

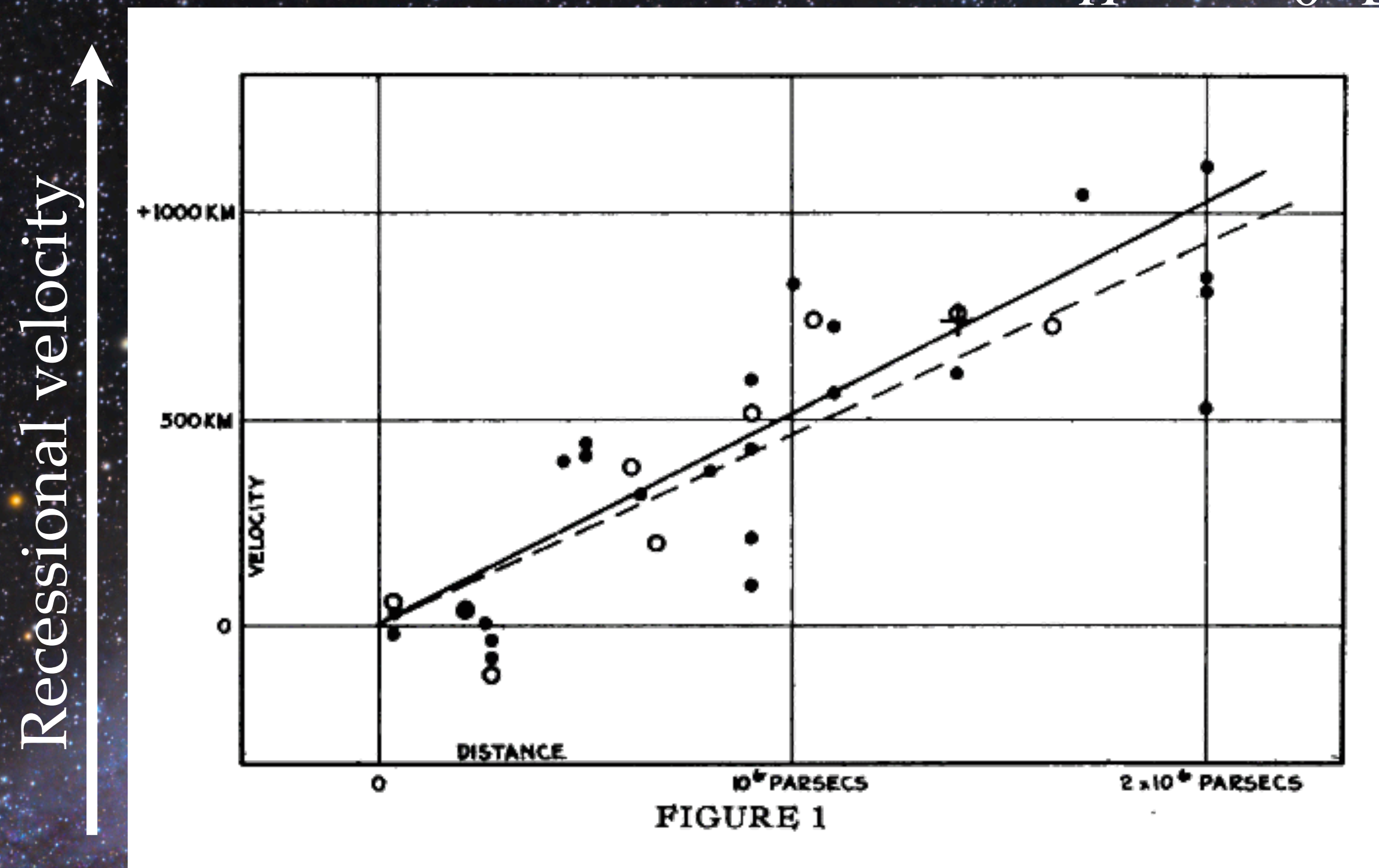


Class 1: Introduction | Observable universe & Dark energy

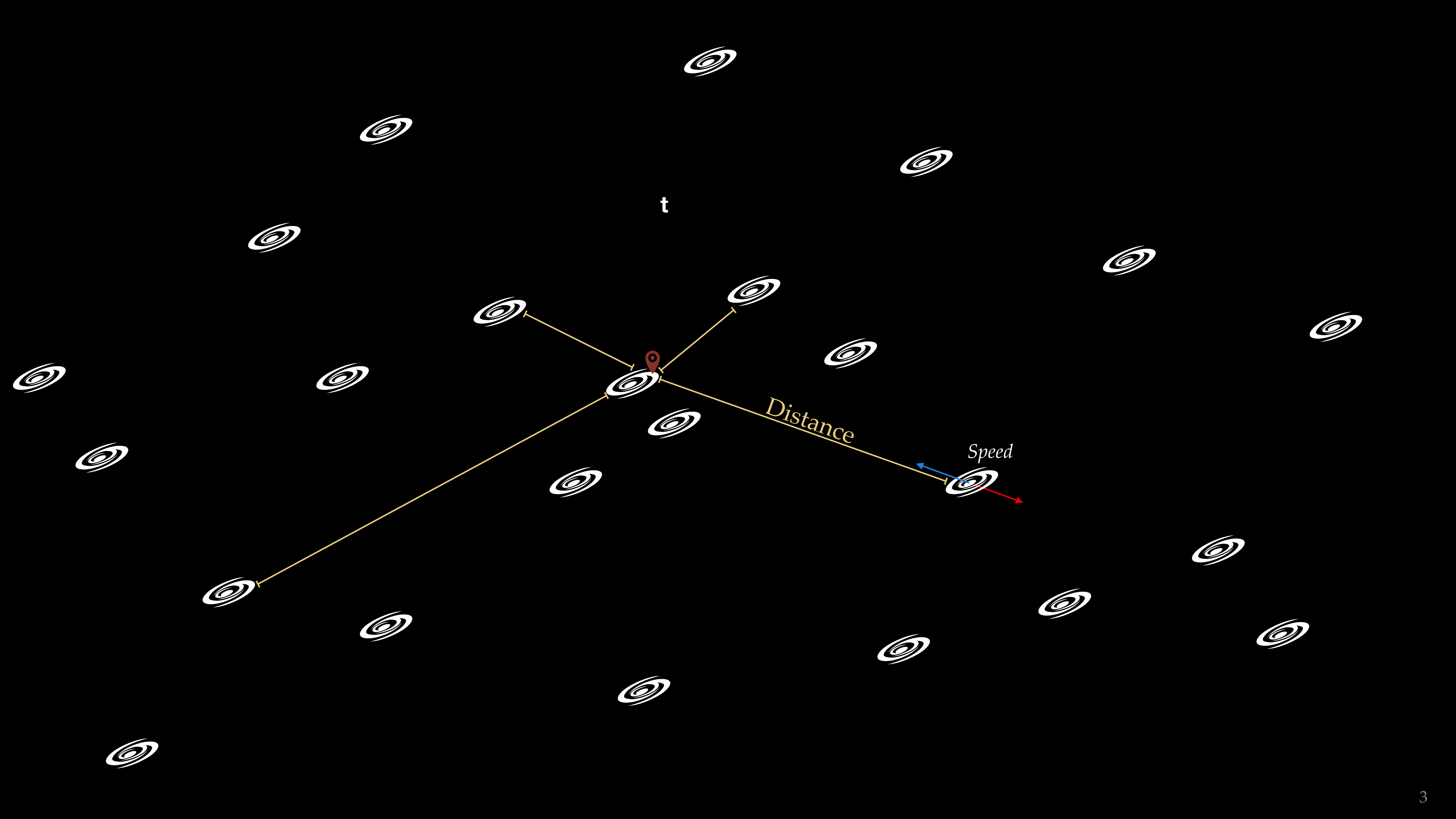
Class 2: Details on Type Ia Supernova cosmology

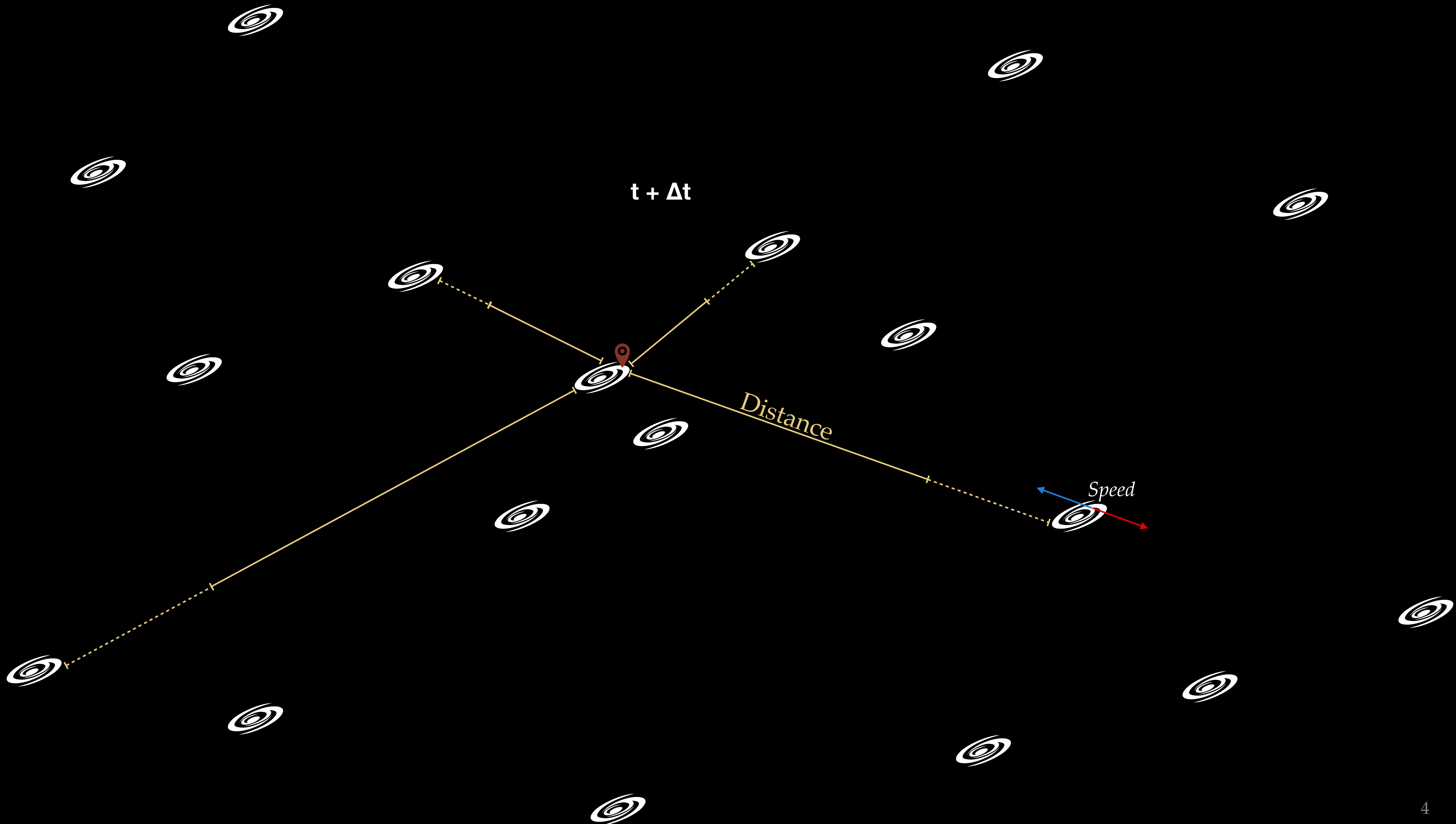
Class 3: The Hubble Constant tension

$$v_H = H_0 d_L$$



The Hubble Constant H_0 : How fast the Universe is currently expanding





Modern Cosmology | H_0 Direct vs. Indirect Measurements

$$H_0 = d_l / v_h$$

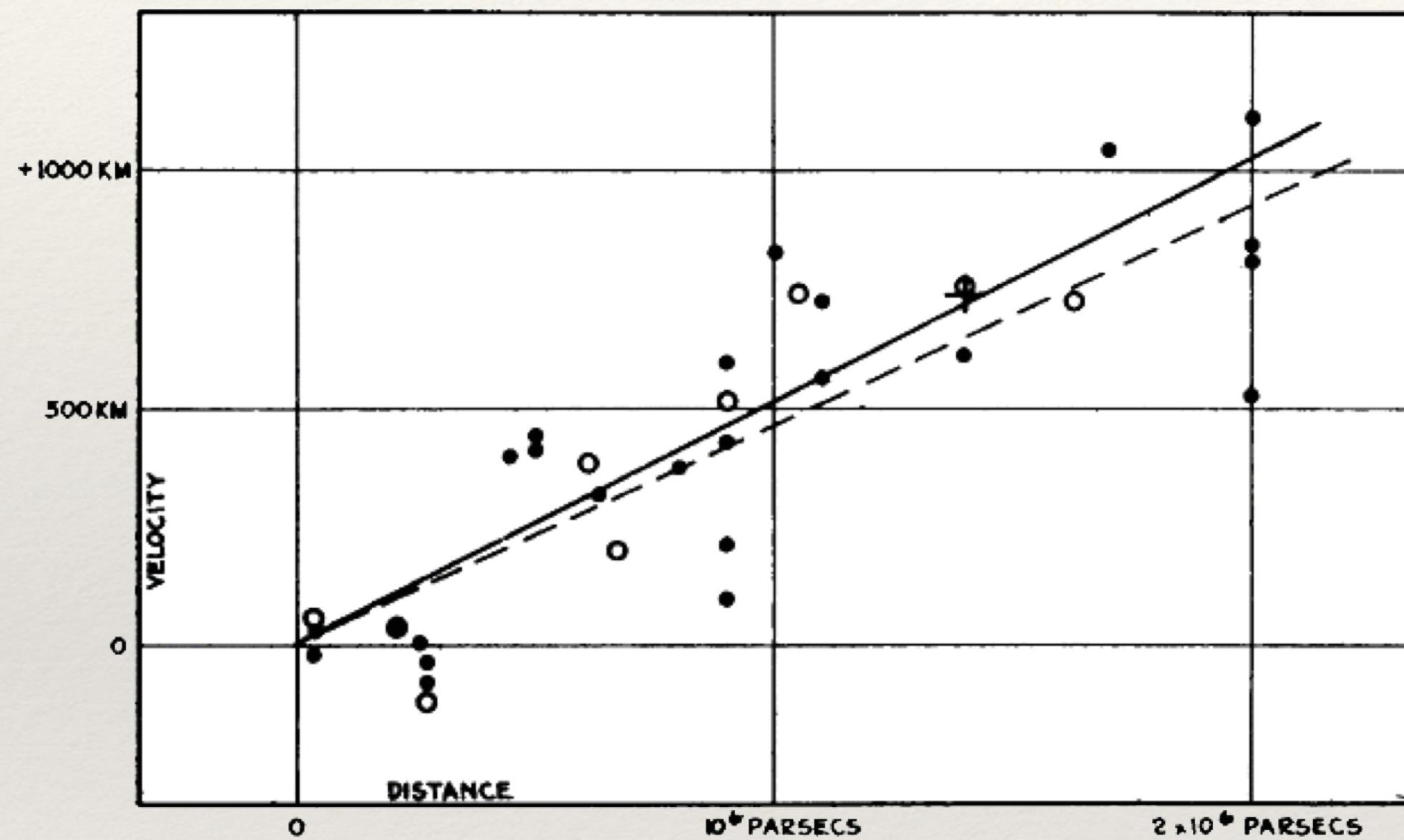
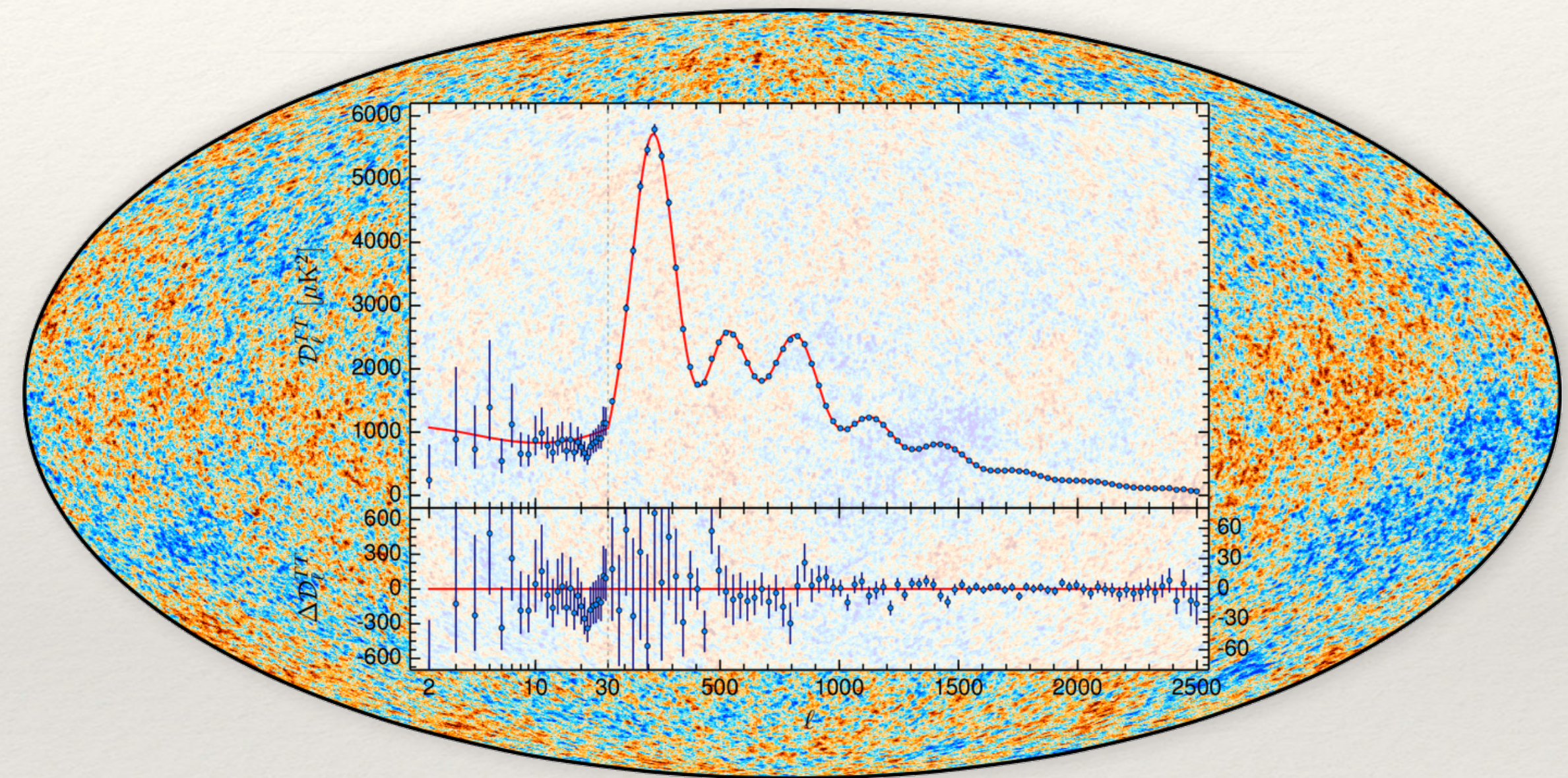


FIGURE 1

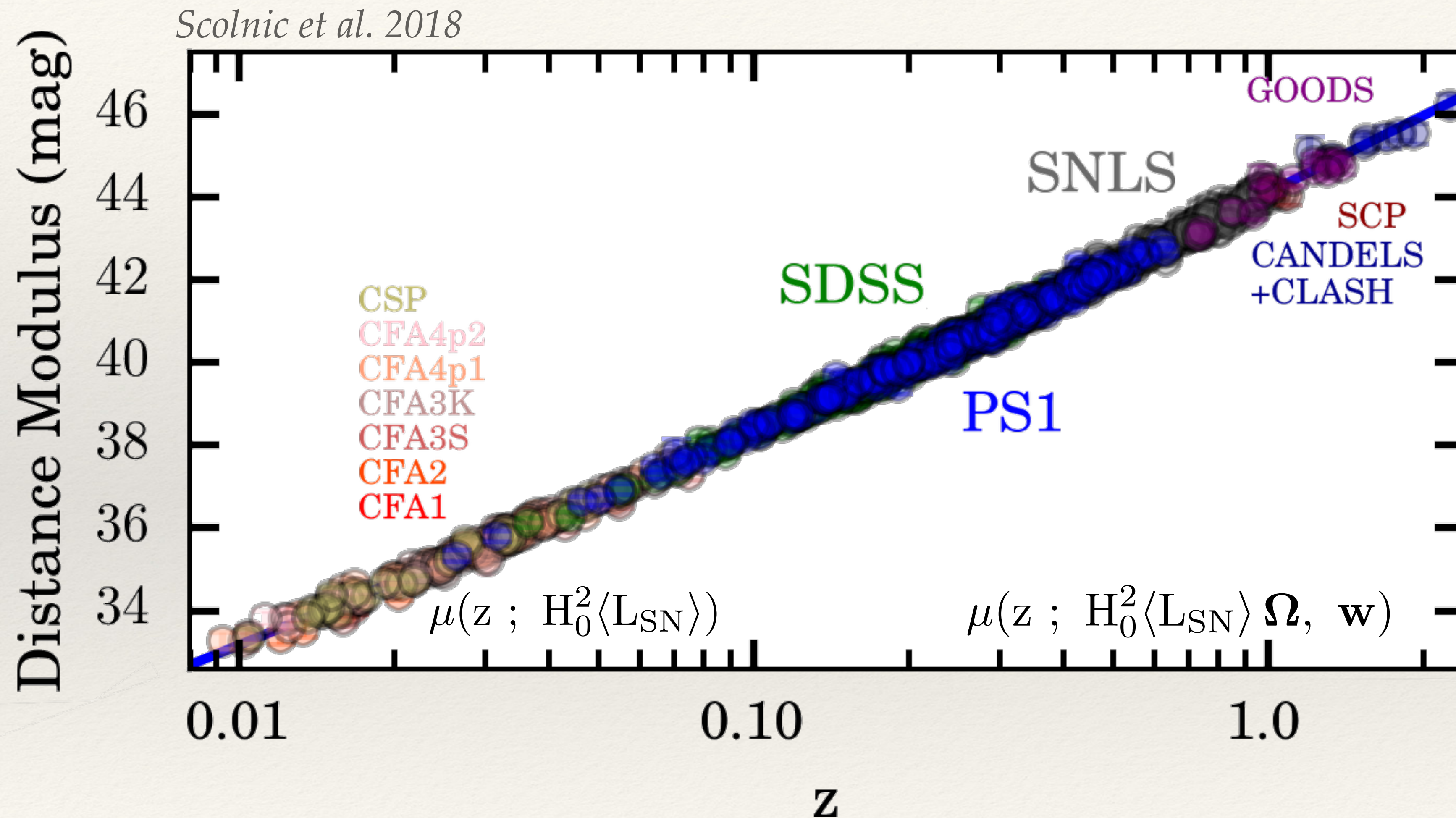
Careful with peculiar velocities

$$H(\underline{z}) = H_0 \times \sqrt{\Omega_r(1 + \underline{z})^4 + \Omega_m(1 + \underline{z})^3 + \Omega_\Lambda(1 + \underline{z})^{3(1+w)}}$$

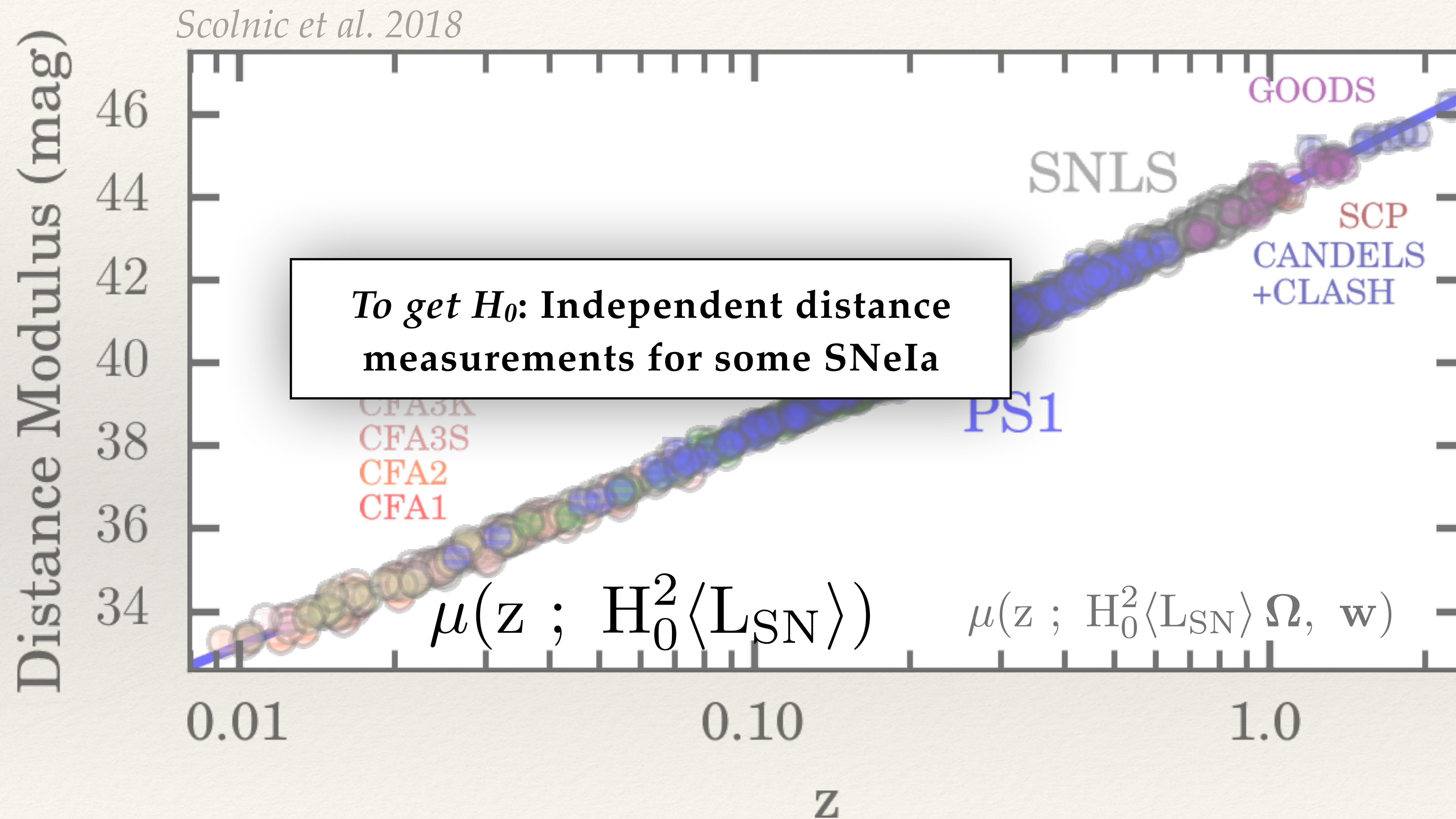


Model dependent

Type Ia Cosmology

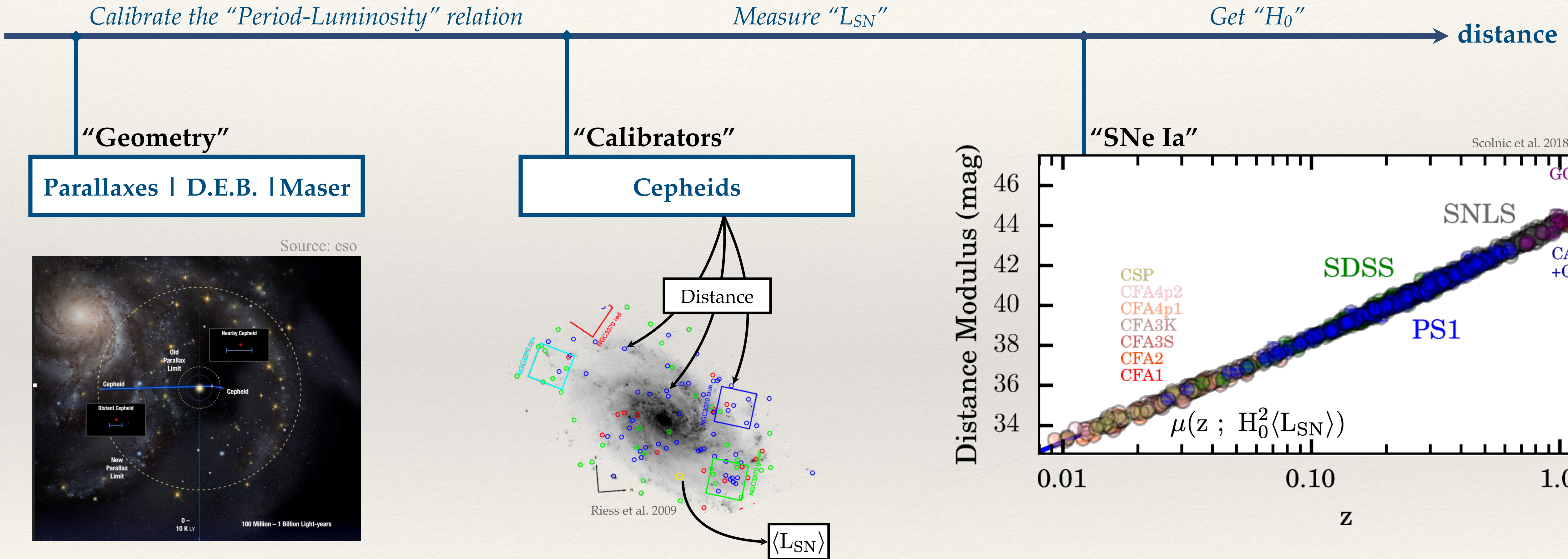


Type Ia Cosmology | Measuring H_0



Direct Distance Ladder | *SHOES*

Get independent distances for SNe Ia



SHOES

Geometrical Distances

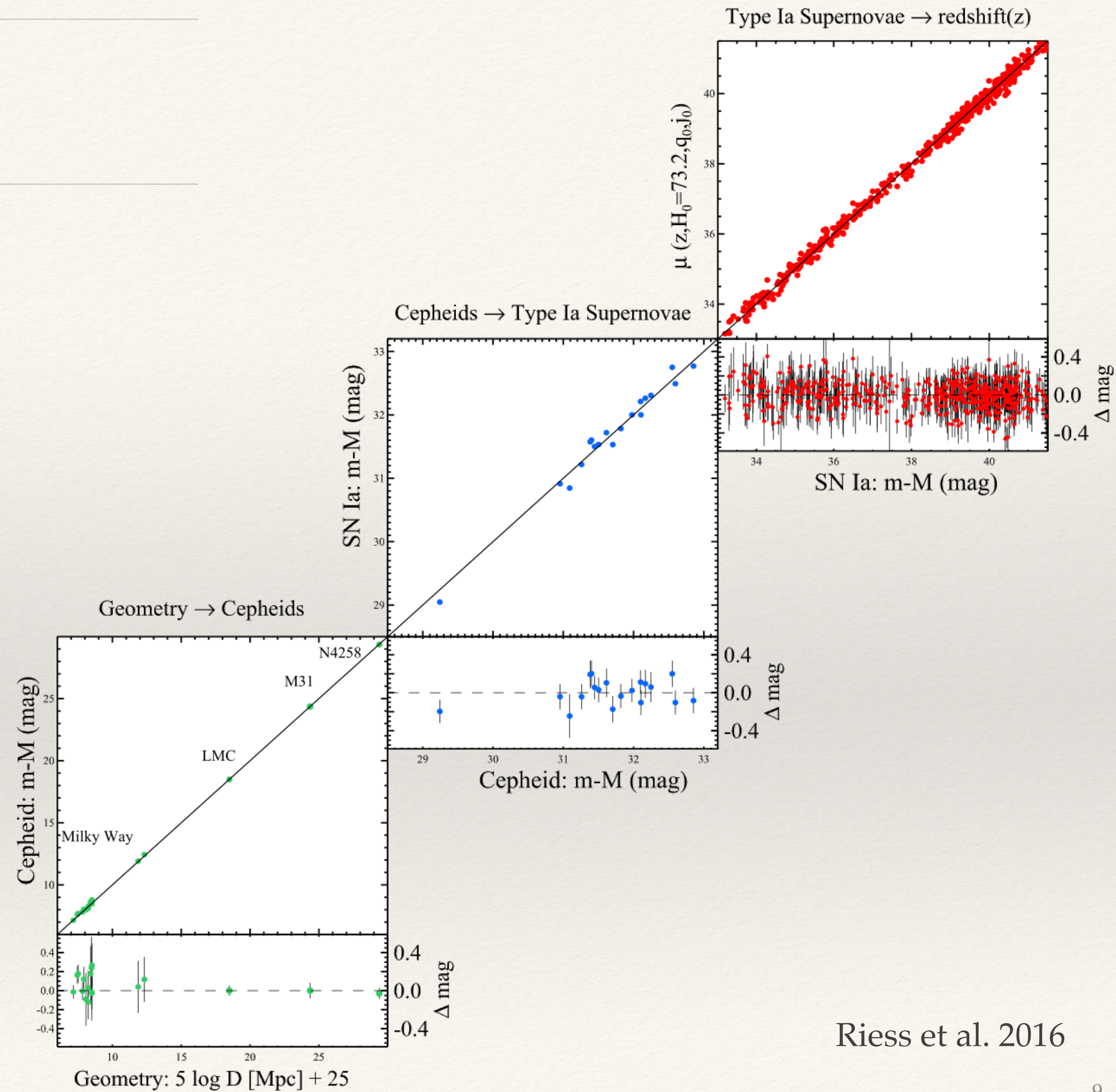
Cepheids

Type Ia Supernovae

10 Mpc

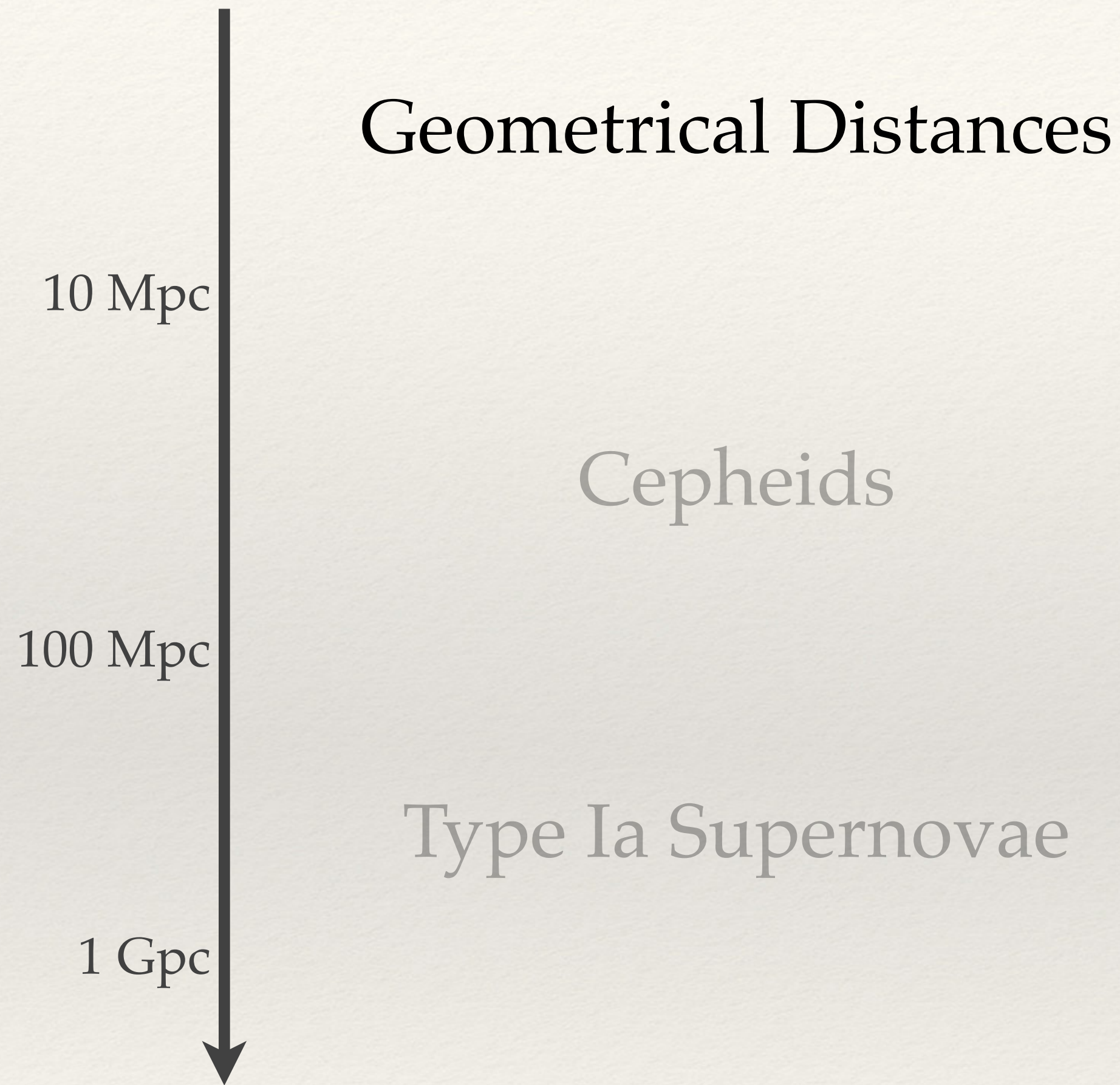
100 Mpc

1 Gpc

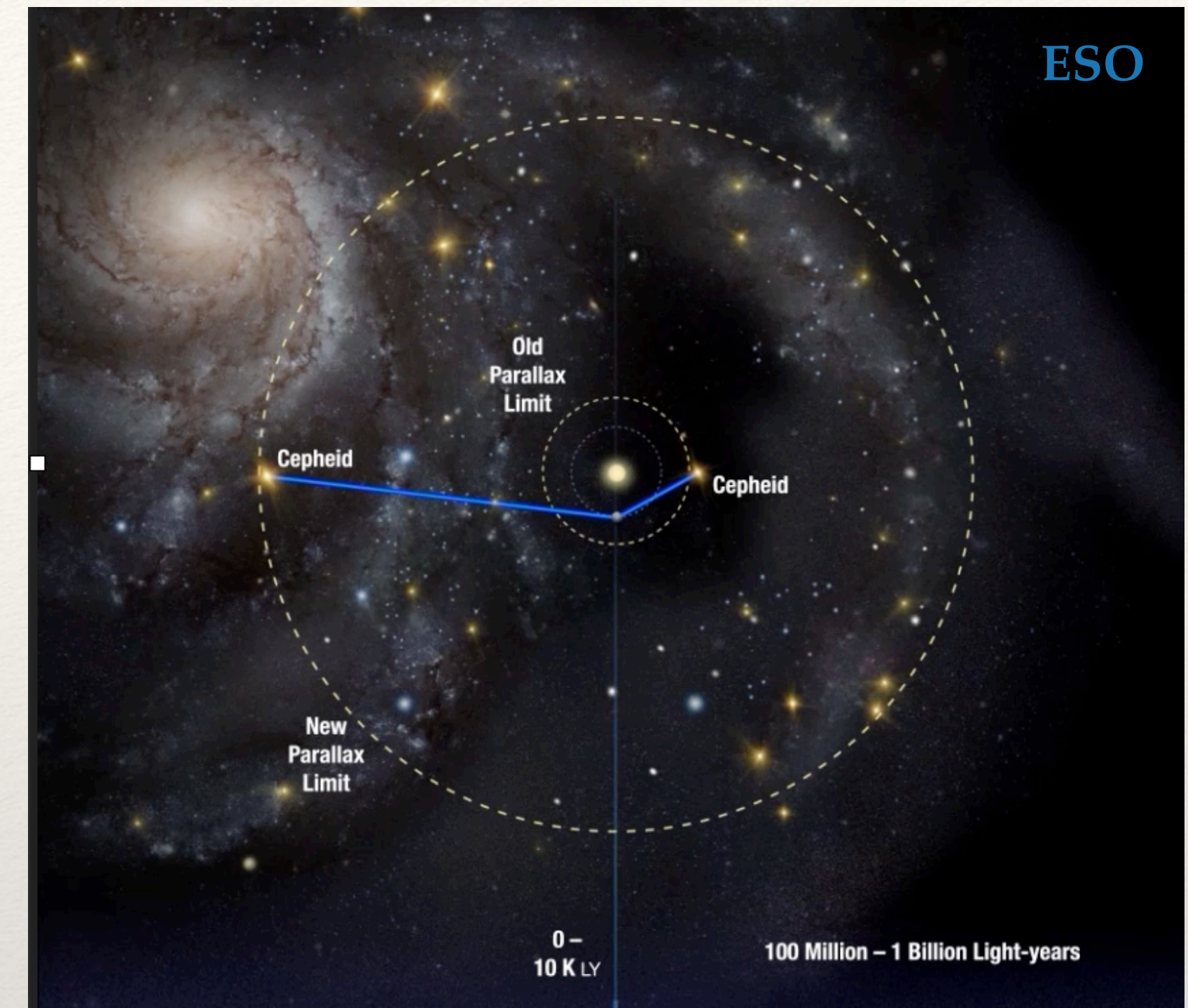


Riess et al. 2016

SHOES



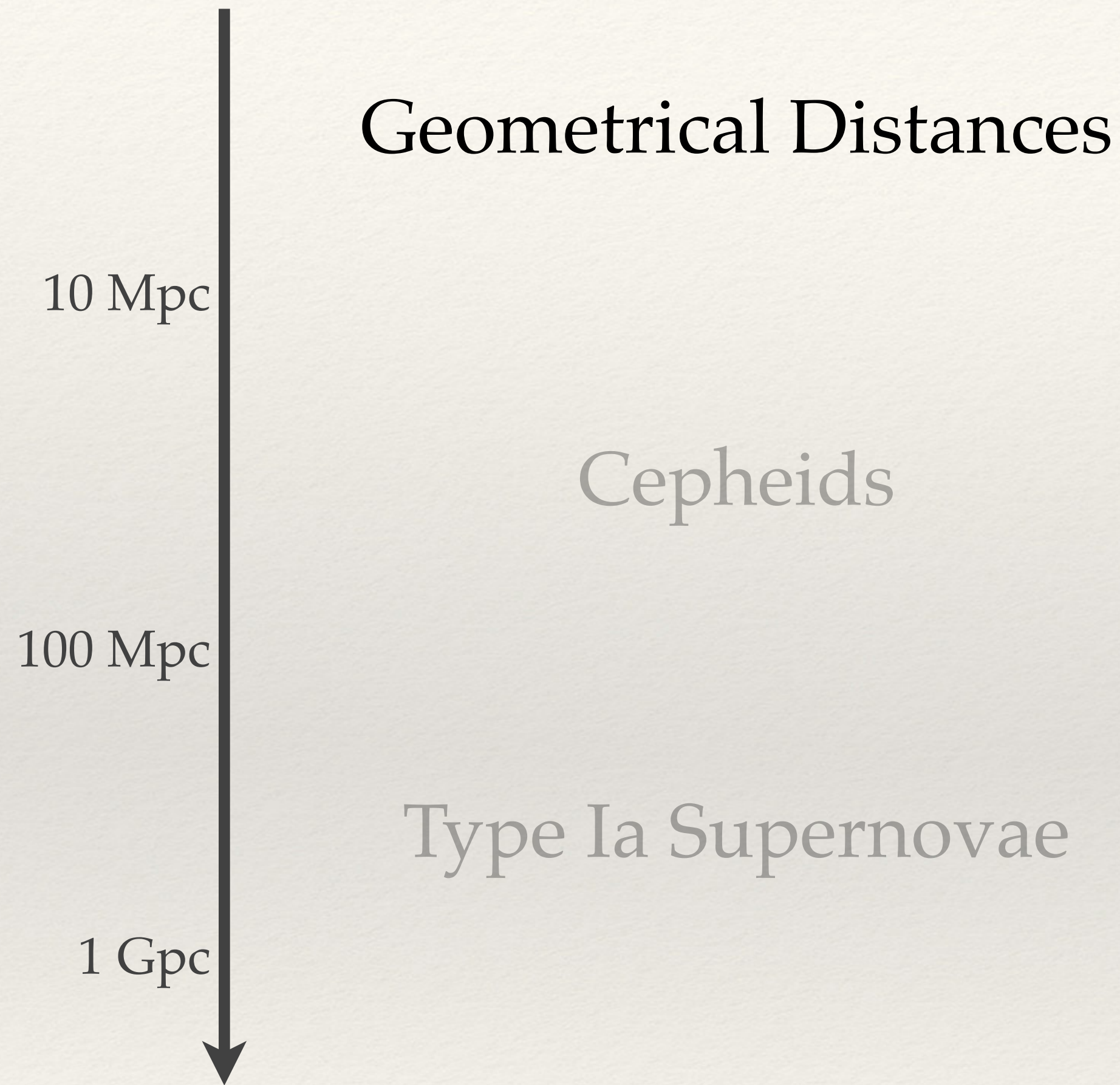
Parallaxes (*Milky Way*)



Detached Eclipsing Binaries (*LMC & M31*)

Mega Maser (*NGC4258*)

SHOES



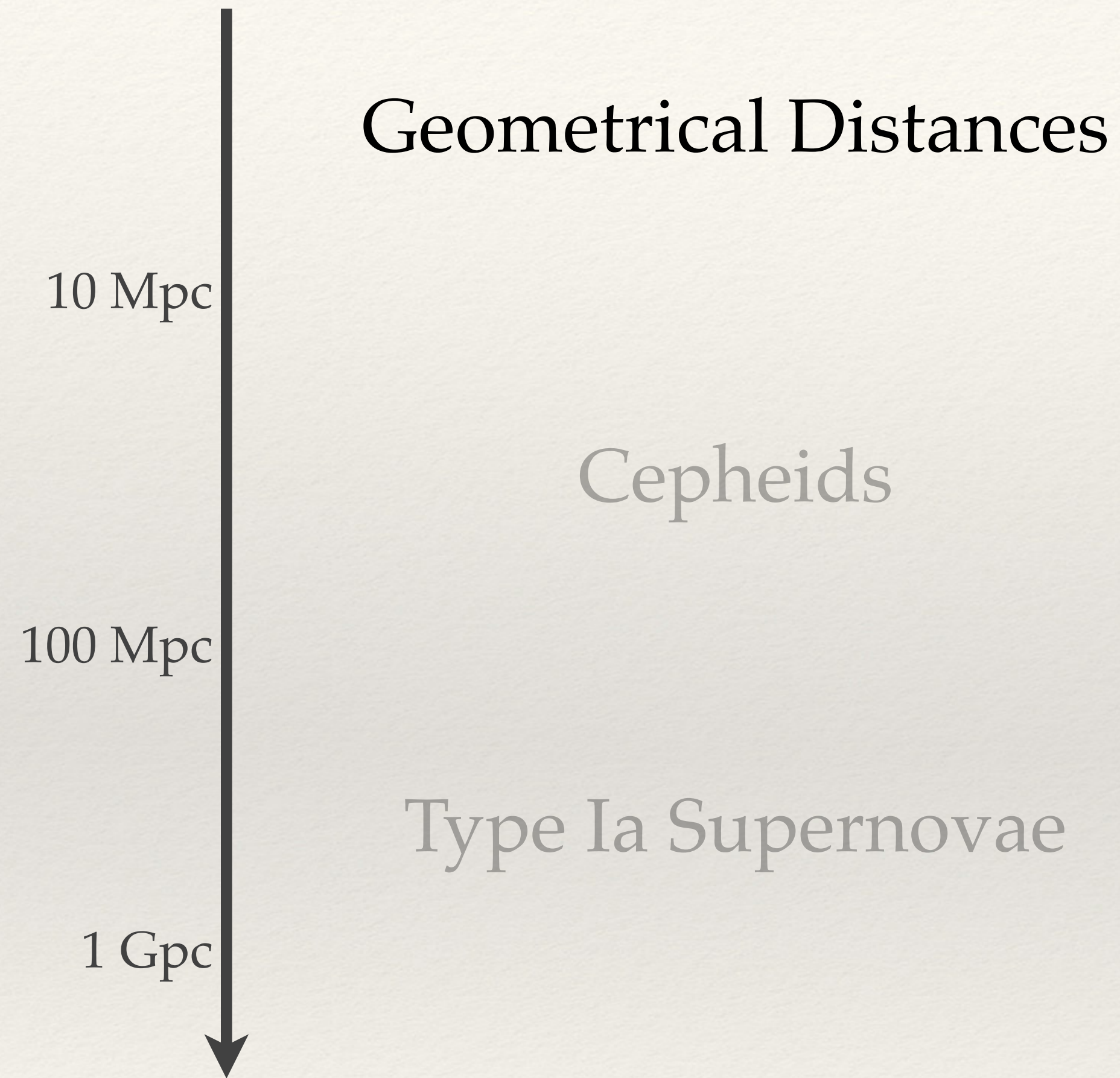
Parallaxes (*Milky Way*)

