Cosmological simulations of "Milky Way-like" galaxies

Overview

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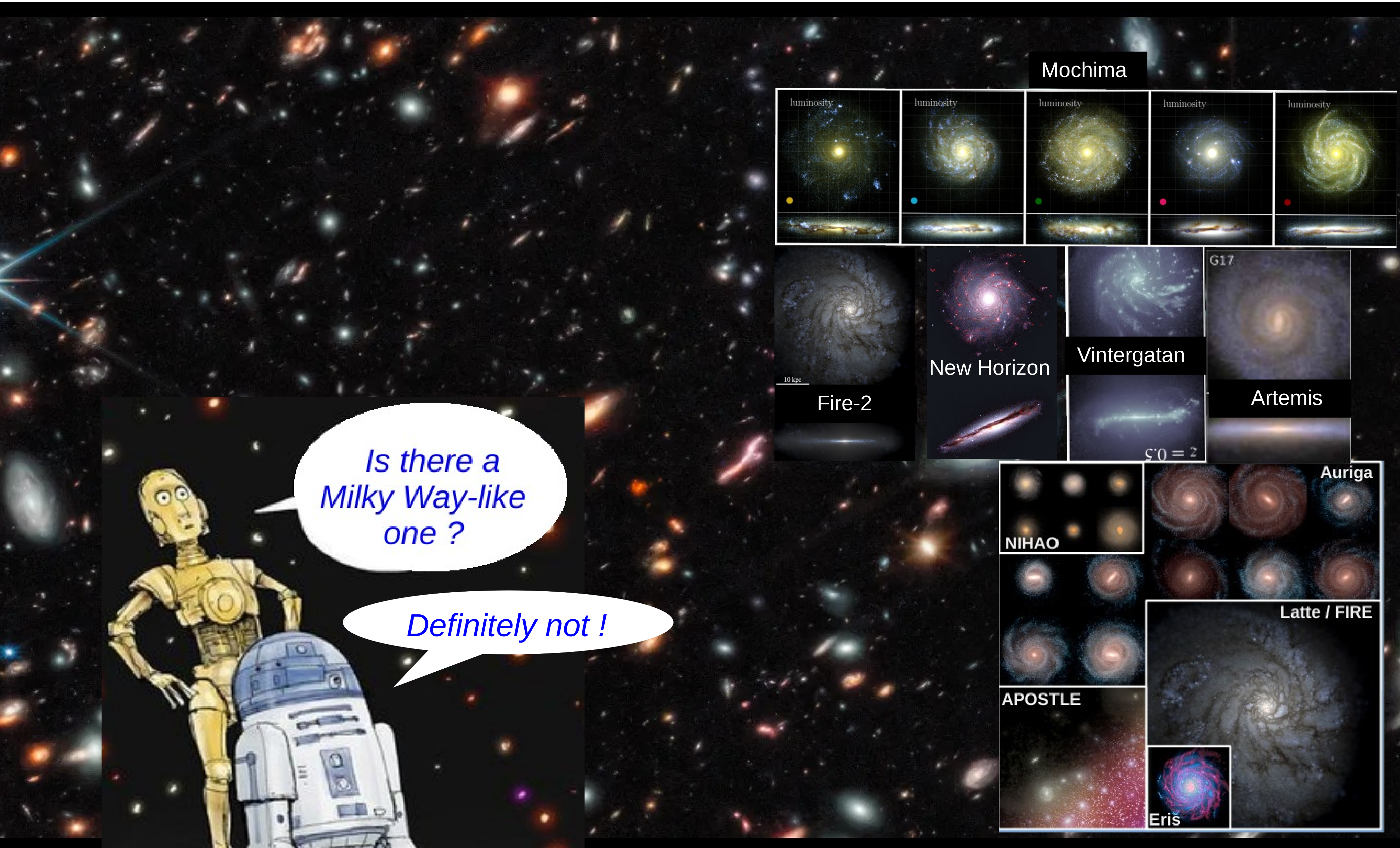
Cosmological simulations of "Milky Way-like" galaxies

Spoiler !



Outline

- *Introduction*

  - Principles

  - Baryonic physics

- *Cosmological Zoom-in of spiral galaxies / Milky Way “analogs”*

  - *- Dark matter only simulations*
  - *- Hydro simulations*
  - *- Selected results on*

    Galaxies

    Dark matter

    Beyond CDM

- *Summary - Conclusion*

Some references :

- **Cosmological Simulations of Galaxy Formation**

  Mark Vogelsberger[1], Federico Marinacci[2], Paul Torrey[3], and Ewald Puchwein[4]

  arXiv:1909.07976

- **Theoretical Challenges in Galaxy Formation**

  Thorsten Naab[1] & Jeremiah P. Ostriker[2,3]

  arXiv:1612.06891

- *GISM 2021 Florent Renaud*

  https://ismgalaxies2021.sciencesconf.org/

GRAVITY: Dark matter (+Stars)

Modeling dark matter

collisionless Boltzmann equation: $\frac{\mathrm{d}f}{\mathrm{d}t} = \frac{\partial f}{\partial t} + \mathbf{v}\frac{\partial f}{\partial \mathbf{r}} - \frac{\partial \Phi}{\partial \mathbf{r}}\frac{\partial f}{\partial \mathbf{v}} = 0$ **Poisson's equation:** $\nabla^2\Phi = 4\pi G \int f \mathrm{d}\mathbf{v}$

The collisionless Boltzmann equation describes the evolution of the phase-space density or distribution function of dark matter, $f = f(\mathbf{r}, \mathbf{v}, t)$, under the influence of the collective gravitational potential, Φ, given by Poisson's equation. The collisionless Boltzmann equation states the conservation of the local phase-space density; i.e. Liouville's theorem.

