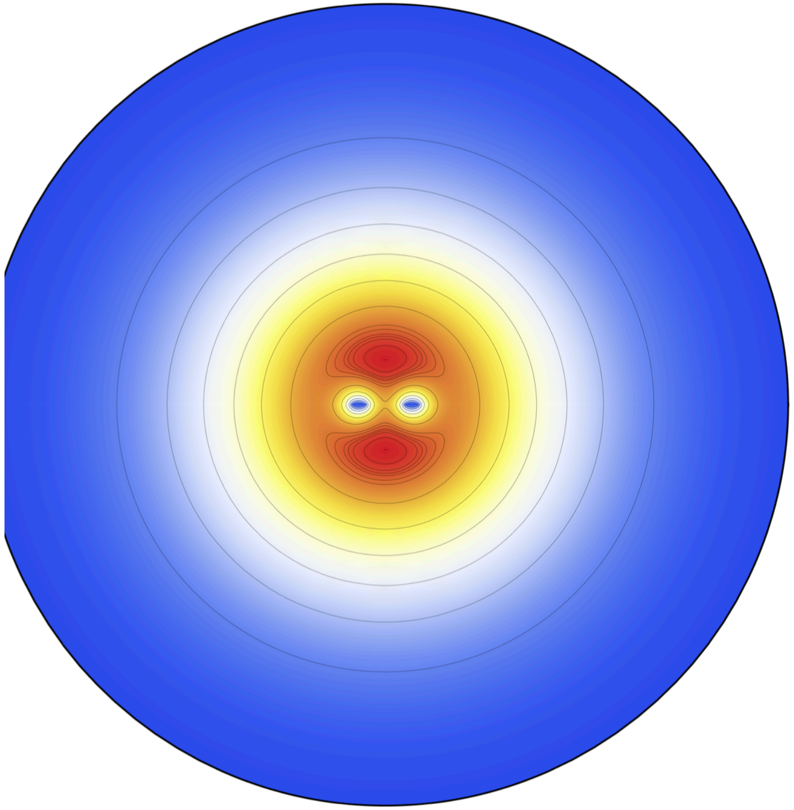
Black hole Binaries in an Expanding Universe

Óscar Dias

STAG RESEARCH CENTRE

UNIVERSITY OF Southampton



OD, Gary Gibbons, Jorge Santos, Benson Way,

2303.07361 (PRL) & 2406.10333 (PRL)

Black holes and their symmetries. University of Tours (July 2025)

→ Introduction & Motivation & short Summary

Our **Universe** appears to be undergoing an **accelerated expansion** due to the presence of a **positive cosmological constant, Λ>0** ...

(**ΛCDM model:** ~69% **vacuum energy** <—> **Λ>0** <—> **dark matter** with EoS p=w ρ, w~-1)

So we should ask:

- What is the phase space of **stationary black hole (BH)** solutions of the Einstein equation in **de Sitter**?
- Are there **other solutions** besides **de Sitter-Schwarzschild** & **de Sitter-Kerr** ?
- Can we have **multi-BHs** (eg **BH binaries**) with **Λ>0** ?

→ **Can we have <u>multi</u>-BHs (eg BH <u>binaries</u>) ?**

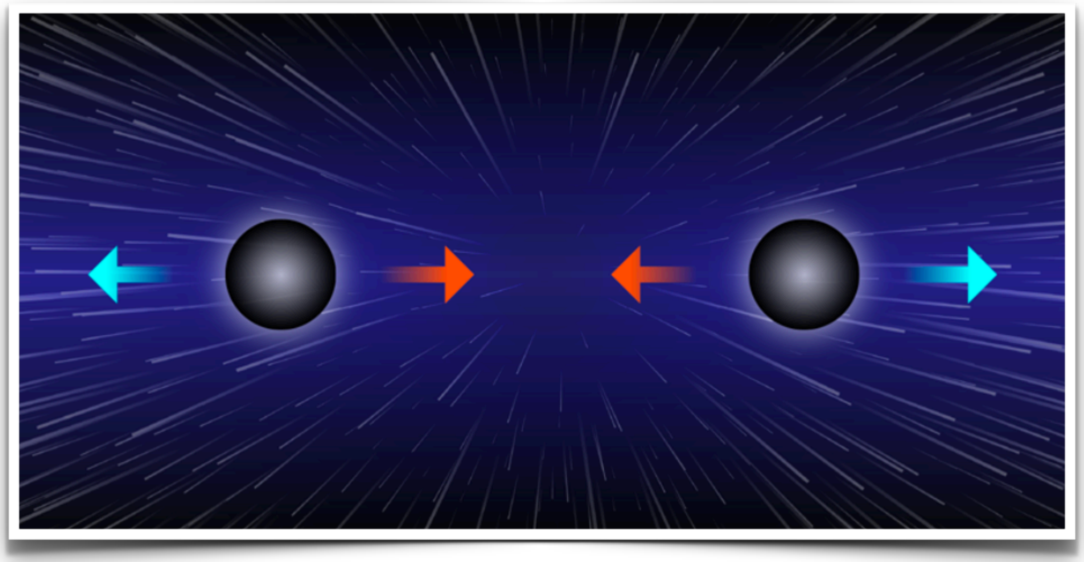
- On one hand, **Newton-Hooke analysis:** **cosmological expansion** should be able to **balance** **gravitational attraction**

- On the other hand, some **mathematical theorems** in the literature **claim uniqueness** of Schwarzschild/Kerr solutions in de Sitter !!!

[LeFloch, Rozoy '10] [Borghini, Chruściel, Mazzieri '19]
[ul Alam, Yu '14]

→ Solve the Einstein equations to **settle the issue**!

- We find that **regular static/stationary BH binaries do exist in de Sitter.**



- <u>**Not**</u> **in conflict** with available **Uniqueness theorems:** we have (explicitly identified) **assumptions** of these theorems that can be **evaded**

→ Λ=0: Uniqueness, No-hair theorems & multi-BHs

- When Λ=0, **Stationarity** => **axisymmetry** [**Hawking '73** and **Wald '92, Chrúsciel '23**]
  - => **No-hair & Uniqueness theorems** [**Kerr '67, Carter '71, Robinson '75**]
  - => BHs are uniquely characterized by their M, J, Q: the **Kerr-Newman BH** family

- For **static** configurations, mathematical theorems **preclude** the existence of regular asymptotically **flat multiple BHs** [**Bunting, Masood-Ul-Alam '97**].

- Asymptotically **flat multi-Kerr BHs**, where their gravitational attraction might be balanced by **spin-spin** interactions, have **been ruled out**. [**Neugebauer, Hennig '10-'14, Chrusciel et al '11**]

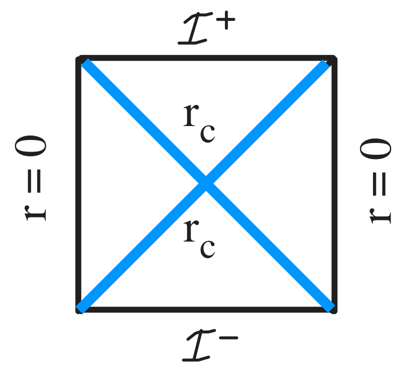
- All Einstein(-Maxwell) **binary (multi-BH)** solutions in **4-dim** found so far have **naked singularities** or **conical singularities** (e.g. **Bach-Weyl** (1922) and **Israel-Khan** (1964)), <ins>except</ins> **Majumdar-Papapetrou** solution (1947)

1. Asymp Flat higher dimensional BHs (rings, Saturns, ultraspinning lumpy BHs) [**Emparan, Reall** '02, …]
2. Exotic matter [**Volkov, Gal'tsov** '89], [**Herdeiro, Radu** '14]
3. Expanding bubbles of nothing. [**Astorino, Emparan, Viganò** '22]

→ What about de Sitter (Λ>0) ? … Uniqueness ?

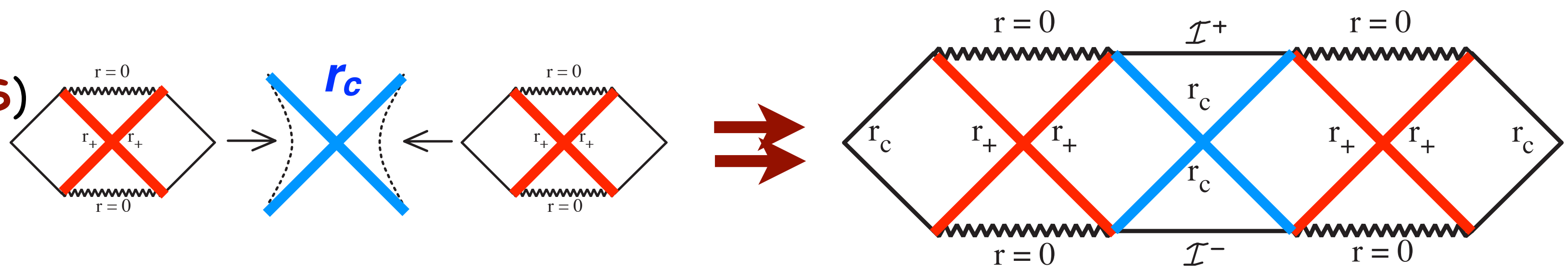
- Our Universe appears to be **expanding & accelerating** due to the presence of a **positive Λ**.
- Einstein equation with a **positive cosmological constant:** $R_{ab} = \frac{3}{\ell^2} g_{ab}$ $\quad \Lambda \equiv 3/\ell^2 > 0$
- We would like to understand the moduli space of **static/stationary BHs** of this theory.
- **For Λ>0 <u>uniqueness</u> of Kerr-dS is <u>not</u> established** ($\exists$ theorems but with assumptions…)
- Spacetimes with a **positive cosmological constant** have **spatial slices that grow exponentially**.
  => at late times, an **inertial observer O** in de Sitter experiences a **cosmological horizon**.
- Region **visible** to O — the **de Sitter static patch** — can be described by a **static** metric:

$$ds^2 = -f\,dt^2 + \frac{dr^2}{f} + r^2(d\theta^2 + \sin^2\theta\, d\phi^2) \qquad f = 1 - \frac{r^2}{\ell^2}$$



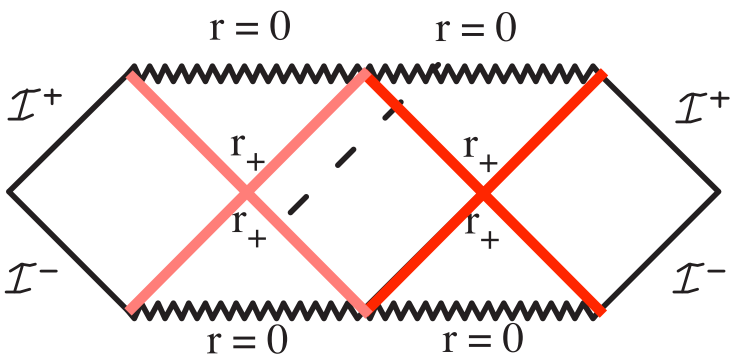
where polar coords are built **around an inertial observer O** placed at r = 0.
Null hypersurface r = r_c = ℓ, is a **cosmological** horizon:
a surface beyond which nothing influences O

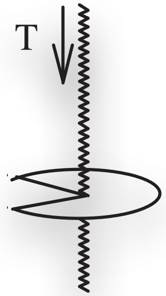
- **Kottler BH (Schw-dS)**
- **Kerr-dS BH**



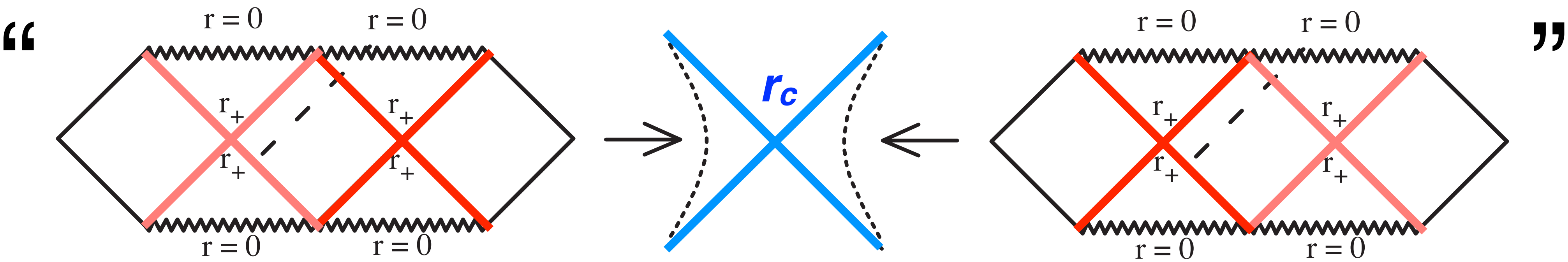
→ Can we have multi-BH, eg BH binaries ?

