

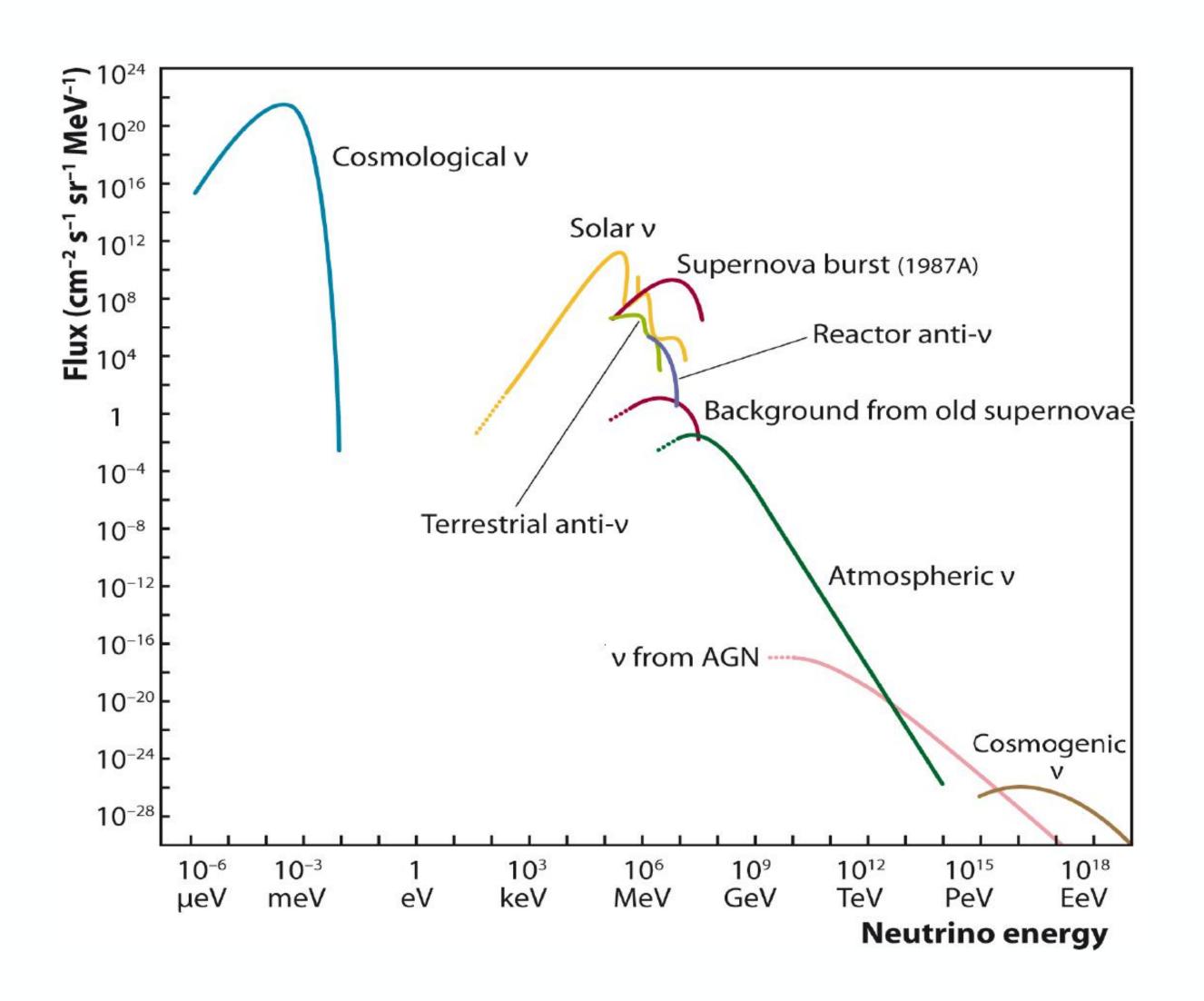




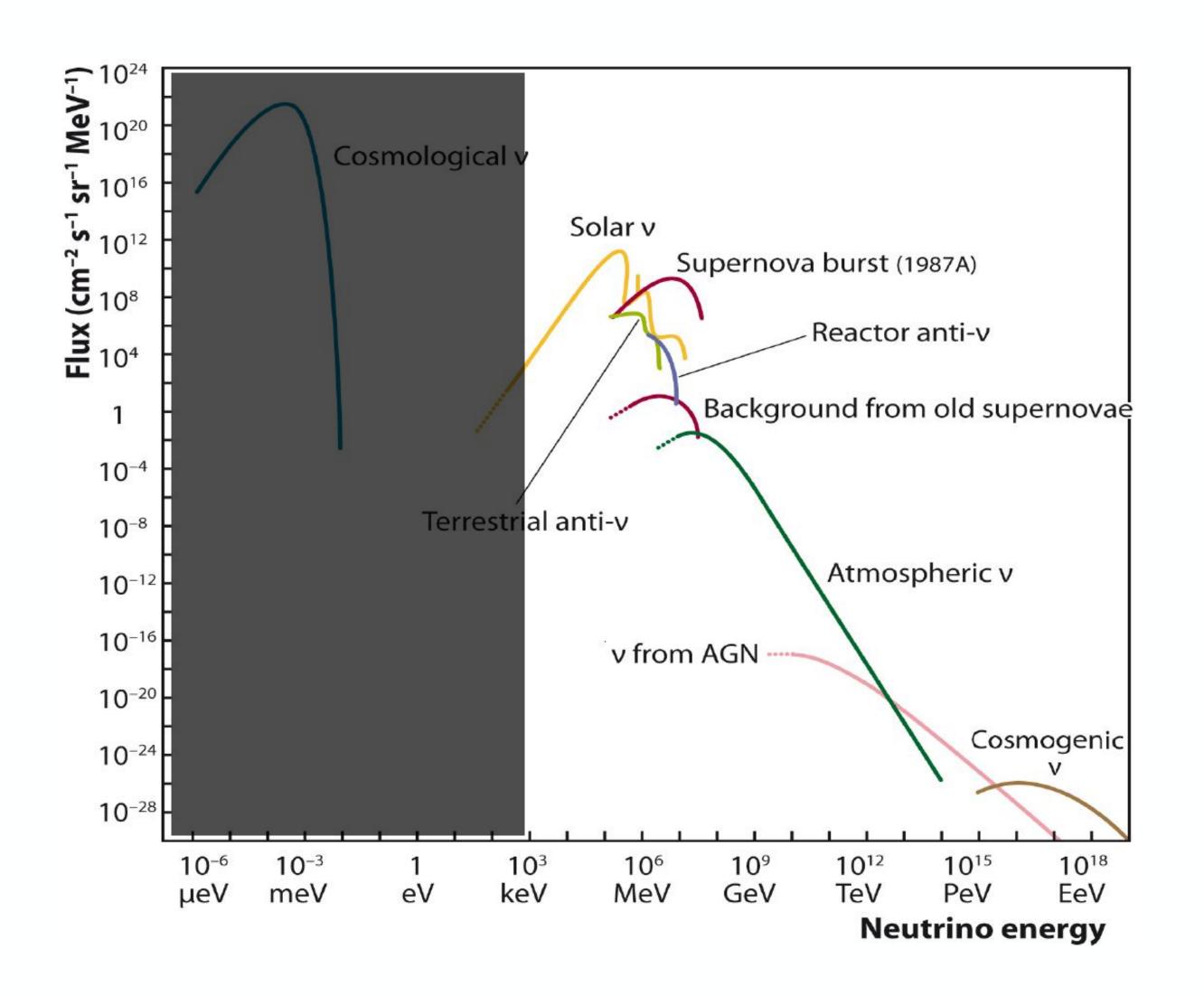
Results from KM3NeT and IceCube

Review on neutrino astronomy at high-energy (and beyond)

"Global" neutrino spectrum:



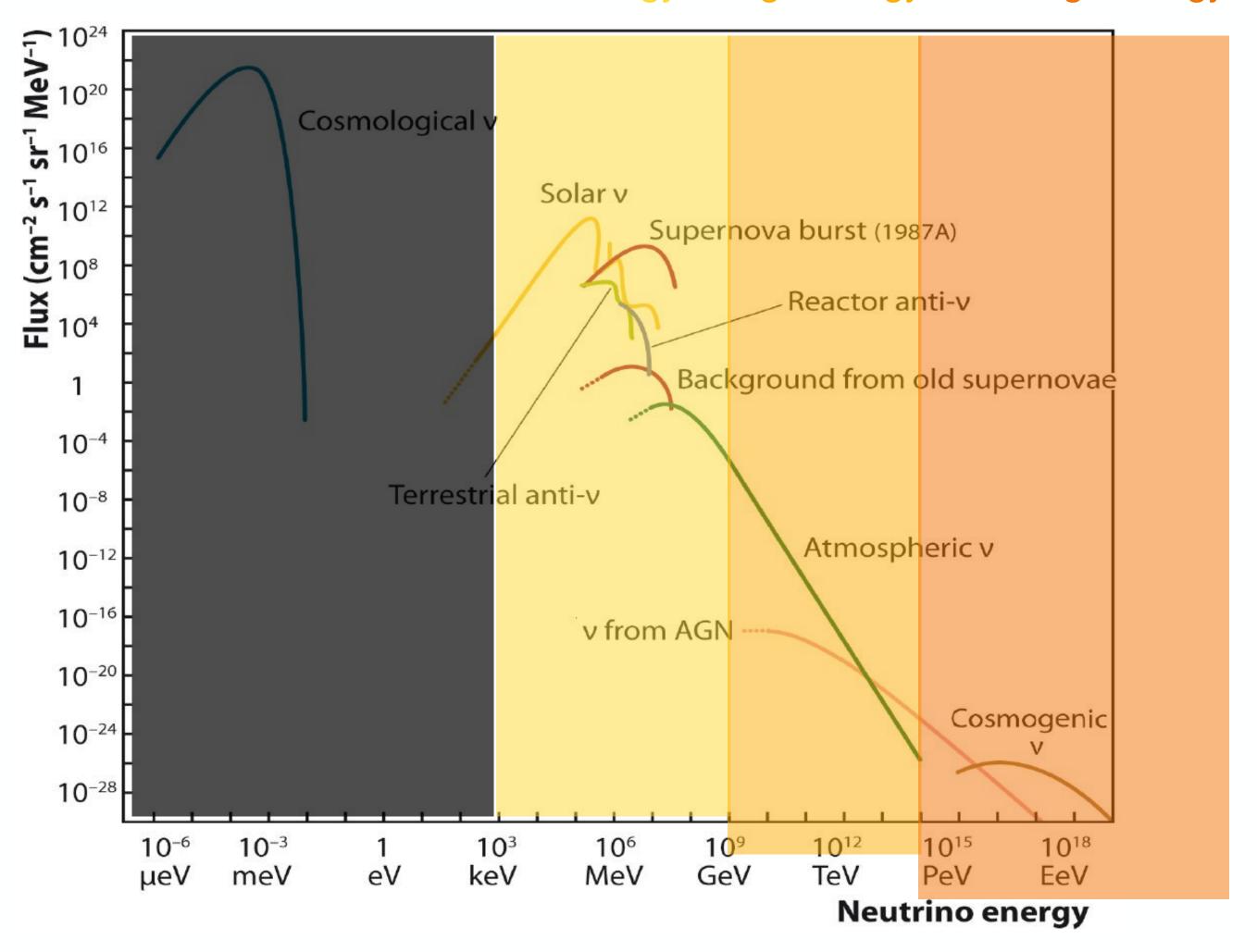
"Global" neutrino spectrum:



"Global" neutrino spectrum:

• 3 energy scales

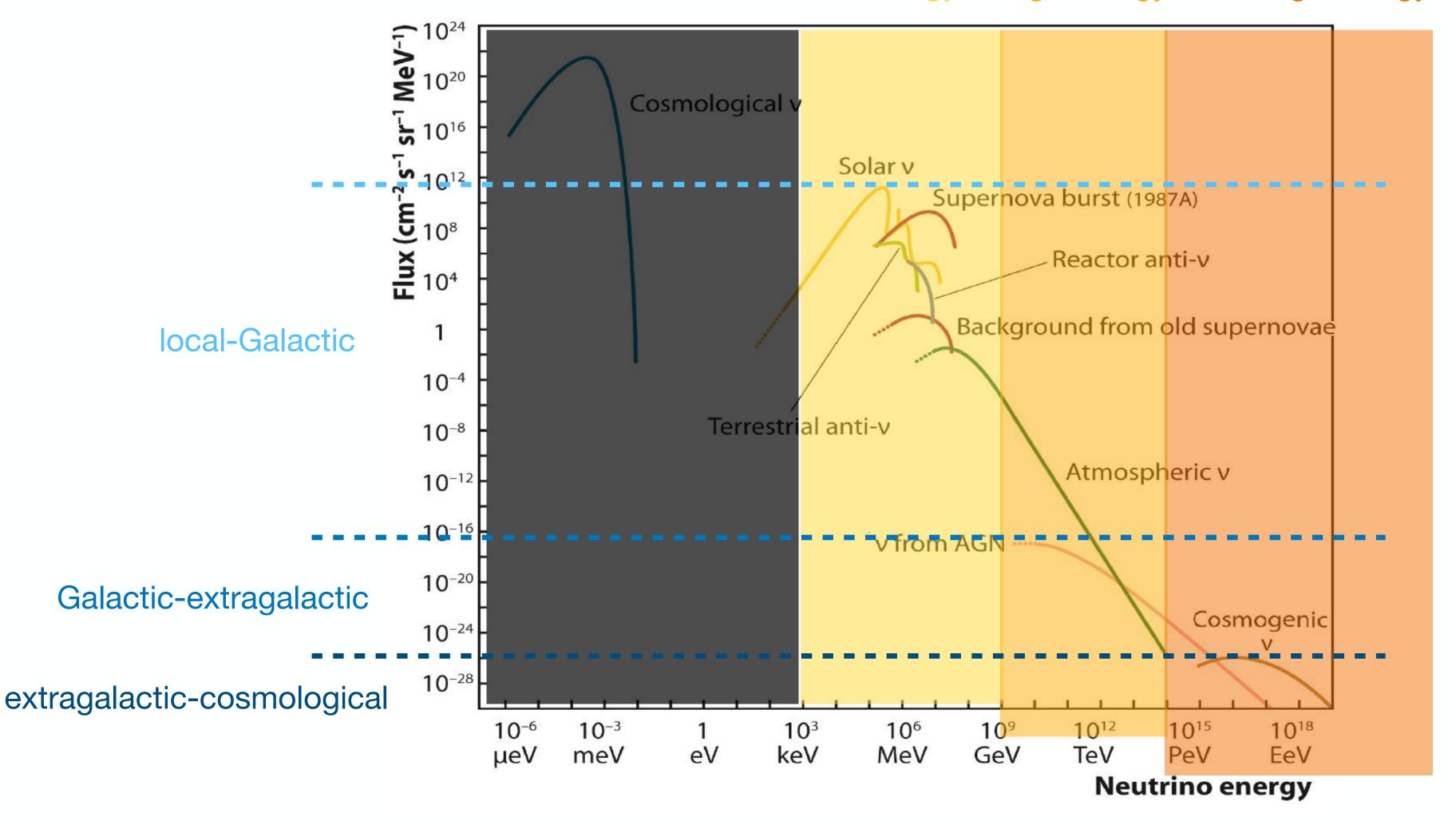
low energy high energy ultra-high energy



low energy high energy ultra-high energy

"Global" neutrino spectrum:

- 3 energy scales
- 3 distance scales

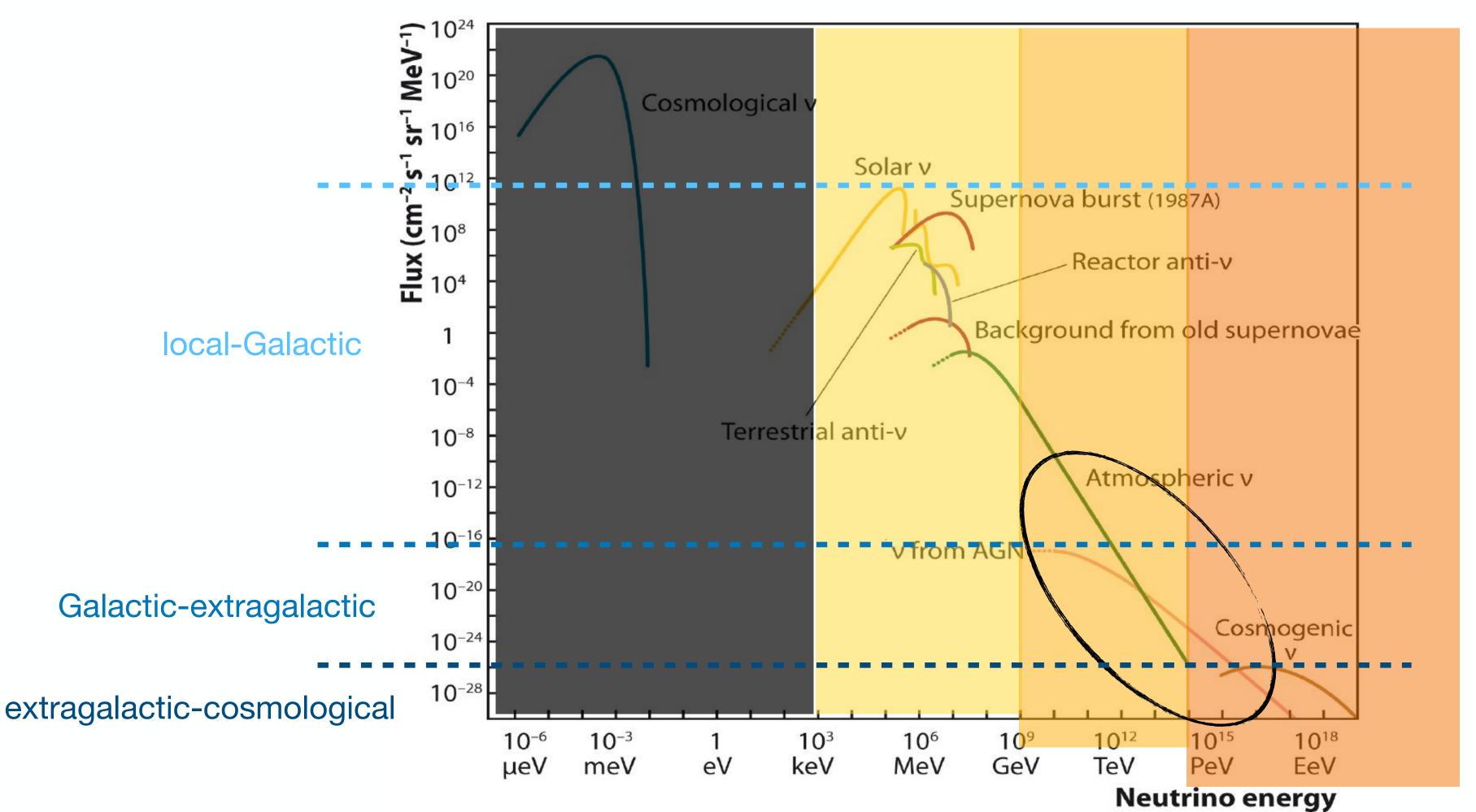


C. Spiering 2020 **2**

low energy high energy ultra-high energy

"Global" neutrino spectrum:

- 3 energy scales
- 3 distance scales

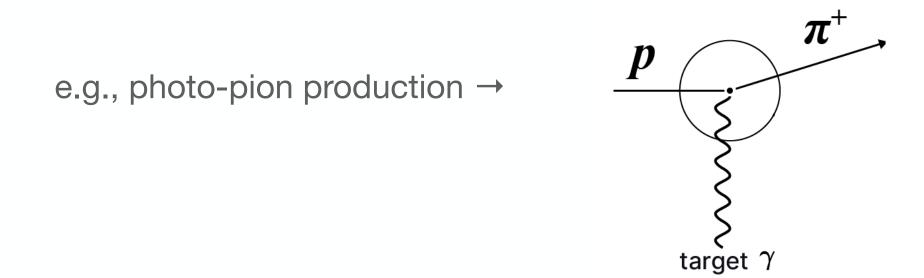


Focus on high energy neutrinos → local to extragalactic sources

One "standard" scenario for high-energy neutrino production:

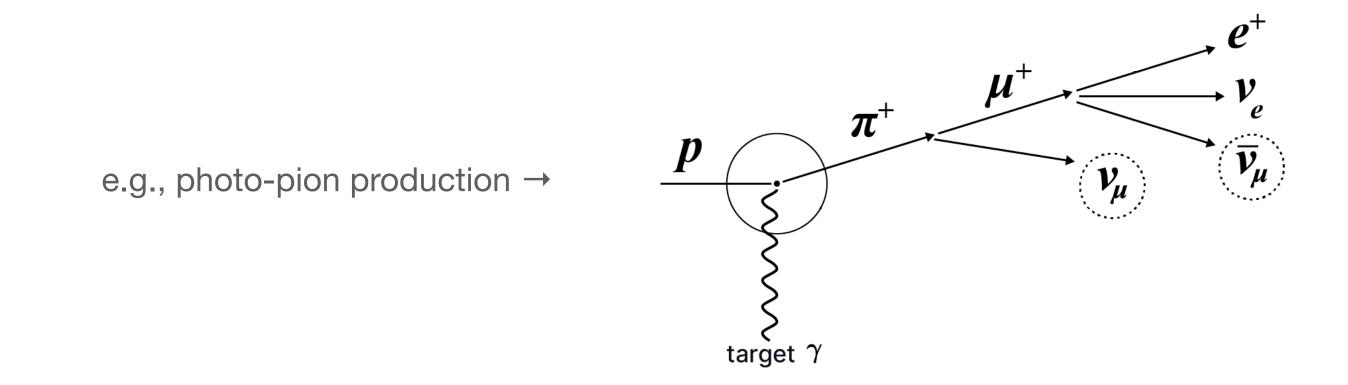
One "standard" scenario for high-energy neutrino production:

• Photonuclear or baryonic interactions → charged mesons



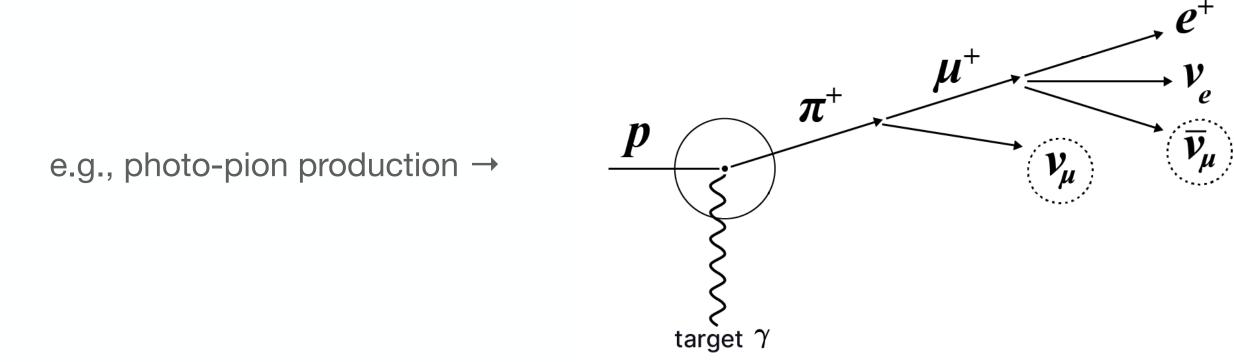
One "standard" scenario for high-energy neutrino production:

- Photonuclear or baryonic interactions → charged mesons
- Decay of charged mesons → neutrinos



One "standard" scenario for high-energy neutrino production:

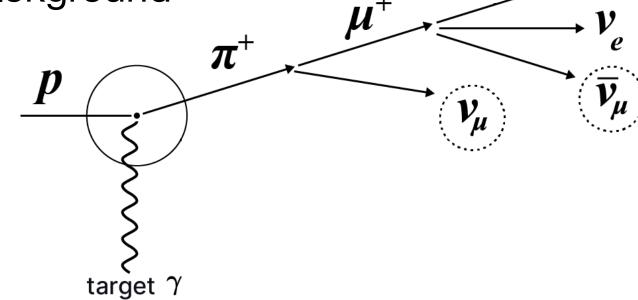
- Photonuclear or baryonic interactions → charged mesons
- Decay of charged mesons → neutrinos
- → Neutrino energy ~5% of primary energy

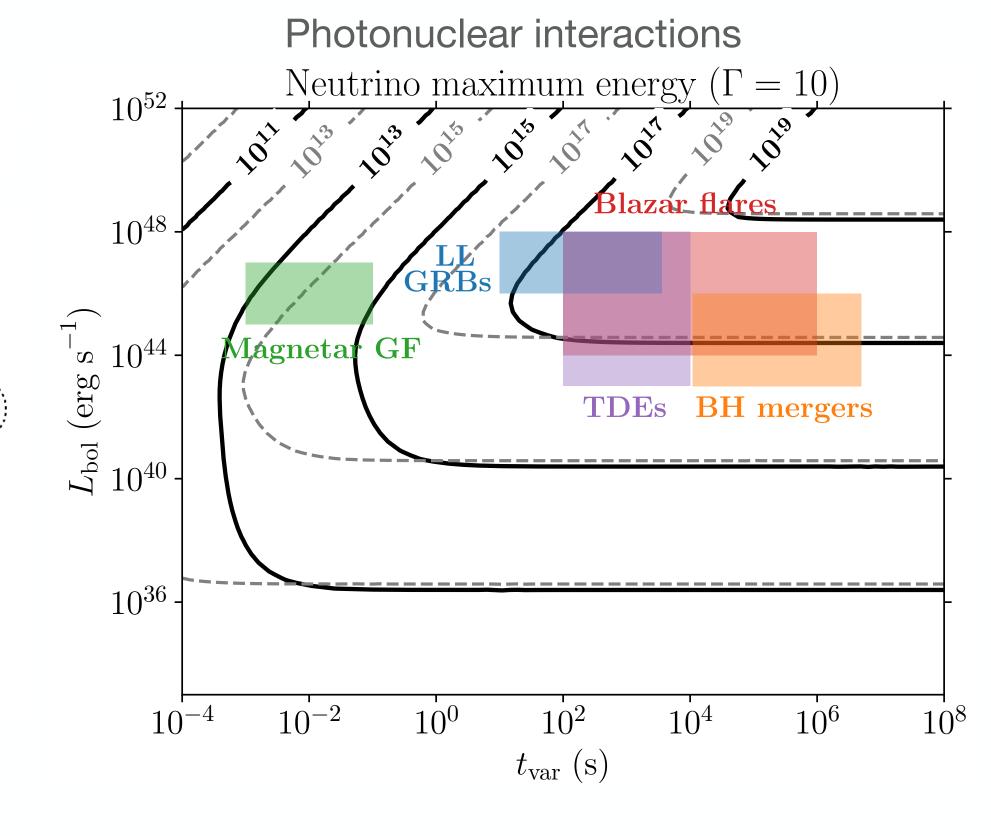


One "standard" scenario for high-energy neutrino production:

- Photonuclear or baryonic interactions → charged mesons
- Decay of charged mesons → neutrinos
- → Neutrino energy ~5% of primary energy
- → Accelerated cosmic rays + photon or baryon background

e.g., photo-pion production →



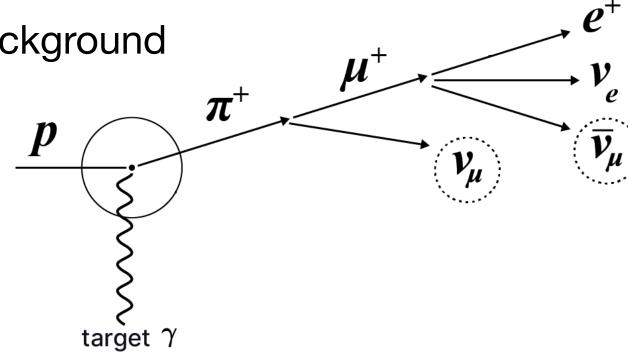


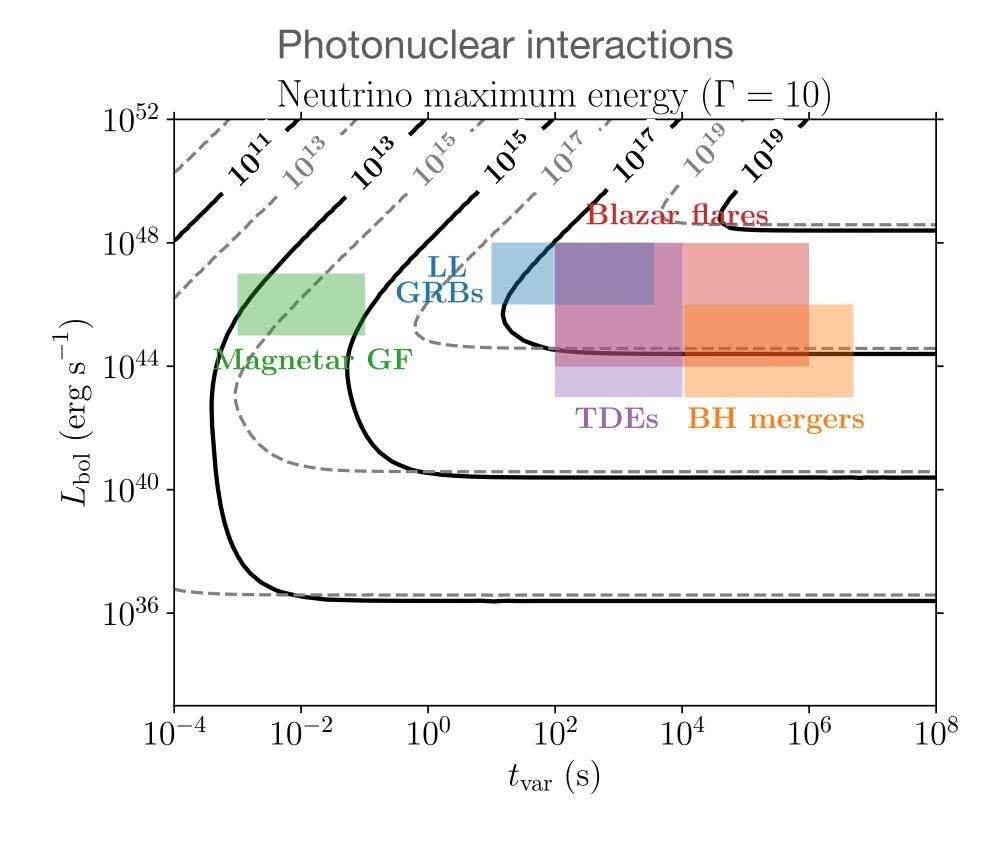
One "standard" scenario for high-energy neutrino production:

- Photonuclear or baryonic interactions → charged mesons
- Decay of charged mesons → neutrinos
- → Neutrino energy ~5% of primary energy
- → Accelerated cosmic rays + photon or baryon background

e.g., photo-pion production →

Co-signatures:

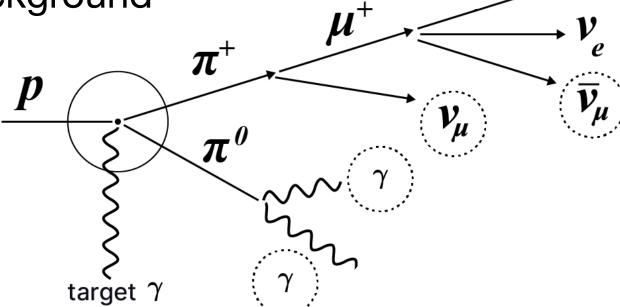




One "standard" scenario for high-energy neutrino production:

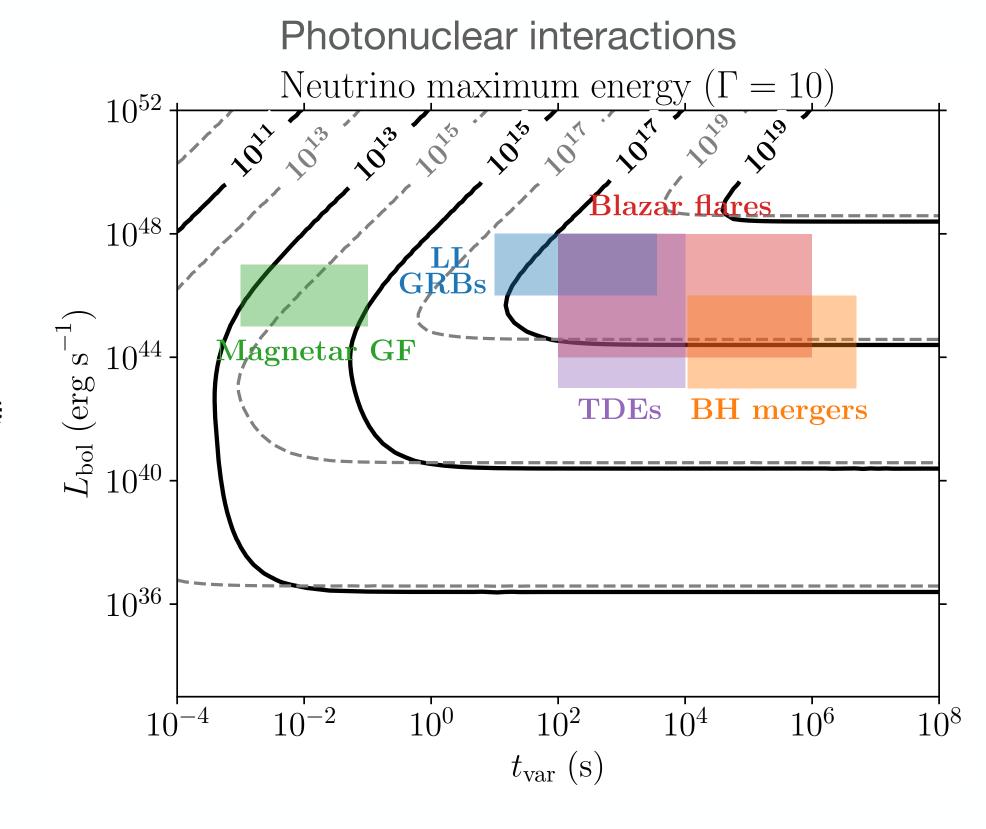
- Photonuclear or baryonic interactions → charged mesons
- Decay of charged mesons → neutrinos
- → Neutrino energy ~5% of primary energy
- → Accelerated cosmic rays + photon or baryon background

e.g., photo-pion production →



Co-signatures:

Decay of neutral mesons → high-energy gamma rays



One "standard" scenario for high-energy neutrino production:

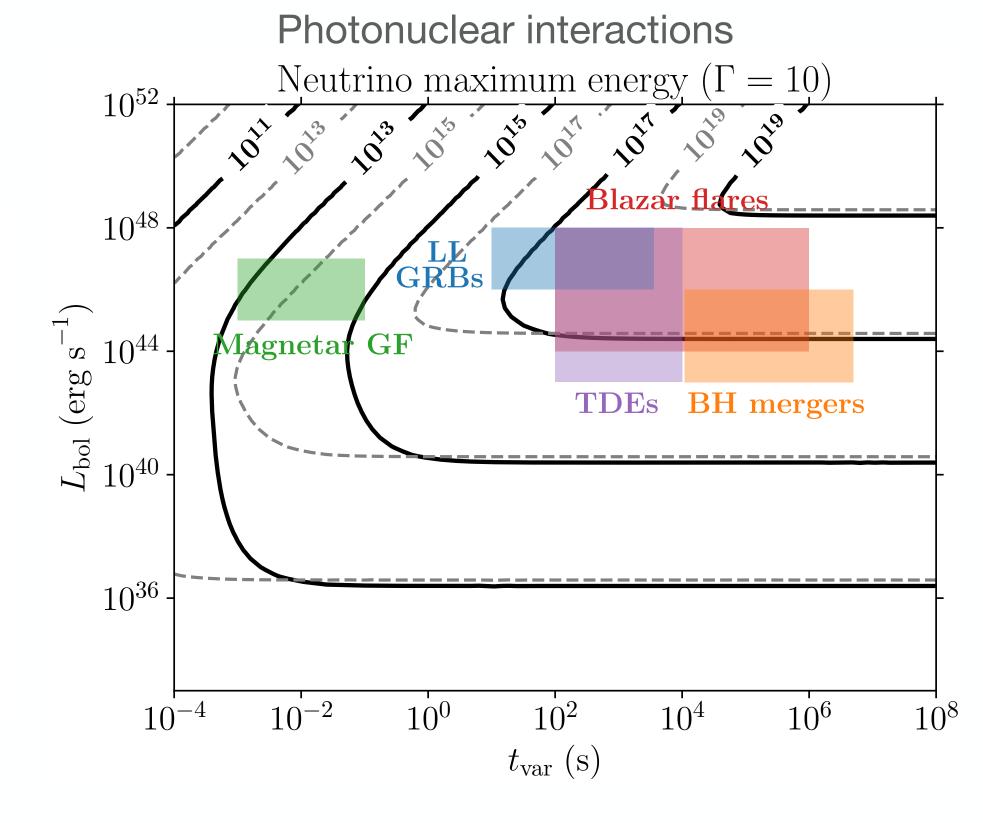
- Photonuclear or baryonic interactions → charged mesons
- Decay of charged mesons → neutrinos
- → Neutrino energy ~5% of primary energy
- → Accelerated cosmic rays + photon or baryon background

e.g., photo-pion production →

$\begin{array}{c|c} p & \pi^+ & \psi_e \\ \hline \pi^0 & \psi_\mu & \overline{\psi}_\mu \end{array}$ target γ

Co-signatures:

- Decay of neutral mesons → high-energy gamma rays
- Radiation of subsequent leptons → gamma rays



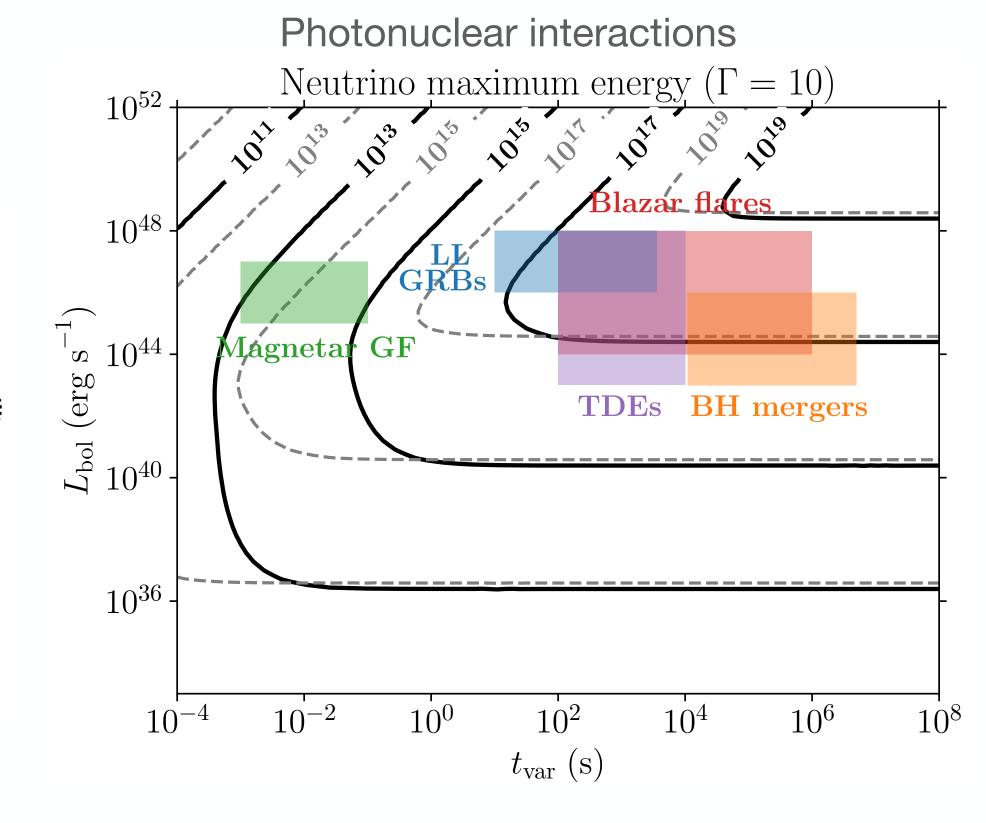
One "standard" scenario for high-energy neutrino production:

- Photonuclear or baryonic interactions → charged mesons
- Decay of charged mesons → neutrinos
- → Neutrino energy ~5% of primary energy
- → Accelerated cosmic rays + photon or baryon background

e.g., photo-pion production →

Co-signatures:

- Decay of neutral mesons → high-energy gamma rays
- Radiation of subsequent leptons → gamma rays
- → Intrinsic connections to gamma rays



One "standard" scenario for high-energy neutrino production:

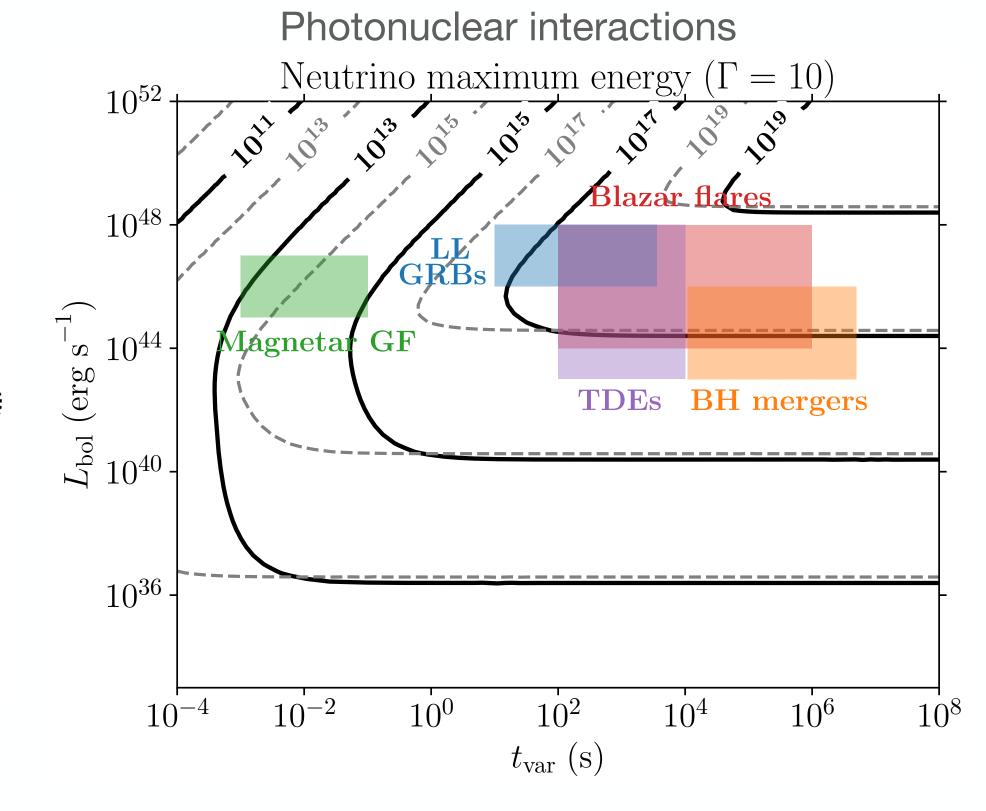
- Photonuclear or baryonic interactions → charged mesons
- Decay of charged mesons → neutrinos
- → Neutrino energy ~5% of primary energy
- → Accelerated cosmic rays + photon or baryon background

e.g., photo-pion production →

ckground μ^+ ν_e ν_μ ν_μ ν_μ ν_μ ν_μ ν_μ

Co-signatures:

- Decay of neutral mesons → high-energy gamma rays
- Radiation of subsequent leptons → gamma rays
- → Intrinsic connections to gamma rays



High-energy neutrinos can probe internal hadronic mechanisms of cosmic-ray and high-energy sources

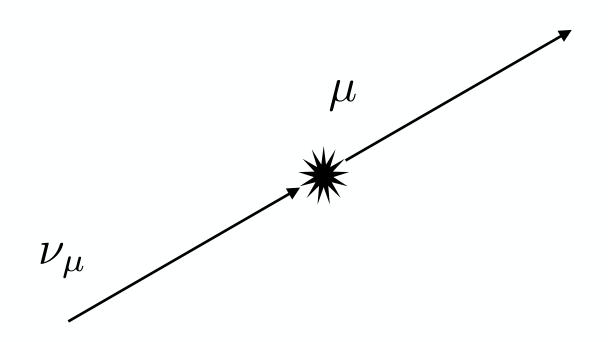
A concept from 1960s (see DUMAND):

Look for charged and neutral current interactions from neutrinos

A concept from 1960s (see DUMAND):

Look for charged and neutral current interactions from neutrinos

$$\rightarrow \nu_l + X \rightarrow Y + l \qquad (l = e, \mu, \tau)$$

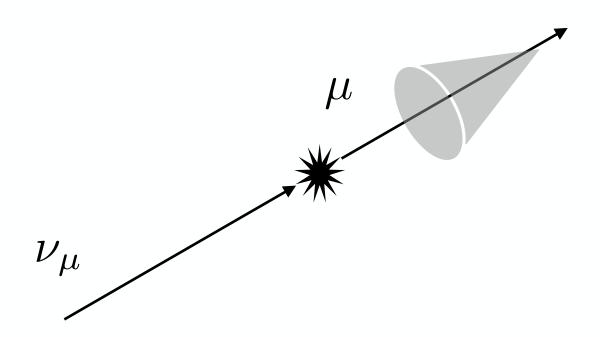


A concept from 1960s (see DUMAND):

Look for charged and neutral current interactions from neutrinos

$$\rightarrow \nu_l + X \rightarrow Y + l \qquad (l = e, \mu, \tau)$$

- The lepton propagates or decays
 - → leaving a Cherenkov track or cascade

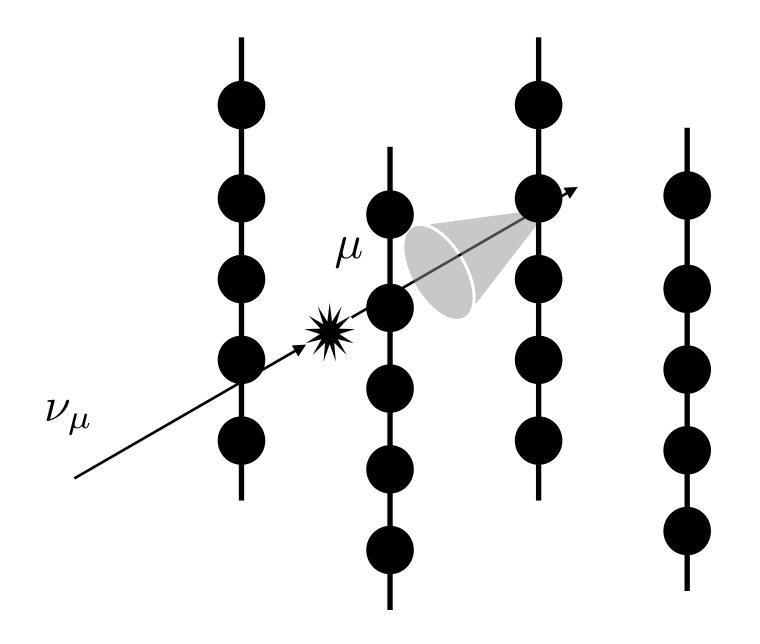


A concept from 1960s (see DUMAND):

Look for charged and neutral current interactions from neutrinos

$$\rightarrow \nu_l + X \rightarrow Y + l \qquad (l = e, \mu, \tau)$$

- The lepton propagates or decays
 - → leaving a Cherenkov track or cascade
- Record the Cherenkov lights with large arrays of PMTs



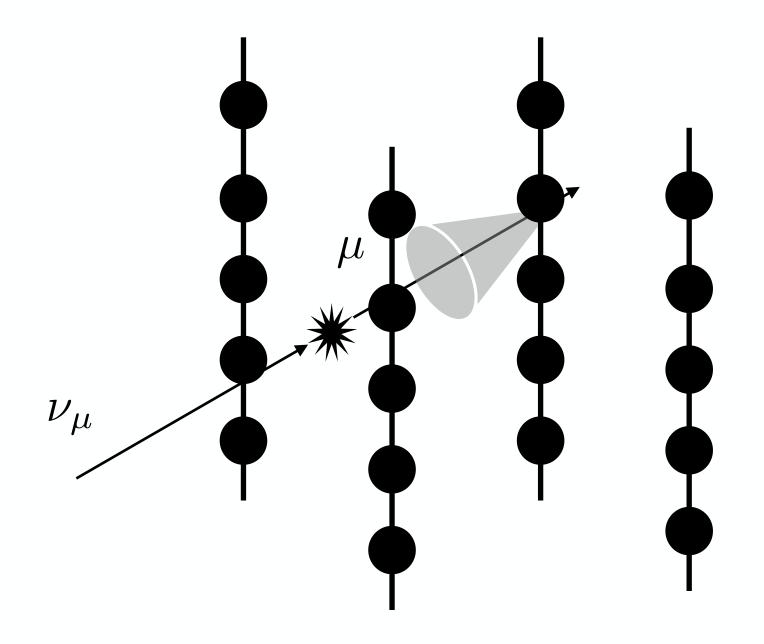
A concept from 1960s (see DUMAND):

Look for charged and neutral current interactions from neutrinos

$$\rightarrow \nu_l + X \rightarrow Y + l \qquad (l = e, \mu, \tau)$$

- The lepton propagates or decays
 - → leaving a Cherenkov track or cascade
- Record the Cherenkov lights with large arrays of PMTs

3 main sources of background:



A concept from 1960s (see DUMAND):

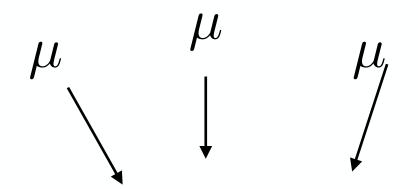
Look for charged and neutral current interactions from neutrinos

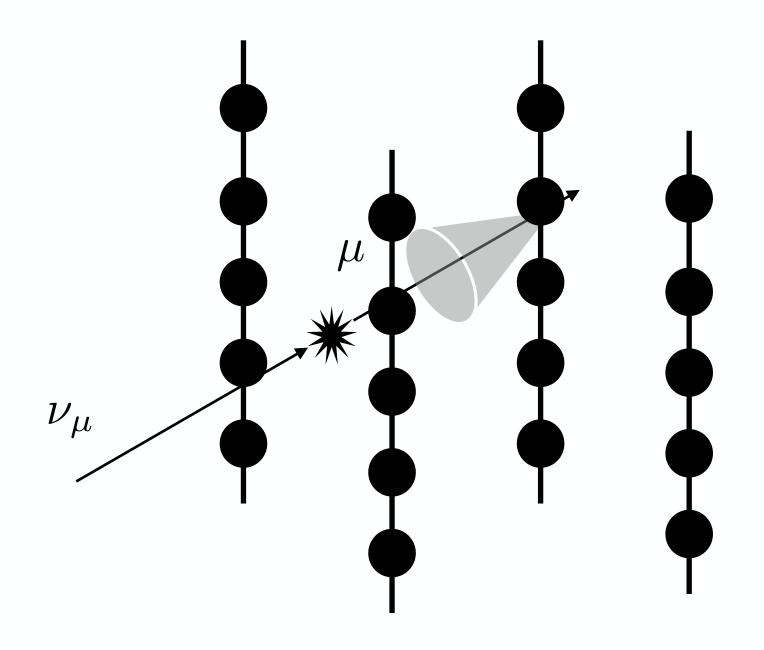
$$\rightarrow \nu_l + X \rightarrow Y + l \qquad (l = e, \mu, \tau)$$

- The lepton propagates or decays
 - → leaving a Cherenkov track or cascade
- Record the Cherenkov lights with large arrays of PMTs

3 main sources of background:

Atmospheric muons





A concept from 1960s (see DUMAND):

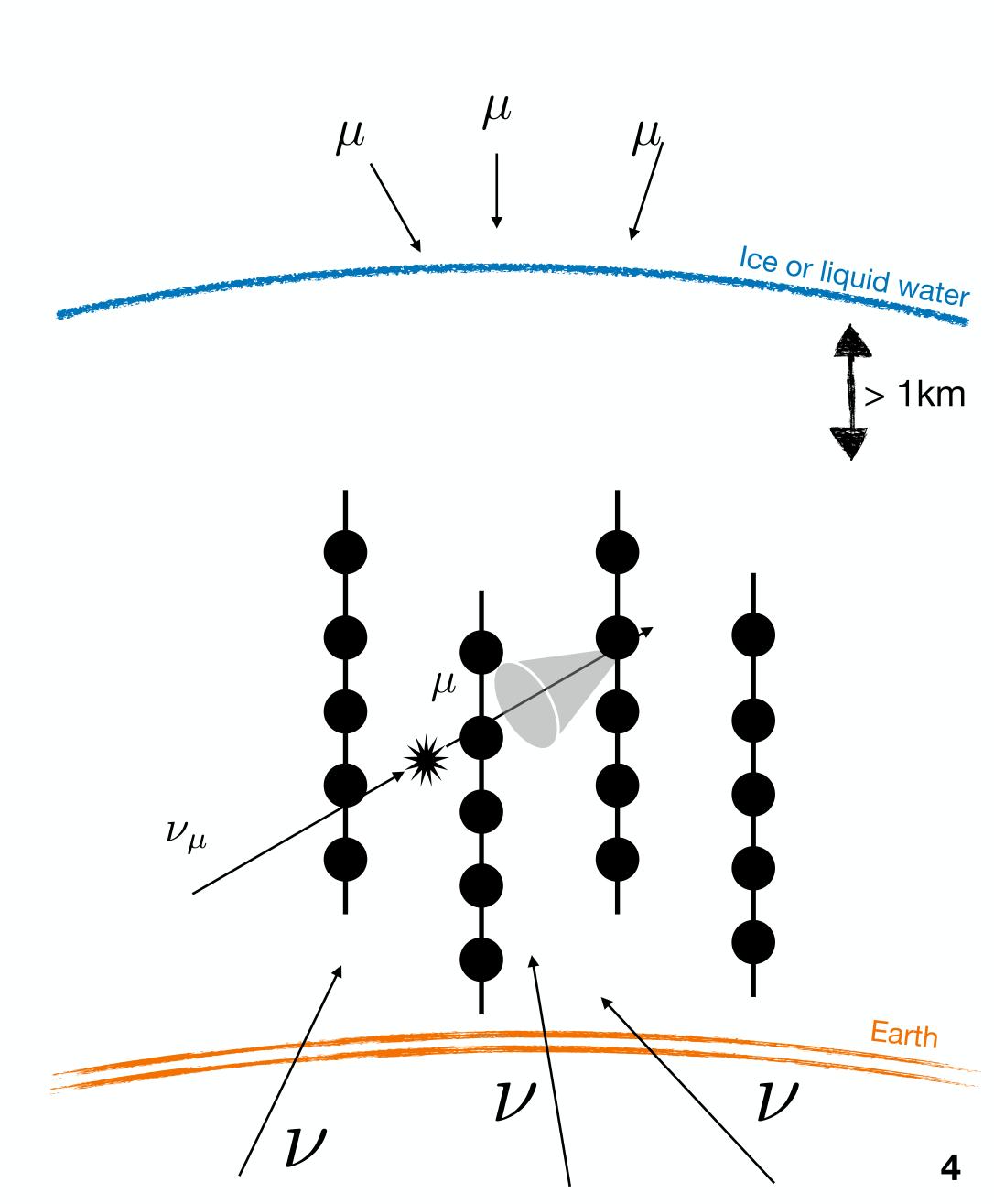
Look for charged and neutral current interactions from neutrinos

$$\rightarrow \nu_l + X \rightarrow Y + l \qquad (l = e, \mu, \tau)$$

- The lepton propagates or decays
 - → leaving a Cherenkov track or cascade
- Record the Cherenkov lights with large arrays of PMTs

3 main sources of background:

• Atmospheric muons → shielding / up-going events



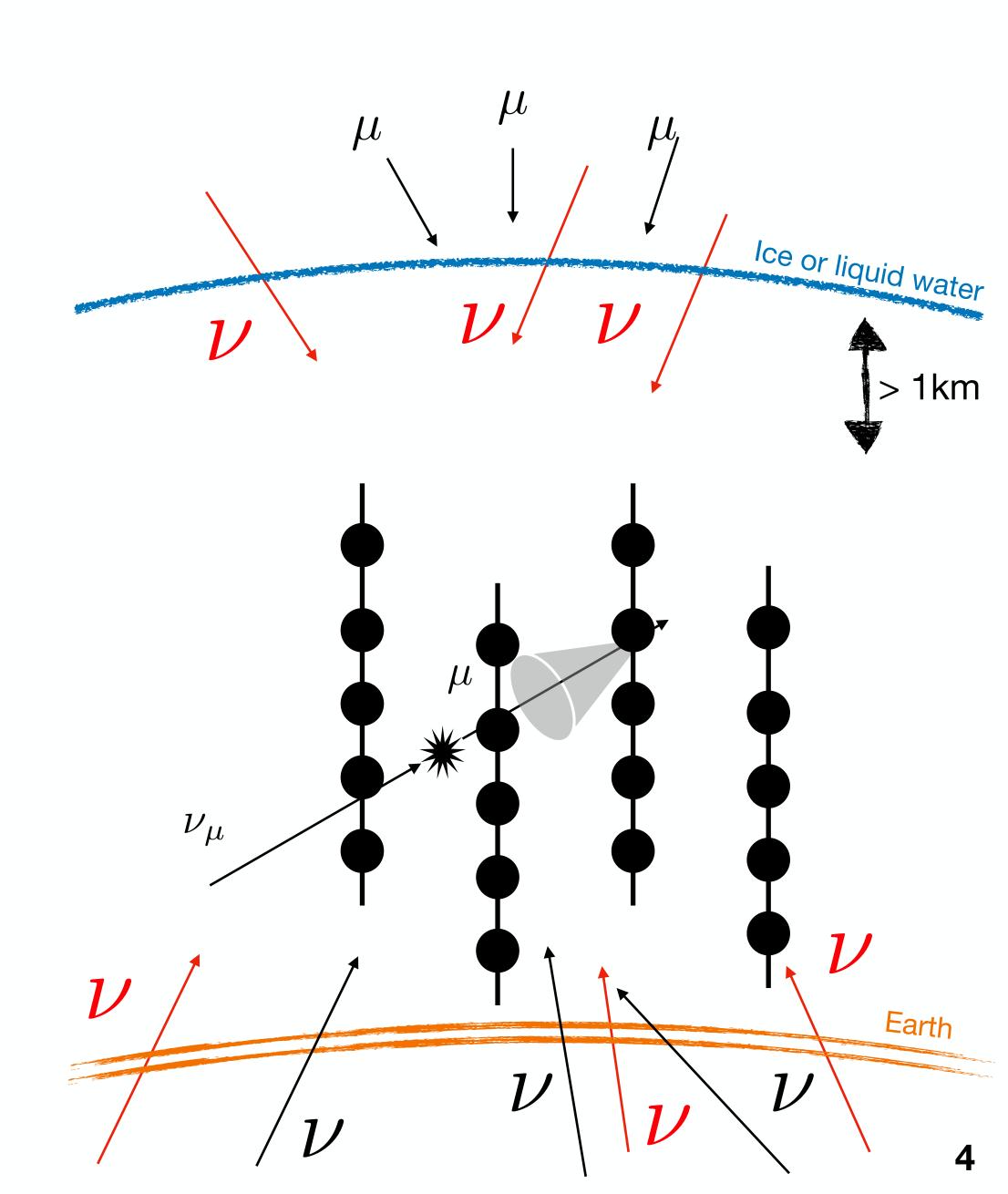
A concept from 1960s (see DUMAND):

Look for charged and neutral current interactions from neutrinos

$$\rightarrow \nu_l + X \rightarrow Y + l \qquad (l = e, \mu, \tau)$$

- The lepton propagates or decays
 - → leaving a Cherenkov track or cascade
- Record the Cherenkov lights with large arrays of PMTs

- Atmospheric muons → shielding / up-going events
- Atmospheric neutrinos



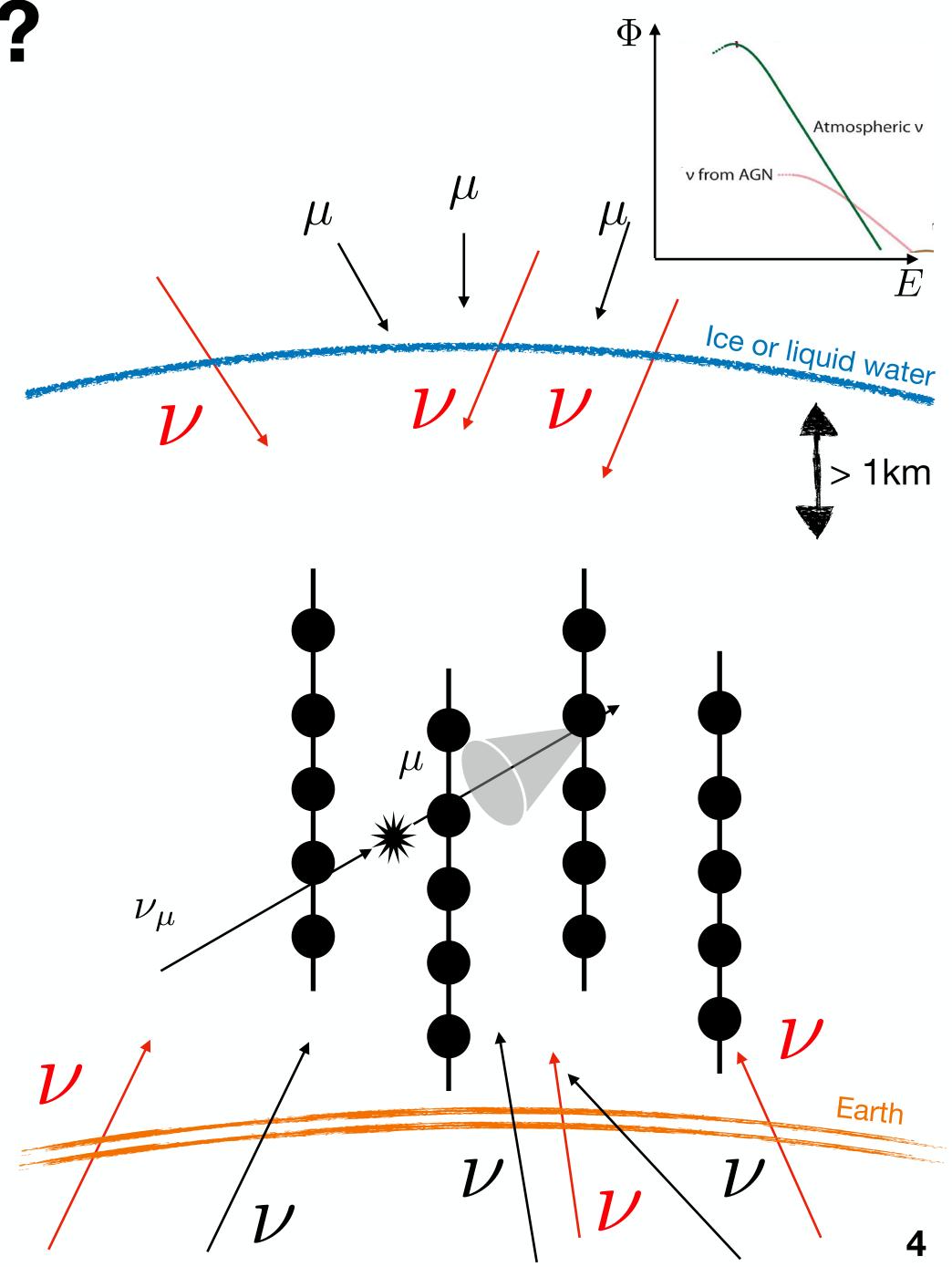
A concept from 1960s (see DUMAND):

Look for charged and neutral current interactions from neutrinos

$$\rightarrow \nu_l + X \rightarrow Y + l \qquad (l = e, \mu, \tau)$$

- The lepton propagates or decays
 - → leaving a Cherenkov track or cascade
- Record the Cherenkov lights with large arrays of PMTs

- Atmospheric muons → shielding / up-going events
- Atmospheric neutrinos → statistical disentangling (spectral shape)



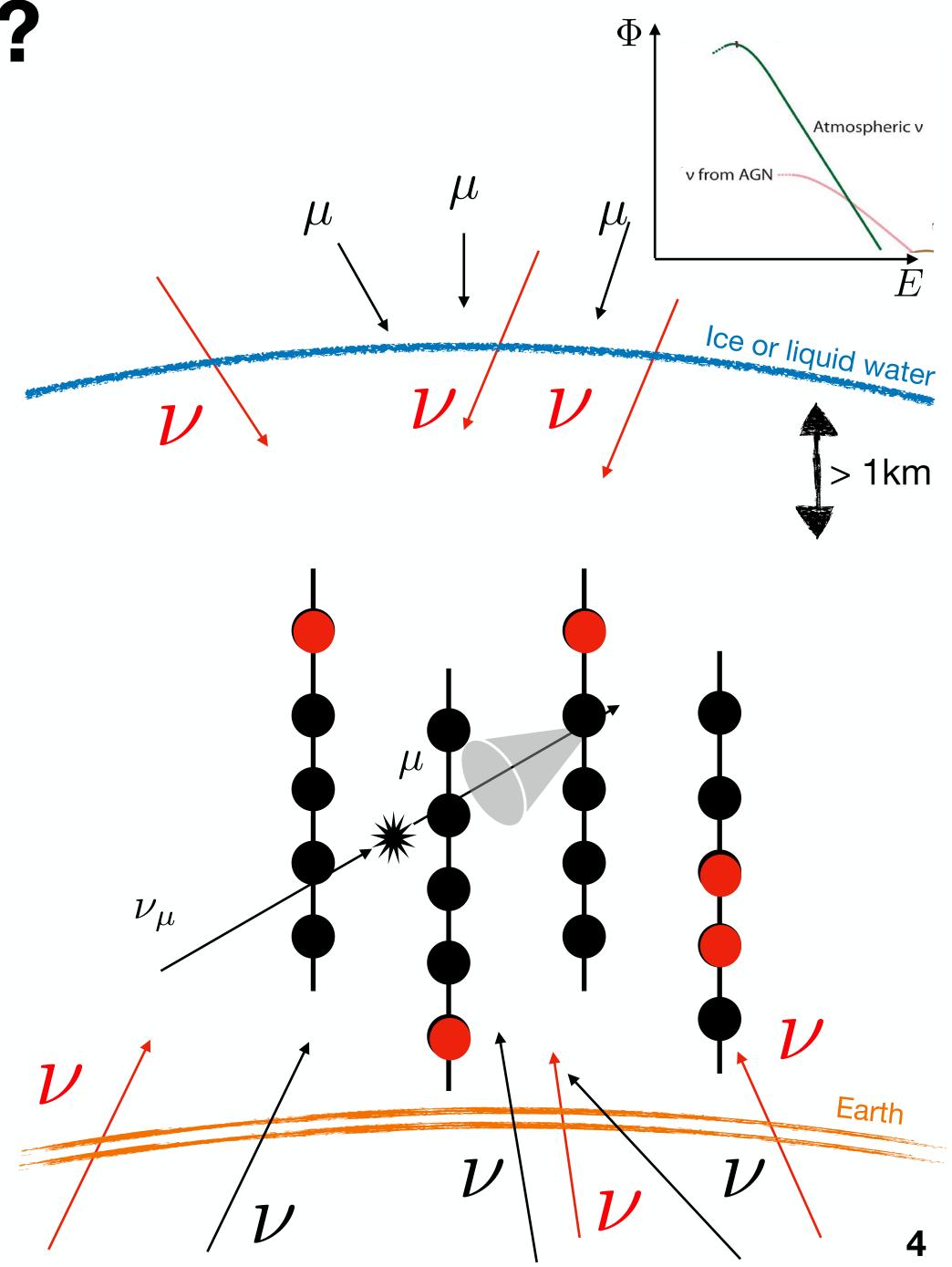
A concept from 1960s (see DUMAND):

Look for charged and neutral current interactions from neutrinos

$$\rightarrow \nu_l + X \rightarrow Y + l \qquad (l = e, \mu, \tau)$$

- The lepton propagates or decays
 - → leaving a Cherenkov track or cascade
- Record the Cherenkov lights with large arrays of PMTs

- Atmospheric muons → shielding / up-going events
- Atmospheric neutrinos → statistical disentangling (spectral shape)
- Random backgrounds (dark, radioactivity...)