HYDRO: Gas

Modeling cosmic gas

Eulerian formulation:

$$\frac{\partial \rho}{\partial t} + \nabla\cdot(\rho\mathbf{v}) = 0$$
$$\frac{\partial \rho\mathbf{v}}{\partial t} + \nabla\cdot(\rho\mathbf{v}\otimes\mathbf{v} + P\mathbb{1}) = 0$$
$$\frac{\partial \rho e}{\partial t} + \nabla\cdot(\rho e + P)\mathbf{v} = 0$$

Lagrangian formulation:

$$\frac{\mathrm{D}\rho}{\mathrm{D}t} = -\rho\nabla\cdot\mathbf{v}$$
$$\frac{\mathrm{D}\mathbf{v}}{\mathrm{D}t} = -\frac{1}{\rho}\nabla P$$
$$\frac{\mathrm{D}e}{\mathrm{D}t} = -\frac{1}{\rho}\nabla\cdot P\mathbf{v}$$

Arbitrary Lagrangian-Eulerian formulation:

$$\frac{\mathrm{d}}{\mathrm{d}t}\int_V \rho \mathrm{d}V = -\int_S \rho(\mathbf{v}-\mathbf{w})\cdot\mathbf{n}\mathrm{d}S$$
$$\frac{\mathrm{d}}{\mathrm{d}t}\int_V \rho\mathbf{v} \mathrm{d}V = -\int_S \rho\mathbf{v}(\mathbf{v}-\mathbf{w})\cdot\mathbf{n}\mathrm{d}S - \int_S P\mathbf{n}\mathrm{d}S$$
$$\frac{\mathrm{d}}{\mathrm{d}t}\int_V \rho e \mathrm{d}V = -\int_S \rho e(\mathbf{v}-\mathbf{w})\cdot\mathbf{n}\mathrm{d}S - \int_S P\mathbf{v}\cdot\mathbf{n}\mathrm{d}S$$

Different forms of the hydrodynamical equations. $\mathrm{D}/\mathrm{d}t \equiv \partial/\partial t + \mathbf{v}\cdot\nabla$ denotes the Lagrangian derivative and $e = u + \mathbf{v}^2/2$ the total energy per unit mass. The equations are closed through $P = (\gamma-1)\rho u$ with $\gamma = 5/3$. For the arbitrary Lagrangian-Eulerian formulation the grid moves with velocity $\mathbf{w}$ and cell volumes evolve as $\mathrm{d}V/\mathrm{d}t = \int_V(\nabla\cdot\mathbf{w})\mathrm{d}V$.

arXiv:1909.07976

Codes

Table 1: Major galaxy formation simulation codes

code name	gravity treatment[a]	hydrodynamics treatment[b]	parallelization technique[c]	code availability[d]	primary reference
ART	PM/ML	AMR	data-based	public	Kravtsov (1997)[27]
RAMSES	PM/ML	AMR	data-based	public	Teyssier (2002)[38]
GADGET-2/3	TreePM	SPH	data-based	public	Springel (2005)[39]
Arepo	TreePM	MMFV	data-based	public	Springel (2010)[40]
Enzo	PM/MG	AMR	data-based	public	Bryan et al. (2014)[41]
ChaNGa[e]	Tree/FM	SPH	task-based	public	Menon et al. (2015)[42–44]
GIZMO[f]	TreePM	MLFM/MLFV	data-based	public	Hopkins et al. (2015)[45]
HACC	TreePM/P^3M	CRK-SPH	data-based	private	Habib et al. (2016)[46]
PKDGRAV3	Tree/FM	–	data-based	public	Potter et al. (2017)[47]
Gasoline2	Tree	SPH	task-based	public	Wadsley et al. (2017)[48]
SWIFT	TreePM/FM	SPH	task-based	public	Schaller et al. (2018)[49]

[a] PM: particle-mesh; TreePM: tree + PM, FM: fast multipole, P^3M: particle-particle-particle-mesh; ML: multilevel; MG: multigrid

[b] SPH: smoothed particle hydrodynamics, CRK-SPH: conservative reproducing kernel smoothed particle hydrodynamics , AMR: adaptive-mesh-refinement, MMFV: moving-mesh finite volume, MLFM/MLFV: mesh-free finite mass / finite volume

[c] data-based: data parallelism focuses on distributing data across different nodes, which operate on the data in parallel; task-based: task parallelism focuses on distributing tasks concurrently performed

[d] private: private code; public: publicly available code (in some cases with limited functionality)

[e] gravity solver is based on PKDGRAV3

[f] based on the GADGET-3 code

arXiv:1909.07976

Baryonic physics

Baryonic physics

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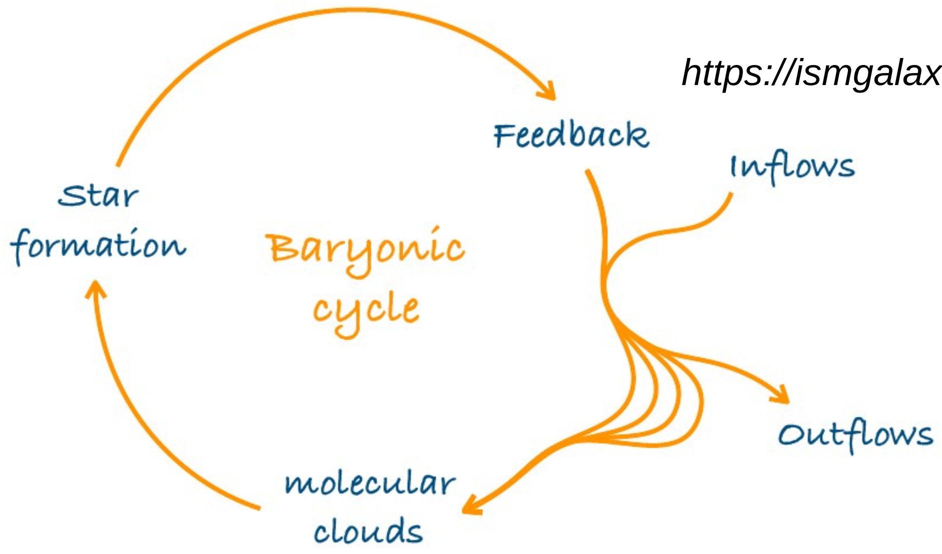
https://ismgalaxies2021.sciencesconf.org/



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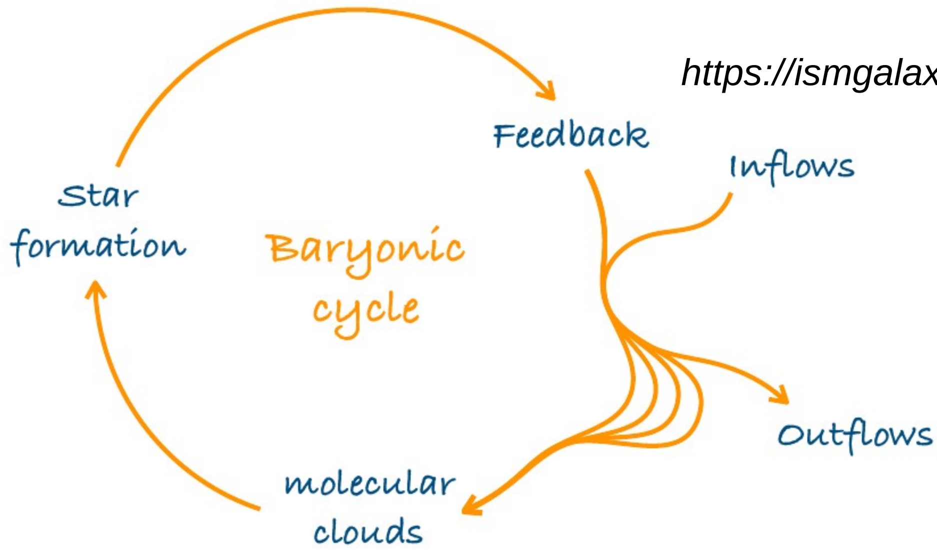


