

Determination of muon multiplicity in ANTARES

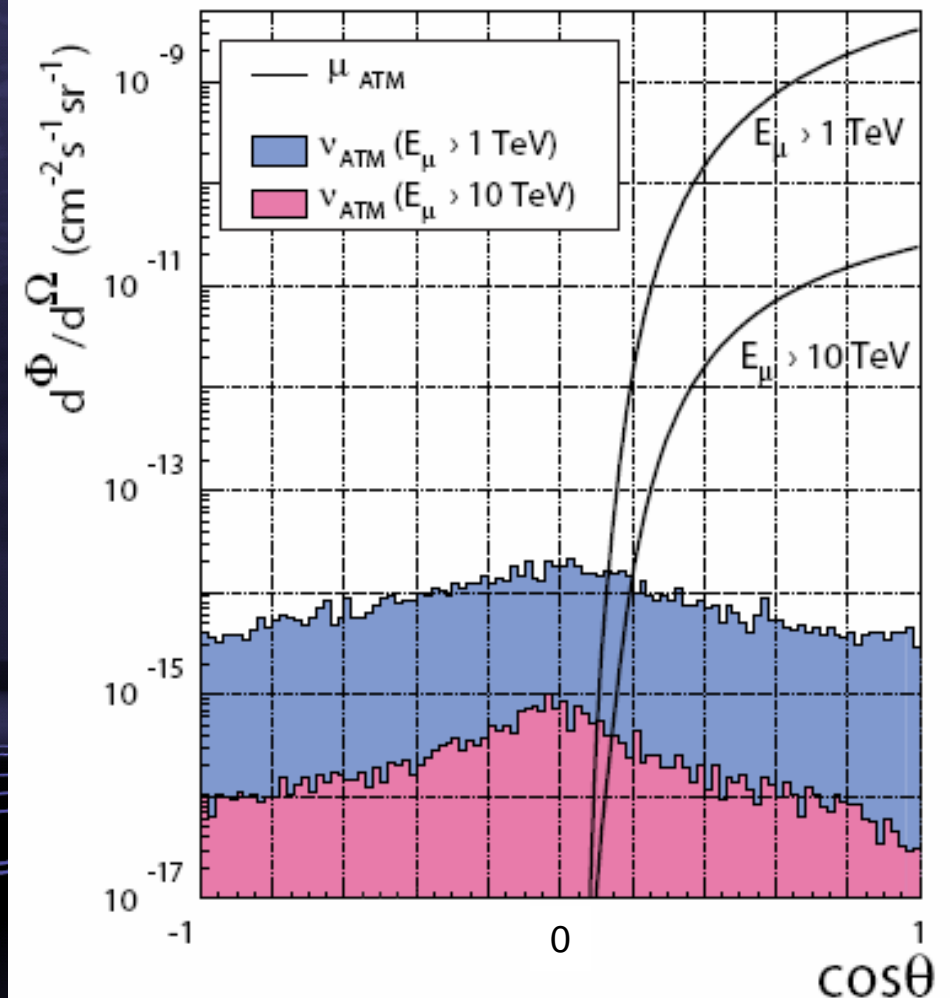
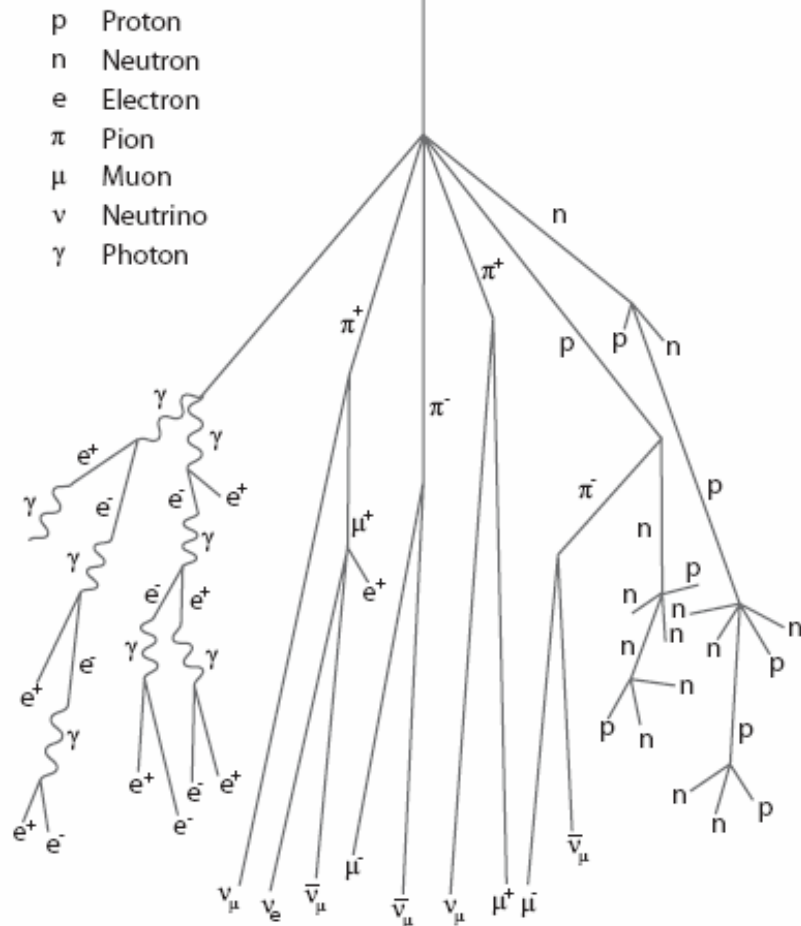
Claire Picq