Detached Eclipsing Binaries
(*LMC & M31*)



Mega Maser (*NGC4258*)

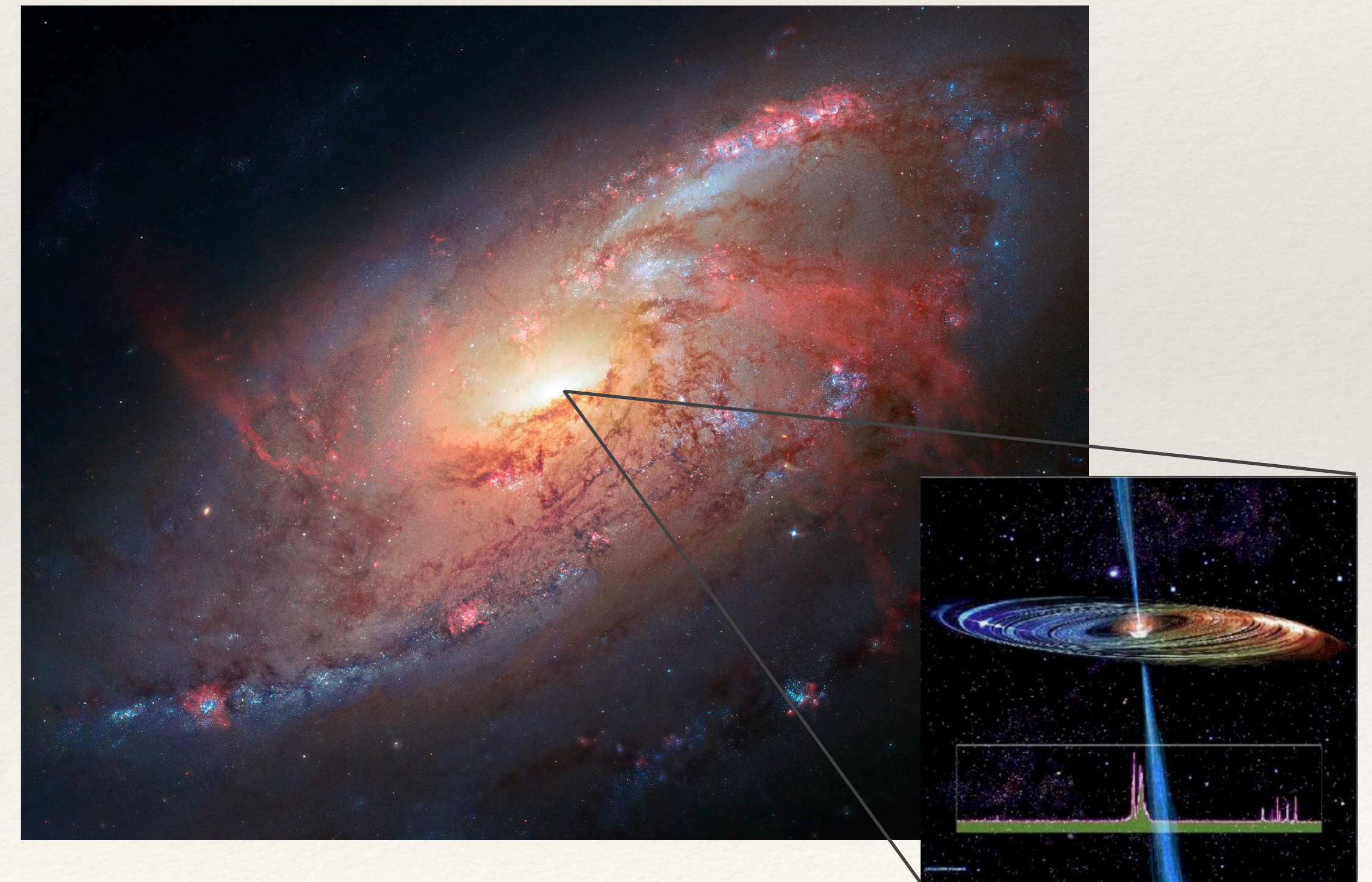
SHOES



Parallaxes (*Milky Way*)

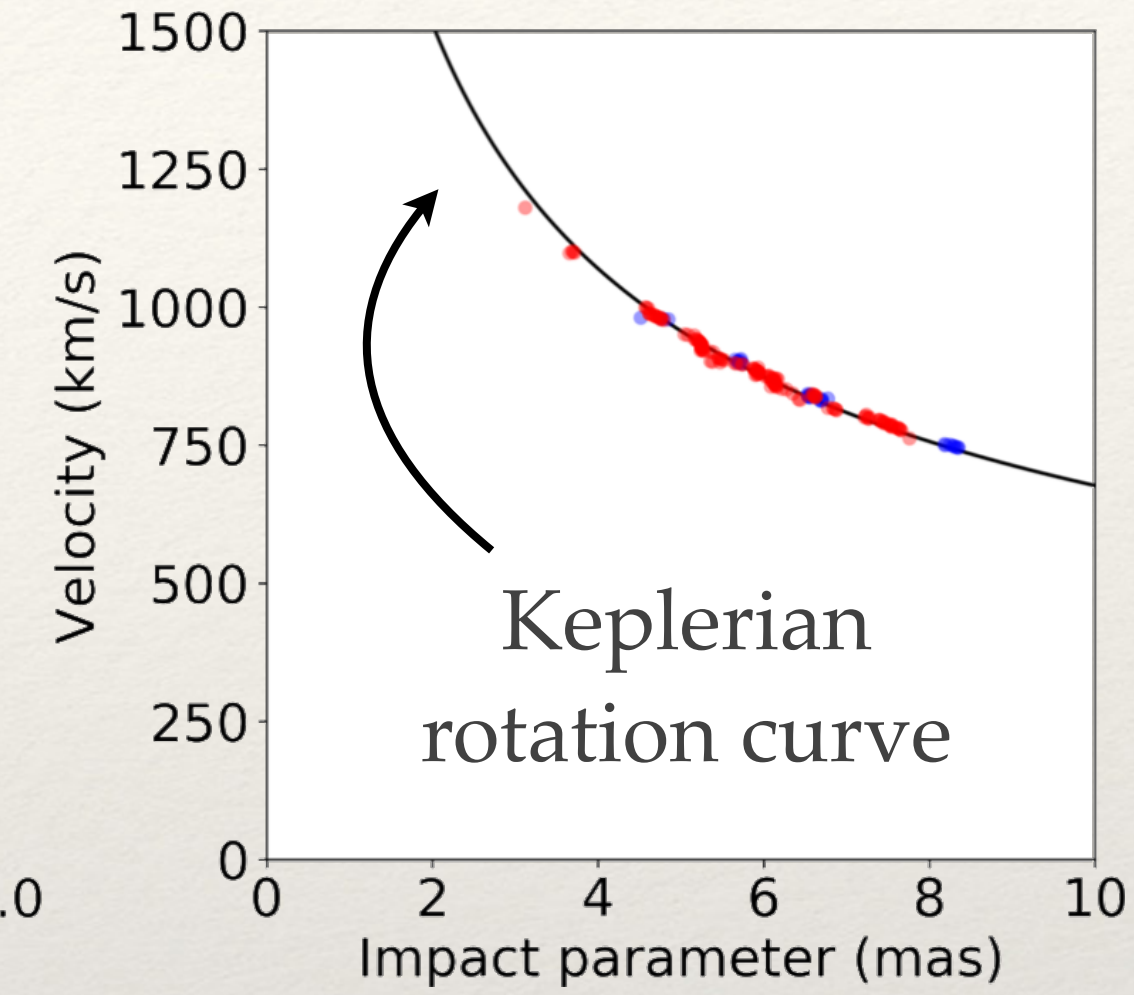
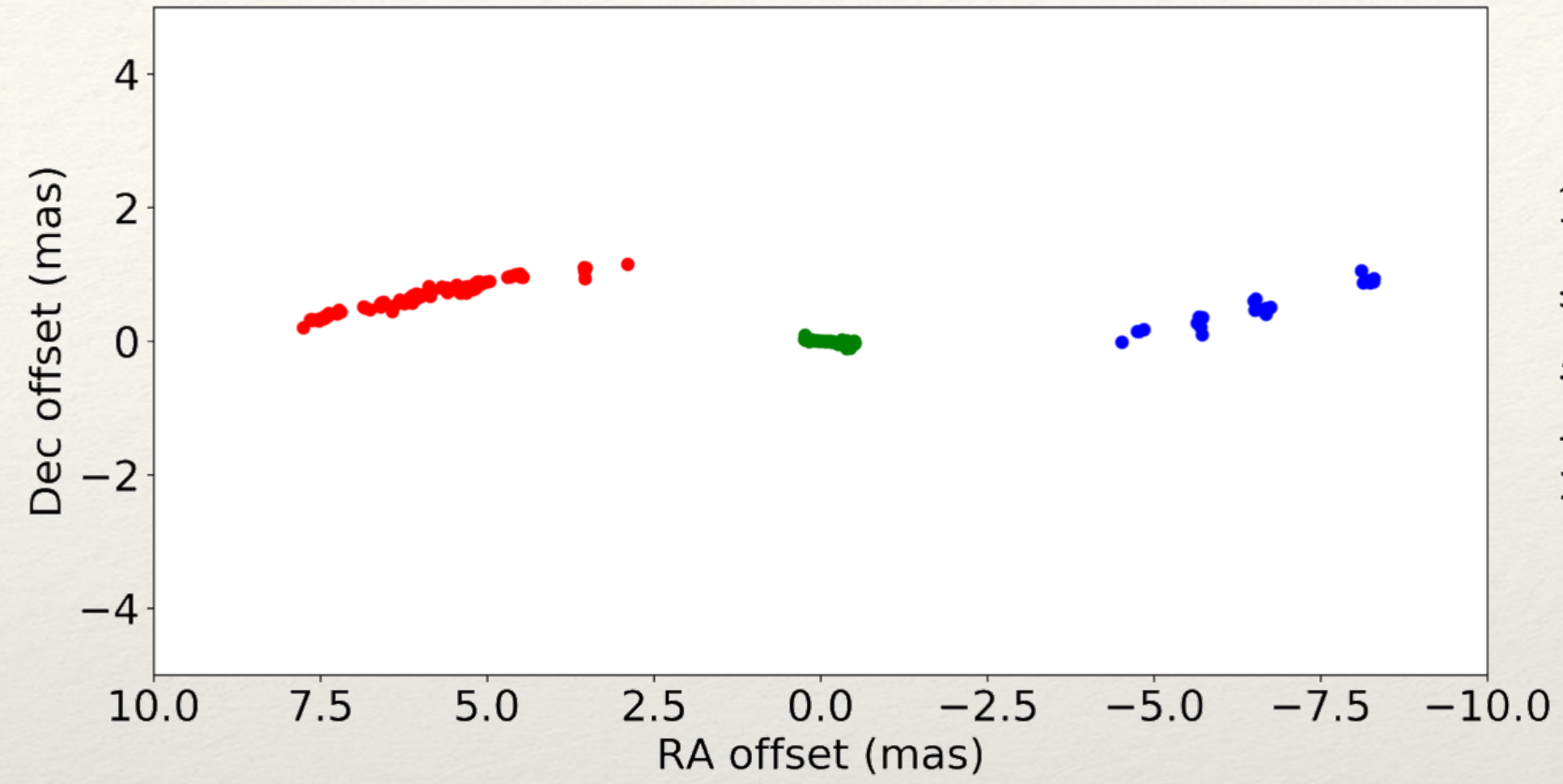
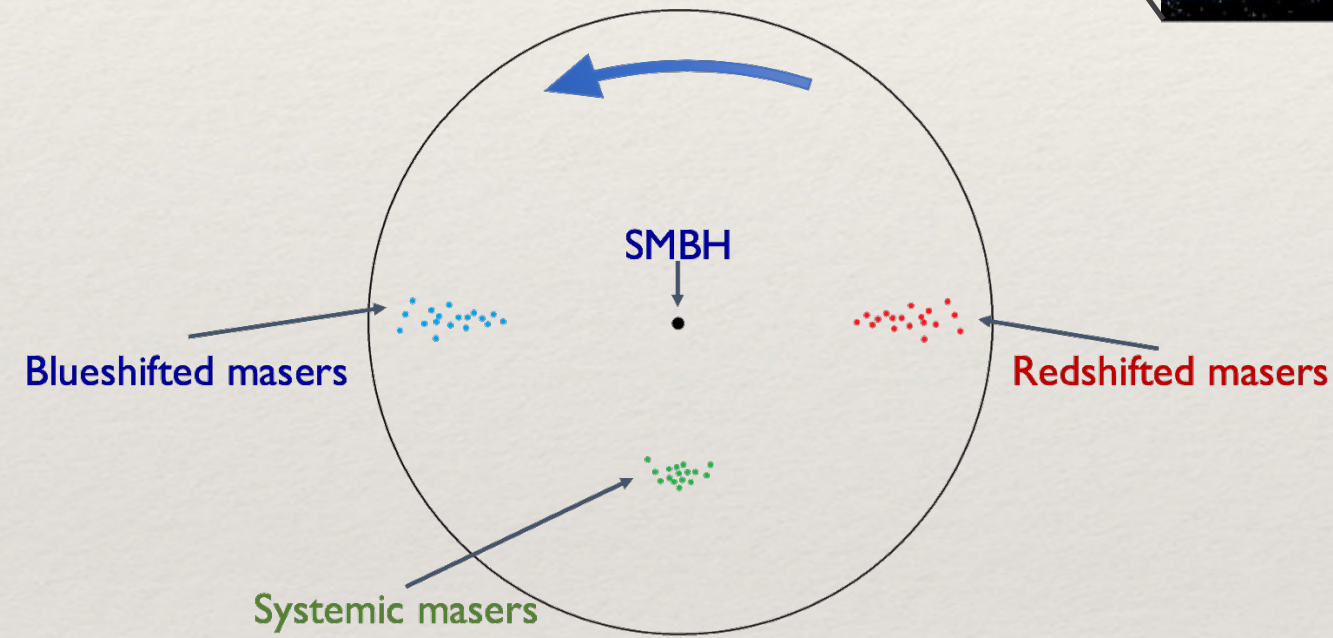
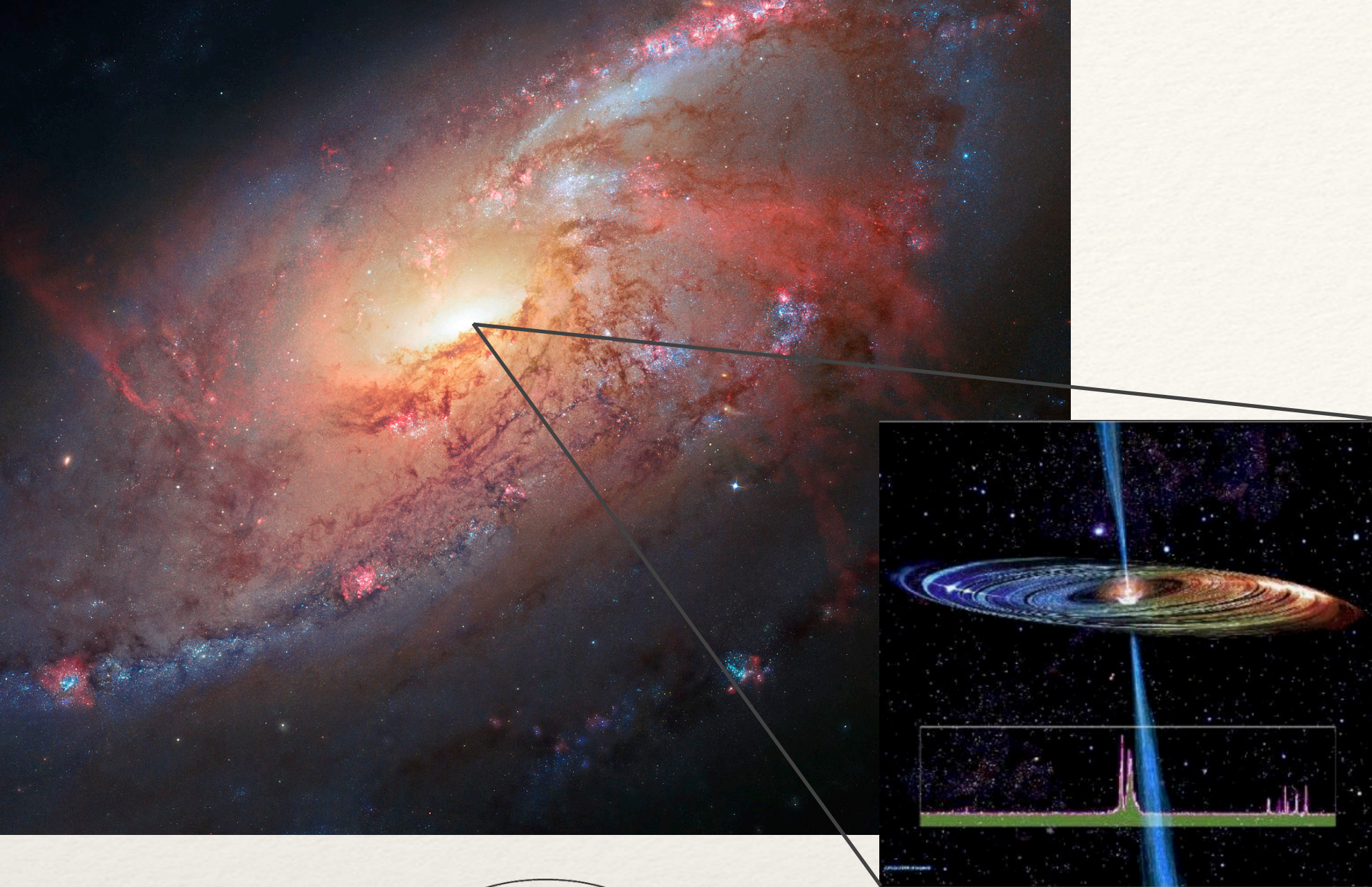
Detached Eclipsing Binaries
(*LMC & M31*)

Mega Maser (*NGC4258*)



Maser

credit: Dom Pesce



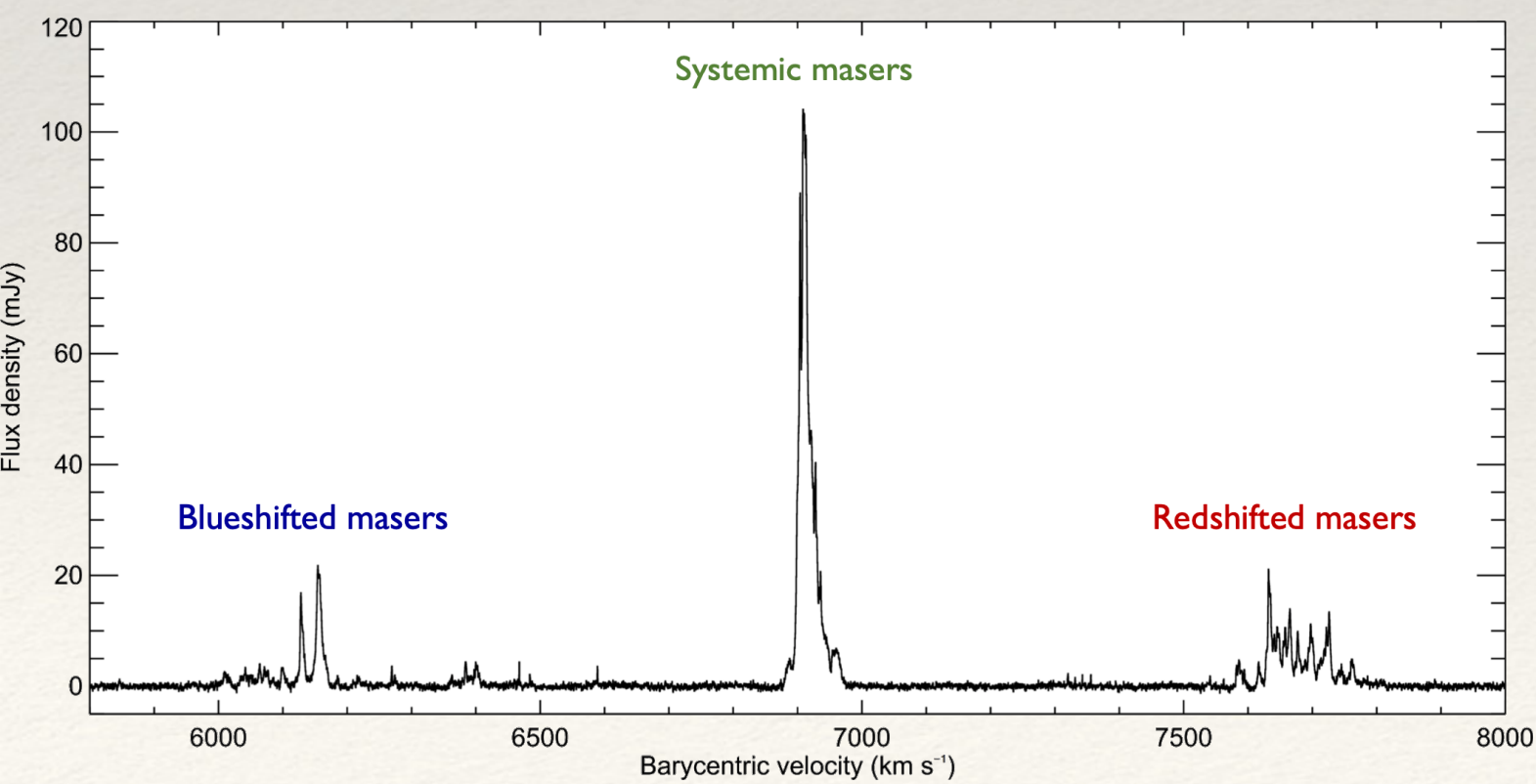
Consider a masing cloud on a circular orbit at (angular) radius θ_r around a central SMBH of mass M and situated at an azimuthal angle φ with respect to the line of sight

Observed (on-sky) position: $\theta = \theta_r \sin(\varphi)$
 Observed (line-of-sight) velocity: $v = v_r \sin(\varphi)$
 Observed (line-of-sight) acceleration: $a = a_r \cos(\varphi)$

$$v_r = \sqrt{\frac{GM}{\theta_r D}} \quad a_r = \frac{v_r^2}{\theta_r D} = \frac{GM}{\theta_r^3 D^2}$$

○ = measured

○ = fit



SHOES

Geometrical Distances

10 Mpc

Cepheids

100 Mpc

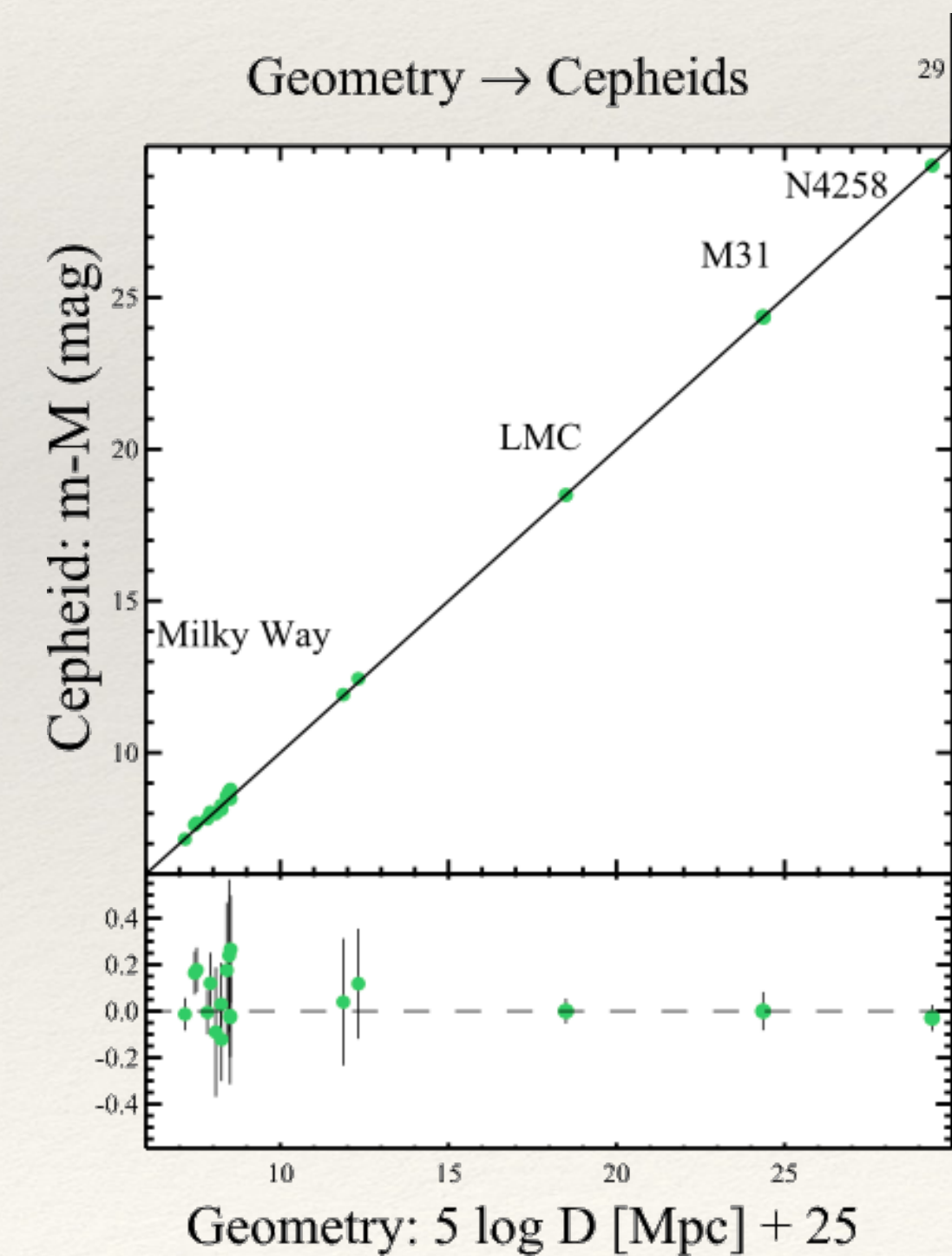
Type Ia Supernovae

1 Gpc

Parallaxes (*Milky Way*)

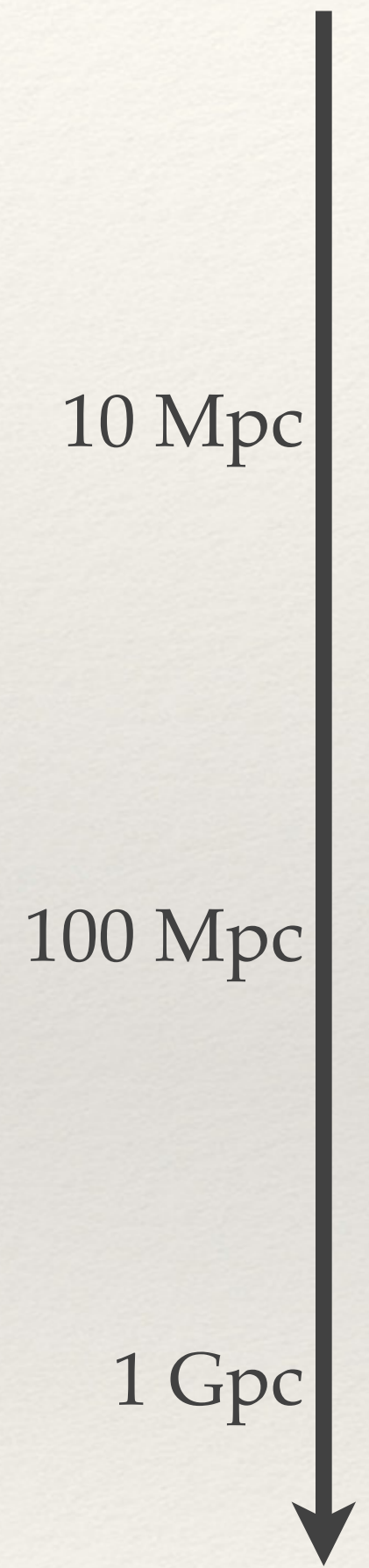
Detached Eclipsing Binaries
(*LMC & M31*)

Mega Maser (*NGC4258*)



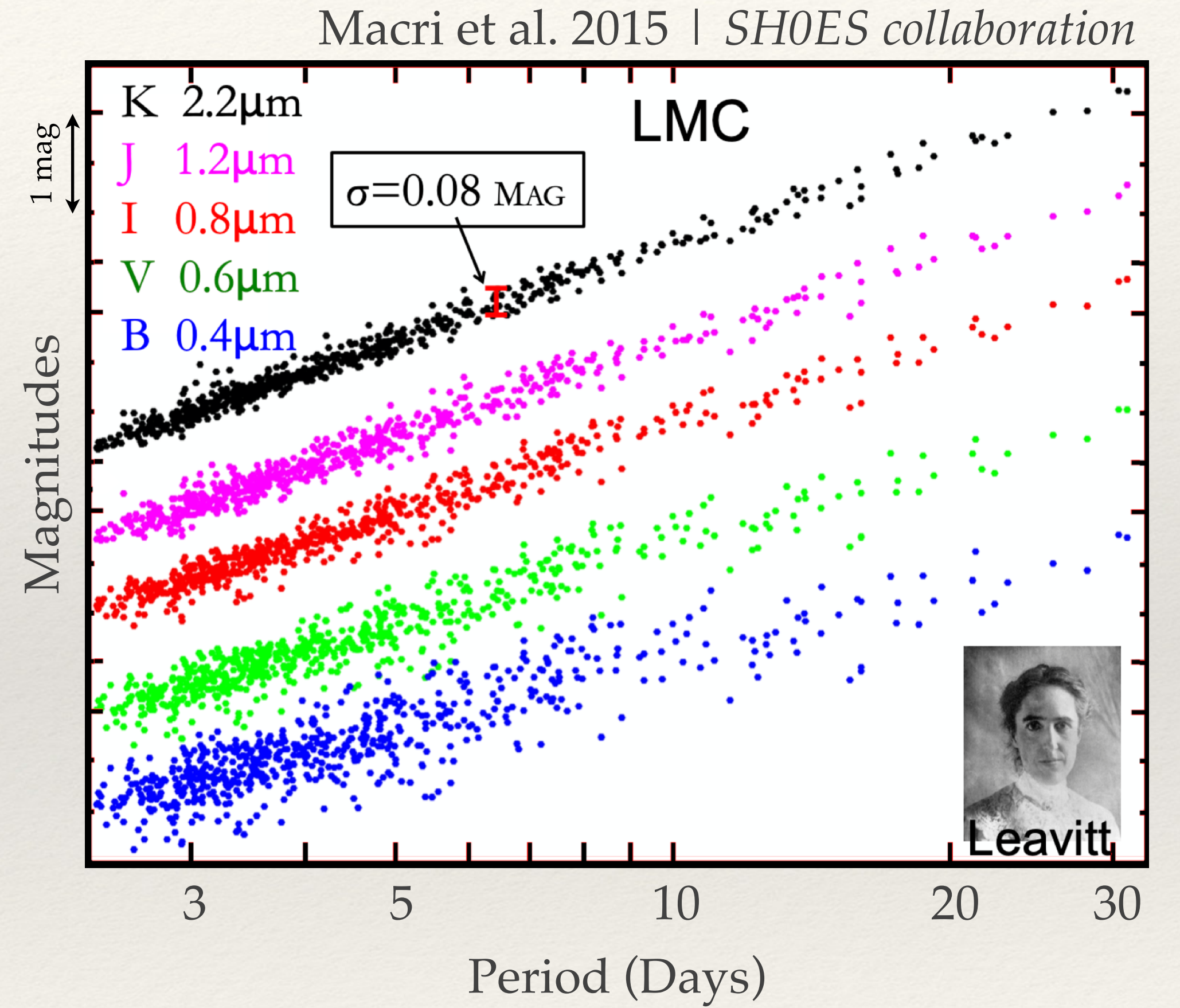
SHOES

Geometrical Distances



Cepheids

Type Ia Supernovae



SHOES

Geometrical Distances

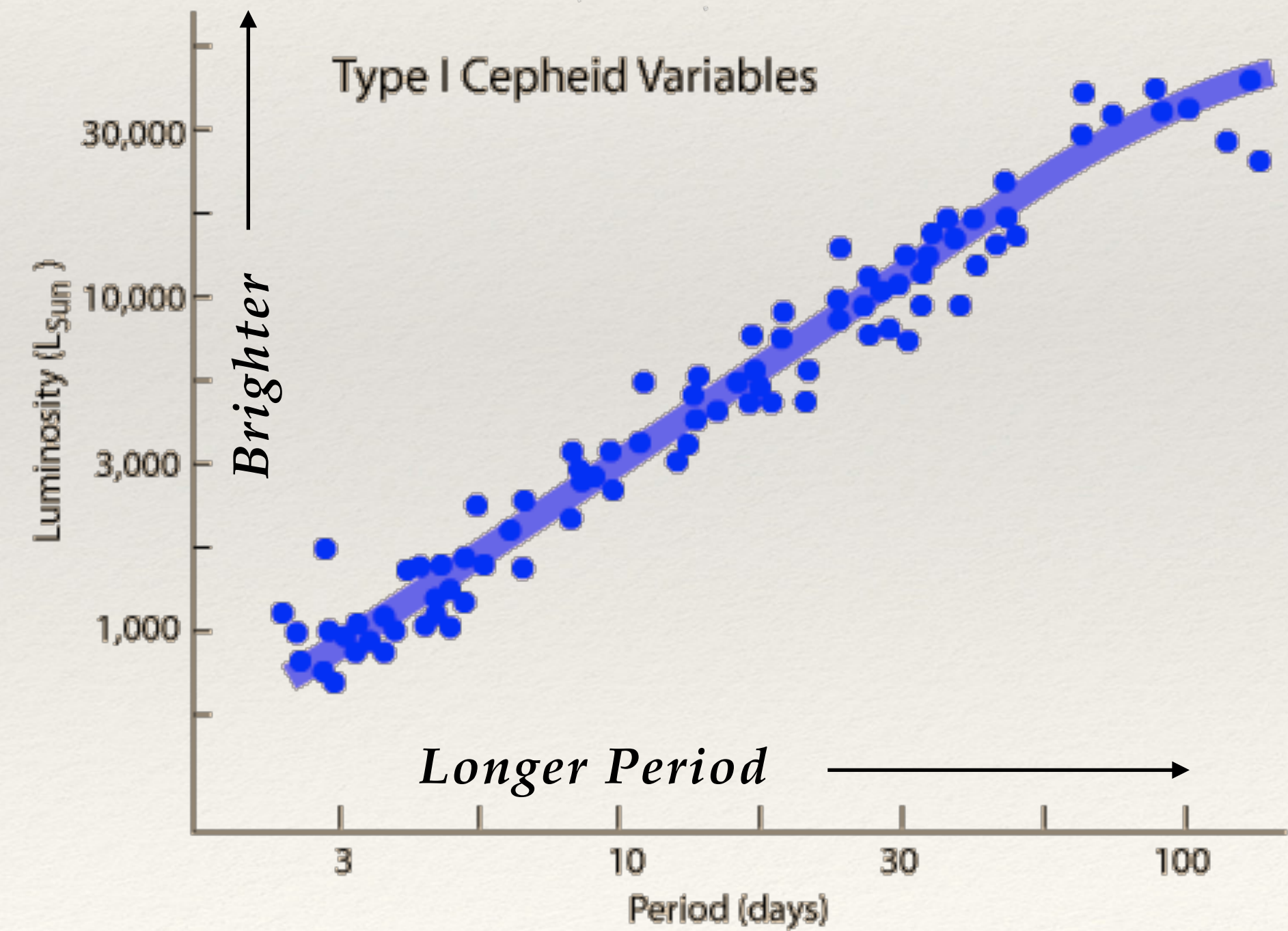
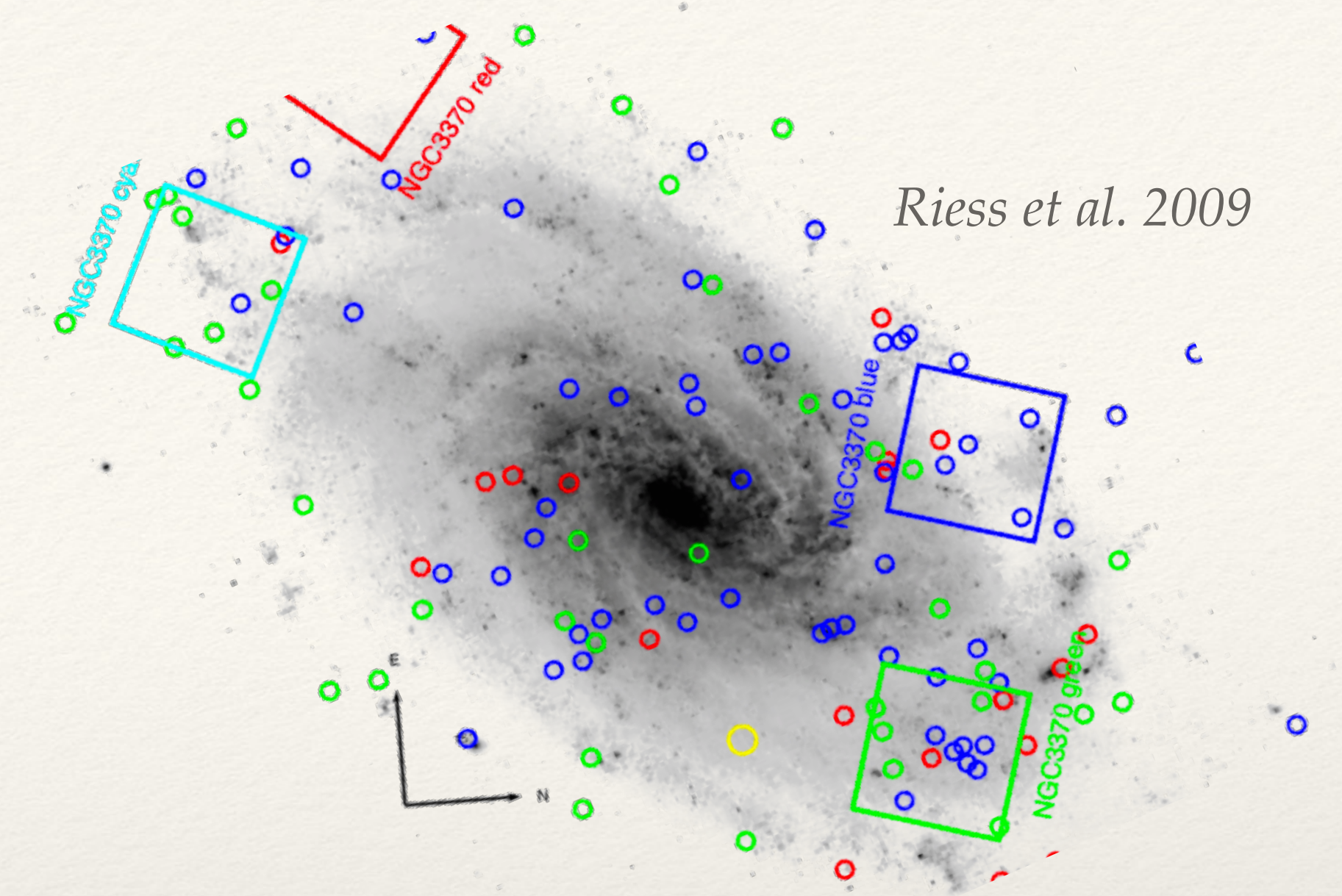
10 Mpc

