Photons as tracers of the curvature of space-time and the mass distribution in the Universe

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A questioned cosmology: ΛCDM

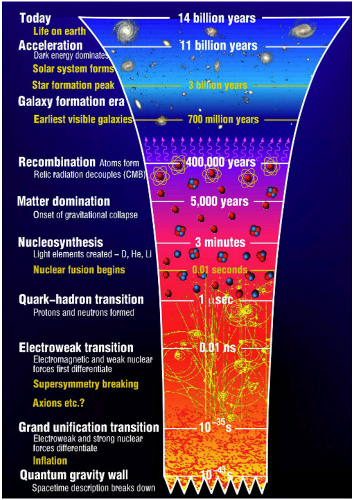


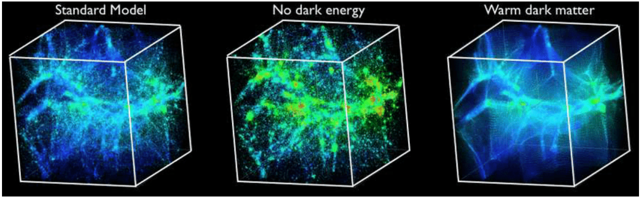
Figure 1: Illustration of the ΛCDM model.

Param.	Desc.
$\Omega_b h^2$	Baryon density
$\Omega_c h^2$	Dark matter density
H_0	Hubble constant
τ	Optical depth (reionization)
n_s	Scalar spectral index
A_s	Power spectrum amplitude

Table 1: The six base parameters of the ΛCDM model

A questioned cosmology: ΛCDM

Probe	**Parameters**	**Measurements**
CMB	H_0, Ω_m, n_s	Temperature anisotropies
SN Ia	H_0	Luminosity distances
BAO	H_0, Ω_m	Galaxy correlation scale
Galaxy clusters	Ω_m, σ_8	Mass, spatial distribution



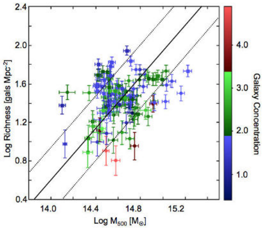
(a) Matter distribution changes with different cosmological models
Credit: K. Heitmann

(b) Abell 370
NASA/ESA

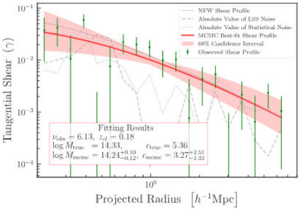
Inferring cosmology from the Universe's mass distribution

One quantity, many ways to measure it

- Mass-richness relations
- X-ray luminosity
- Number counts
- Sunyaev-Zel'dovich signal
- Weak lensing
- In simulations: Friends-of-Friends, M_{200}



(a) Example of a mass-richness relation for galaxy clusters. (Rettura 2017)



(b) Example of a shear profile around a galaxy cluster. (Chen et al. 2020)

Inferring cosmology from the Universe's mass distribution

Credit: Rubin Observatory/NSF/J. Pinto

(a) The principle of gravitational lensing.

Credit: NASA/ESA

(b) Weak and strong lensing in Abell 370.

What does lensing probe?

Lensing is a direct probe of the mass profiles and mass distribution of galaxies and galaxy clusters.

Weak lensing: a brief overview

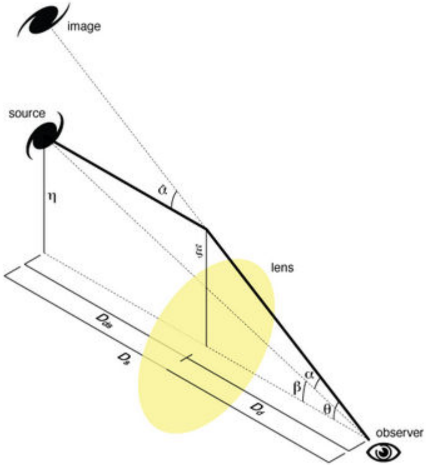


Figure 5: Definition of lensing angles and distances.
Figure by Michael SACHS

Lens equation

$$\vec{\beta} = \vec{\theta} - \frac{D_{ds}}{D_s}\hat{\alpha}(D_d\vec{\theta})$$

Magnification matrix

$$\frac{\partial \beta_i}{\partial \theta_j} = \delta_{ij} - \frac{\partial \alpha_i}{\partial \theta_j} = A_{ij}$$