DSM/IRFU/SPP-CEA Saclay and APC Paris 7



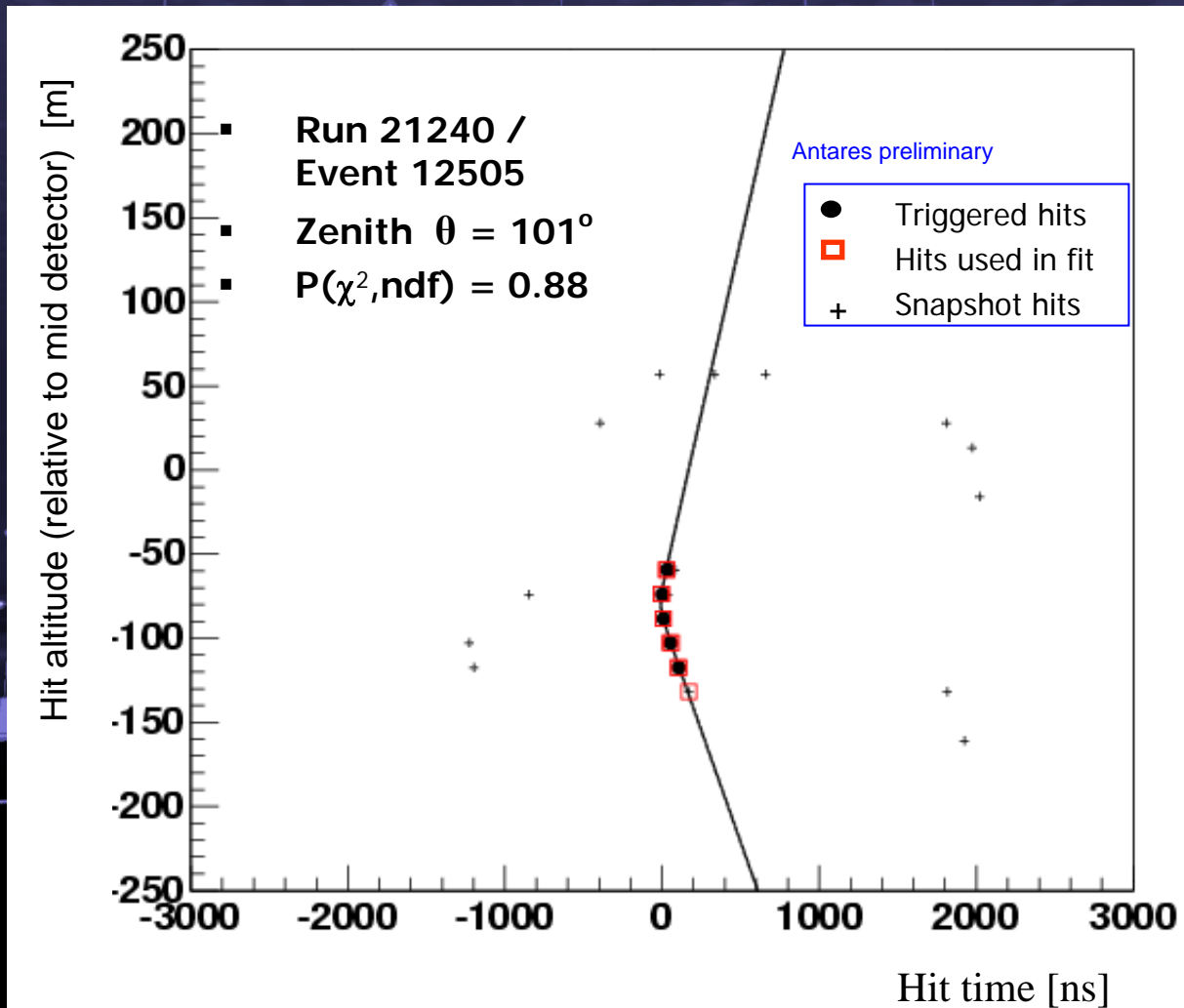
Atmospheric muons

Primary cosmic ray

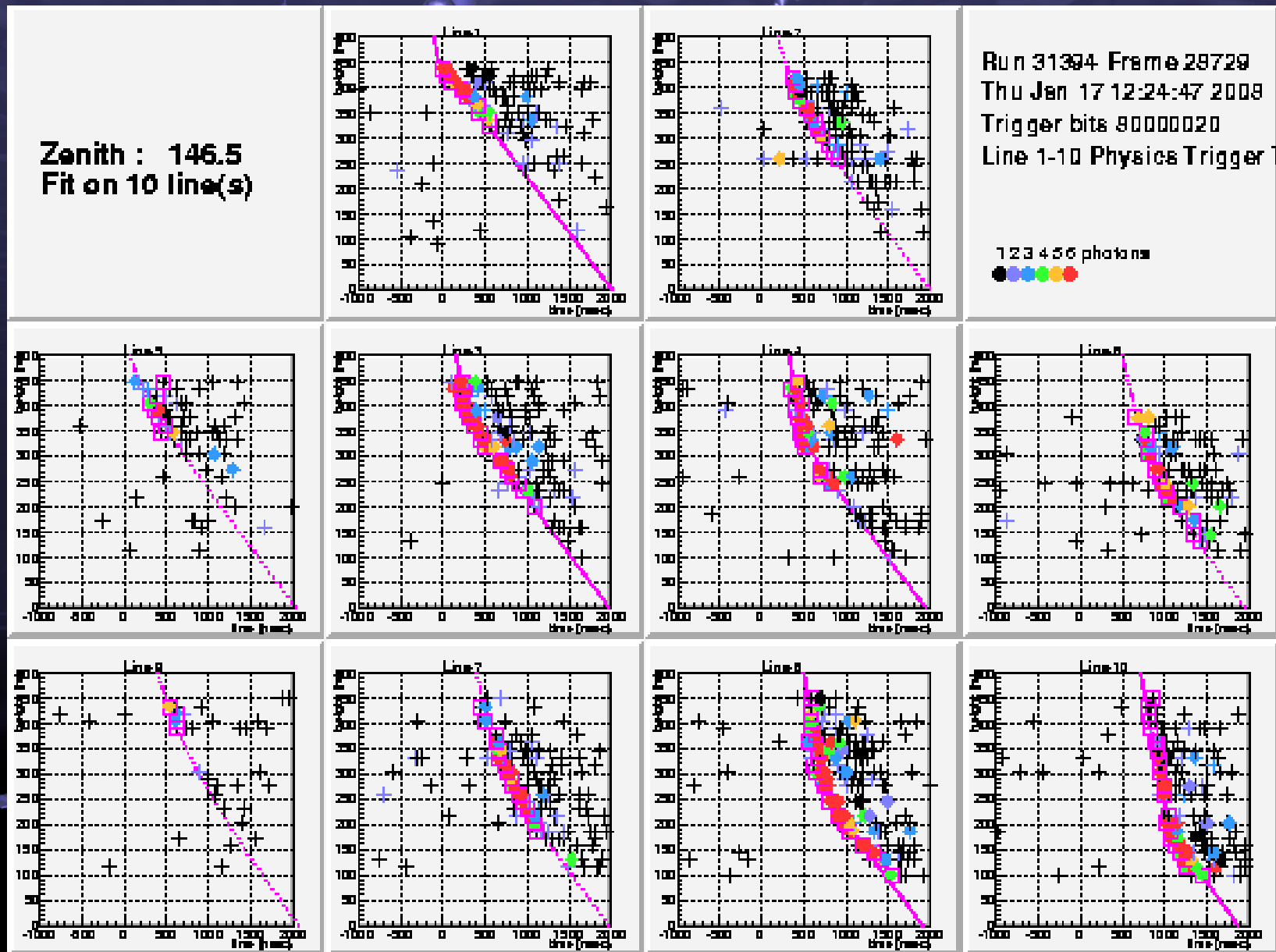


Atmospheric muon

likelihood maximization to estimate the zenith angle



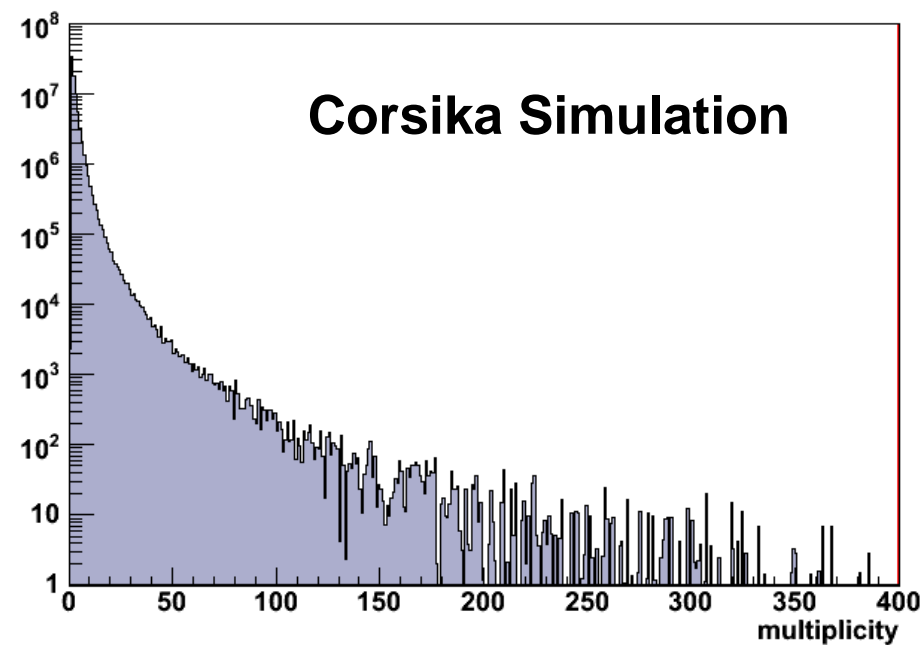
Online display: multi muon event



Goals of the study

- Isolate single muons for systematic studies purpose
- Estimator of multiplicity
- Measure the down-going muon flux