Cepheids

100 Mpc

Type Ia Supernovae

1 Gpc



SHOES

SHOES collaboration

Geometrical Distances

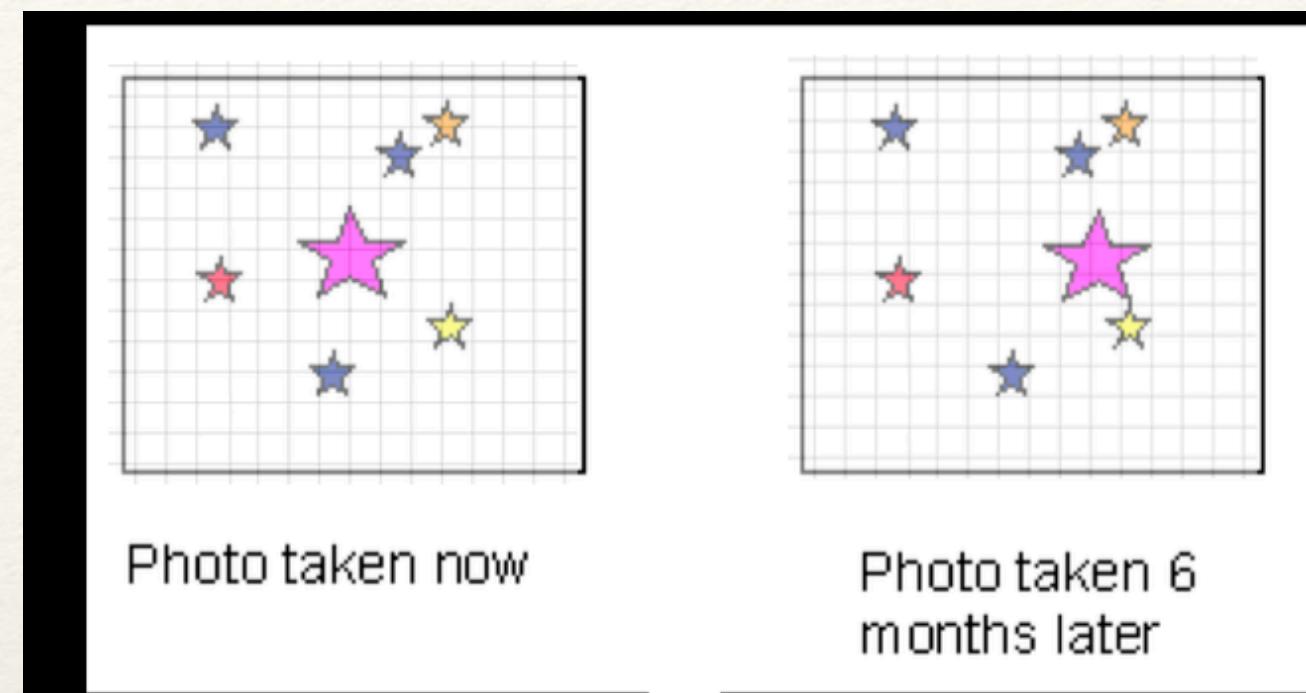
10 Mpc

Cepheids

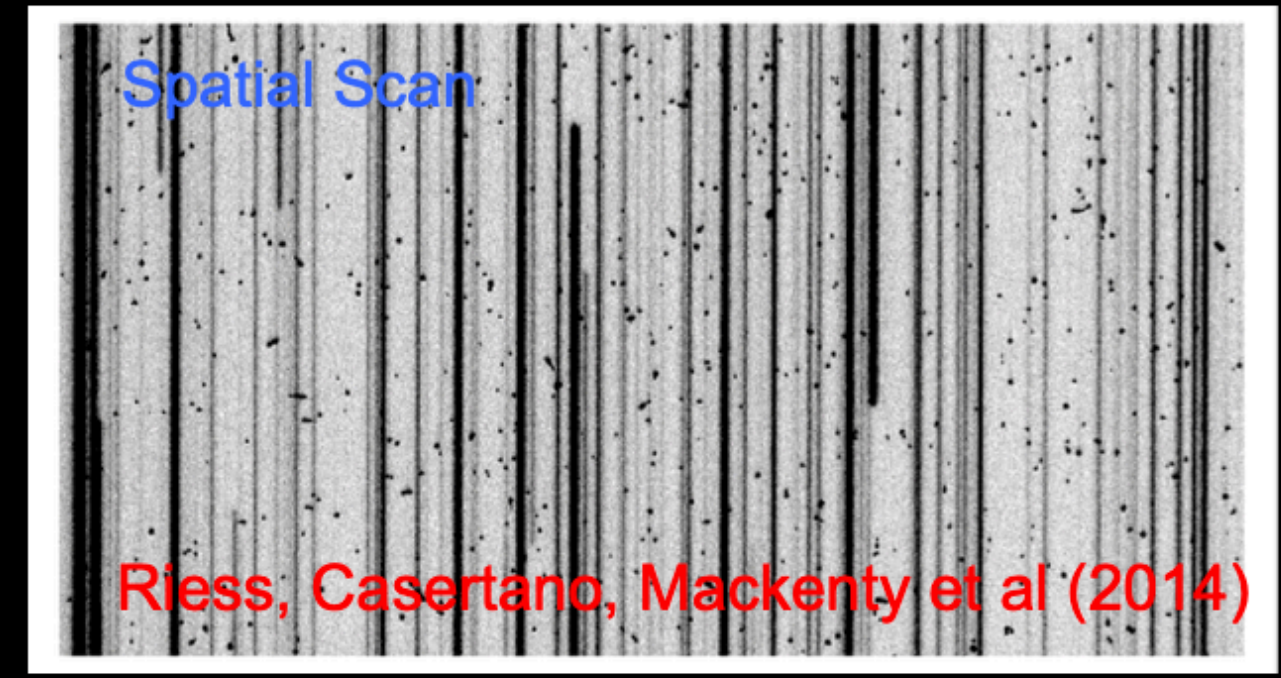
100 Mpc

Type Ia Supernovae

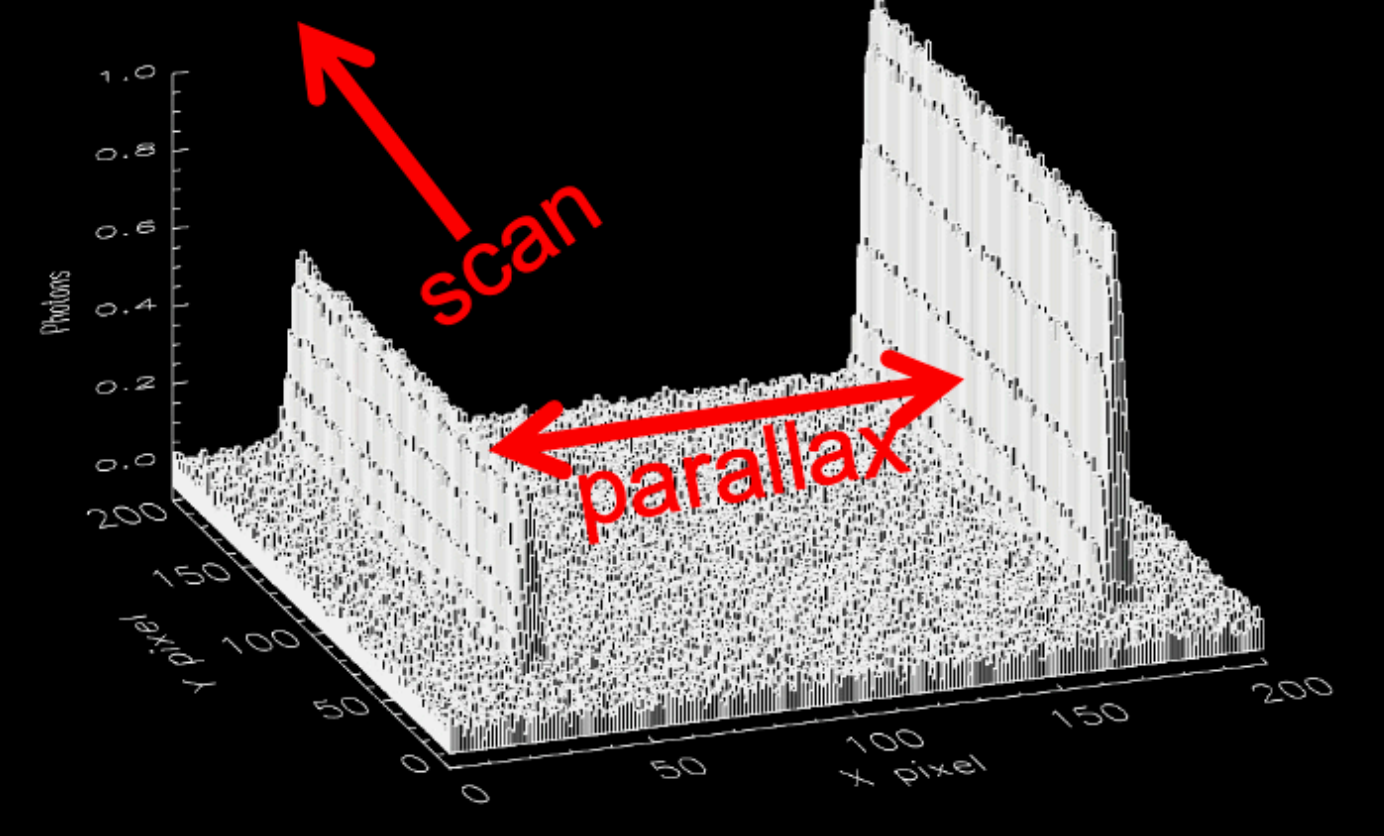
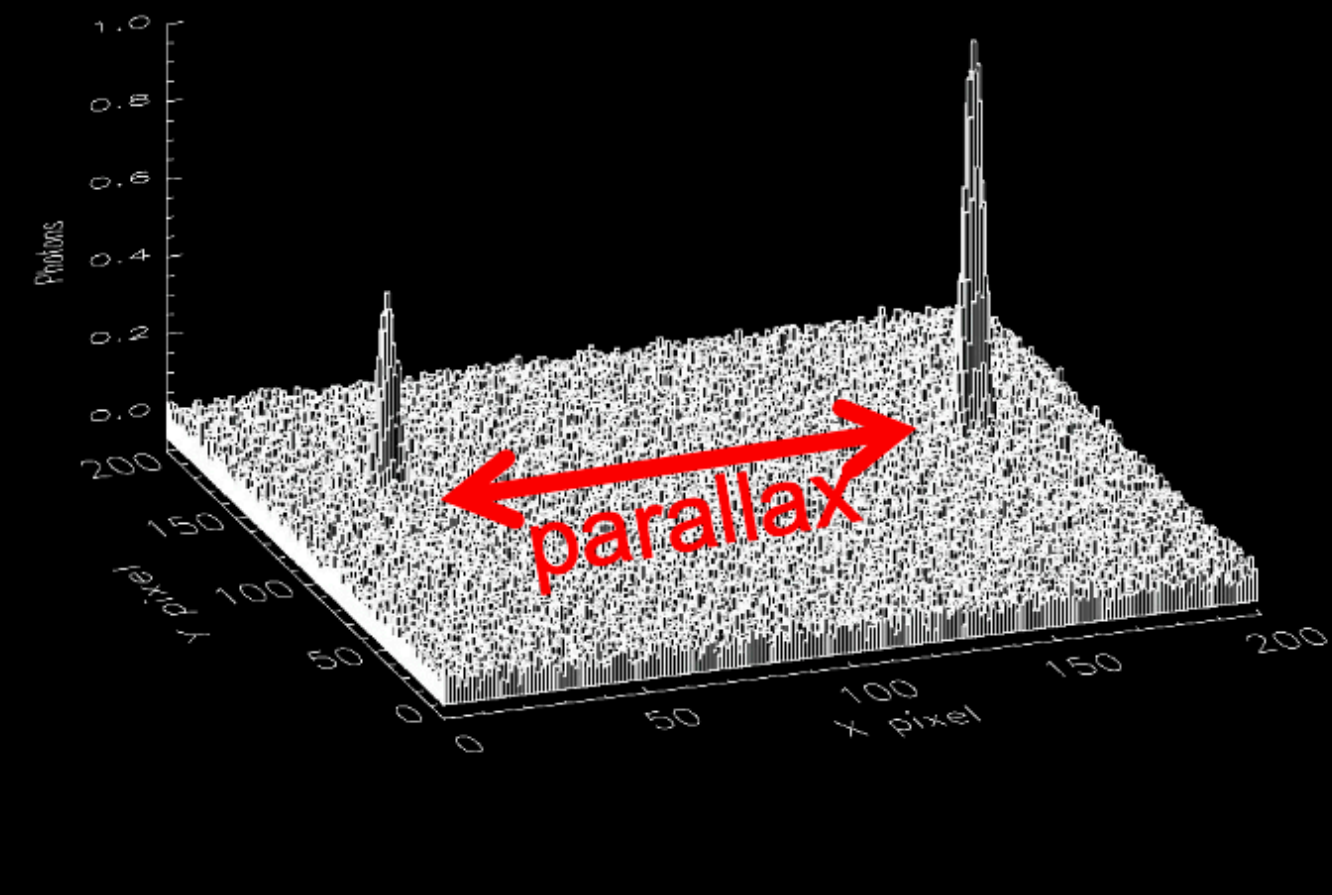
1 Gpc



Imaging: astrometry $\sigma_\theta = 0.01$ pix
HST: 0.4mas, $\sim 1\sigma$ @ 2 kpc



Scanning, $\sigma_\theta = 0.01/\sqrt{N}$ samples pix

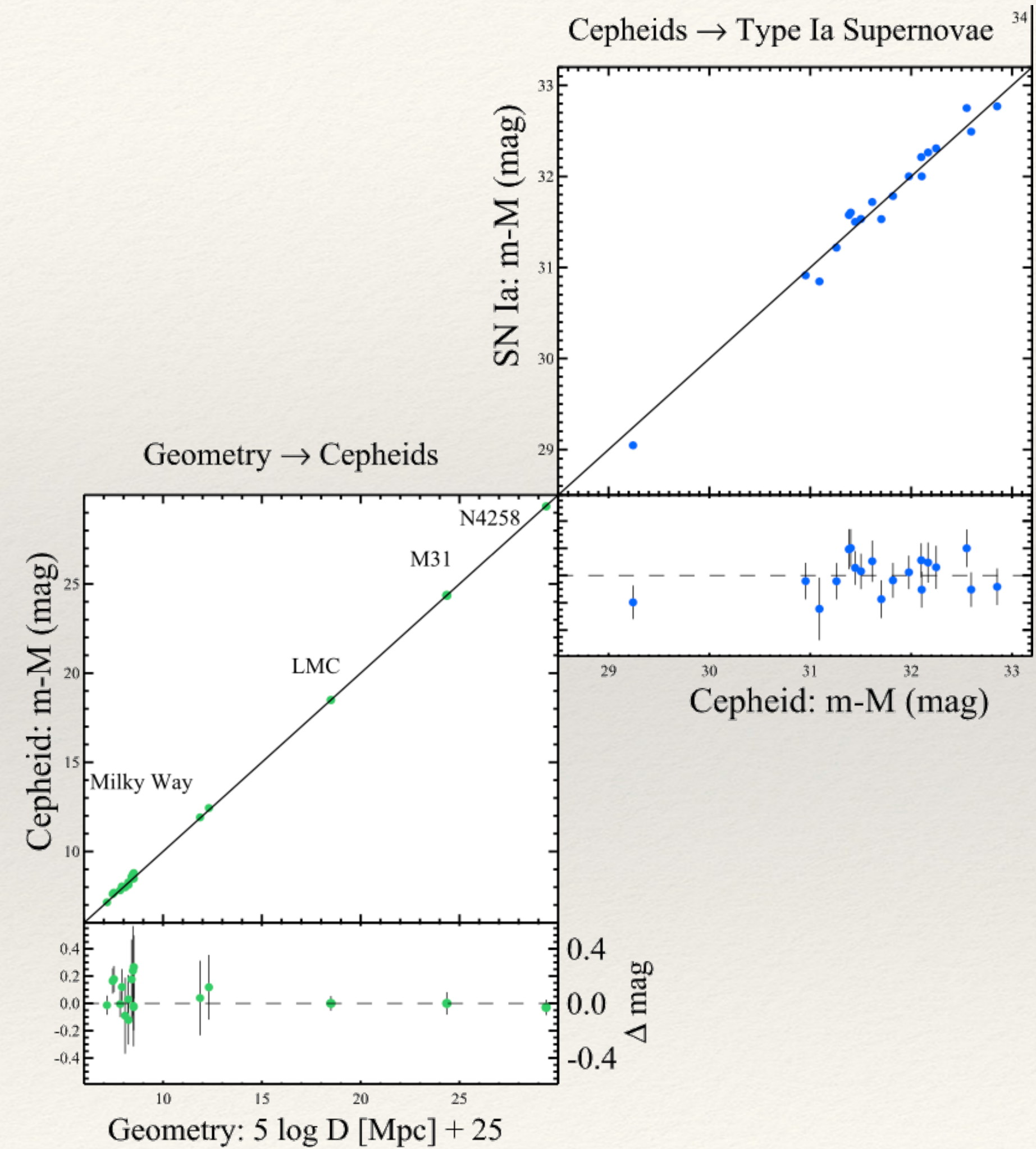


SHOES

Geometrical Distances

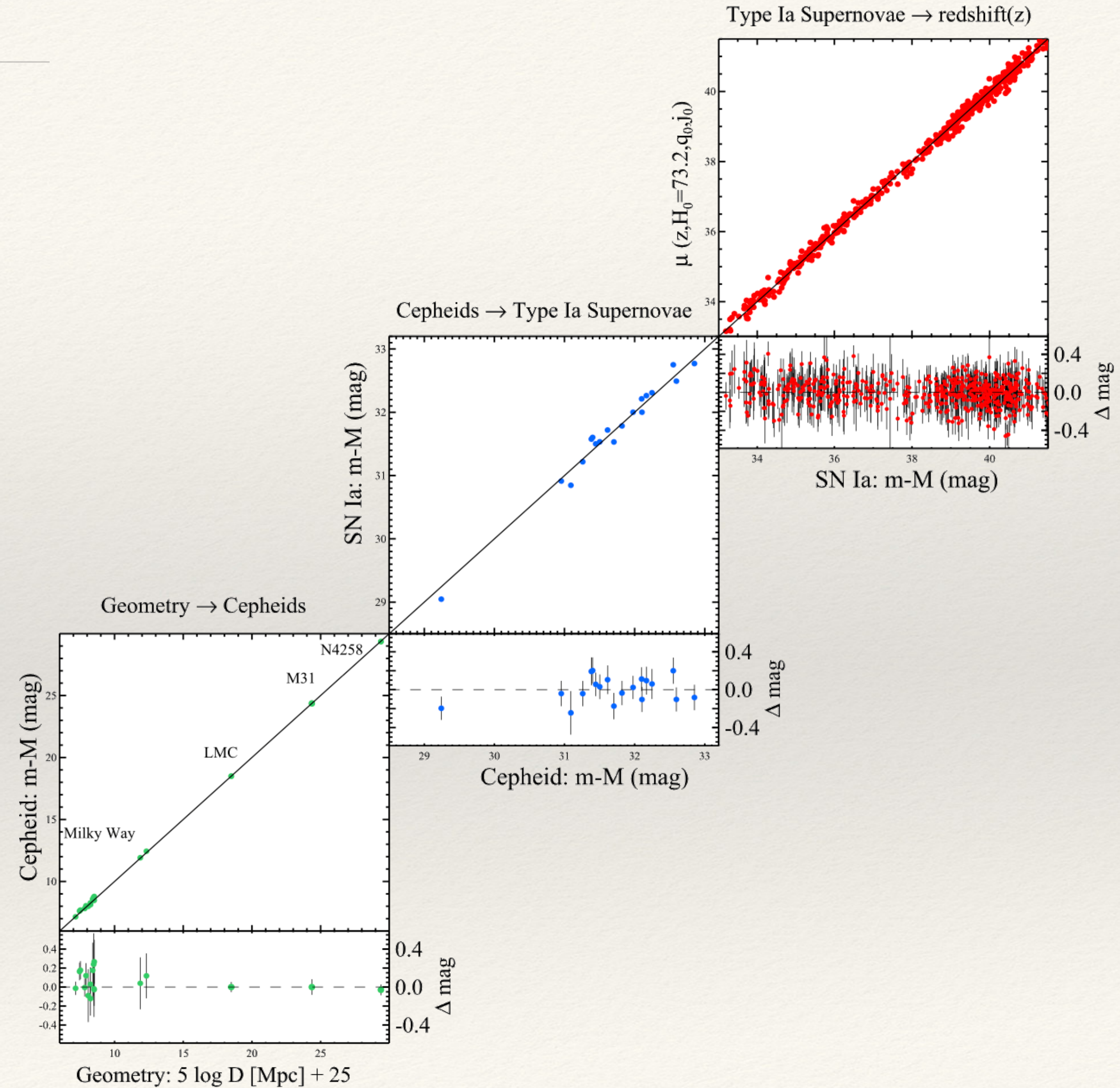
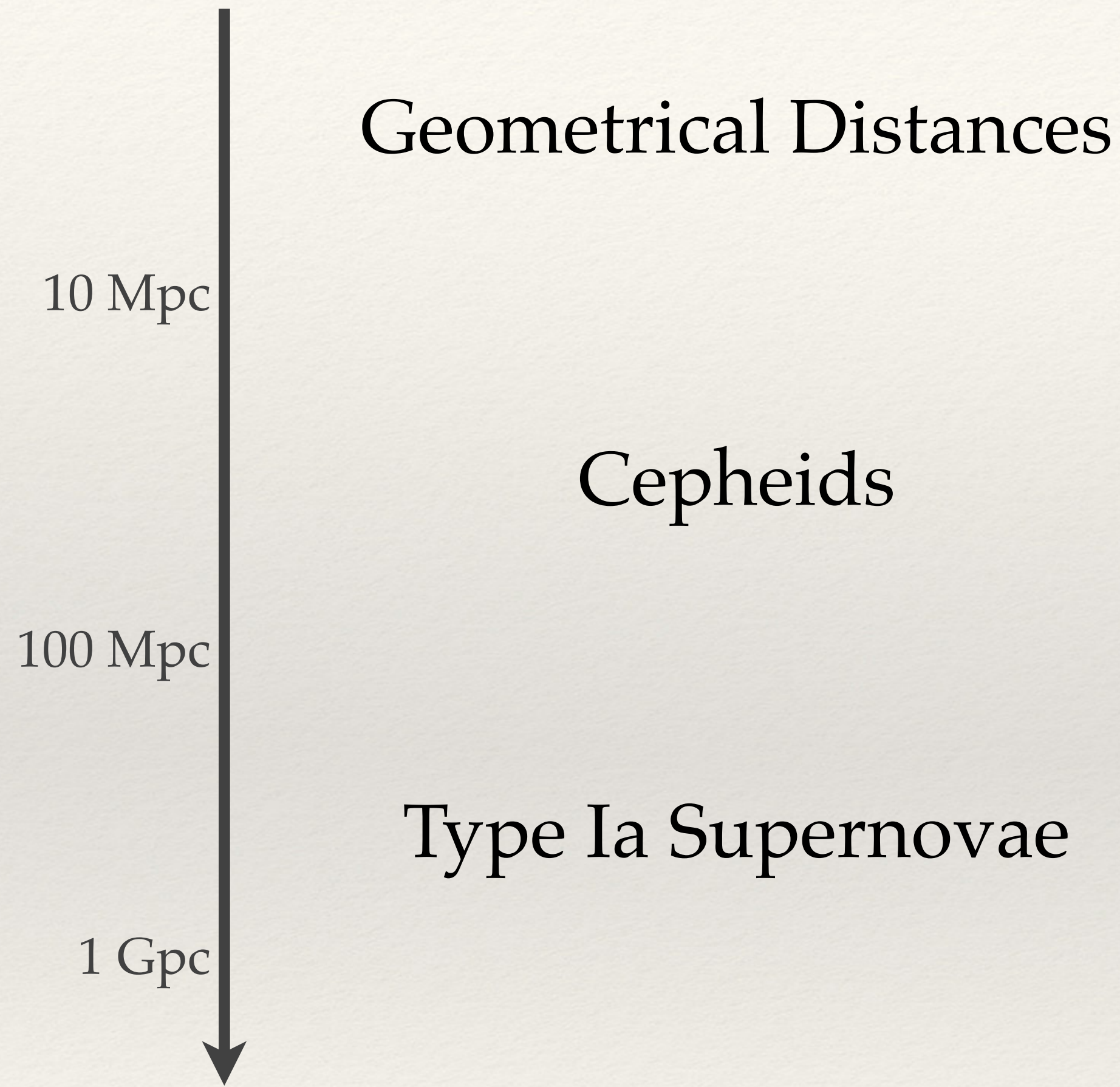
Cepheids

Type Ia Supernovae



SHOES

cf. Supernova Cosmology Class



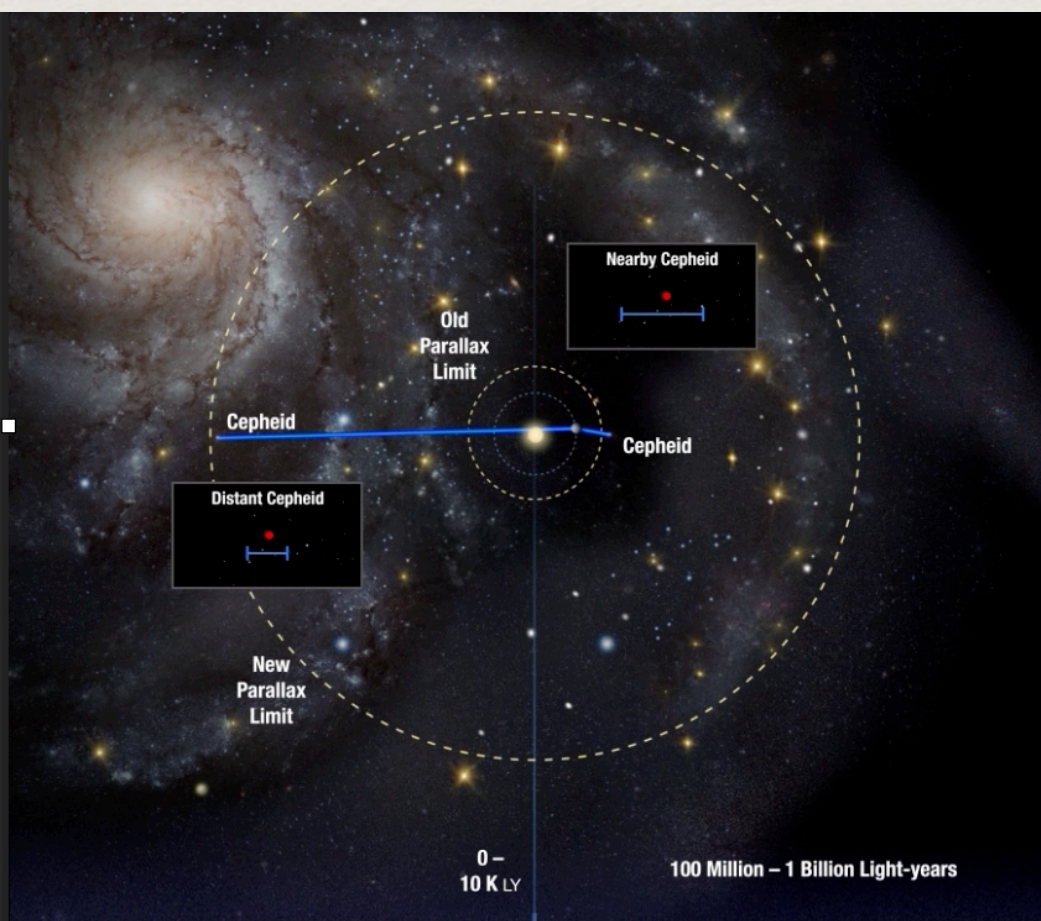
Direct Distance Ladder | *SHOES*

Get independent distances for SNe Ia



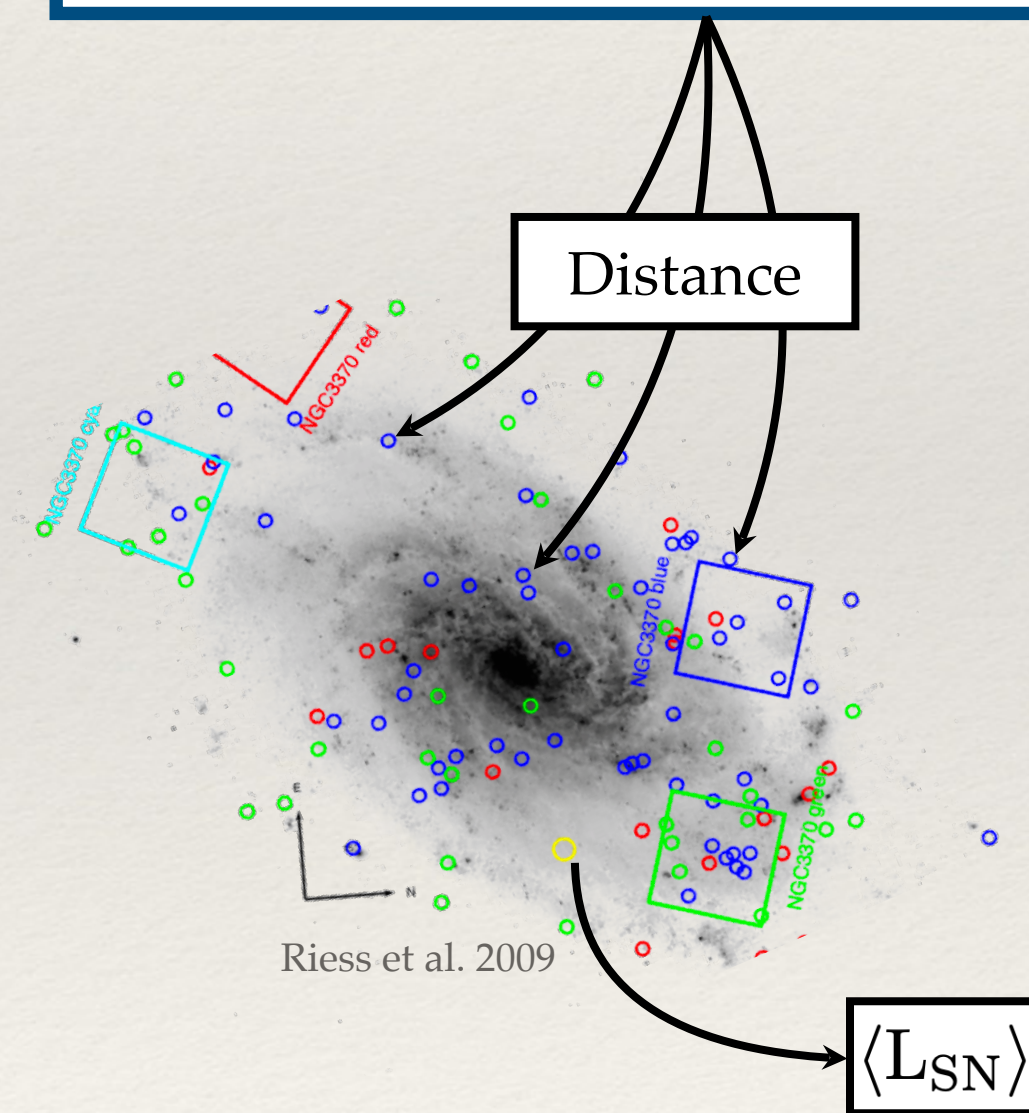
"Geometry"

Parallaxes | D.E.B. | Maser

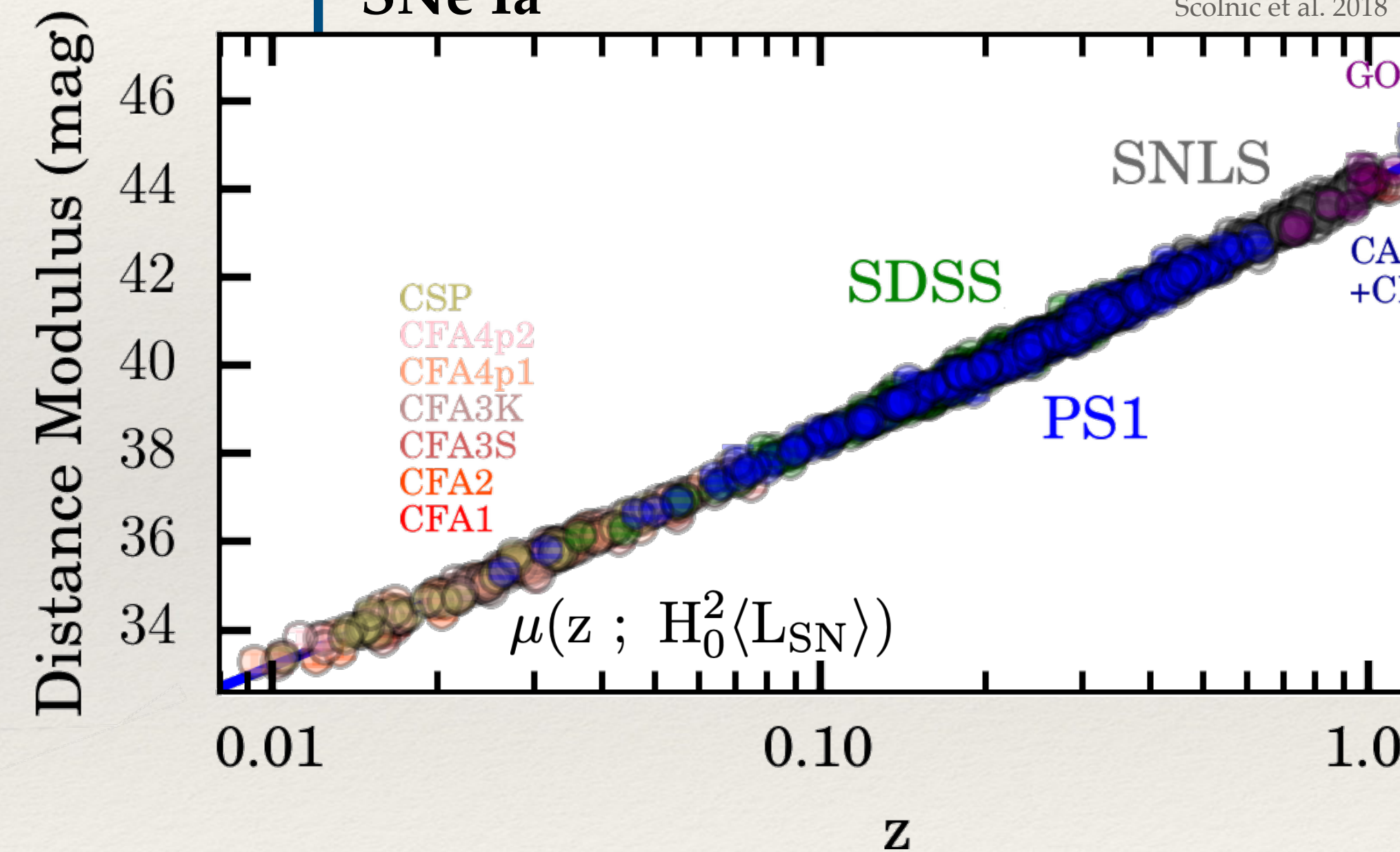


"Calibrators"

Cepheids



"SNe Ia"



Riess et al. 2022

$$H_0 = 73.0 \pm 1.0 \text{ km s}^{-1} \text{ Mpc}^{-1}$$

Modern Cosmology | H_0 Direct vs. Indirect Measurements

$$H_0 = d_l / v_h$$

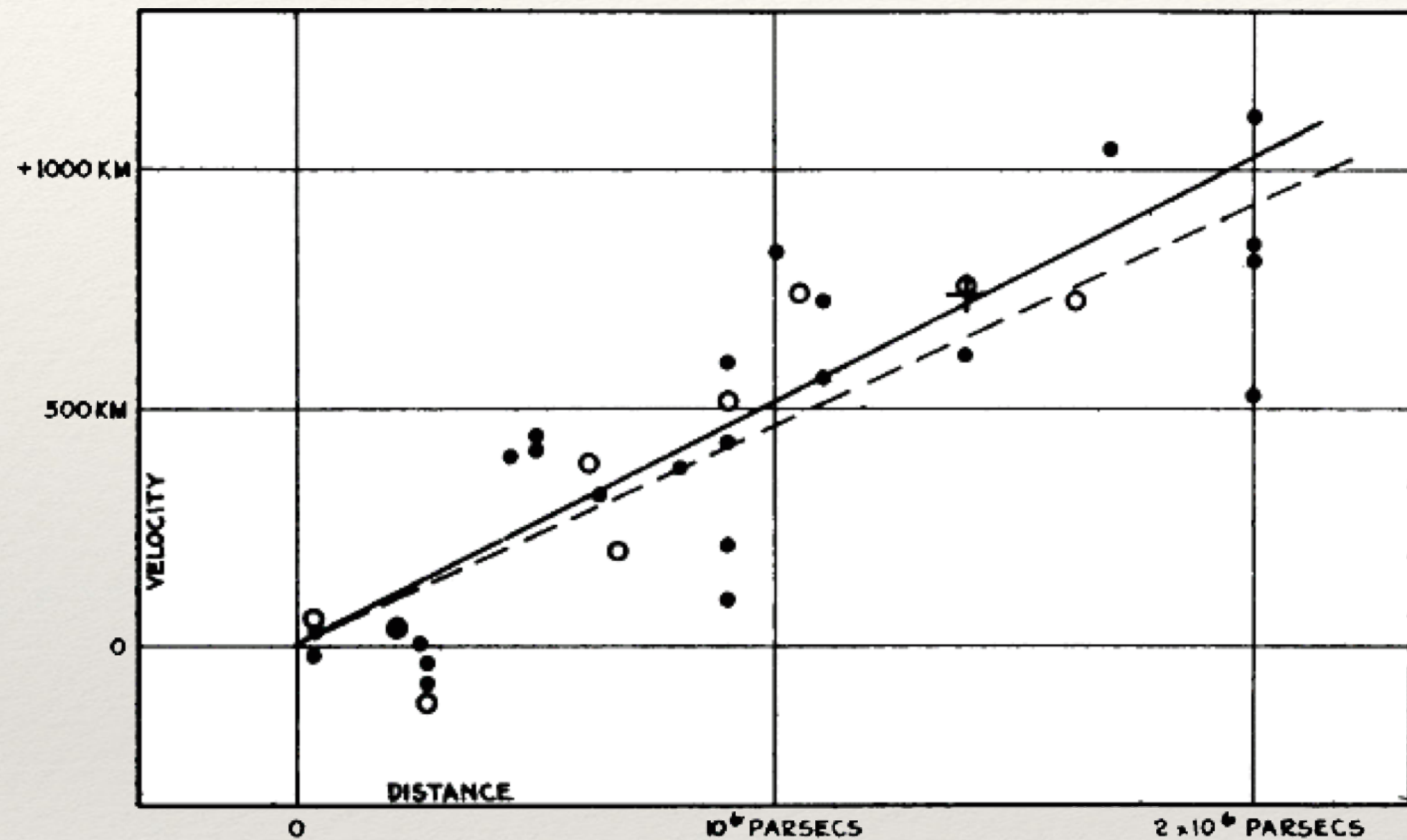
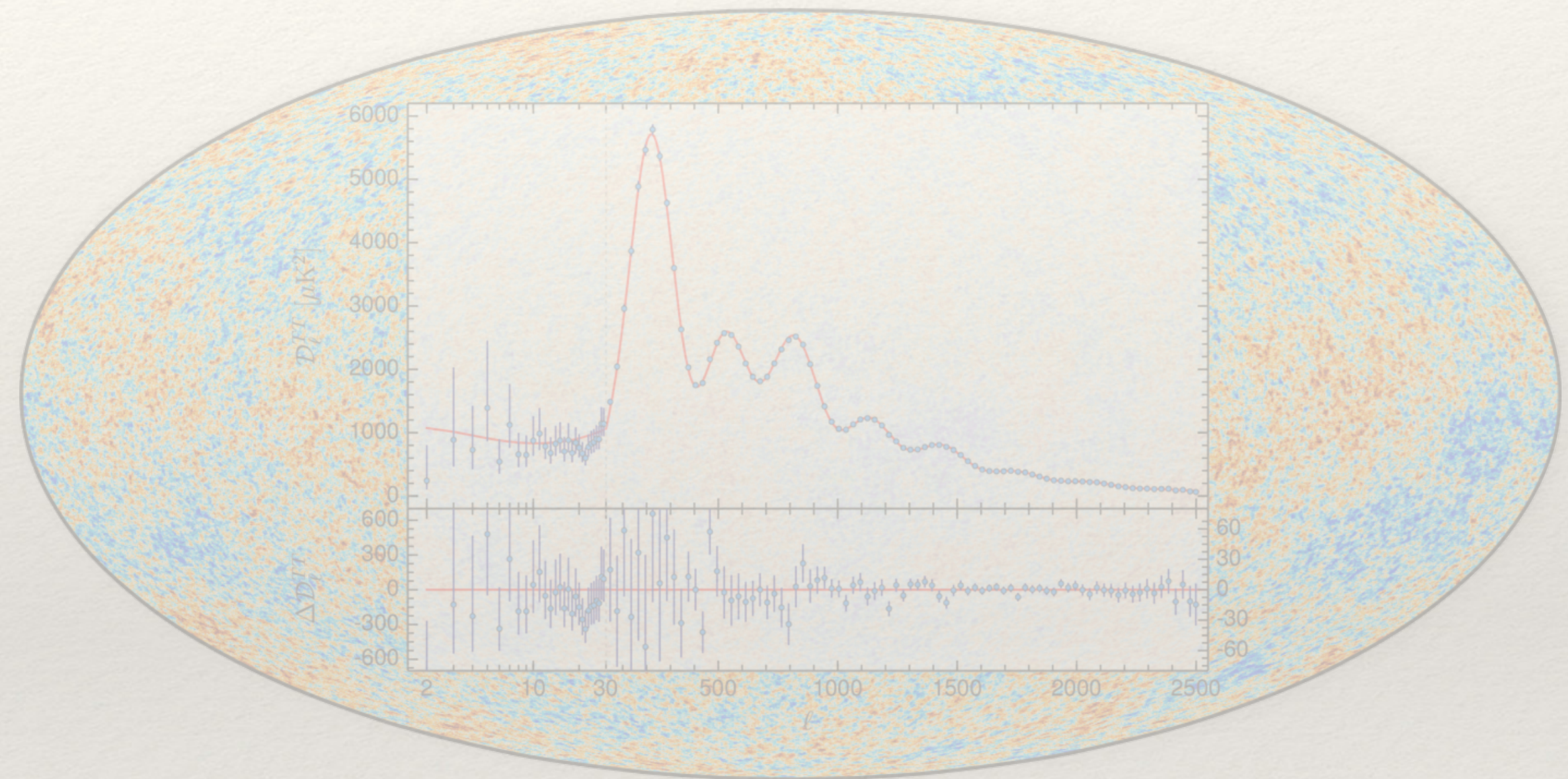


FIGURE 1

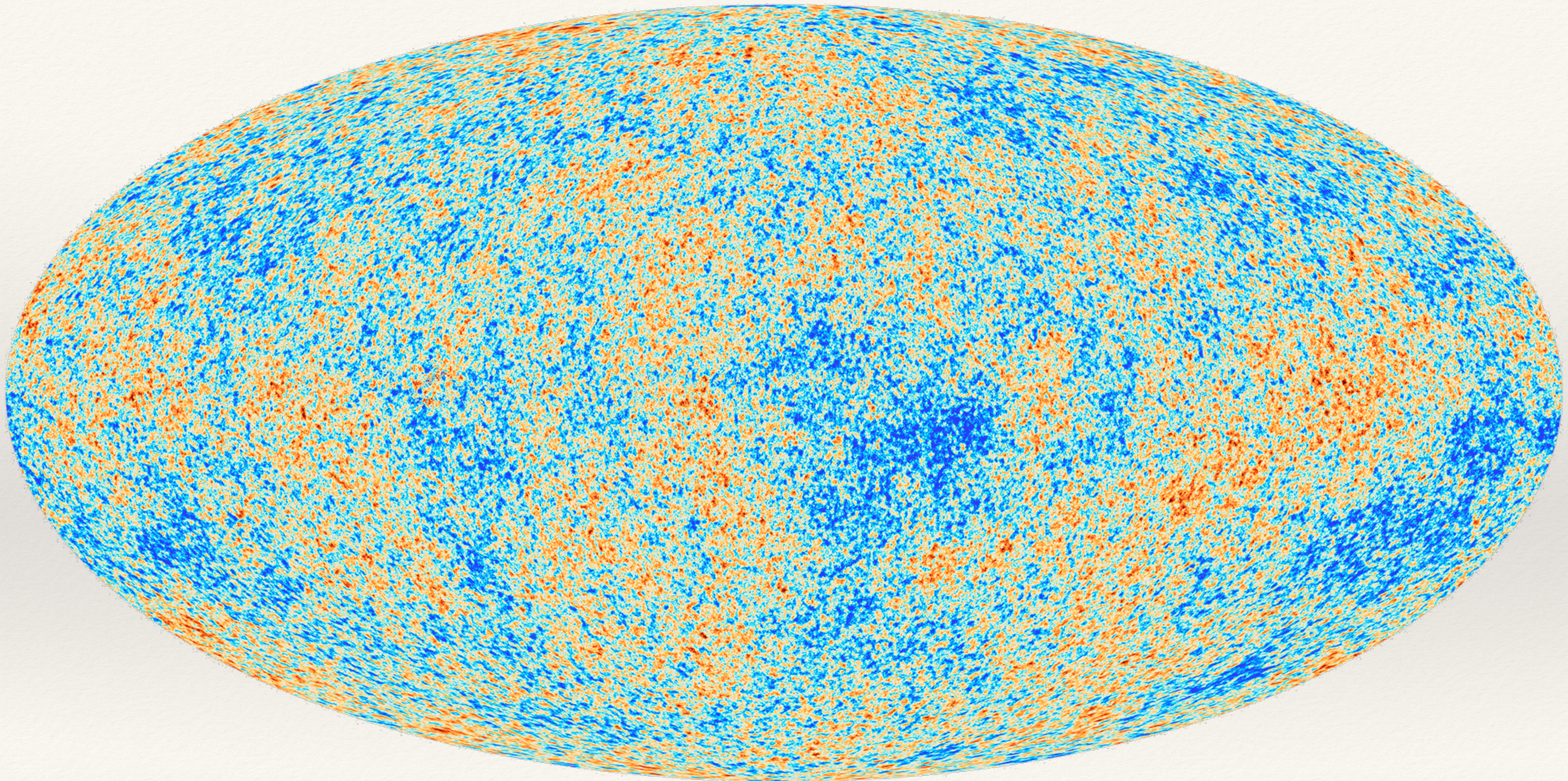
Geometry → Cepheids → SNe Ia

$$H_0 = 73.0 \pm 1.0 \text{ km s}^{-1} \text{ Mpc}^{-1}$$

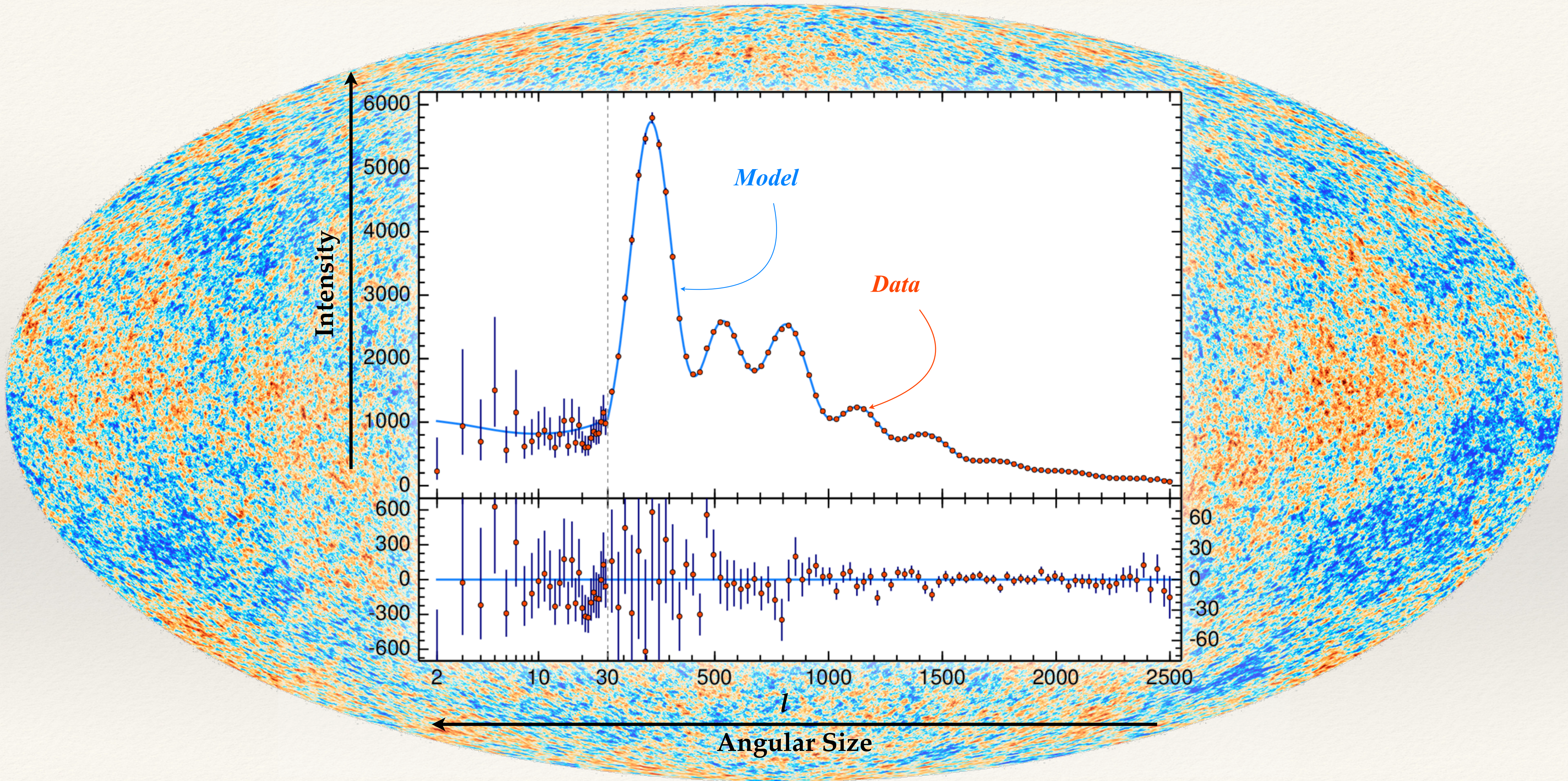
$$H(z) = H_0 \times \sqrt{\Omega_r(1+z)^4 + \Omega_m(1+z)^3 + \Omega_\Lambda(1+z)^{3(1+w)}}$$



Model dependent



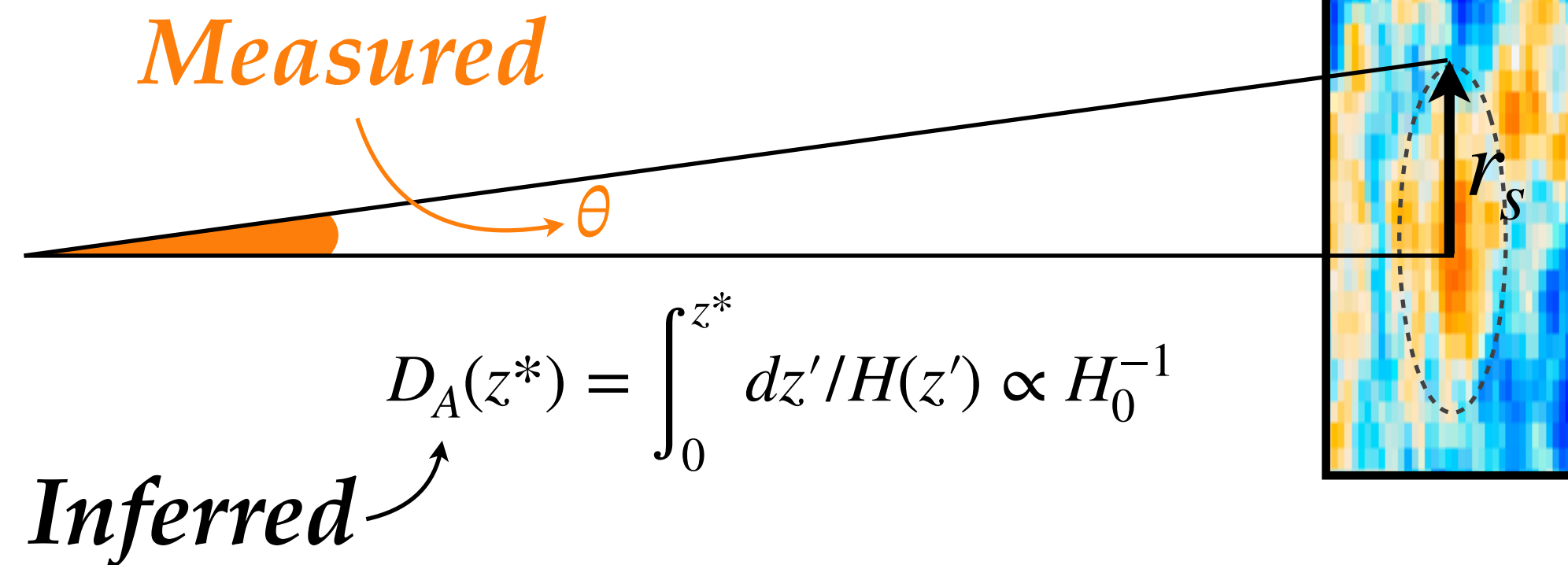
Planck 2020



Deriving H_0

$$H(z) = H_0 \times \sqrt{\Omega_r(1+z)^4 + \Omega_m(1+z)^3 + \Omega_\Lambda(1+z)^{3(1+w)}}$$