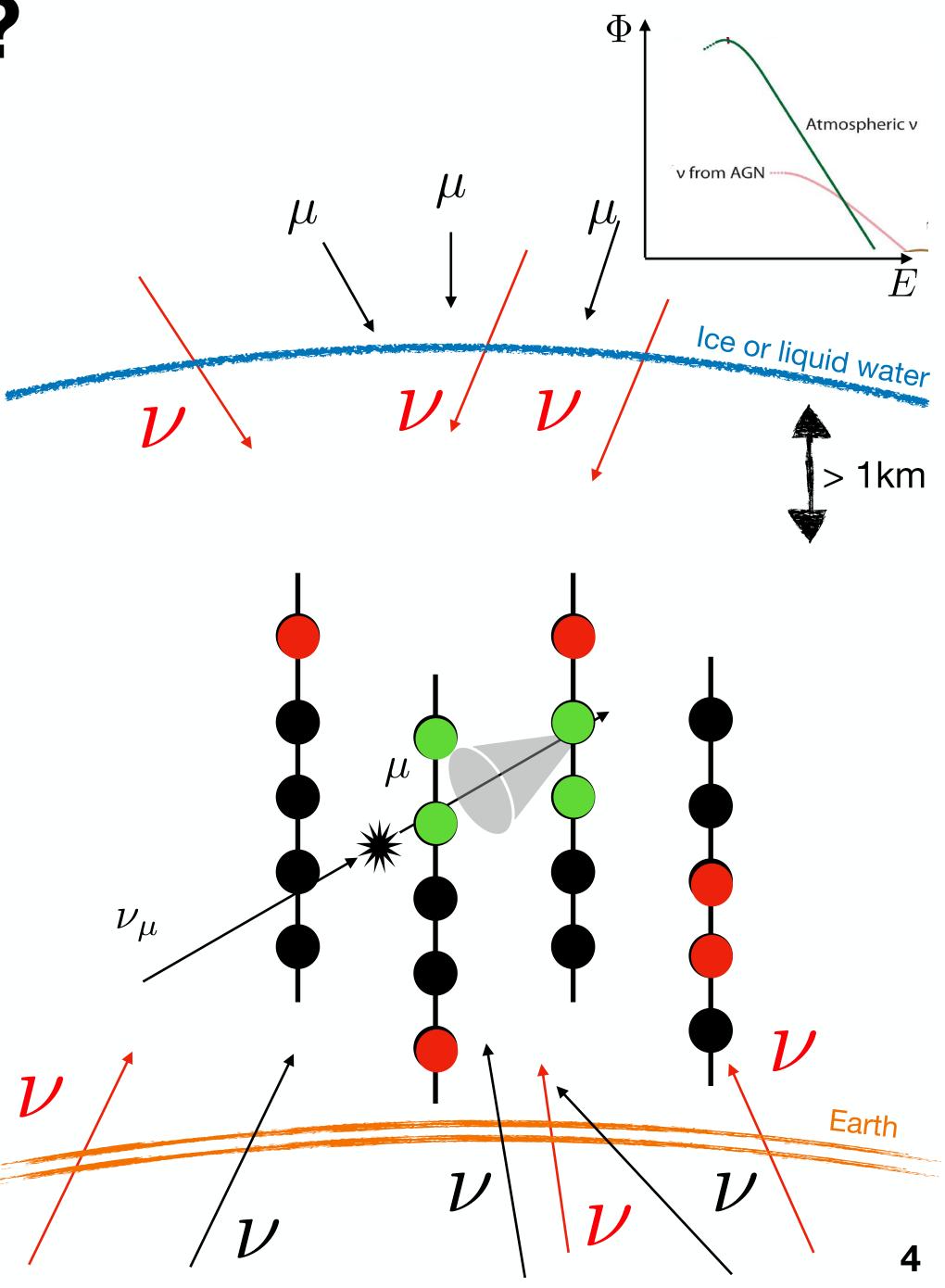
A concept from 1960s (see DUMAND):

Look for charged and neutral current interactions from neutrinos

$$\rightarrow \nu_l + X \rightarrow Y + l \qquad (l = e, \mu, \tau)$$

- The lepton propagates or decays
 - → leaving a Cherenkov track or cascade
- Record the Cherenkov lights with large arrays of PMTs

- Atmospheric muons → shielding / up-going events
- Atmospheric neutrinos → statistical disentangling (spectral shape)
- Random backgrounds (dark, radioactivity...) → veto with local coincidences



A concept from 1960s (see DUMAND):

Look for charged and neutral current interactions from neutrinos

$$\rightarrow \nu_l + X \rightarrow Y + l \qquad (l = e, \mu, \tau)$$

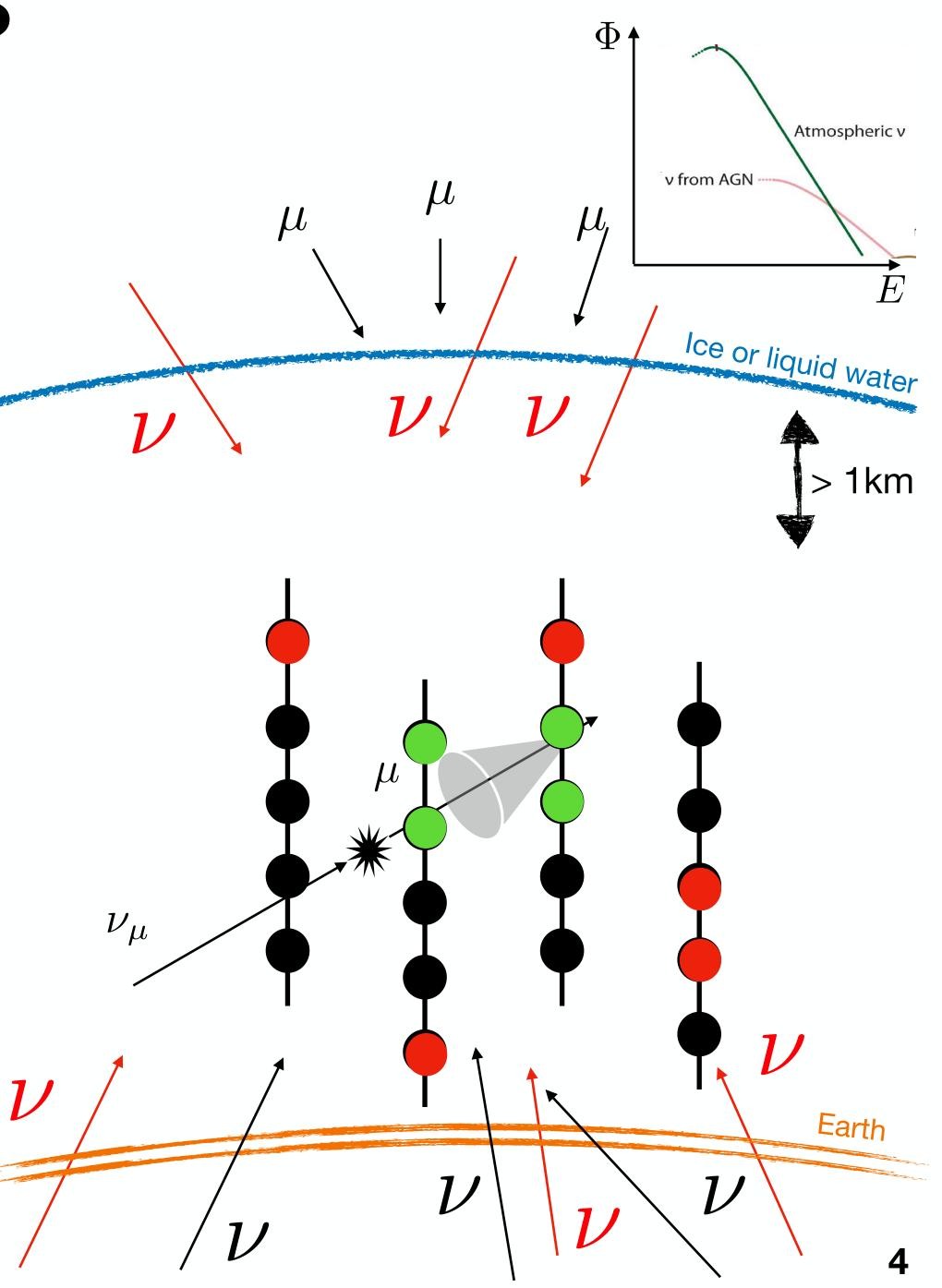
- The lepton propagates or decays
 - → leaving a Cherenkov track or cascade
- Record the Cherenkov lights with large arrays of PMTs

3 main sources of background:

- Atmospheric muons → shielding / up-going events
- Atmospheric neutrinos → statistical disentangling (spectral shape)
- Random backgrounds (dark, radioactivity...) → veto with local coincidences

Main challenge: low fluxes + weakly interacting particle

→ large effective volume → large scale detectors

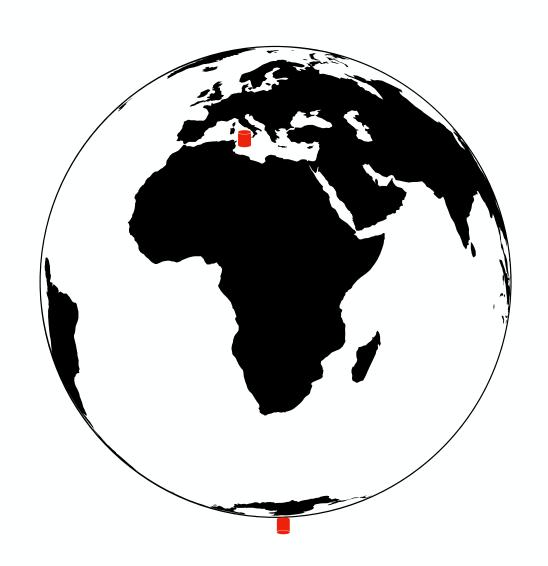


~1Gton of instrumented water target (Ice or liquid)

minimum energy: 100-300GeV

peak energy: 100TeV

• sensitivity: 10⁻⁸-10⁻⁹ GeV cm⁻²s⁻¹ sr⁻¹



~1Gton of instrumented water target (Ice or liquid)

• minimum energy: 100-300GeV

peak energy: 100TeV

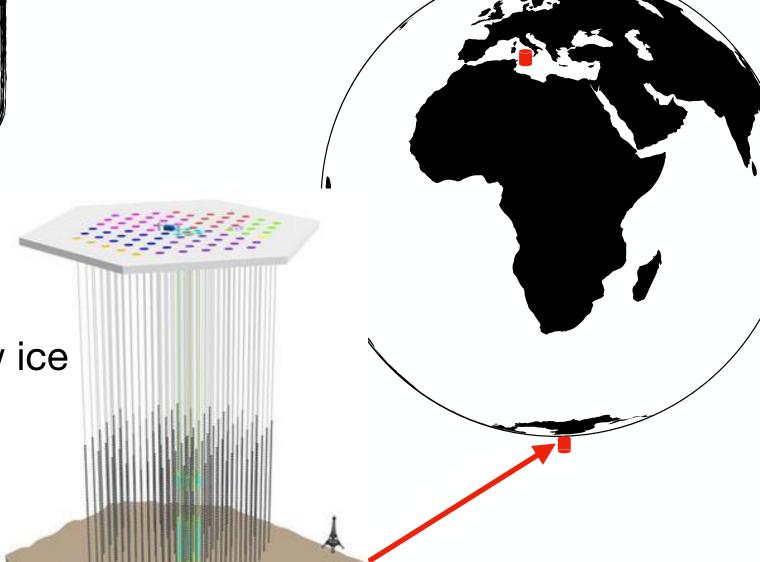
• sensitivity: 10⁻⁸-10⁻⁹ GeV cm⁻²s⁻¹ sr⁻¹

IceCube (South Pole):

FoV: towards Galactic ridge

angular resolutions: 0.4° (10°)

completion date: 2010





- ~2500m below ice
- 86 strings
- 5160 DOMs

~1Gton of instrumented water target (Ice or liquid)

• minimum energy: 100-300GeV

peak energy: 100TeV

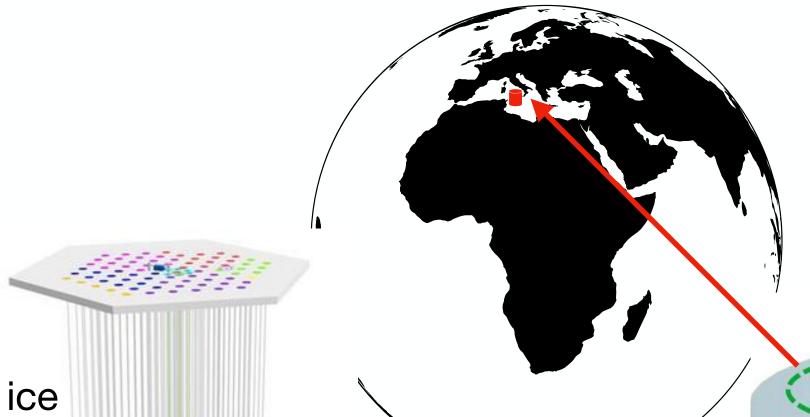
• sensitivity: 10⁻⁸-10⁻⁹ GeV cm⁻²s⁻¹ sr⁻¹

IceCube (South Pole):

FoV: towards Galactic ridge

angular resolutions: 0.4° (10°)

completion date: 2010



KM3NeT-ARCA (Mediterranean sea):

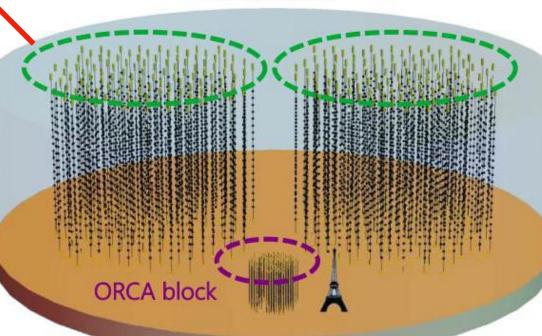
FoV: towards Galactic center

• angular resolutions: 0.1° (1.5°)

• completion date: ~2028



- ~2500m below ice
- 86 strings
- 5160 DOMs



ARCA blocks

- ~3500m b.s.l
- 345 strings
- 6210 DOMs
- 192510 PMTs



~1Gton of instrumented water target (Ice or liquid)

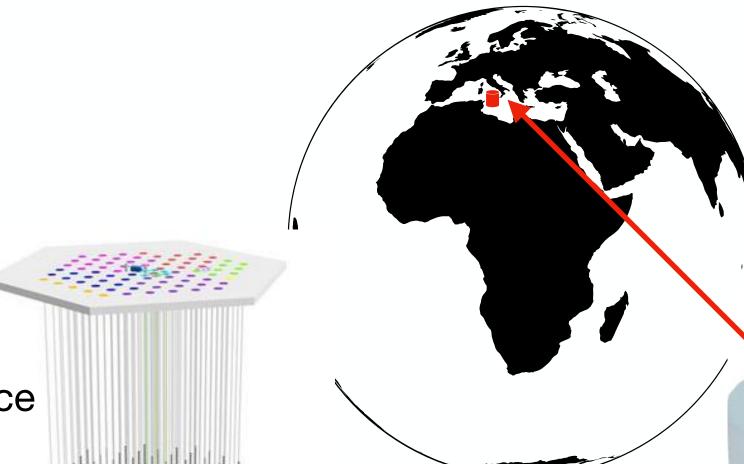
minimum energy: 100-300GeV

peak energy: 100TeV

• sensitivity: 10⁻⁸-10⁻⁹ GeV cm⁻²s⁻¹ sr⁻¹

IceCube (South Pole):

- FoV: towards Galactic ridge
- angular resolutions: 0.4° (10°)
- completion date: 2010

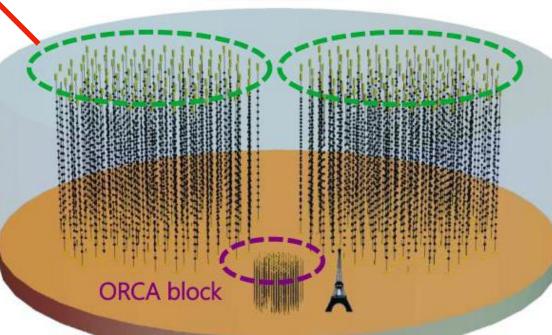




- angular resolutions: 0.1° (1.5°)
- completion date: ~2028



- ~2500m below ice
- 86 strings
- 5160 DOMs



ARCA blocks

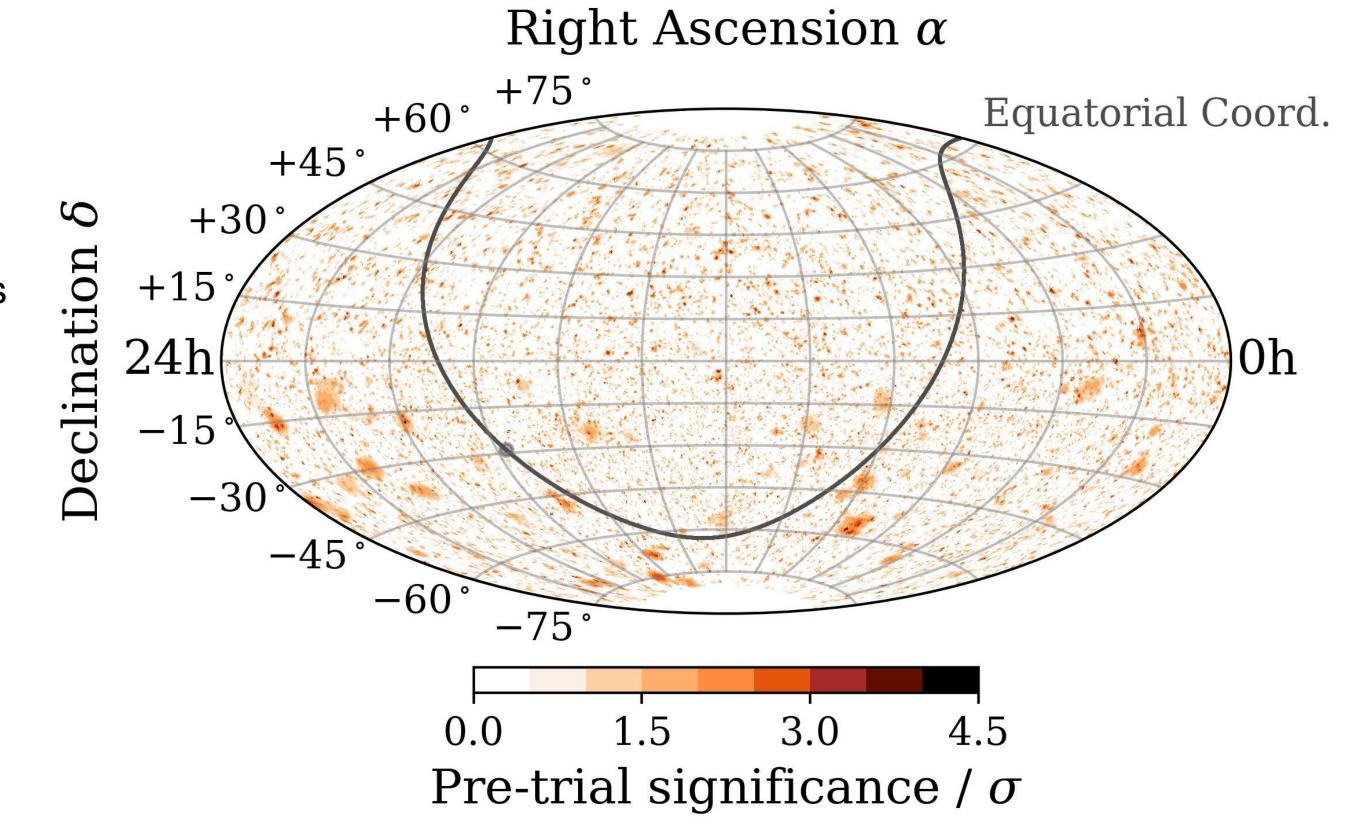
What did we see?!

- ~3500m b.s.l
- 345 strings
- 6210 DOMs
- 192510 PMTs

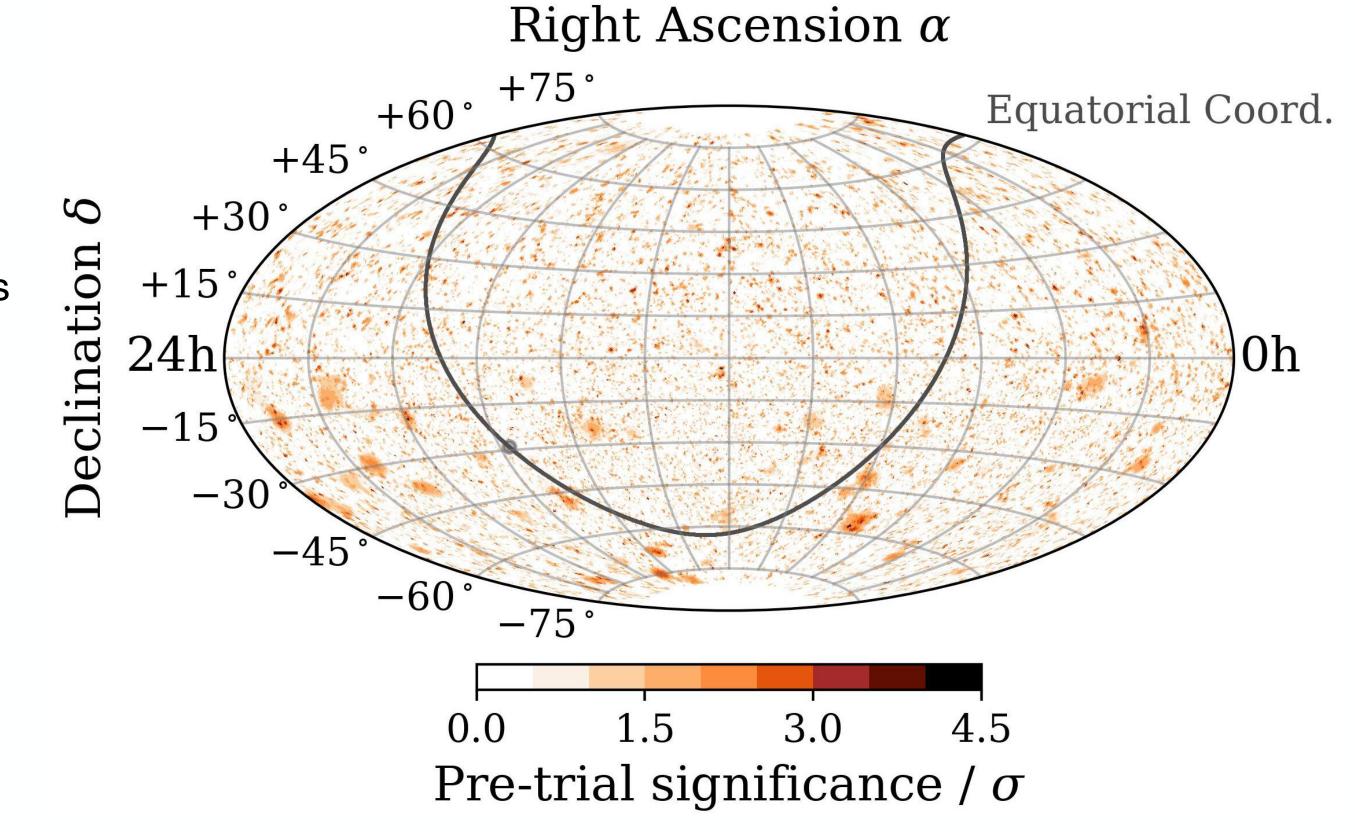


A diffuse flux from TeV to PeV

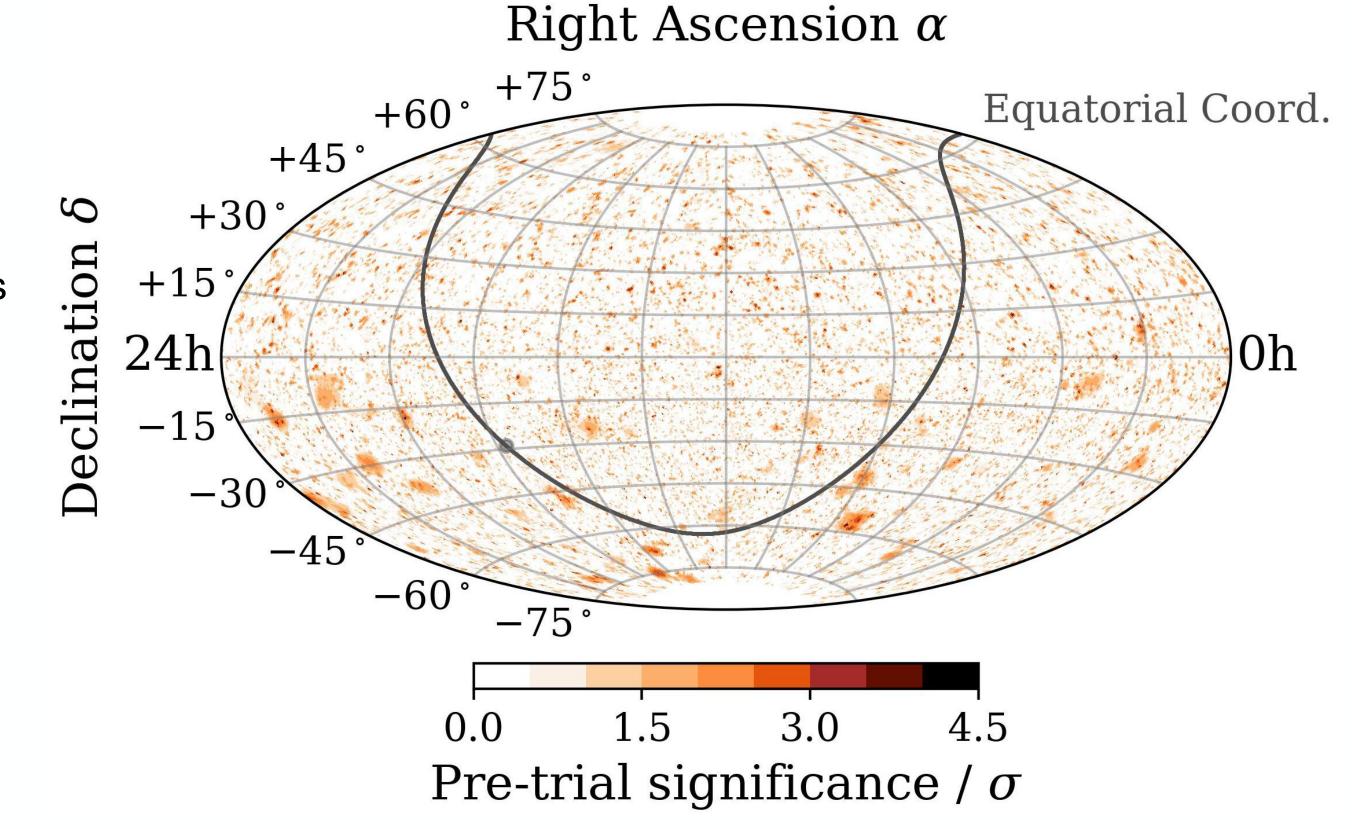
- Detection of a diffuse astrophysical flux of high-energy neutrinos
 - \rightarrow now confirmed >7 σ (since 2013)



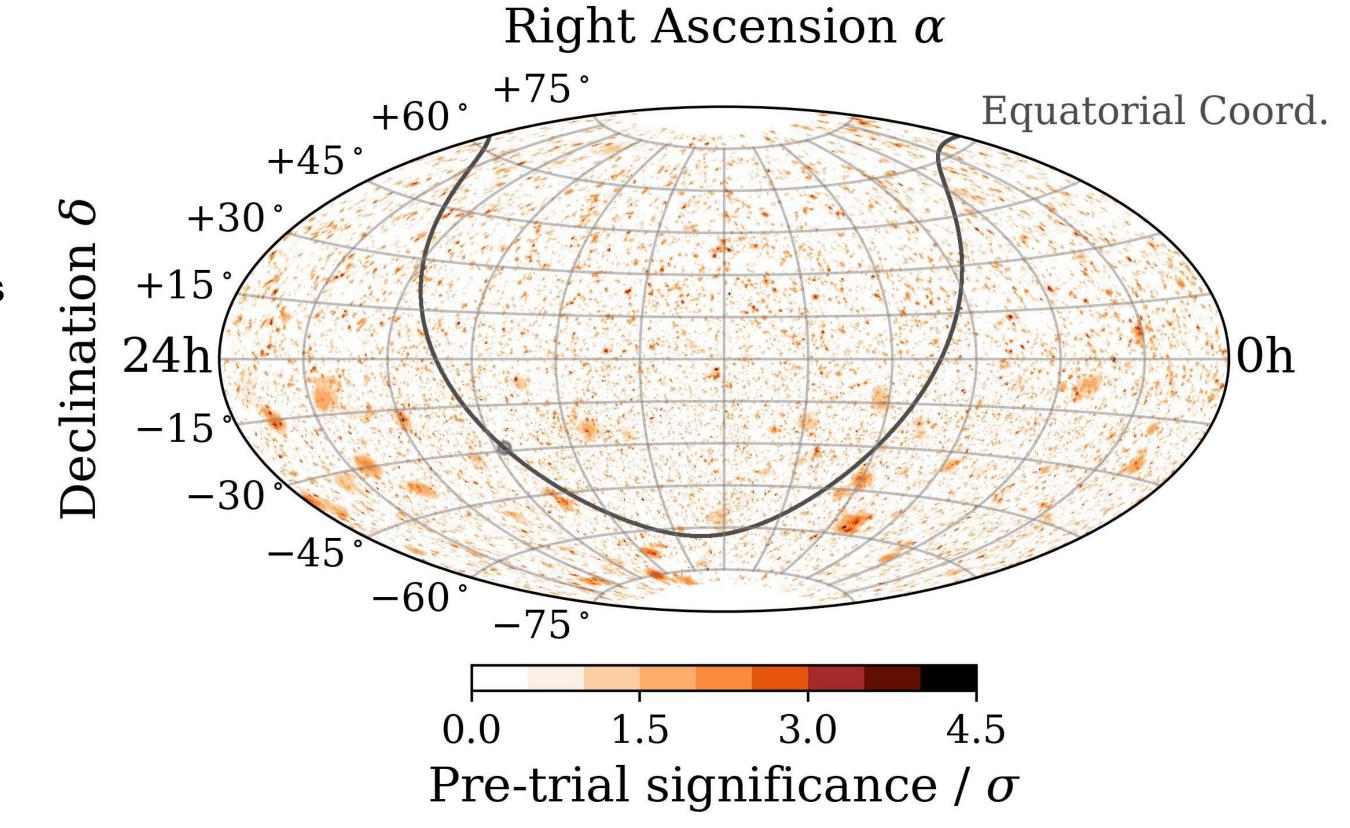
- Detection of a diffuse astrophysical flux of high-energy neutrinos
 - \rightarrow now confirmed >7 σ (since 2013)
- Isotropic → extragalactic origin (mostly)