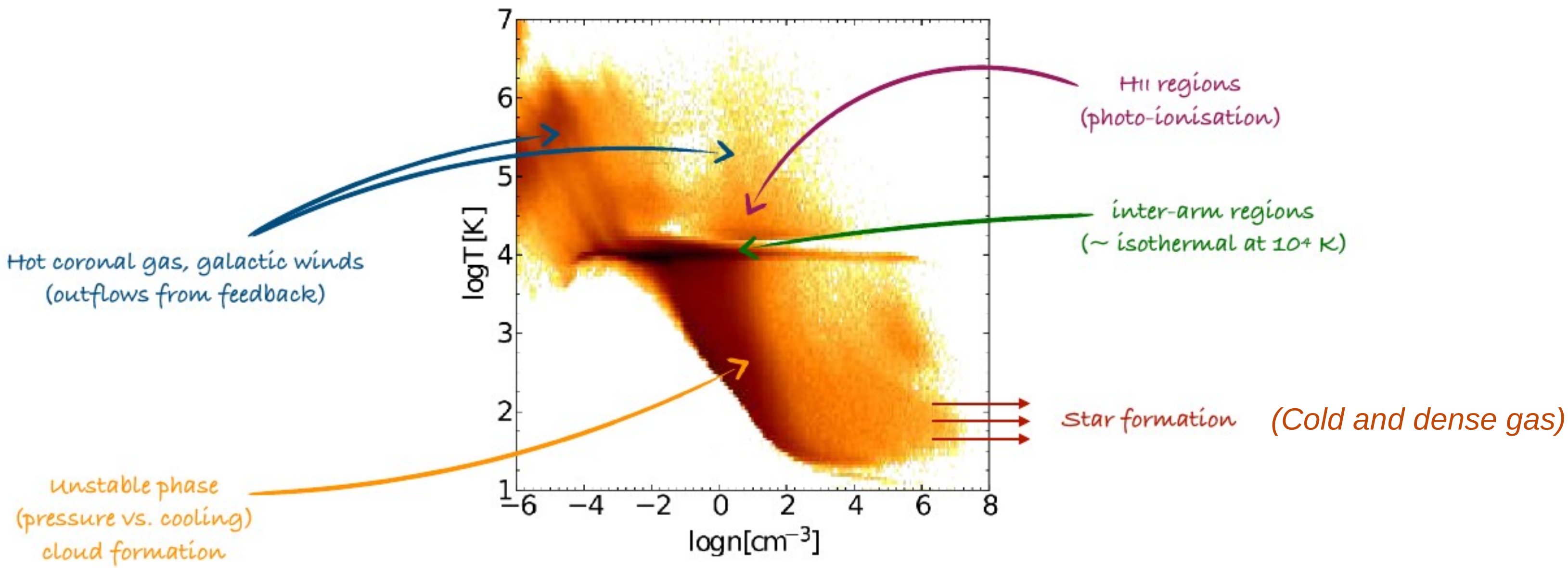
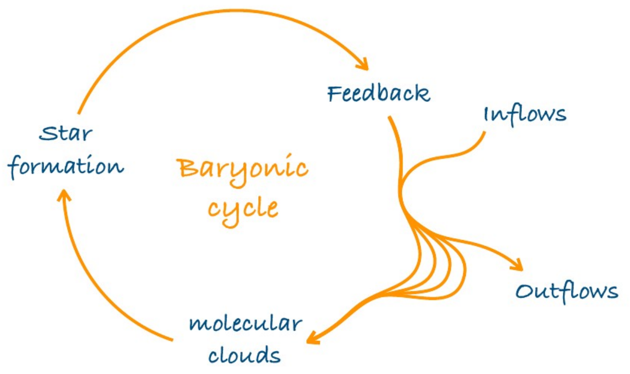


Fig: phase diagram
(adapted from Marinacci et al. 2019)

Baryonic physics

- *Sub resolution effective modeling/recipes*
- *Calibration, parameters, resolution dependent*



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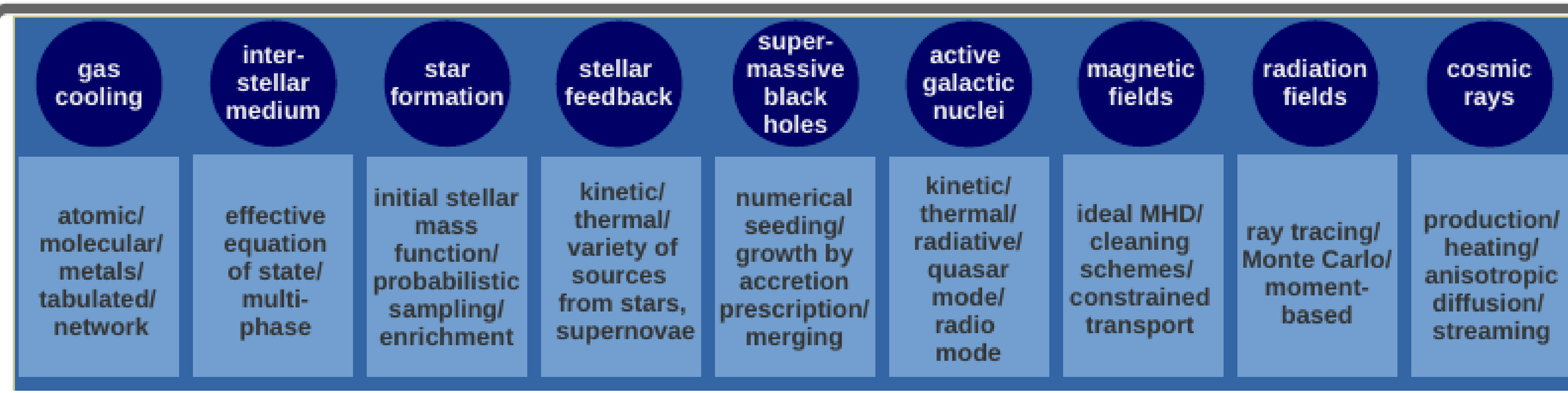

arXiv:1909.07976

(some) Relevant baryonic physics processes/models (for MW-size galaxies)

- *Star formation:*

  ISM conditions, physical state of gas: cold dense gas

  - Kennicut-Schmidt law (Shmidt 1959), star formation above a density threshold, constant efficiency (~1%)

  - Multifree-fall: efficiency = function of gas properties (density, turbulence …)

  (Federath & Klessen 2012, Padoan & Nordlund 2011, Henebelle & Chabrier 2011)

  Q: Universal IMF ? Impact of spirals, bars ? Environment ? Interaction and mergers ? Turbulence description Redshift dependence ? Multi-scale and multi-physics topic.