H₀ from CMB



$r_s \ll$ the sound horizon »

Calculated

$$r_s = \int_{z_s}^{\infty} \frac{c_s(z)}{H(z)} dz$$

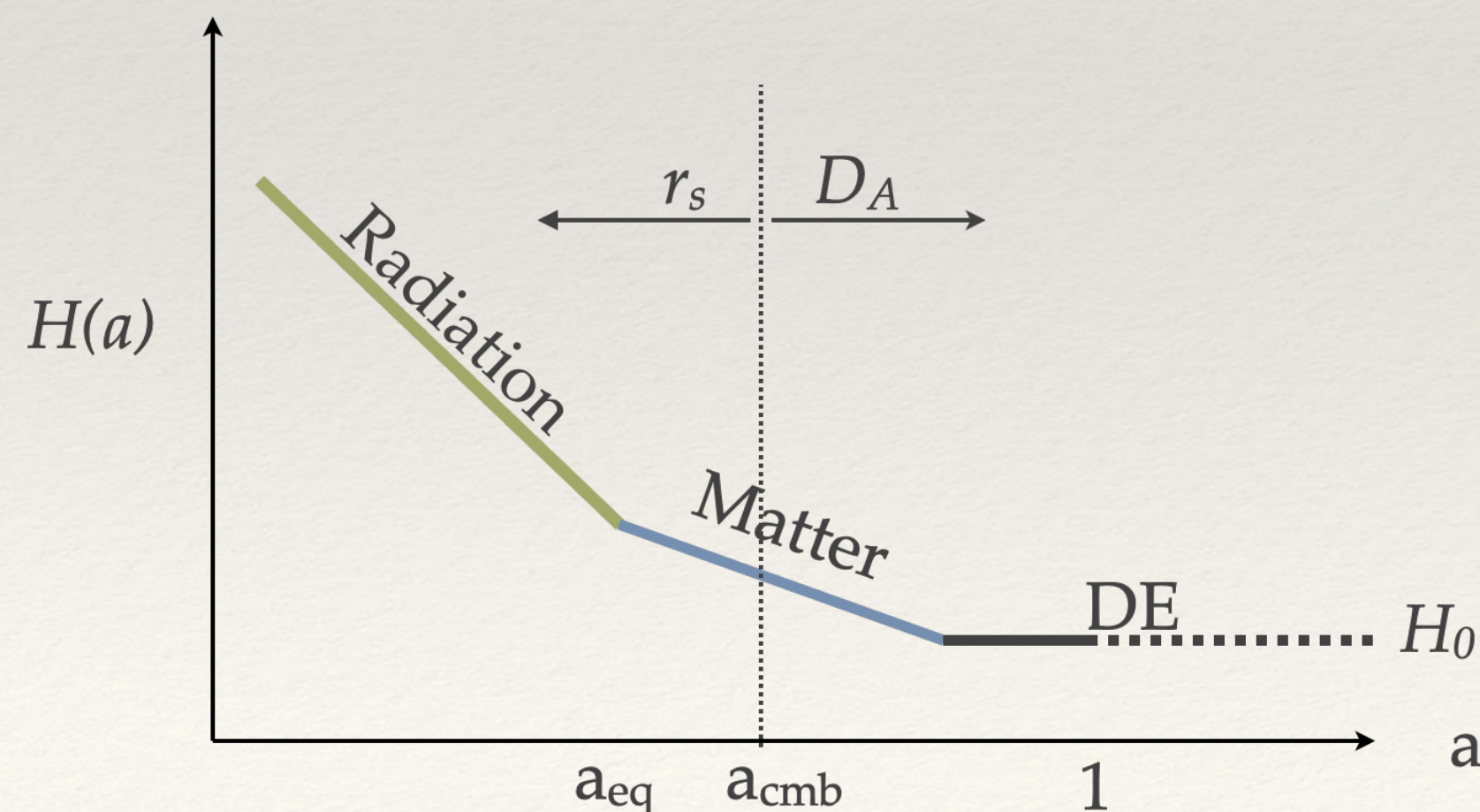
Radiation driven

$$H^2(z) = \frac{8\pi G}{3} (\rho_\gamma + \rho_\nu + \rho_m)$$

r_s from baryon and matter density (radiation)

θ from observation

$H(z)$ from $D_A = r_s / \theta$



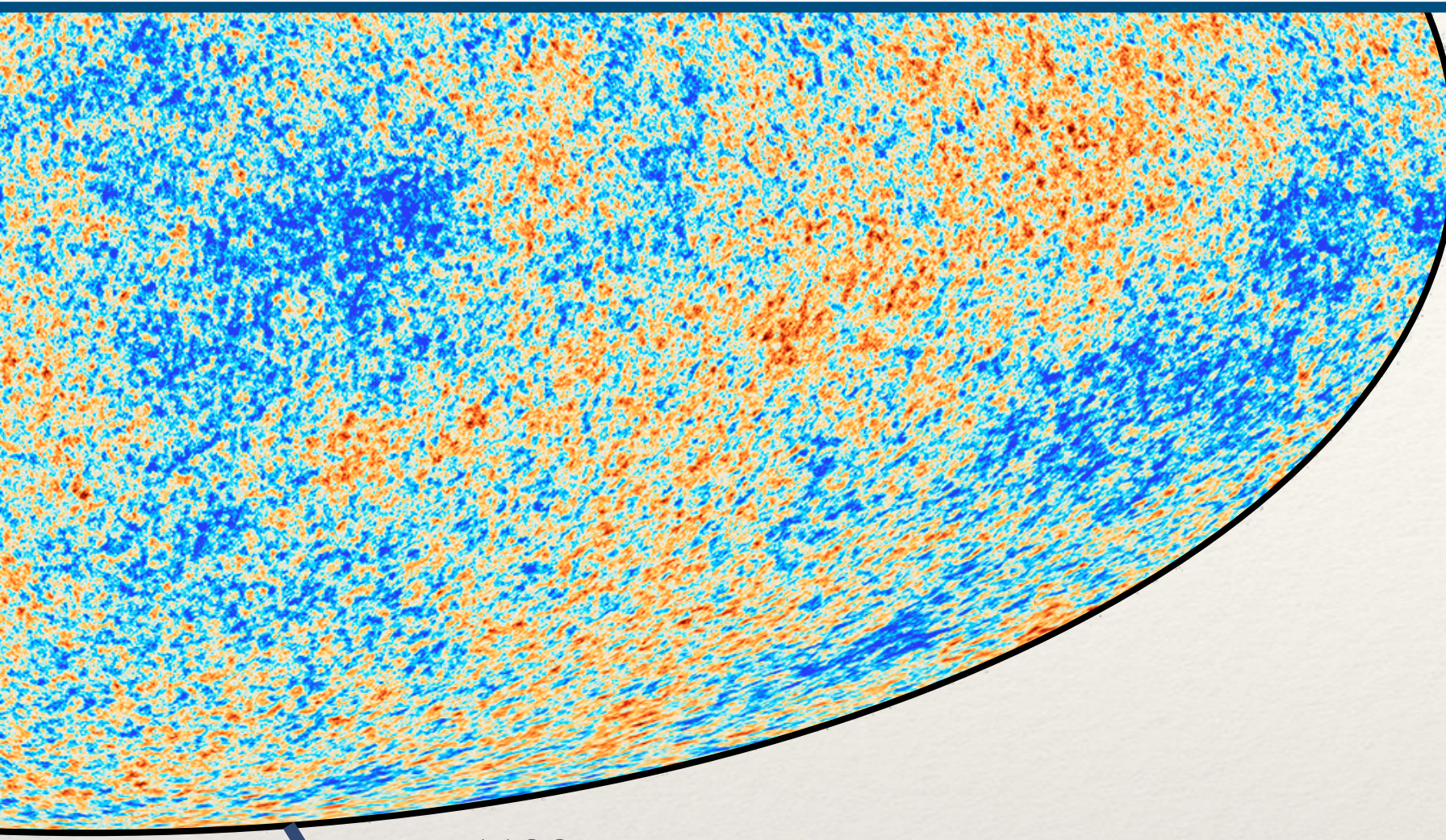
We can think of the estimation of H_0 from CMB data as proceeding in three steps:

- 1) determine the baryon density and matter density to allow for calculation of r_s ,
- 2) infer θ_s from the spacing between the acoustic peaks to determine the comoving angular diameter distance to last scattering $D_A = r_s / \theta_s$,
- 3) adjust the only remaining free density parameter in the model so that D_A gives this inferred distance.

With this last step complete we now have $H(z)$ determined for all z , including $z = 0$.

Indirect determination of H_0

Planck et al. 2020



$z \sim 1100$

**THE MODEL
CONSTRAINS H_0**

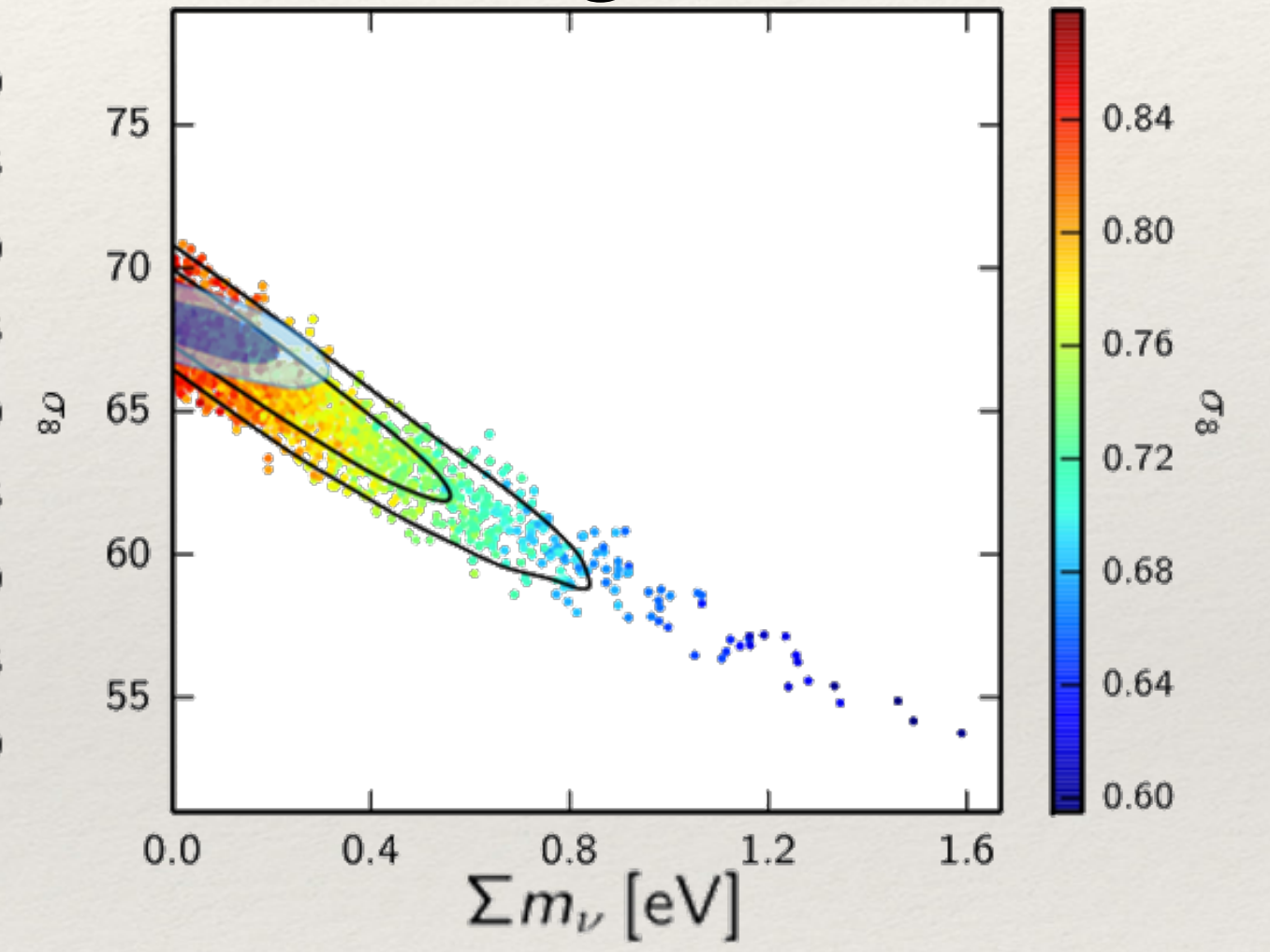
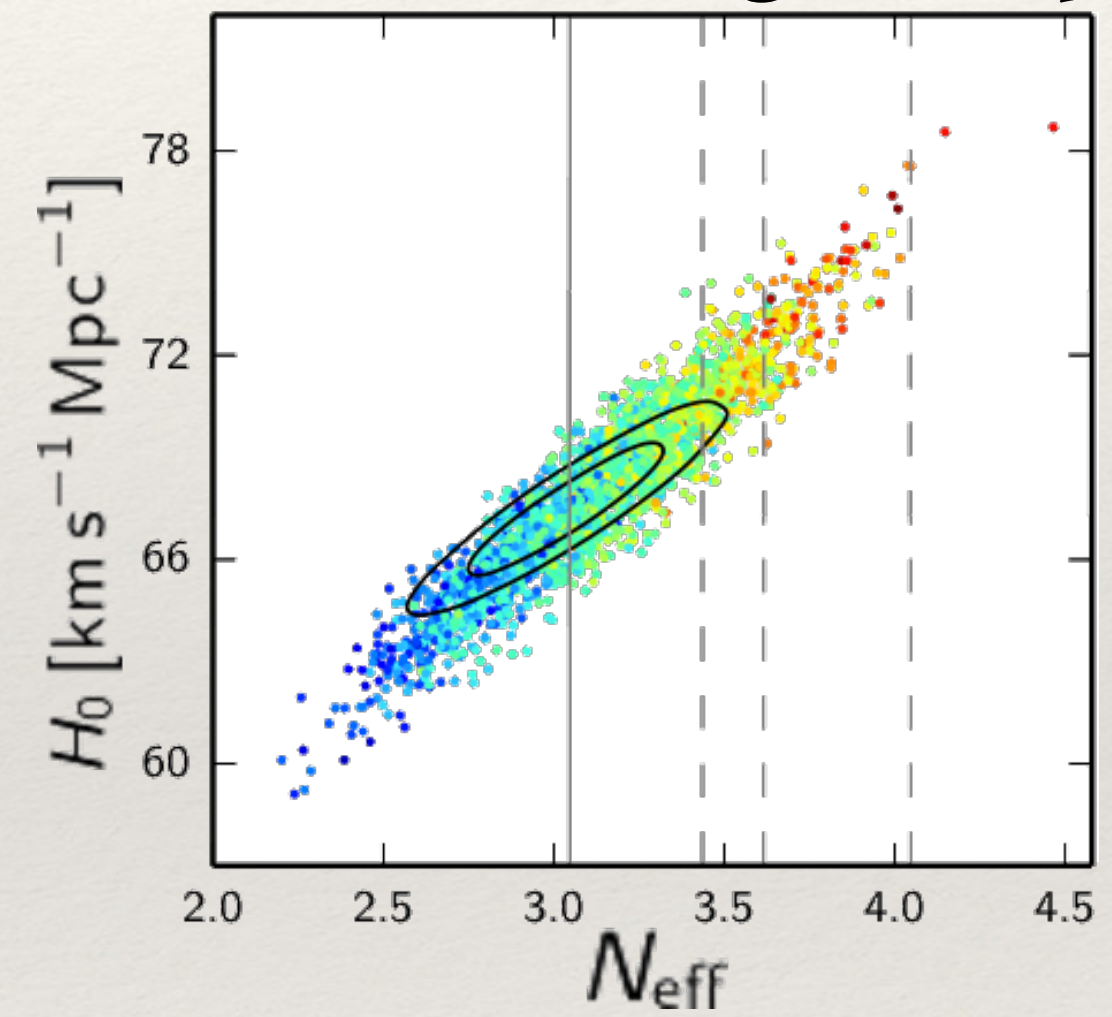
$z \sim 0$

$$H_0 = 67.4 \pm 0.5 \text{ km s}^{-1} \text{ Mpc}^{-1}$$

— based on Λ CDM —

*Test the concordance
model Λ CDM*

Change the parameters, change H_0

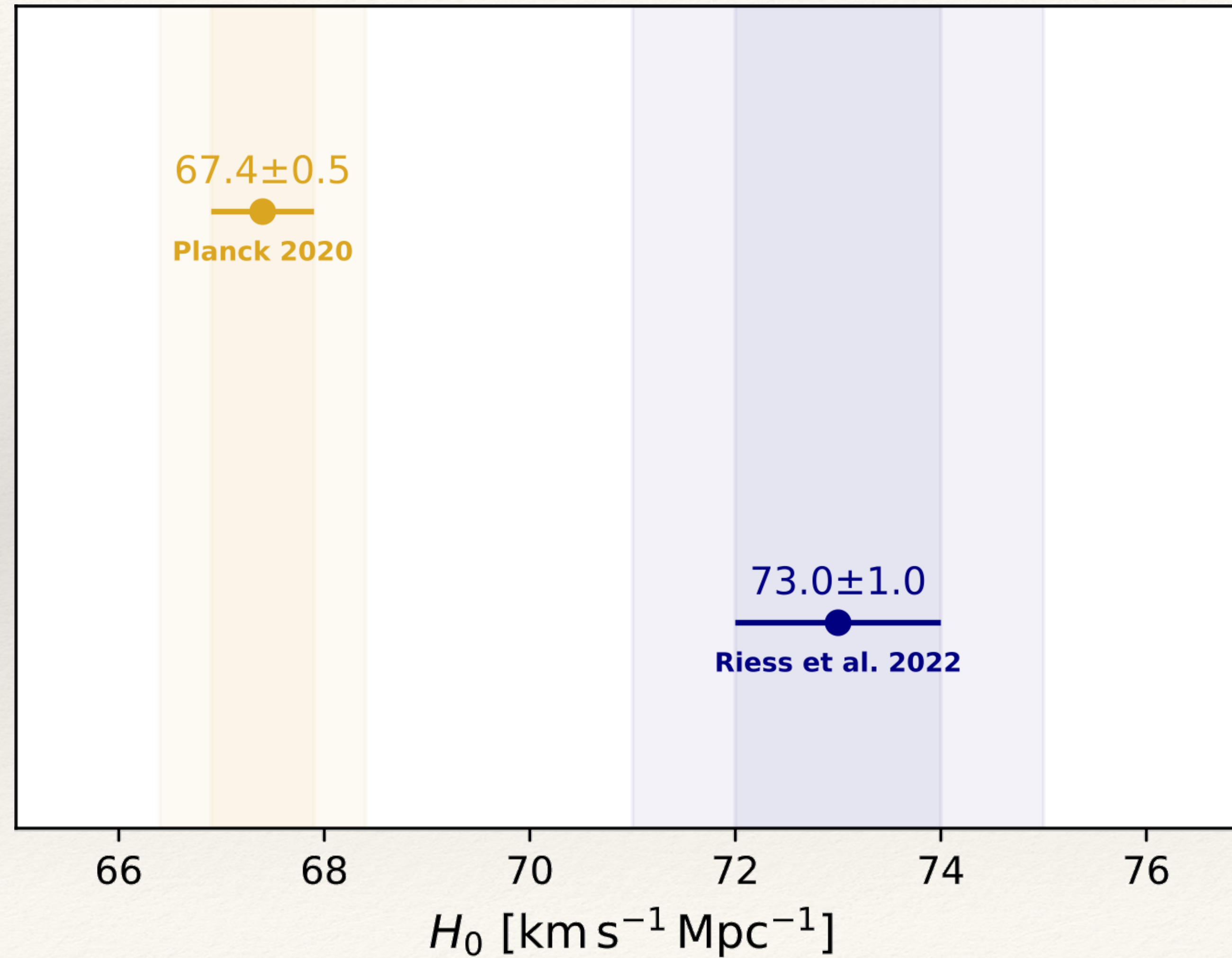


Illustrative plots from Planck 2015

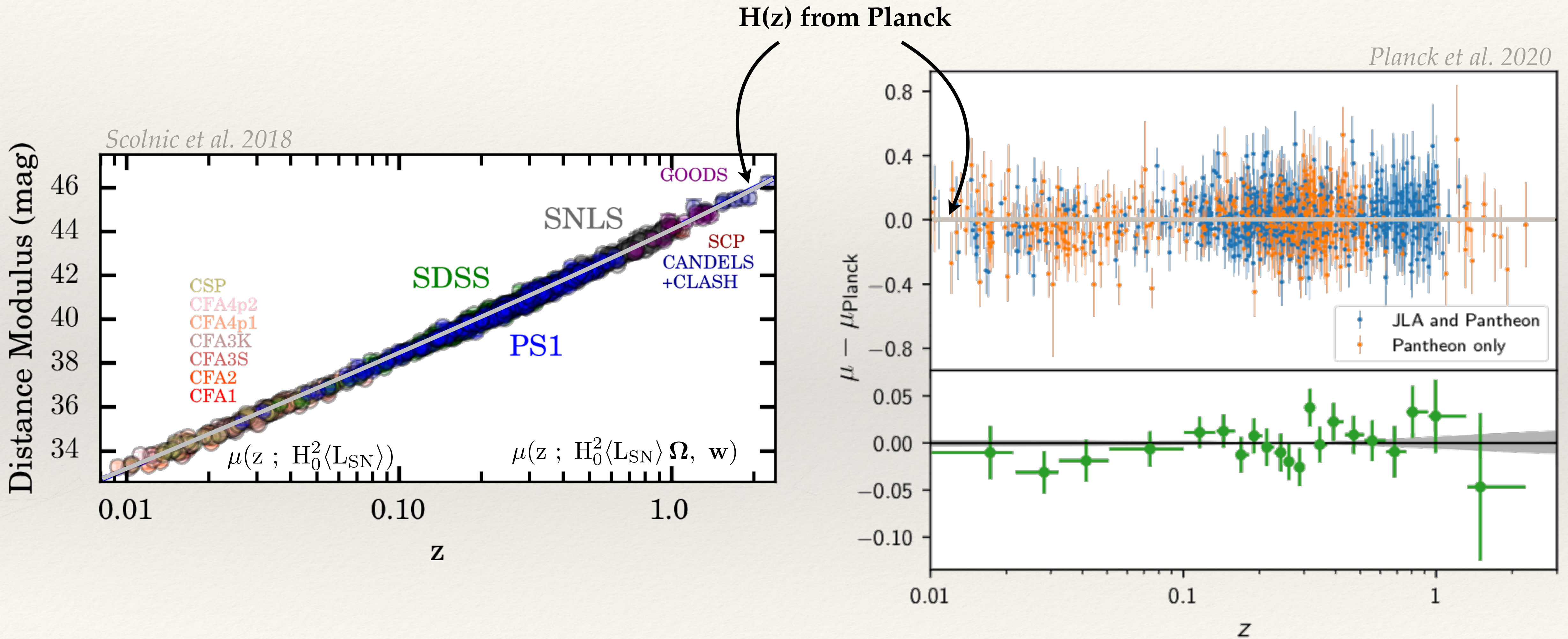
H_0 Tension | *SHOES* vs. *Planck*

Λ CDM

direct



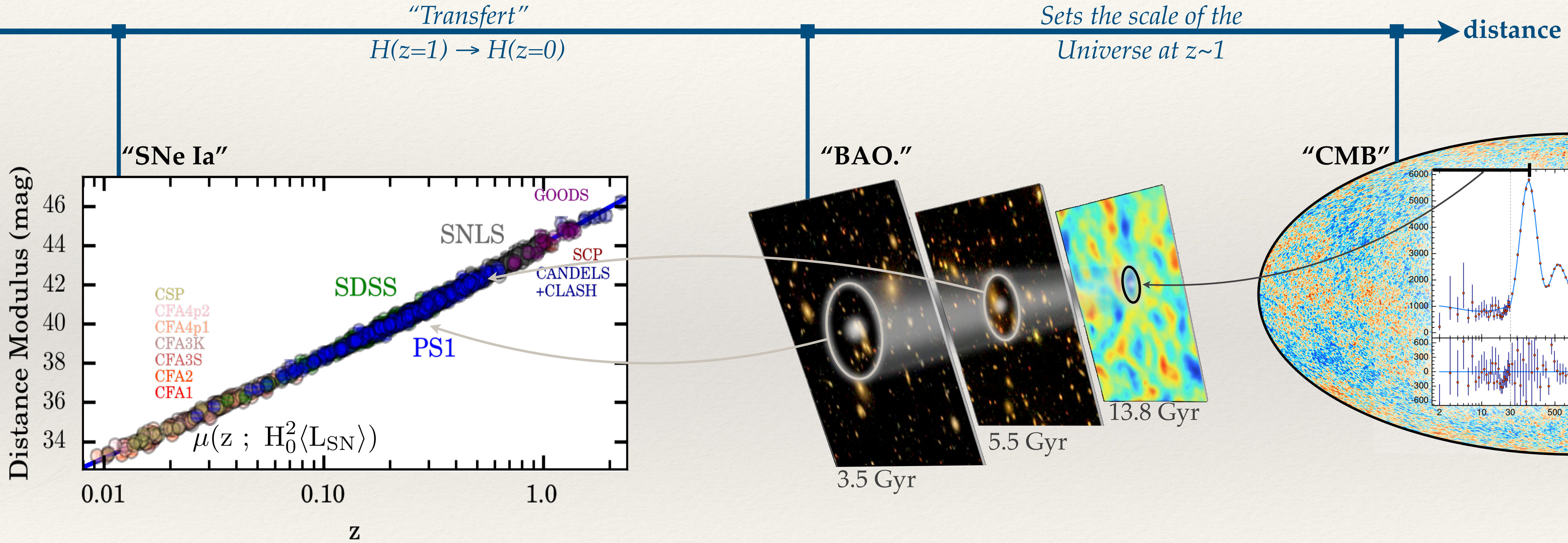
Are Supernovae & CMB in tension ? *No!*



Inverse Distance Ladder

See also e.g.:
Aubourg et al. 2015 • Macaulay et al. 2018

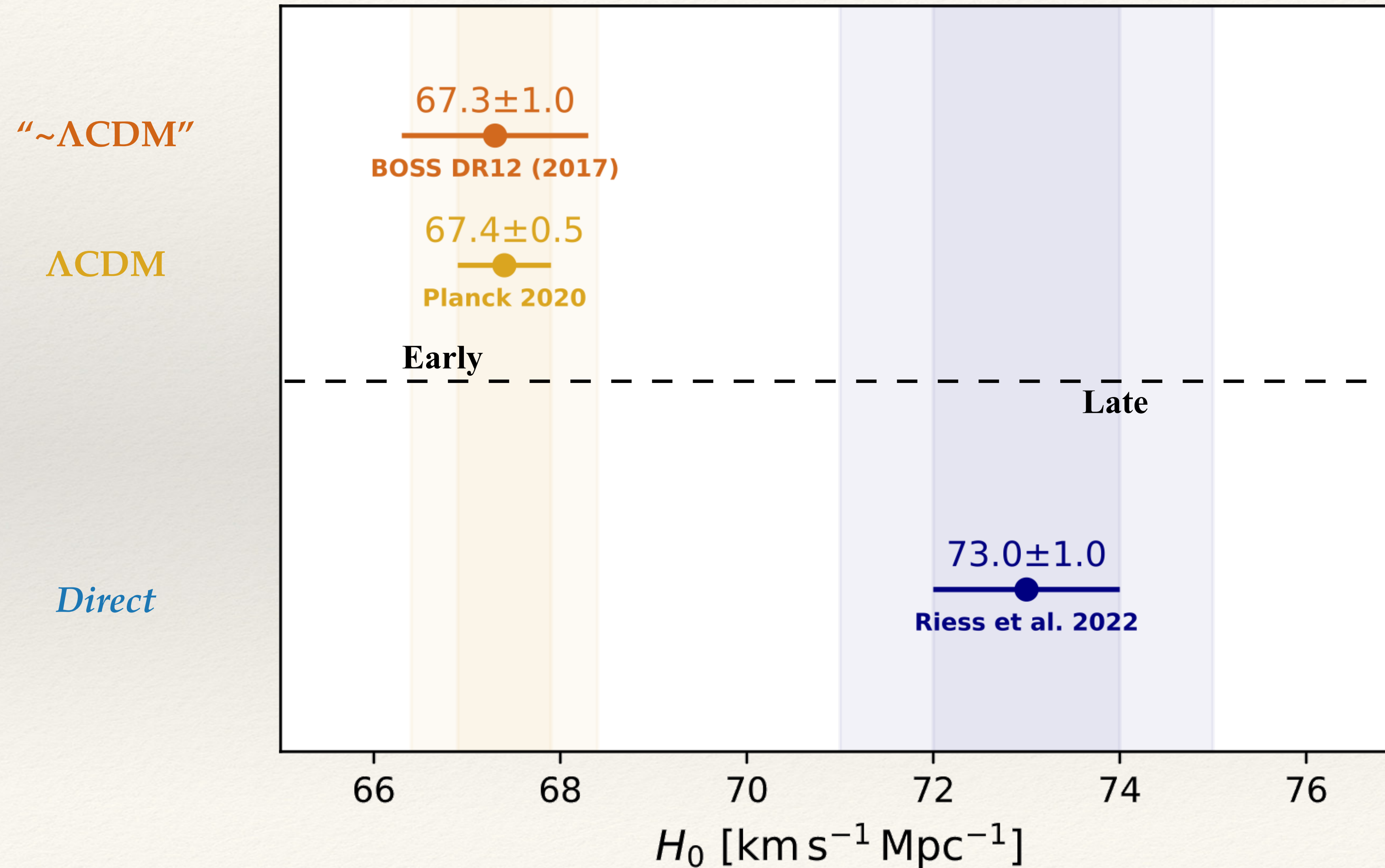
Get independent distances for SNe Ia



BOSS DR12 | Alam et al. 2017

$H_0 = 67.3 \pm 1.0 \text{ km s}^{-1} \text{ Mpc}^{-1}$

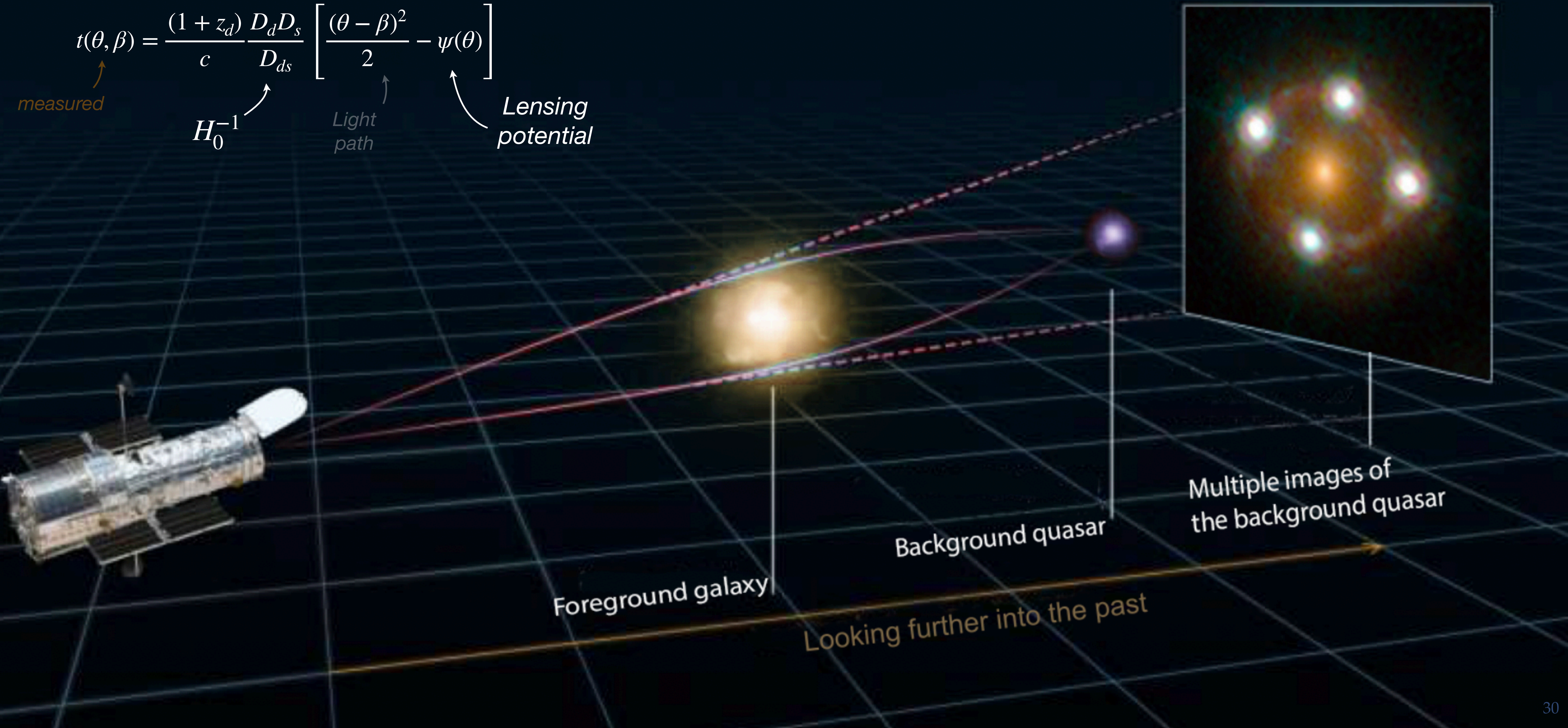
H_0 Tension | *Early vs. Late physics*



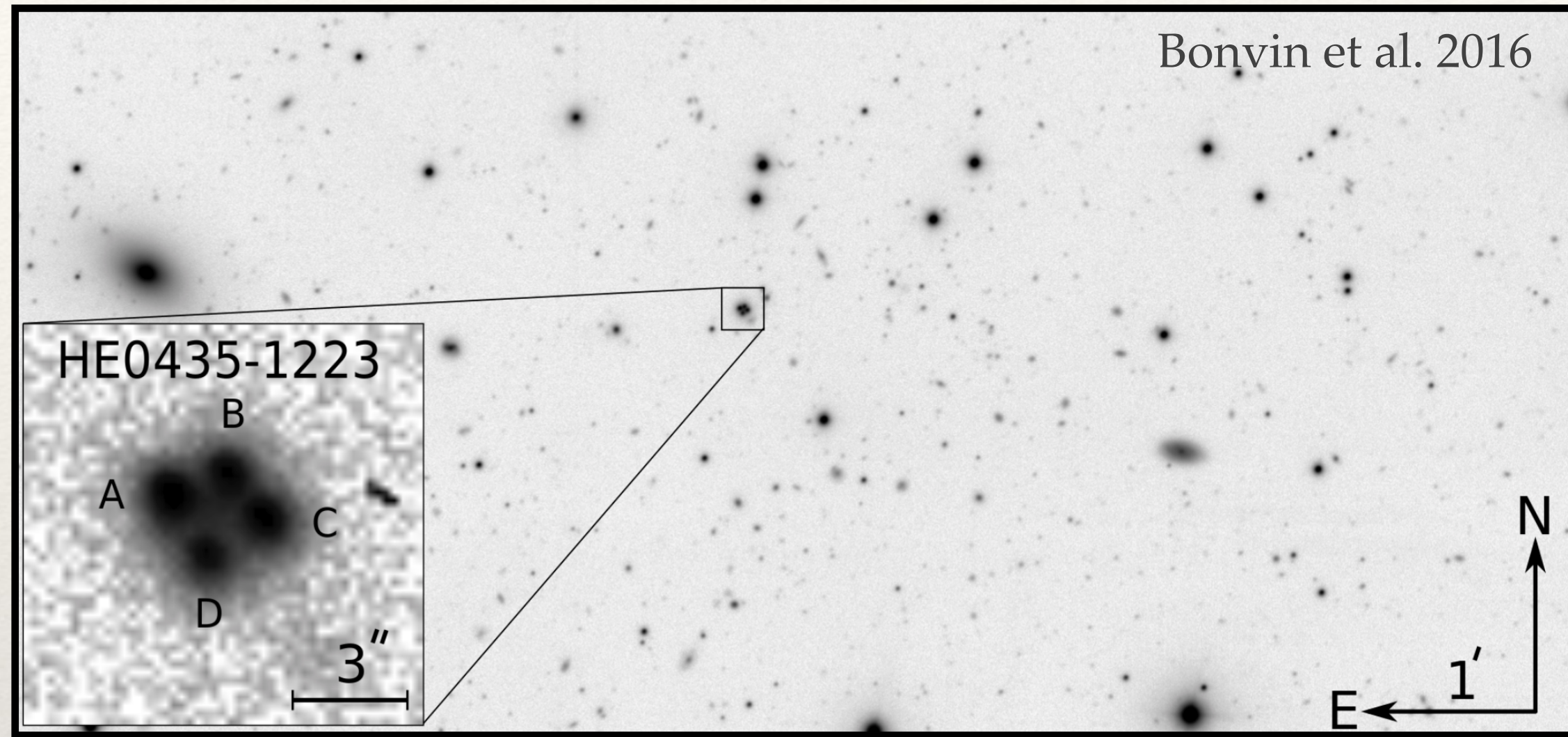
Time Delay Cosmology

$$t(\theta, \beta) = \frac{(1+z_d) D_d D_s}{c D_{ds}} \left[\frac{(\theta - \beta)^2}{2} - \psi(\theta) \right]$$

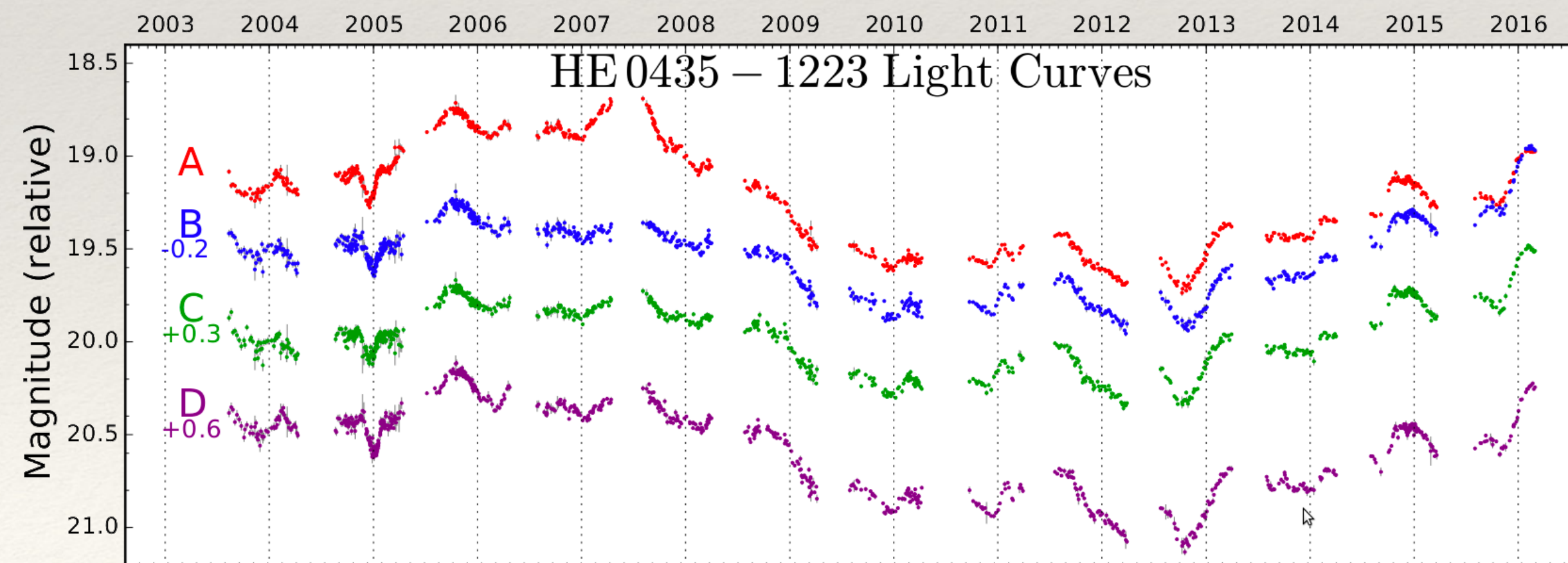
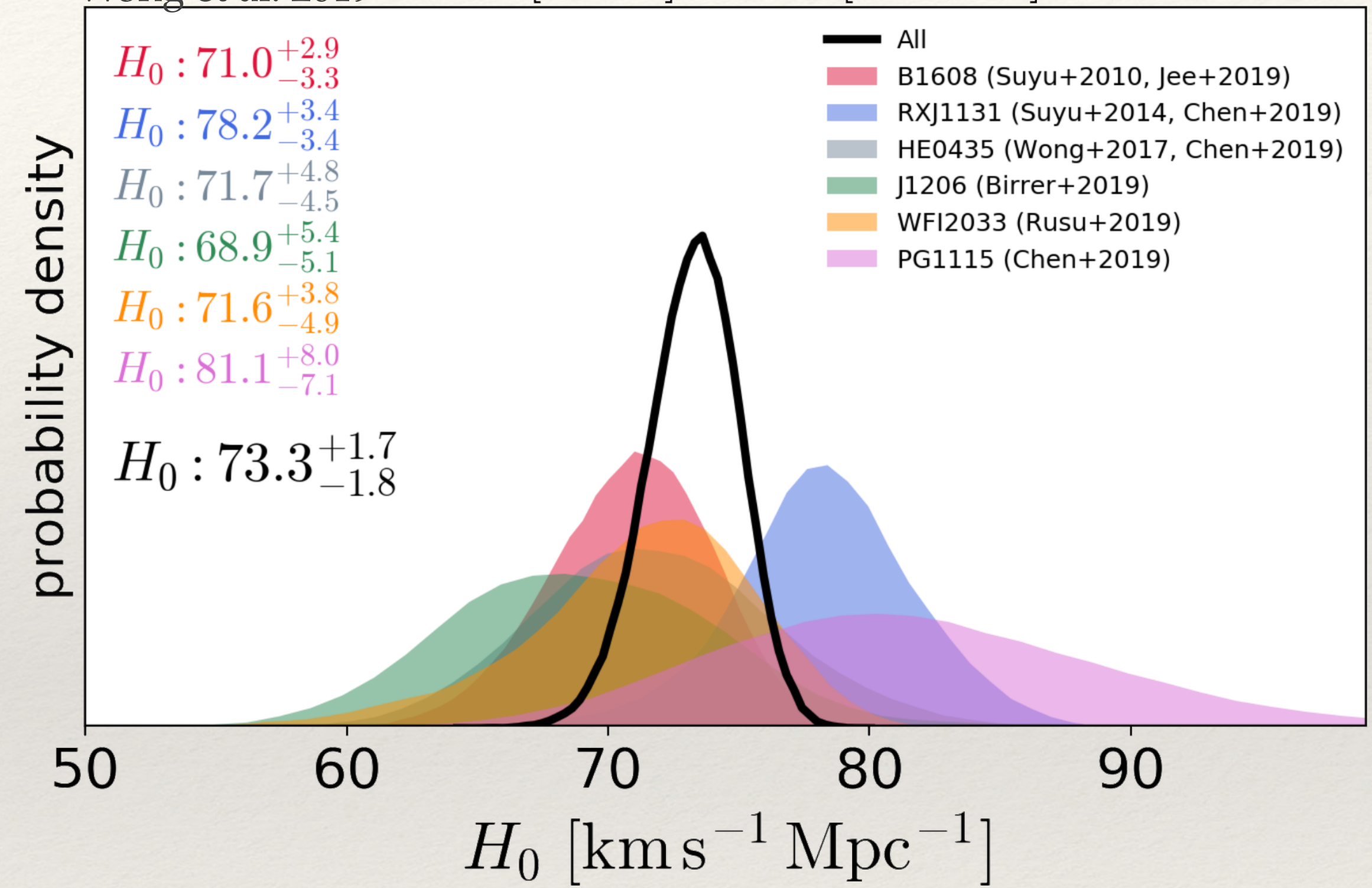
measured \uparrow $t(\theta, \beta)$
 H_0^{-1} \uparrow $\frac{(1+z_d) D_d D_s}{c D_{ds}}$
Light path \uparrow $\frac{(\theta - \beta)^2}{2}$
Lensing potential \uparrow $\psi(\theta)$



H_0 from Strong Lensing



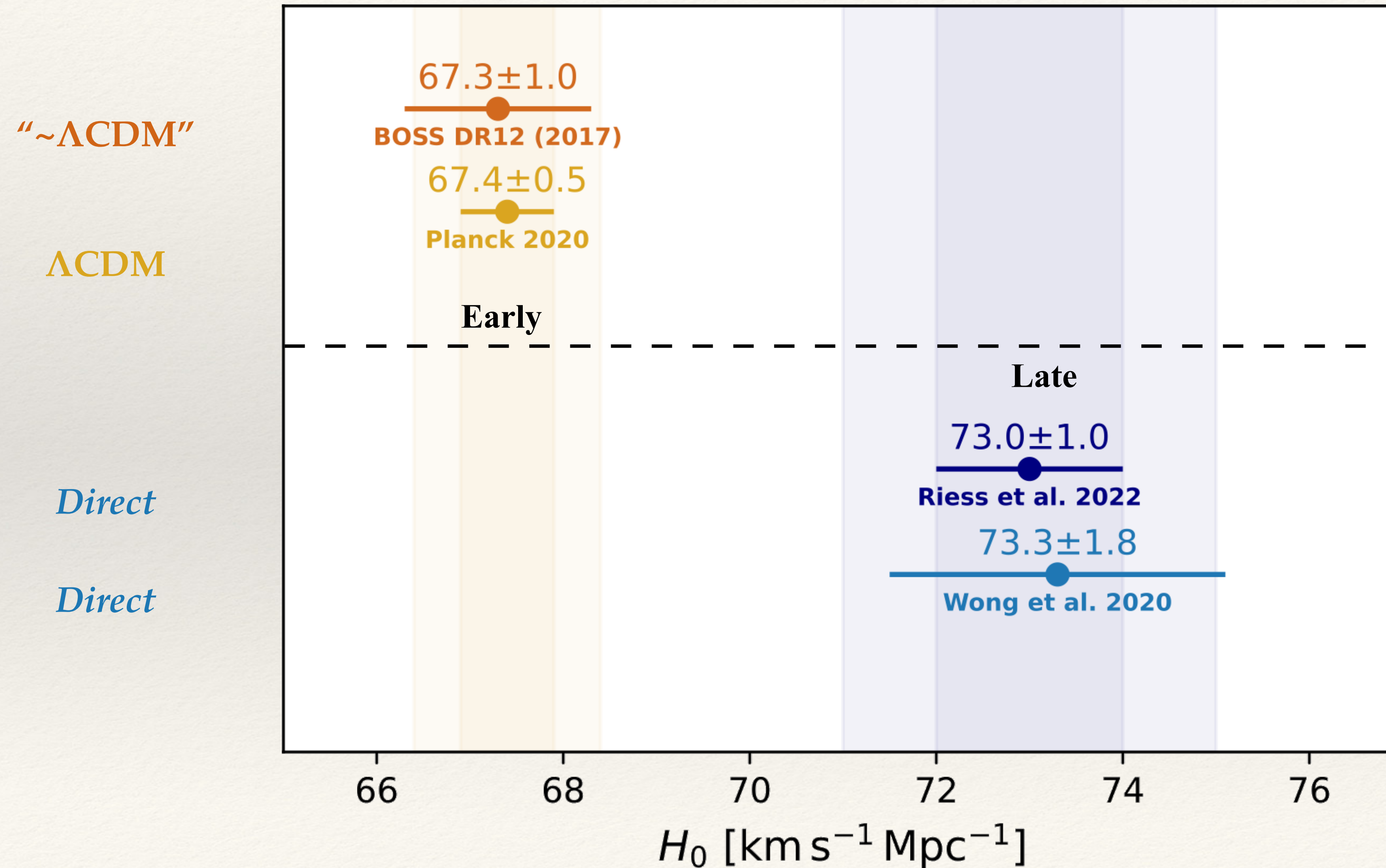
Wong et al. 2019 $H_0 \in [0, 150]$ $\Omega_m \in [0.05, 0.5]$