Multiplicity definition:
1 muon has given at
least 6 hits



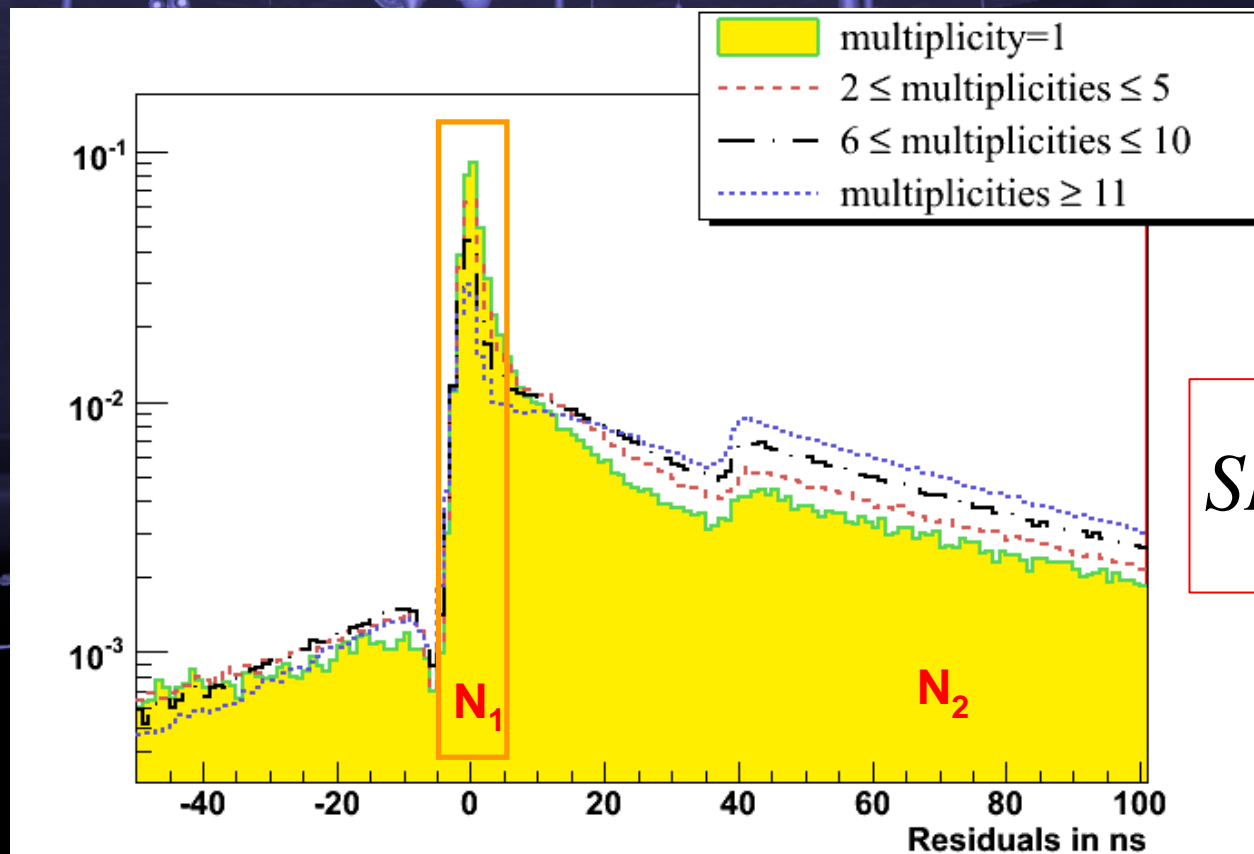
Time residuals

Time residual = reconstructed – real time

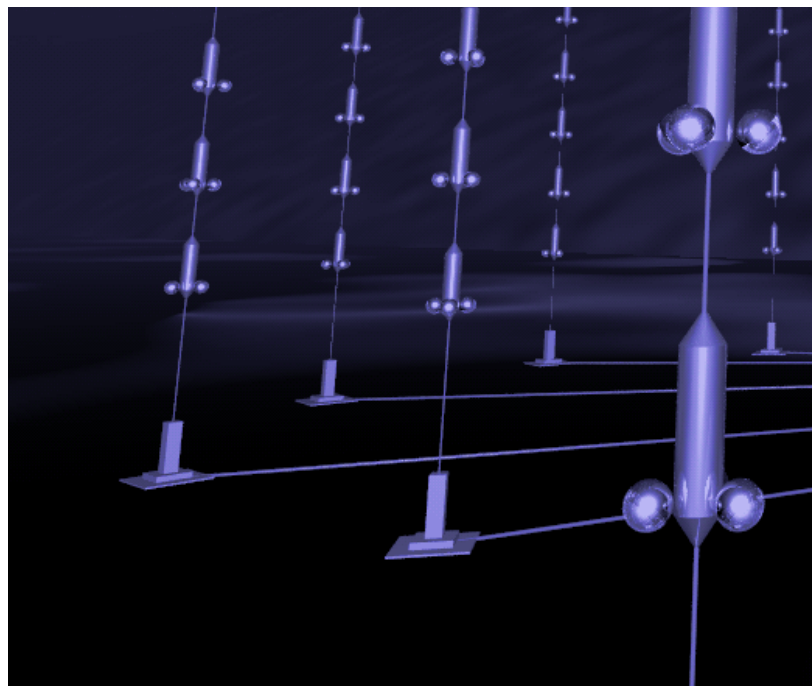