- *Stellar feedback*

  Death of heavy stars

  release energy and momentum (thermally, kinematically)

  - Delayed Cooling: stop cooling (Teyssier et 2013, Dubois et al 2015)

  - Mechanical FB: mimic Sedov blast phases (Kimm & Cen 2014)

  Q: Coupling to galactic scale ? Drift of stars ? Expansion and volume of SN bubbles ?

- *AGN feedback ?*

  - BH growth $\propto$ Bondi accretion (and < Eddington rate)

  - AGN released power $\propto$ BH growth : quasar thermal and radio jet modes

  (Dubois et al 2014)

  Q: Centering of BH ? Eddington limit ?

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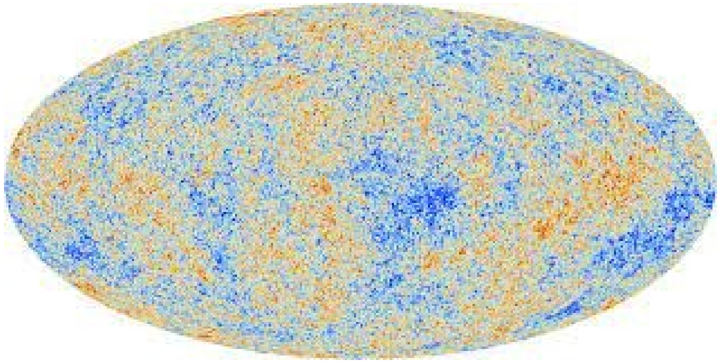
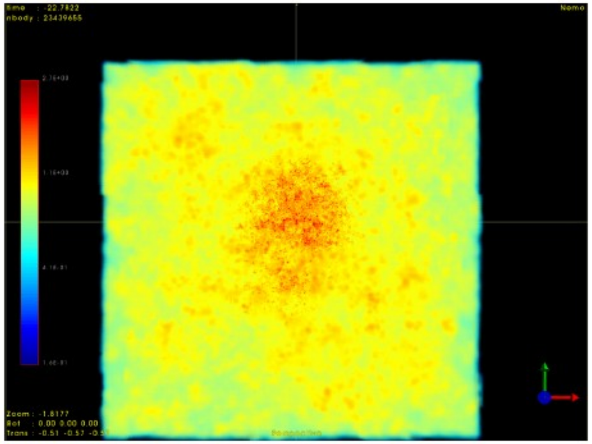
Effective models, parameters, calibration, resolution in (cosmological) simulations ?

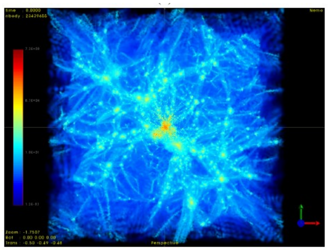
Cosmological simulations

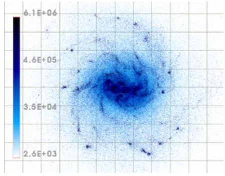
CMB → Initial conditions

Zel'Dovich 1970, Linear Perturbation Theory
Bertchinger 2001
Hahn,Abel 2011
MUSIC, MONOPHONIC









This talk

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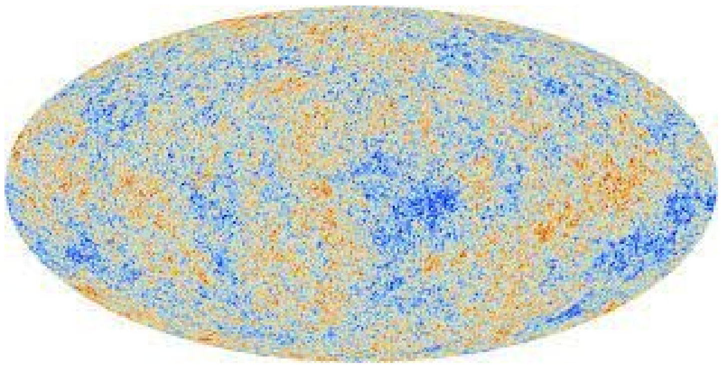
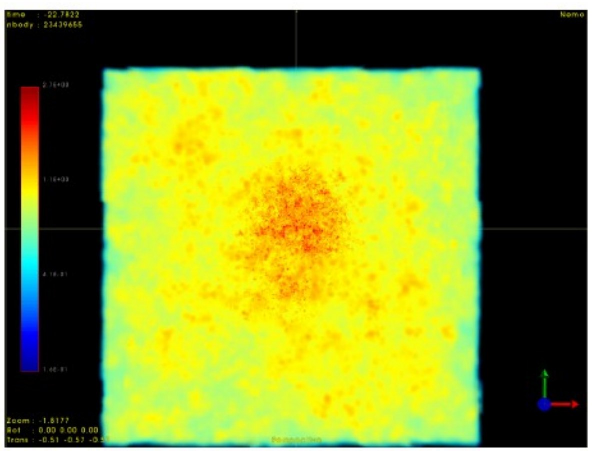
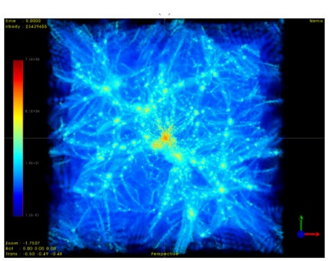
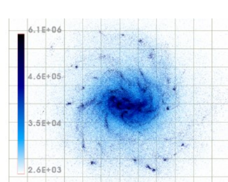