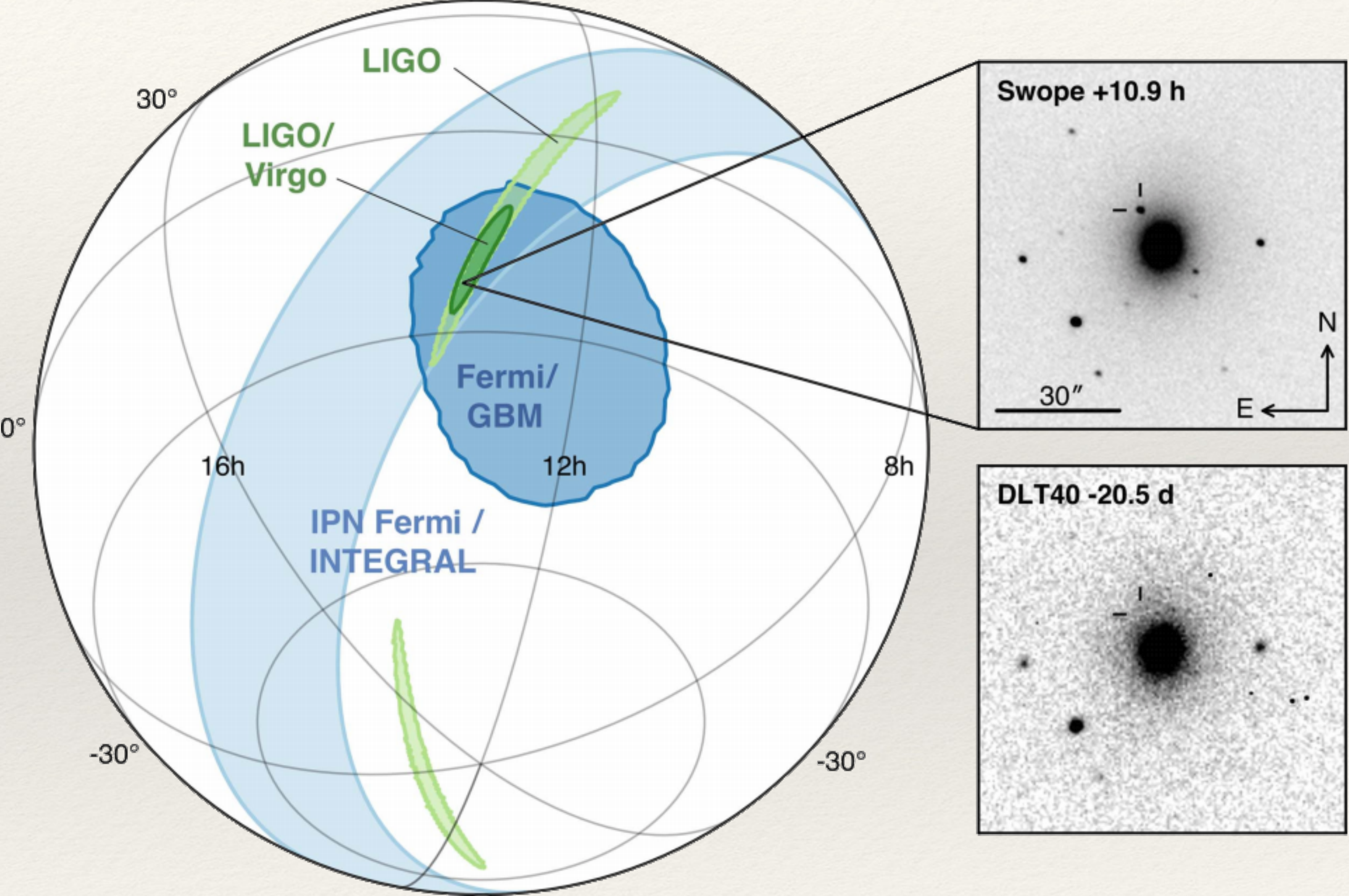
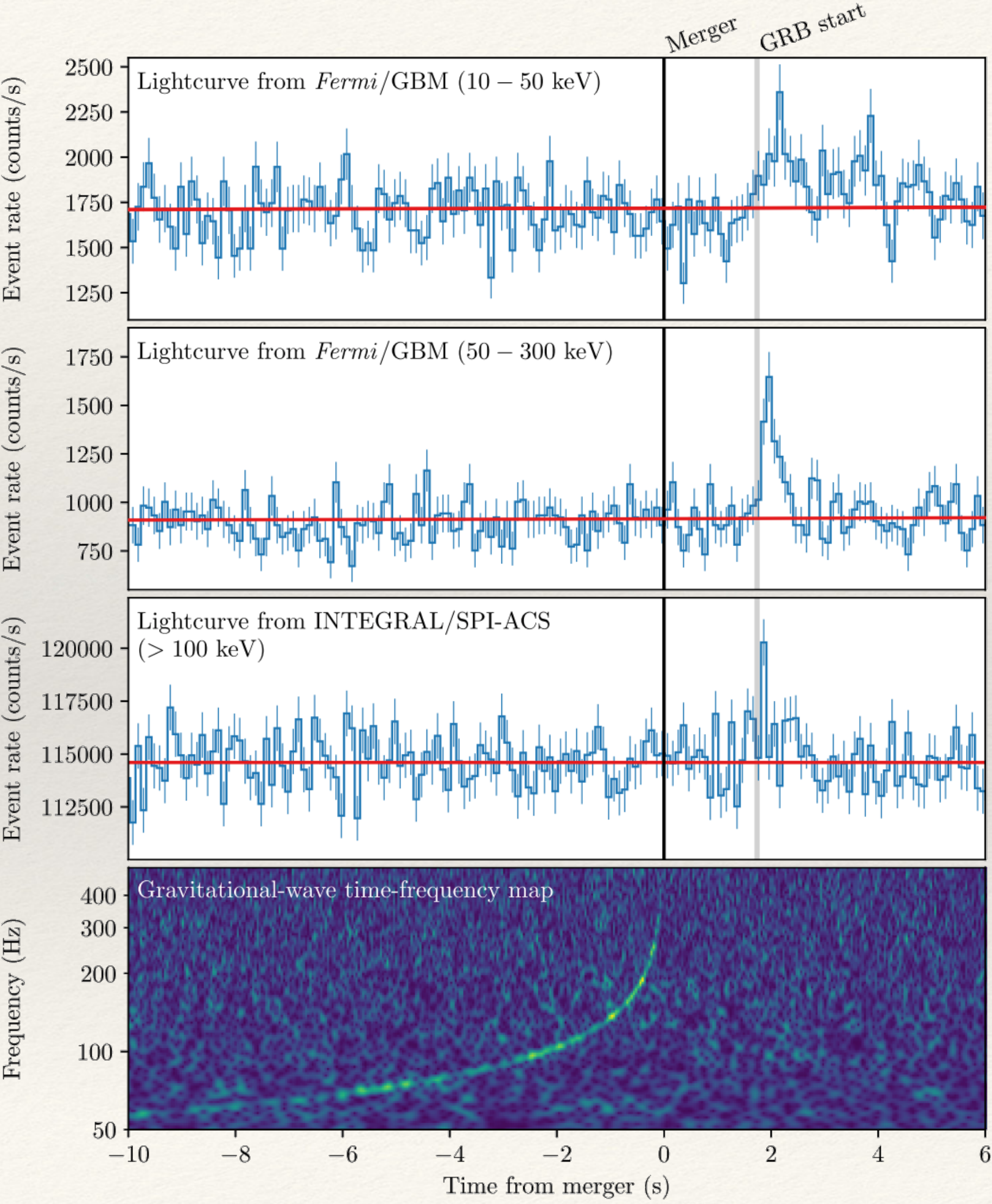
$$\Delta t = \frac{1}{c} D_{\Delta t} \phi_{lens} \propto H_0^{-1}$$

Obtained from lensing mass model

H_0 Tension | *At least 2 independent systematics*



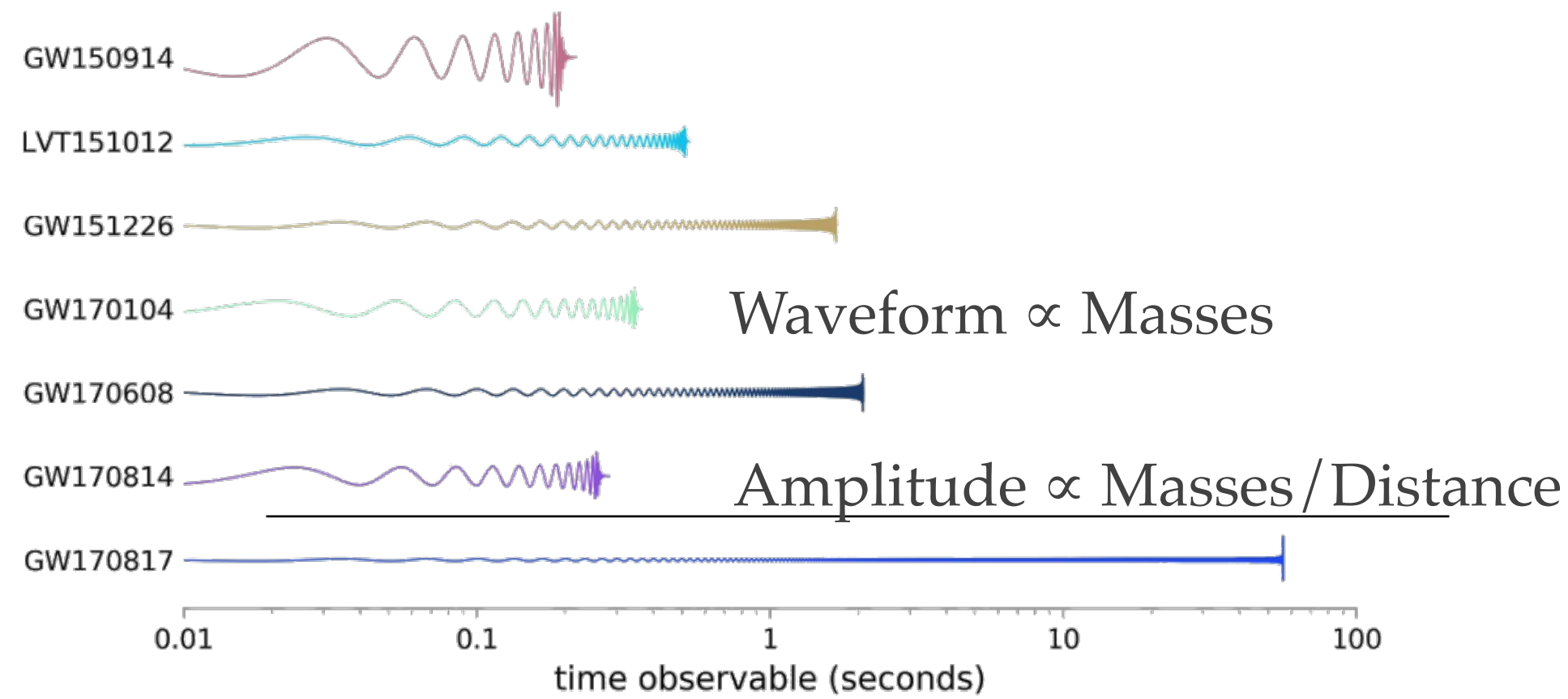
Standard Sirens



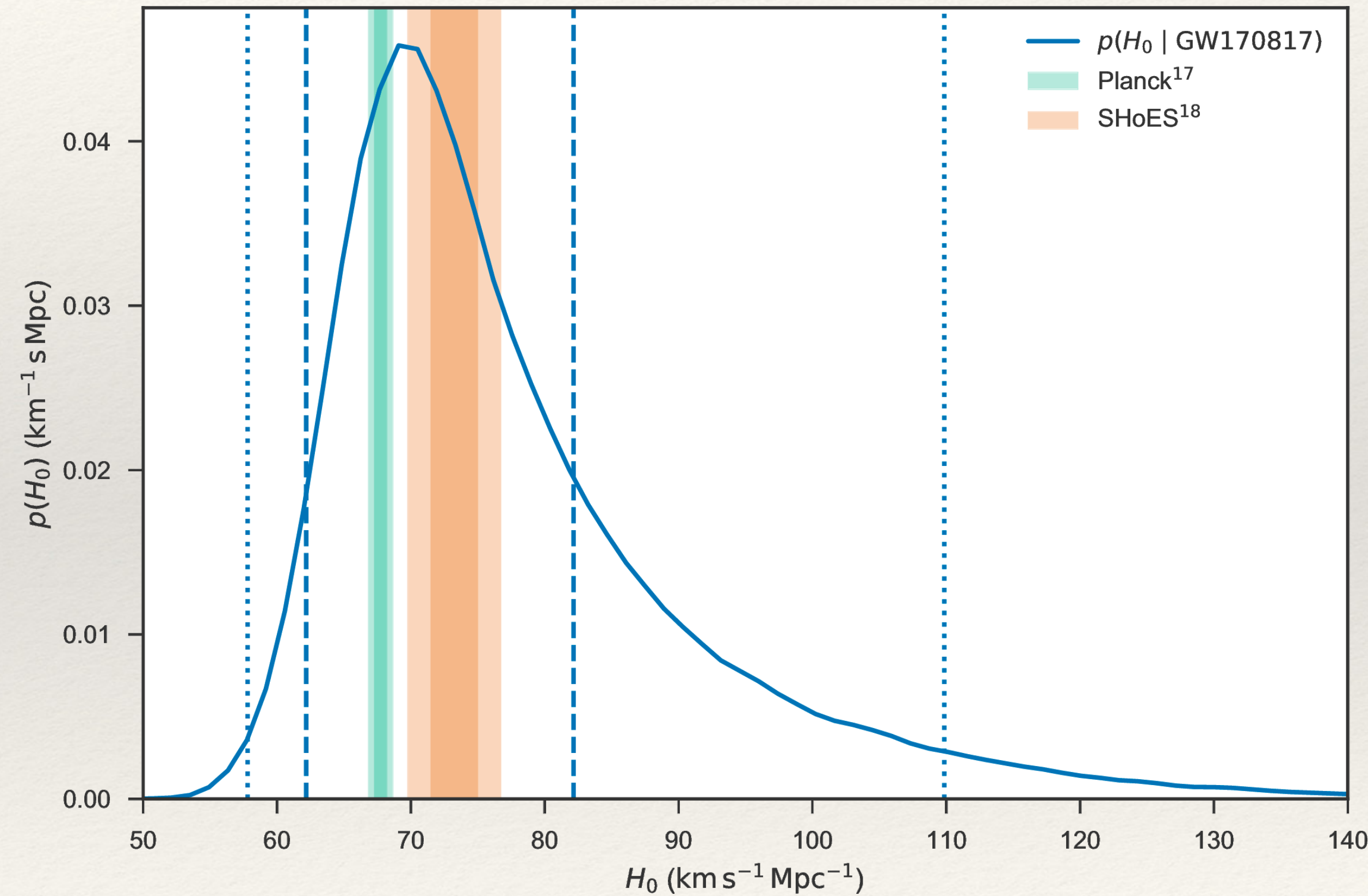
Gravitational Waves & ElectroMagnetism | H_0

Gravitational Waves

d_L

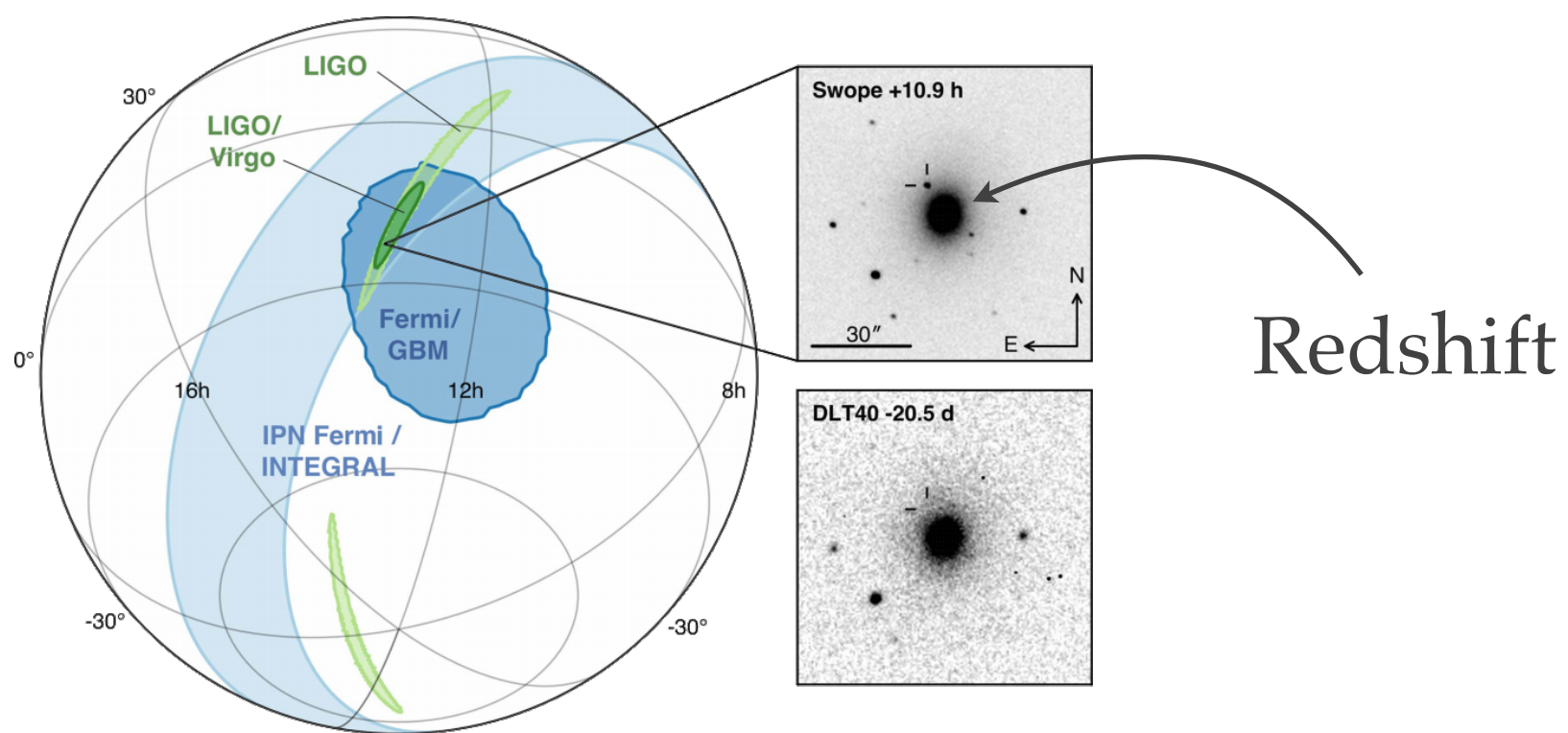


$$v_H = H_0 d_L$$

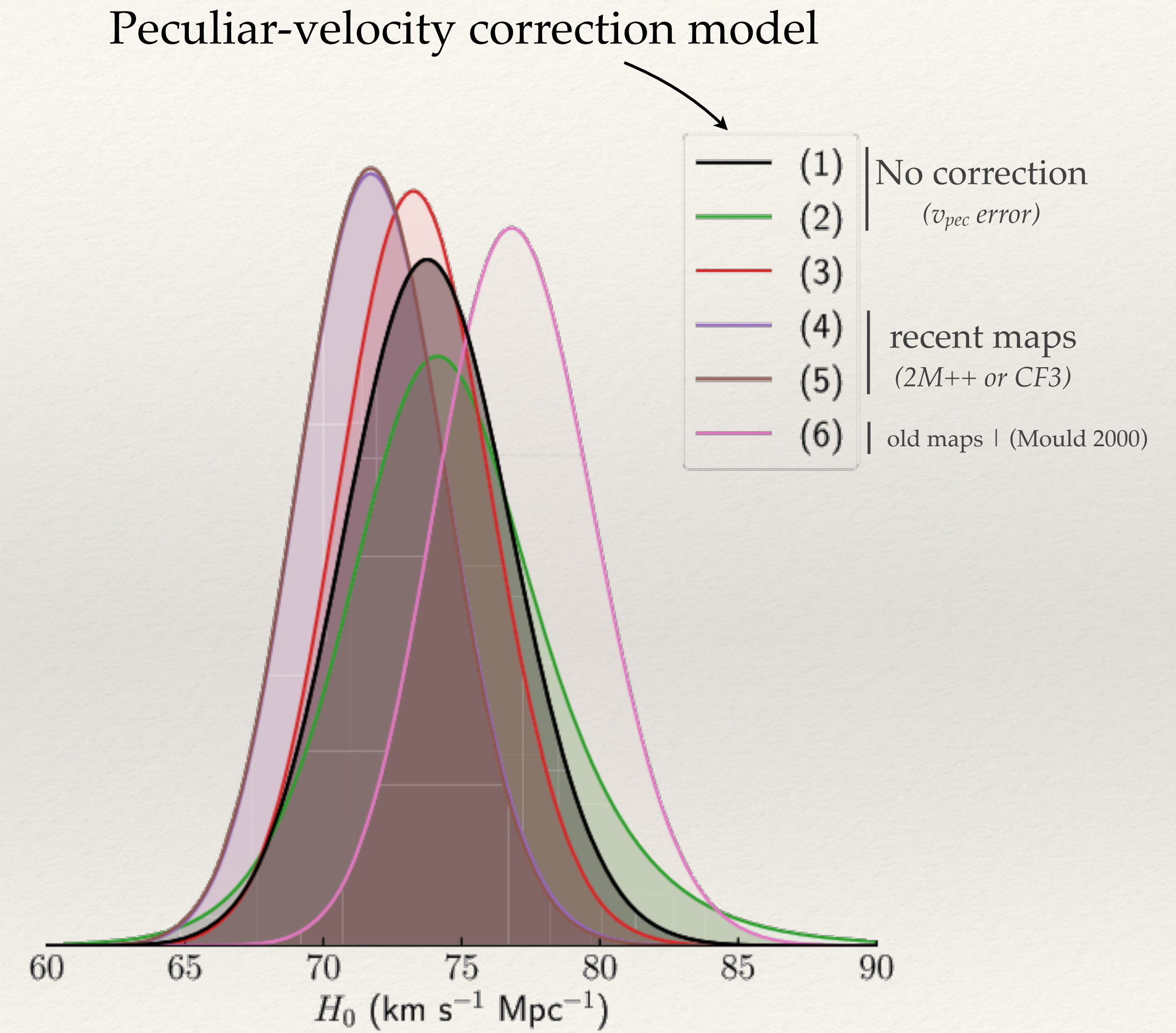
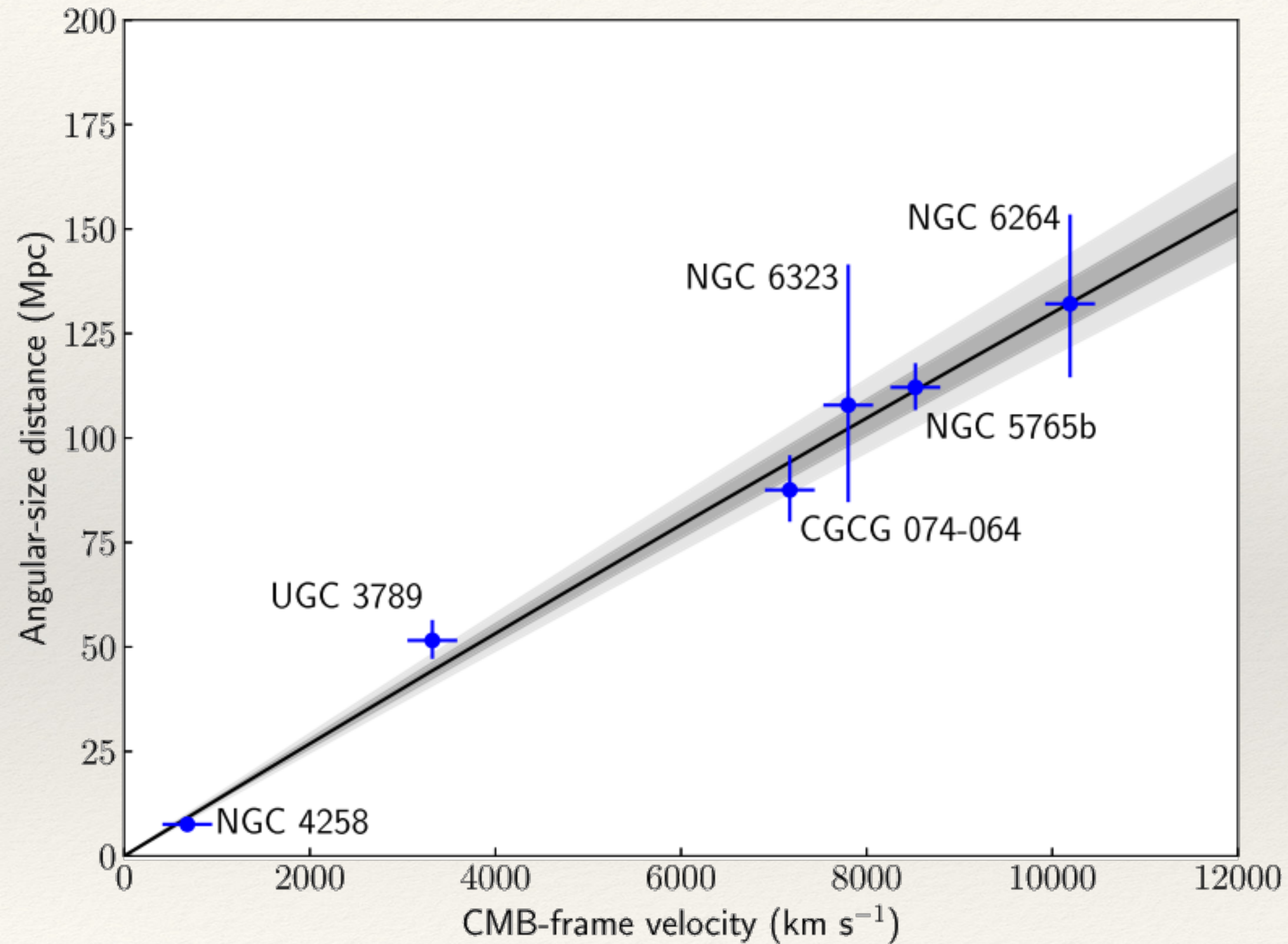


Optical Counterpart

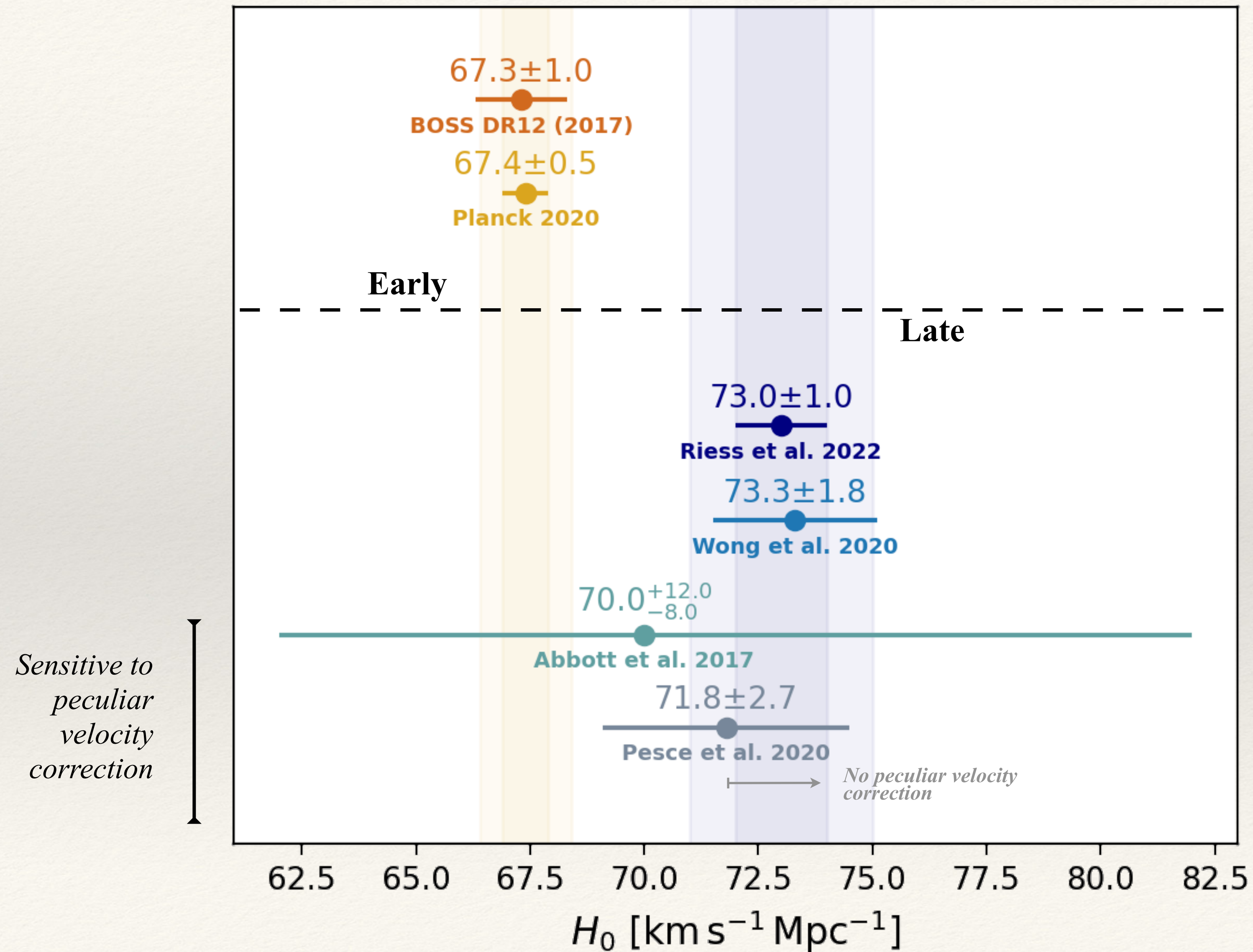
z



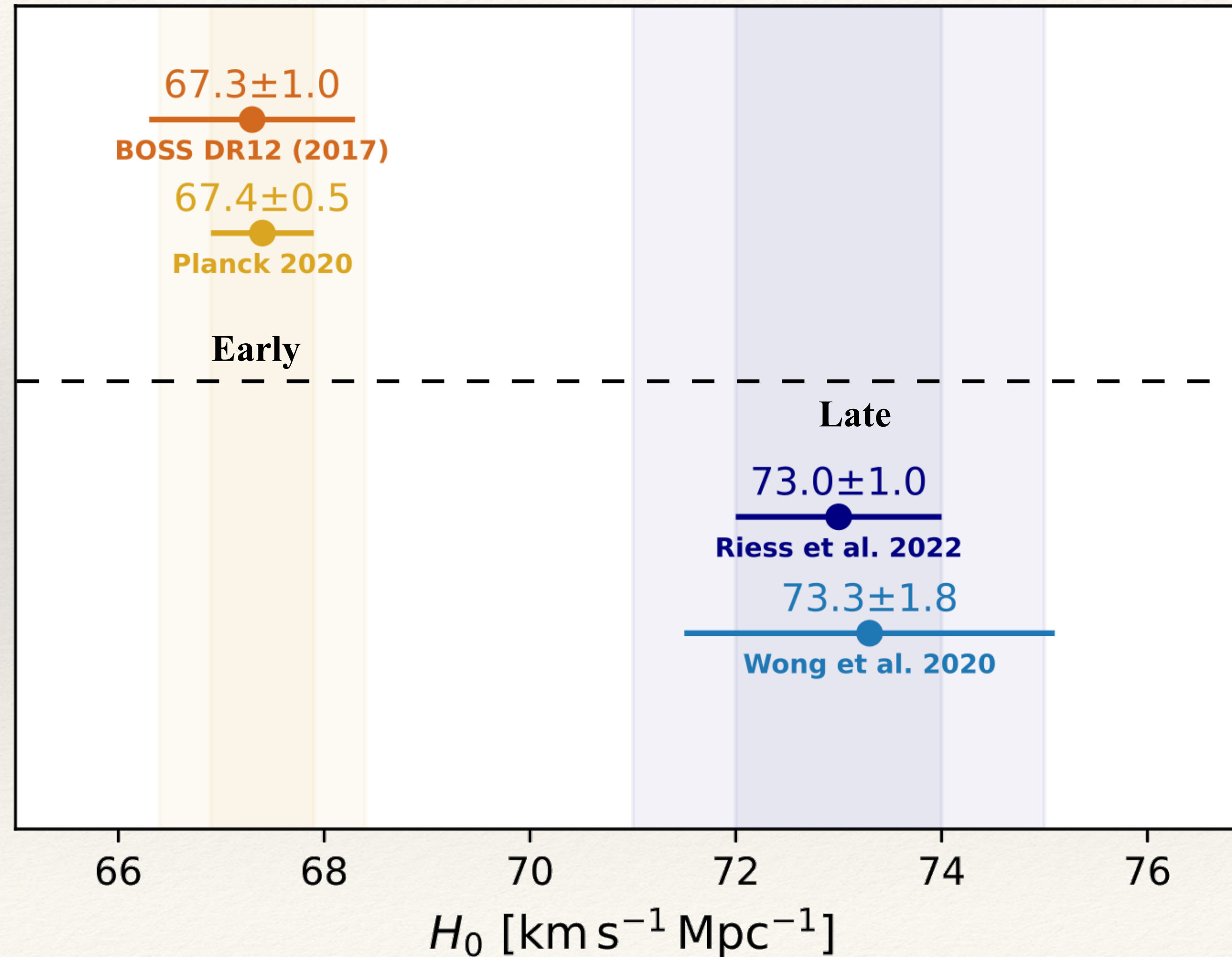
Mega Maser cosmology project



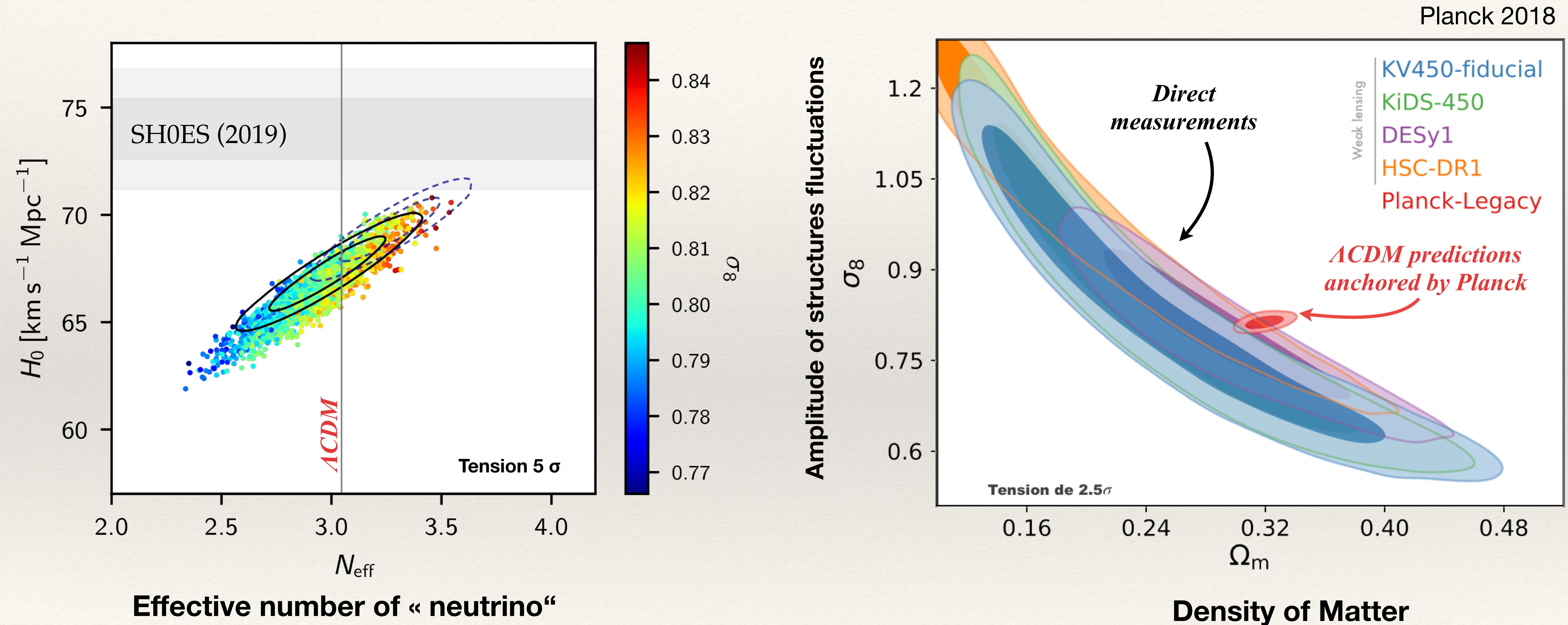
H_0 Tension | *alternative probes are still quite far*



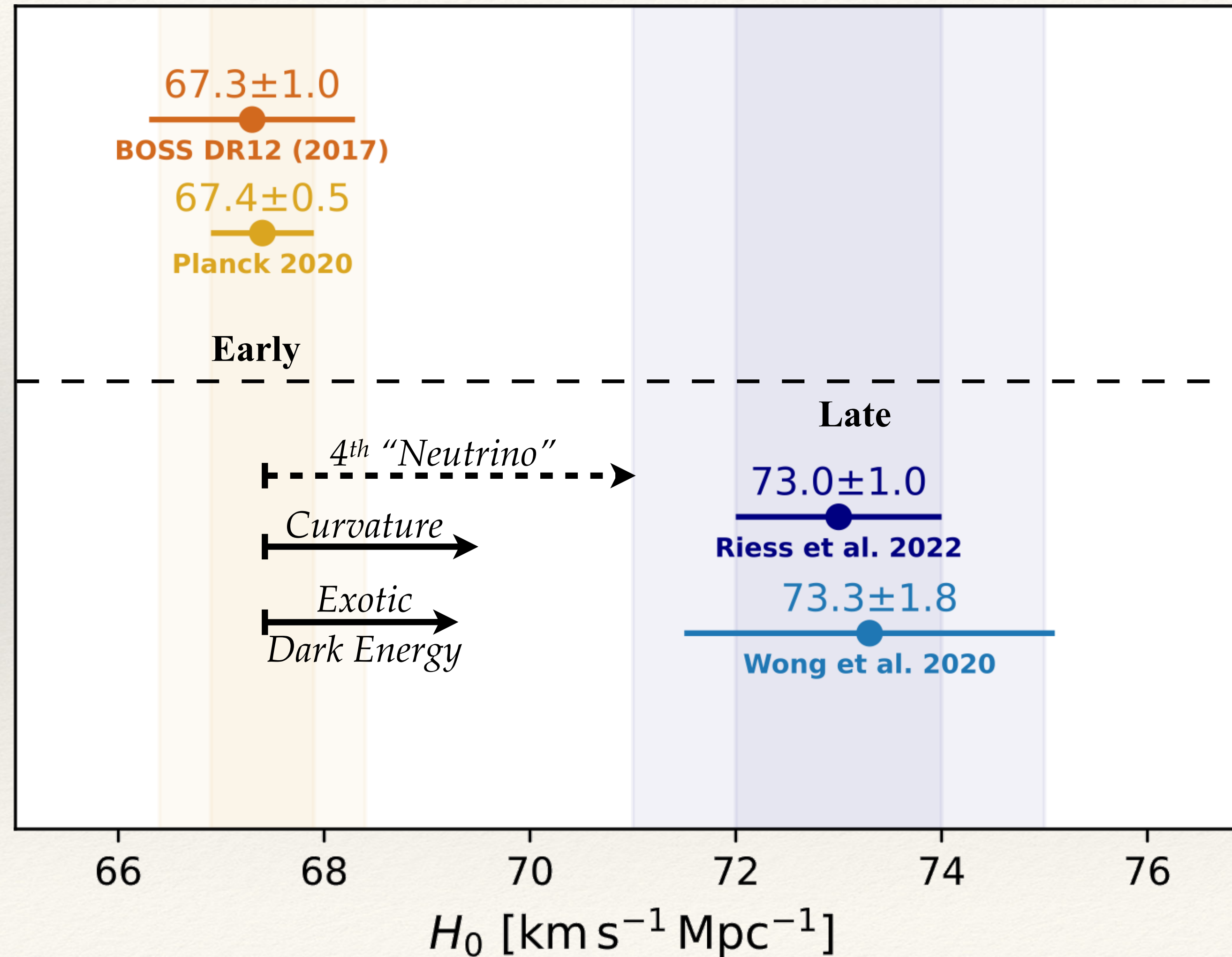
H_0 Tension | *At least 2 independent systematics*



Tensions In Cosmology | *Changing the model*



H_0 Tension | *At least 2 independent systematics*



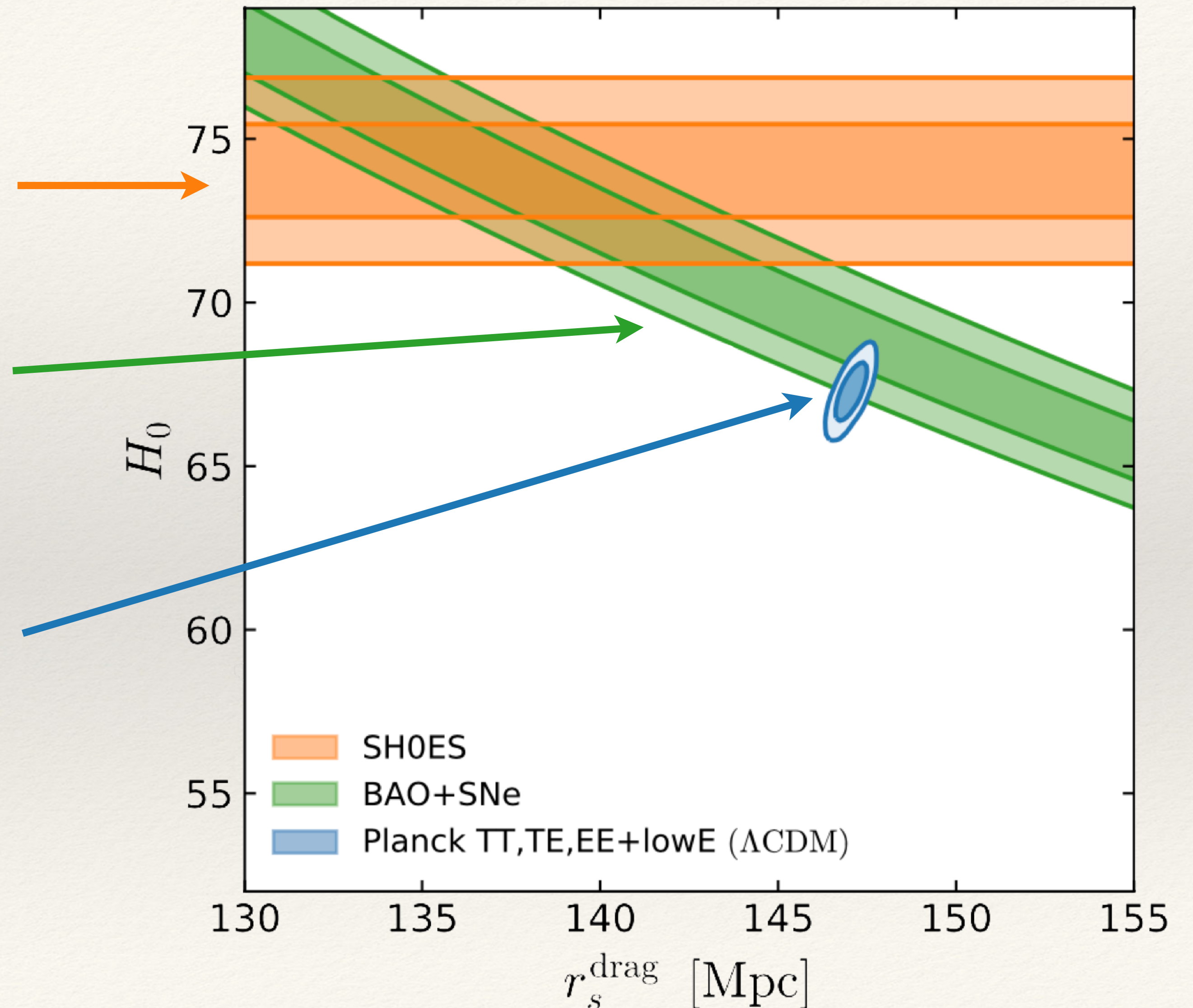
H_0 tension or r_s tension ?