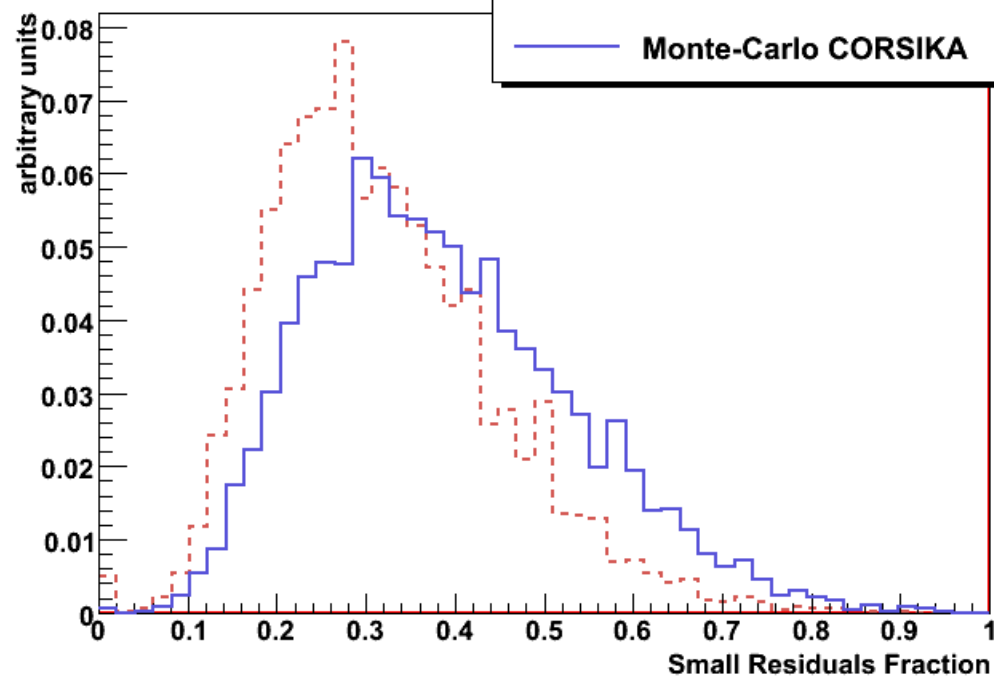
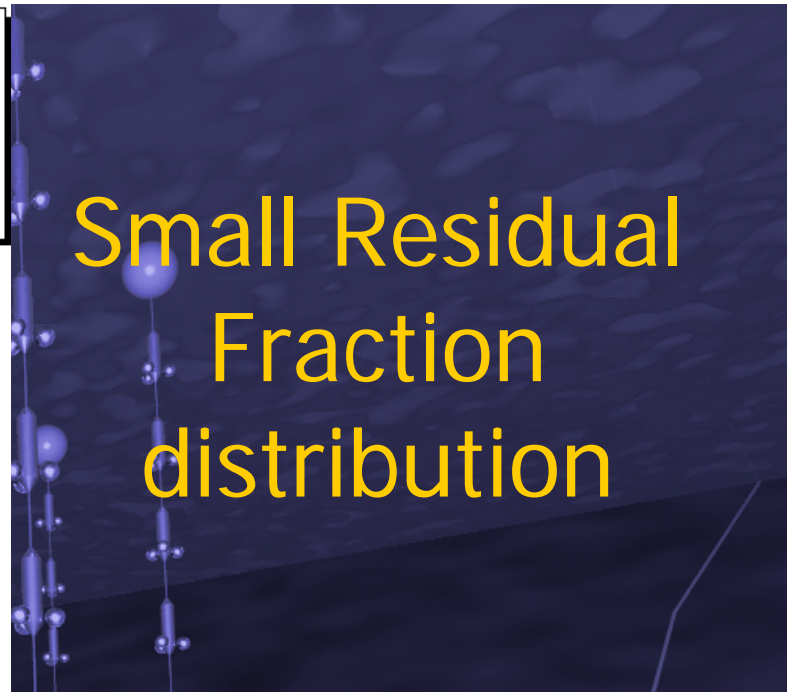
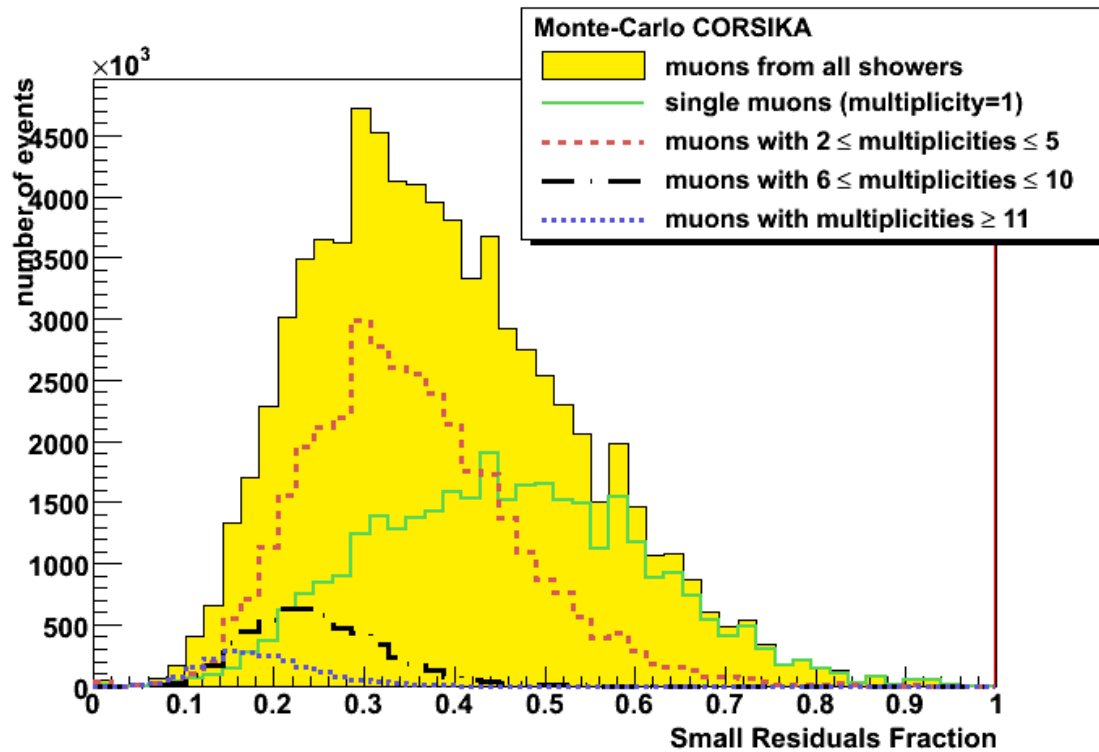
N_1 = number of hits with time residuals in $[-5 \text{ ns}; 5 \text{ ns}]$

