

Non-Minimal Kaluza–Klein Dark Matter

Henrik Melb us

Royal Institute of Technology, Stockholm, Sweden

EPS-HEP 2011

Outline

1 Kaluza–Klein Dark Matter

2 Phenomenology

- Collaborators:

- ▶ Johan Bonnevier
- ▶ Mattias Blenow
- ▶ Alexander Merle
- ▶ Tommy Ohlsson

Universal Extra Dimensions

- One of the most popular extra-dimensional extensions of the SM
- One or more flat extra dimensions
- All SM fields promoted to higher-dimensional fields

Universal Extra Dimensions

- One of the most popular extra-dimensional extensions of the SM
- One or more flat extra dimensions
- All SM fields promoted to higher-dimensional fields
- Remnant of translational invariance
⇒ Conservation of KK parity $(-1)^n$
- Lower bound: $R^{-1} > 300 \text{ GeV}$

Kaluza–Klein Dark Matter

- KK parity \Rightarrow the Lightest Kaluza–Klein Particle (LKP) is stable
 - ▶ Analogous to the LSP in SUSY models with R parity

Kaluza–Klein Dark Matter

- KK parity \Rightarrow the Lightest Kaluza–Klein Particle (LKP) is stable
 - ▶ Analogous to the LSP in SUSY models with R parity
- $m_{\text{LKP}} \simeq R^{-1} = \mathcal{O}(1 \text{ TeV})$
- Possible WIMP DM candidate

The LKP

- The identity of the LKP depends on the UED mass spectrum
- Minimal UED: B^1 is the LKP (Cheng, Matchev, Schmaltz [hep-ph/0204342])
 - ▶ Compare to assumptions regarding SUSY breaking terms
- Non-minimal UEDs could give other LKPs (Flacke, Menon, Phalen [0811.1598])
- Possible WIMP DM candidates:
 - ▶ B^1
 - ▶ Z^1
 - ▶ H^1
 - ▶ ν^1

The LKP

- The identity of the LKP depends on the UED mass spectrum
- Minimal UED: B^1 is the LKP (Cheng, Matchev, Schmaltz [hep-ph/0204342])
 - ▶ Compare to assumptions regarding SUSY breaking terms
- Non-minimal UEDs could give other LKPs (Flacke, Menon, Phalen [0811.1598])
- Possible WIMP DM candidates:
 - ▶ B^1
 - ▶ Z^1
 - ▶ H^1
 - ▶ ν^1 — Excluded by direct detection (Servant, Tait [hep-ph/0209262])

Relic abundance

- The masses giving the correct relic abundance have been calculated to different levels of precision:

Relic abundance

- The masses giving the correct relic abundance have been calculated to different levels of precision:
- B^1
 - ▶ Coannihilations with all level-1 KK particles: $m_{B^1} \simeq 500 - 1600$ GeV
(Burnell, Kribs [hep-ph/0509118], Kong, Matchev [hep-ph/0509119])
 - ▶ Adding resonances from second KK-level shifts the preferred mass to a somewhat higher value
(Belanger, Kakizaki, Pukhov [1012.2577])

Relic abundance

- The masses giving the correct relic abundance have been calculated to different levels of precision:
- B^1
 - ▶ Coannihilations with all level-1 KK particles: $m_{B^1} \simeq 500 - 1600$ GeV
(Burnell, Kribs [hep-ph/0509118], Kong, Matchev [hep-ph/0509119])
 - ▶ Adding resonances from second KK-level shifts the preferred mass to a somewhat higher value
(Belanger, Kakizaki, Pukhov [1012.2577])
- Z^1
 - ▶ Coannihilations with all level-1 KK particles: $m_{Z^1} \simeq 1800 - 2500$ GeV
(Arrenberg, Baudis, Kong, Matchev, Yoo [0805.4210])

Relic abundance

- The masses giving the correct relic abundance have been calculated to different levels of precision:
 - B^1
 - ▶ Coannihilations with all level-1 KK particles: $m_{B^1} \simeq 500 - 1600$ GeV
(Burnell, Kribs [hep-ph/0509118], Kong, Matchev [hep-ph/0509119])
 - ▶ Adding resonances from second KK-level shifts the preferred mass to a somewhat higher value
(Belanger, Kakizaki, Pukhov [1012.2577])
 - Z^1
 - ▶ Coannihilations with all level-1 KK particles: $m_{Z^1} \simeq 1800 - 2500$ GeV
(Arrenberg, Baudis, Kong, Matchev, Yoo [0805.4210])
 - h^1
 - ▶ Without coannihilations: $m_{h^1} \simeq 3$ TeV (preliminary)