Knox & Millea et al. 2019

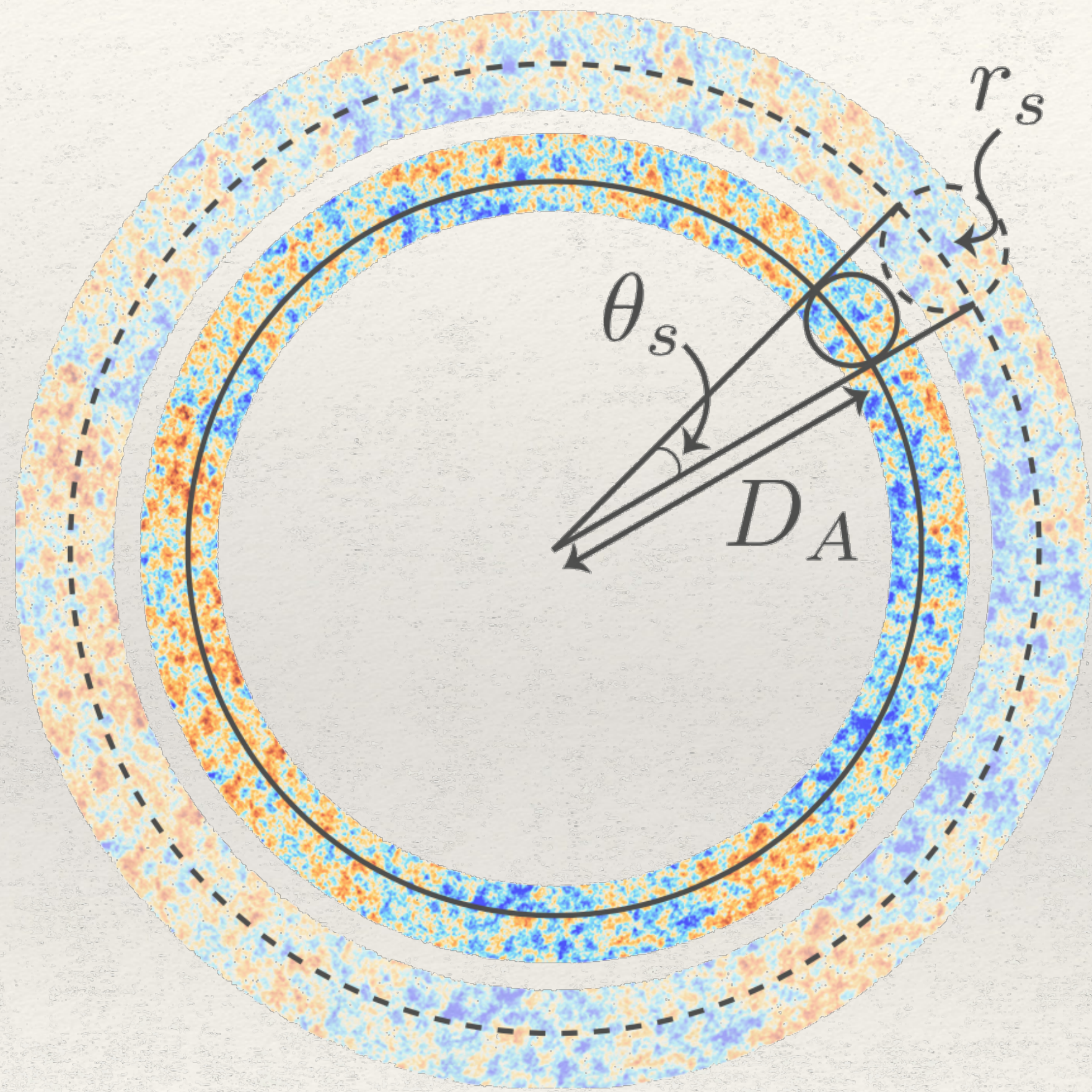
SH0ES 2019 (Cepheids & SNe Ia)
(No assumption of Λ CDM)

BAO + SNe Ia
(No assumption of Λ CDM | 5d spline)

Planck
(Assumes Λ CDM)

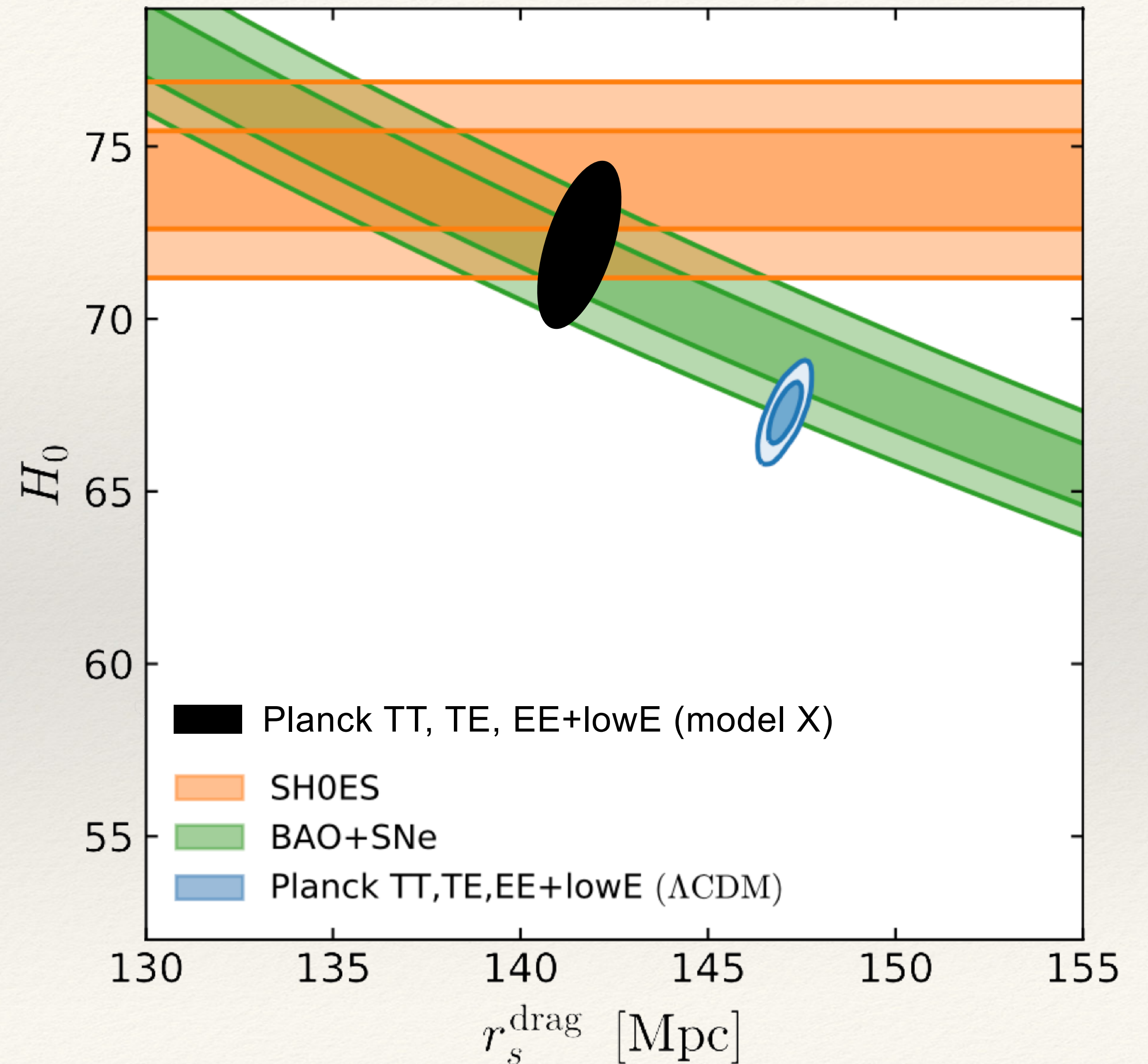


Then what about New Fundamental Physics ?



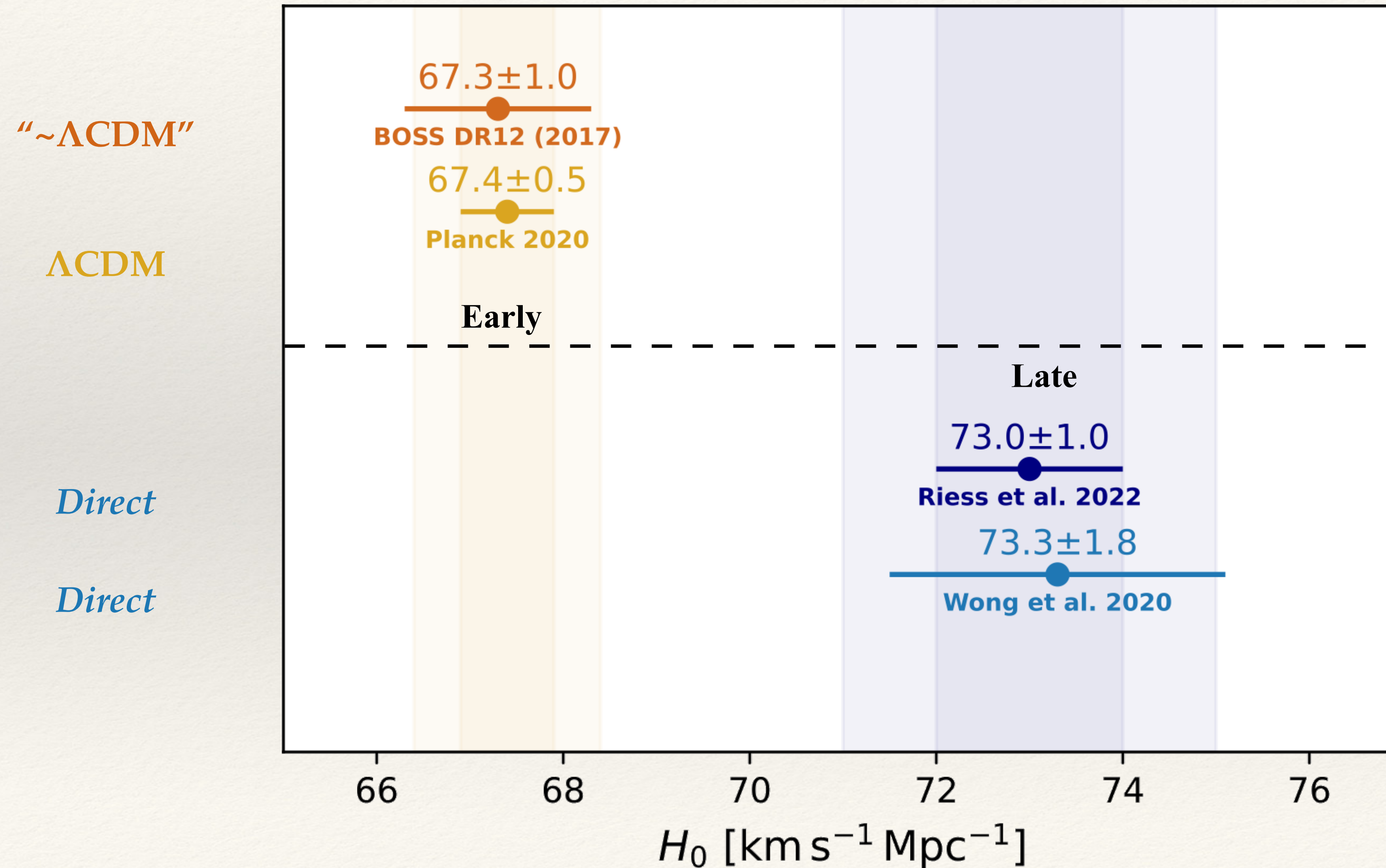
T. Smith | V. Poulin

Knox & Millea et al. 2019



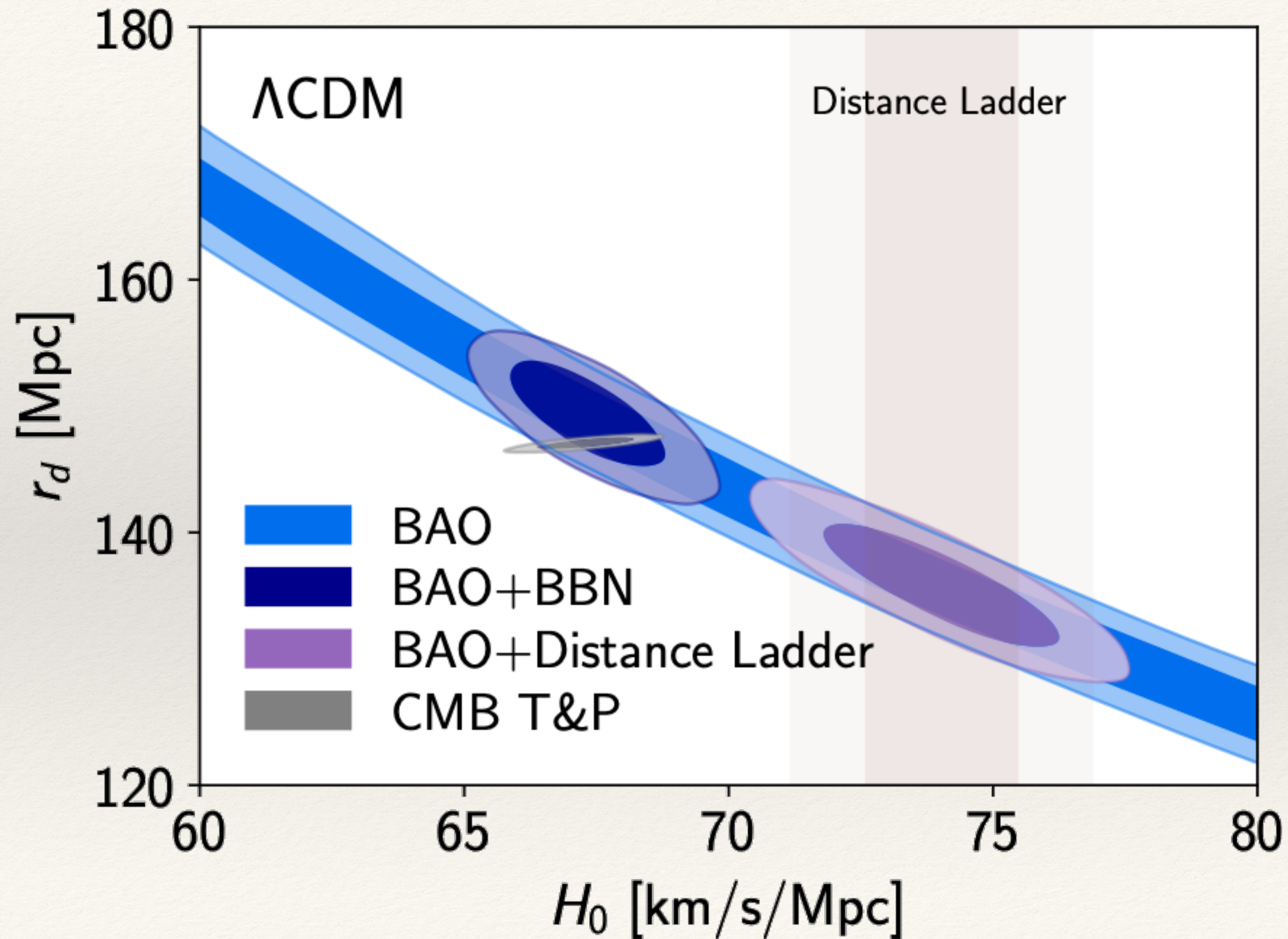
...or it's systematics

H_0 Tension | *At least 2 independent systematics*

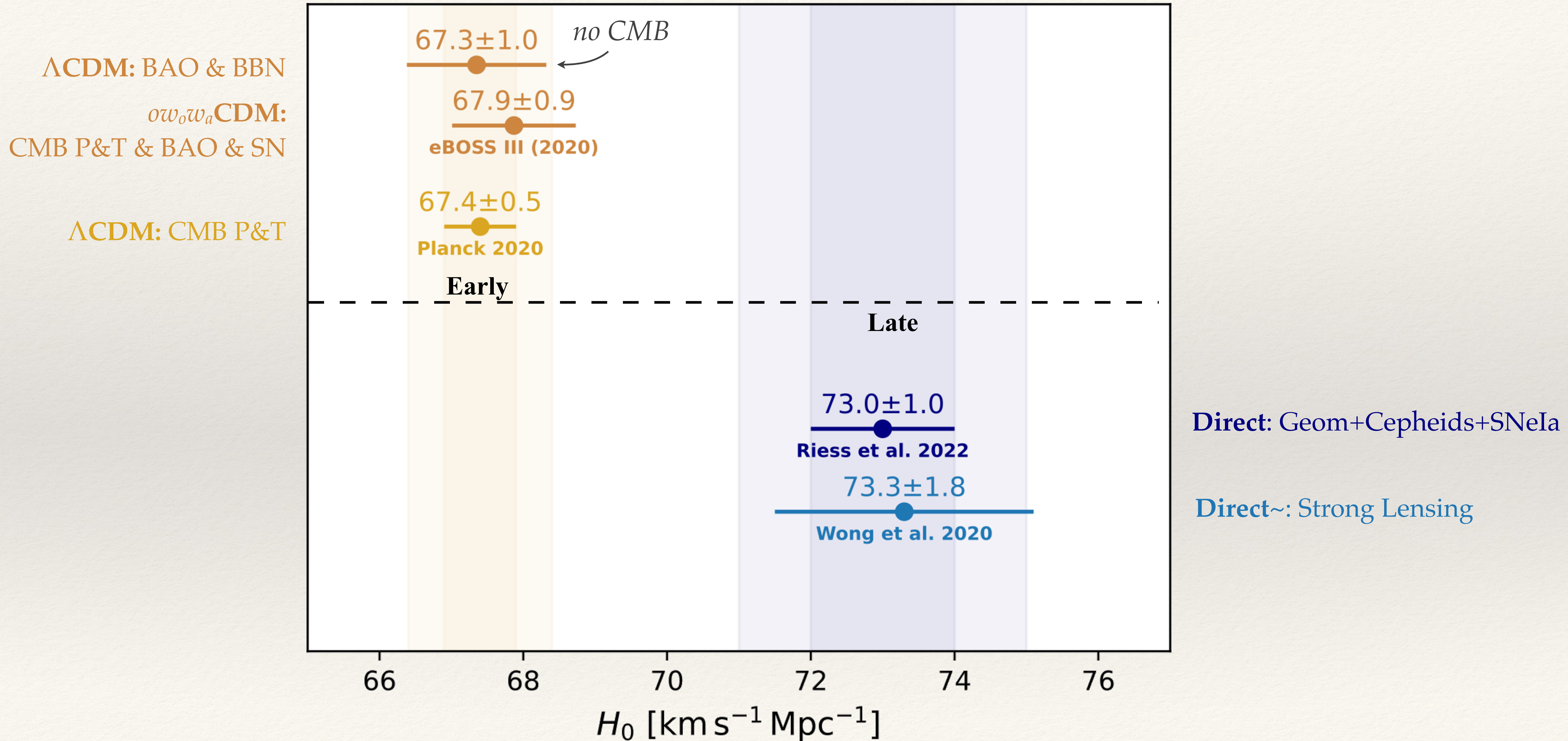


No CMB Data | *Big Band Nucleosynthesis*

eBOSS collab. 2020



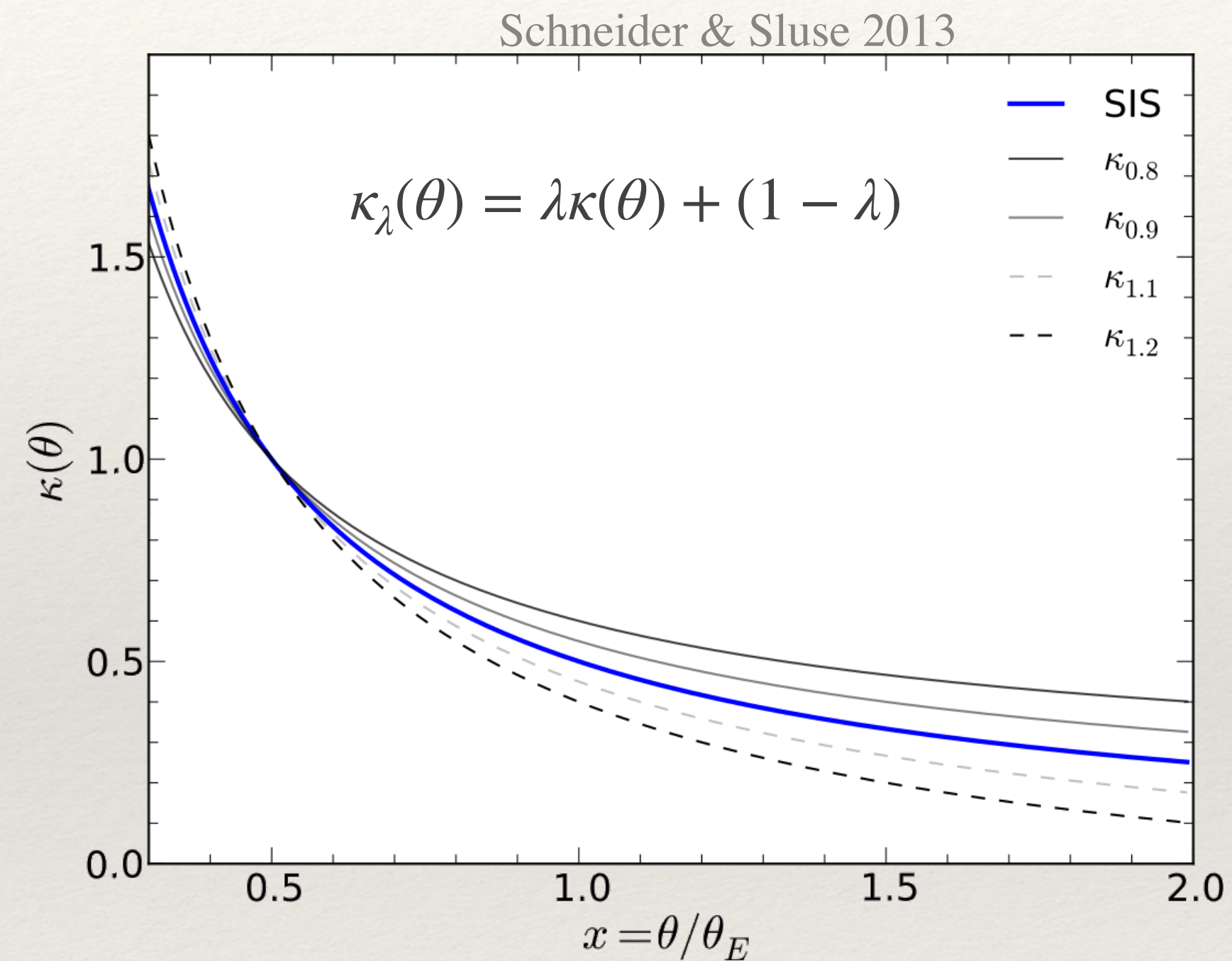
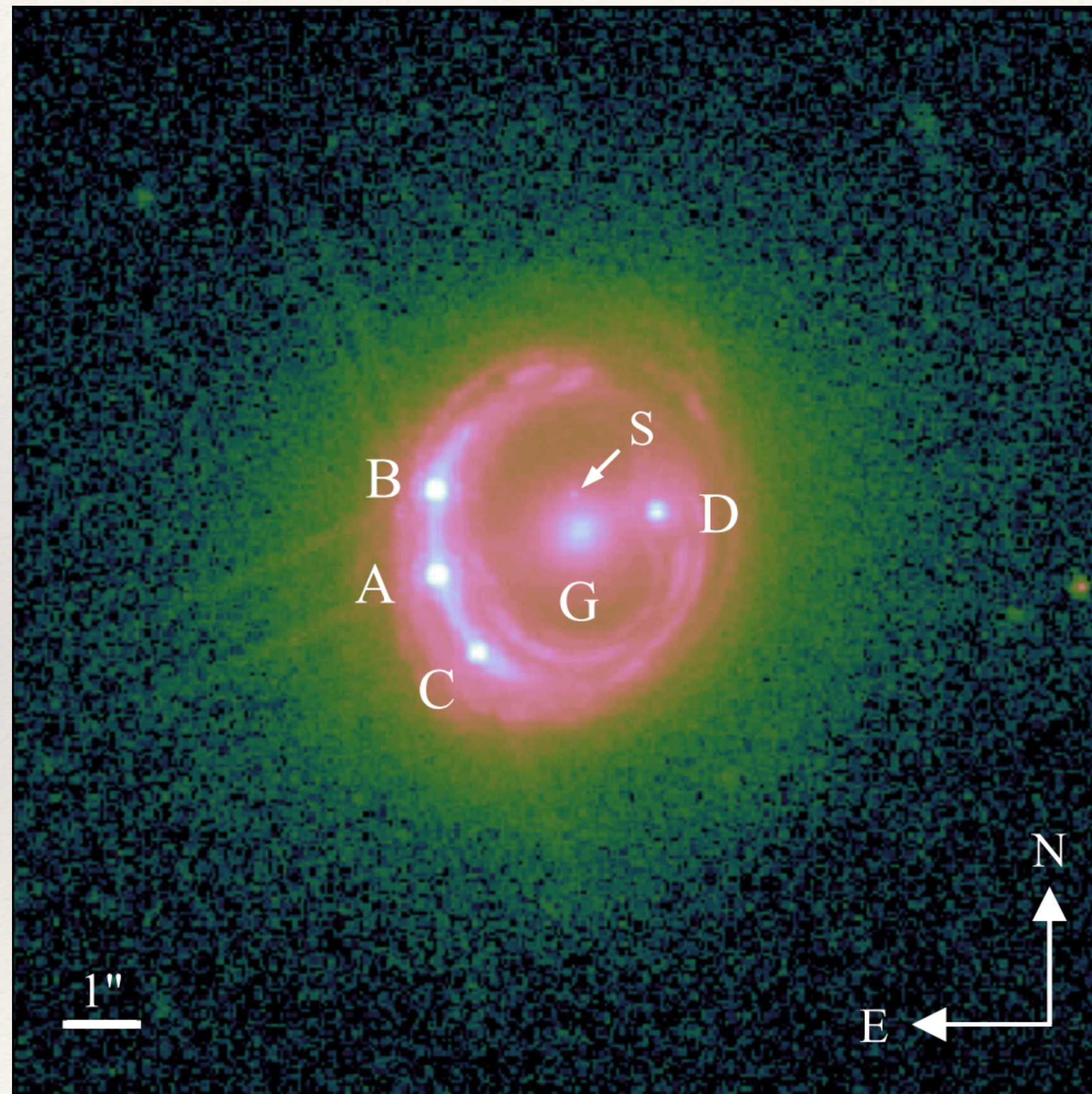
H_0 Tension | *At least 2 independent (late) systematics*



Strong Lensing | *systematics in density profile*

The mass-sheet degeneracy:

from the observed image positions and flux ratios, one cannot distinguish between the original κ and any κ_λ



“if the SIS density profile provides a good fit to the lensing data, an equally good fit is obtained by all the κ_λ .”

But:

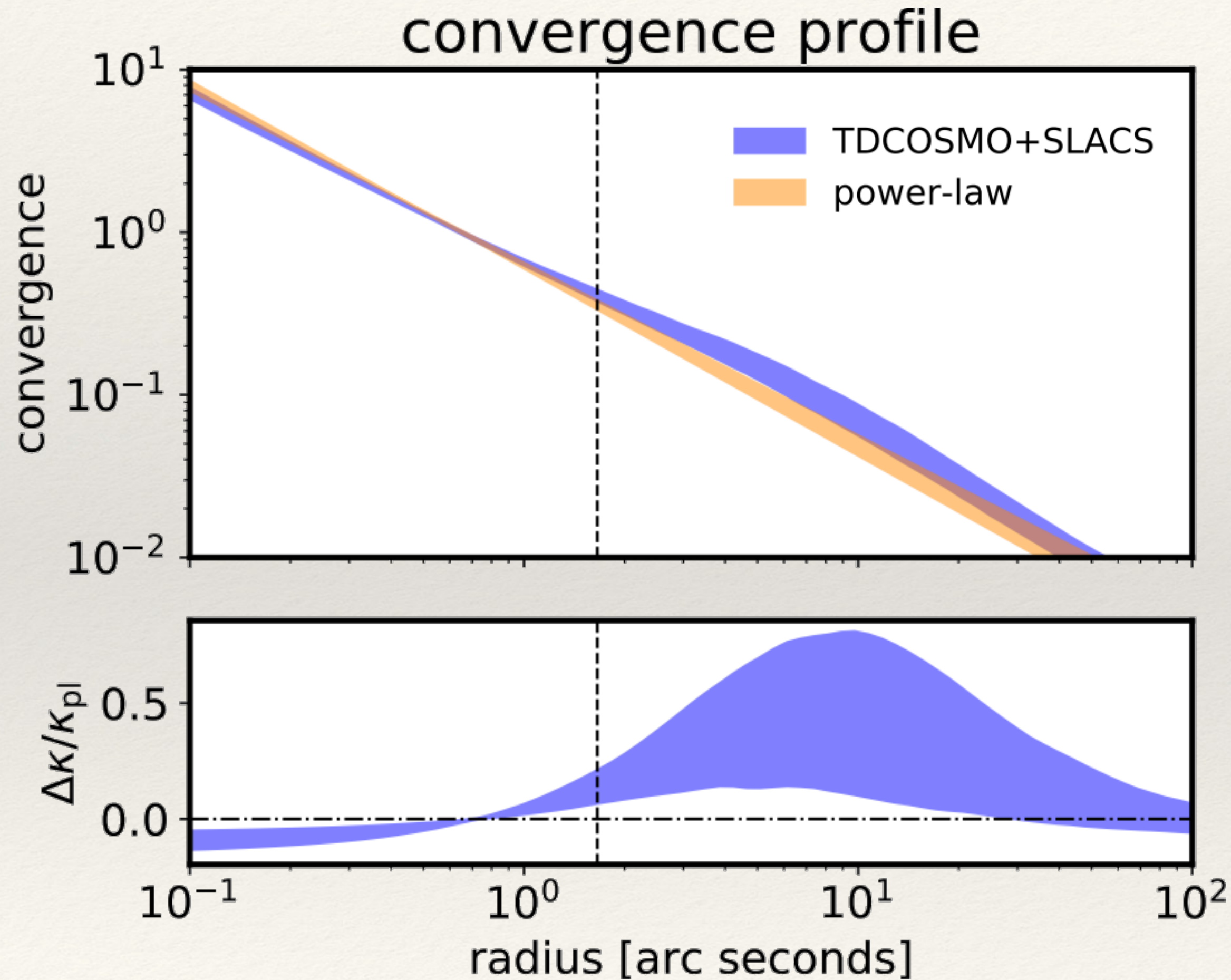
$$H_0\Delta t \rightarrow \lambda H_0\Delta t$$

Good news:

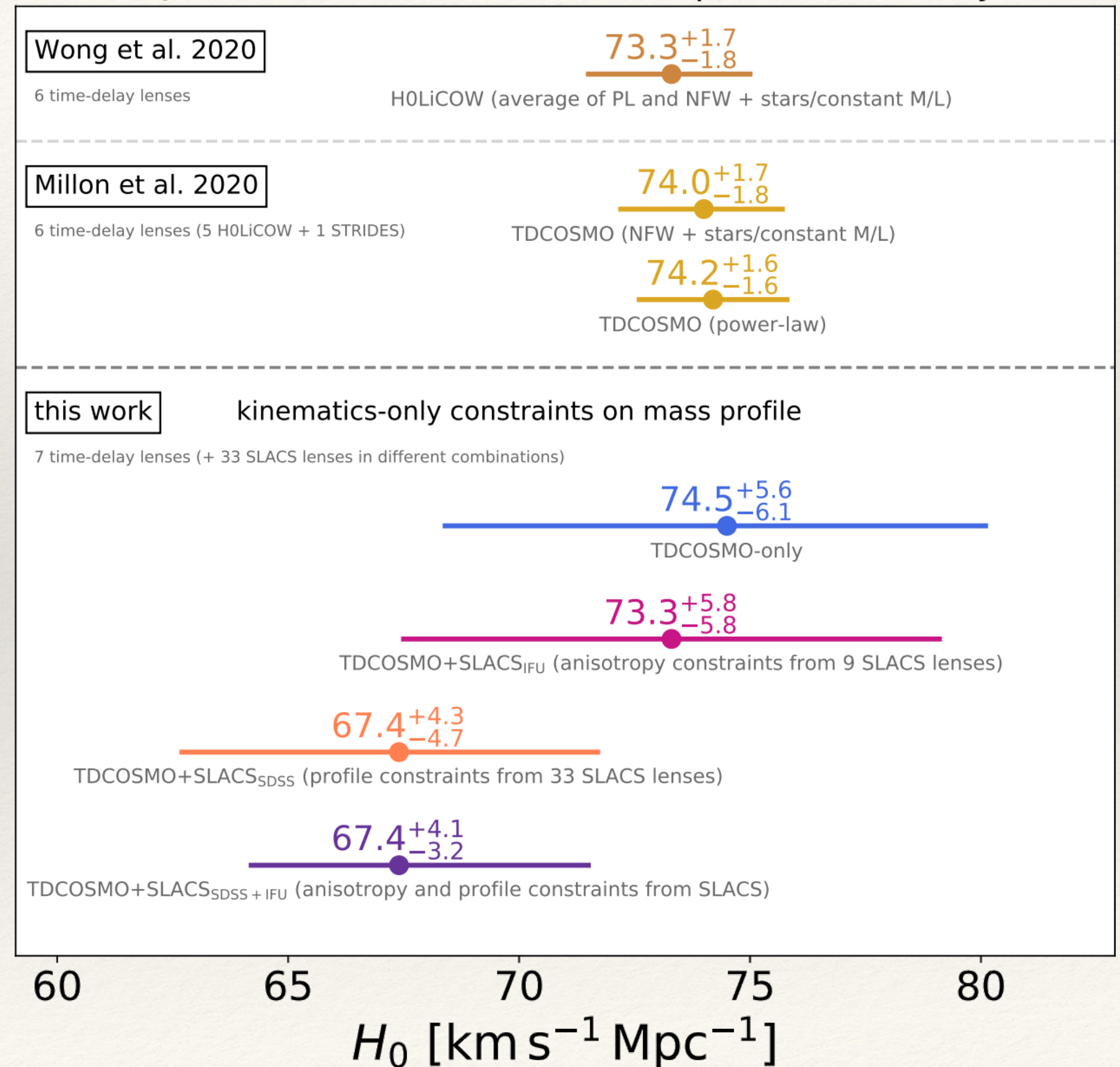
The mass sheet degeneracy can be broken by spatially resolved velocity dispersion measurements

Strong Lensing | *systematics in density profile*

Birrer et al. 2020

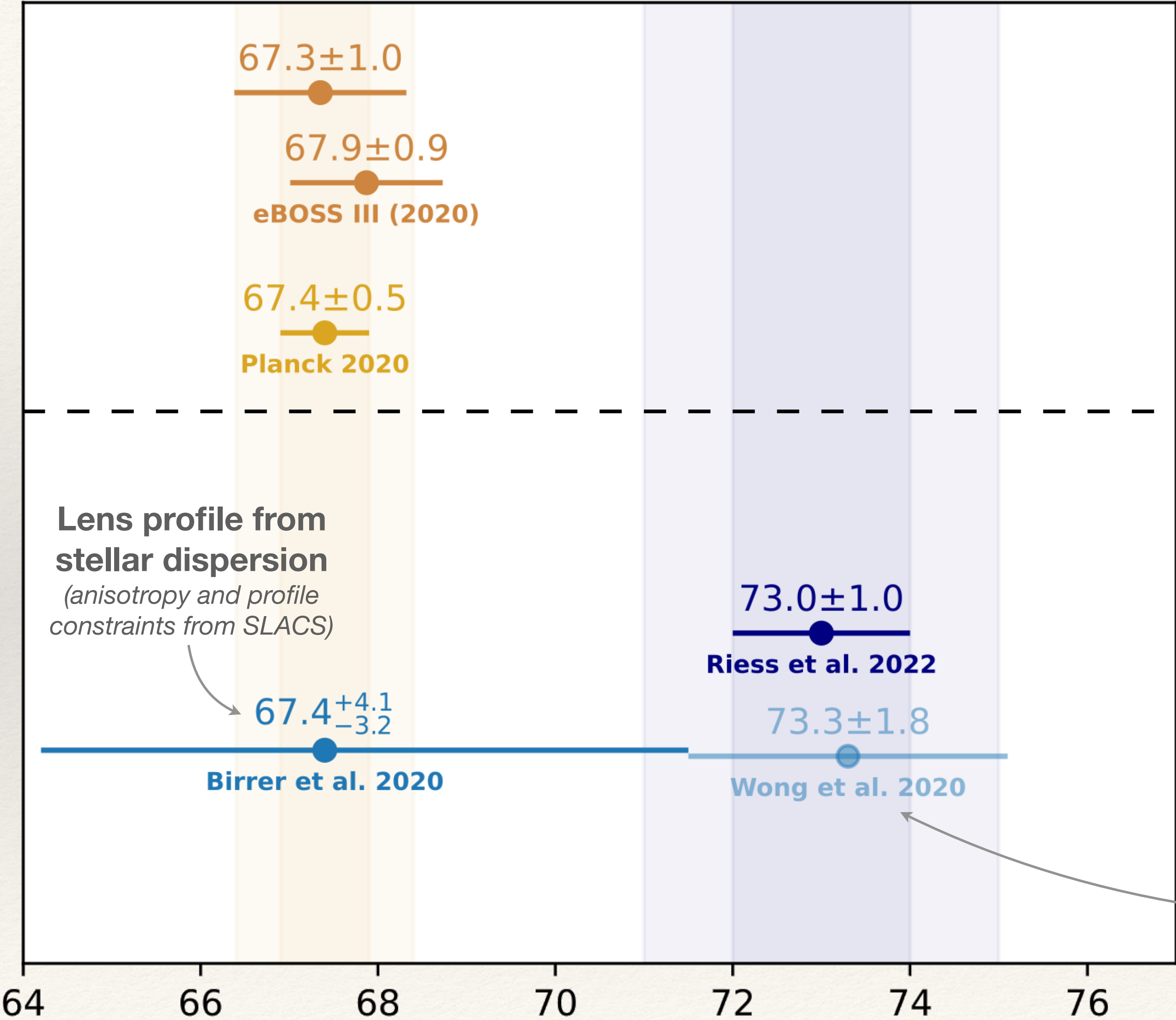


H_0 measurements in flat Λ CDM - performed blindly



H_0 Tension | *Highly debated systematics in strong lensing*

Λ CDM: BAO & BBN
 $\sigma\omega_0\omega_a$ CDM:
 CMB P&T & BAO & SN
 Λ CDM: CMB P&T



Lens profile from stellar dispersion
(anisotropy and profile constraints from SLACS)

Direct: Geom+Cepheids+SNeIa

Direct~: Strong Lensing

Power-law & star constant M/L & Dark matter halo
(See Suyu 2009)

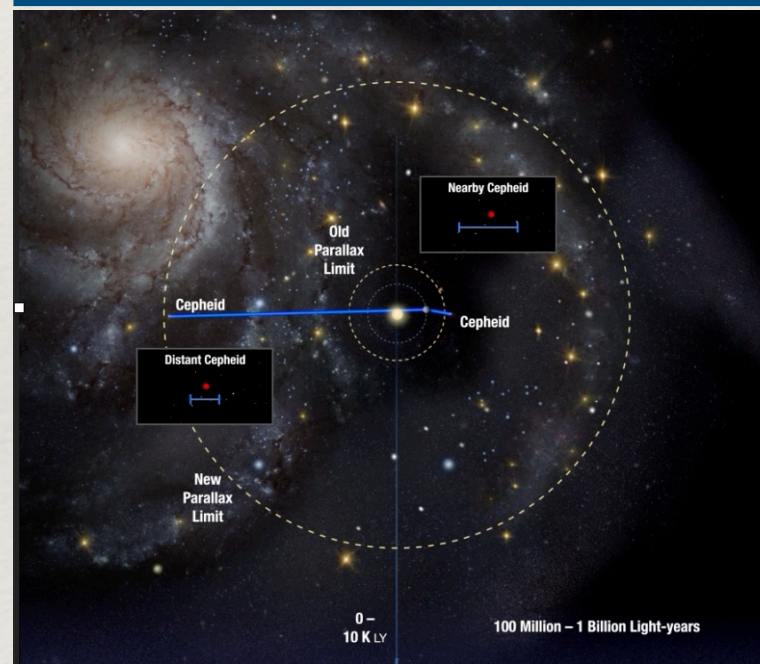
Direct Distance Ladder | TRGB

Get independent distances for SNe Ia



"Geometry"

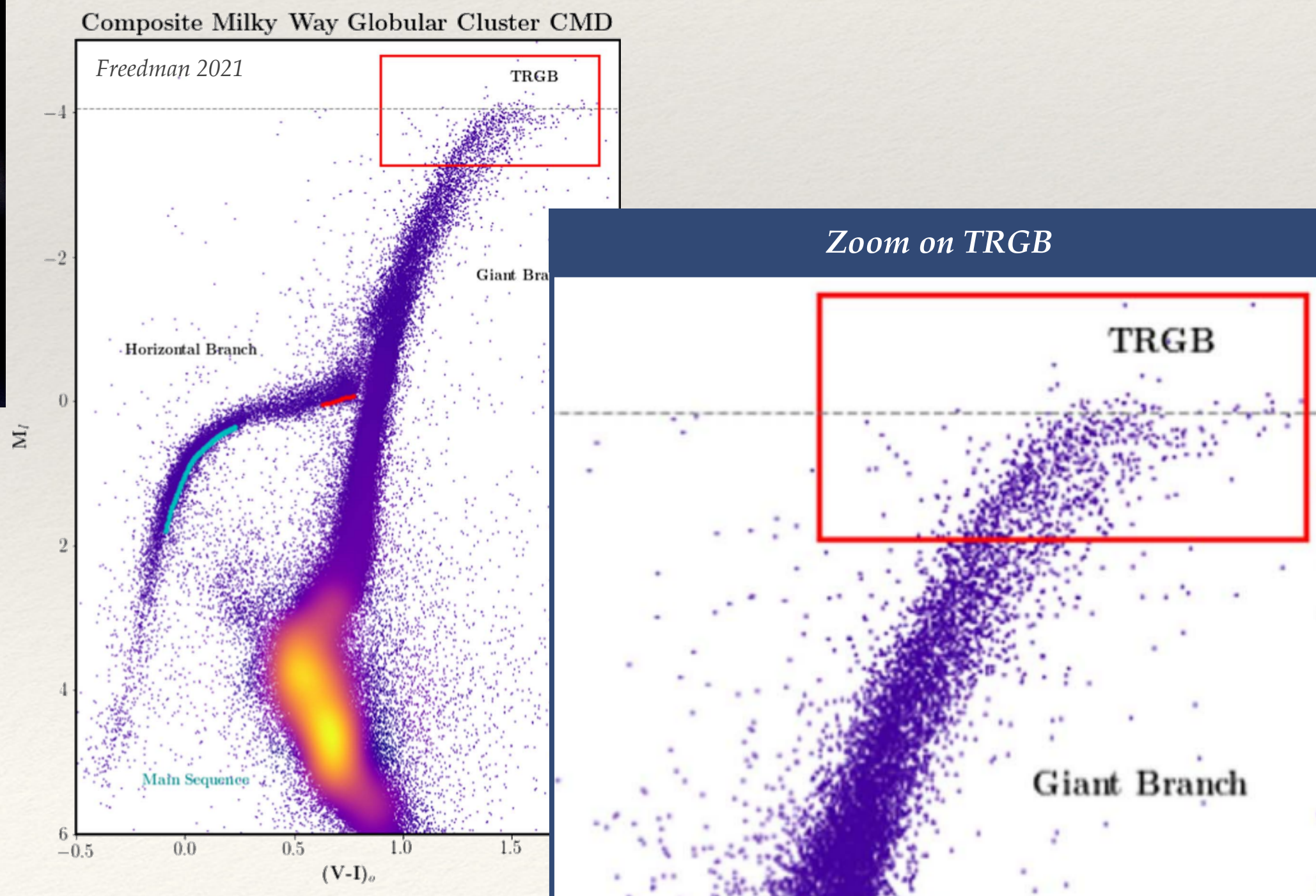
Parallaxes | D.E.B. | Maser



Source: eso

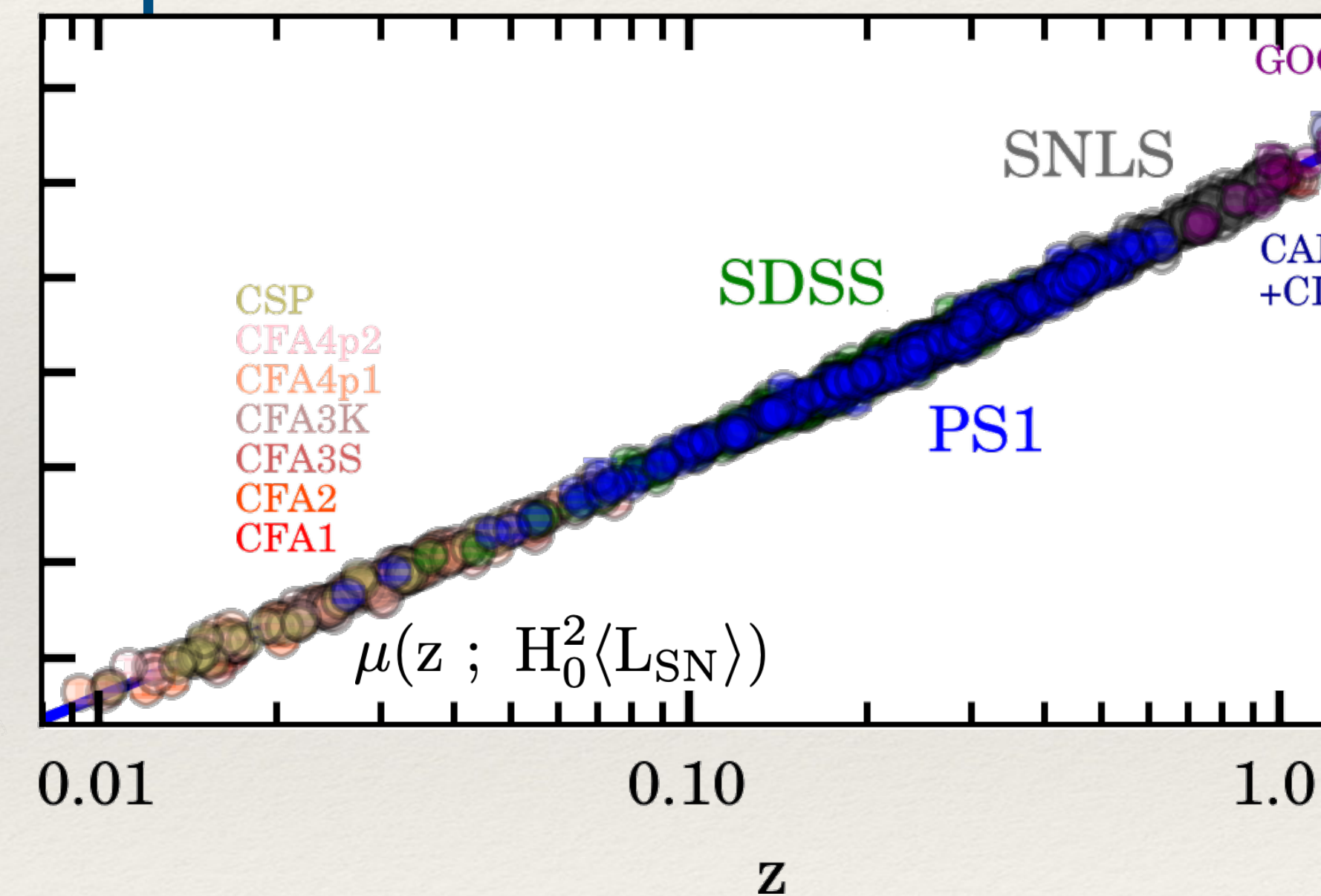
"Calibrators"

TRGB



"SNe Ia"