- **Λ=0** **Bach-Weyl** (1922) or **Israel-Khan** (1964) solution. But it has **conical singularities:**





- We want a **BH binary** **without** **conical singularities: maybe possible with Λ>0 ?**



→ Start with **Newtonian** analysis: consider a configuration of **N small BHs in de Sitter** space

- **Newton-Hooke** equations of motion: $$m_a \frac{d^2\mathbf{x}_a}{dt^2} - m_a \frac{\mathbf{x}_a}{\ell^2} = -\sum_{b\neq a}^{b=N} \frac{m_a m_b(\mathbf{x}_a - \mathbf{x}_b)}{|\mathbf{x}_a - \mathbf{x}_b|^3}$$

- **Static** solutions exist when: $$\frac{\mathbf{x}_a}{\ell^2} = \sum_{b\neq a}^{b=N} \frac{m_b(\mathbf{x}_a - \mathbf{x}_b)}{|\mathbf{x}_a - \mathbf{x}_b|^3} \quad (1) \qquad \Lambda \equiv 3/\ell^2 > 0.$$

- **Two equal mass BHs** aligned along z axis and separated by a **distance d:**

$$N = 2, \quad x_1 = -x_2 = \tfrac{d}{2}\,\hat{e}_z, \quad m_a = m_b = M$$

- Then (1) yields: $$\frac{d^3}{\ell^3} = \frac{r_+}{\ell} \quad \Rightarrow \quad \frac{d}{\ell} = \frac{1}{(4\pi\ell\,T_+)^{1/3}} \quad (2) \qquad r_+ = 2M, \quad T_+ = (4\pi r_+)^{-1}$$

- Require **validity** of **Newton + Hooke** approxs (**BHs inside** a **single cosmological** horizon):

$$r_+ \ll d\,, \quad d \ll \ell \quad (\Rightarrow r_+ \ll \ell\,)$$

- These conditions are consistent with Newton-Hooke equilibrium condition (2):

$$\frac{r_+}{\ell} = \frac{d^3}{\ell^3} \ll 1 \Rightarrow \begin{cases} d \ll \ell \\ \frac{r_+}{d} = \frac{d^2}{\ell^2} \ll 1 \end{cases}$$

=> ∃ **static de Sitter binaries** with **small BHs** are **consistent** with **Newton-Hooke theory.**

→ Going beyond the Newton-Hooke approximation: General Relativity (GR) solution

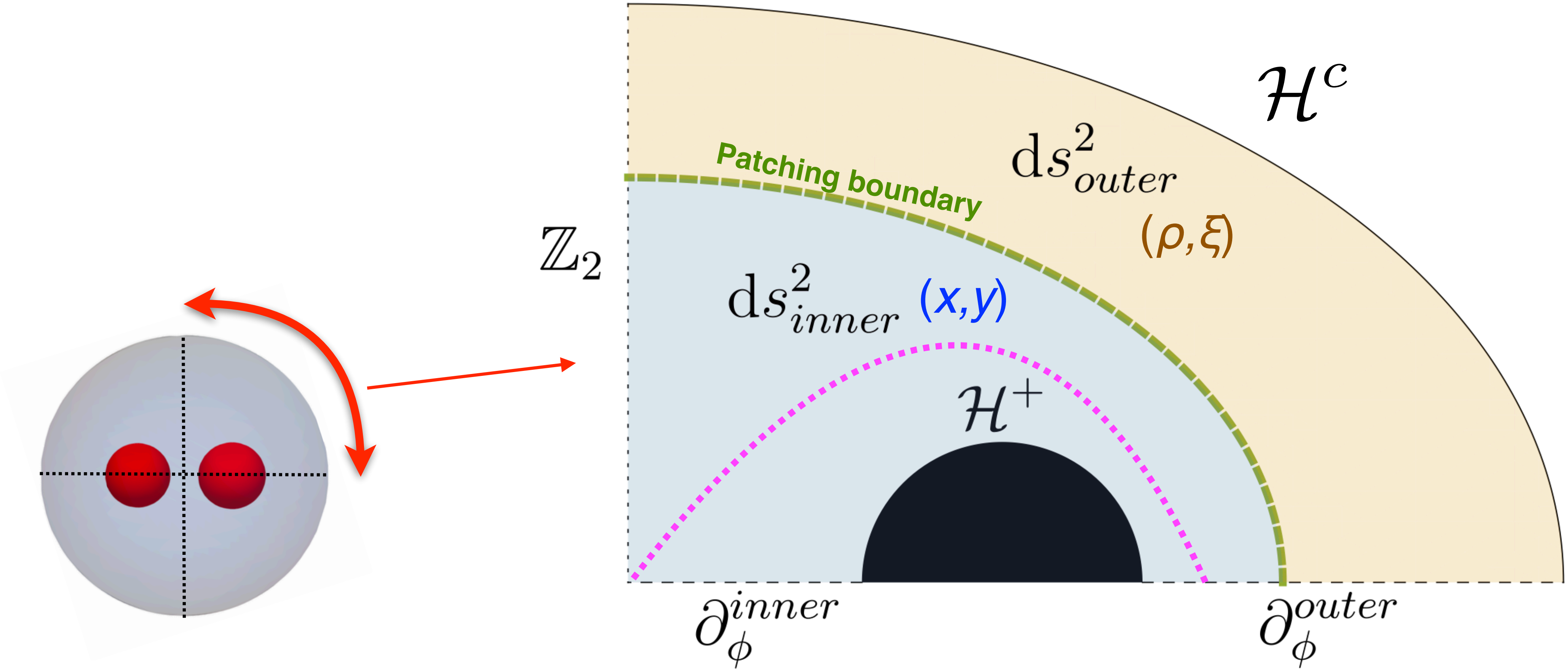
We find exact solutions to this 2-body problem in GR with $\Lambda > 0$ using numerics.

Use **Einstein-deTurck formulation of GR:**

[Headrick, Kitchen,Wiseman '09] [Review: OD, Santos,Way '15]

- Solve instead $G^{\mathrm{H}}_{ab} \equiv R_{ab} - \nabla_{(a}\xi_{b)} = \dfrac{3}{\ell^2} g_{ab}$ $\qquad \Lambda \equiv 3/\ell^2 > 0$.
- **De Turck vector** ξ can be arbitrary. We choose: $\xi^a \equiv g^{bc}\left[\Gamma^a_{bc}(g) - \Gamma^a_{bc}(\bar{g})\right]$
- $\bar{g}$ is a **reference metric** of choice: it must have the <u>same</u> **asymptotics** & **causal structure** as g.
- <u>**Advantage:**</u> **Principal symbol** of $G^{\mathrm{H}}_{ab} = 0$ is simply $\mathcal{P} \sim g^{ab}\partial_a\partial_b$

  PS=wave operator (restatement of well-posedness of Cauchy problem in Einstein eqn; Choquet-Bruhat)
  Imposing certain symmetries => wave operator —>Laplacian operator => Hyperbolic —> Elliptic problem
- For stationary problems, $G^H = 0$, together with appropriate **BCs**, yields a set of **Elliptic PDEs!**
- Ultimately, we want to solve $R_{ab} = \dfrac{3}{\ell^2} g_{ab}$ & thus we **want solutions** of $G^H = 0$ that have $\xi=0$.
- Find a solution, and check that $\xi \rightarrow 0$ in the **continuum** limit:
  Ellipticity (local uniqueness) **guarantees** that solutions w/ $\xi \neq 0$ will <u>**not**</u> be **nearby** those w/ $\xi=0$.



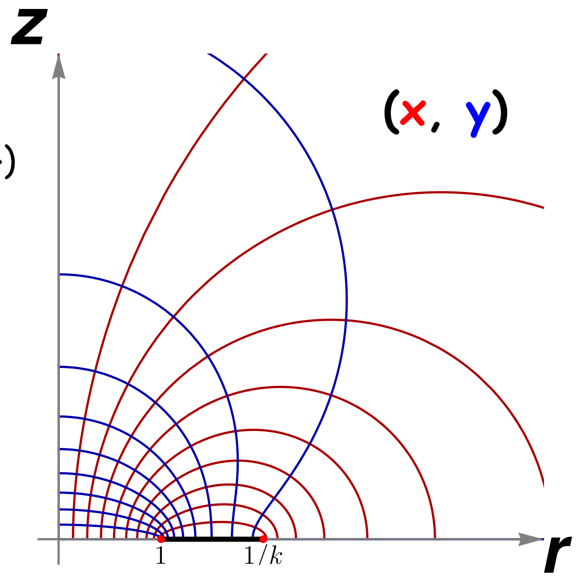
- **Outer region:** near (single) cosmological horizon, solution looks like **de Sitter** space:

  Polar Weyl coordinates **(ρ,ξ) coords**

$$z = \rho\,\xi\sqrt{2-\xi^2}$$
$$r = \rho(1-\xi^2)$$

- **Inner region:** solution looks like **warped Israel-Khan** (cylindrical Weyl {r,z})

  but <u>without</u> conical singularity: ring-like **(x,y)** coords



- Inner region is **pentagonal** (5 boundaries)

  => so split it into 2 squared (4 boundaries) sub-regions

→ Choosing a good reference metric $\bar{g}$ & metric ansatz with patching

(de Turck reference metric should have same causal structure & symmetries as solution we look for)

1) de Turck reference metric:

Israel-Khan without conical singularity; **(x,y)**

$$
\begin{aligned}
\mathrm{d}s^2_{\mathrm{ref}} &= \frac{\ell^2}{g_+^2}\left\{-f g_-^2\, F\, \mathrm{d}t^2 + \frac{\lambda^2}{m^2\Delta_{xy}^2}\left[p^2\left(\frac{4\mathrm{d}x^2}{(2-x^2)\Delta_x} + \frac{4\mathrm{d}y^2}{(2-y^2)\Delta_y}\right) + y^2(2-y^2)(1-y^2)^2\, s\, \mathrm{d}\phi^2\right]\right\} \\
&= \frac{\ell^2}{g_+^2}\left\{-f g_-^2\, F\, \mathrm{d}t^2 + \frac{\lambda^2 h}{f}\left[\mathrm{d}\rho^2 + \rho^2\left(\frac{4\mathrm{d}\xi^2}{2-\xi^2} + \frac{(1-\xi^2)^2}{h}\, s\, \mathrm{d}\phi^2\right)\right]\right\}.
\end{aligned}
$$

de Sitter space: **(ρ,ξ) coords**

$$s = 1 - \alpha(1-y^2)^2$$

$$\alpha = \cdots$$

2) metric ansatz with patching:

$$
\begin{aligned}
\mathrm{d}s^2 &= \frac{\ell^2}{g_+^2}\Bigg\{ - f g_-^2\, F\, \mathcal{T}\, \mathrm{d}t^2 + \frac{\lambda^2}{m^2\Delta_{xy}^2}\Bigg[w^2\left(\frac{4\mathcal{A}\,\mathrm{d}x^2}{(2-x^2)\Delta_x} + \frac{4\mathcal{B}}{(2-y^2)\Delta_y}\Big(\mathrm{d}y - x\,(1-x^2)\,y\,(2-y^2)(1-y^2)\mathcal{F}\,\mathrm{d}x\Big)^2\right) \\
&\qquad + y^2(2-y^2)(1-y^2)^2\, s\, \mathcal{S}\, \mathrm{d}\phi^2\Bigg]\Bigg\} \\
&= \frac{\ell^2}{g_+^2}\left\{ - f g_-^2\, F\, \widetilde{\mathcal{T}}\, \mathrm{d}t^2 + \frac{\lambda^2 h}{f}\left[\widetilde{\mathcal{A}}\,\mathrm{d}\rho^2 + \rho^2\left(\frac{4\widetilde{\mathcal{B}}}{2-\xi^2}\Big(\mathrm{d}\xi - \xi\,(2-\xi^2)(1-\xi^2)\,\rho\,\widetilde{\mathcal{F}}\,\mathrm{d}\rho\Big)^2 + \frac{(1-\xi^2)^2}{h}\, s\, \widetilde{\mathcal{S}}\,\mathrm{d}\phi^2\right)\right]\right\}
\end{aligned}
$$