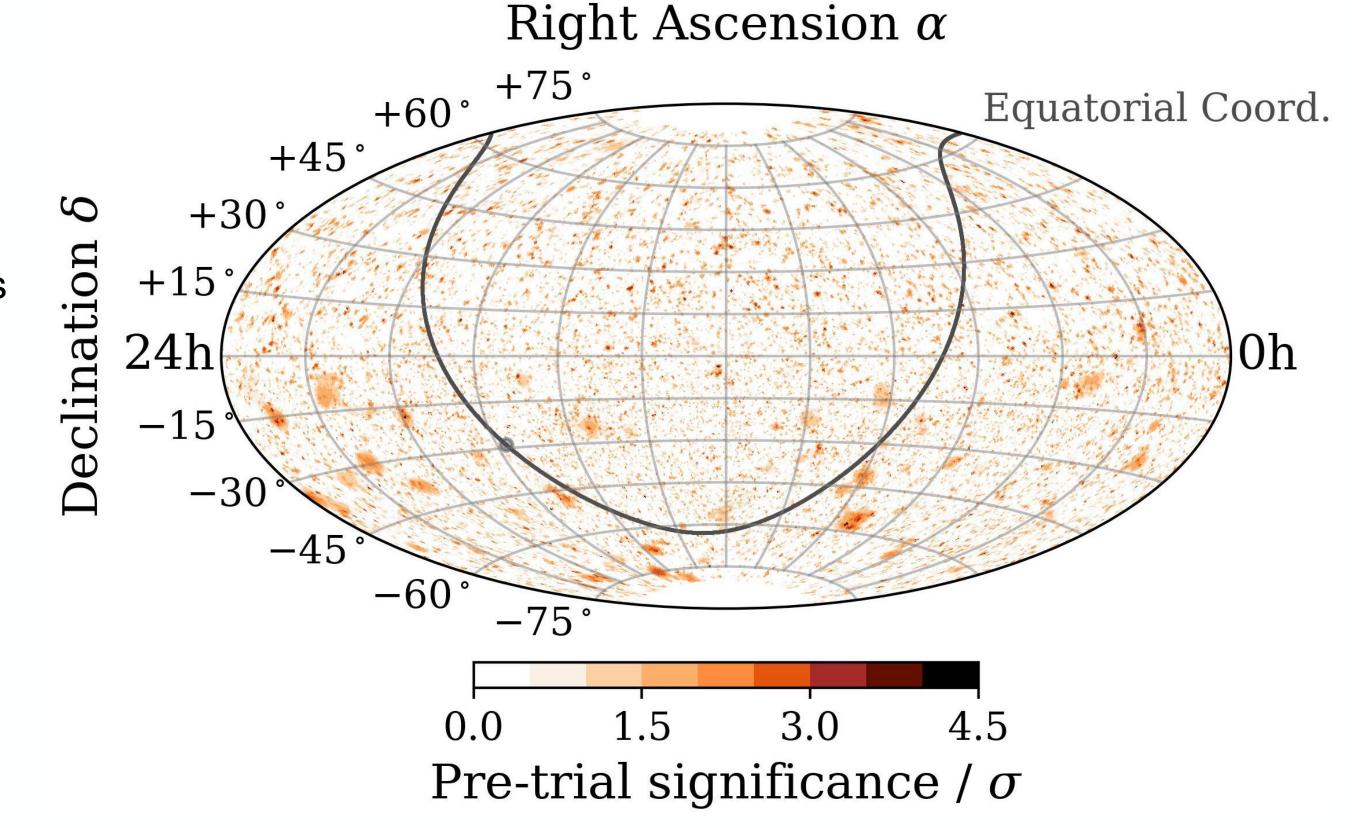
- Detection of a diffuse astrophysical flux of high-energy neutrinos
 → now confirmed >7 σ (since 2013)
- Isotropic → extragalactic origin (mostly)
- Produced by a population of un-resolved sources



- Detection of a diffuse astrophysical flux of high-energy neutrinos
 → now confirmed >7 σ (since 2013)
- Isotropic → extragalactic origin (mostly)
- Produced by a population of un-resolved sources
- Blazars and GRBs excluded as the main sources



- Detection of a diffuse astrophysical flux of high-energy neutrinos
 → now confirmed >7 σ (since 2013)
- Isotropic → extragalactic origin (mostly)
- Produced by a population of un-resolved sources
- Blazars and GRBs excluded as the main sources



Some directions are slightly more significant → nearly resolved sources?

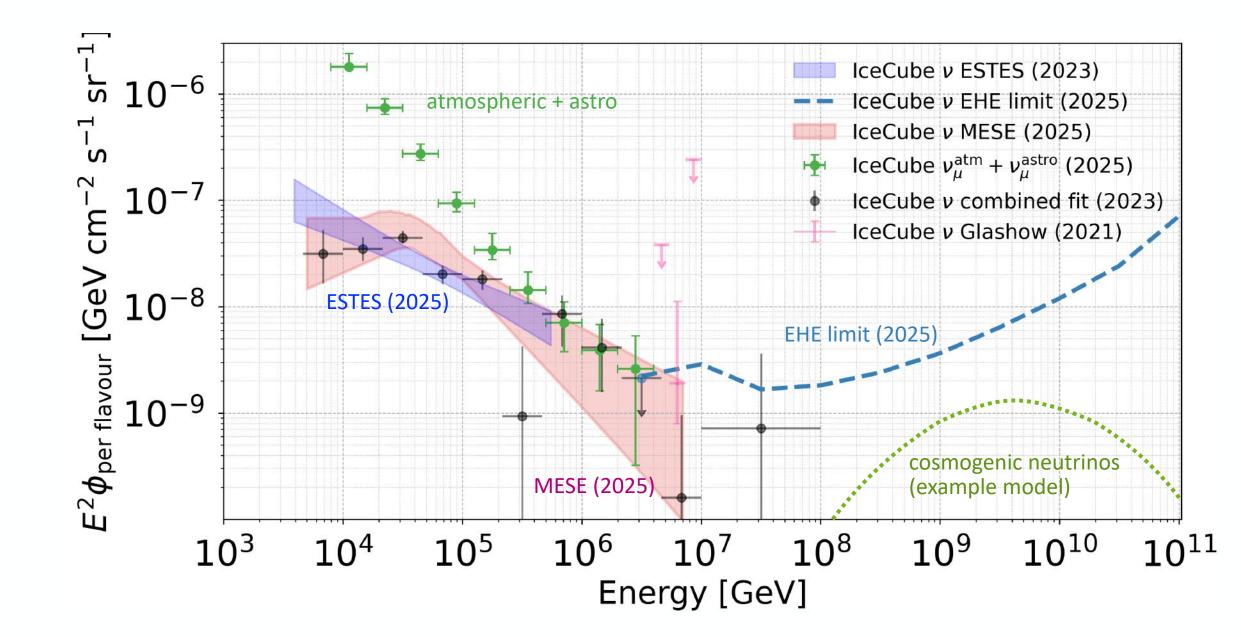
2025

2025

More statistic and new selection of events (ESTE → MESE)

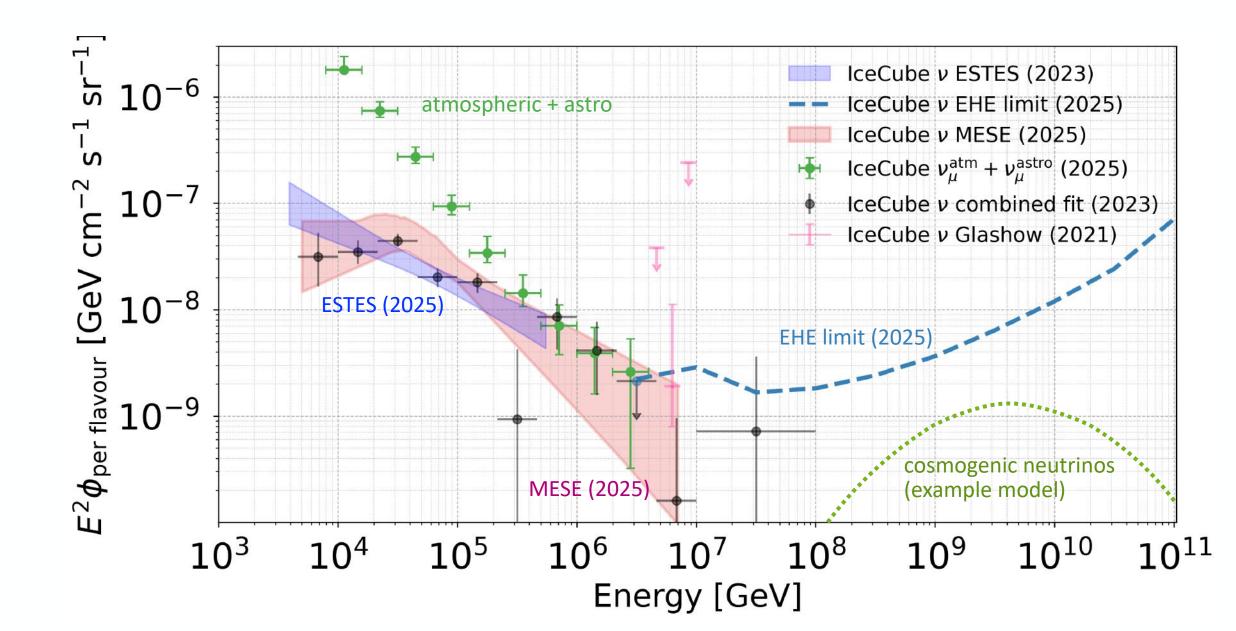
2025

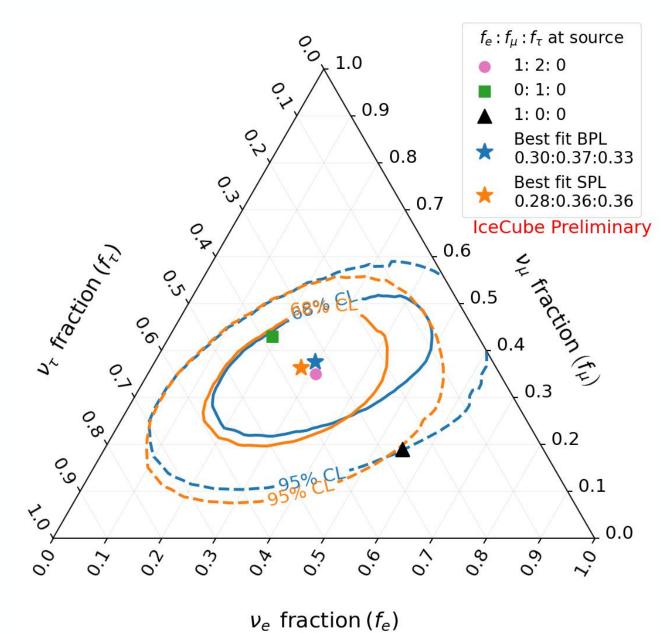
- More statistic and new selection of events (ESTE → MESE)
- Hardening below ~30 TeV
 - \rightarrow Best fitted by BPL at 4.7 σ or log parabola at 4.2 σ
 - → Population or production effect?



2025

- More statistic and new selection of events (ESTE → MESE)
- Hardening below ~30 TeV
 - \rightarrow Best fitted by BPL at 4.7 σ or log parabola at 4.2 σ
 - → Population or production effect?
- Composition:
 - Pion beam and muon damp mechanisms → compatible (68% C.L.)
 - However pion beam favoured over muon damp
 - Neutron decay → excluded (94.8% C. L.)

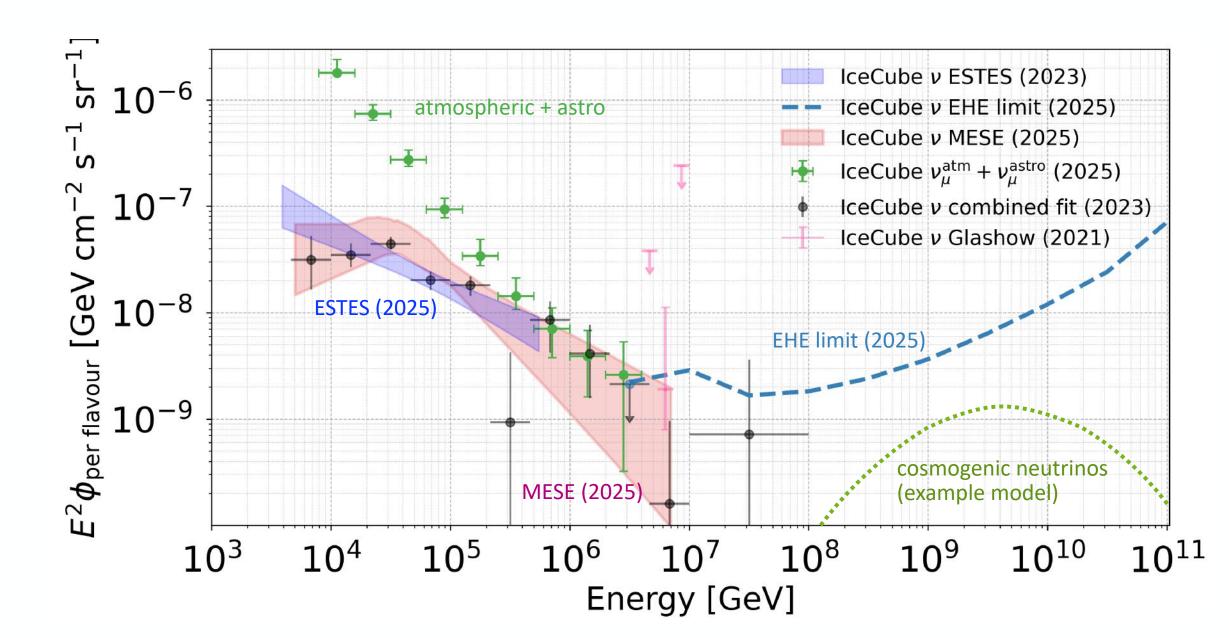


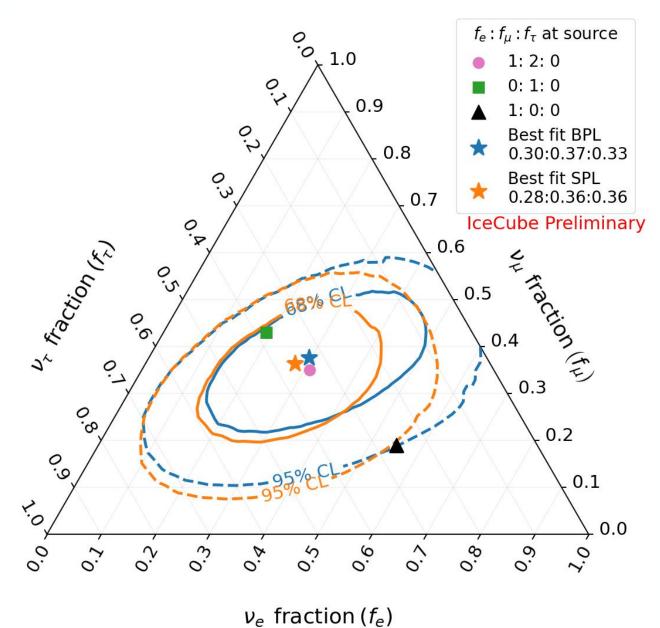


2025

- More statistic and new selection of events (ESTE → MESE)
- Hardening below ~30 TeV
 - \rightarrow Best fitted by BPL at 4.7 σ or log parabola at 4.2 σ
 - → Population or production effect?
- Composition:
 - Pion beam and muon damp mechanisms → compatible (68% C.L.)
 - However pion beam favoured over muon damp
 - Neutron decay → excluded (94.8% C. L.)

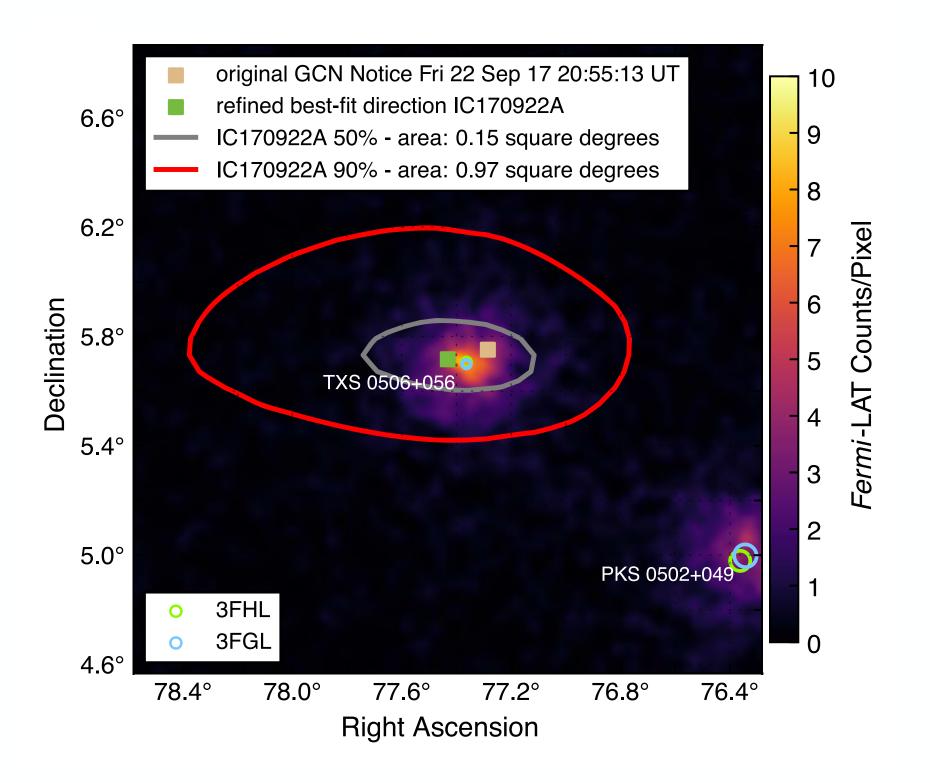
The sources of the diffuse neutrino flux remain unknown





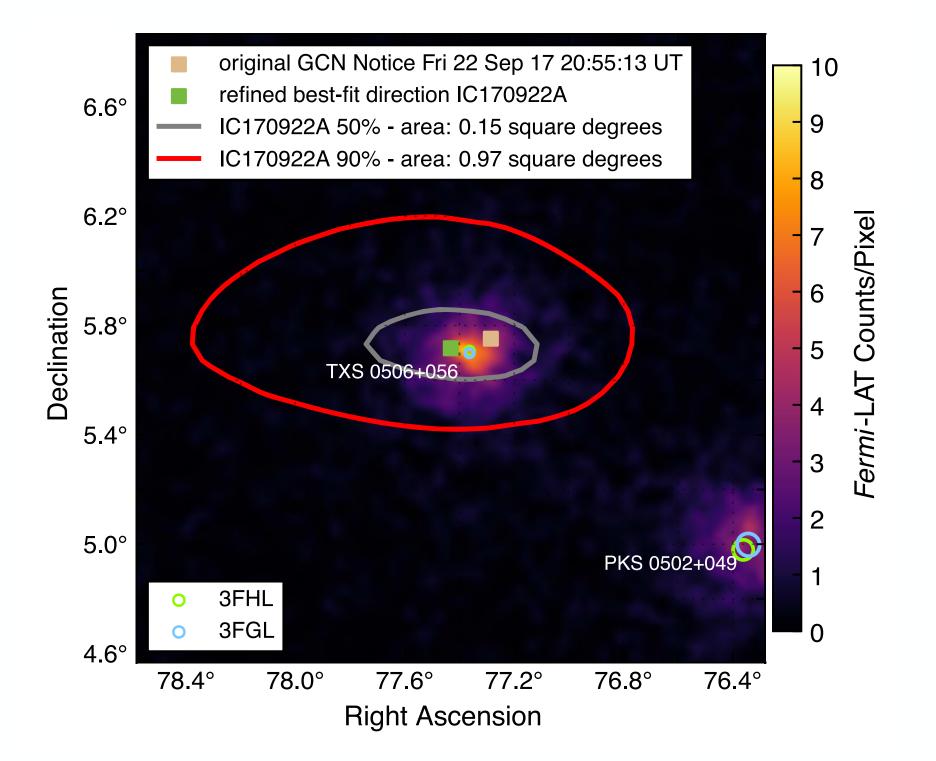
2018

• Detection of a neutrino event in coincidence with a blazar flare



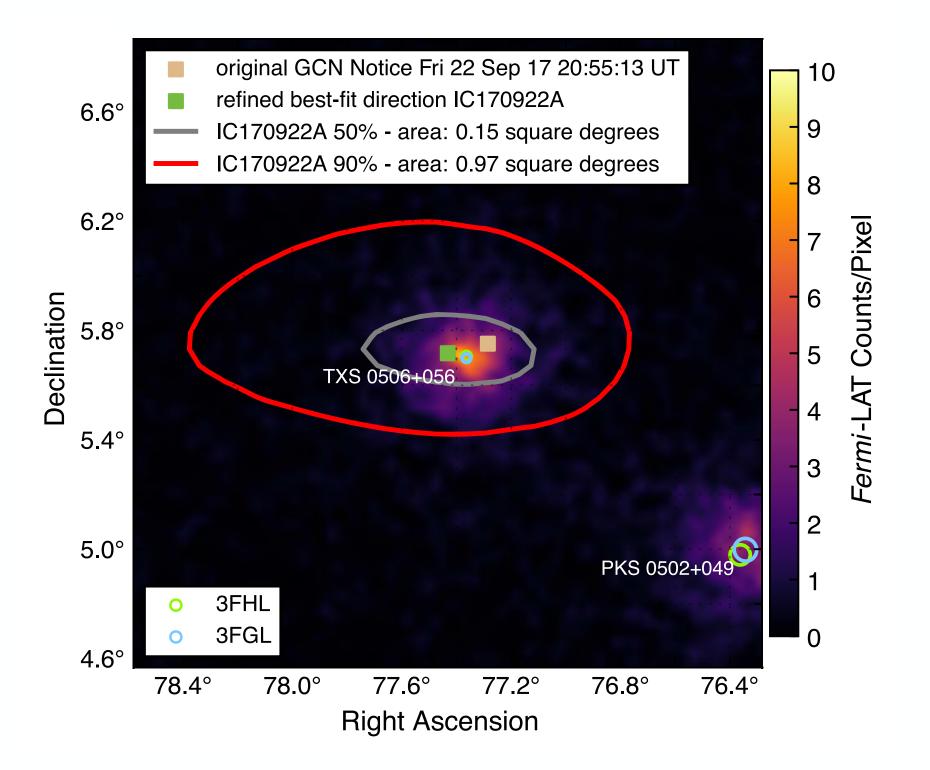
2018

- Detection of a neutrino event in coincidence with a blazar flare
- Exact detection is debated \rightarrow significance close to 3 σ



2018

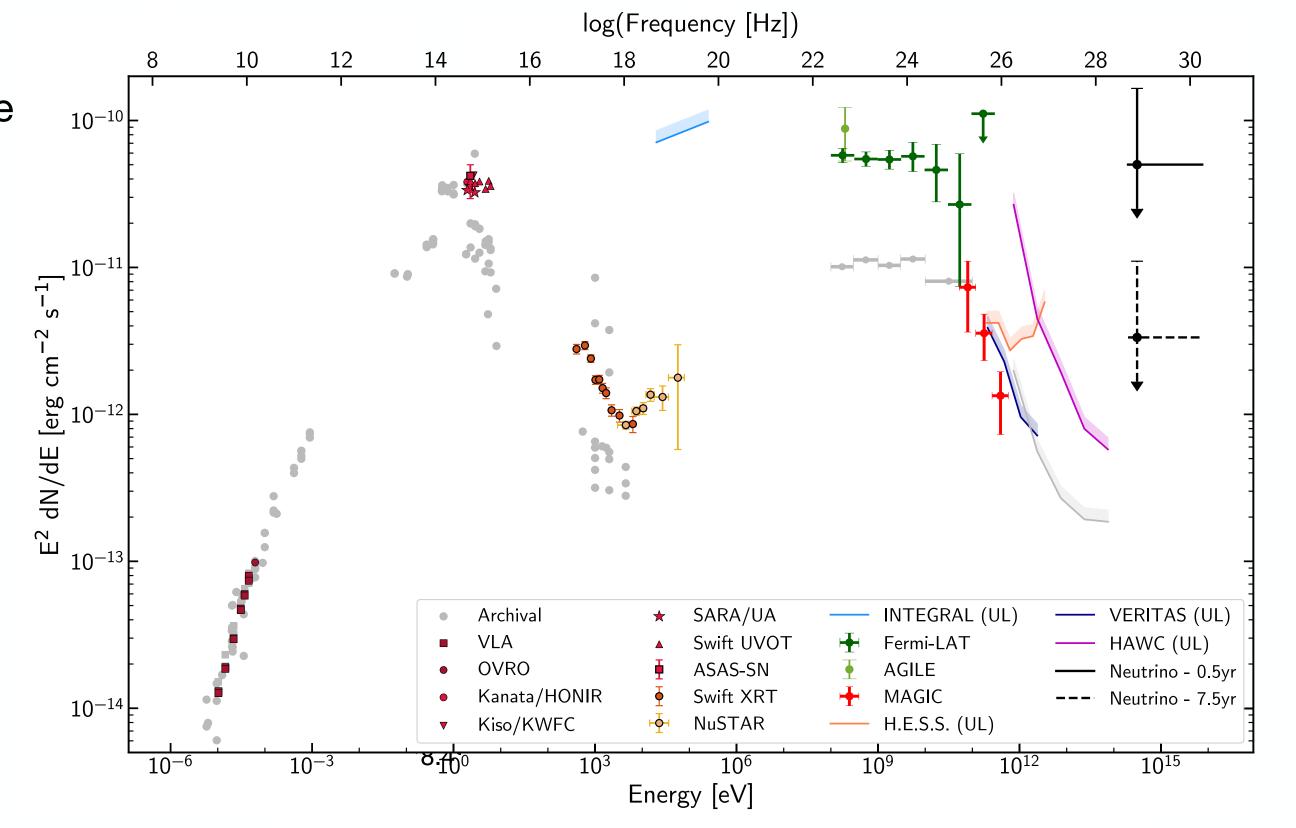
- Detection of a neutrino event in coincidence with a blazar flare
- Exact detection is debated \rightarrow significance close to 3 σ
- First connection between neutrinos and AGNs



8

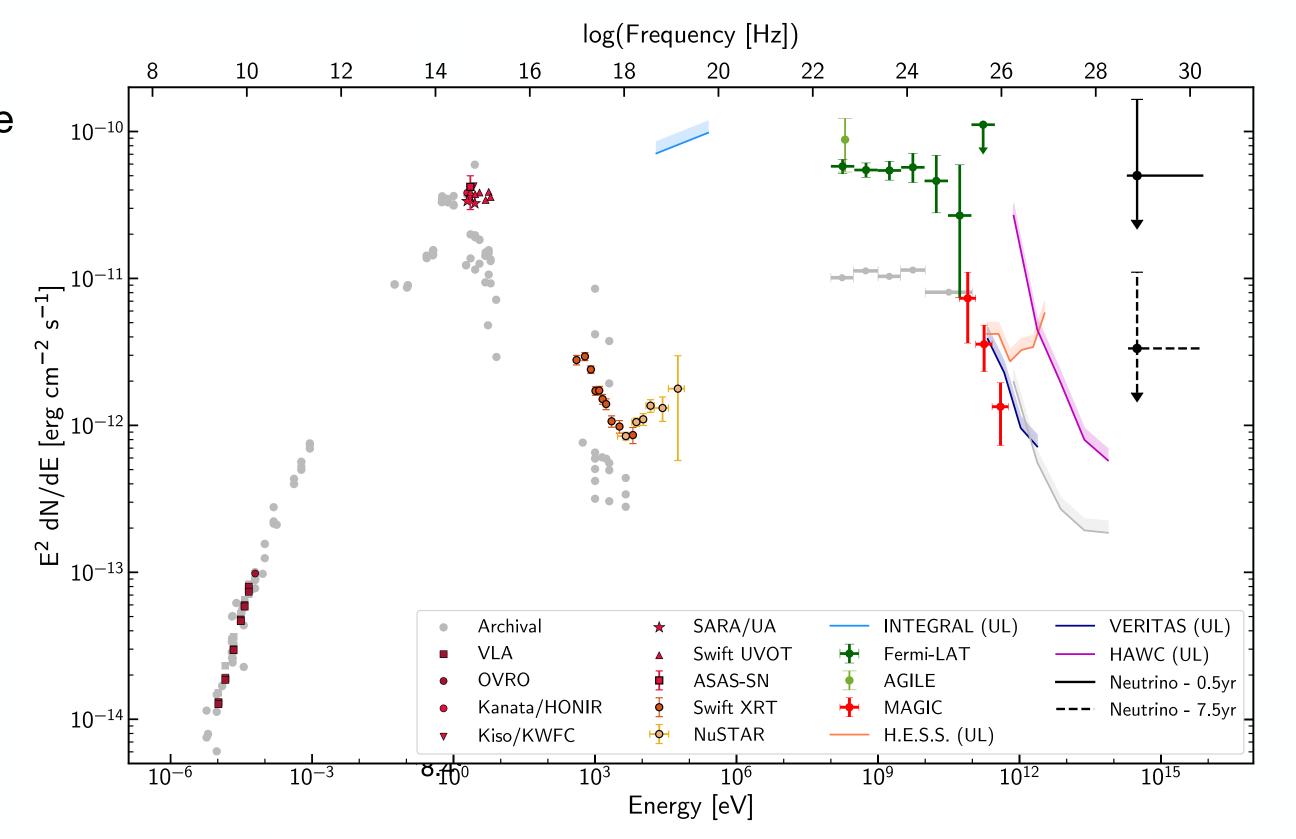
2018

- Detection of a neutrino event in coincidence with a blazar flare
- Exact detection is debated → significance close to 3 σ
- First connection between neutrinos and AGNs
- Gamma-ray flare → jet related neutrino production ?



2018

- Detection of a neutrino event in coincidence with a blazar flare
- Exact detection is debated \rightarrow significance close to 3 σ
- First connection between neutrinos and AGNs
- Gamma-ray flare → jet related neutrino production ?