N_2 = number of hits with time residuals in $[-50 \text{ ns}; 100 \text{ ns}]$

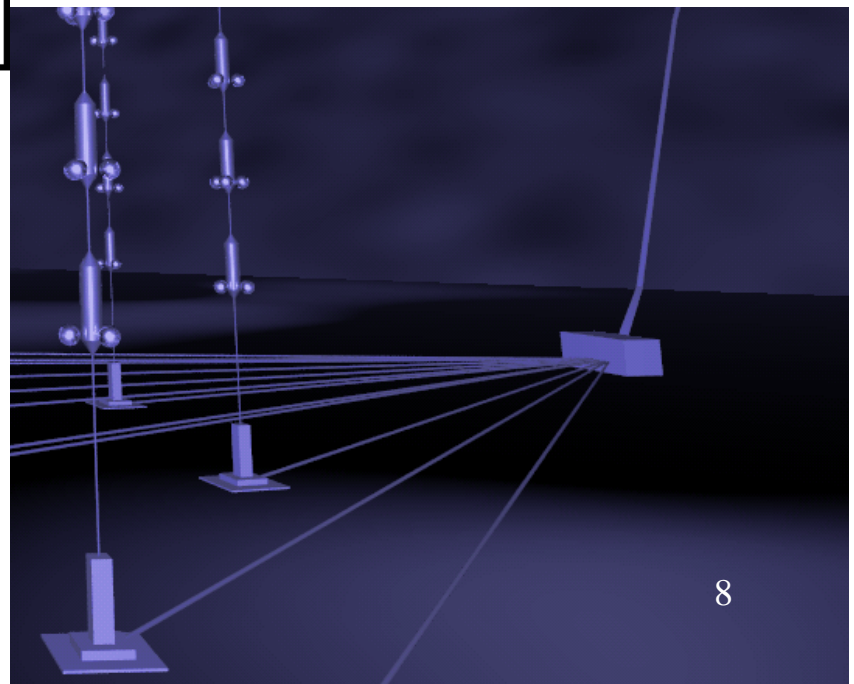
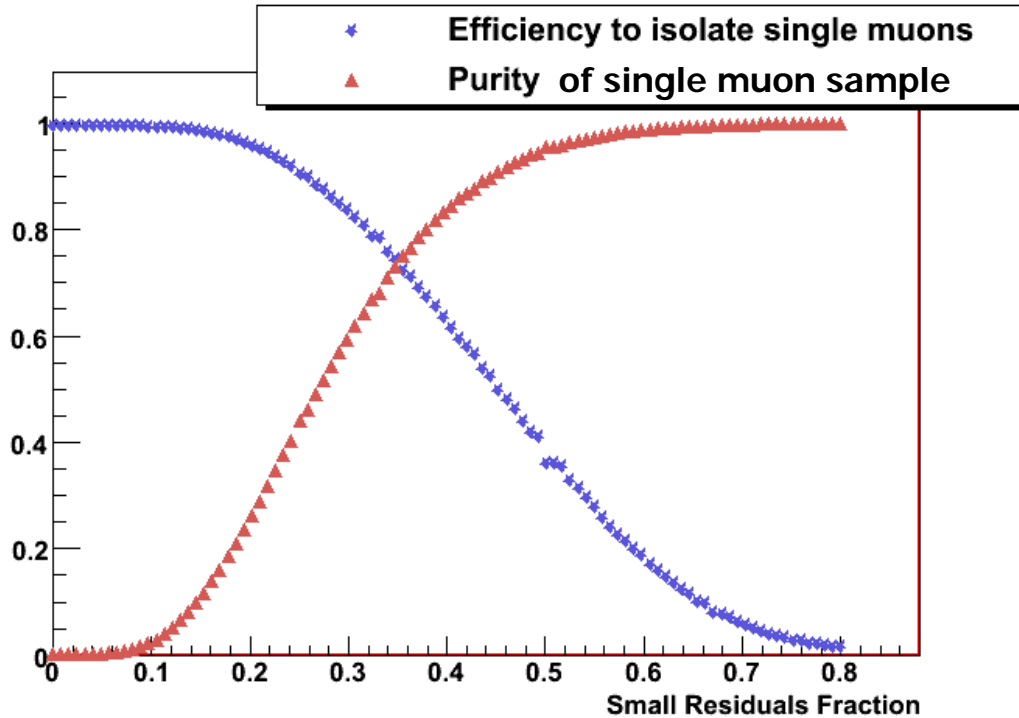
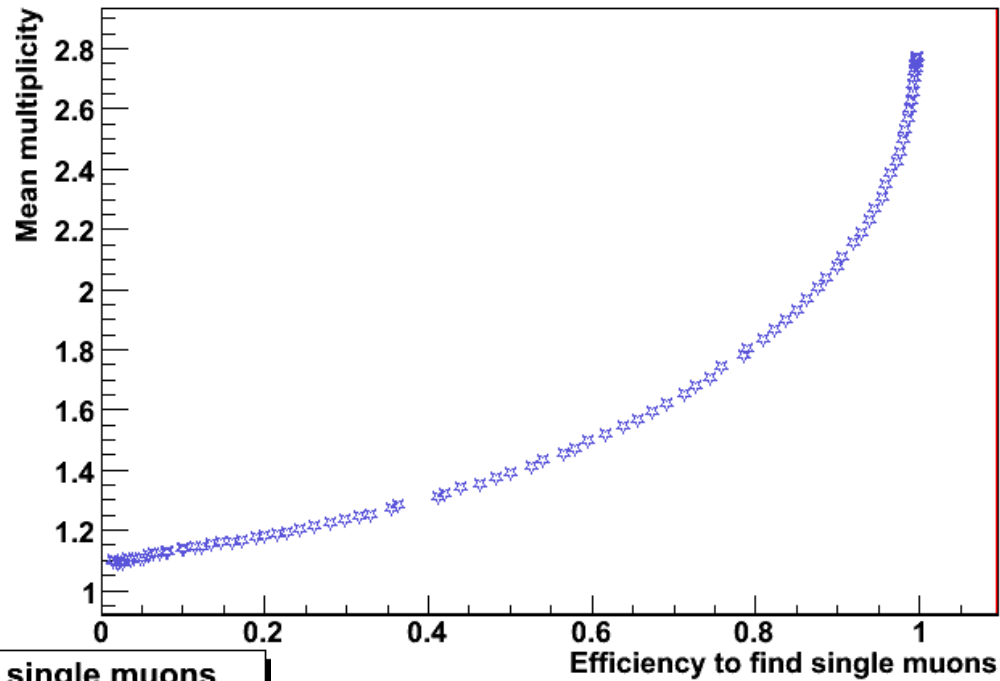
SRF = Small Residual Fraction