<u>**Our mission**</u>: find the **unknown functions** $\{\mathcal{T},\mathcal{A},\mathcal{B},\mathcal{F},\mathcal{S}\}_{(x,y)}$ $\{\widetilde{\mathcal{T}},\widetilde{\mathcal{A}},\widetilde{\mathcal{B}},\widetilde{\mathcal{F}},\widetilde{\mathcal{S}}\}_{(\rho,\xi)}$

by **solving** the **Einstein-de Turck EoM** **(ξ=0)**
subject to the appropriate physical **Boundary Conditions**

We know the map:
ρ(x,y), ξ(x,y)

$$\rho = \frac{\sqrt{y^2(2-y^2)+x^2(2-x^2)(1-y^2)^2}}{\sqrt{(1-y^2)^2+k^2x^2(2-x^2)y^2(2-y^2)}}\,,$$

$$\xi = \sqrt{1 - \frac{(1-x^2)y\sqrt{2-y^2}(1-}{\sqrt{y^2(2-y^2)+x^2(2-x^2)(1-y^2)^2}}}$$

- <u>Our mission</u>: find the **unknown functions** $\{\mathcal{T},\mathcal{A},\mathcal{B},\mathcal{F},\mathcal{S}\}_{(x,y)}$ $\{\widetilde{\mathcal{T}},\widetilde{\mathcal{A}},\widetilde{\mathcal{B}},\widetilde{\mathcal{F}},\widetilde{\mathcal{S}}\}_{(\rho,\xi)}$

  We know the map: $\rho(x,y)$, $\xi(x,y)$

  by **solving** the **Einstein-de Turck EoM** (ξ=0)

  subject to the appropriate (regularity) physical **Boundary Conditions**

- **Numerical method:**

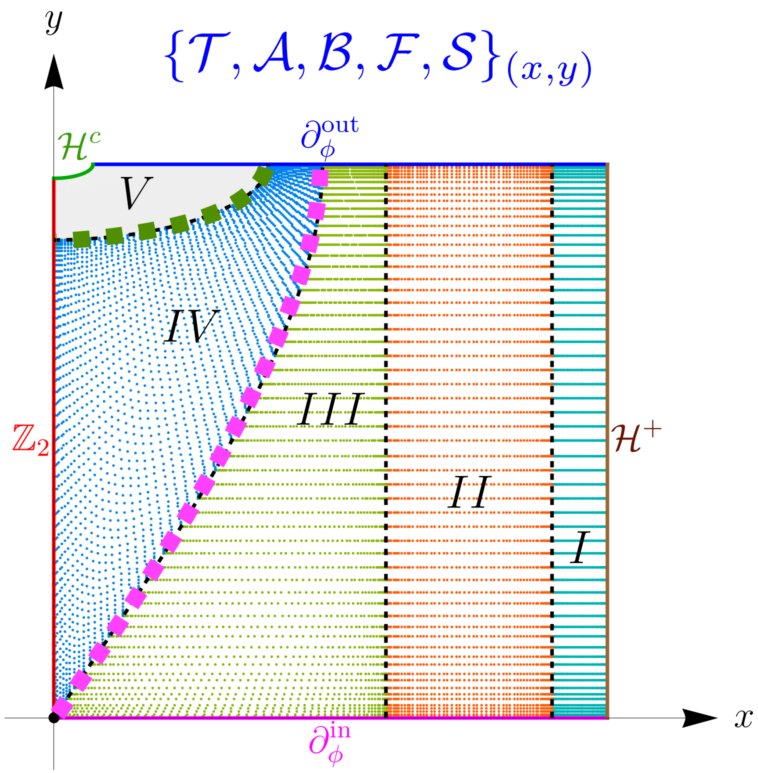
[Review: OD, Santos,Way 1510.02804]

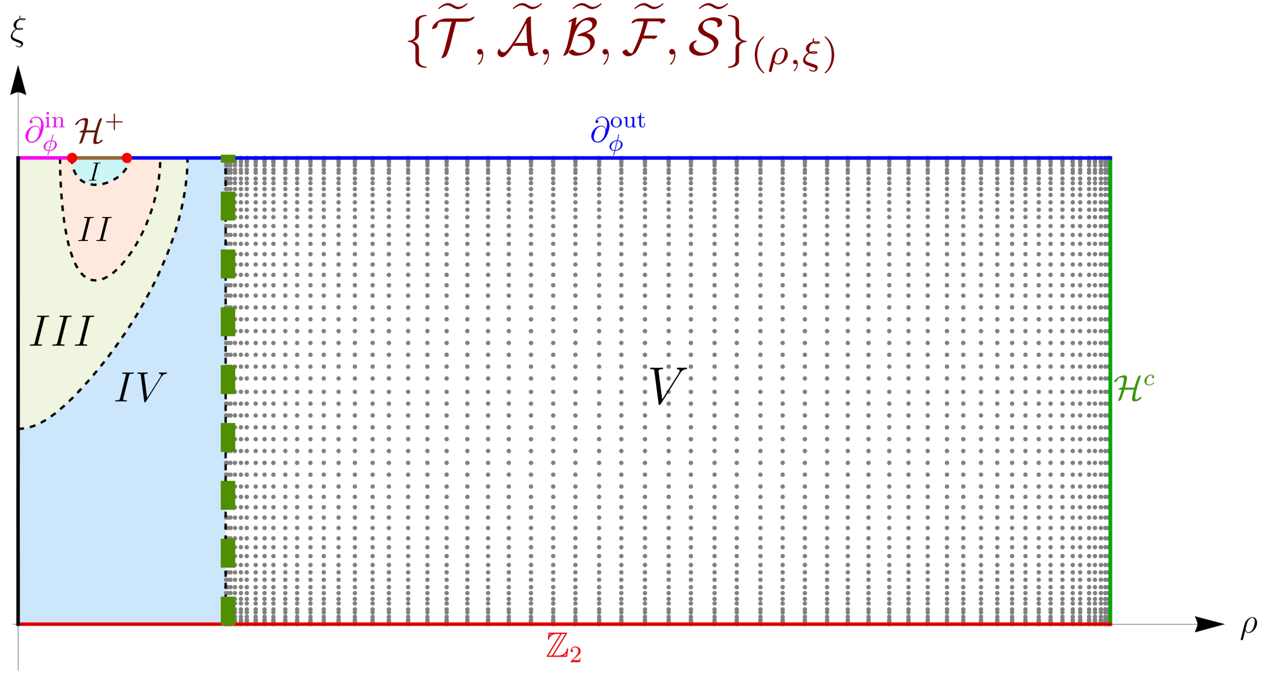
Use a **Newton-Raphson algorithm** with **pseudospectral grid**.

Also use **transfinite interpolation** to complete the **patching**.

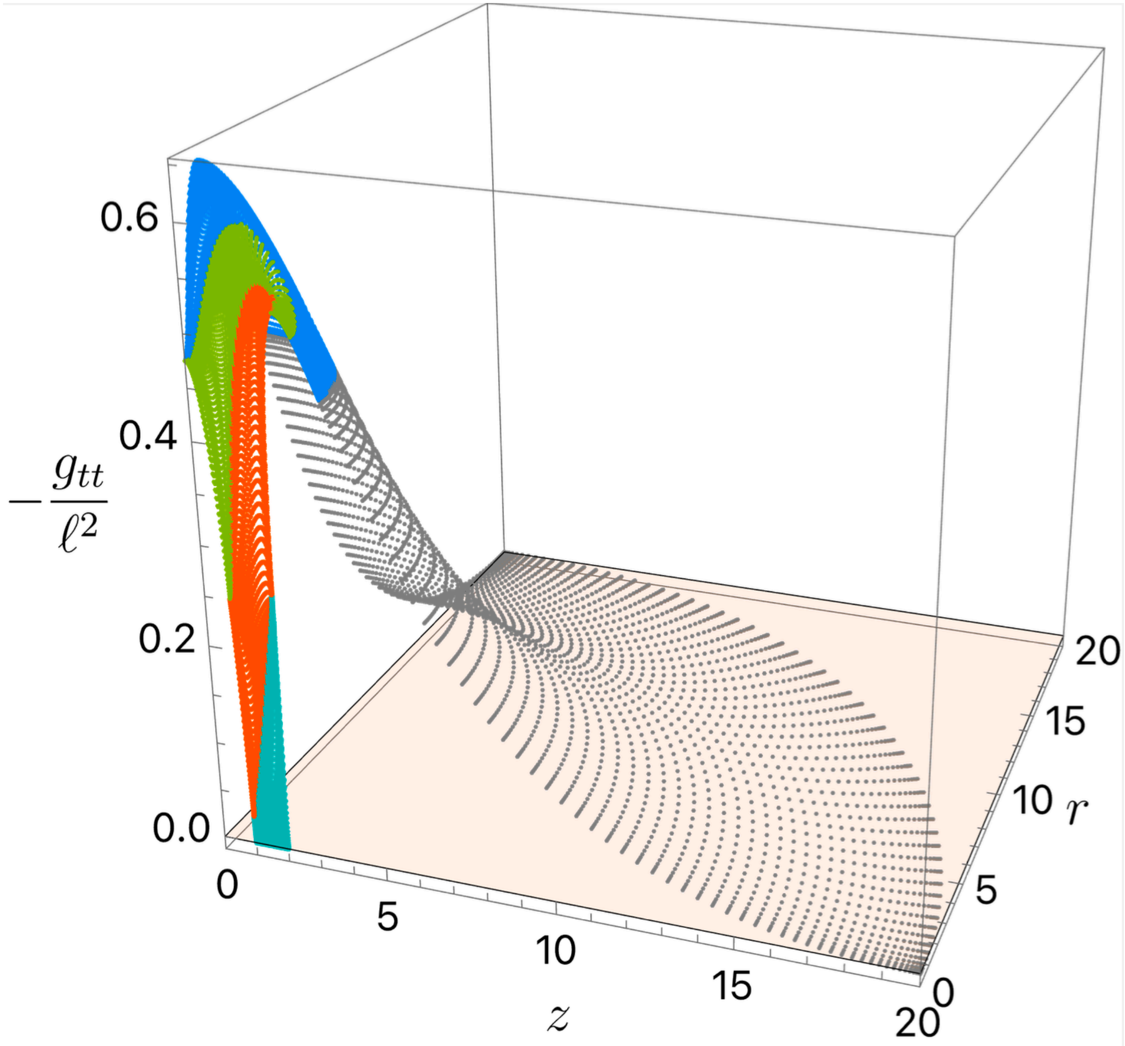
@ patching boundary, require:

1) matching of **two line elements**, & **2) matching** of the **normal derivative** across patch bdry