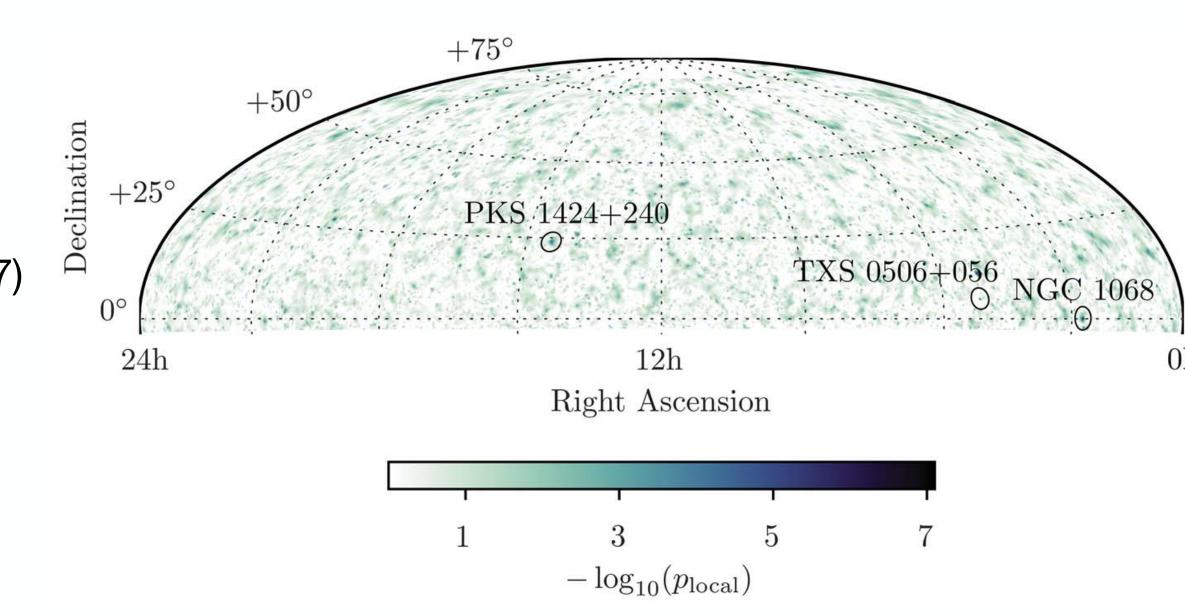


8

Wether or not the neutrino event is associated to the gamma-ray events it has shown the importance of multi-messenger alerts

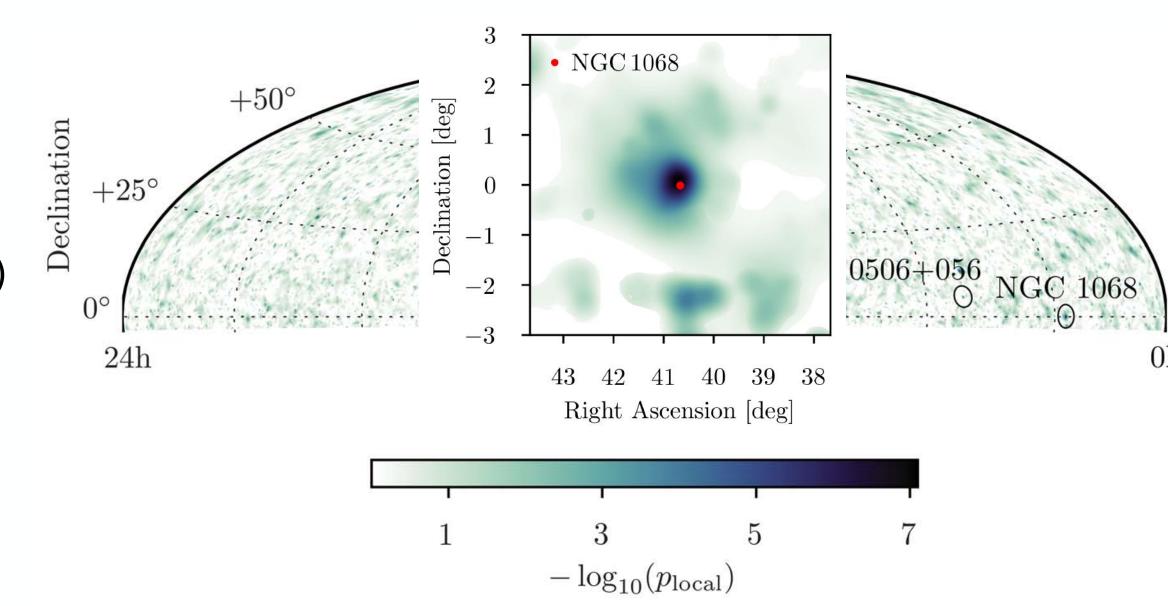
2022

- Discovery of a neutrino excess in the direction of the AGN NGC1068 (M77)
 - \rightarrow First source with a significance > 3 σ (4.2 σ)



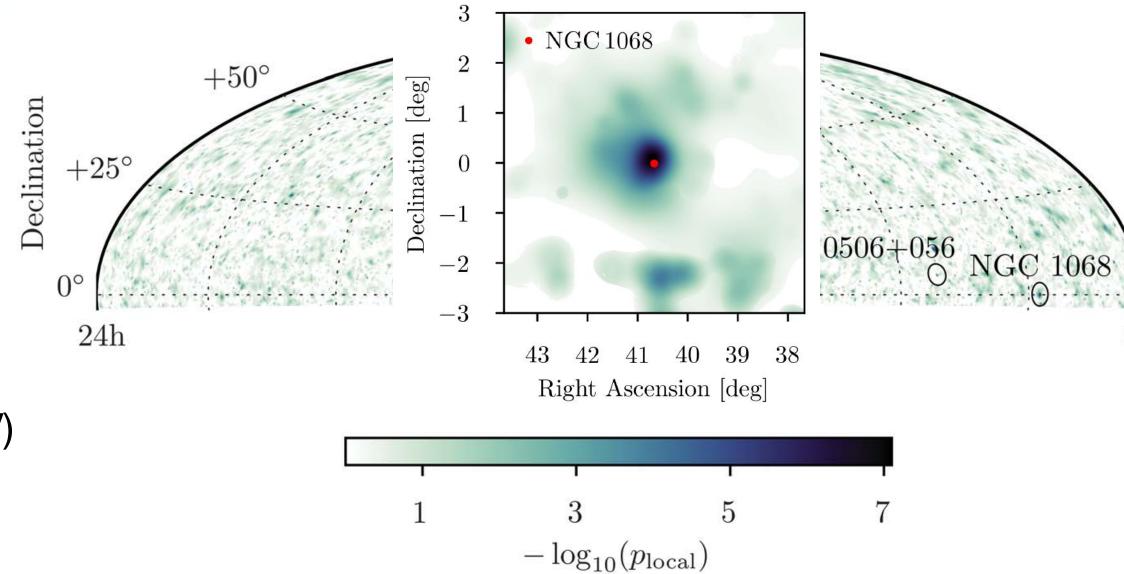
2022

- Discovery of a neutrino excess in the direction of the AGN NGC1068 (M77)
 - \rightarrow First source with a significance > 3 σ (4.2 σ)



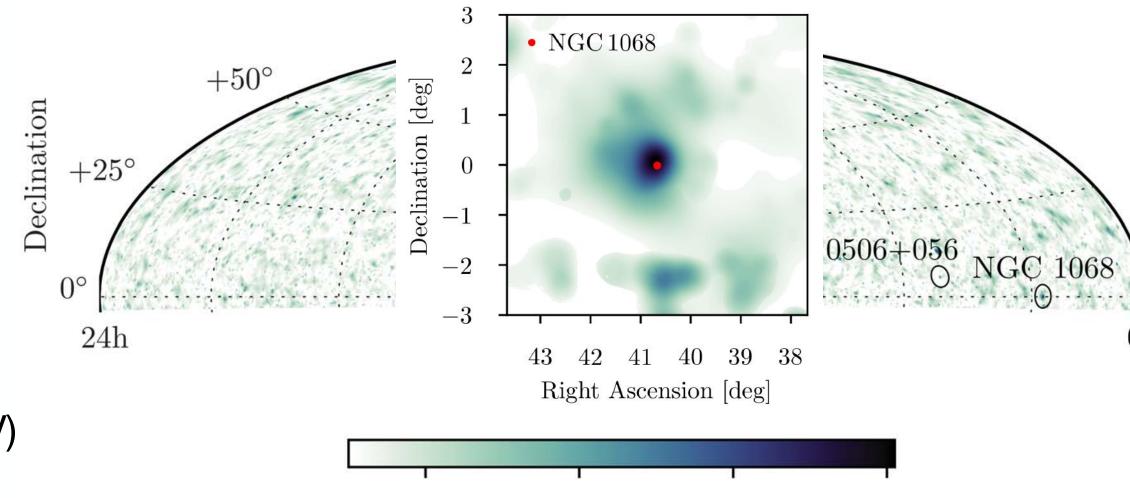
2022

- Discovery of a neutrino excess in the direction of the AGN NGC1068 (M77)
 - \rightarrow First source with a significance > 3 σ (4.2 σ)
- Neutrino flux larger than expected from gamma-ray fluxes (100MeV-100GeV)
 - → Suggests an opaque environment around the source



2022

- Discovery of a neutrino excess in the direction of the AGN NGC1068 (M77)
 - \rightarrow First source with a significance > 3 σ (4.2 σ)
- Neutrino flux larger than expected from gamma-ray fluxes (100MeV-100GeV)
 - → Suggests an opaque environment around the source
- What causes the opacity (BH, absorption, ...) is still unclear

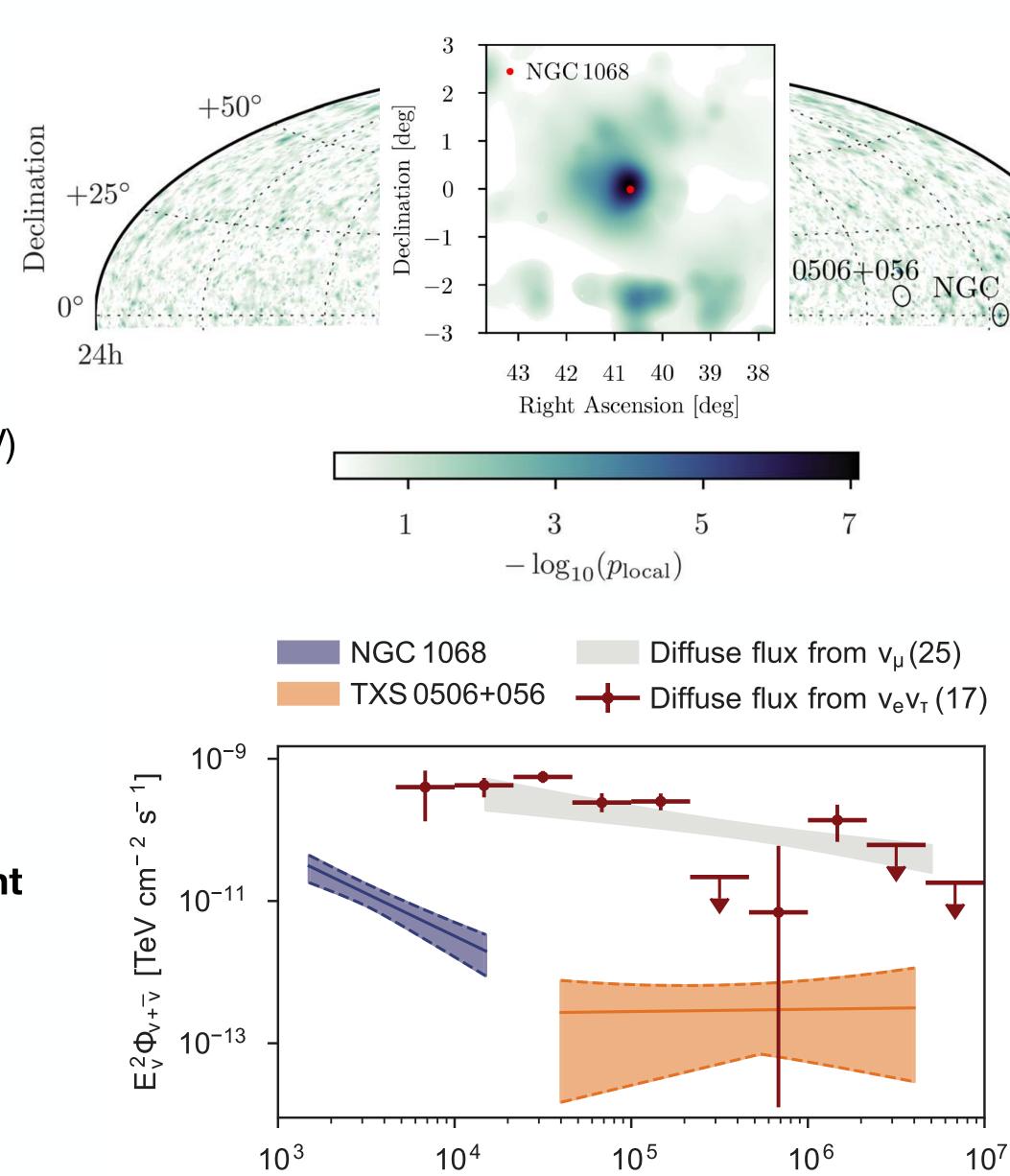


 $-\log_{10}(p_{\text{local}})$

2022

- Discovery of a neutrino excess in the direction of the AGN NGC1068 (M77)
 - \rightarrow First source with a significance > 3 σ (4.2 σ)
- Neutrino flux larger than expected from gamma-ray fluxes (100MeV-100GeV)
 - → Suggests an opaque environment around the source
- What causes the opacity (BH, absorption, ...) is still unclear

The first clearly identified source is a Seyfert II galaxy with very different properties than TXS0506+056

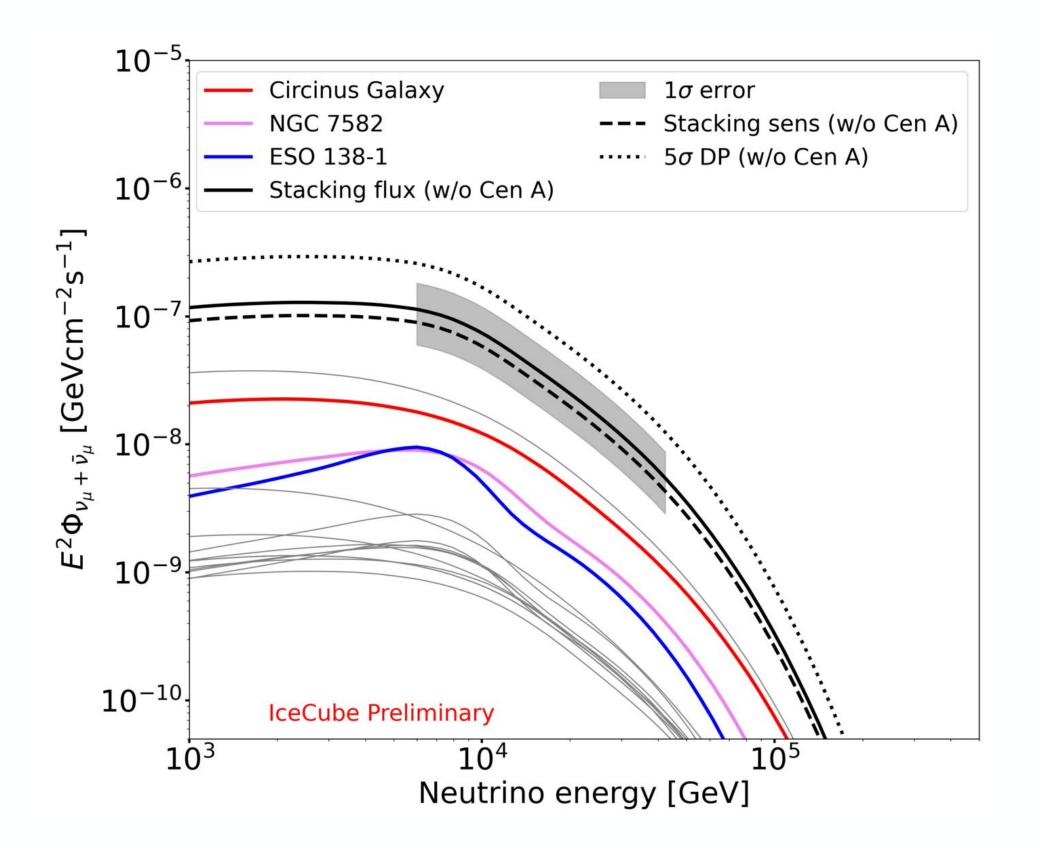


E_v [GeV]

2025

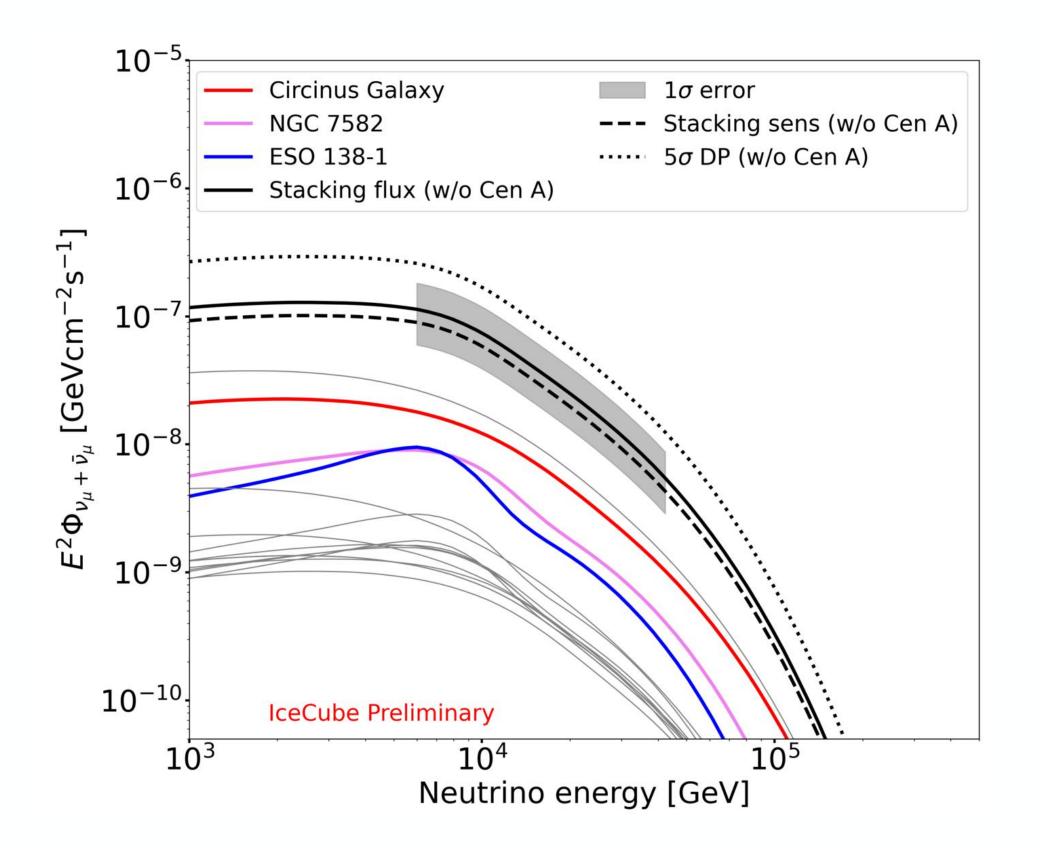
2025

Stacking analysis on Seyfert galaxy catalogs → ~ 3 σ



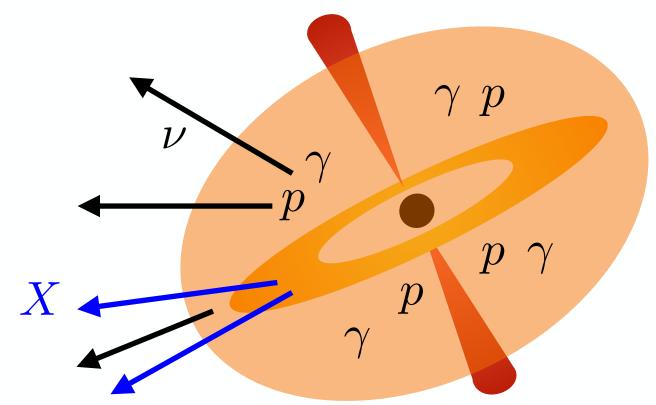
2025

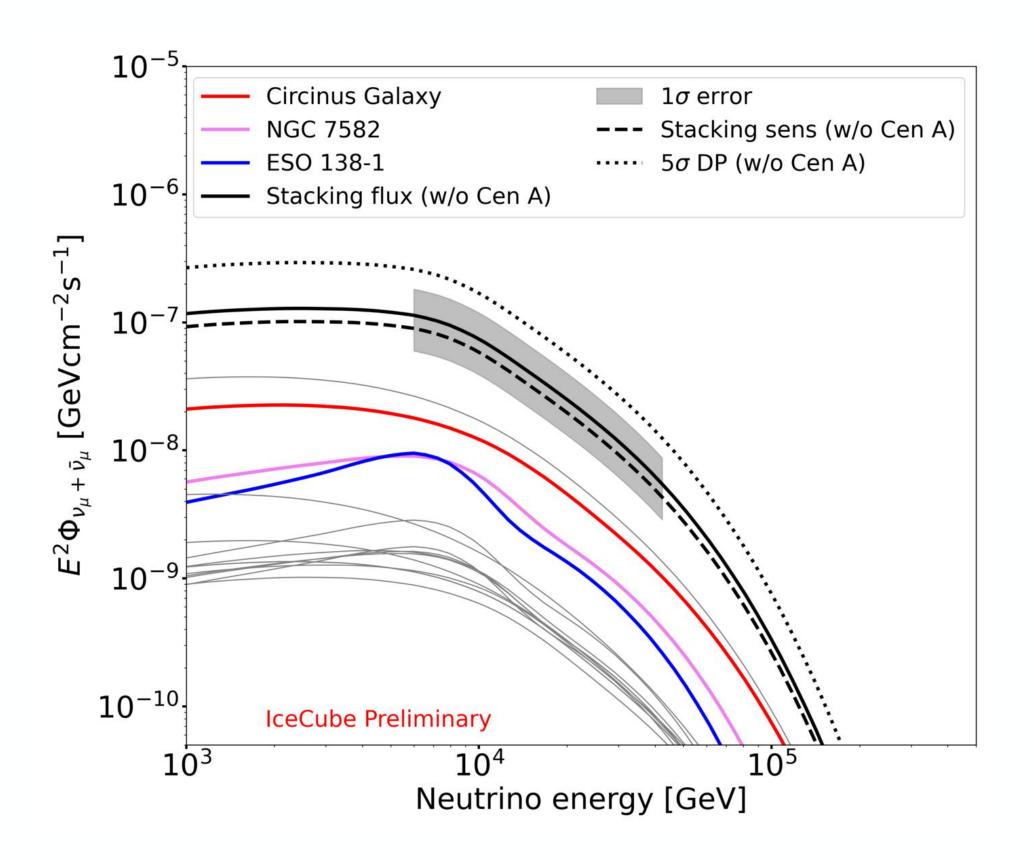
- Stacking analysis on Seyfert galaxy catalogs → ~ 3 σ
- X-ray bright Seyfert galaxies are among the most significant sources



2025

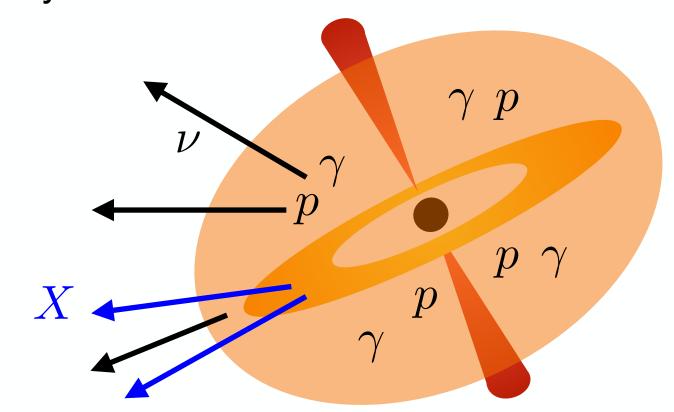
- Stacking analysis on Seyfert galaxy catalogs → ~ 3 σ
- X-ray bright Seyfert galaxies are among the most significant sources
- Hints of disk-corona interplay in neutrino emissions
 - → Xray emission from accretion disk and interaction of nuclei in corona where gamma rays are absorbed

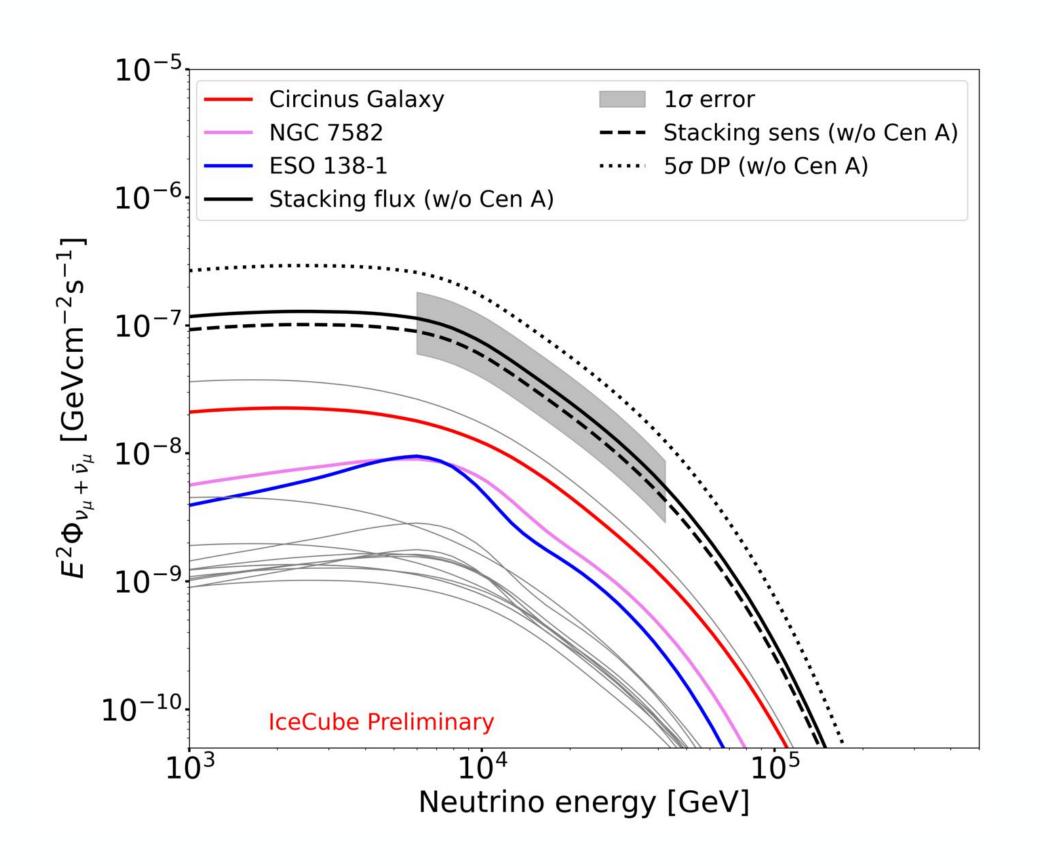




2025

- Stacking analysis on Seyfert galaxy catalogs → ~ 3 σ
- X-ray bright Seyfert galaxies are among the most significant sources
- Hints of disk-corona interplay in neutrino emissions
 - → Xray emission from accretion disk and interaction of nuclei in corona where gamma rays are absorbed

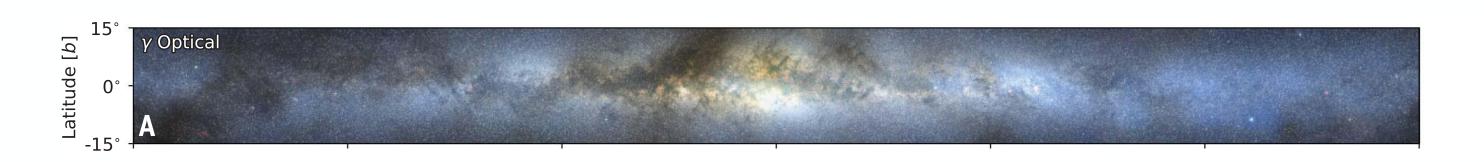


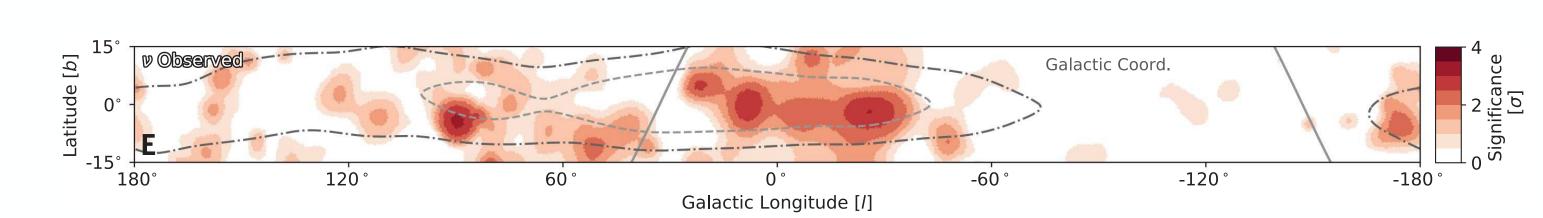


Obscured AGNs seem to play a significant role in the origin of the diffuse HE neutrino flux

2023

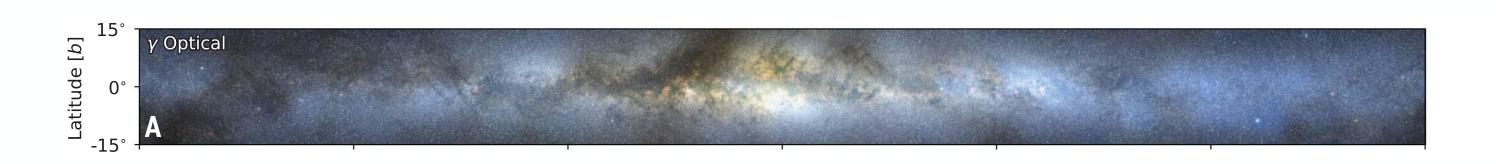
• Identification of a neutrino contribution from the Galactic Plane ($\sim 4.5 \sigma$)!