Weak lensing hypotheses

- WL is perturbative
- Born approximation
- Thin & single lens approximations

Weak lensing: convergence, shear, and flexion

First order approximation of the magnification matrix

$$A = \begin{pmatrix} 1-\kappa & 0 \\ 0 & 1-\kappa \end{pmatrix} - \gamma \begin{pmatrix} \cos(2\phi) & \sin(2\phi) \\ \sin(2\phi) & -\cos(2\phi) \end{pmatrix}$$

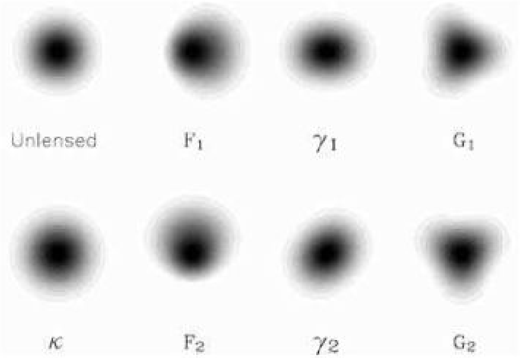


Figure 6: Effects of the different lensing fields on a Gaussian galaxy of radius 1 arcsec. 10% convergence/shear and 0.28 arcsec^{-1} flexion (which is a very high value for this quantity, chosen only to visualize) are applied. From Bacon et al. 2006

Weak lensing: going beyond the hypotheses

Weak lensing hypotheses

- WL is perturbative: all lensing effects are small
- Born approximation: small angles, evaluating deflections transversally
- Thin & single lens approximations: instant and successive deflections

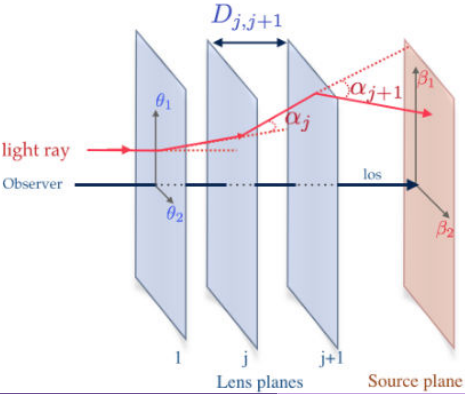


Figure 7: Raytracing in current cosmological simulations. Some hypotheses made are the thin lens and single lens approximations, along with the Born approximation. This framework of analysis can not generate strong lensing effects.
Credit: C.Gouin

From "geometric" to relativistic lightcones

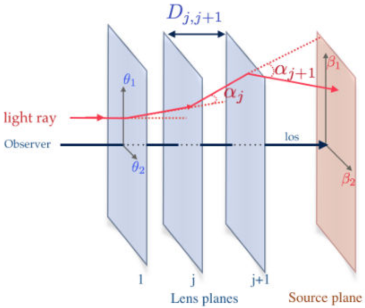
Current lightcone extraction

- Limited accuracy
- Oversimplified deflections
- Excluded strong lensing
- Non linearities?

→

Towards relativistic lightcones

- Full relativistic treatment
- Continuous deflections
- Weak and strong lensing
- GR handles non-linearities



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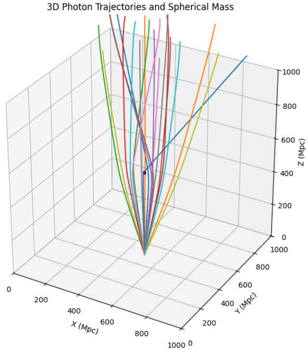


The homogeneous and isotropic Universe

The Friedmann-Lemaitre-Robertson-Walker metric

$$ds^2 = -c^2 dt^2 + a^2(t)\left[\frac{dr^2}{1-kr^2} + r^2(d\theta^2 + \sin^2\theta\, d\phi^2)\right]$$

where:

- $a(t)$: scale factor
- k: curvature parameter ($k = 0, +1, -1$)

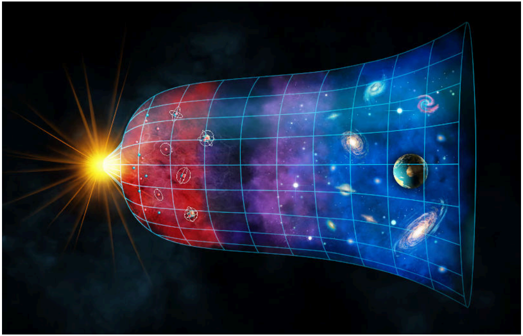


Figure 8: The story of the Universe: from the Big Bang to the present day.
Credit: Natalie Mayer