$$SRF = \frac{N_1}{N_2}$$



Efficiency and Purity



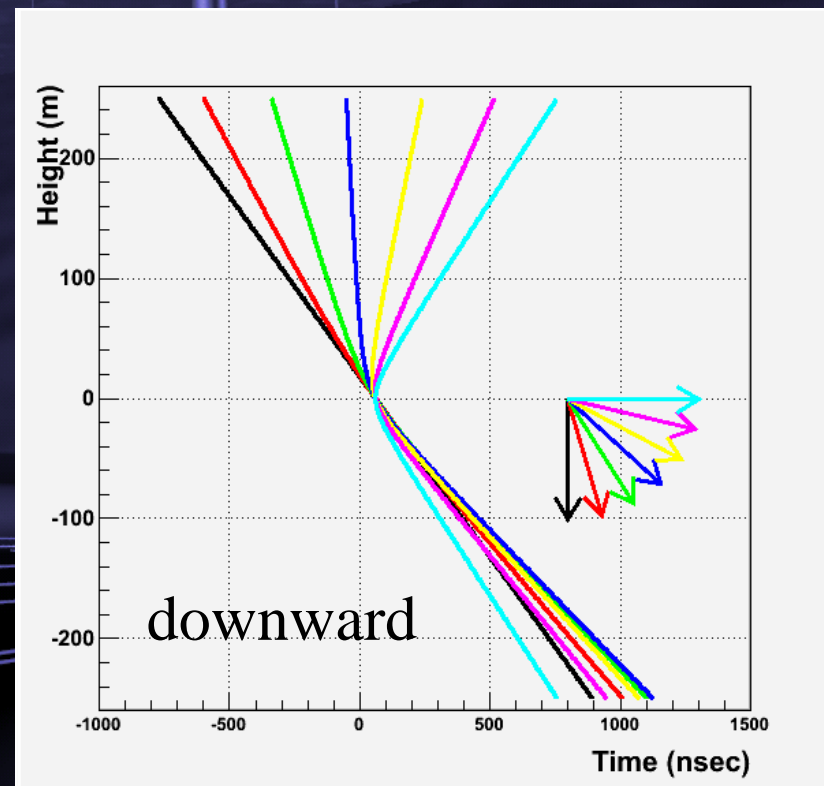
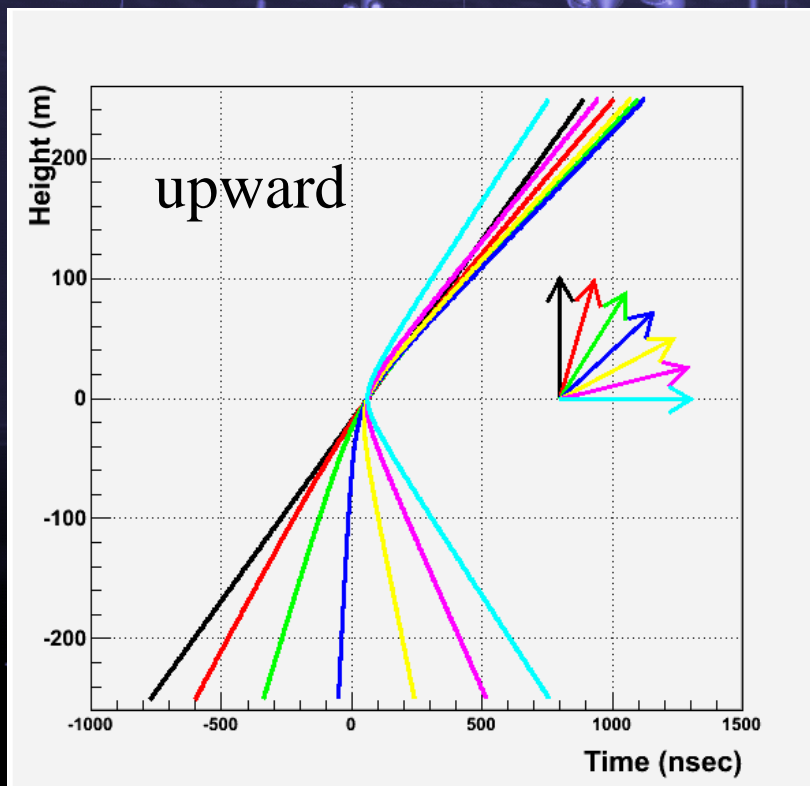
Backup