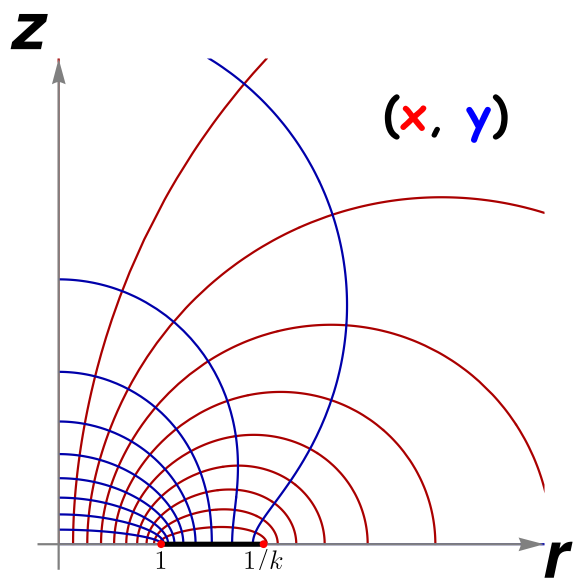


Testing patching: g_{tt} & $g_{\phi\phi}$ are gauge invariant since ∂_t and ∂_ϕ are KVFs



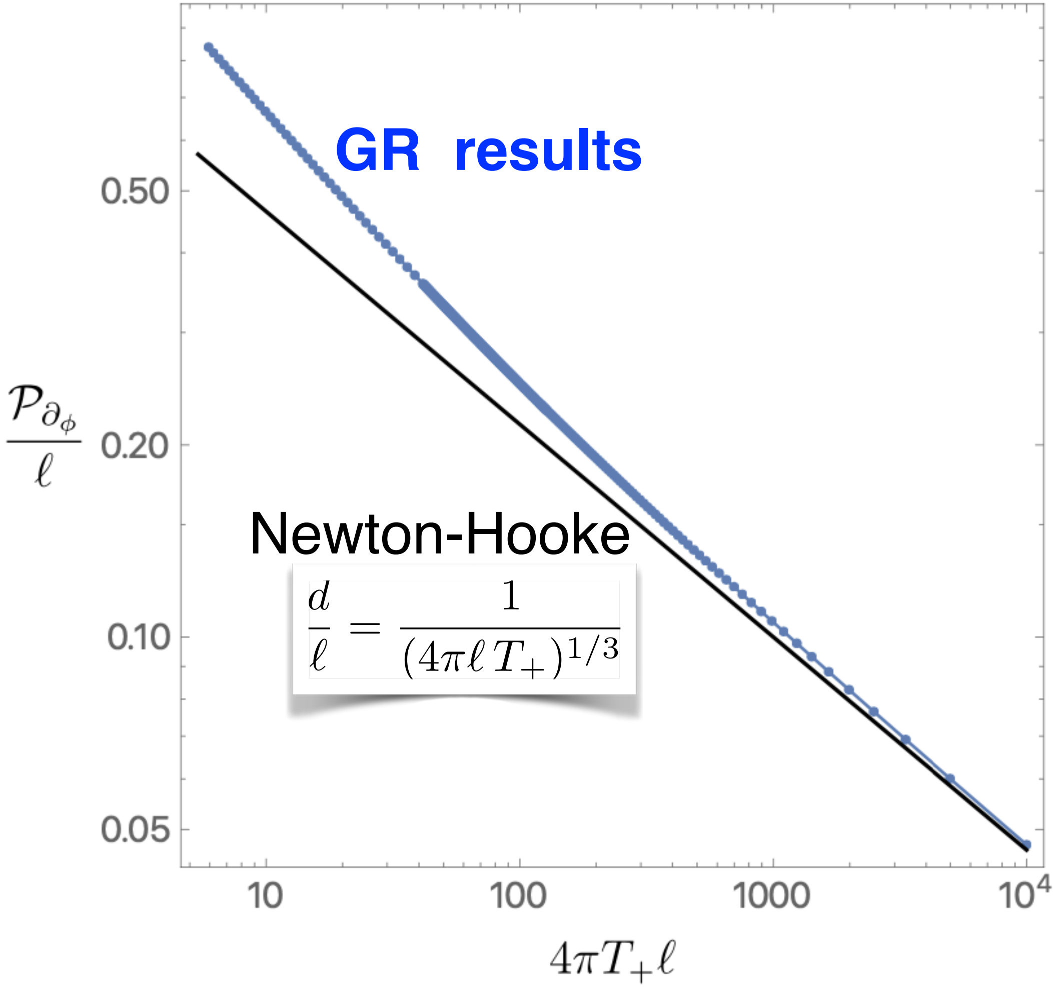
Recall **Bach-Weyl (Israel-Khan) cylindrical-Weyl coord {r,z}** and its **rod-structure** where:

1) the rotation axis and the BH horizons are all located at r = 0
2) there is a Z_2 symmetry



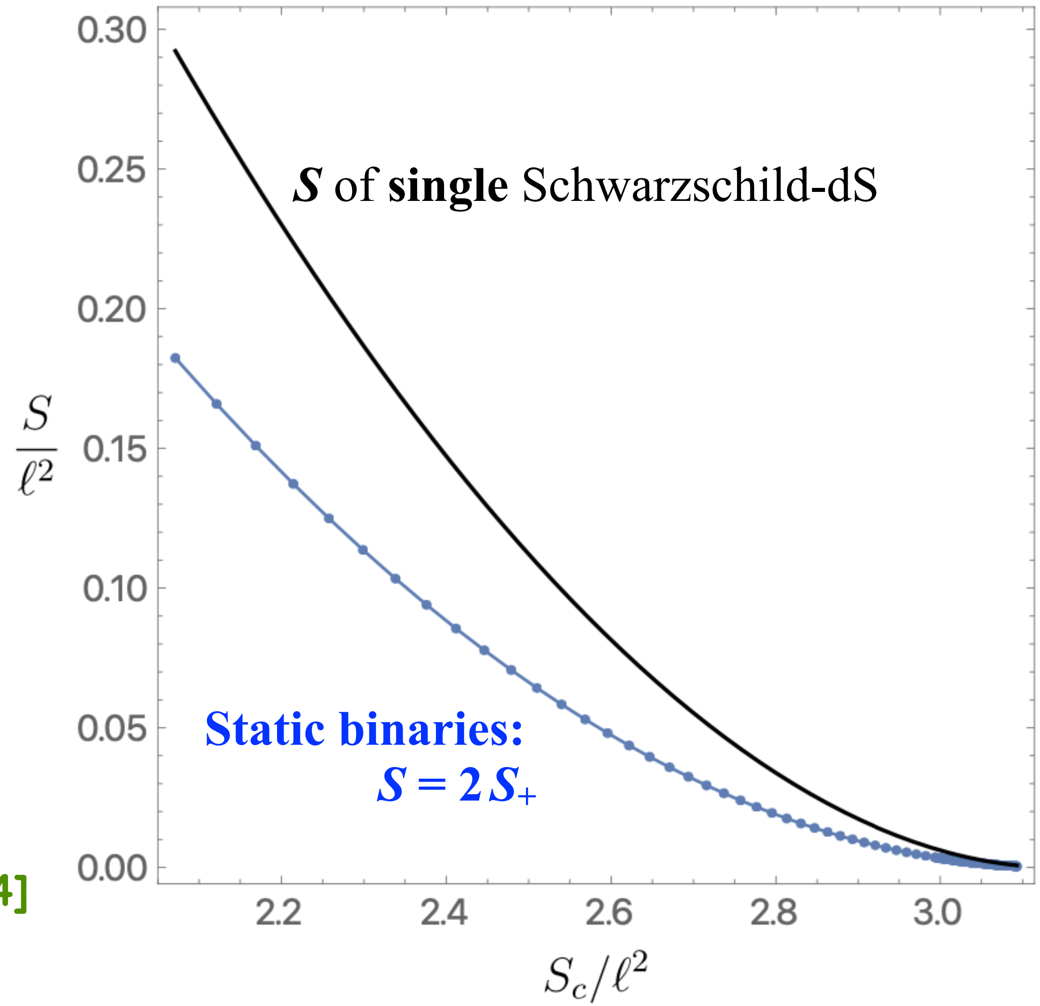
→ Properties of static de Sitter BH binaries

Proper distance between the **BH horizons** versus the BH **temperature**:



NON-Uniqueness in 4D!

Total BH entropy versus the **cosmological horizon entropy**:



First law of thermodynamics:

$$-T_c\,\mathrm{d}S_c = 2\,T_+\,\mathrm{d}S_+$$

[Hawking,Gibbons '74]

Our data satisfies it up to 0.01%

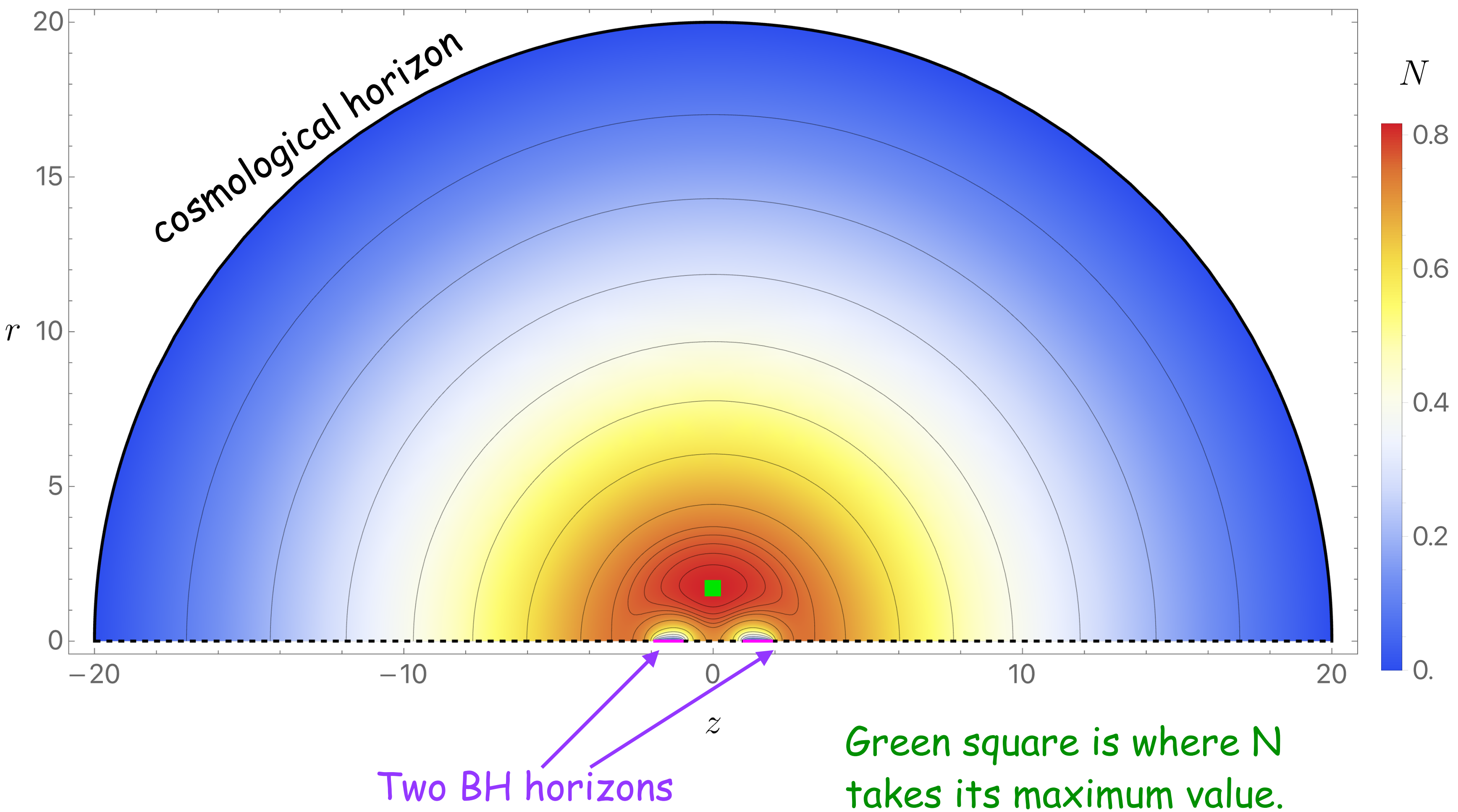
→ Properties of static de Sitter BH binaries