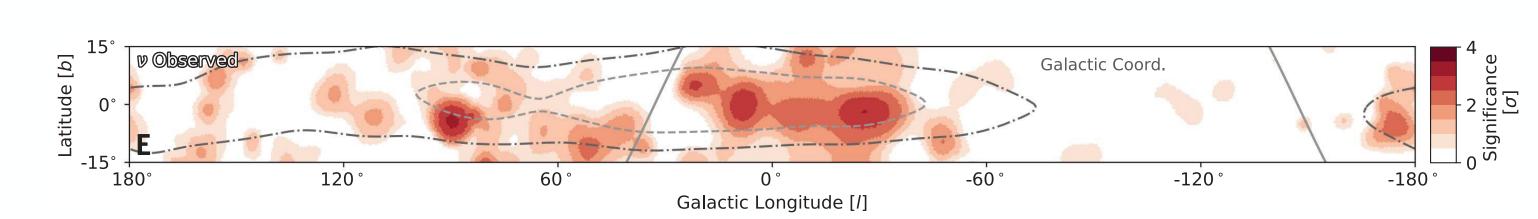




2023

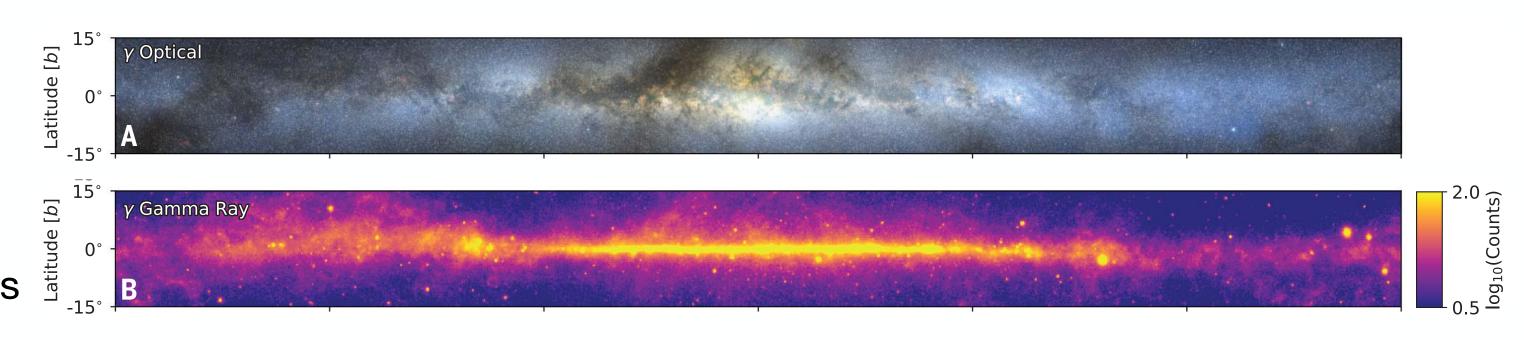
- Identification of a neutrino contribution from the Galactic Plane ($\sim 4.5 \sigma$)!
- Southern sky drowned in atmospheric muons for IceCube → new ML analysis required

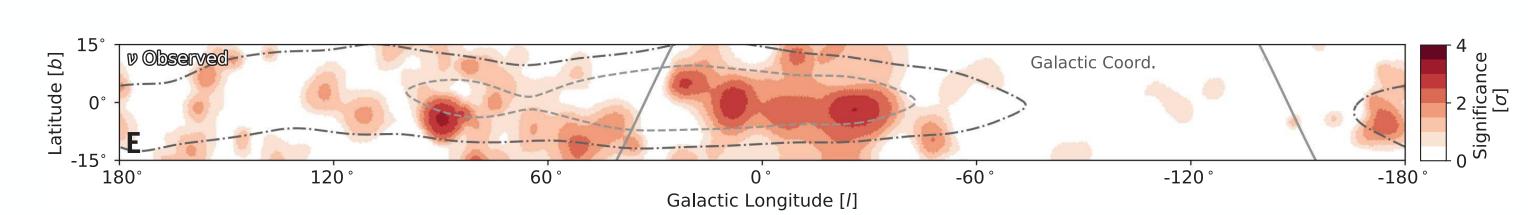




2023

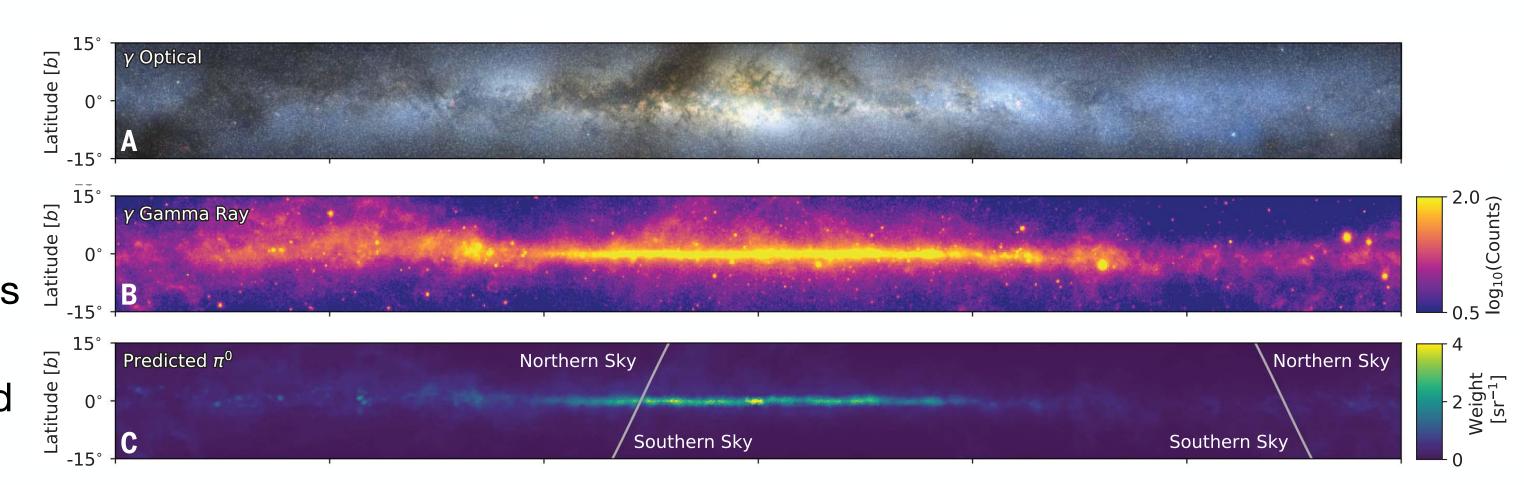
- Identification of a neutrino contribution from the Galactic Plane ($\sim 4.5 \sigma$)!
- Southern sky drowned in atmospheric muons for IceCube → new ML analysis required
- Neutrinos expected from Fermi and HESS observations

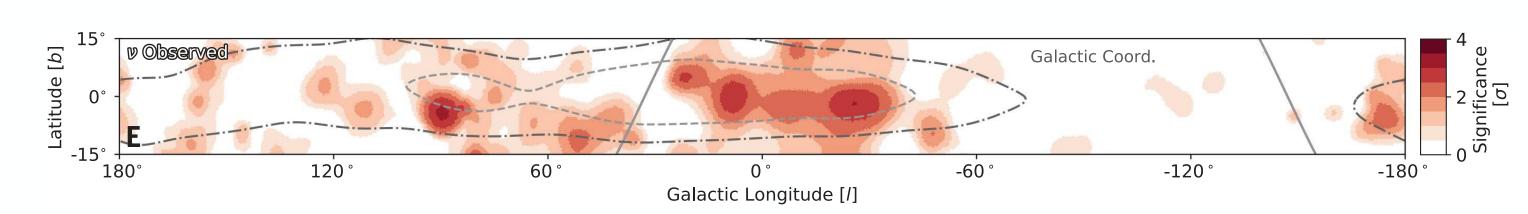




2023

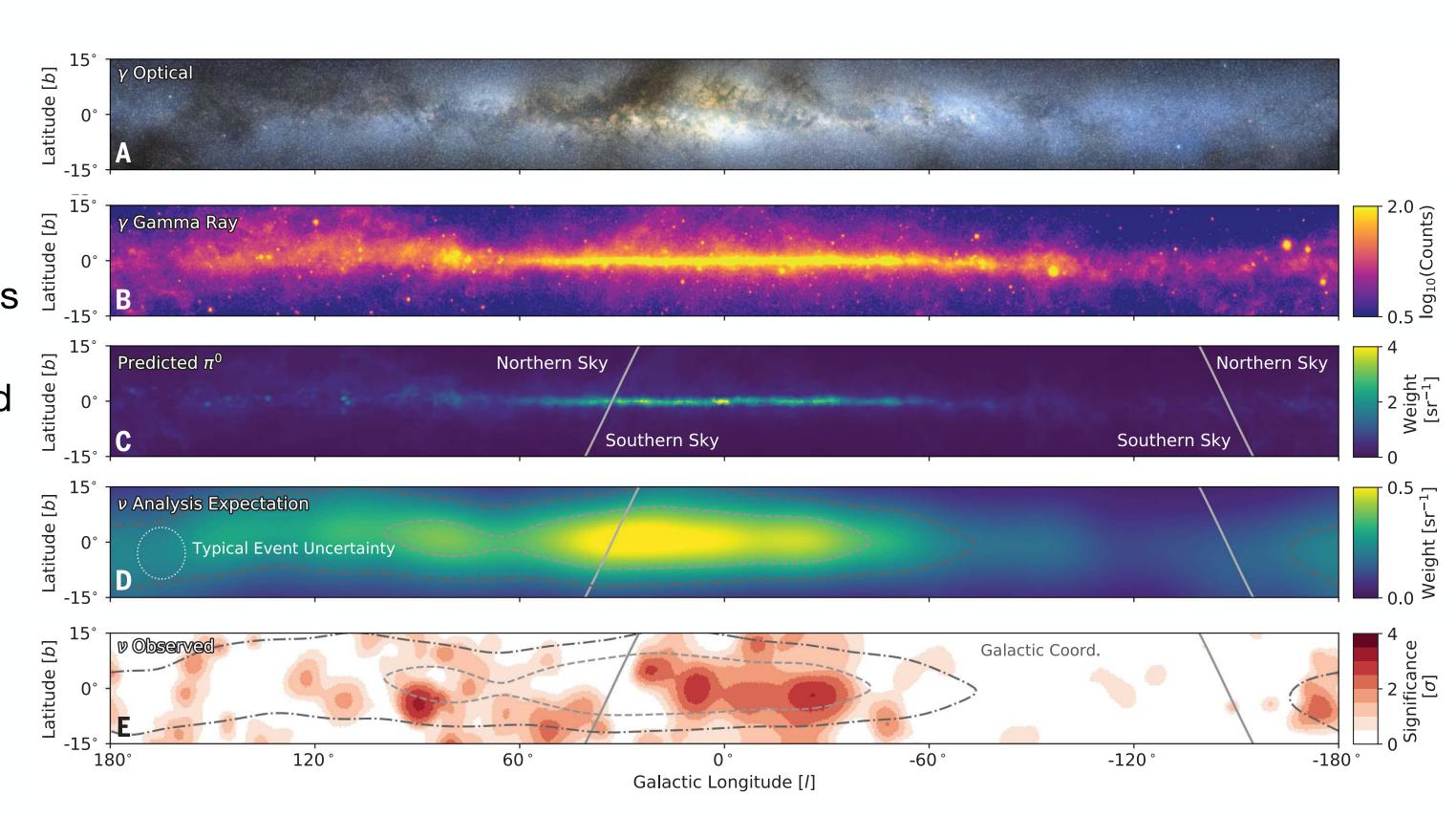
- Identification of a neutrino contribution from the Galactic Plane ($\sim 4.5 \sigma$)!
- Southern sky drowned in atmospheric muons for IceCube → new ML analysis required
- Neutrinos expected from Fermi and HESS observations
- From gamma rays π^0 are inferred $\to \pi^{\pm}$ are extrapolated





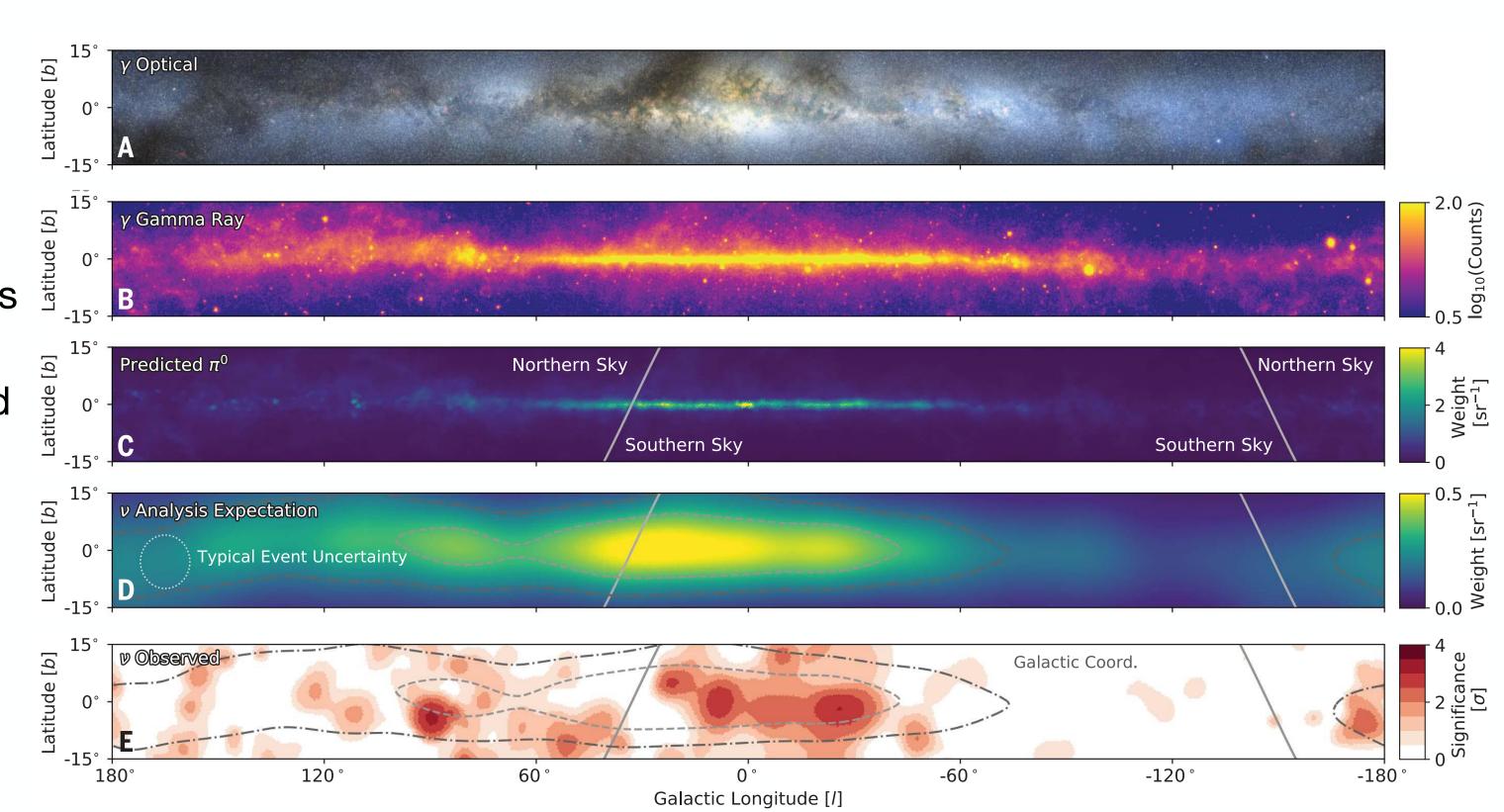
2023

- Identification of a neutrino contribution from the Galactic Plane ($\sim 4.5 \sigma$)!
- Southern sky drowned in atmospheric muons for IceCube → new ML analysis required
- Neutrinos expected from Fermi and HESS observations
- From gamma rays π^0 are inferred $\to \pi^{\pm}$ are extrapolated
- Expected neutrino flux derived + instrumental effects



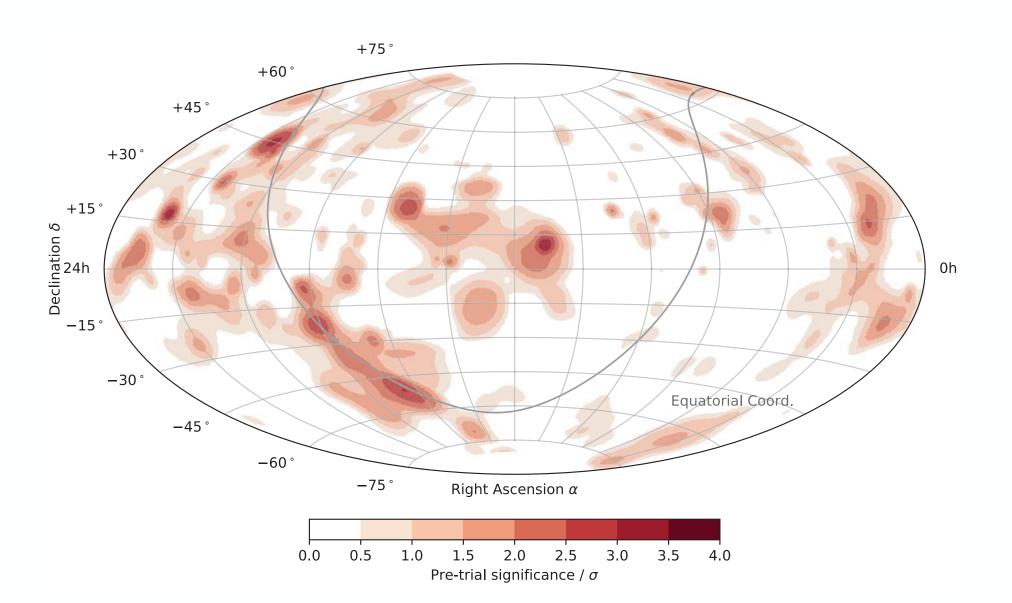
2023

- Identification of a neutrino contribution from the Galactic Plane ($\sim 4.5 \sigma$)!
- Southern sky drowned in atmospheric muons for IceCube → new ML analysis required
- Neutrinos expected from Fermi and HESS observations
- From gamma rays π^0 are inferred $\to \pi^{\pm}$ are extrapolated
- Expected neutrino flux derived + instrumental effects



Galactic TeV neutrinos → 6-13% of all diffuse flux above 30 TeV

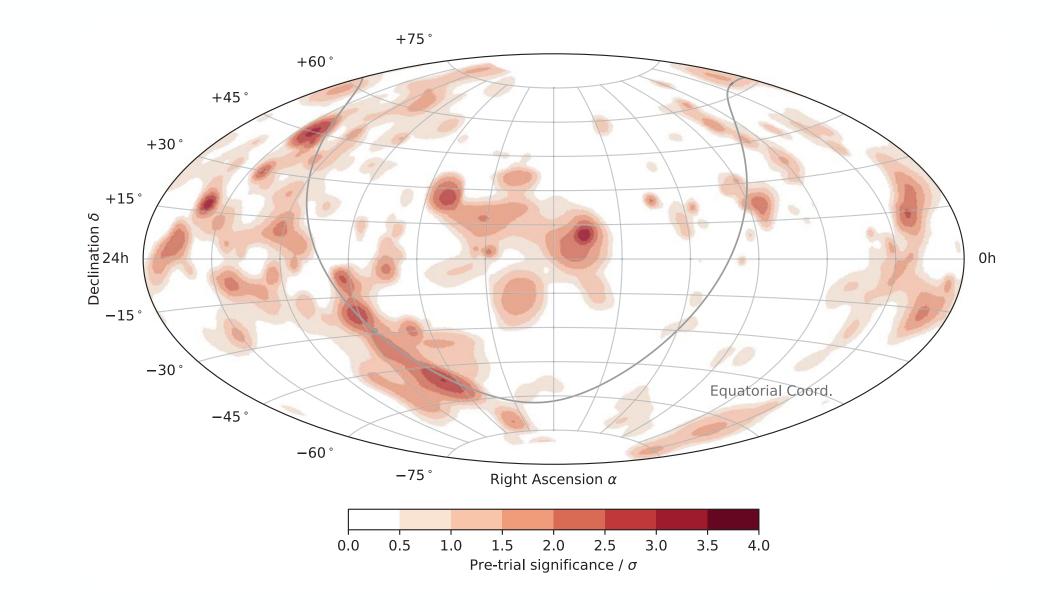
2025

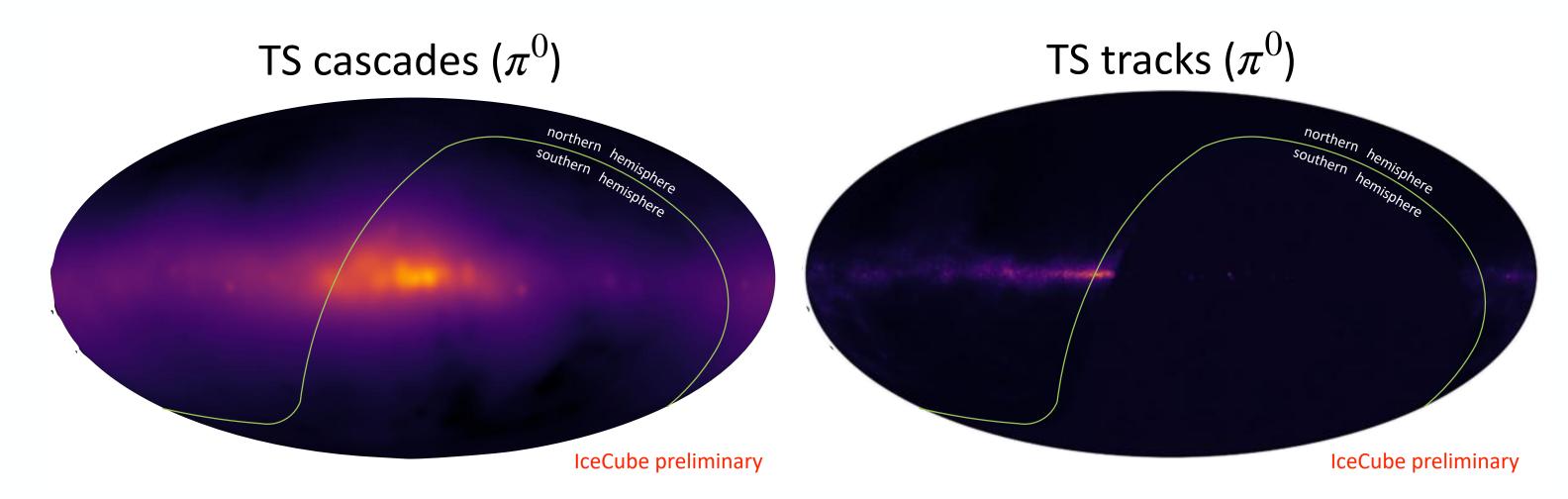


ICRC2025(1130, 1219)

2025

- 2.5 yr more data
- Addition of muon tracks (starting events + up-going)
 - \rightarrow detection $> 5 \sigma$



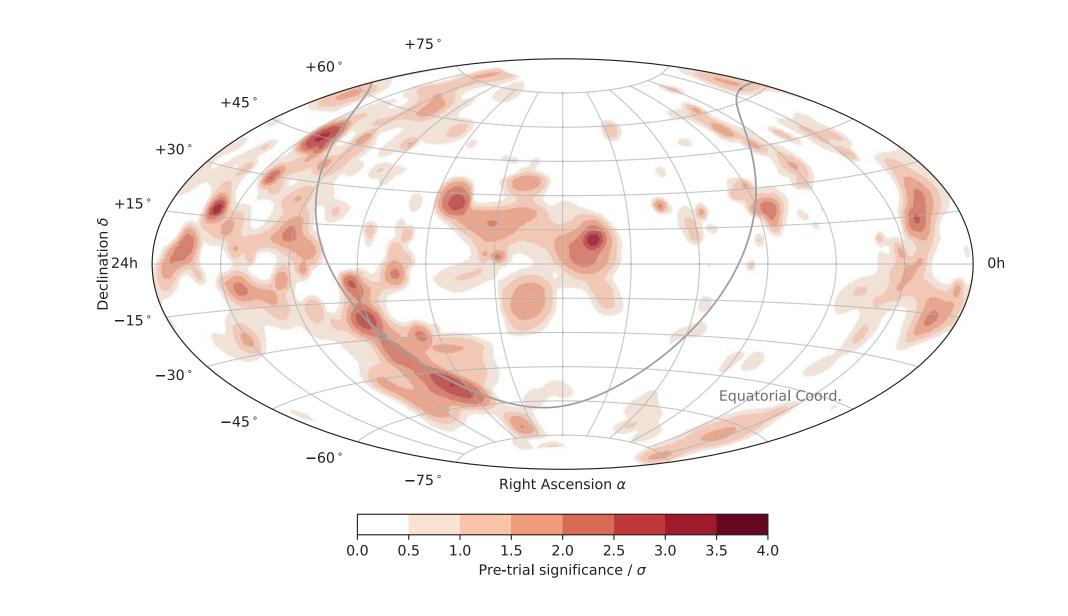


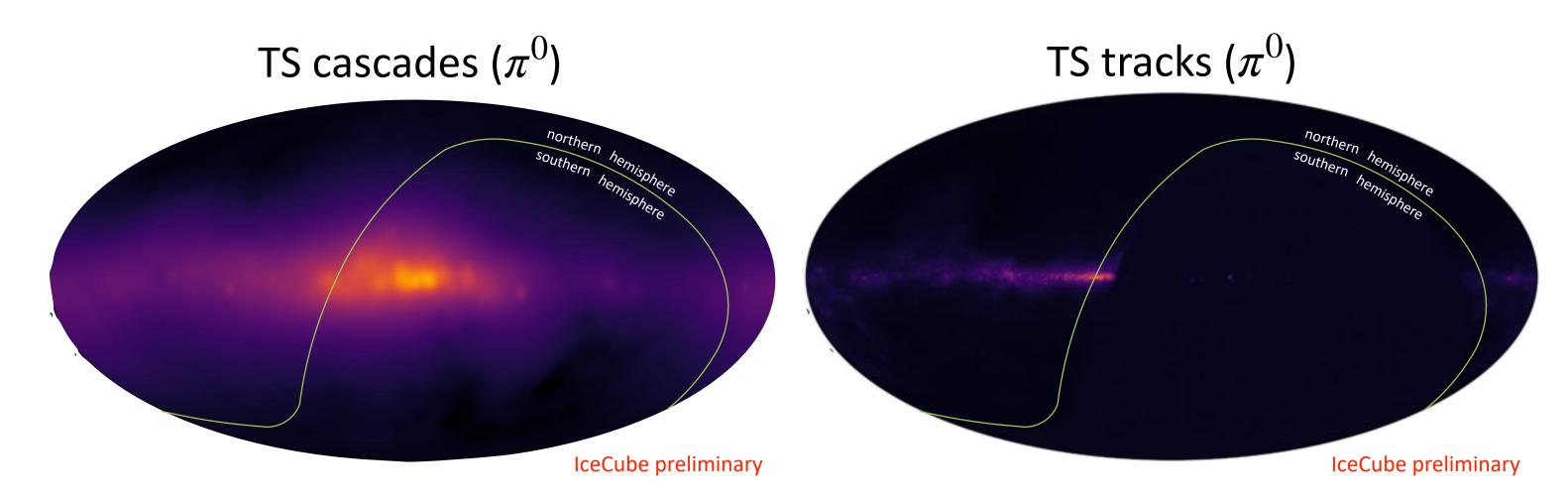
ICRC2025(1130, 1219)

A "local" source: the Galactic plane

2025

- 2.5 yr more data
- Addition of muon tracks (starting events + up-going)
 - \rightarrow detection > 5 σ
- Flux in inner Galaxy is 2x higher





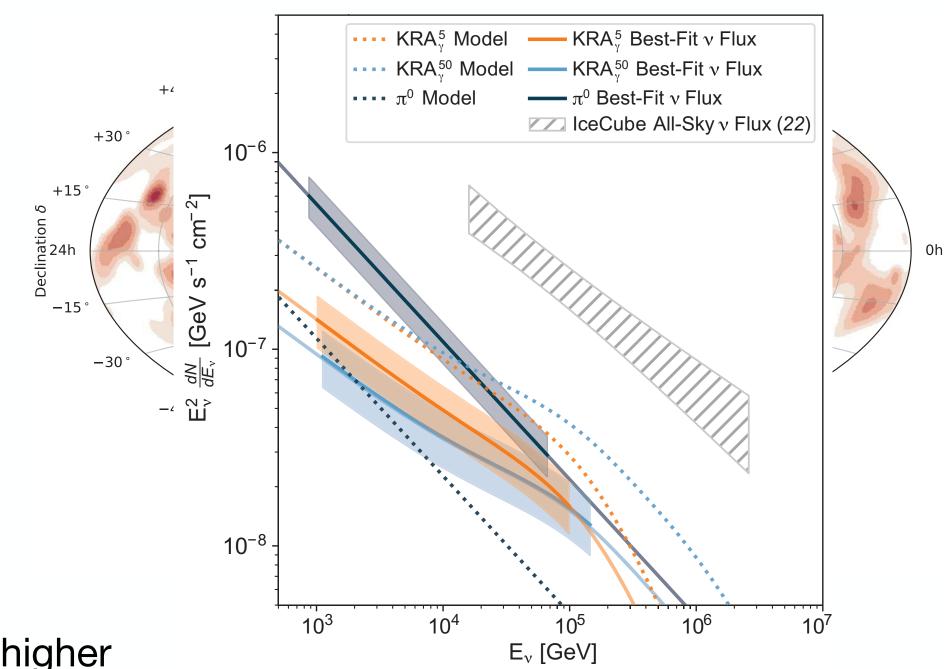
ICRC2025(1130, 1219)

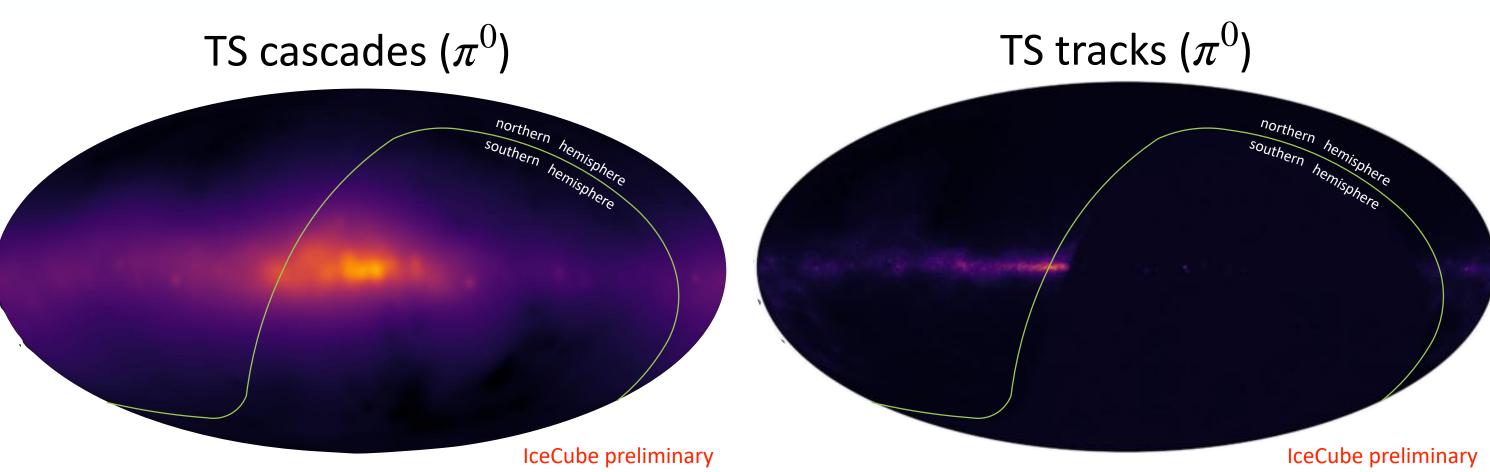
A "local" source: the Galactic plane

2025

- 2.5 yr more data
- Addition of muon tracks (starting events + up-going)
 - \rightarrow detection > 5 σ
- Flux in inner Galaxy is 2x higher

• The best model remains Fermi π^0 but the observed flux is still x5 higher than predicted by the model





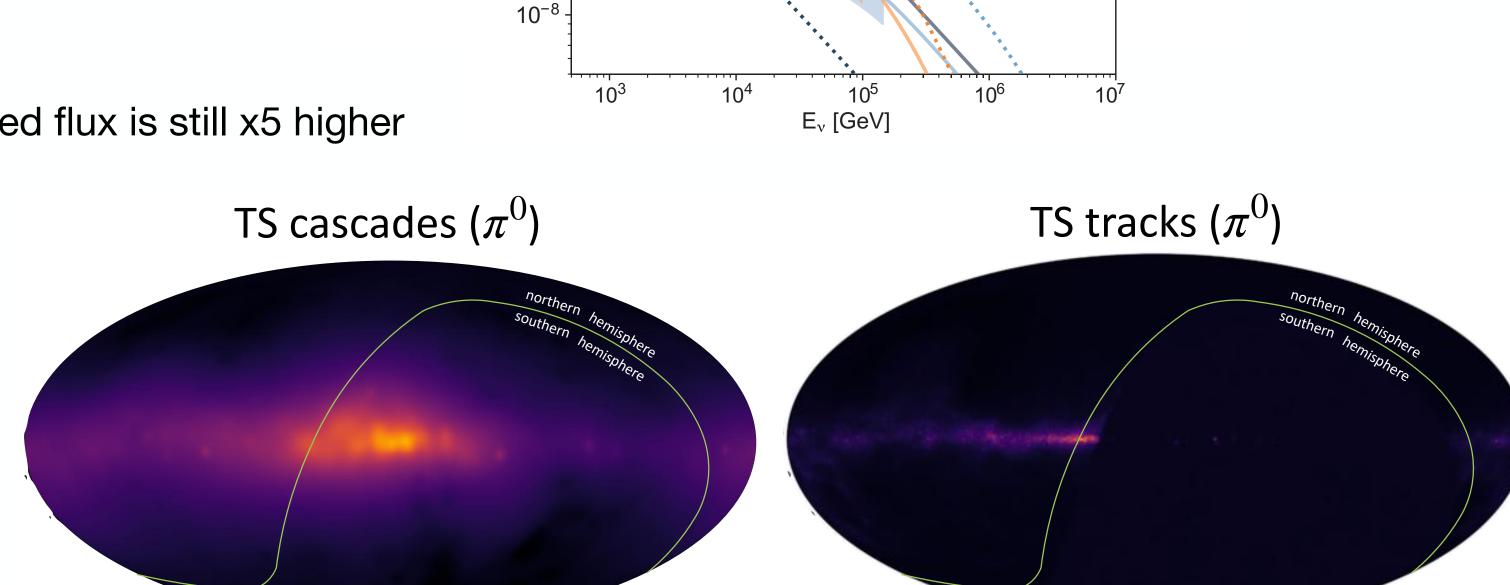
ICRC2025(1130, 1219)

A "local" source: the Galactic plane

2025

- 2.5 yr more data
- Addition of muon tracks (starting events + up-going)
 - \rightarrow detection > 5 σ
- Flux in inner Galaxy is 2x higher
- The best model remains Fermi π^0 but the observed flux is still x5 higher

than predicted by the model



IceCube preliminary

KRA⁵_ν Best-Fit ν Flux

KRA⁵⁰ Best-Fit v Flux

IceCube All-Sky v Flux (22)

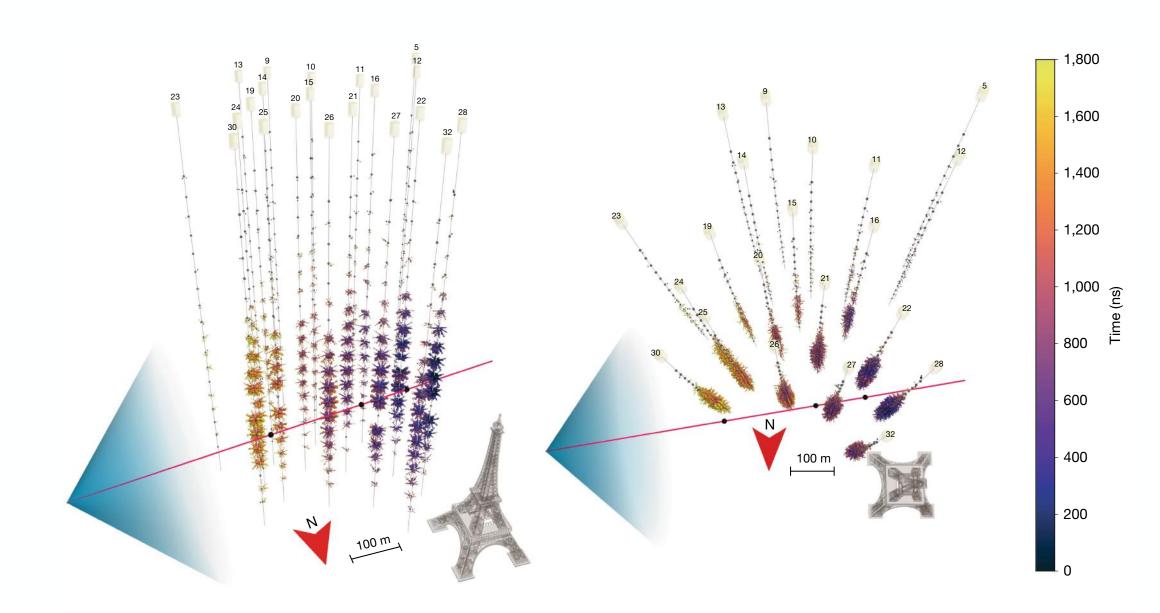
IceCube preliminary

No individual source identified (yet?)