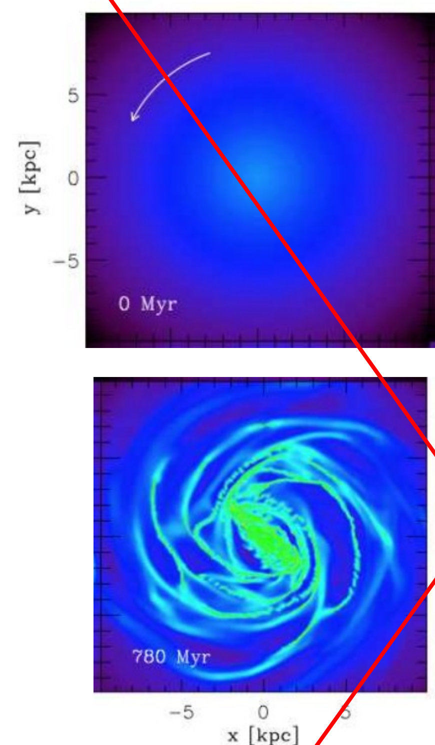


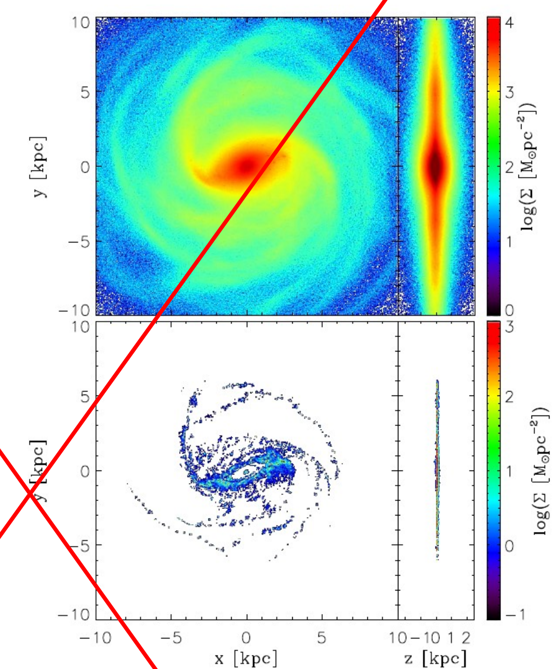


This talk

Isolated/Idealized simulations

Choose/tune Initial conditions





arXiv:1307.5639

Renaud et al 2013, 2021, Grisdale 2017 …

Size~ 100 kpc
Res ~ 0.1-10 pc

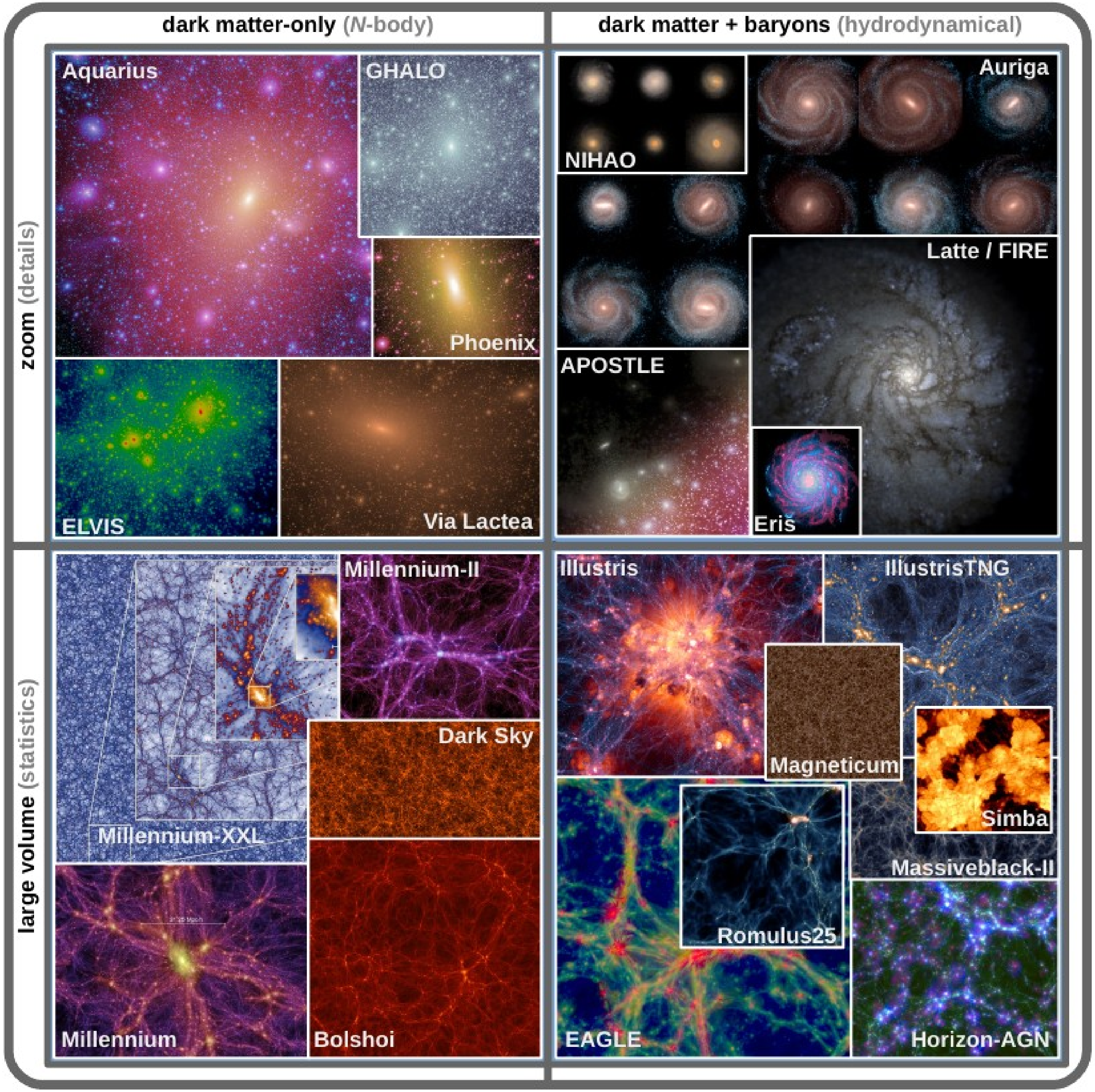
+ Control parameters
+ resolution

- Environment (mergers, gas in/outflows)
- Initial conditions

Not in this talk

Cosmological Simulations of Galaxy Formation

Mark Vogelsberger[1], Federico Marinacci[2], Paul Torrey[3], and Ewald Puchwein[4]