- **Not in conflict** with available Uniqueness theorems:

  we have (explicitly identified) **assumptions** of these theorems that can be **evaded**

[LeFloch, Rozoy '10] [Borghini, Chruściel, Mazzieri '19]
[ul Alam, Yu '14]

Contour plot showing the level sets of the lapse function $N = \sqrt{-g_{tt}}$



Spinning Black hole binaries

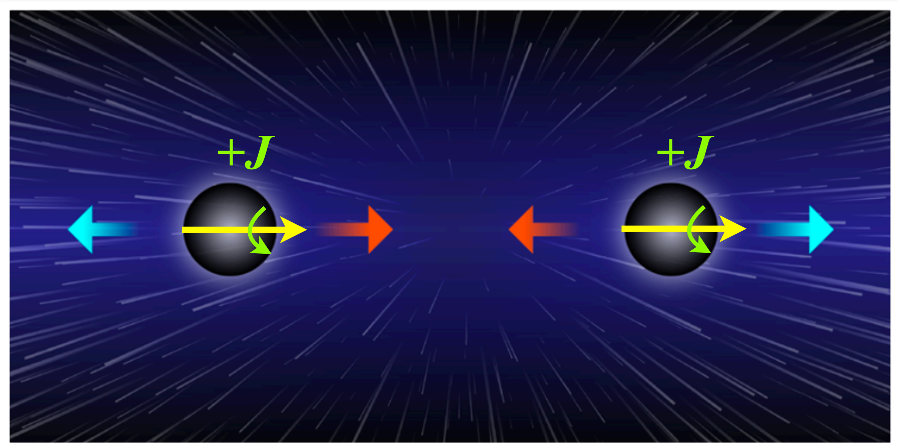
- Add **spin** (not orbital angular momentum) to BHs:

  Cosmological expansion + Grav attraction + Spin-Spin interaction [Wald '72]

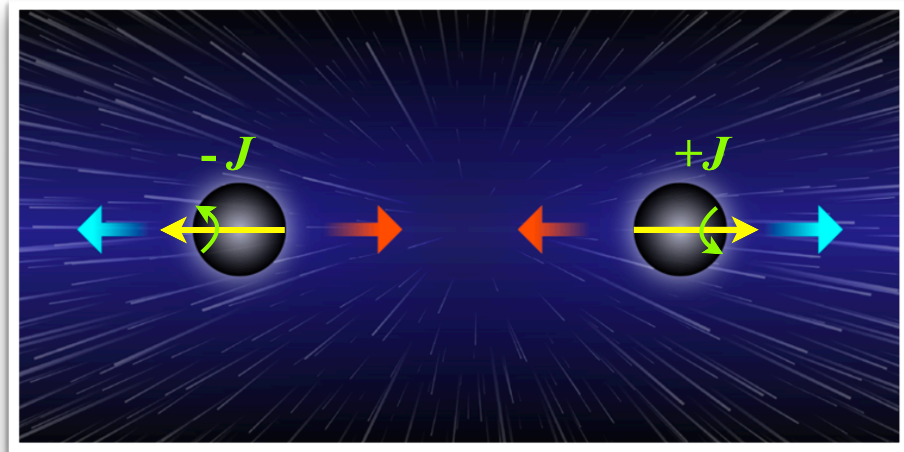
- Can we have stationary (no quadrupole momentum, no radiation) **spinning** BH binaries ?

- Can Spin-Spin interaction **stabilise** the binaries (alike in molecular systems) ?

Aligned [Neuman BC W(-x)=+W(x)]
Repulsive Spin-Spin interaction



Anti-aligned [Dirichlet BC W(-x)=-W(x)]
Attractive Spin-Spin interaction



OD, Jorge Santos, Benson Way, 2406.10333

→ Spinning binaries within Newton-Hooke + Spin-Spin interaction

- **Newton-Hooke** + **Spin-Spin interaction** [Wald '72] equations of motion:

$$m_a \frac{\mathrm{d}^2 \mathbf{x}_a}{\mathrm{d}t^2} = m_a \frac{\mathbf{x}_a}{\ell^2} + \nabla\left(\frac{m_a\, m_b}{|\mathbf{r}_{ab}|}\right) + \nabla\left[\frac{\mathcal{S}_a \cdot \mathcal{S}_b}{|\mathbf{r}_{ab}|^3} - \frac{3\,(\mathcal{S}_a \cdot \mathbf{r}_{ab})\,(\mathcal{S}_b \cdot \mathbf{r}_{ab})}{|\mathbf{r}_{ab}|^5}\right]$$

$$\Lambda \equiv 3/\ell^2 > 0.$$

- **Stationary** solutions exist when: $\frac{\mathrm{d}^2 \mathbf{x}_a}{\mathrm{d}t^2} = 0$ (1)

- **Two equal mass BHs** aligned along z axis and separated by a **distance d**:

$$N = 2, \quad x_1 = -x_2 = \frac{d}{2}\,\hat{e}_z, \quad m_a = m_b = M$$

$$\mathcal{S}_{1,2} = m\,\sigma_{1,2}\,\mathbf{e}_z \qquad \sigma_2 = \gamma\sigma_1 \equiv \gamma\sigma$$

$\gamma = -1$ (atractive SS)

$\gamma = +1$ (repulsive SS)

- Then (1) yields:

$$\frac{d^3}{\ell^3} = \frac{2m}{\ell}\left(1 - 6\,\gamma\,\frac{\sigma^2}{d^2}\right) \quad \Rightarrow \quad \frac{d}{\ell} \simeq \frac{1}{(4\pi T_+ \ell)^{\frac{1}{3}}}\left\{1 - \left[\frac{1}{3} + \frac{2\gamma}{(4\pi T_+ \ell)^{4/3}}\right]\frac{\Omega_+^2}{(4\pi T_+)^2}\right\} \quad (2)$$

$$2\,m = \frac{1}{2\pi T_+ + \sqrt{4\pi^2 T_+^2 + \Omega_+^2}}$$

$$\sigma = \frac{m\,\Omega_+}{\sqrt{4\pi^2 T_+^2 + \Omega_+^2}}\,.$$

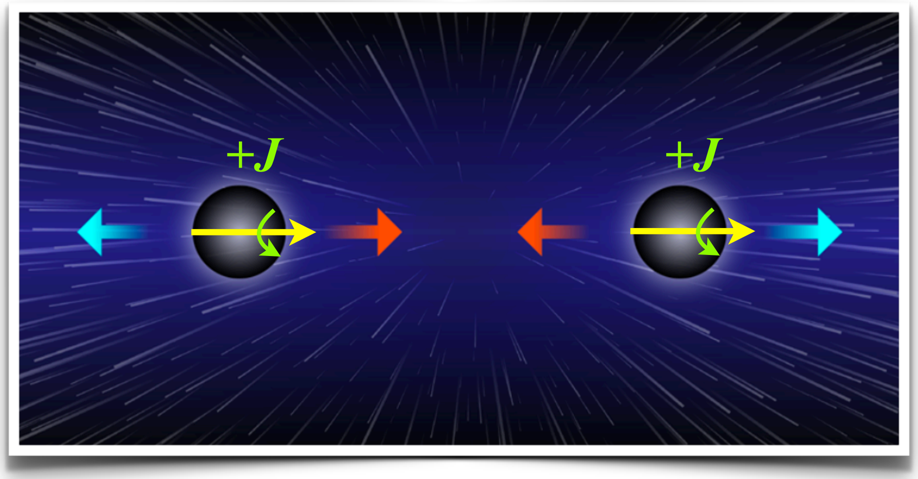
- **Equilibrium condition (2)** can fall within the regime of **validity of Newton-Hook theory:**

$$r_+ \ll d \ll \ell \ \text{(i.e. large } T_+ \ell)$$

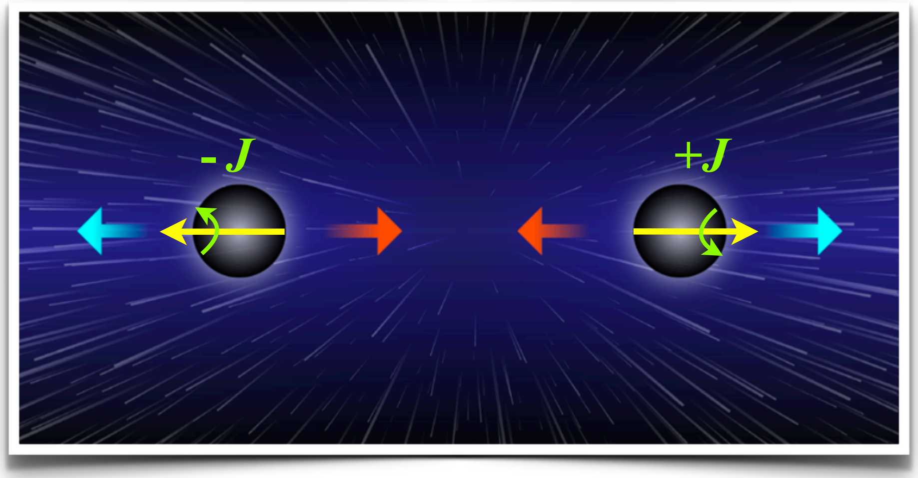
→ Choosing a good reference metric $\bar{g}$ & metric ansatz with BCs

$$
\begin{aligned}
\mathrm{d}s^2 = \frac{\ell^2}{g_+^2}\Bigg\{ &- f g_-^2\, F\, \mathcal{T}\, \mathrm{d}t^2 + \frac{\lambda^2}{m^2 \Delta_{xy}^2}\Bigg[w^2 \left(\frac{4\mathcal{A}\, \mathrm{d}x^2}{(2-x^2)\Delta_x} + \frac{4\mathcal{B}}{(2-y^2)\Delta_y}\left(\mathrm{d}y - x\,(1-x^2)\,y\,(2-y^2)(1-y^2)\mathcal{F}\,\mathrm{d}x\right)^2 \right) \\
&+ y^2(2-y^2)(1-y^2)^2\, s\, \mathcal{S}\, \left(\mathrm{d}\phi - g_-^2\, w\, \mathcal{W}\, \mathrm{d}t\right)^2 \Bigg]\Bigg\} \\
= \frac{\ell^2}{g_+^2}\Bigg\{ &- f g_-^2\, F\, \widetilde{\mathcal{T}}\, \mathrm{d}t^2 + \frac{\lambda^2 h}{f}\Bigg[\widetilde{\mathcal{A}}\, \mathrm{d}\rho^2 + \rho^2 \left(\frac{4\widetilde{\mathcal{B}}}{2-\xi^2}\left(\mathrm{d}\xi - \xi\,(2-\xi^2)(1-\xi^2)\,\rho\, \widetilde{\mathcal{F}}\,\mathrm{d}\rho\right)^2 + \frac{(1-\xi^2)^2}{h}\, s\, \widetilde{\mathcal{S}}\left(\mathrm{d}\phi - g_-^2\, w\, \widetilde{\mathcal{W}}\, \mathrm{d}t\right)^2 \right)\Bigg]\Bigg\}
\end{aligned}
$$