Perturbing the homogeneous and isotropic Universe

Perturbation of the FLRW metric

$$\mathrm{d}s^2 = -c^2(1 + 2\phi/c^2)\mathrm{d}t^2 - 2c \cdot a(t)B_i \mathrm{d}x^i \mathrm{d}t$$
$$- a^2(t)\left[(1 - 2\psi)\gamma_{ij} + 2E_{ij}\right]\mathrm{d}x^i \mathrm{d}x^j$$

where:

- ϕ: Newtonian potential
- B_i: vector potential
- ψ: spatial curvature perturbation
- E_{ij}: tensor perturbation
- γ_{ij}: spatial metric (flat, spherical, or hyperbolic)

From perturbations to trajectories

$$ds^2 = g_{\mu\nu} dx^\mu dx^\nu \Rightarrow \Gamma^\mu_{\alpha\beta} = \frac{1}{2} g^{\mu\nu} \left(\partial_\alpha g_{\nu\beta} + \partial_\beta g_{\nu\alpha} - \partial_\nu g_{\alpha\beta}\right)$$

The Christoffel symbols: geometry within dynamics

From this...

$$\Gamma^{\mu}_{\alpha\beta} = \frac{1}{2} g^{\mu\nu} \left(\partial_\alpha g_{\nu\beta} + \partial_\beta g_{\nu\alpha} - \partial_\nu g_{\alpha\beta} \right)$$

...

$$\Gamma^{0}_{ii} = \frac{1}{2} \Big[g^{00}(2\partial_i g_{i0} - \partial_0 g_{ii}) + g^{01}(2\partial_i g_{i1} - \partial_1 g_{ii}) + g^{02}(2\partial_i g_{i2} - \partial_2 g_{ii}) + g^{03}(2\partial_i g_{i3} - \partial_3 g_{ii}) \Big].$$

...

...to this!

$$\Gamma^{0}_{ii} = \frac{a\dot{a}}{c^2} + \frac{2a\dot{a}}{c^4}(\Phi + \Psi) - \frac{a^2}{c^4}\frac{\partial \Psi}{\partial t}$$

The Christoffel symbols: geometry within dynamics

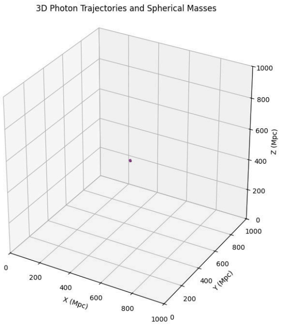
Γ^0_{00}	$\frac{1}{c^2}\frac{\partial\Phi}{\partial t}$
Γ^i_{ii}	$-\frac{1}{c^2}\frac{\partial\Psi}{\partial x^i}$
Γ^0_{ii}	$\frac{a\dot{a}}{c^2}+\frac{2a\dot{a}}{c^4}(\Phi+\Psi)-\frac{a^2}{c^4}\frac{\partial\Psi}{\partial t}$
Γ^i_{00}	$\frac{1}{a^2}\frac{\partial\Phi}{\partial x^i}$
$\Gamma^0_{0i}=\Gamma^0_{i0}$	$\frac{1}{c^2}\frac{\partial\Phi}{\partial x^i}$
$\Gamma^i_{i0}=\Gamma^i_{0i}$	$\frac{\dot{a}}{a}-\frac{1}{c^2}\frac{\partial\Psi}{\partial t}$
Γ^i_{jj}	$\frac{1}{c^2}\frac{\partial\Psi}{\partial x^i}$
$\Gamma^i_{ij}=\Gamma^i_{ji}$	$-\frac{1}{c^2}\frac{\partial\Psi}{\partial x^j}$

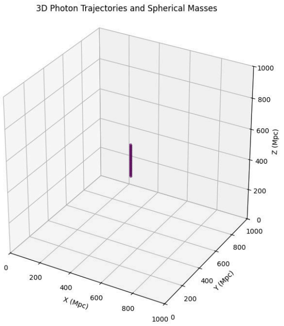
Table 2: The perturbed Christoffel symbols in a flat spacetime, with vector and tensor perturbations set to zero. The terms $\dot{a}$ and a are the time derivative of the scale factor and the scale factor respectively.