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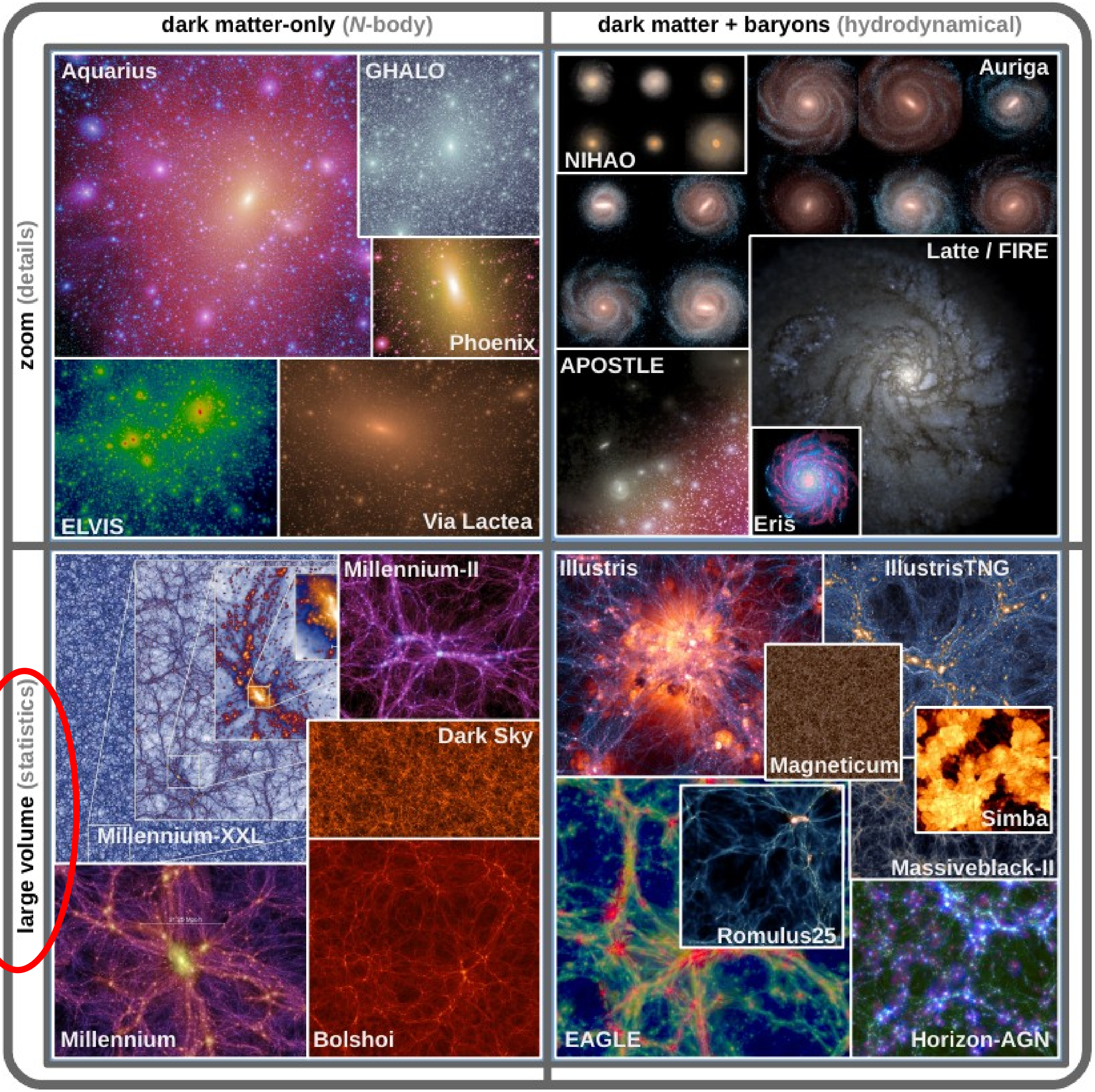

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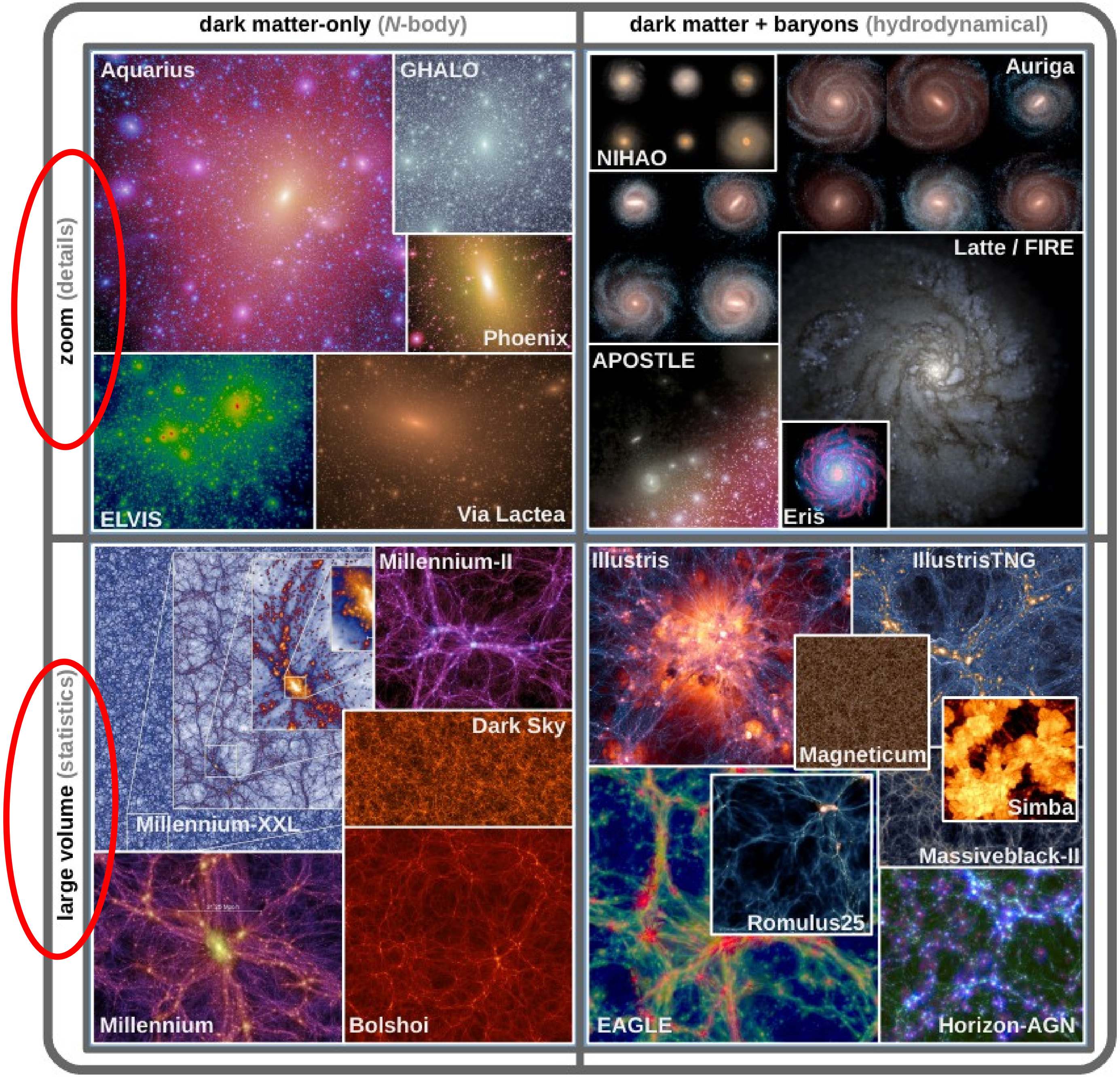
This talk