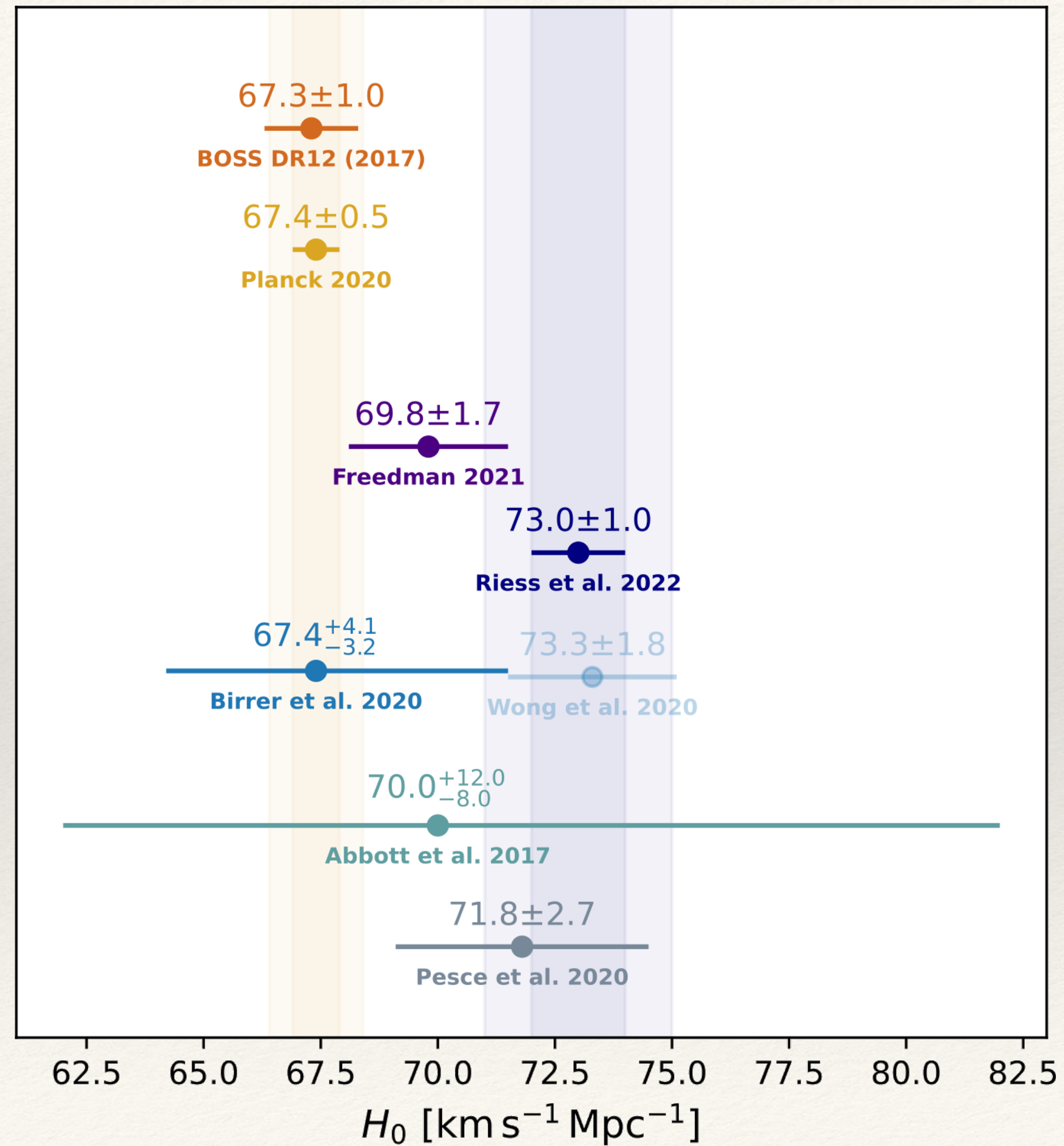
Distance Modulus (mag)



Freedman et al. 2021

$$H_0 = 69.8 \pm 0.6 \text{ (stat)} \pm 1.6 \text{ (sys)} \text{ km s}^{-1} \text{ Mpc}^{-1}$$

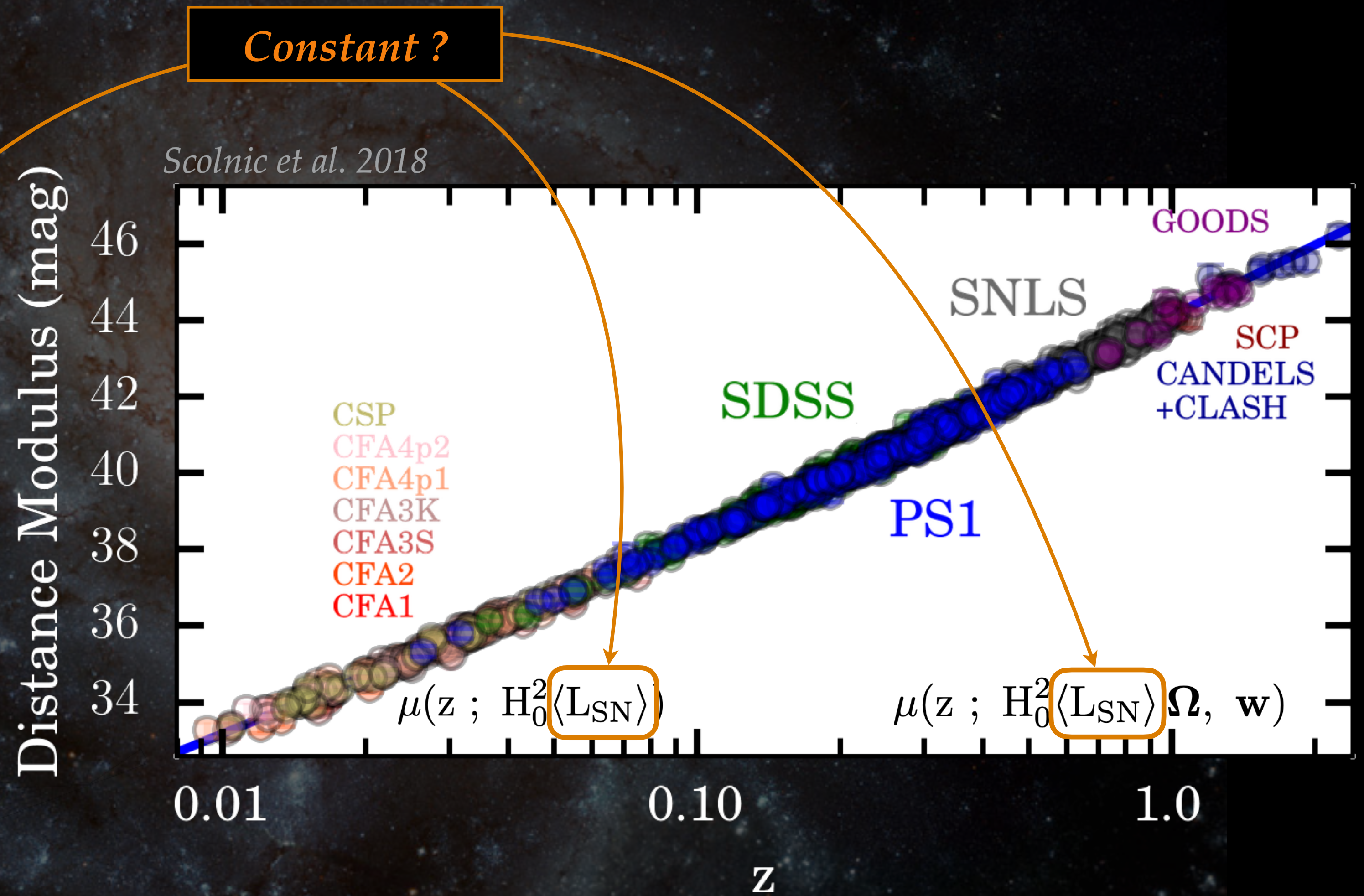
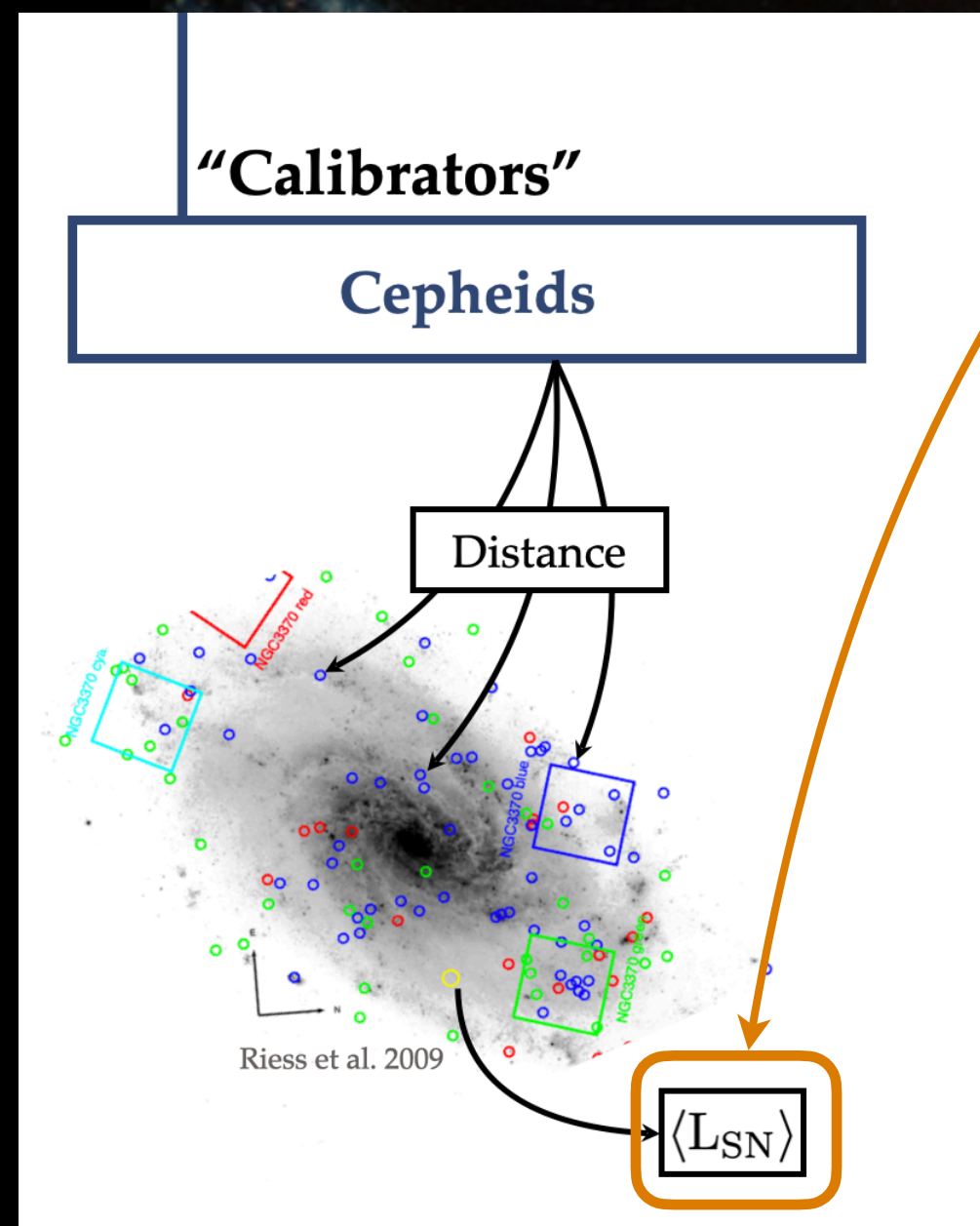
H_0 Tension | *TRGB* vs. *Cepheid*





SN2011fe →

The Progenitor issue | *Astrophysical biases*



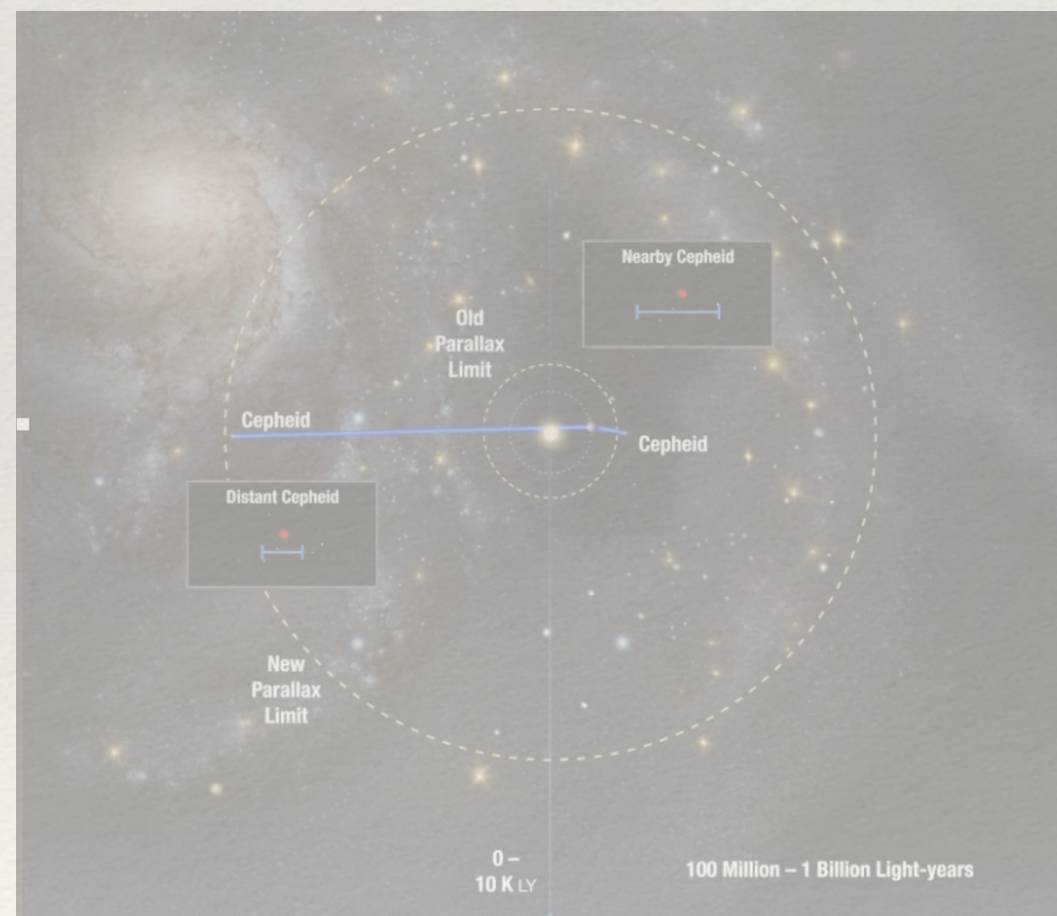
SN2011fe →

Direct Distance Ladder | *SHOES*

Calibrate the
"Period-Luminosity" relation

"Geometry"

Parallaxes | D.E.B. | Maser



The SNe Ia "matching" problem

Measure " L_{SN} "

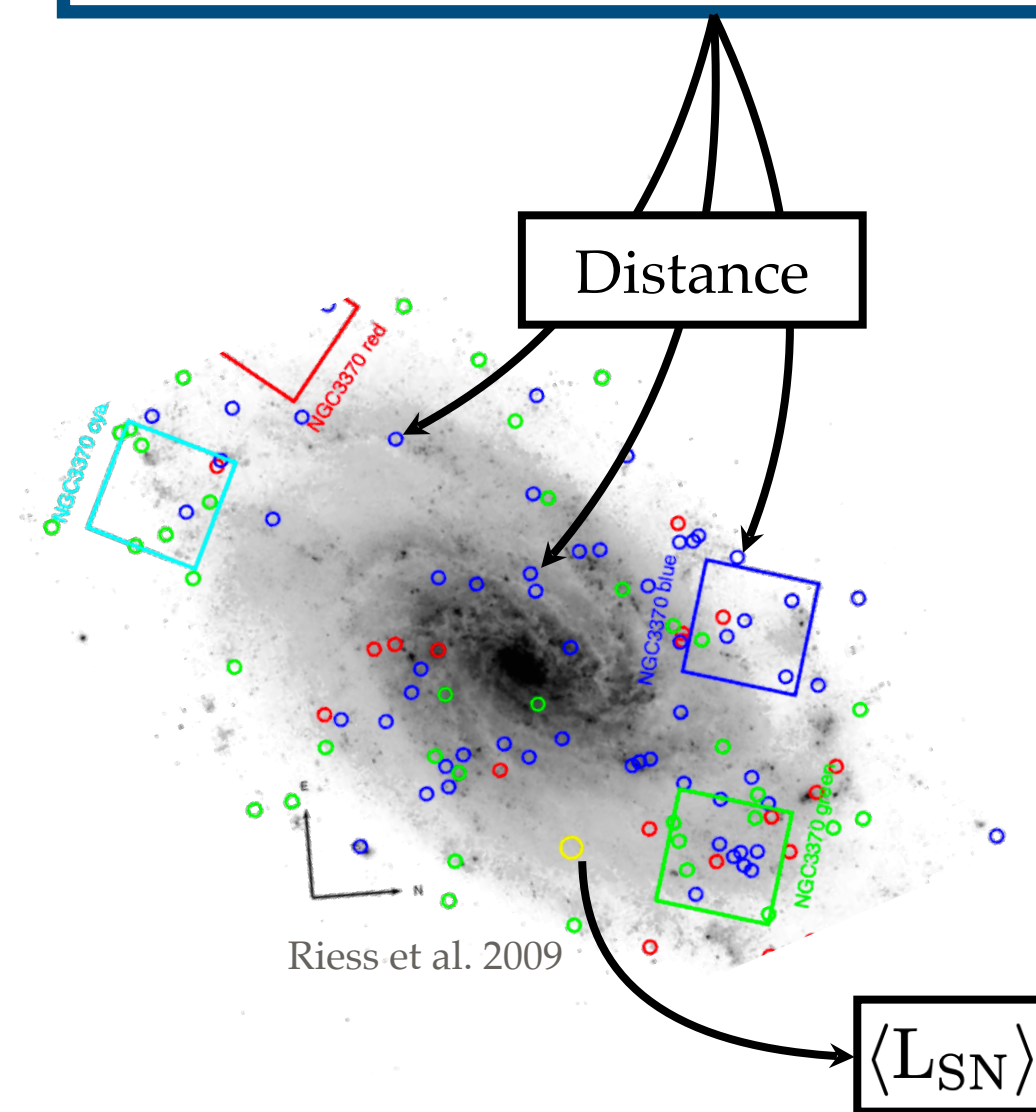
Get " H_0 "

distance

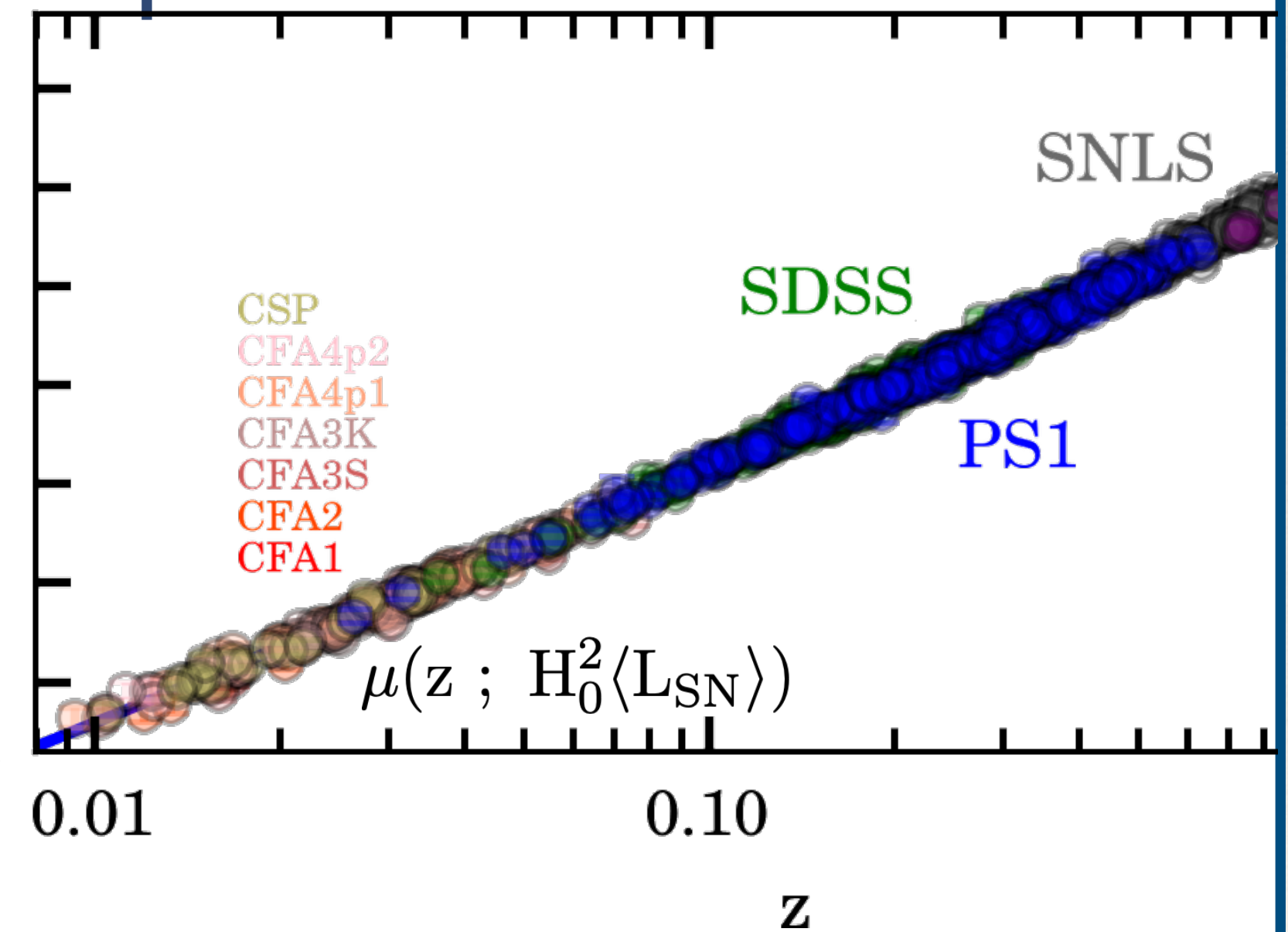
"Calibrators"

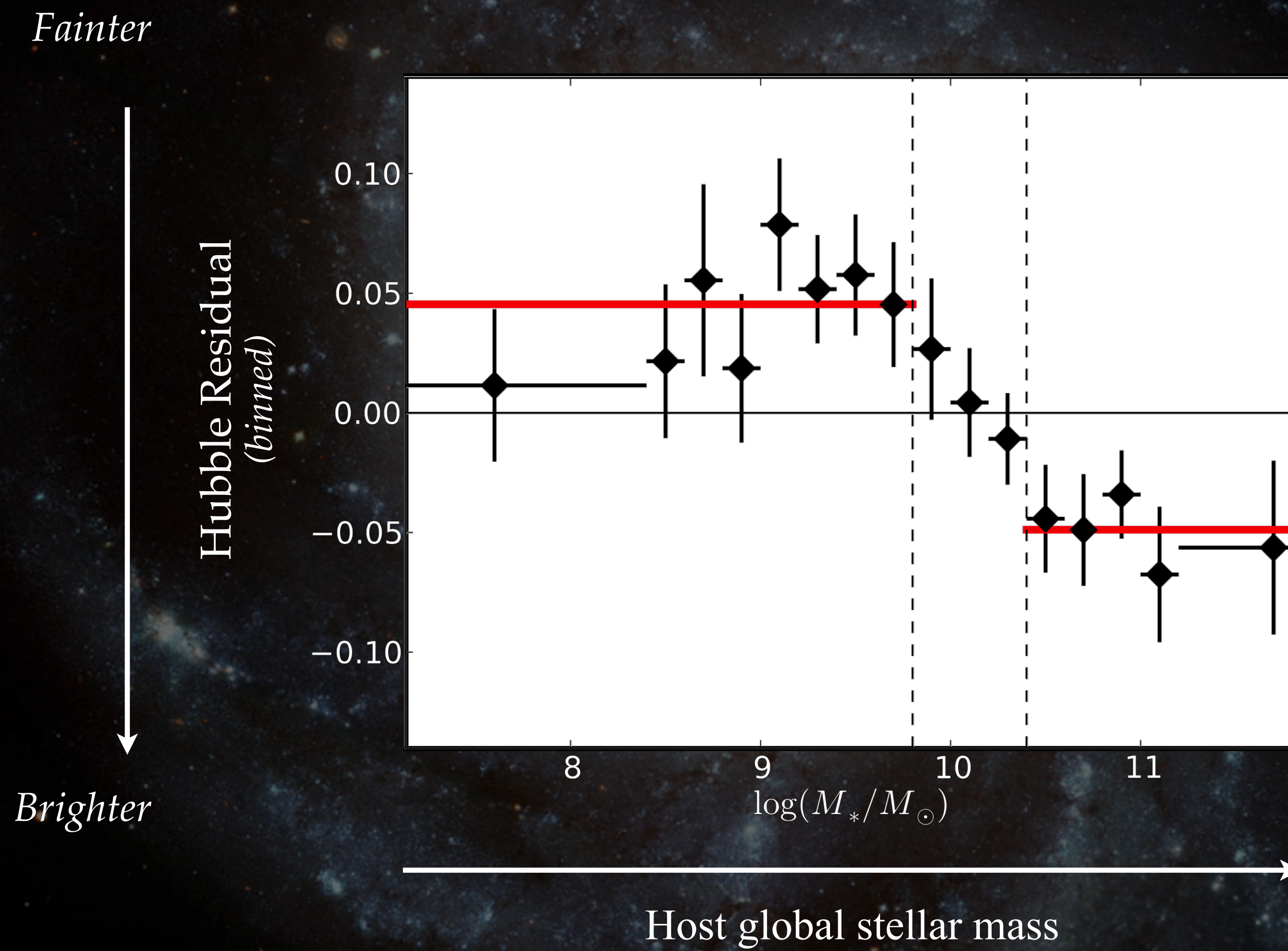
Cepheids

"SNe Ia"



Distance Modulus (mag)

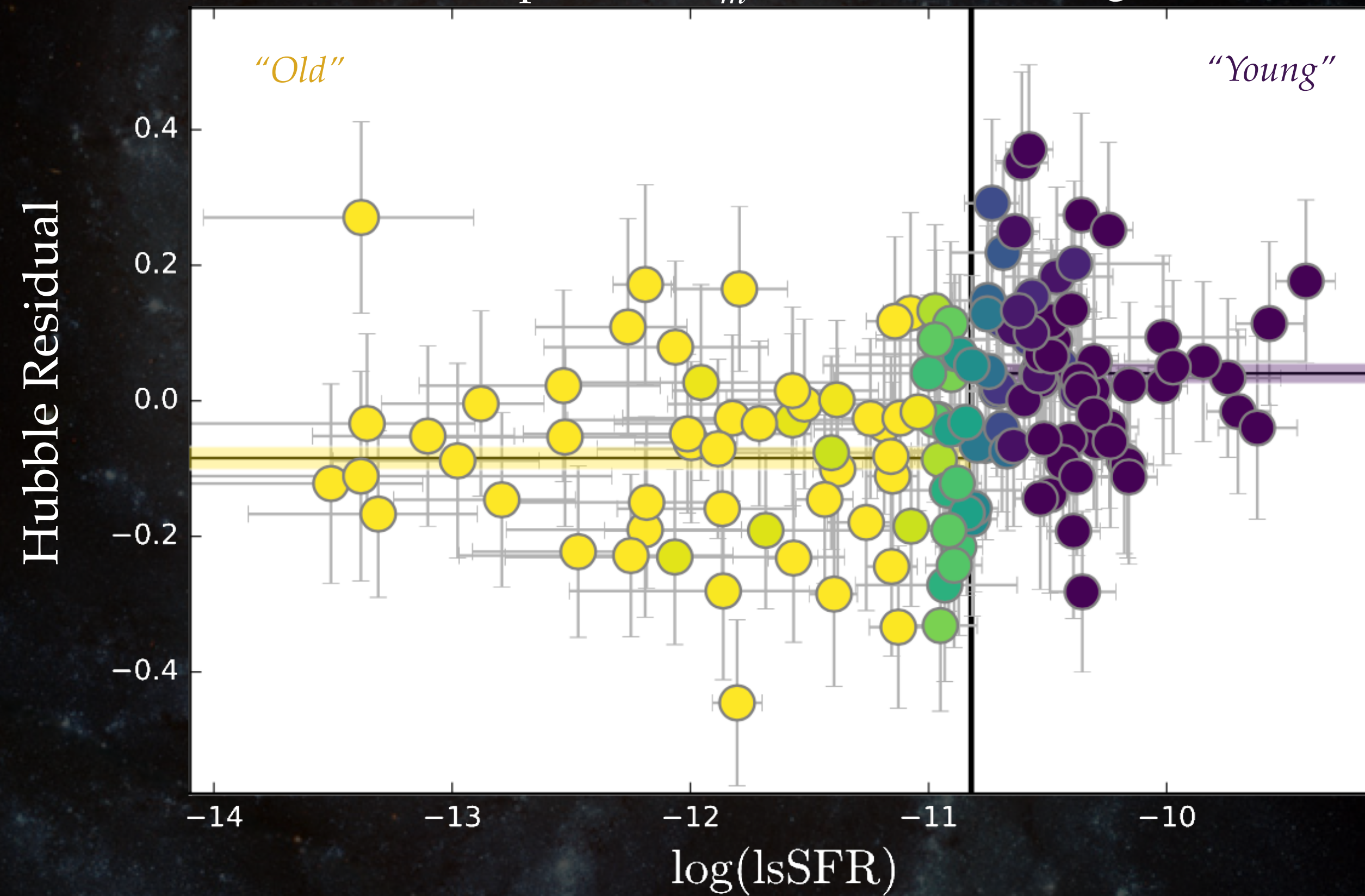




"No" young stars \longrightarrow *High fraction of young stars*

Fainter

Amplitude: $\Delta_m = 0.16 \pm 0.03$ mag



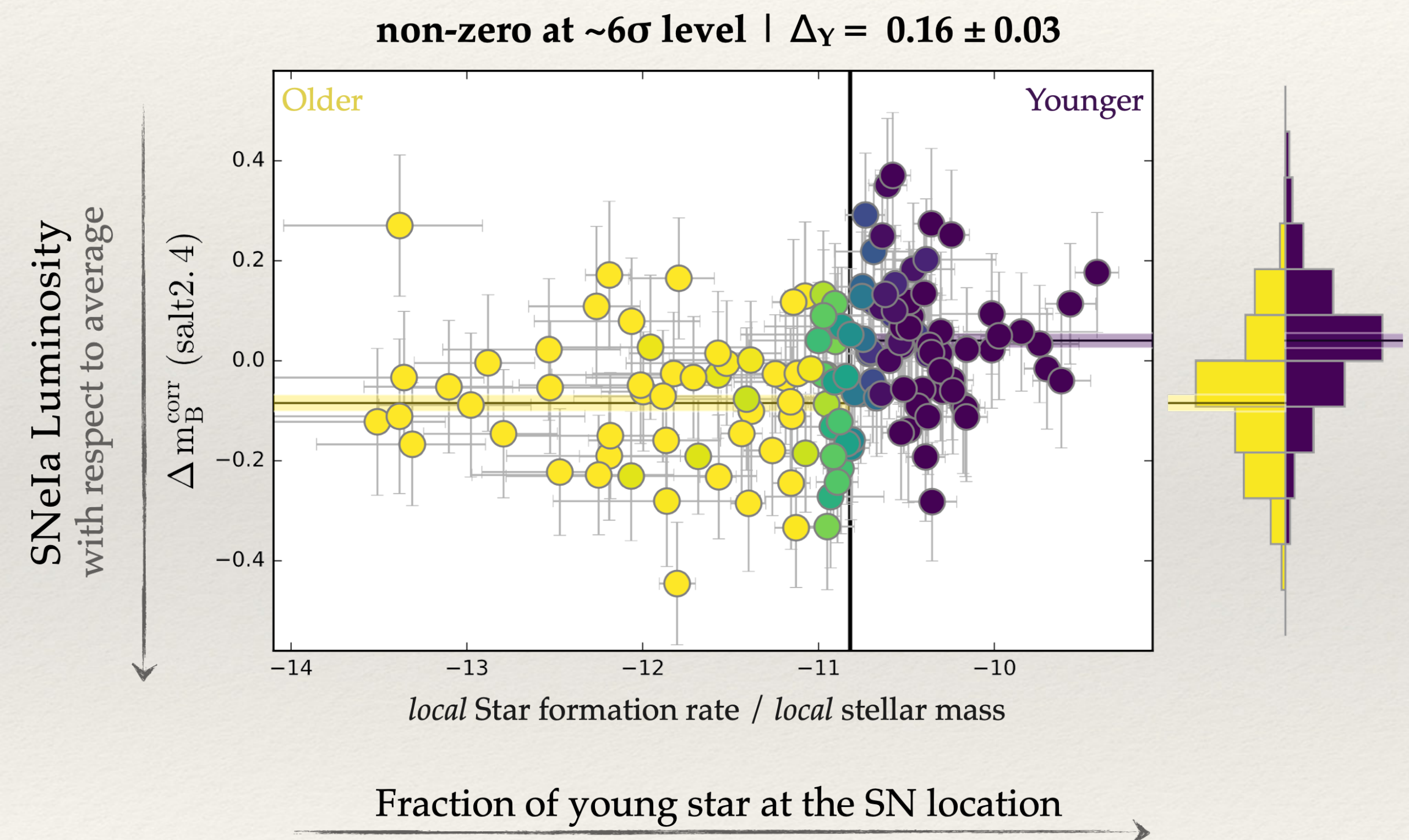
Brighter



$$\text{lsSFR} \propto \frac{\# \text{ Young Stars}}{\# \text{ Old Stars}}$$

The Age Step & H_0

Rigault et al. 2015, 2020



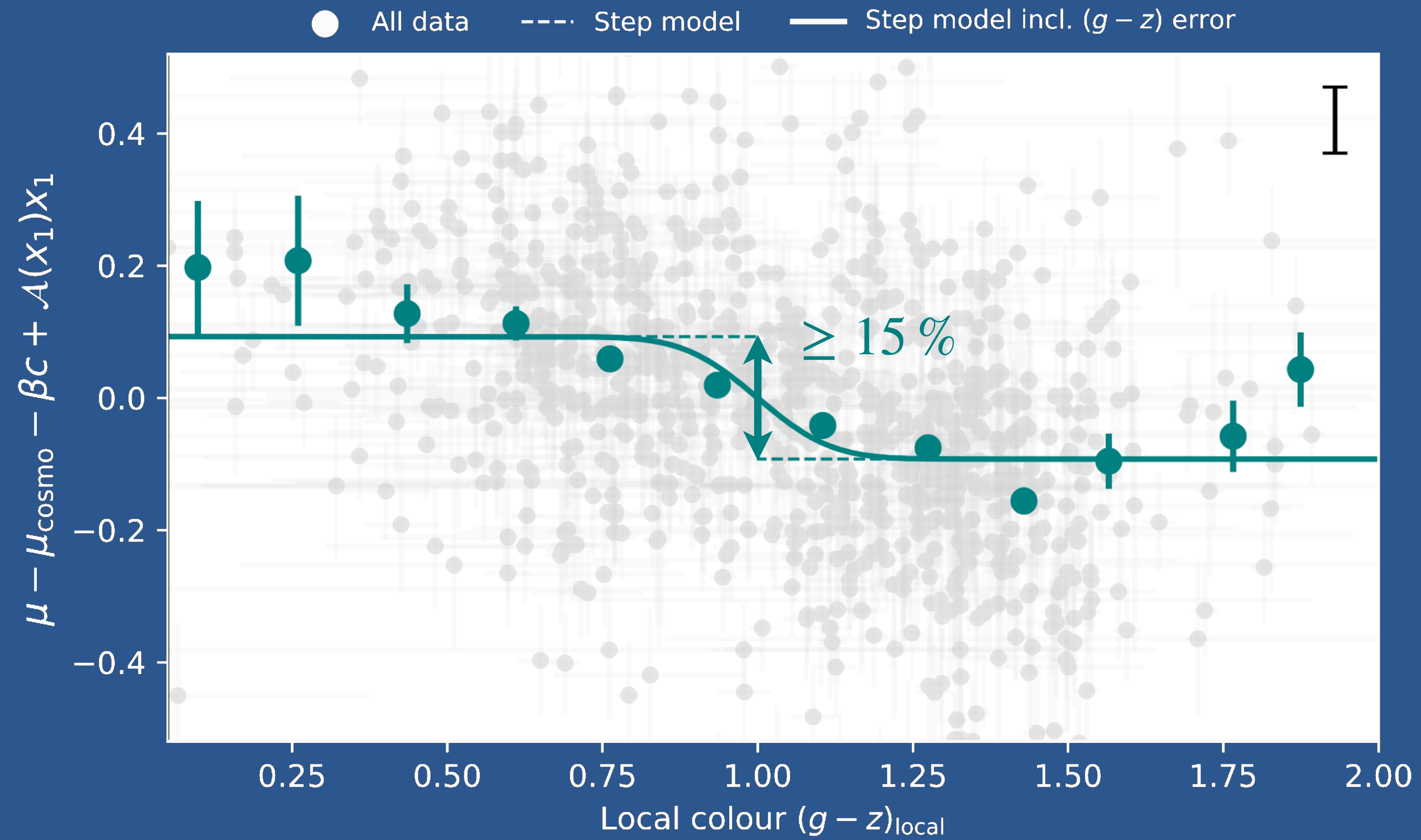
Impact on H_0 of difference in SN Population

Magnitude offset between the two SNe Ia populations

$$\log(H_0^{\text{corr}}) = \log(H_0) - \frac{1}{5} \Delta f_y \times \Delta_\gamma$$

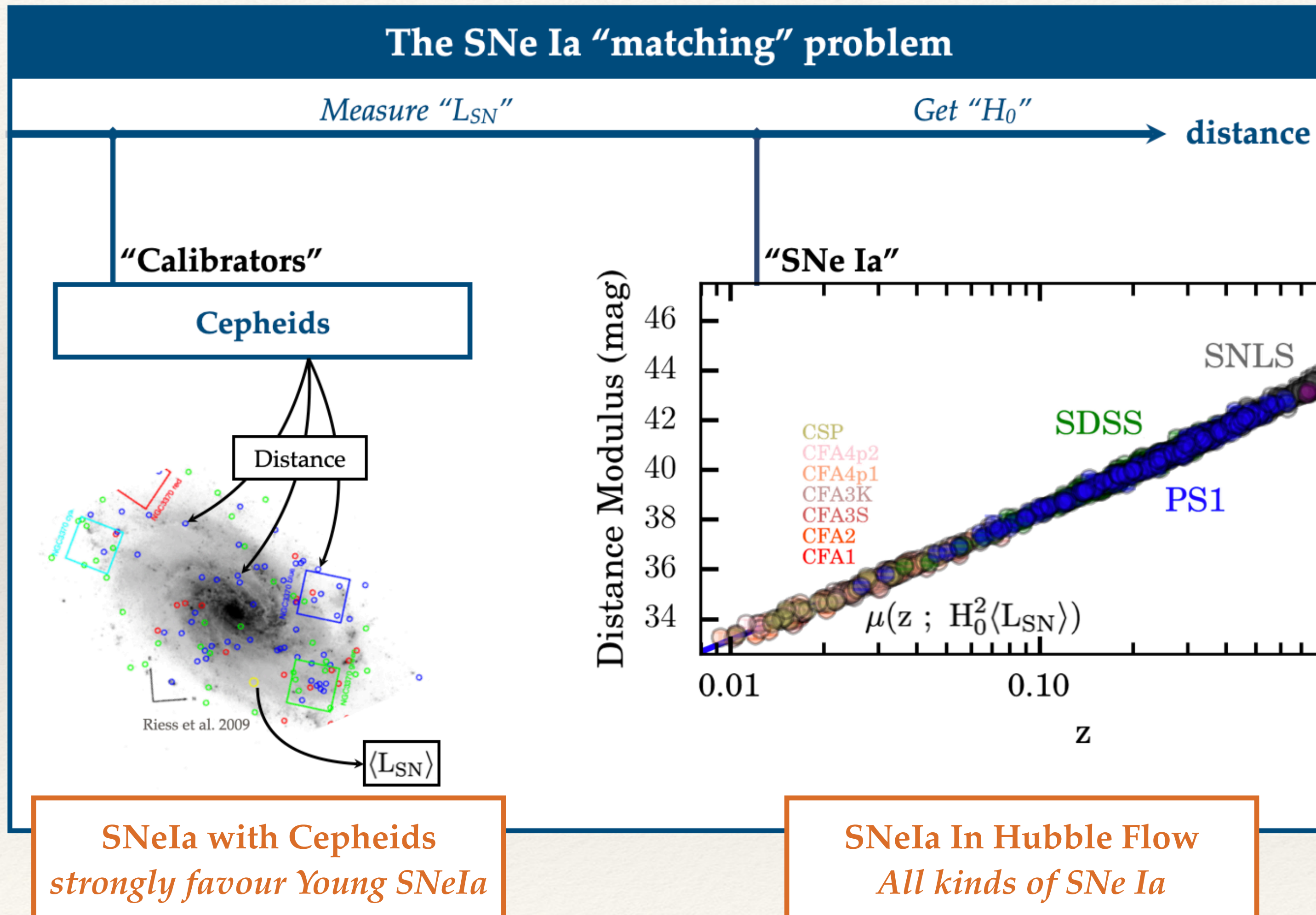
Relative fraction of Young SNe Ia between the Cepheid and HubbleFlow samples

$\sim 0.15 \text{ mag}$



Astrophysical Bias affecting H_0

Rigault et al. 2015



3% bias on H_0

So a $2 \text{ km s}^{-1} \text{ Mpc}^{-1}$ shift

Total current SH0ES error budget
 $1.04 \text{ km s}^{-1} \text{ Mpc}^{-1}$

SH0ES “corrected”
 $\sim 71 \pm 1.5 \text{ km s}^{-1} \text{ Mpc}^{-1}$

Rigault et al. in prep. | Rigault et al. 2015, 2020

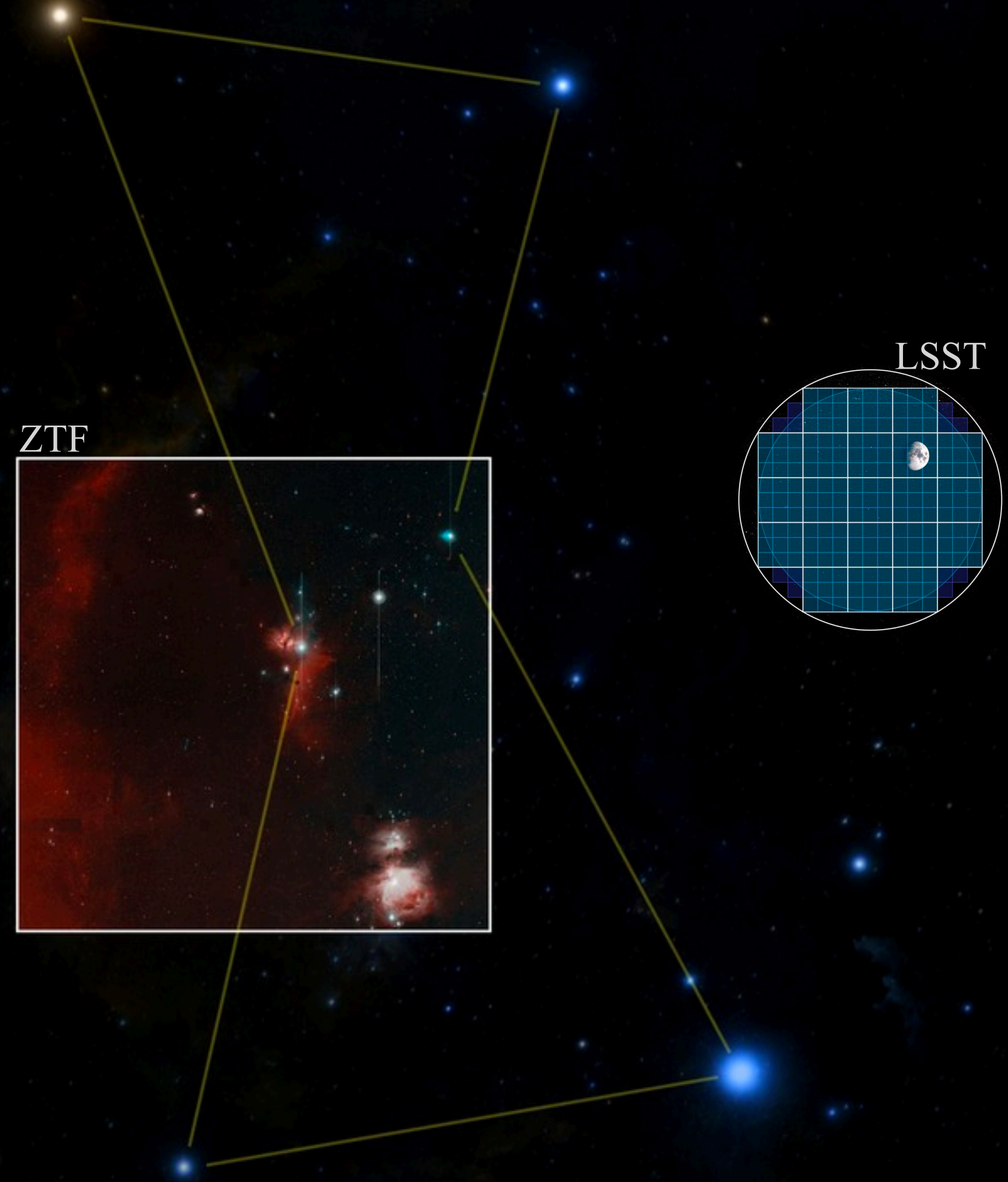
SH0ES rebuttal

“If we mimic the Cepheids selection function and only take Hubble flow SNe Ia from *Spiral* hosts, H_0 reduces by 0.5%”

Riess et al. 2022 | Riess et al. 2016, 2019

What's next ?