ICRC2025(1130, 1219)

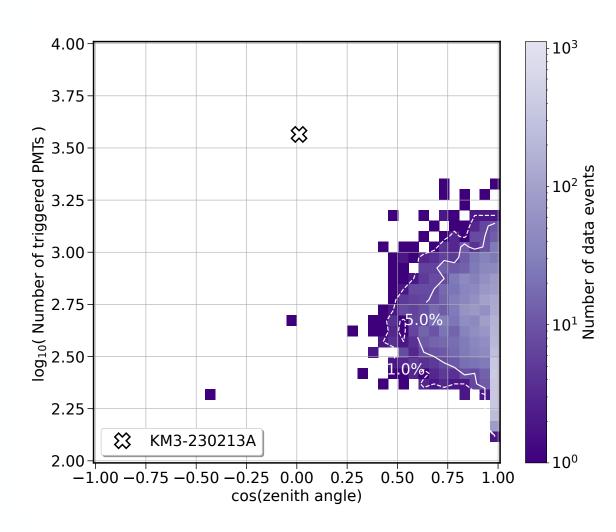
2023

Outstanding event : 28,086 PMT hits → clearest track ever



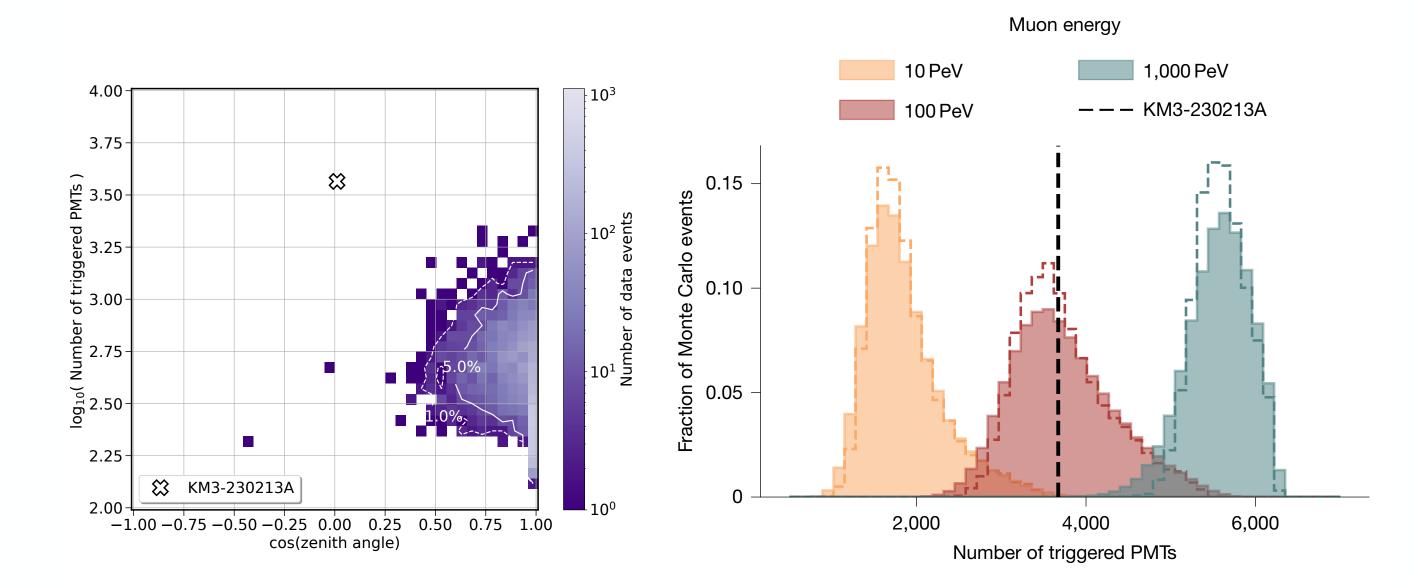
2023

Outstanding event : 28,086 PMT hits → clearest track ever



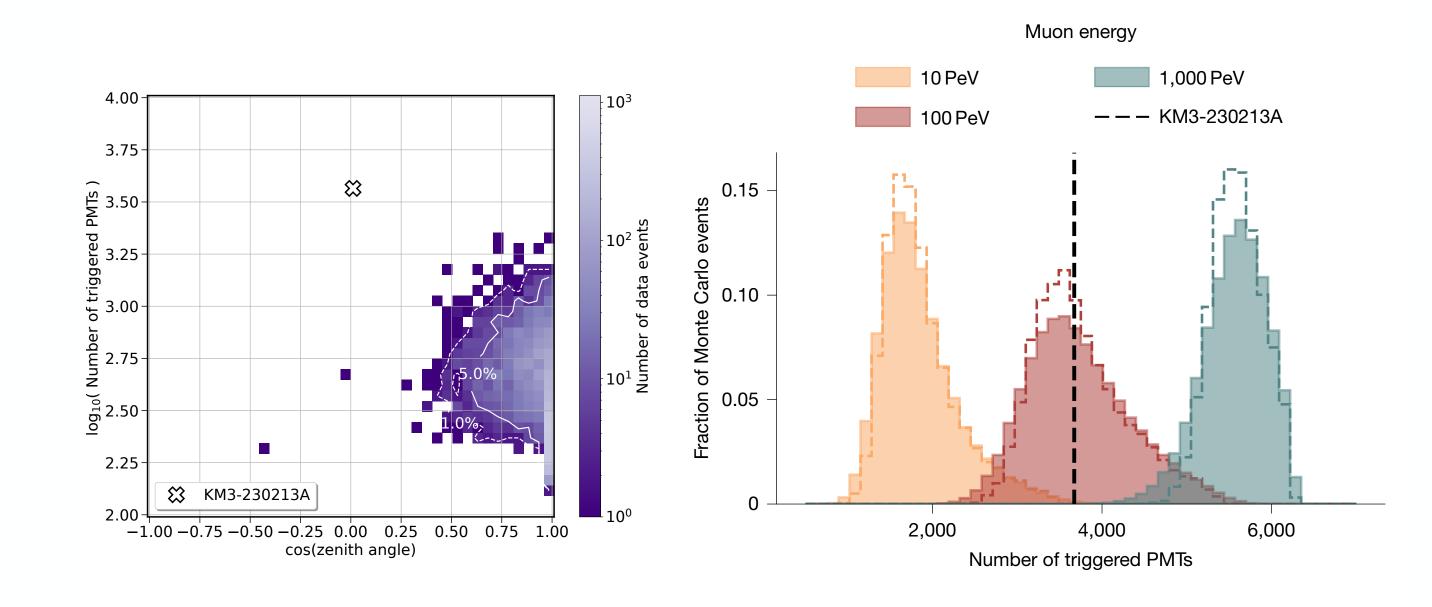
2023

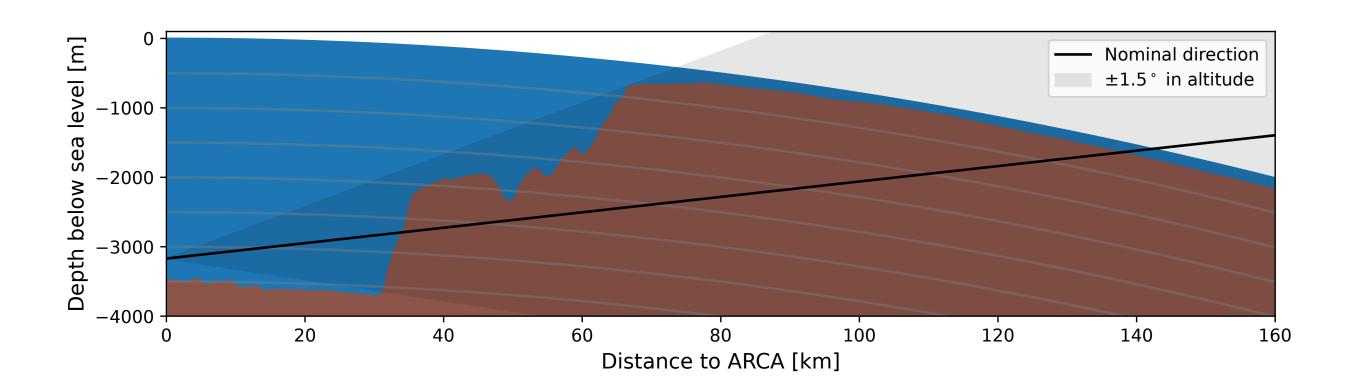
- Outstanding event: 28,086 PMT hits → clearest track ever
- Detected muon with E = 120 (+110/-60) PeV
 - → UHE event



2023

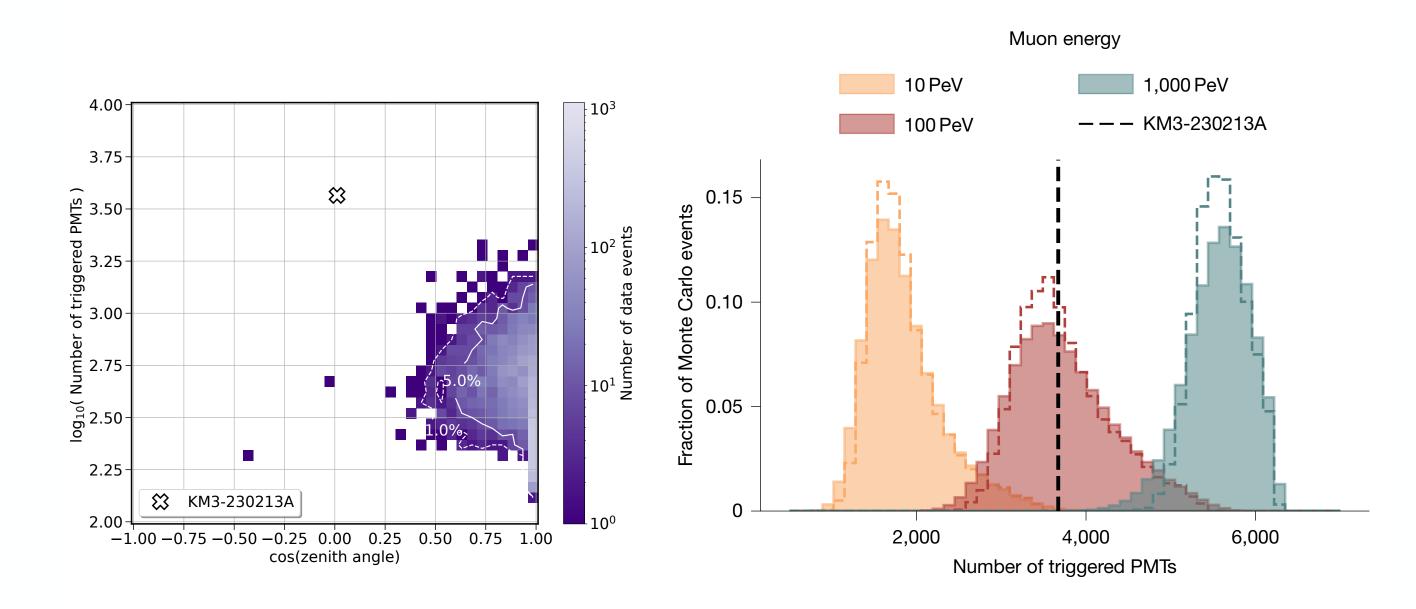
- Outstanding event: 28,086 PMT hits → clearest track ever
- Detected muon with E = 120 (+110/-60) PeV
 - → UHE event
- Direction near horizon → through the continental plateau
 - → Neutrino event

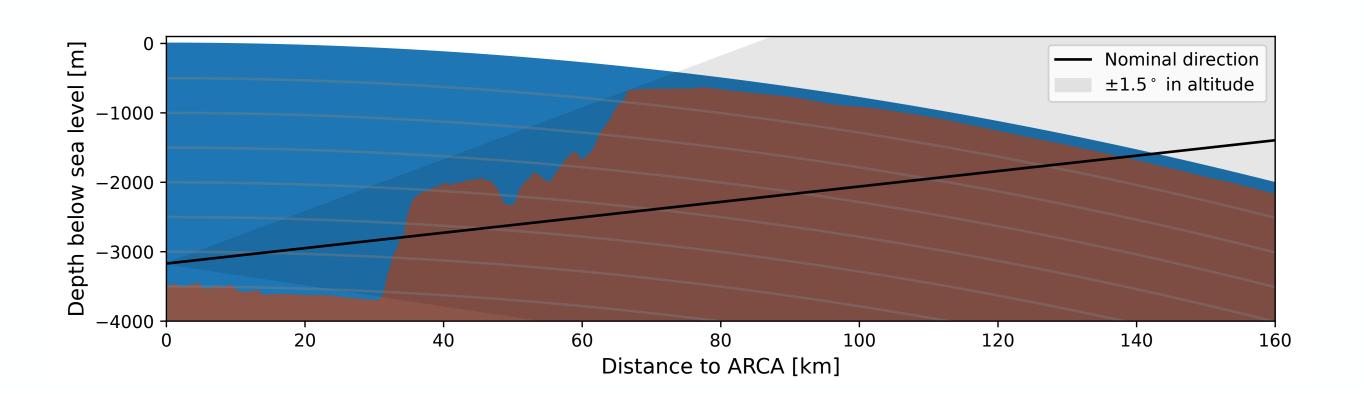




2023

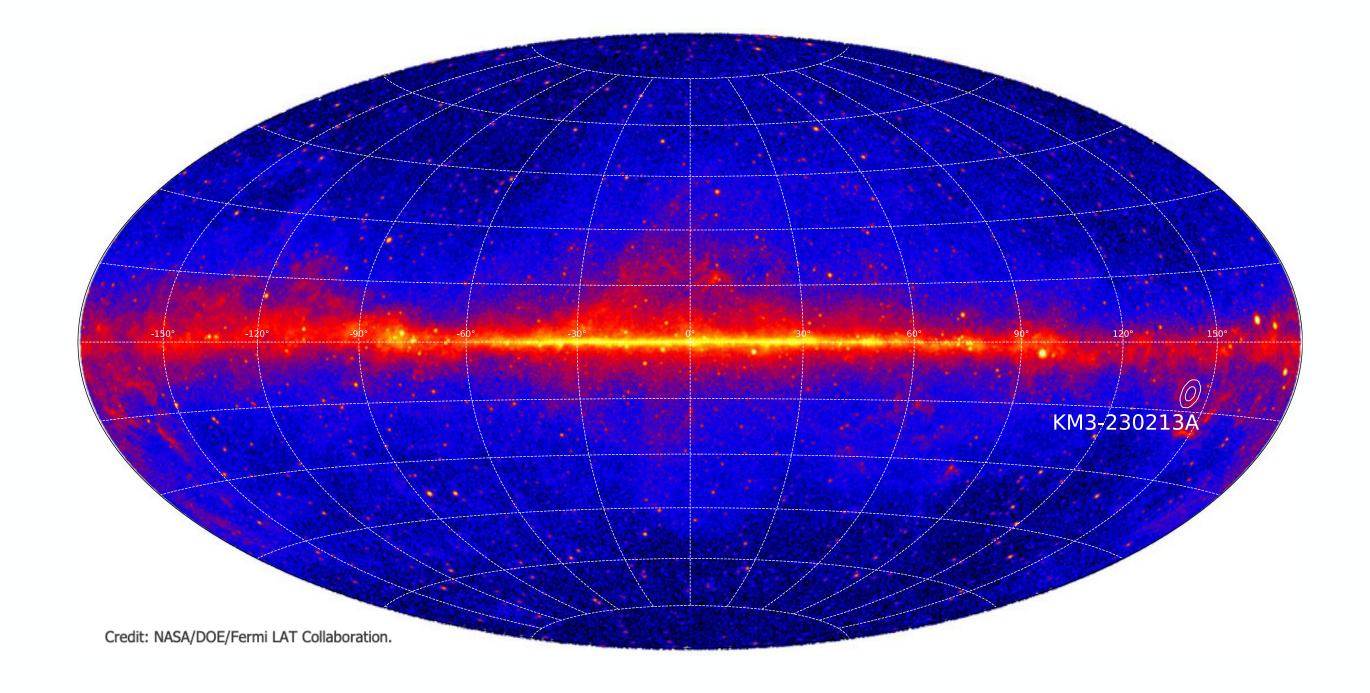
- Outstanding event : 28,086 PMT hits → clearest track ever
- Detected muon with E = 120 (+110/-60) PeV
 - → UHE event
- Direction near horizon → through the continental plateau
 - → Neutrino event
- \rightarrow UHE neutrino with E = 220 (+570/-110) PeV





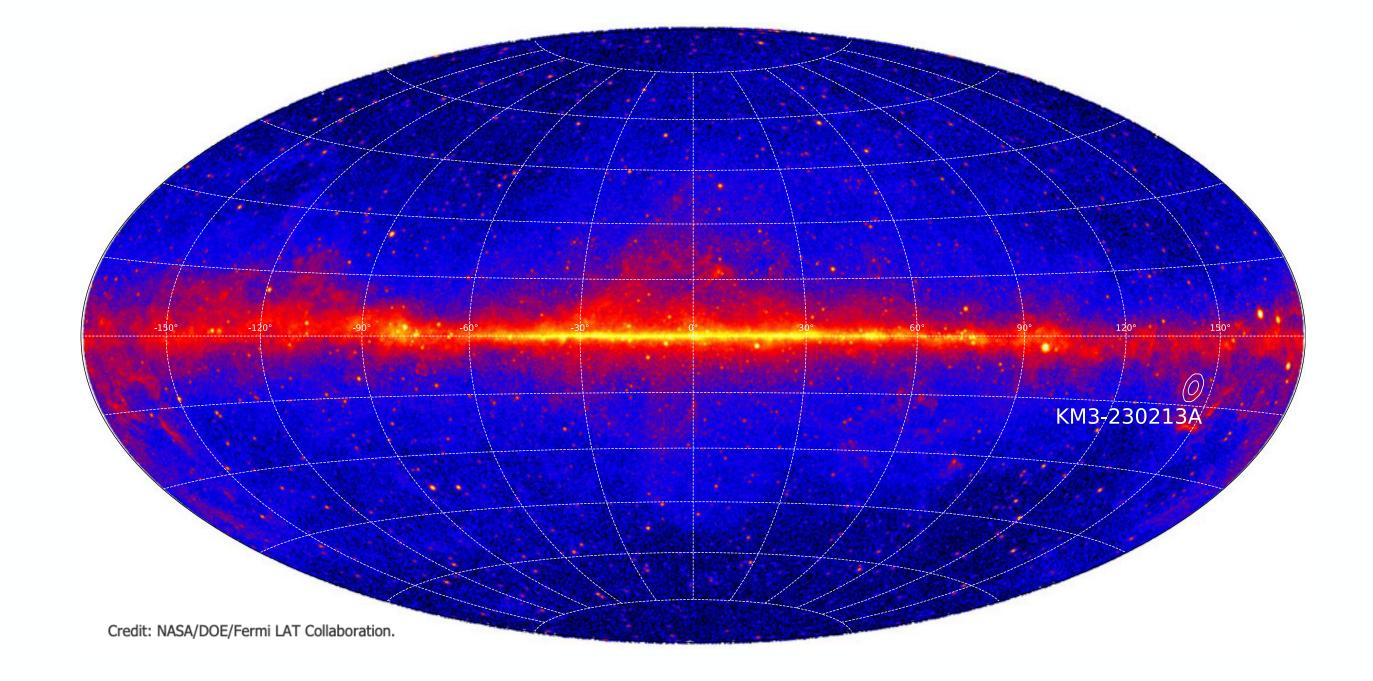
2024-2025

• From the astrophysical point of view → where is it from?



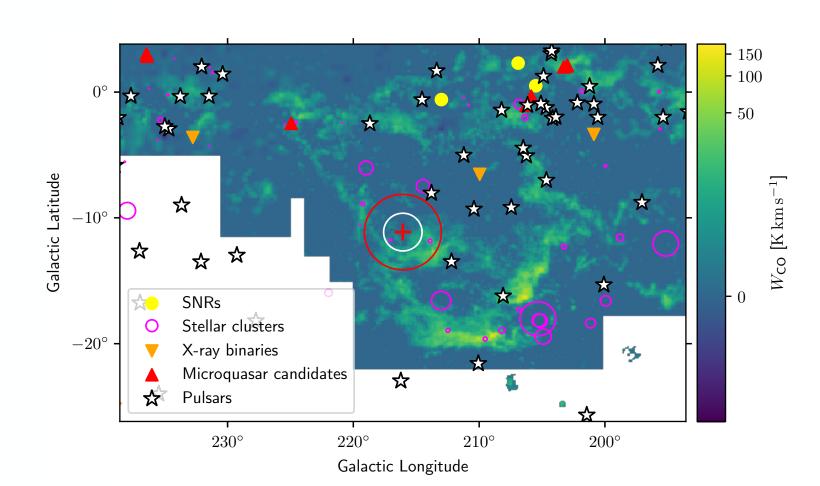
2024-2025

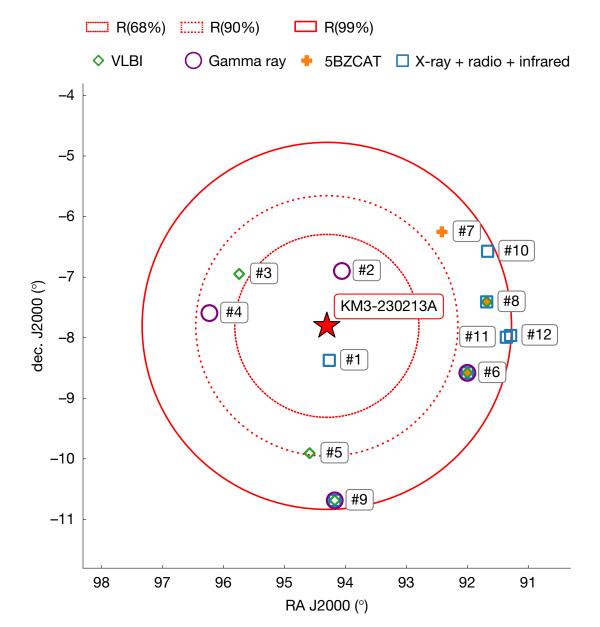
- From the astrophysical point of view → where is it from?
 - No counterparts (from radio to gamma)



2024-2025

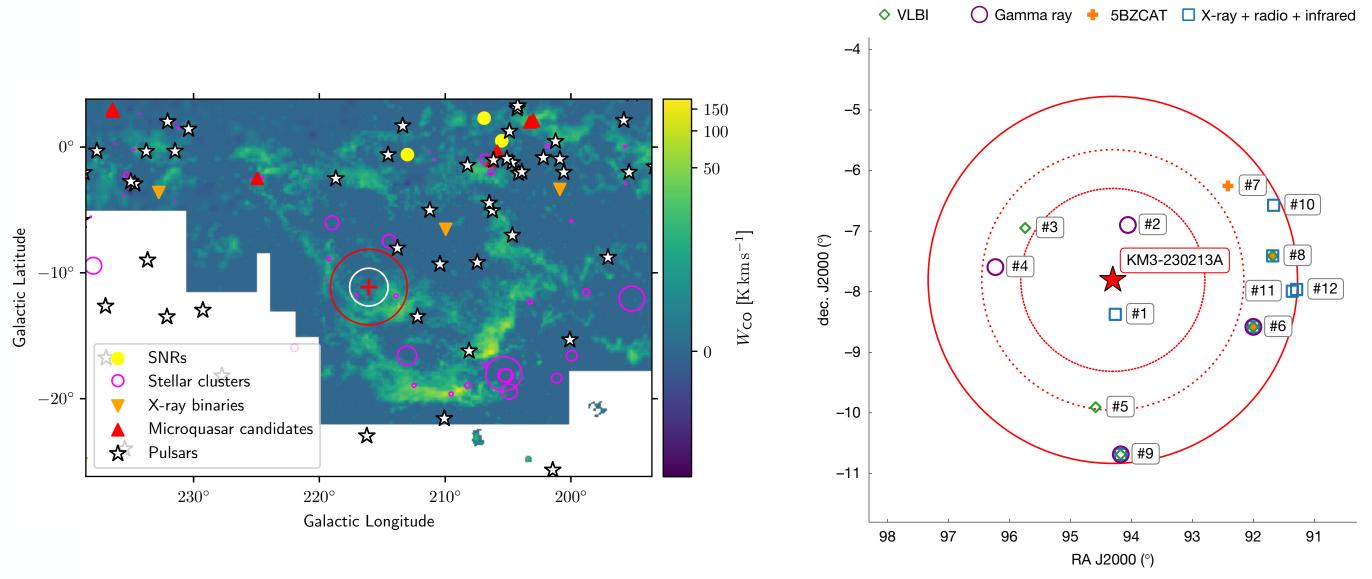
- From the astrophysical point of view → where is it from?
 - No counterparts (from radio to gamma)
 - No specific direction (a few blazar in the LOS and a galactic molecular cloud)

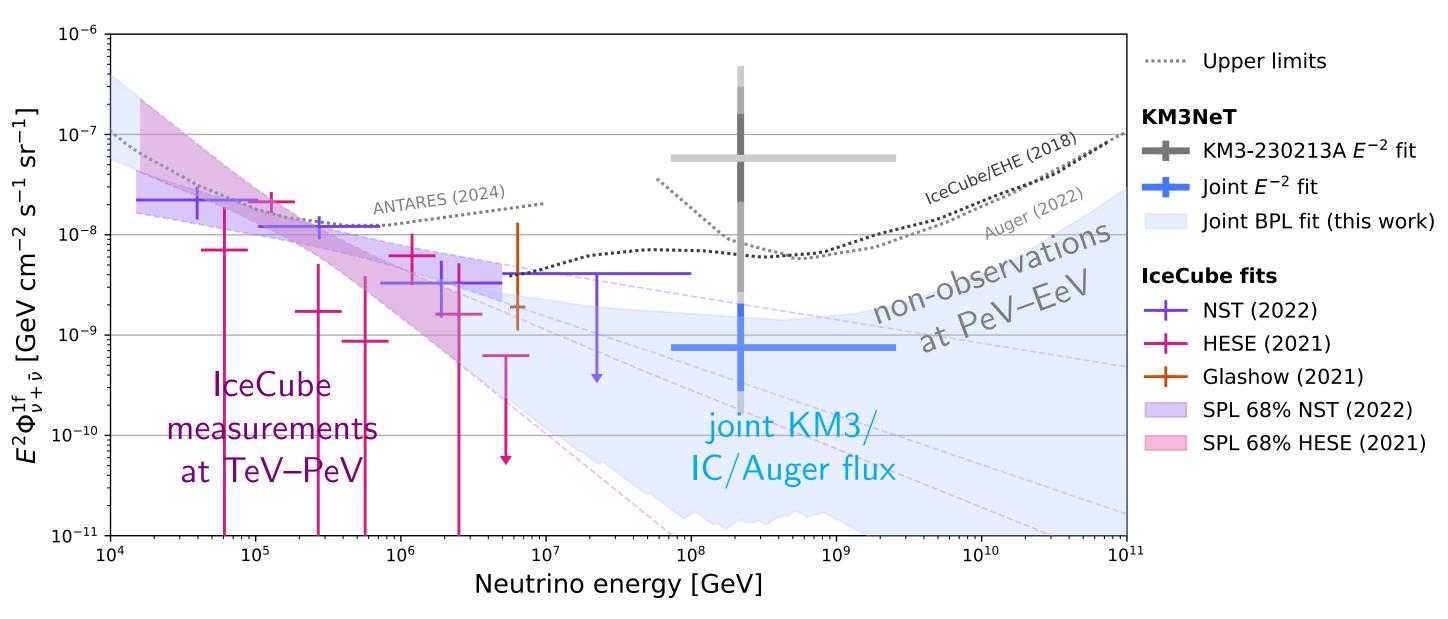




2024-2025

- From the astrophysical point of view → where is it from?
 - No counterparts (from radio to gamma)
 - No specific direction (a few blazar in the LOS and a galactic molecular cloud)
 - Can it be cosmogenic? maybe but difficult to accommodate with UL from IC and Auger

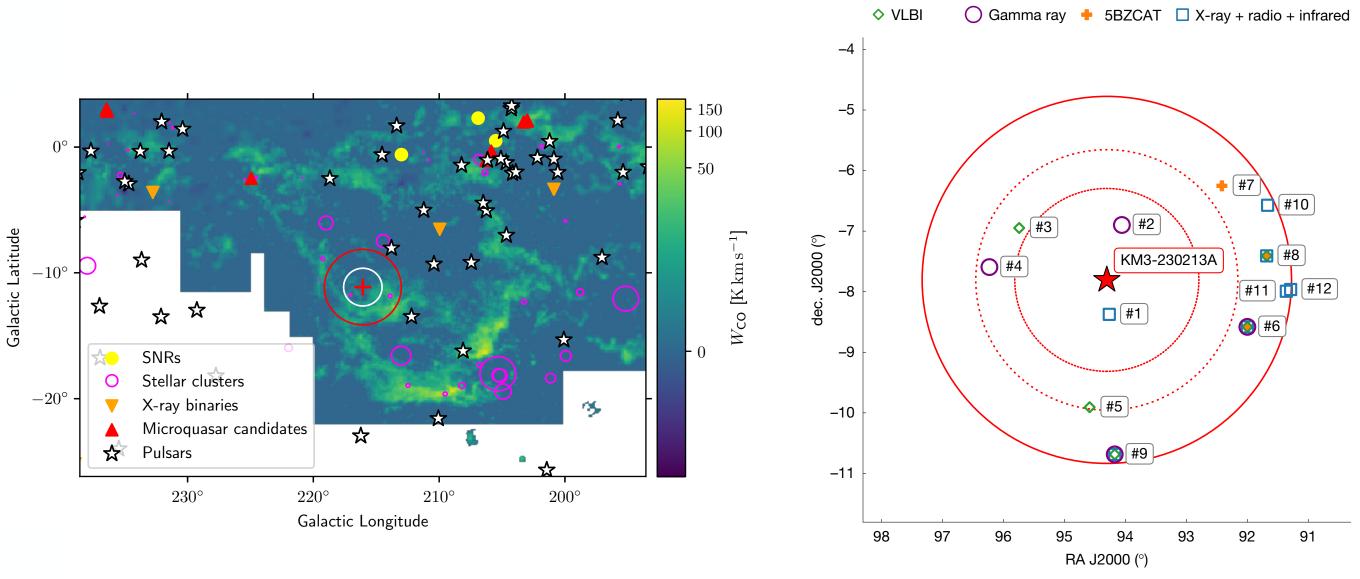


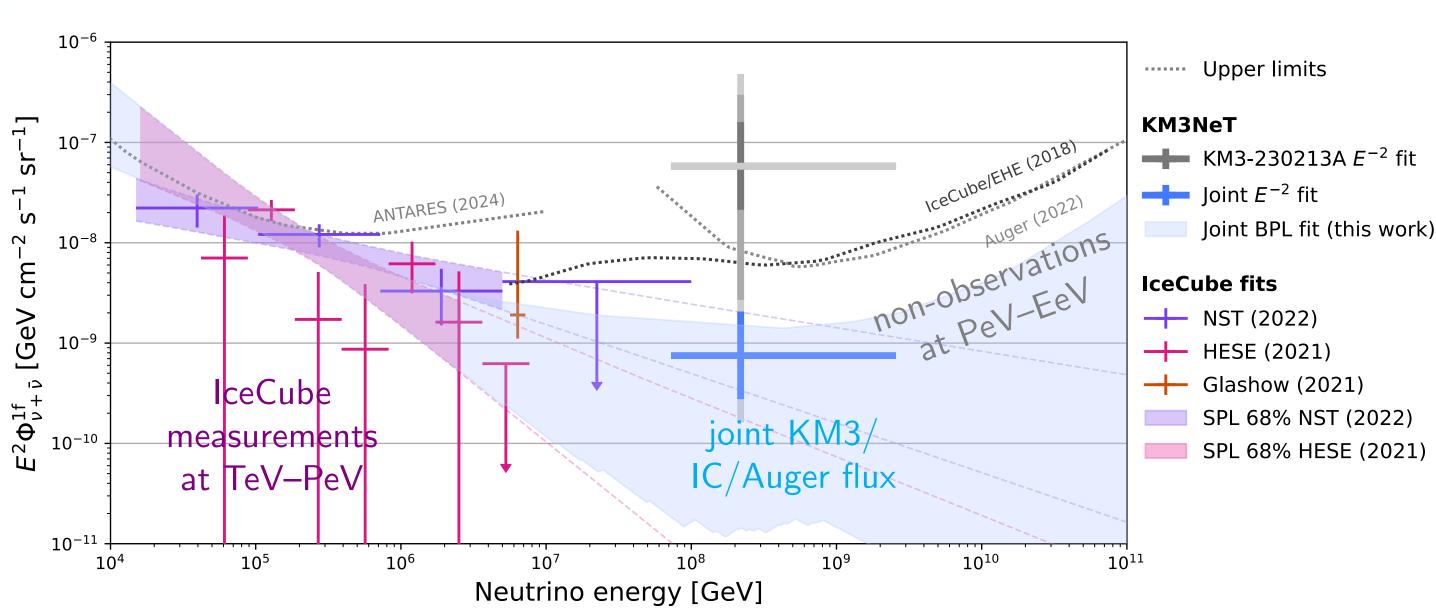


2024-2025

- From the astrophysical point of view → where is it from?
 - No counterparts (from radio to gamma)
 - No specific direction (a few blazar in the LOS and a galactic molecular cloud)
 - Can it be cosmogenic? maybe but difficult to accommodate with UL from IC and Auger