Aligned [Neuman BC W(-x)=+W(x)]
Repulsive Spin-Spin interaction



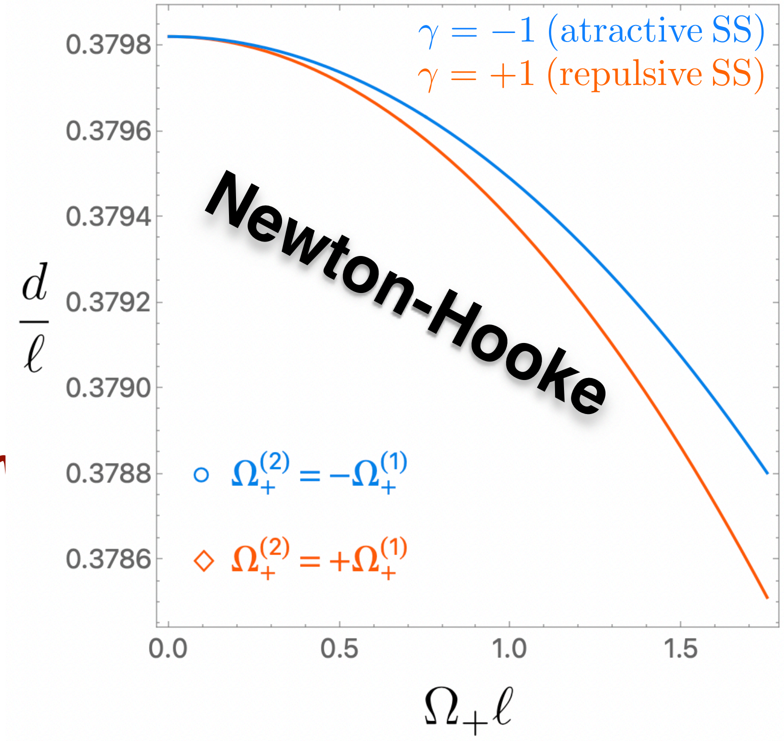
Anti-aligned [Dirichlet BC W(-x)=-W(x)]
Attractive Spin-Spin interaction

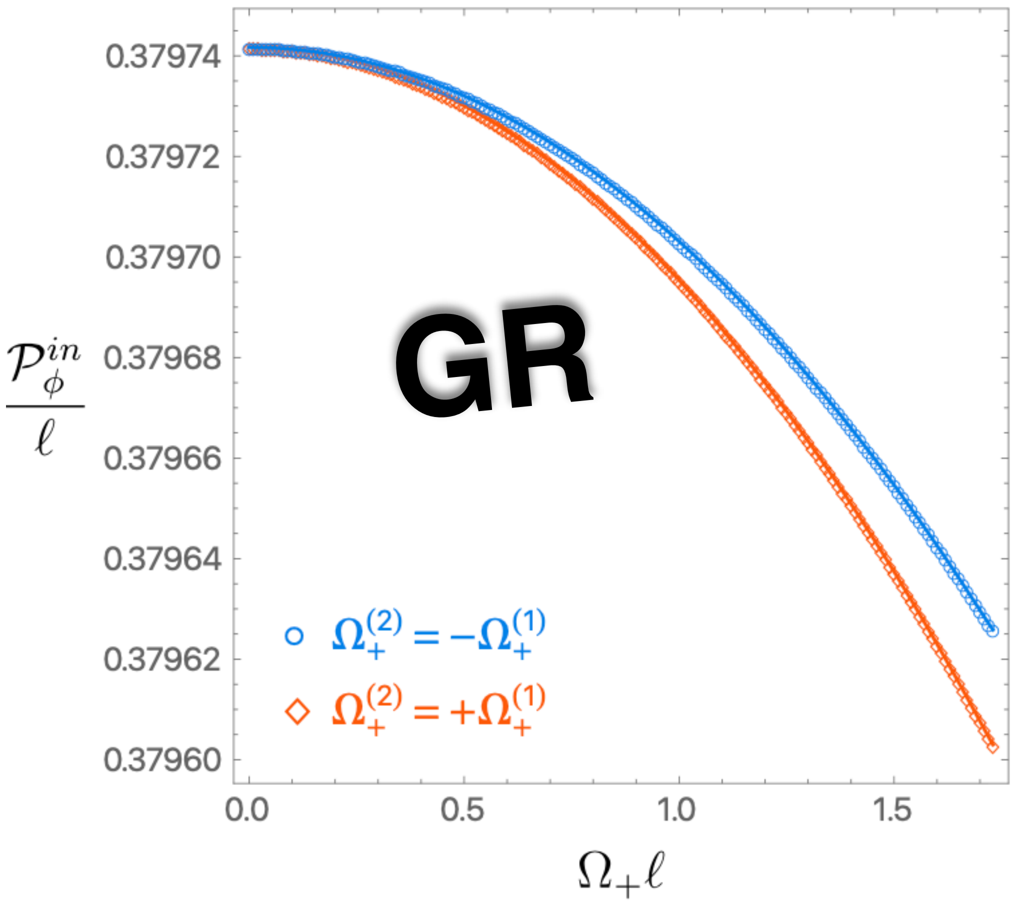


→ Properties of SPINNING de Sitter BH binaries

- $d(\Omega_+)$ curve for the aligned binary (γ = 1) is always **below** the curve for the anti-aligned case (γ = −1).

- **Spin-spin forces are repulsive for aligned spins**
  => BHs need to be closer apart to remain in equilibrium
  (for fixed gravitational and cosmological forces).
  The opposite is true for anti-aligned spins.

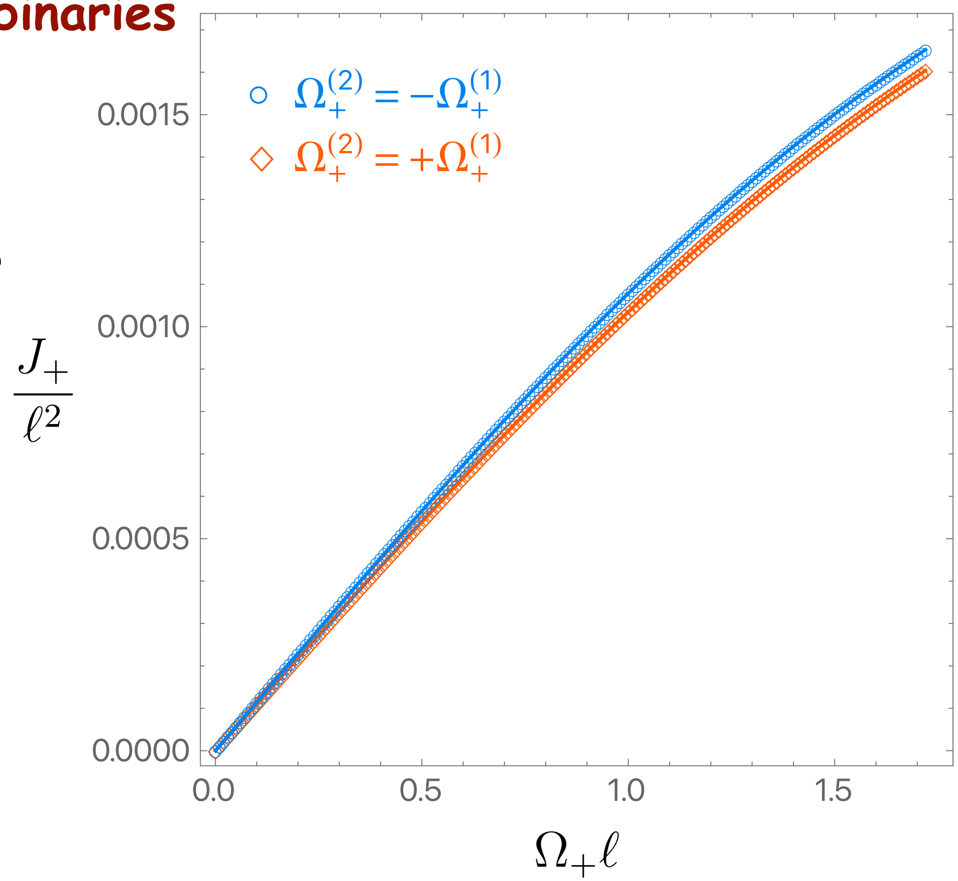


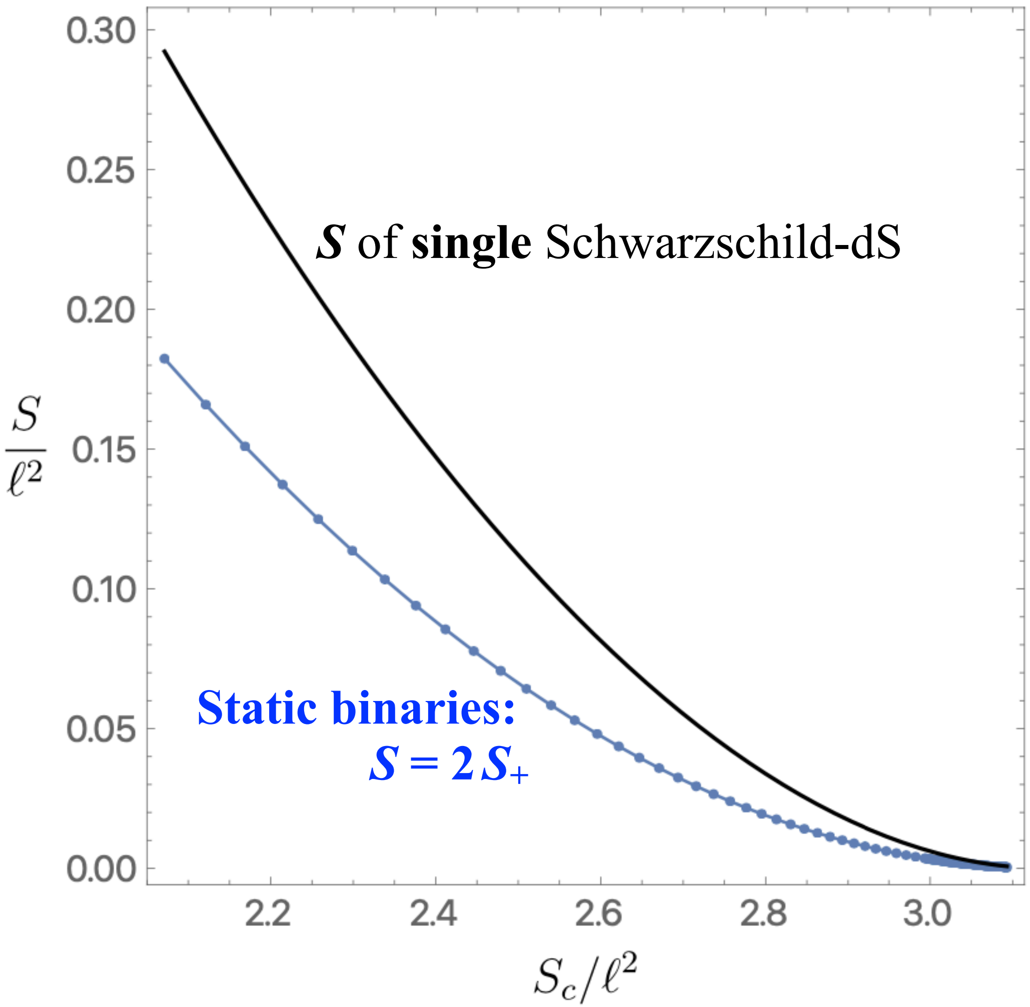


→ Properties of SPINNING de Sitter BH binaries

- Binaries are thermodynamically unstable
  For a given S_c and J_{tot}, spinning binaries have lower total event horizon S than dS Schw/Kerr

