3
0.2
10^{-1}
1
10
E_ν [MeV]
P_{ee} for LMA
^{7}Be: Borexino
^{8}B: Borexino, (> 3 MeV)
^{8}B: Borexino (> 5 MeV)
^{8}B: SNO (> 4 MeV)
pp: all solar ν experiments



0.004
0.003
0.002
0.001
Δ m²₃₂ (eV²)
0.8
0.85
0.9
0.95
1
sin² 2 θ₂₃
90% C.L.
T2K 2ν
MINOS 2ν
SK 1+2+3+4 L/E 2ν
SK 1+2+3+4 Zenith 2ν
SK 1+2+3+4 Zenith 3ν

Neutrino candidates rate