Focusing on MW-size haloes



arXiv:1909.07976

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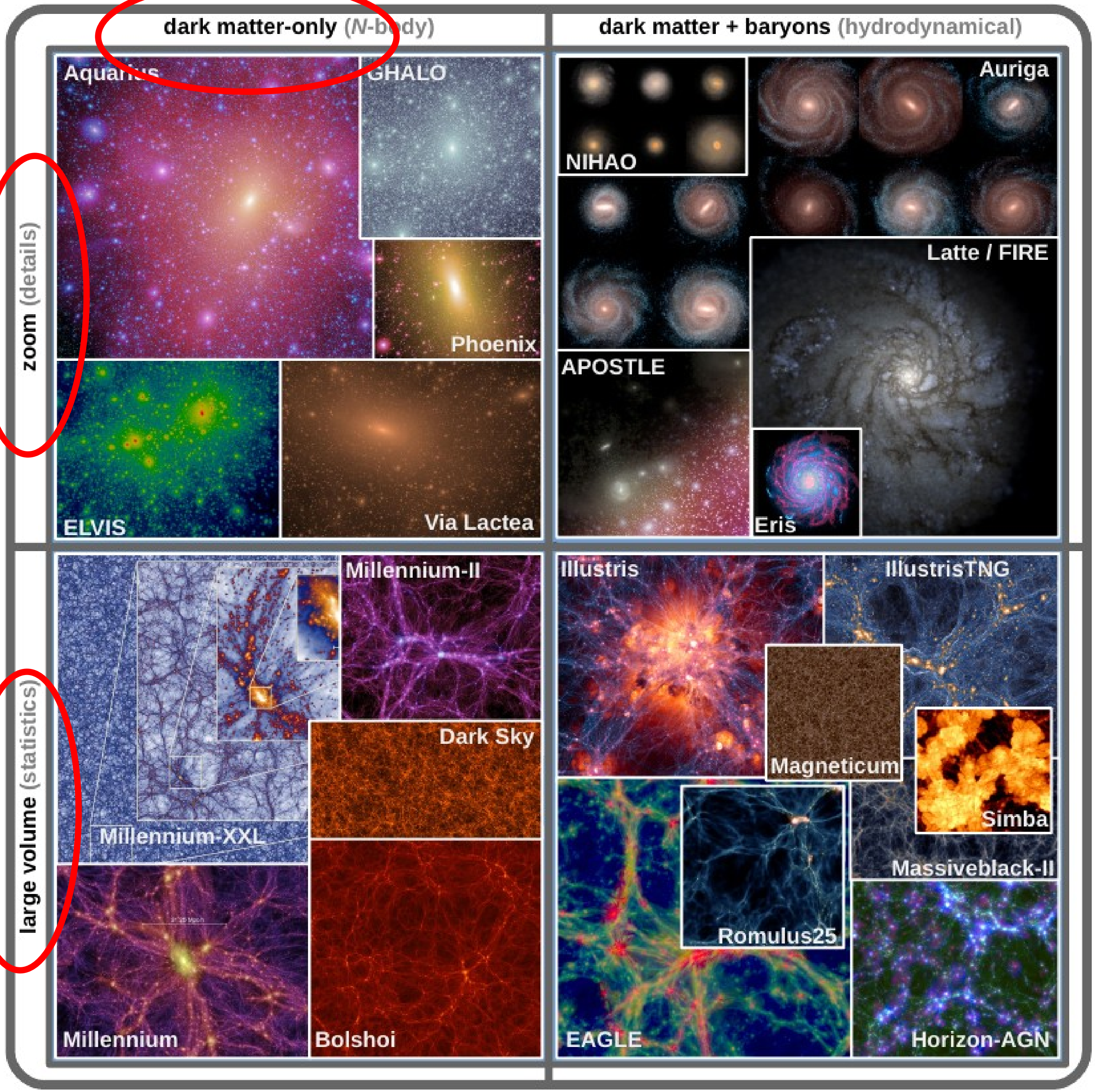