- **Continuous Non-Uniqueness**





First law of thermodynamics:

$$-T_c\,\mathrm{d}S_c = 2\,T_+\,\mathrm{d}S_+ + 2\,\Omega_+\,\mathrm{d}J_+$$

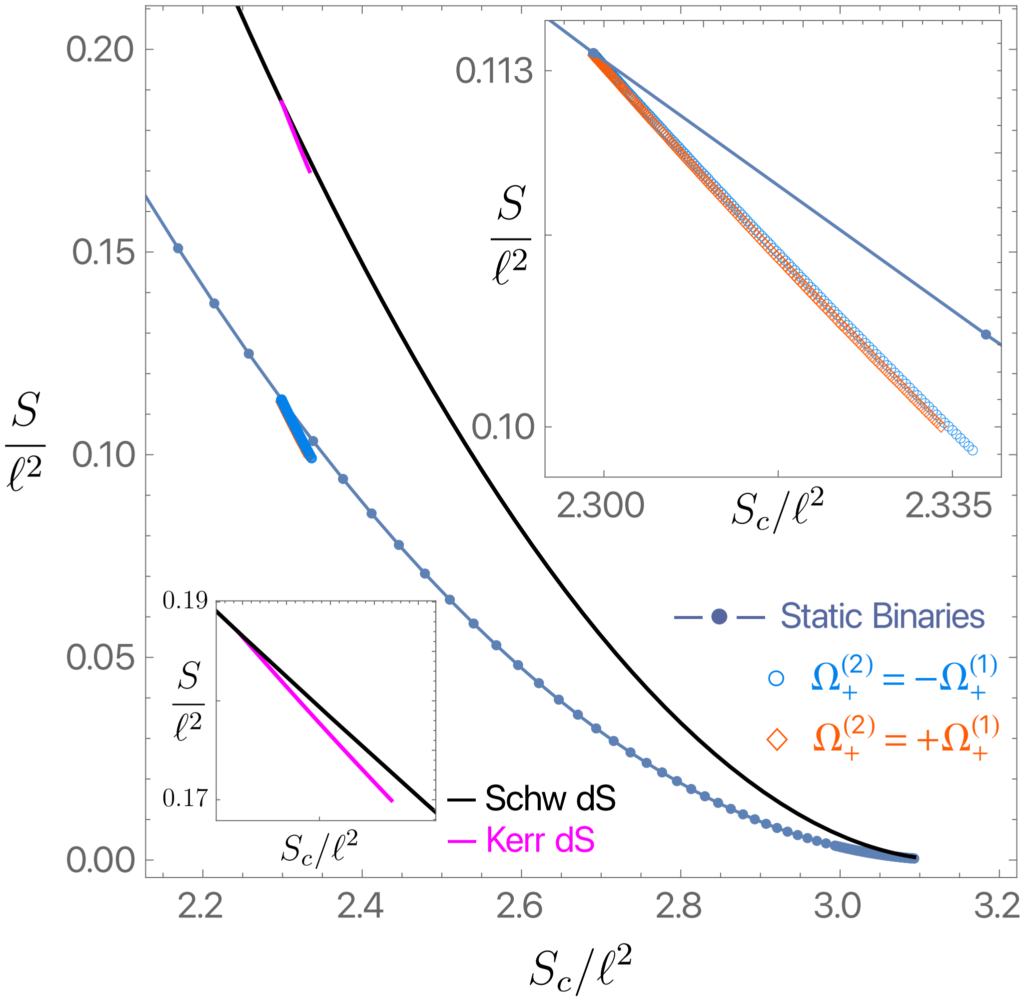
[Hawking,Gibbons '74]

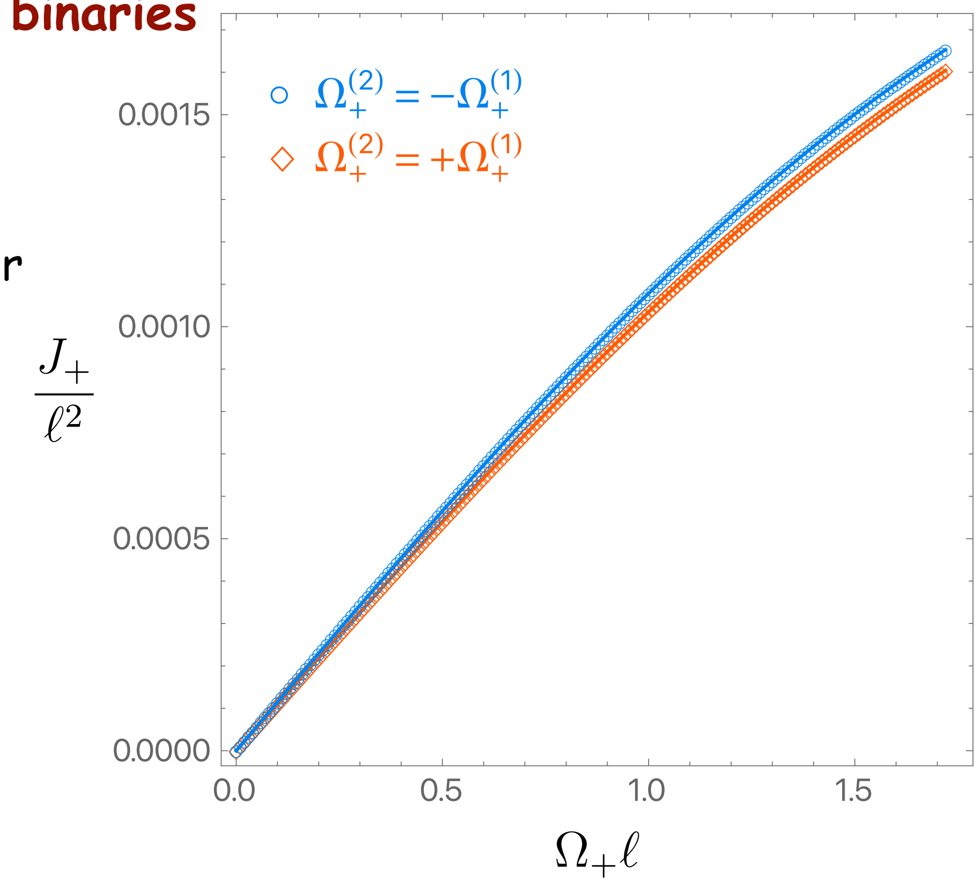
Our data satisfies it up to 0.01%

→ Properties of SPINNING de Sitter BH binaries

- Binaries are thermodynamically unstable
  For a given S_c and J_{tot}, spinning binaries have lower total event horizon S than dS Schw/Kerr

- **Continuous Non-Uniqueness**





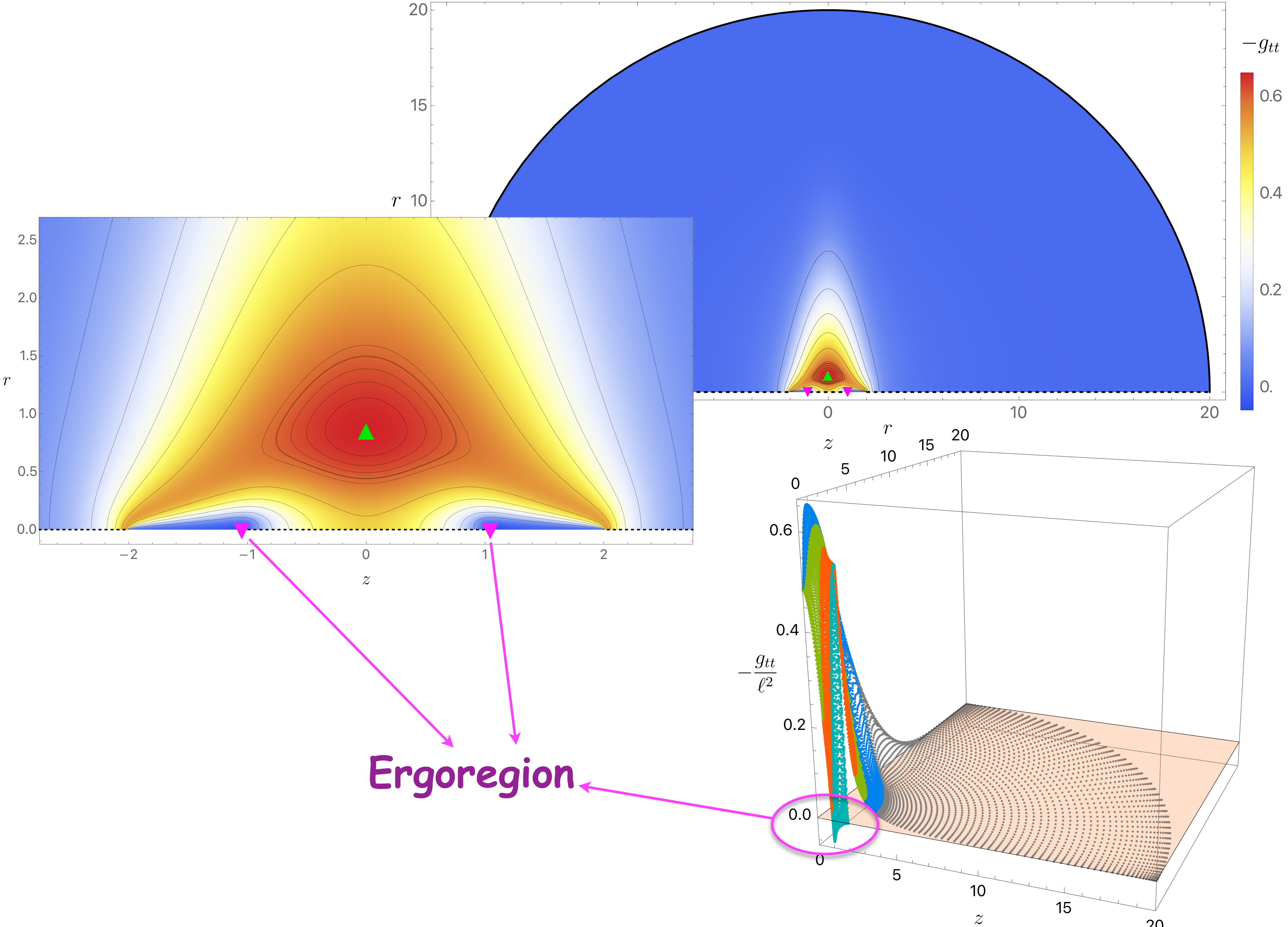
First law of thermodynamics:

$$-T_c\,\mathrm{d}S_c = 2\,T_+\,\mathrm{d}S_+ + 2\,\Omega_+\,\mathrm{d}J_+$$

[Hawking, Gibbons '74]

Our data satisfies it up to 0.01%

→ Properties of SPINNING de Sitter BH binaries



→ Outlook

- Our binaries are **thermodynamically unstable**: but, under small perturbations, the BH pair necessarily needs to merge into a single BH or fly apart? Or can it be **dynamically stable** ?

- **Future**: study the **dynamical stability** of spinning binaries by perturbing our stationary solutions.