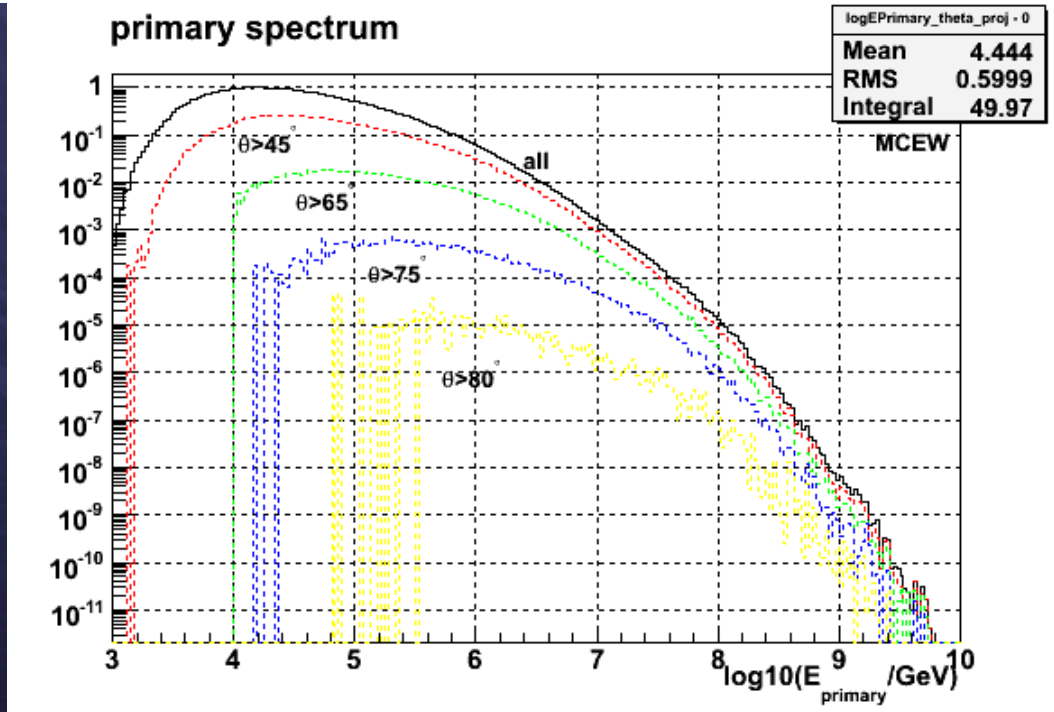
Reconstructed events

Height (z) vs time

⇒ Function of the zenith angle and distance

Many algorithms used by the collaboration 1D, 3D, χ^2 , ML

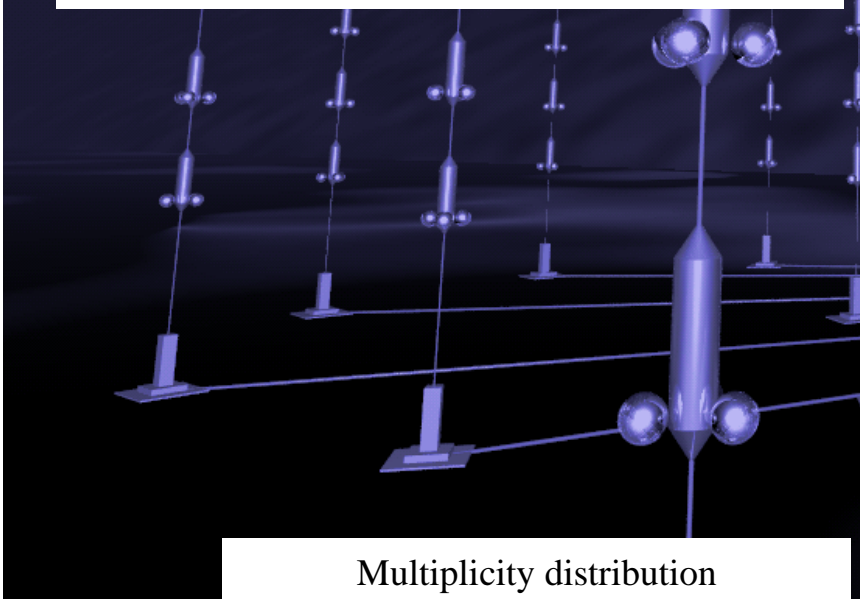
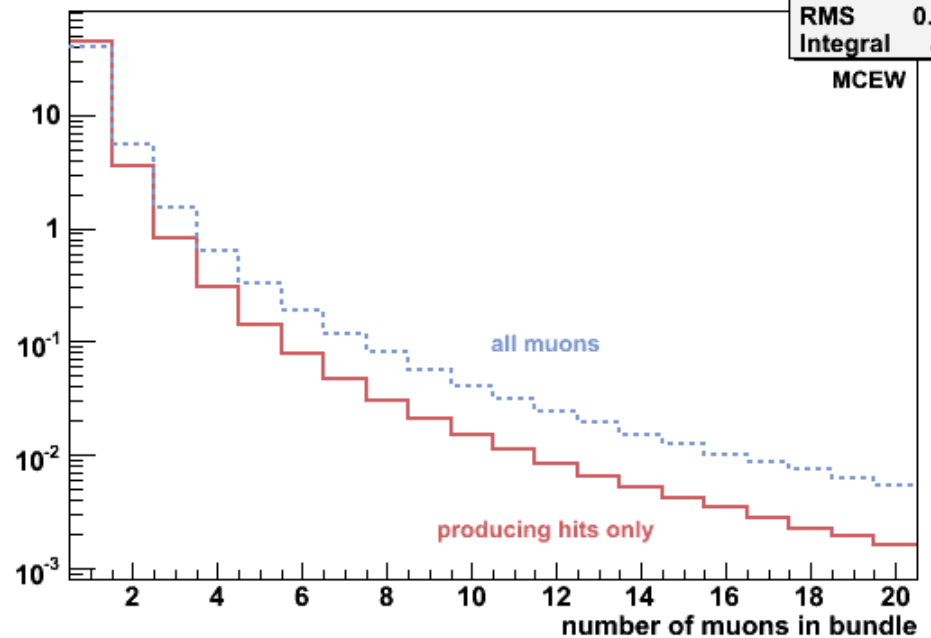