Direct Detection

- B^1 :
 - ▶ $\sigma_{B^1, p} \simeq 10^{-6}$ pb for $m_{B^1} = 1$ TeV (Cheng, Feng, Matchev [hep-ph/0207125])
 - ▶ Possibly detectable, depending on the KK mass spectrum

Direct Detection

- B^1 :
 - ▶ $\sigma_{B^1,p} \simeq 10^{-6}$ pb for $m_{B^1} = 1$ TeV (Cheng, Feng, Matchev [hep-ph/0207125])
 - ▶ Possibly detectable, depending on the KK mass spectrum
- Z^1 :
 - ▶ $\sigma_{Z^1,p} \simeq 10^{-8}$ pb for $m_{B^1} = 2$ TeV (Arrenberg, Baudis, Kong, Matchev, Yoo [0805.4210])
 - ▶ Much more difficult to detect

Direct Detection

- B^1 :
 - ▶ $\sigma_{B^1,p} \simeq 10^{-6}$ pb for $m_{B^1} = 1$ TeV (Cheng, Feng, Matchev [hep-ph/0207125])
 - ▶ Possibly detectable, depending on the KK mass spectrum
- Z^1 :
 - ▶ $\sigma_{Z^1,p} \simeq 10^{-8}$ pb for $m_{B^1} = 2$ TeV (Arrenberg, Baudis, Kong, Matchev, Yoo [0805.4210])
 - ▶ Much more difficult to detect
- H^1 :
 - ▶ Unknown, but presumably small due to the large preferred mass

Neutrinos From the Sun

- Detection of neutrinos from the annihilation of WIMPs trapped in the Solar core

Neutrinos From the Sun

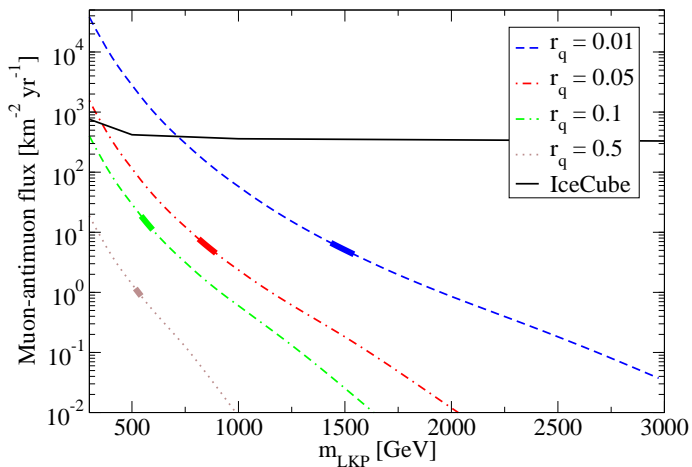
- Detection of neutrinos from the annihilation of WIMPs trapped in the Solar core
- B^1 :
 - ▶ Large branching ratios into leptons
 - ▶ Possibly observable at IceCube (Hooper, Kribs [hep-ph/0208261])

Neutrinos From the Sun

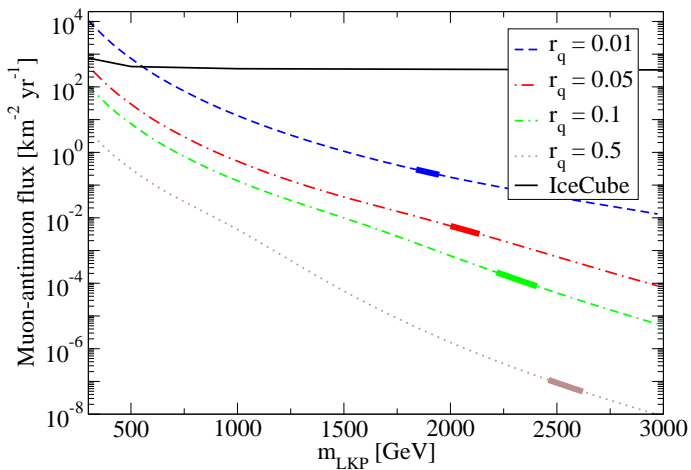
- Detection of neutrinos from the annihilation of WIMPs trapped in the Solar core
- B^1 :
 - ▶ Large branching ratios into leptons
 - ▶ Possibly observable at IceCube (Hooper, Kribs [hep-ph/0208261])
- Z^1 :
 - ▶ Large branching ratio into $W^+ W^-$
 - ▶ Not observable, due to suppression of the rate by its high mass (Flacke, Menon, Hooper, Dan, Freese [0908.0899], Blennow, H. M., Ohlsson [0910.1588])