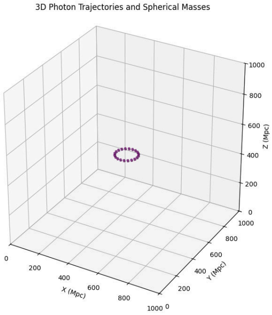
The numerical tool EXCALIBUR

Geodesic equation in General Relativity: photon trajectory

$$\frac{d^2x^\mu}{d\lambda^2} + \Gamma^\mu_{\alpha\beta}\frac{dx^\alpha}{d\lambda}\frac{dx^\beta}{d\lambda} = 0$$







EXCALIBUR: **EX**act **CA**lculation of **LI**ght **B**ending **U**sing **R**elativity

A Python prototype for relativistic raytracing in a cosmological context.

Geodesic equation in General Relativity: photon trajectory

$$\frac{d^2x^\mu}{d\lambda^2} + \Gamma^\mu_{\alpha\beta}\frac{dx^\alpha}{d\lambda}\frac{dx^\beta}{d\lambda} = 0$$

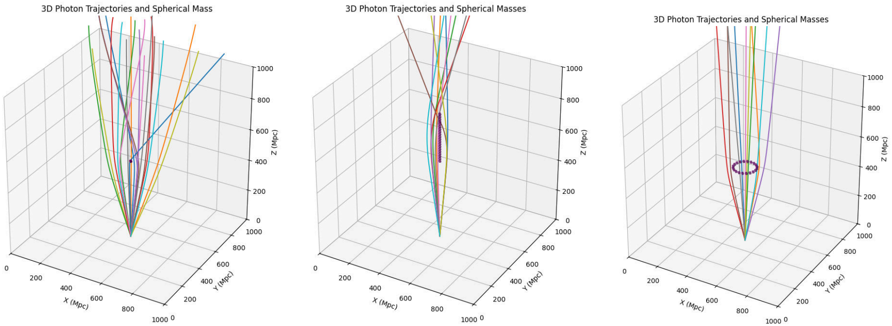


Figure 10: Simulations of geodesics deflected by masses in the Universe. Figures from the LAPP internship project. The sum of masses is always $10^{20} M_\odot$. From left to right: a single mass, a line of masses, and a ring-shaped mass distribution.

Relativistic light propagation and gravitational lensing

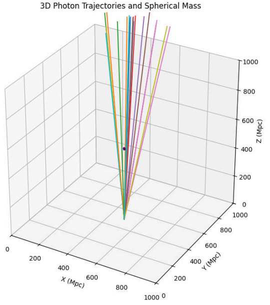


Figure 11: Simulation of geodesics deflected by a realistic mass ($10^{15} M_\odot$) in the Universe. Figure from the LAPP internship project.

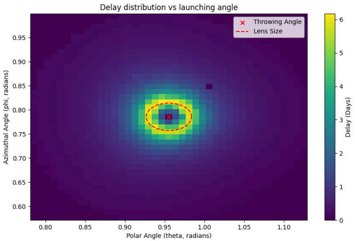
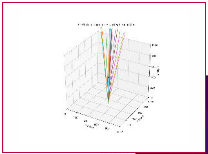
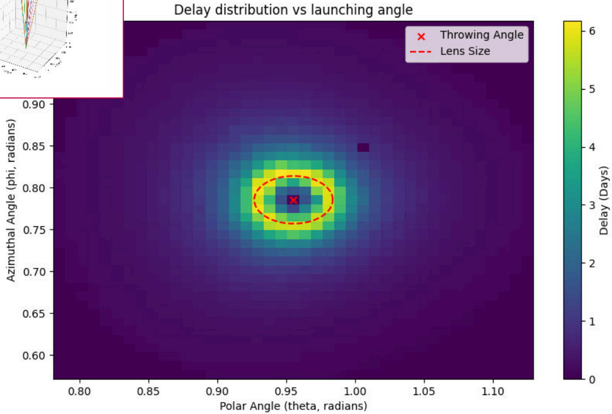


Figure 12: Time delay map for a realistic mass of $10^{15} M_\odot$ with a radius of 25 Mpc. The delay is measured between a deflected photon and a photon that would have traveled in a straight line. Figure from the LAPP internship project.