- The **spin-spin interactions** act on shorter length scales, & **might provide a mechanism for stabilizing binaries** in some windows of parameters, **alike it stabilizes molecules**

→ Back-of-the-envelop analysis within Newton-Hooke approximation:

- **Static binary:**

V(r) = −1/r −r²/2 ←—Schrödinger potential

=> single **maximum** (r=1)

=> **unstable equilibrium**

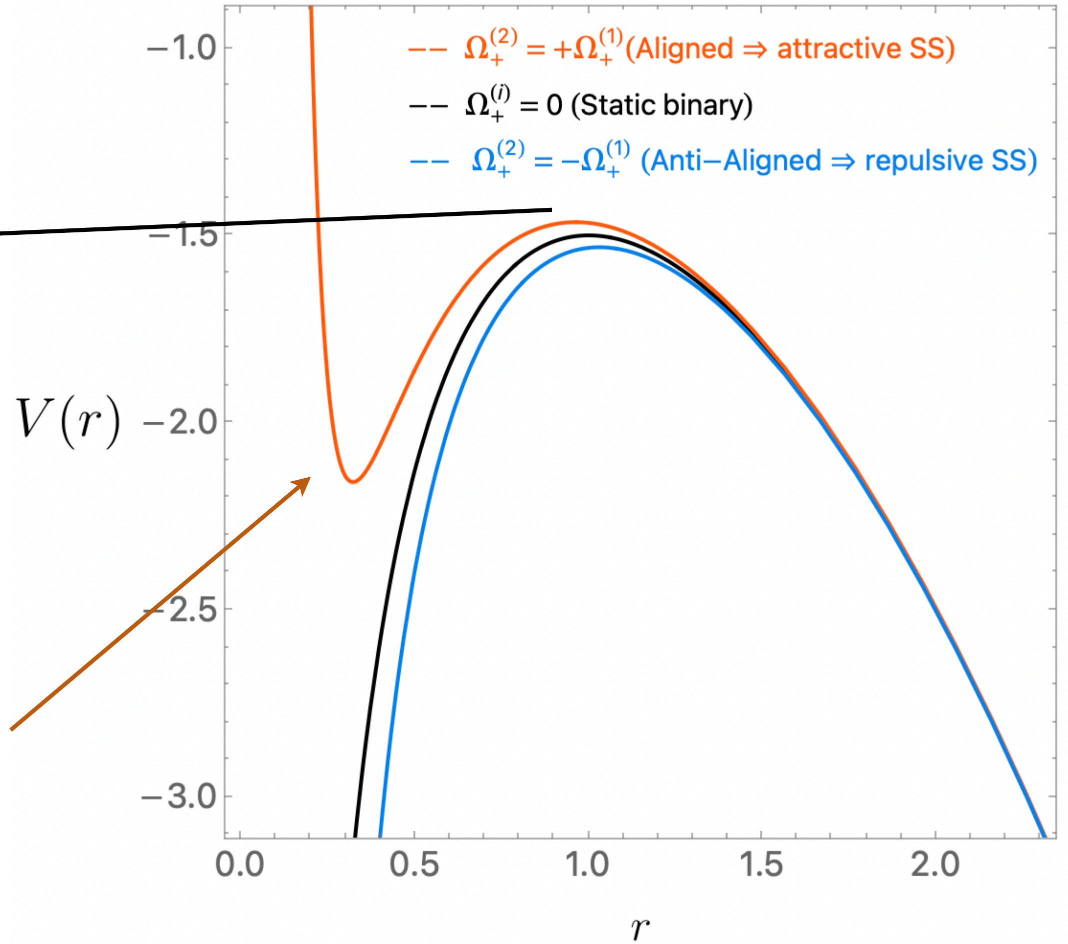
- **Spinning binary:** Add a $1/r^3$ spin-spin term,

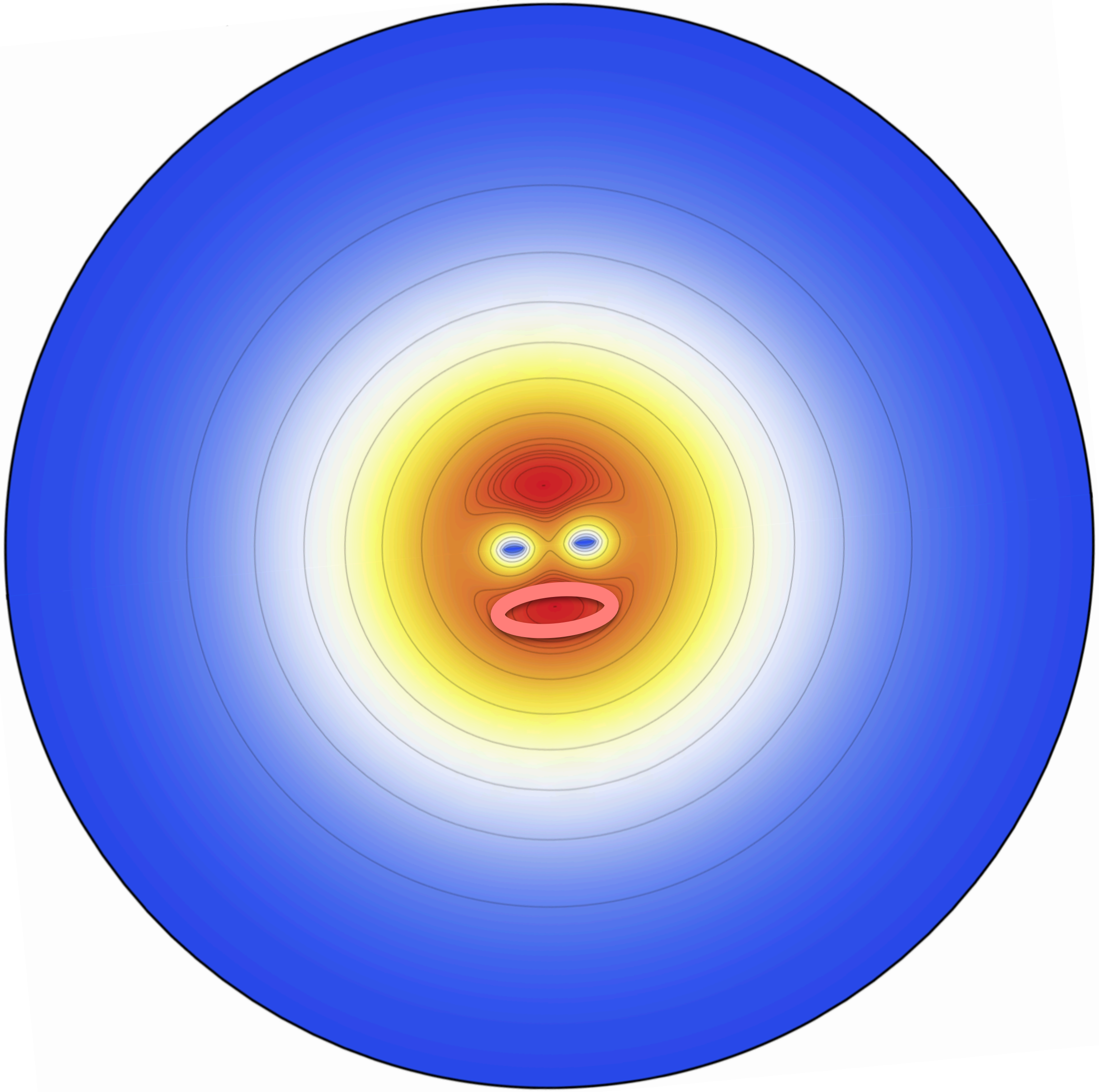
$V_{spin}(r)=V(r)-\gamma/r^3$

For $\gamma>1/30$ (spin-spin repulsive interaction)

$V_{spin}(r)$ has **local minimum**

=> a **STABLE equilibrium point**





Merci!

→ Choosing a good reference metric $\bar{g}$

A) near the event horizons

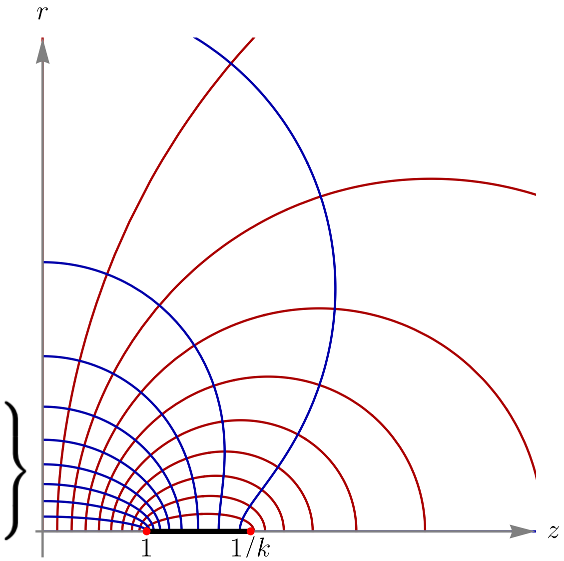
- For binaries well within the **cosmological horizon**, ie **near the event horizons**, the solution should be well approximated by a **Bach-Weyl 1922 (Israel-Khan 1964)** but **without conical singularity:**

$$ds^2 = \ell^2 \left[-f dt^2 + \frac{\lambda^2}{f} [h(dr^2 + dz^2) + \mathbf{s}\, r^2 d\phi^2] \right]$$

- Introduce a **Schwarz-Christoffel map** from **cylindrical-Weyl** to **ring-like coordinates (x, y)**:

$$r = \frac{(1-x^2)\sqrt{1-k^2x^2(2-x^2)}\,y\sqrt{2-y^2}(1-y^2)}{(1-y^2)^2 + k^2x^2(2-x^2)y^2(2-y^2)}$$

$$z = \frac{x\sqrt{2-x^2}\sqrt{(1-y^2)^2 + k^2y^2(2-y^2)}}{(1-y^2)^2 + k^2x^2(2-x^2)y^2(2-y^2)}$$



Lines of constant x & y, along with the **rod structure** of Bach-Weyl

$$ds^2 = \ell^2 \left\{ -f dt^2 + \frac{\lambda^2}{m^2 \Delta_{xy}^2} \left[p^2 \left(\frac{4dx^2}{(2-x^2)\Delta_x} + \frac{4dy^2}{(2-y^2)\Delta_y} \right) + \mathbf{s}\, y^2(2-y^2)(1-y^2)^2 d\phi^2 \right] \right\}$$

$$\mathbf{s} = 1 - \boldsymbol{\alpha}(1-y^2)^2$$

- Wish to **join the Bach-Weyl** solution **with** a **de Sitter horizon.** In anticipation, we **write the Bach-Weyl** solution in **polar-Weyl coordinates (ρ, ξ)**

$$z = \rho\,\xi\sqrt{2-\xi^2}$$
$$r = \rho(1-\xi^2)$$

$$ds^2 = \ell^2 \left\{ -f dt^2 + \frac{\lambda^2 h}{f} \left[d\rho^2 + \rho^2 \left(\frac{4d\xi^2}{2-\xi^2} + s\,\frac{(1-\xi^2)^2}{h} d\phi^2 \right) \right] \right\}$$

B) near the cosmological horizon

- **Closer** to the **cosmological horizon**, we would like the metric to **look like pure de Sitter:**

$$\mathrm{d}s^2 = -\left(1-\frac{R^2}{\ell^2}\right)\mathrm{d}\tau^2 + \frac{\mathrm{d}R^2}{1-\frac{R^2}{\ell^2}} + r^2(\mathrm{d}\theta^2 + \sin^2\theta \mathrm{d}\phi^2)$$

- Introduce **isotropic coordinates:** $\frac{R}{\ell} = \frac{\lambda\,\rho}{1+\frac{\lambda^2\rho^2}{4}}\,, \qquad \sin\theta = 1-\xi^2\,, \qquad \tau = \ell\, t$

$$\mathrm{d}s^2 = \frac{\ell^2}{g_+^2}\left\{-g_-^2\mathrm{d}t^2 + \lambda^2\left[\mathrm{d}\rho^2 + \rho^2\left(\frac{4\mathrm{d}\xi^2}{2-\xi^2} + (1-\xi^2)^2\mathrm{d}\phi^2\right)\right]\right\} \qquad g_\pm = 1 \pm \frac{\lambda^2\rho^2}{4}$$

- In these coords, the **de Sitter horizon** is at **ρ = 2/λ** (where $g_-^2 = 0$)
  & has a constant temperature of **Tc = 1/(2π).**

- **de Sitter space** in isotropic coords **resembles** **Bach-Weyl solution** in polar-Weyl coords:

$$\mathrm{d}s^2 = \ell^2\left\{-f\mathrm{d}t^2 + \frac{\lambda^2 h}{f}\left[\mathrm{d}\rho^2 + \rho^2\left(\frac{4\mathrm{d}\xi^2}{2-\xi^2} + \boldsymbol{s}\,\frac{(1-\xi^2)^2}{h}\mathrm{d}\phi^2\right)\right]\right\} \qquad f,h\big|_{\rho\gg 1} \to 1$$