Zwicky Transient Facility (ZTF) is acquiring ~1000 SNeIa per year at $z < 0.1$ since 2018



ZTF H_0

Pure ZTF

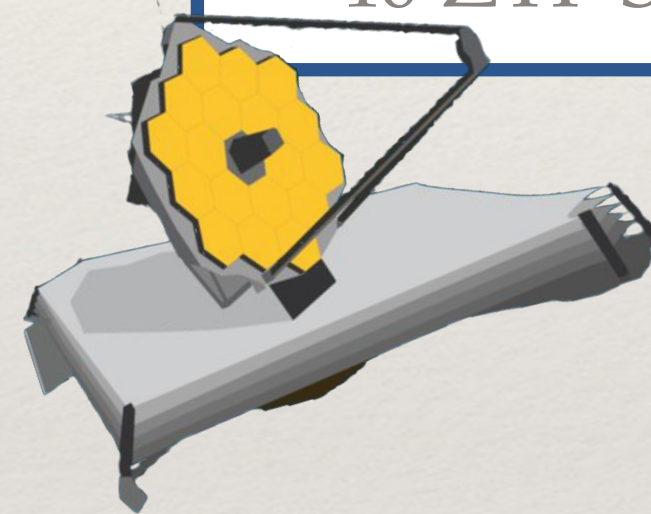
Pure Volume limited

Get an independent measurement of L_{SN}

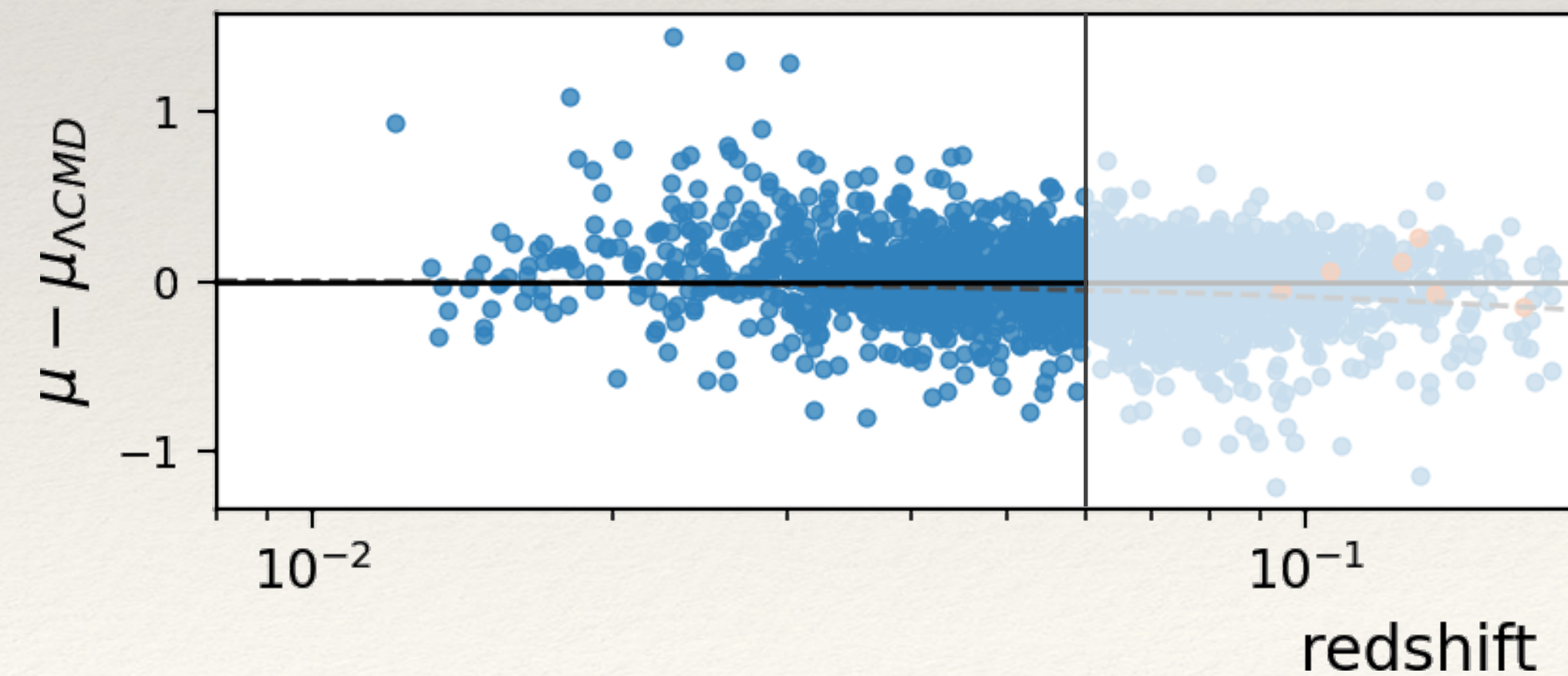
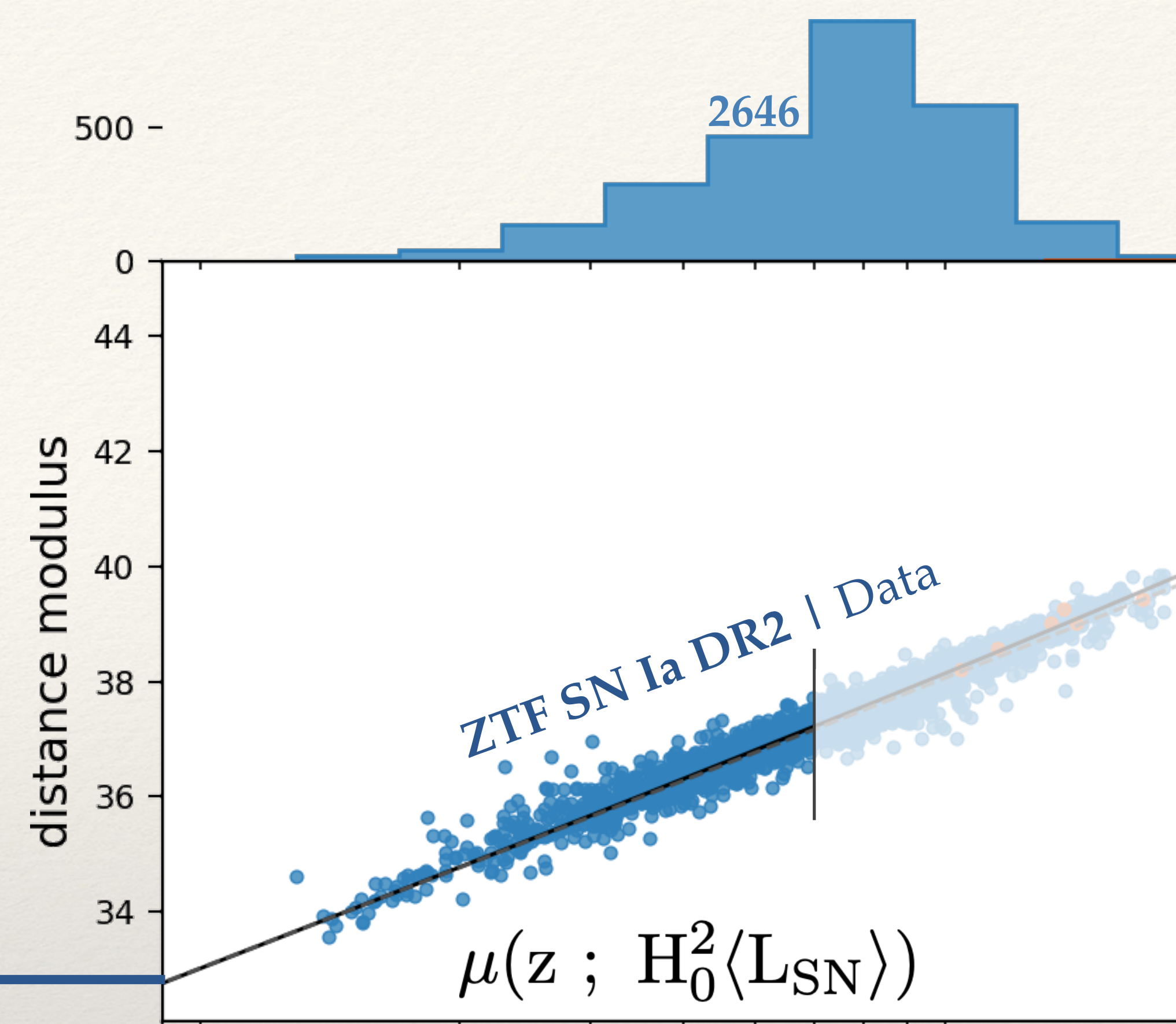
No selection effect

Self-consistent calibration

Measure L_{SN}
<i>Volume limited</i> ZTF SNeIa < 60 Mpc
Tip Red Giant Branch (doable in any galaxy)
~40 ZTF SNe Ia <i>need JWST</i>



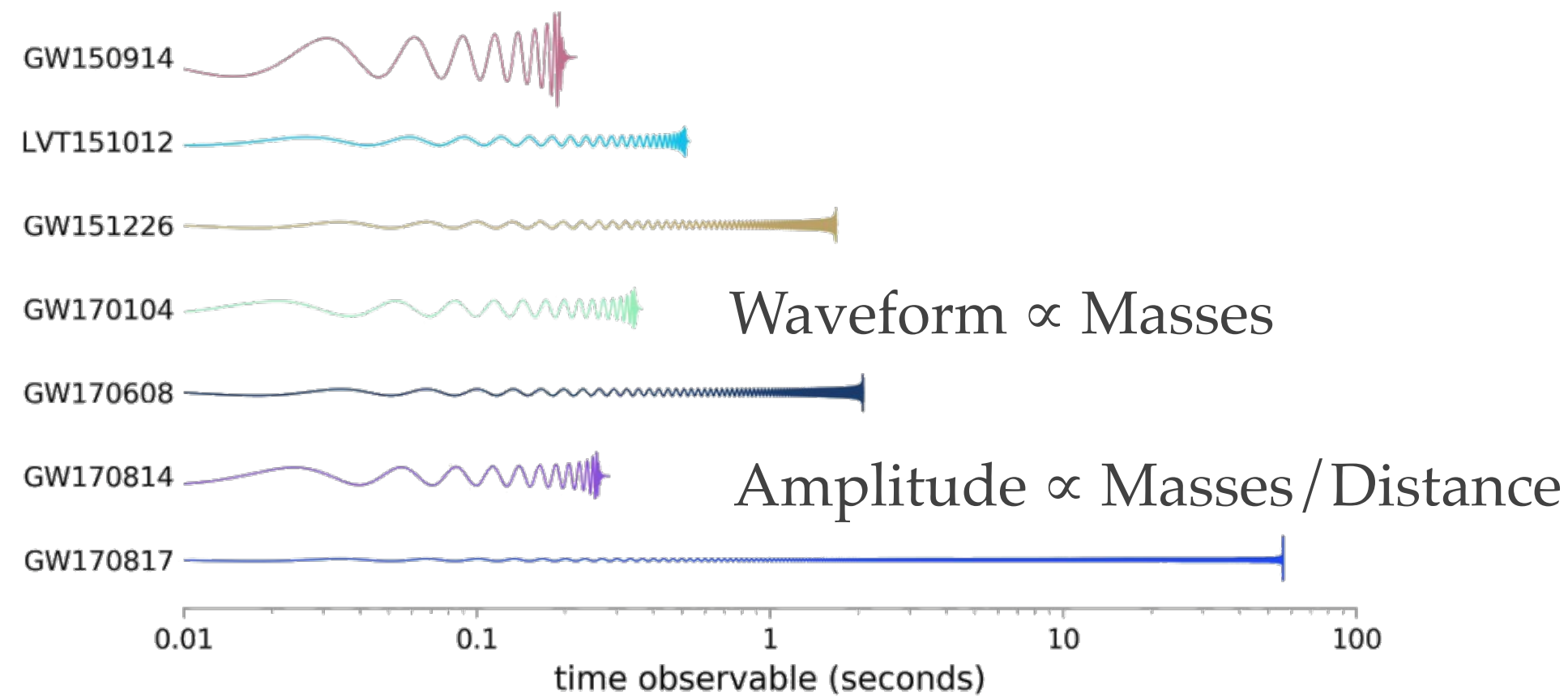
Get H_0
<i>Volume limited</i> ZTF SNeIa $z < 0.06$
$O(1000)$ ZTF SNe Ia <i>ready</i>



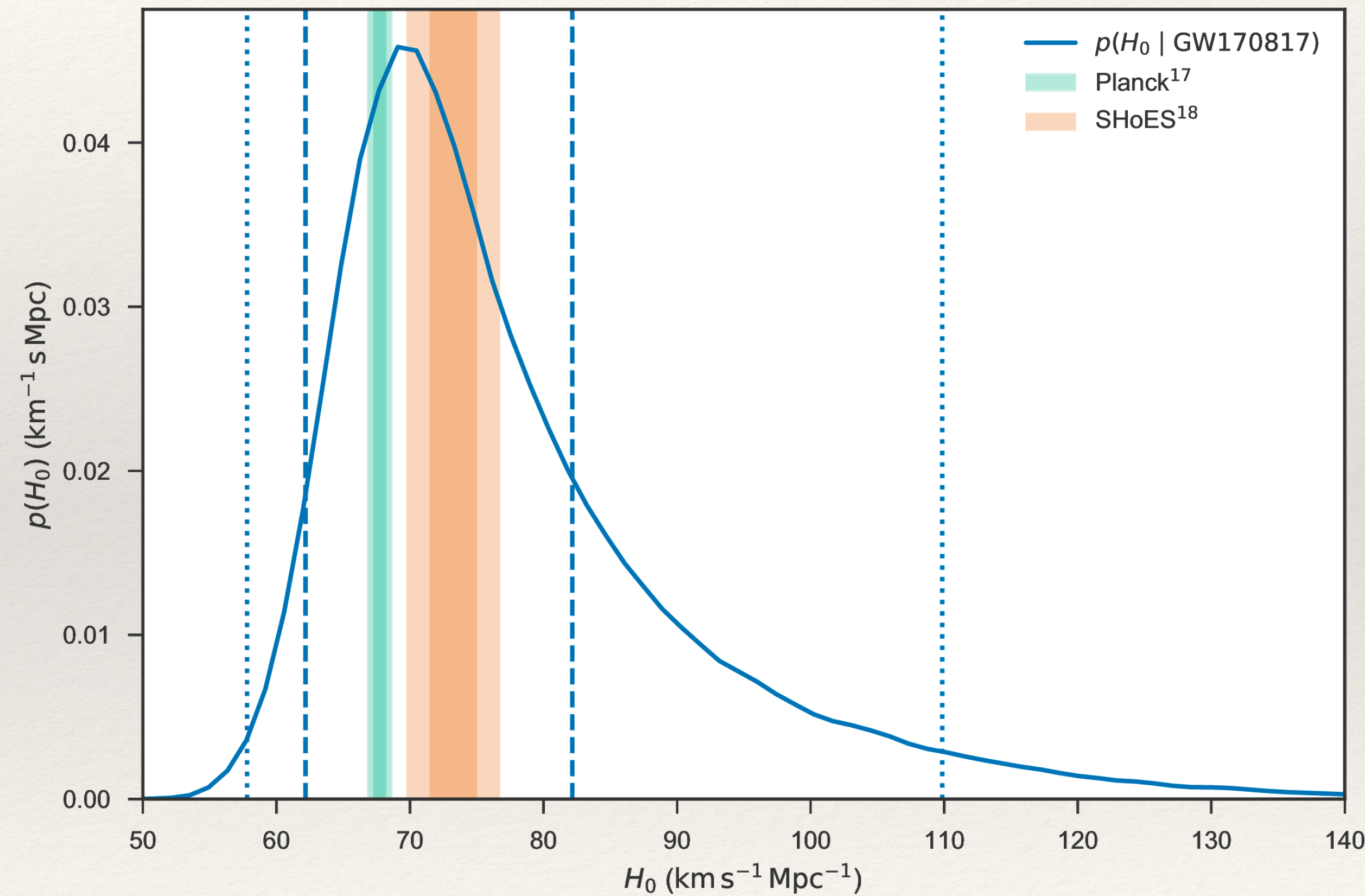
Gravitational Waves & ElectroMagnetism | H_0

Gravitational Waves

d_L

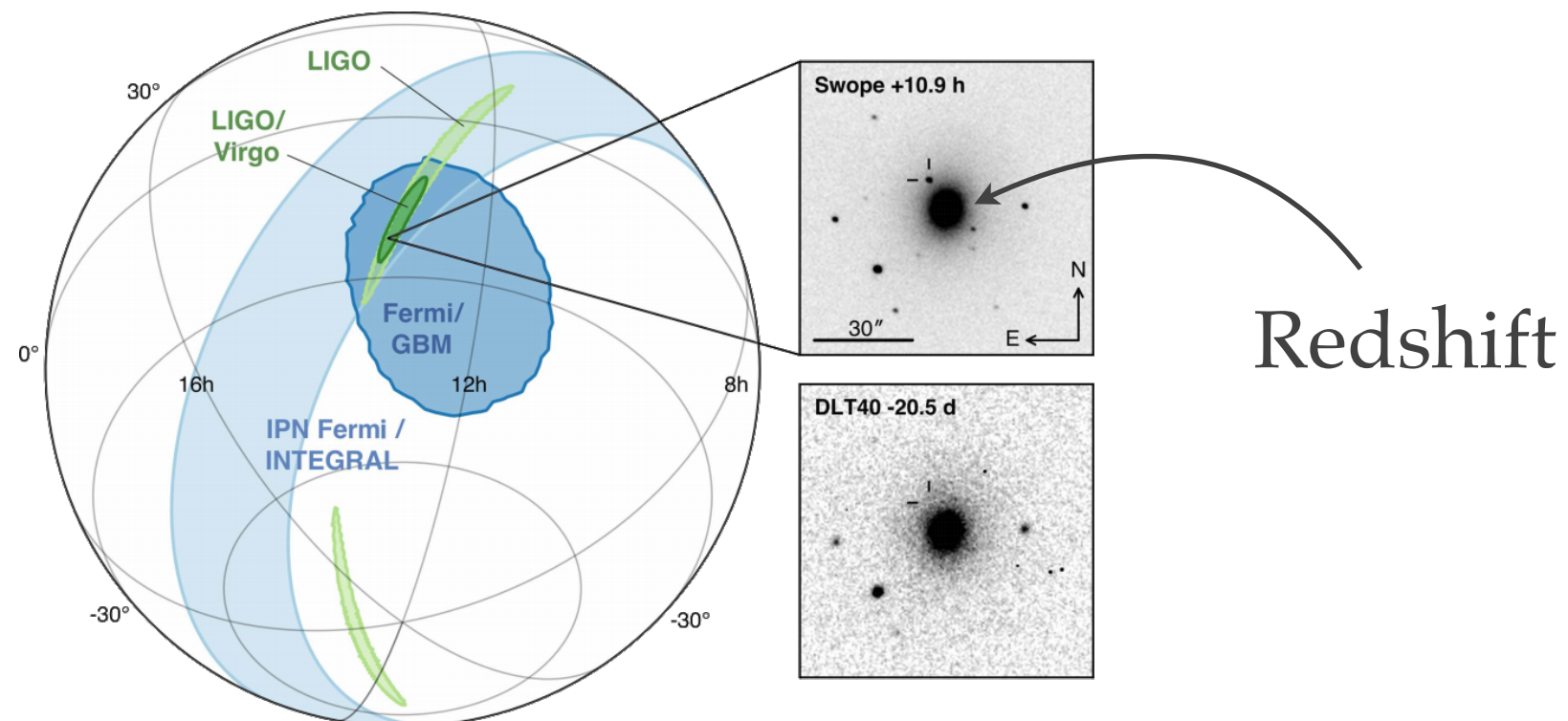


$$v_H = H_0 d_L$$



Optical Counterpart

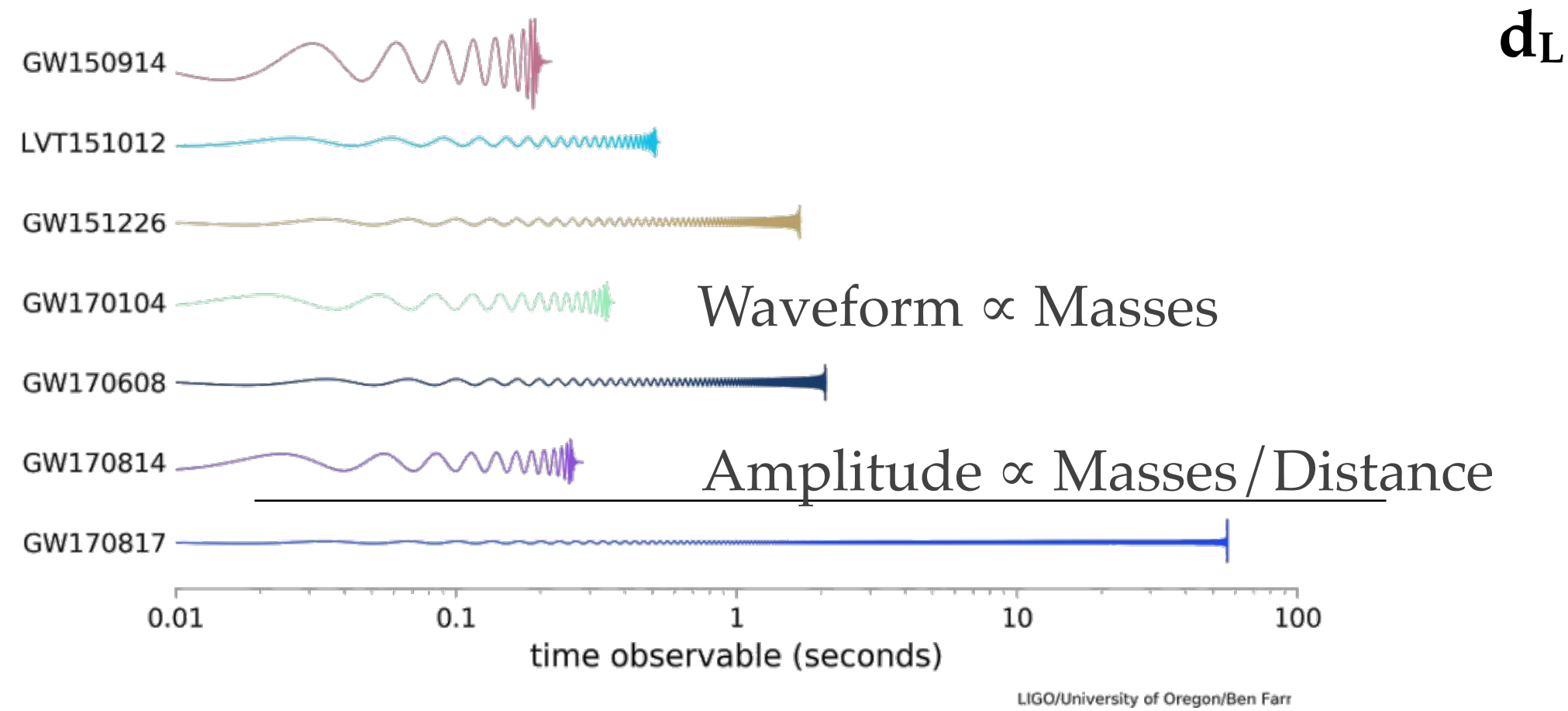
z



Gravitational Waves & ElectroMagnetism | H_0

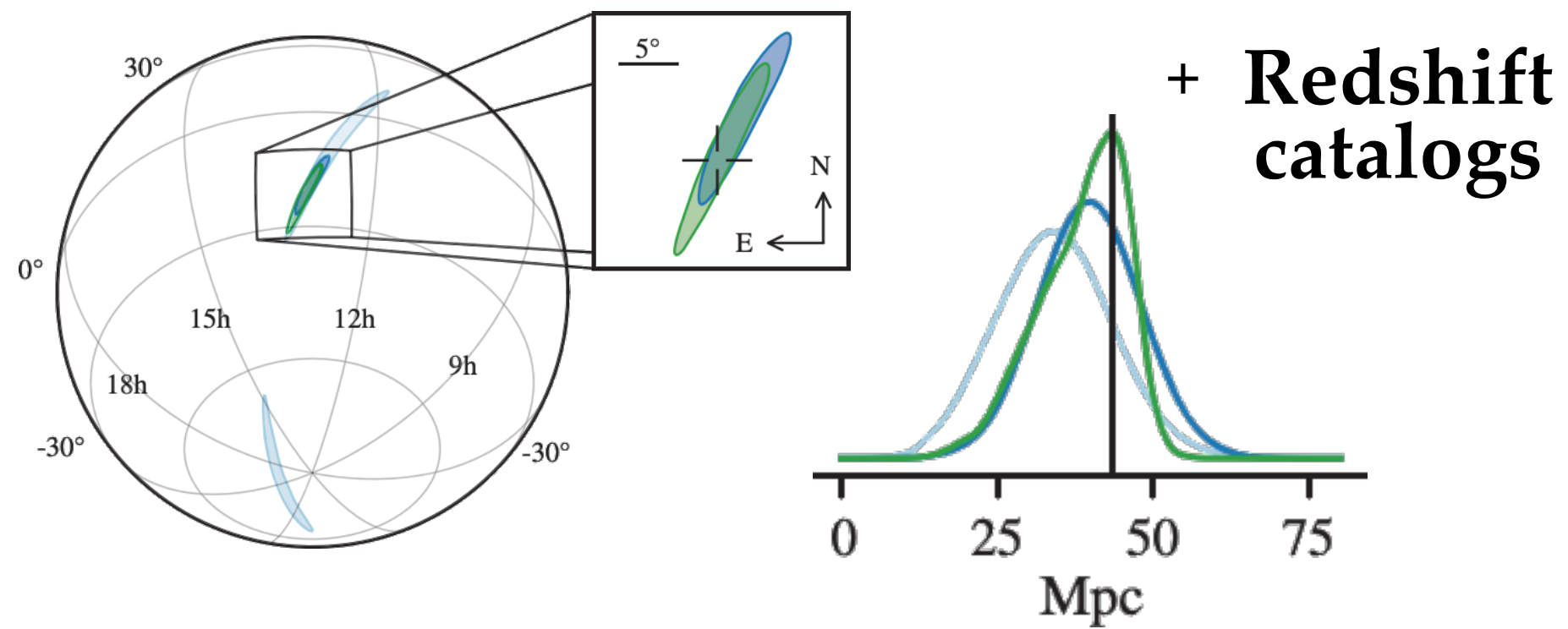
Works with any Merger

Gravitational Waves

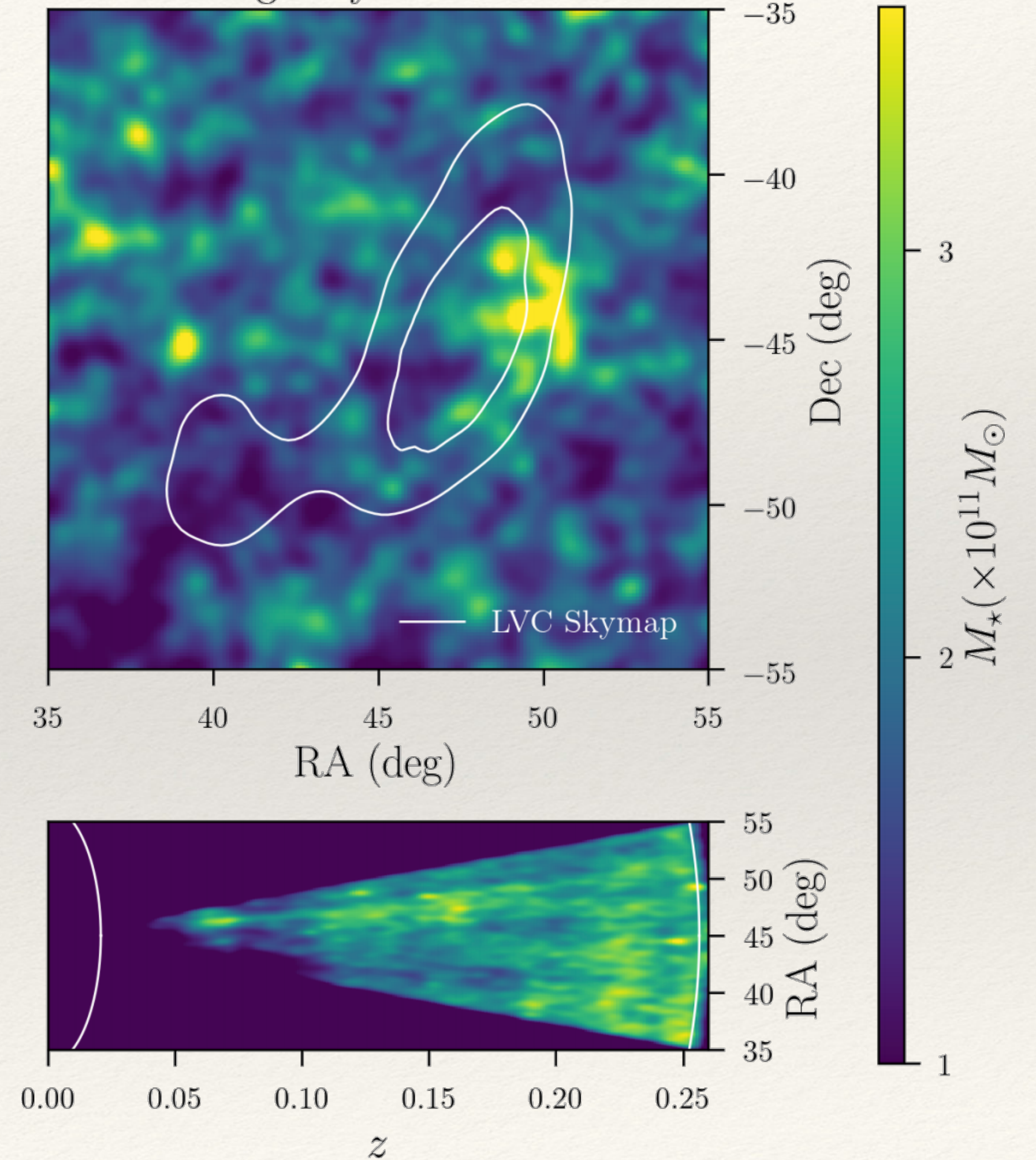


Optical Counterpart

— Statistical —

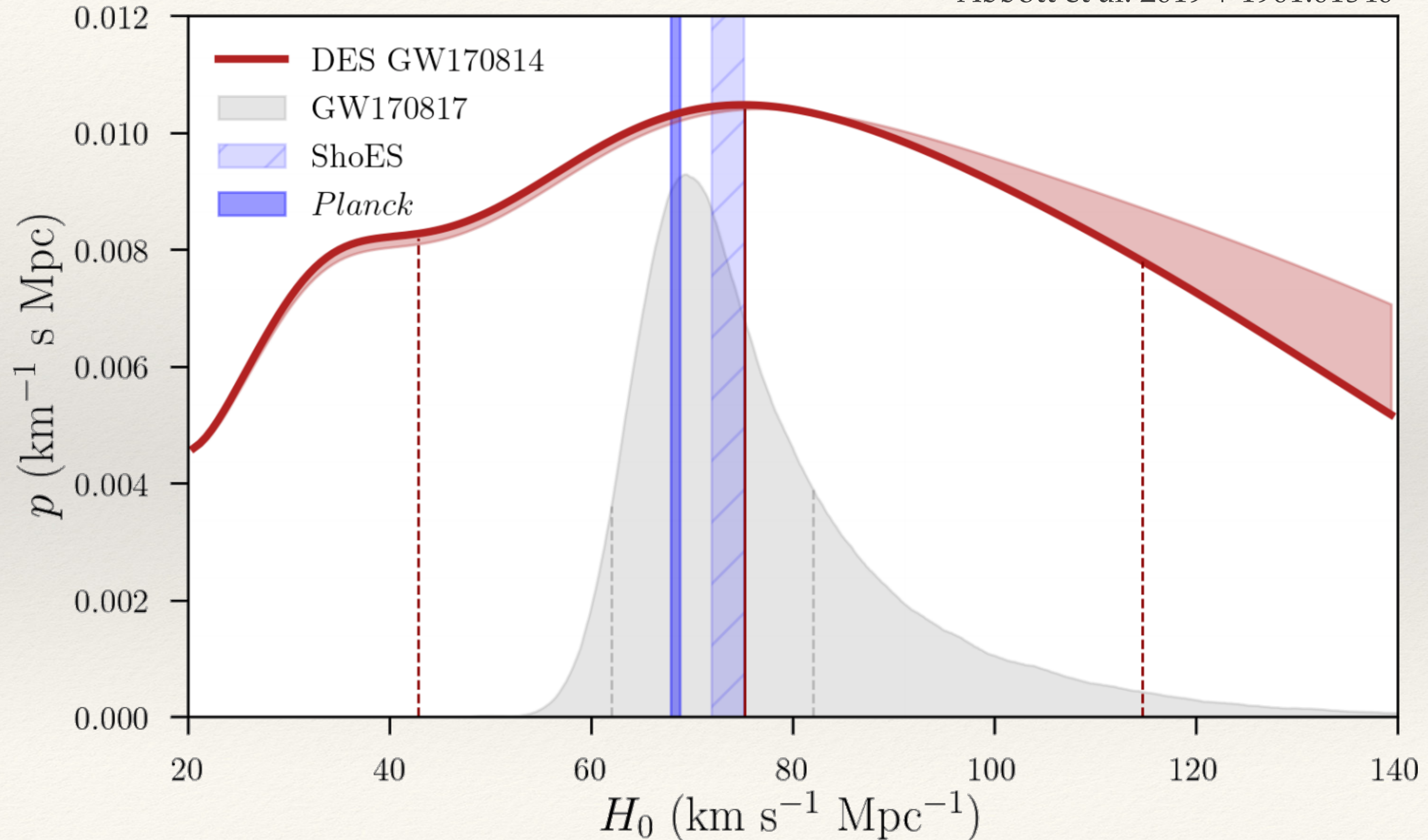


DES galaxy distribution

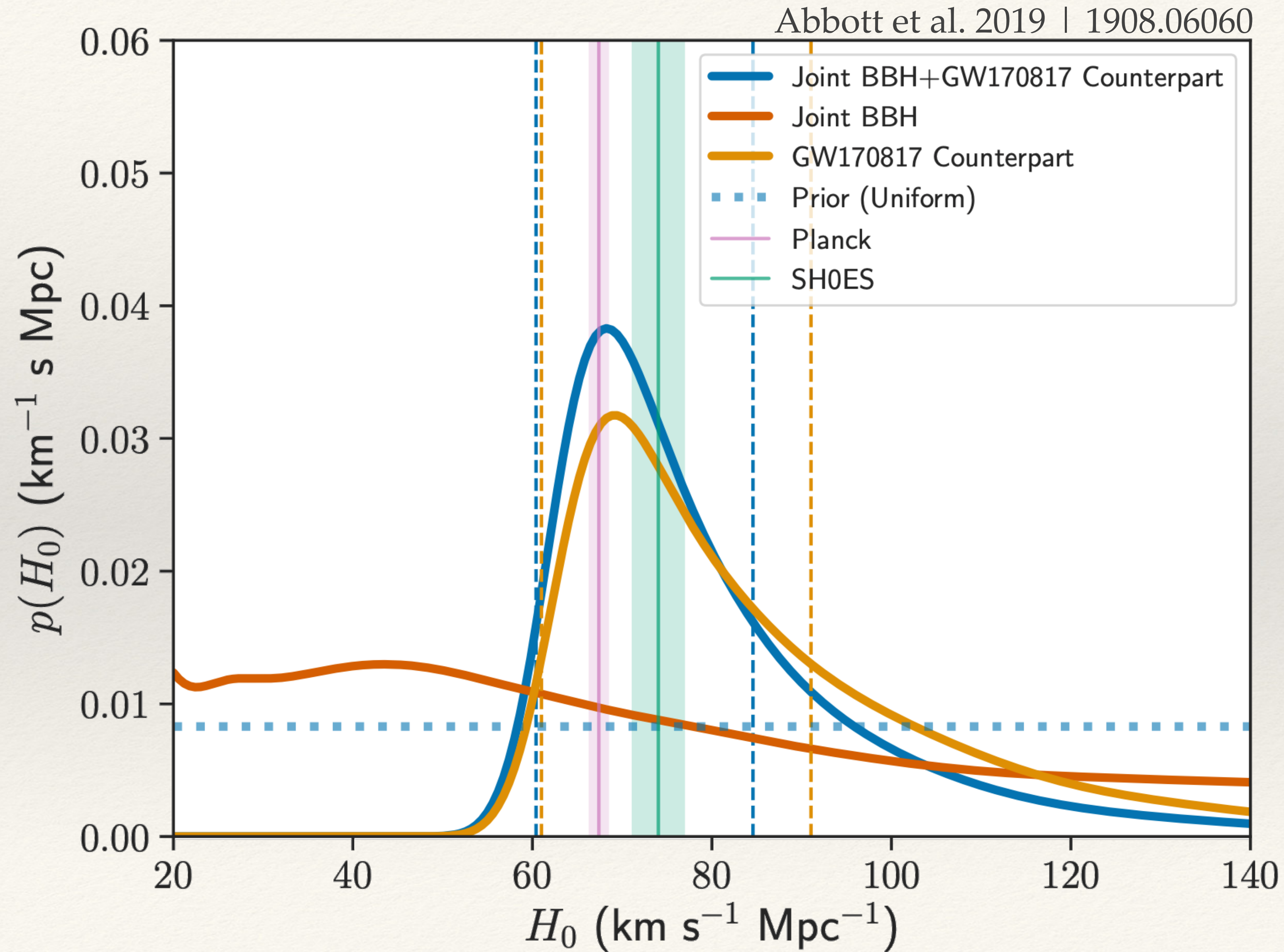


Direct measurement of H_0 | *without counterpart*

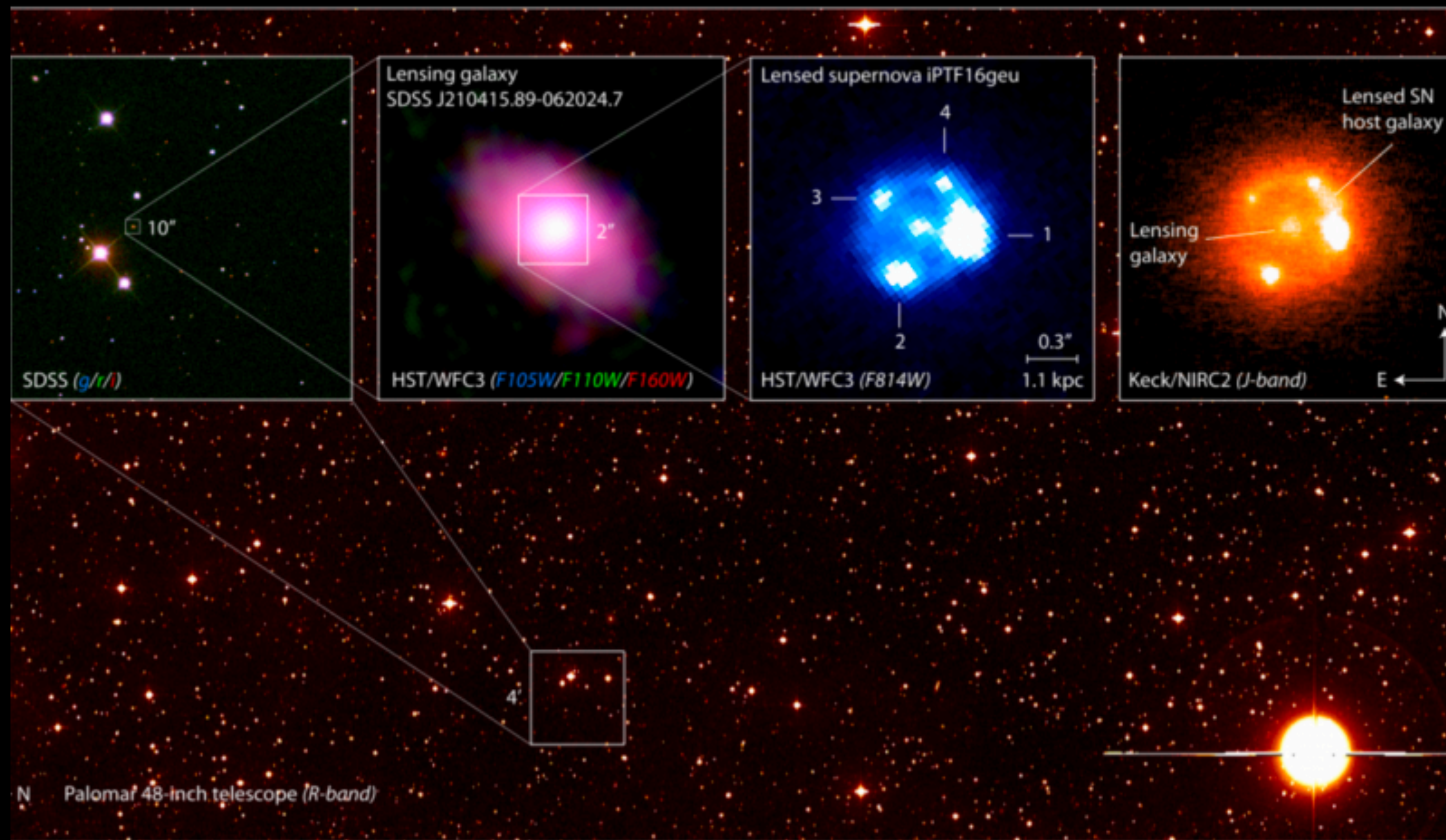
Abbott et al. 2019 | 1901.01540



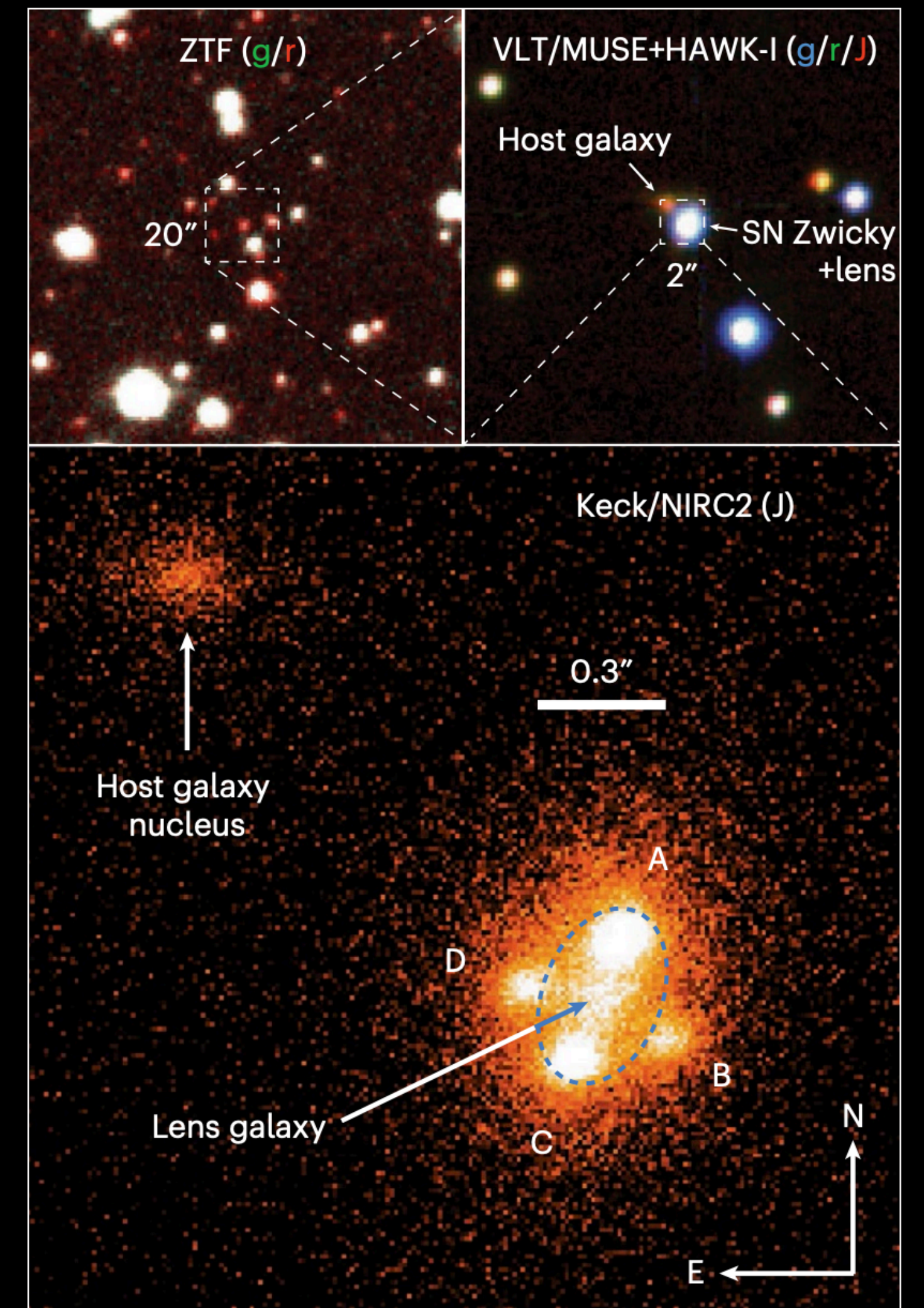
Direct measurement of H_0 | *without counterpart*



iPTF16geu | Goobar et al. 2017

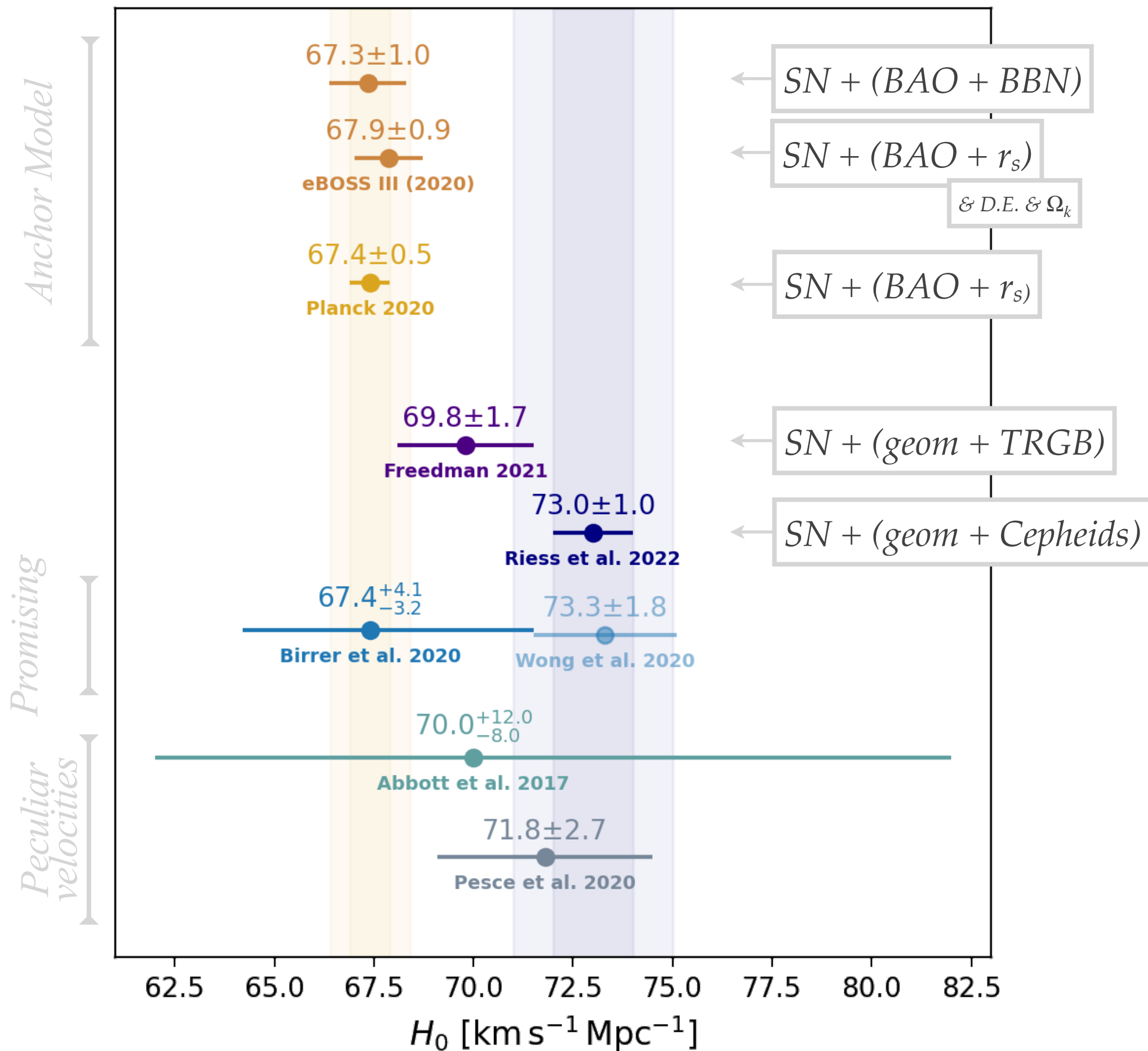


SN Zwicky | Goobar et al. 2023



Conclusion | *Hubble-Lemaître Constant*

H_0 Tension



— New fundamental Physics

No simple solutions so far
These solving H_0 break σ_8

— Type Ia Supernovae a key for H_0

Understanding their systematics is of paramount importance for cosmology
ZTF is about to change the game

— Systematic Uncertainties

Must be multiple sources
e.g. : age-bias for SNe Ia
& lensing modeling for strong lensing