Neutrinos From the Sun

- Detection of neutrinos from the annihilation of WIMPs trapped in the Solar core
- B^1 :
 - ▶ Large branching ratios into leptons
 - ▶ Possibly observable at IceCube (Hooper, Kribs [hep-ph/0208261])
- Z^1 :
 - ▶ Large branching ratio into $W^+ W^-$
 - ▶ Not observable, due to suppression of the rate by its high mass (Flacke, Menon, Hooper, Dan, Freese [0908.0899], Blennow, H. M., Ohlsson [0910.1588])
- H^1 :
 - ▶ Scalar, too strongly constrained by direct detection to be observable in indirect neutrino signals

Neutrinos From the Sun — B^1 

● (M. Blennow, H. M., T. Ohlsson [0910.1588])

Neutrinos from the Sun — Z^1 

● (M. Blennow, H. M., T. Ohlsson [0910.1588])

Monoenergetic Gamma-Rays

- Monoenergetic peak on top of the continuous spectrum

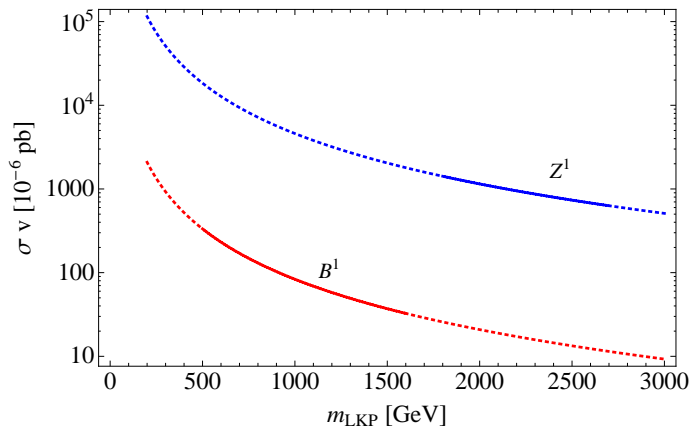
Monoenergetic Gamma-Rays

- Monoenergetic peak on top of the continuous spectrum
- B^1 :
 - ▶ Potentially observable by future high-resolution Air Cherenkov Telescopes (Bergström, Bringmann, Eriksson, Gustafsson [hep-ph/0412001])
 - ▶ Additional contribution from $Z\gamma$ and ZH is small (Bertone, Jackson, Shaughnessy, Tait, Vallinotto [1009.5107])

Monoenergetic Gamma-Rays

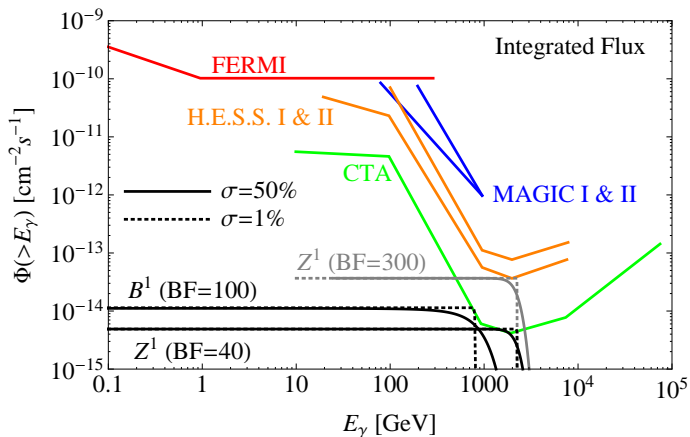
- Monoenergetic peak on top of the continuous spectrum
- B^1 :
 - ▶ Potentially observable by future high-resolution Air Cherenkov Telescopes (Bergström, Bringmann, Eriksson, Gustafsson [hep-ph/0412001])
 - ▶ Additional contribution from $Z\gamma$ and ZH is small (Bertone, Jackson, Shaughnessy, Tait, Vallinotto [1009.5107])
- Z^1 :
 - ▶ Similar conclusions as for B^1 , even though Z^1 is twice as massive (J. Bonnevier, H. M., A. Merle, T. Ohlsson [1104.1430])

Monoenergetic Gamma-Rays



- (J. Bonnevier, H. M., A. Merle, T. Ohlsson [1104.1430])

Monoenergetic Gamma-Rays



- (J. Bonnevier, H. M., A. Merle, T. Ohlsson [1104.1430])

Summary

- Possible non-minimal KKDM candidates are the B^1 , Z^1 and H^1
- B^1 : Possibly detectable in direct detection experiments, as well as indirect neutrino and gamma-ray searches
- Z^1 : Heavier than B^1 , and hence harder to detect, but perhaps possible in gamma-ray searches
- H^1 : Even heavier, and more difficult to detect

The End

- The End!
- Thank you for listening!