Monte Carlo 5 lines (no trigger)

Figure 1: Energy spectra of cosmic ray particles ("primaries") detected with an ANTARES line at one hit level (left plot) and trigger level (right plot). MC simulation. The colored curves correspond to the cuts on incident polar angle, as indicated on the plots.

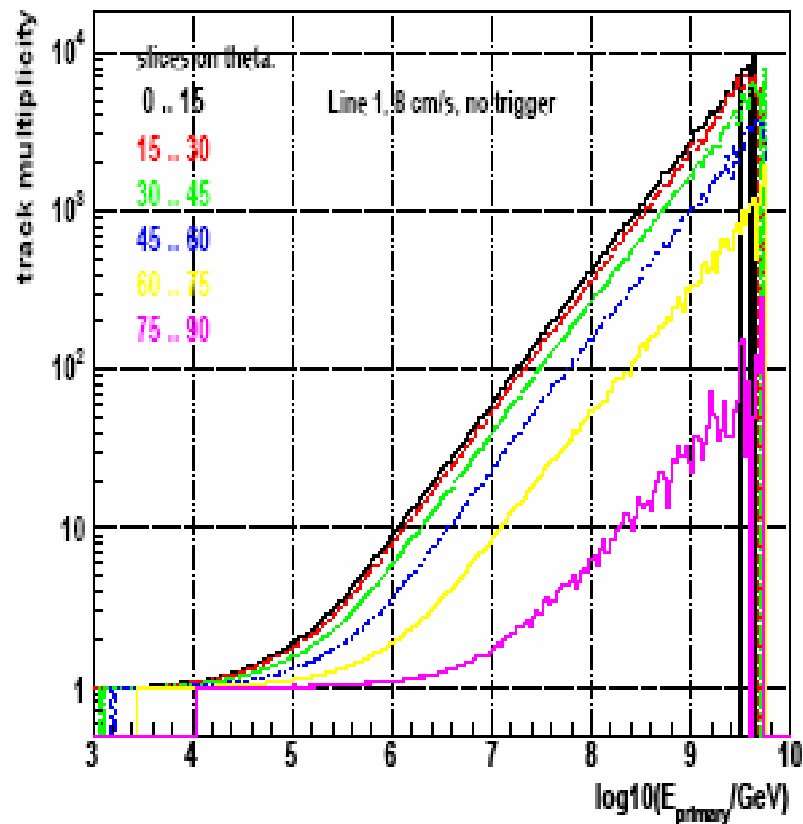
track multiplicity distribution



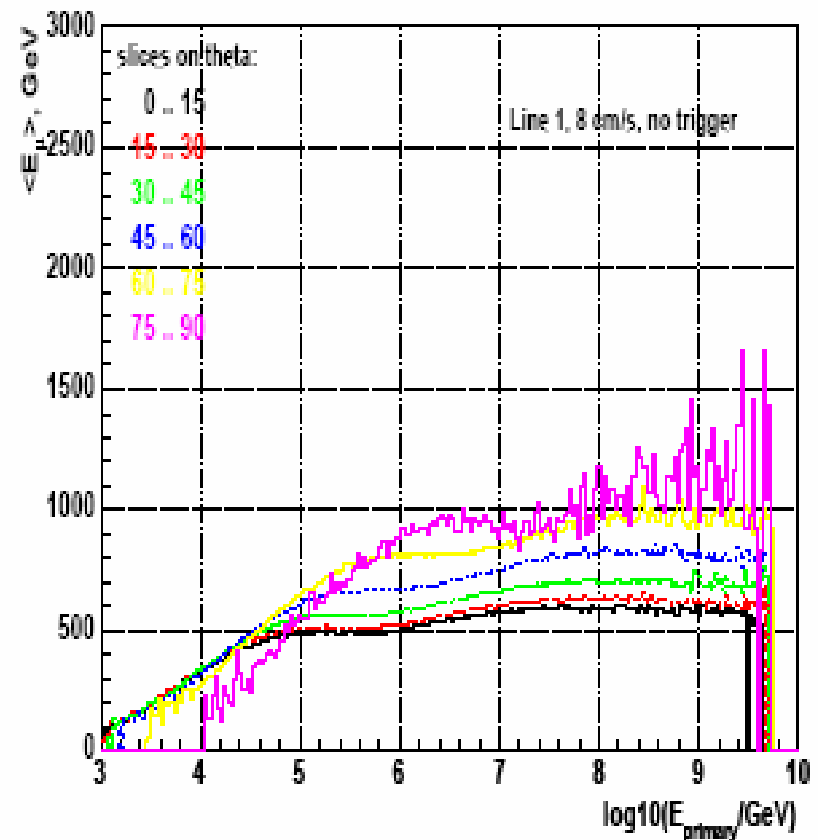
Multiplicity distribution

Monte Carlo 1 lines D. Zaborov

track multiplicity: θ slices



average muon energy



Multi muon event