Relativistic light propagation and gravitational lensing

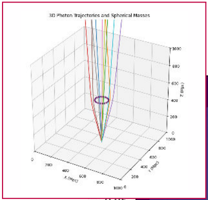


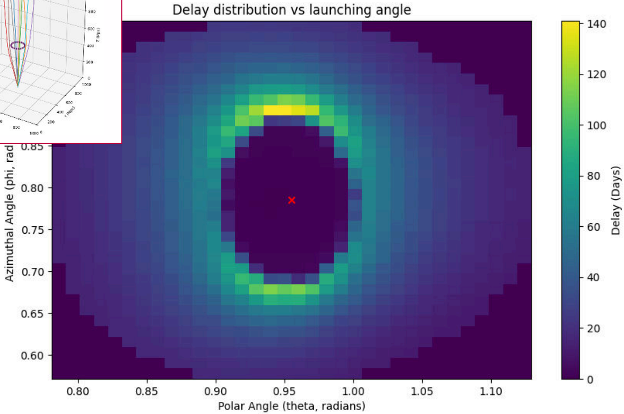


A weak but measurable effect

The time delay is of the order of a few days, which is very small compared to the time it takes for a photon to propagate (a few billion years).

Relativistic light propagation and gravitational lensing

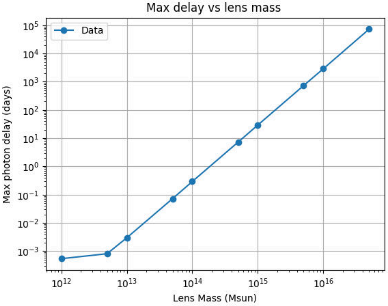




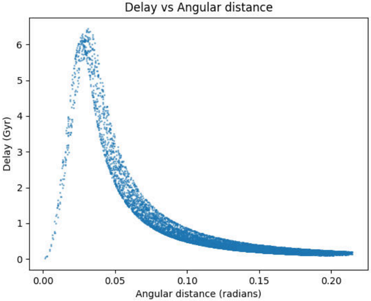
Different delay profiles for different mass distributions

The ring shape is clearly visible and rays that pass through the center are barely delayed, compared to the ones that pass closer to the ring.

Towards mass and radius estimations



(a) Possibility to estimate the mass of the cluster based on the measured delay



(b) Possibility to estimate the radius of the cluster based on the measured delay

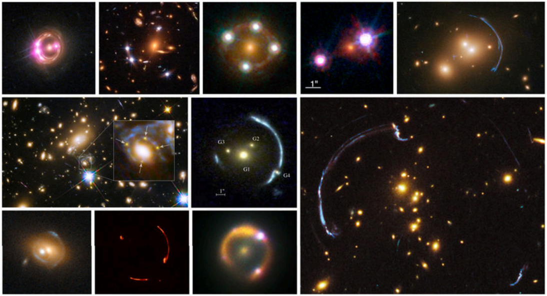
Figure 13: First attempts at a calibration of mass and radius based on time delay measurements. Here, considered halos are spherical and assume constant density profiles. The plot on the right assumes a 25 Mpc radius and a mass of $10^{15} M_\odot$

LSST data: a new era for cosmology

Credit: Paulo Lago

(a) The LSST telescope in Chile

Credit: Euclid/NASA/ESA/NRAO/AUI/NSF



Perspectives: the numerical tool EXCALIBUR

Scaling **EXCALIBUR**

Translating from Python to C++ for performance. Scaling the code to run on modern cosmological simulations and make it ready to handle large datasets.

Going beyond the prototype

- Include the expansion of the Universe
- Include redshift studies
- Consider specific effects (moving sources,...)
- Extract lensing observables

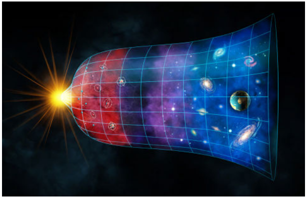
Figure 14: CC-IN2P3, Lyon

Perspectives: a unified framework for relativistic effects

Credit: NASA

Galaxies' peculiar velocities

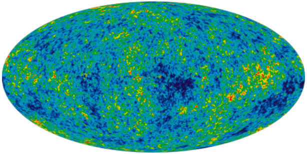
Credit: Andrea Danti



Universe expansion

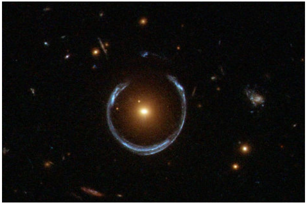
$$G_{\mu\nu} = \frac{8\pi G}{c^4} T_{\mu\nu}$$

Credit: WMAP



Integrated Sachs-Wolfe

Credit: ESA/HST



Gravitational lensing

A unified theoretical framework

Studying relativistic effects related to light propagation in cosmology.