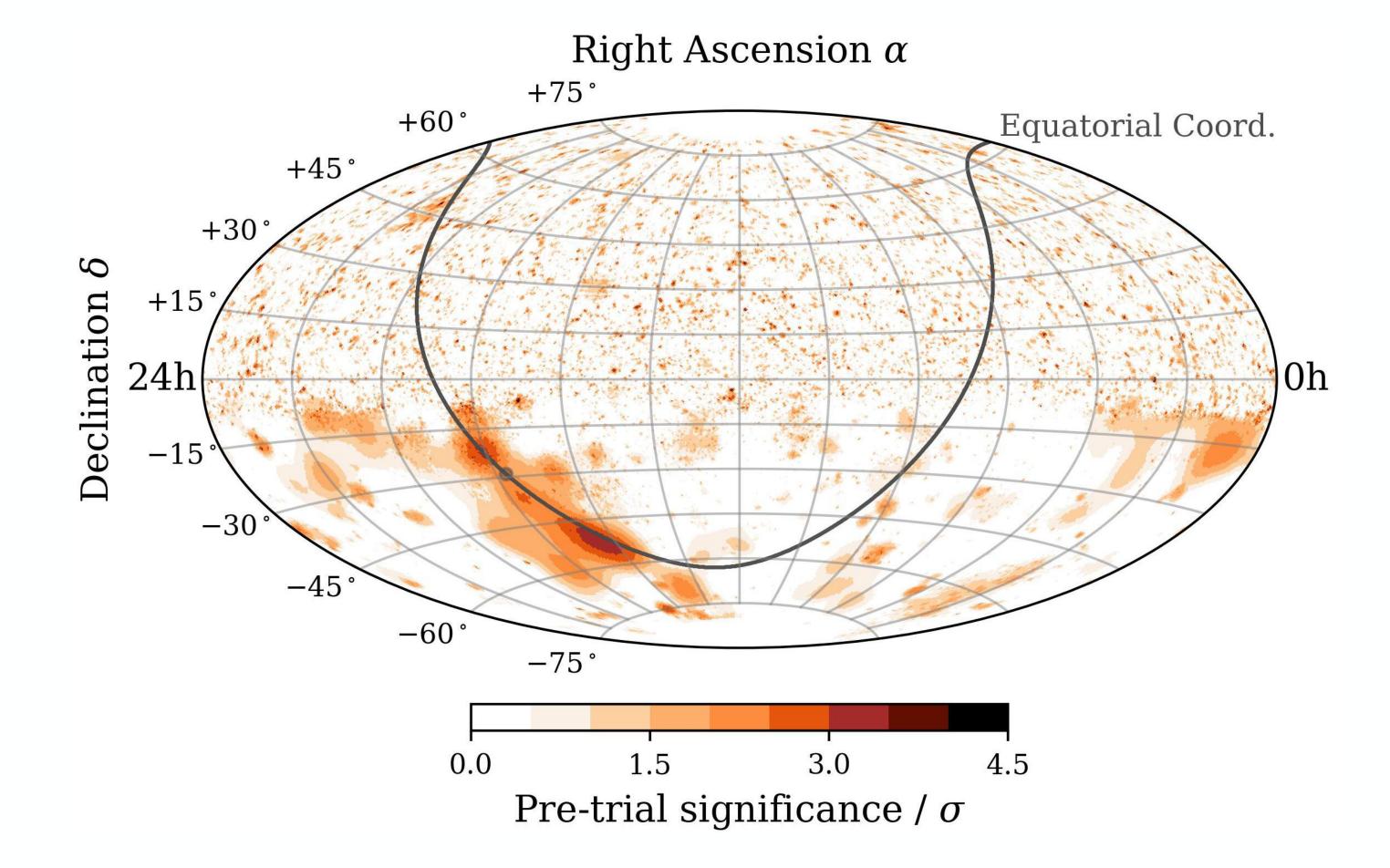
Tension ~2.5 σ with IceCube which makes this event even more interesting!



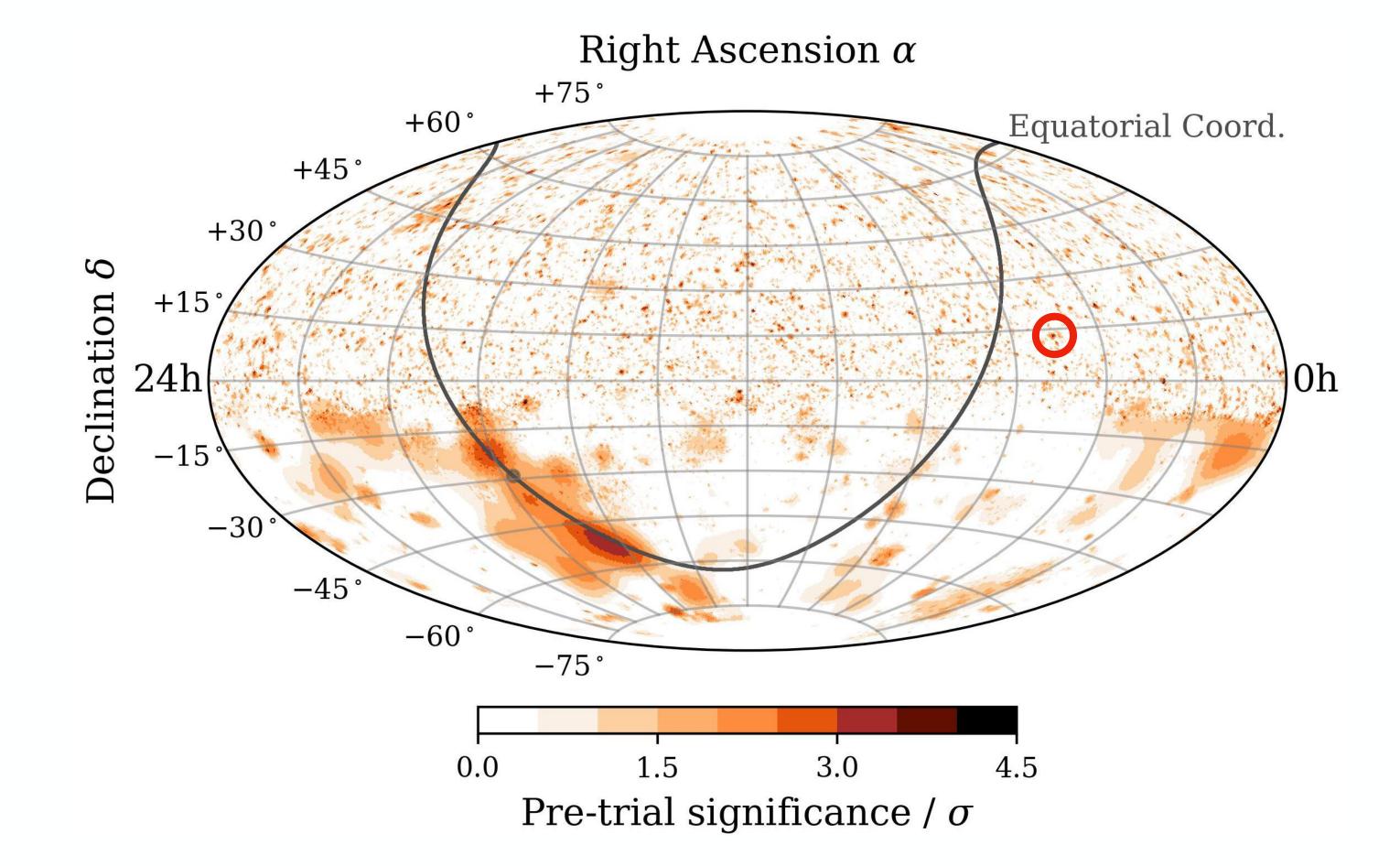


Four neutrino contributions identified:

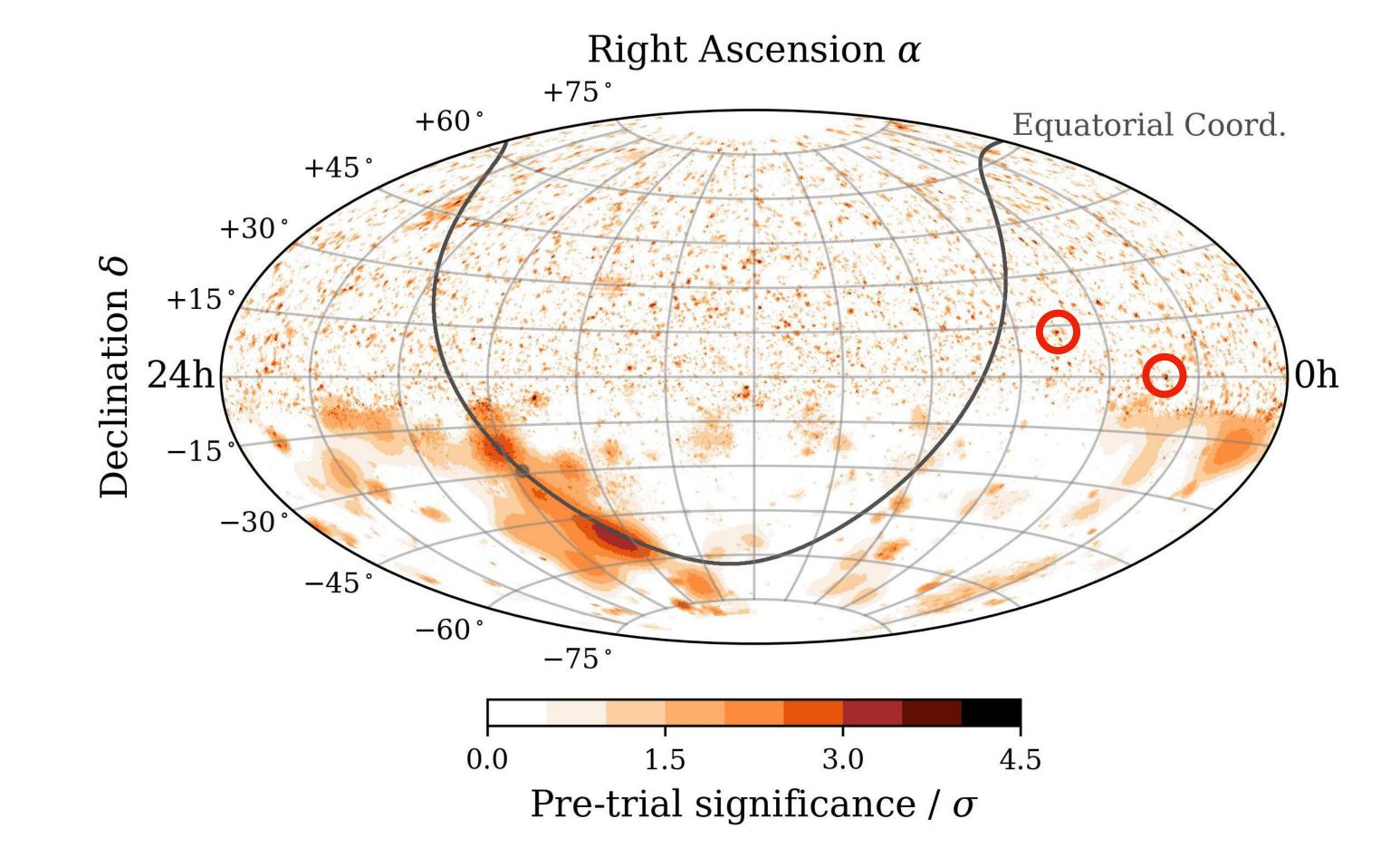
Diffuse flux (TeV-PeV)



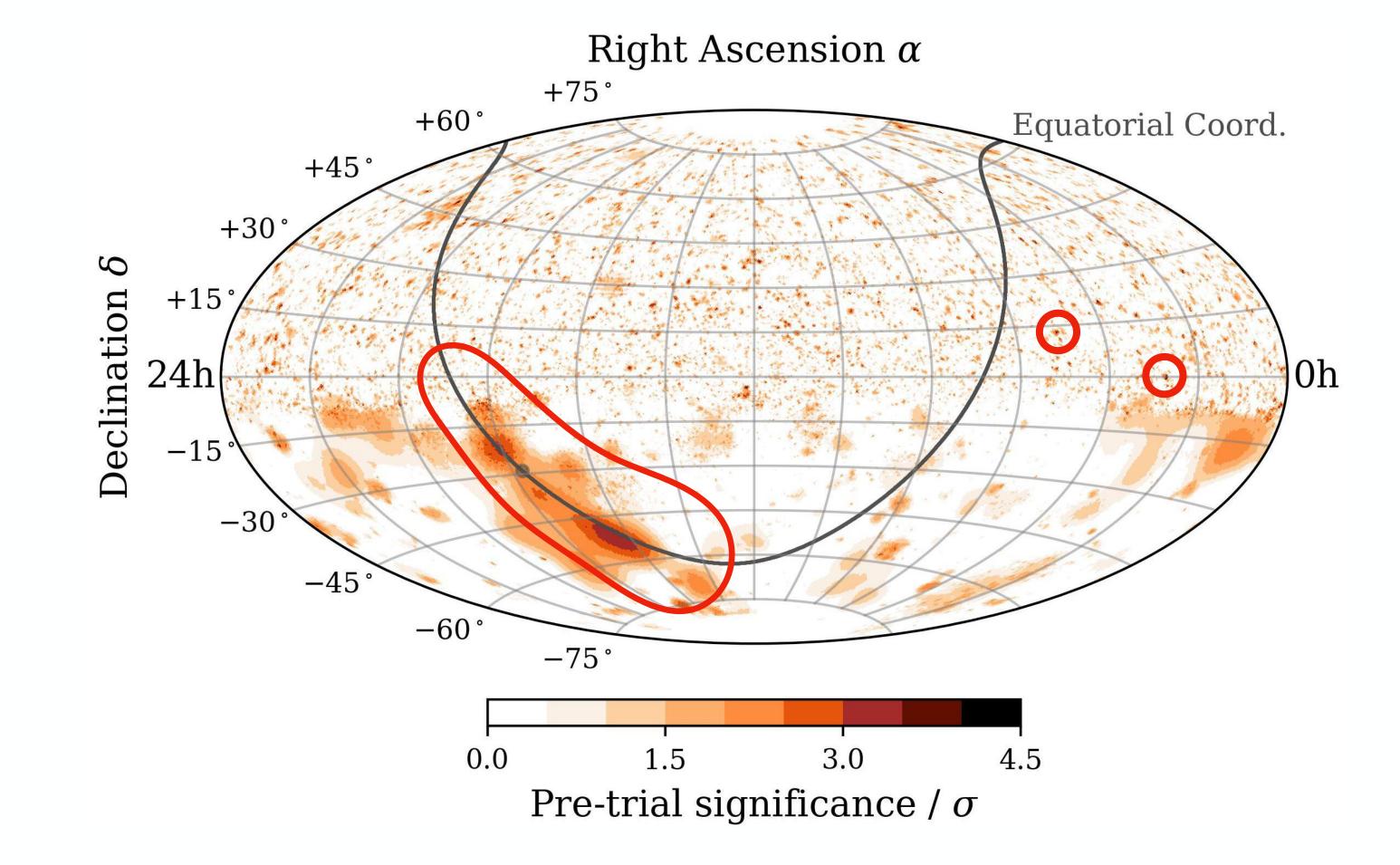
- Diffuse flux (TeV-PeV)
- TXS0506+056 (~ PeV)



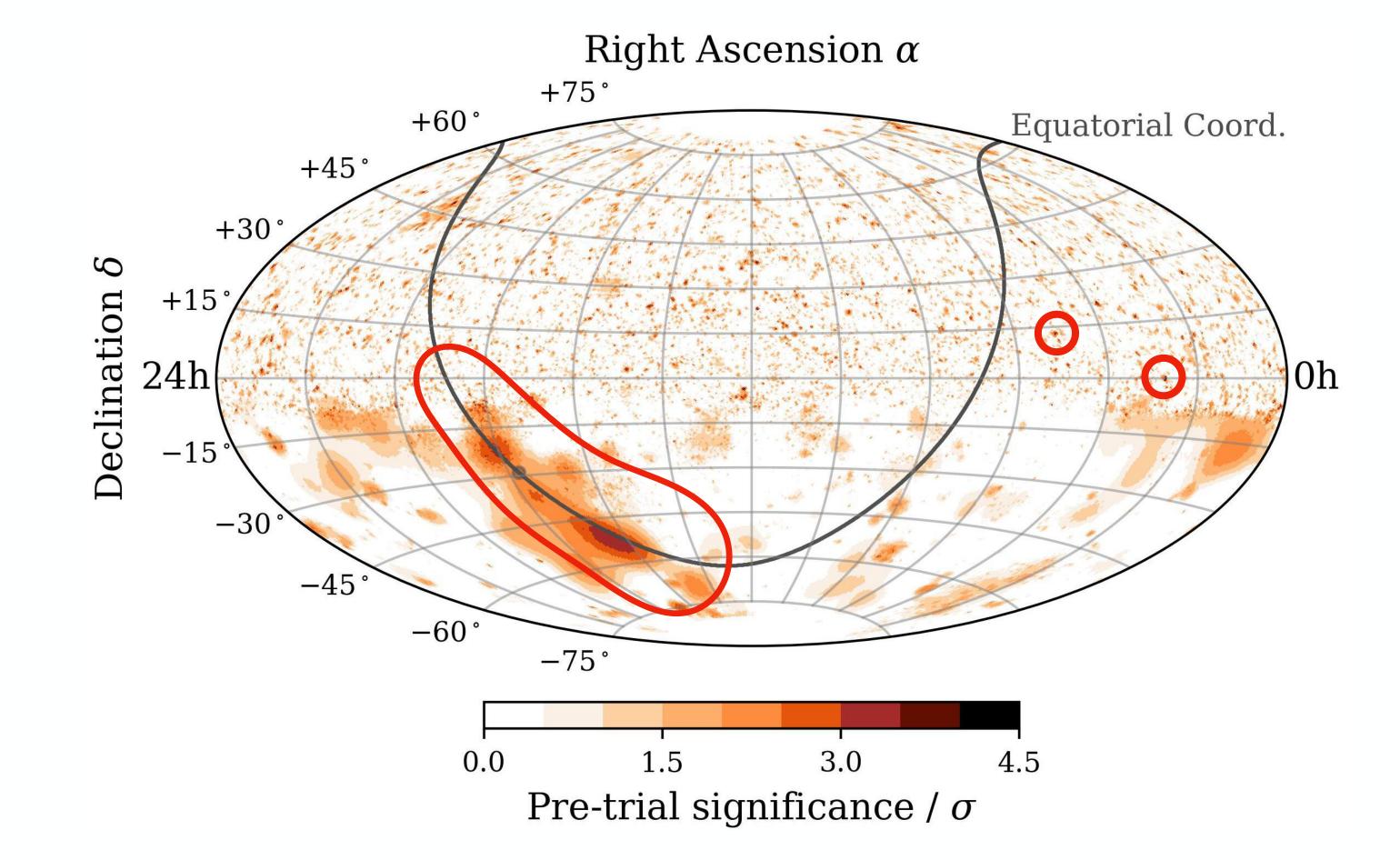
- Diffuse flux (TeV-PeV)
- TXS0506+056 (~ PeV)
- NGC1068 (~ TeV)



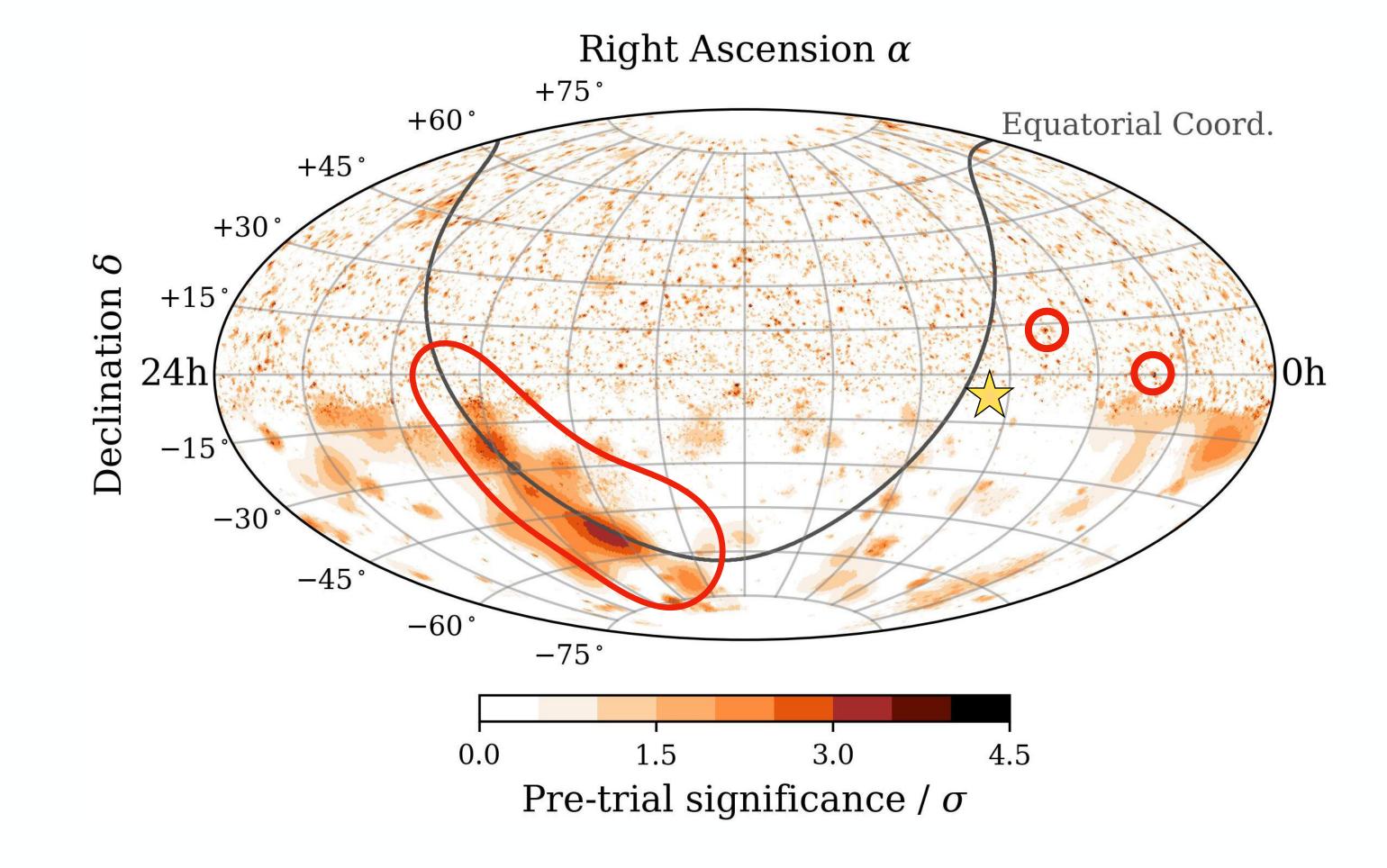
- Diffuse flux (TeV-PeV)
- TXS0506+056 (~ PeV)
- NGC1068 (~ TeV)
- Galactic plane (TeV-PeV)



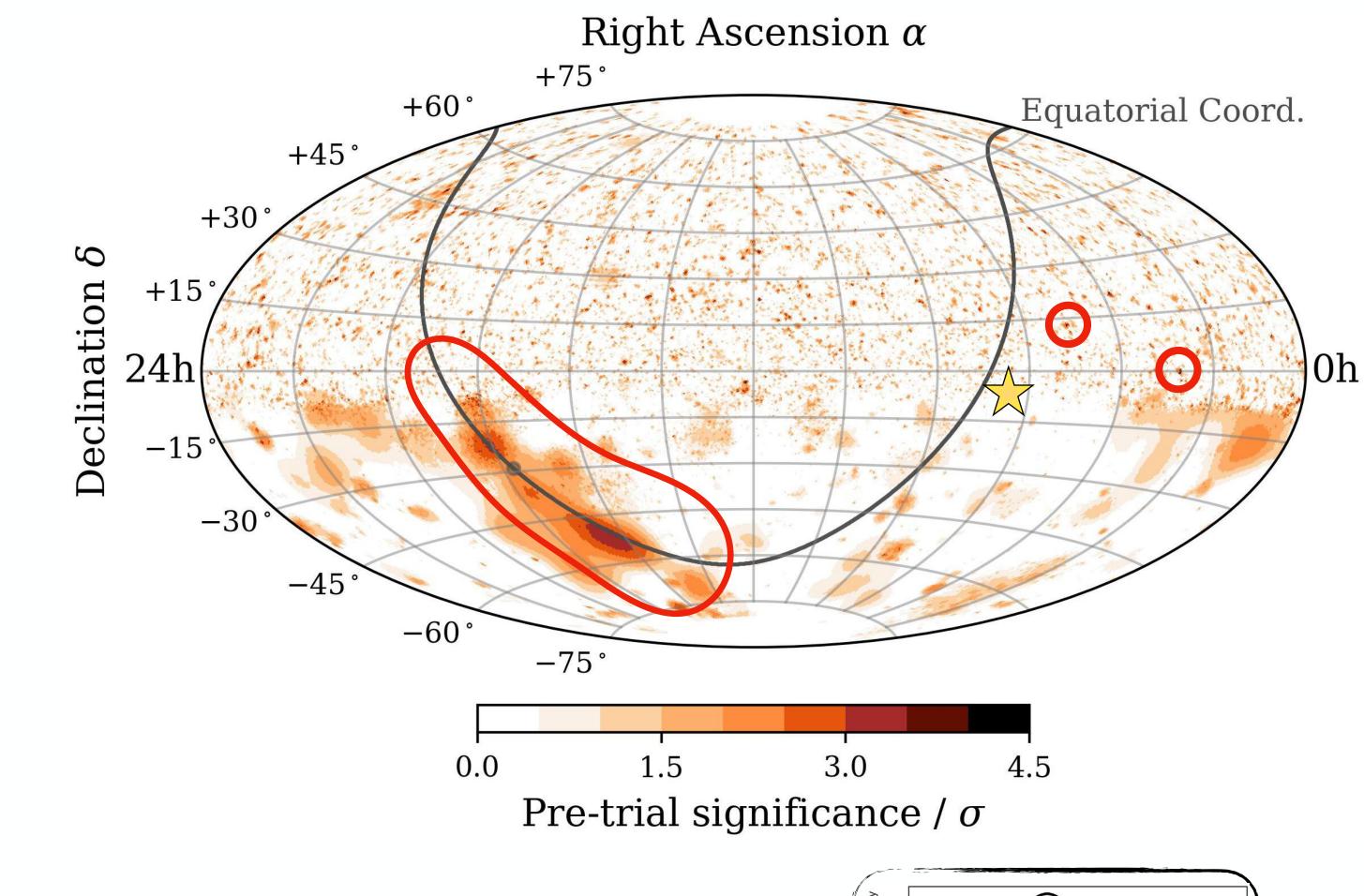
- Diffuse flux (TeV-PeV)
- TXS0506+056 (~ PeV)
- NGC1068 (~ TeV)
- Galactic plane (TeV-PeV)
- → Predominance of AGNs (blazars + Seyferts)

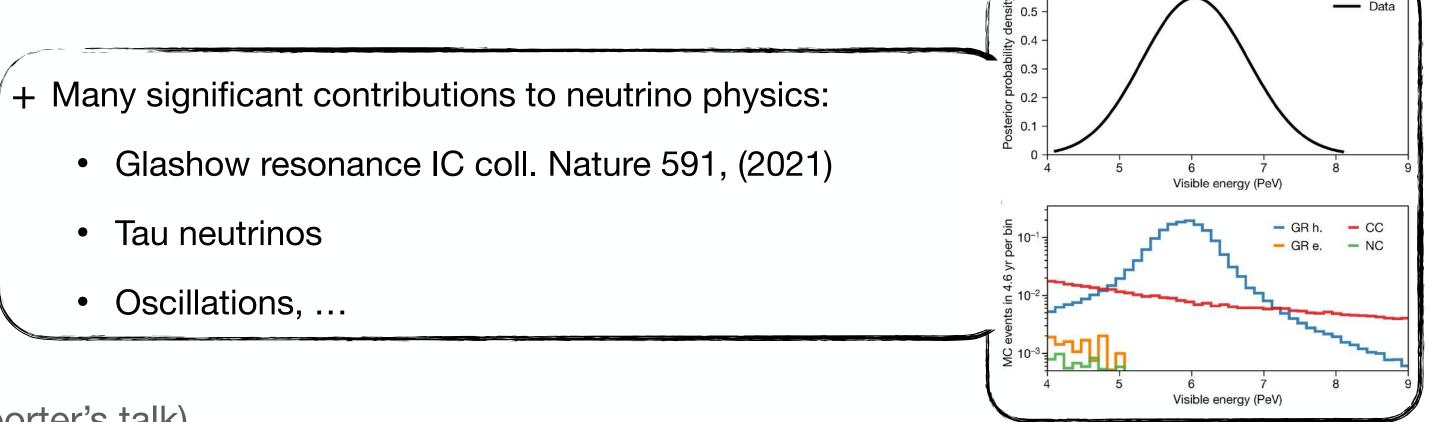


- Diffuse flux (TeV-PeV)
- TXS0506+056 (~ PeV)
- NGC1068 (~ TeV)
- Galactic plane (TeV-PeV)
- → Predominance of AGNs (blazars + Seyferts)
- ☆First UHE neutrino!



- Diffuse flux (TeV-PeV)
- TXS0506+056 (~ PeV)
- NGC1068 (~ TeV)
- Galactic plane (TeV-PeV)
- → Predominance of AGNs (blazars + Seyferts)
- ☆First UHE neutrino!





At high energy: combination of IceCube and KM3NeT data

- Diffuse flux: more resolved sources and better measurement of spectrum features
- Steady sources: stacking analysis of AGNs should reach the discovery level
- Galactic plane: what sources are hidden there? → KM3NeT FoV!
- Transients: new multi-messenger events thanks to the fast development of alerts and follow-up networks

At high energy: combination of IceCube and KM3NeT data

- Diffuse flux: more resolved sources and better measurement of spectrum features
- Steady sources: stacking analysis of AGNs should reach the discovery level
- Galactic plane: what sources are hidden there? → KM3NeT FoV!
- Transients: new multi-messenger events thanks to the fast development of alerts and follow-up networks

At ultra-high energy:

- KM3NeT proved the existence of UHE neutrinos!
- Radio-detection experiments like GRAND/HERON will take the searches to the next level!

At high energy: combination of IceCube and KM3NeT data

- Diffuse flux: more resolved sources and better measurement of spectrum features
- Steady sources: stacking analysis of AGNs should reach the discovery level
- Galactic plane: what sources are hidden there? → KM3NeT FoV!
- Transients: new multi-messenger events thanks to the fast development of alerts and follow-up networks

At ultra-high energy:

- KM3NeT proved the existence of UHE neutrinos!
- Radio-detection experiments like GRAND/HERON will take the searches to the next level!



At high energy: combination of IceCube and KM3NeT data

- Diffuse flux: more resolved sources and better measurement of spectrum features
- Steady sources: stacking analysis of AGNs should reach the discovery level
- Galactic plane: what sources are hidden there? → KM3NeT FoV!
- Transients: new multi-messenger events thanks to the fast development of alerts and follow-up networks

At ultra-high energy:

- KM3NeT proved the existence of UHE neutrinos!
- Radio-detection experiments like GRAND/HERON will take the searches to the next level!