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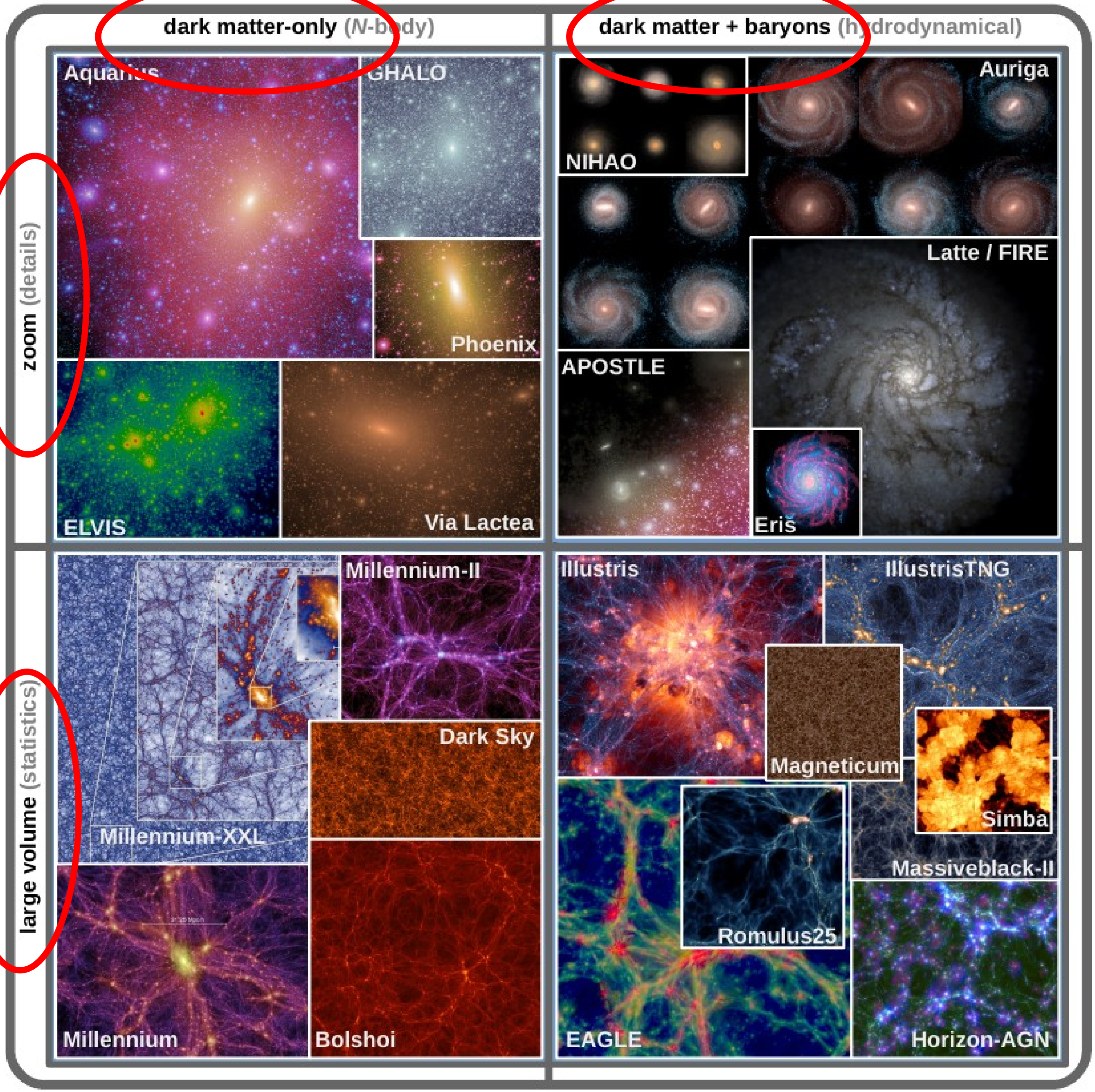


arXiv:1909.07976

Big volume simulations

~ 1 Gpc – (100)s Mpc boxes

Big volume simulations

Dark matter only (DMO)

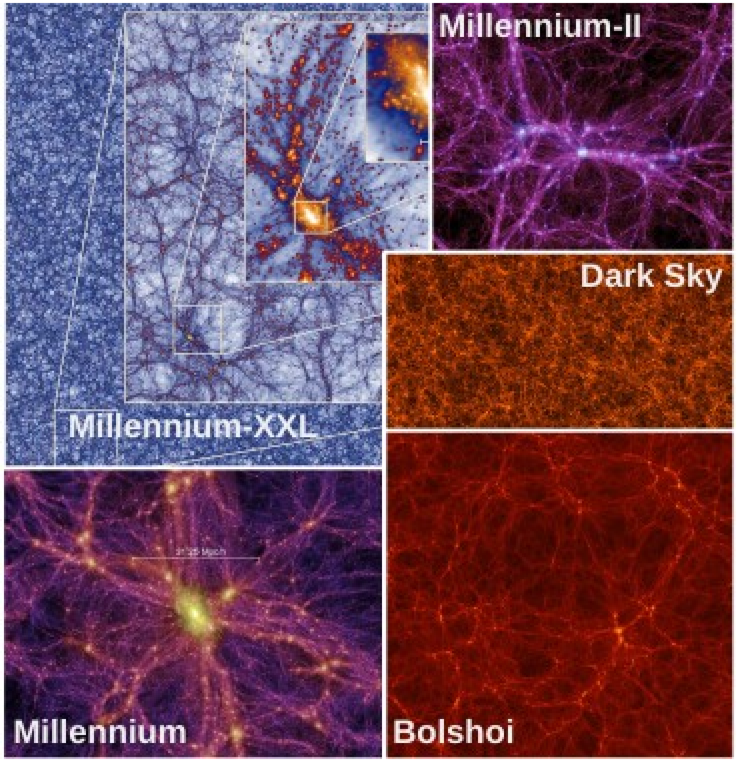
- Cosmic web (filaments, voids, halos ...)

- Large scale structure (matter distribution)

- Halo mass function

- Cosmological scenario

....



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Big volume simulations

Dark matter only (DMO)

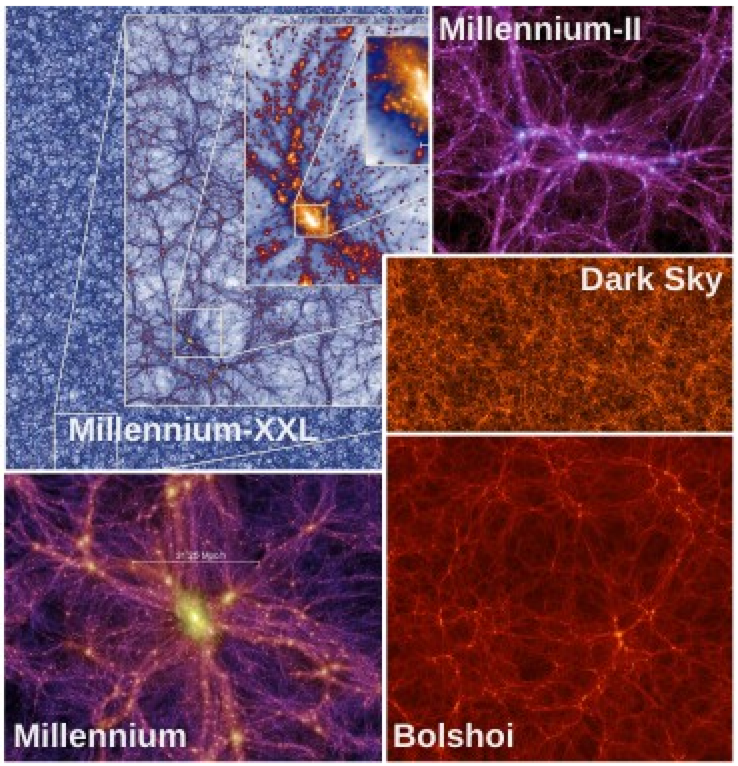
- *Cosmic web (filaments, voids, halos ...)*
- *Large scale structure (matter distribution)*
- *Halo mass function*
- *Cosmological scenario*

....

Hydrodynamical

- *Galaxy population*
- *Stellar(-to-halo) mass function*
- *Gas around galaxies*
- *Clustering*
- *Scaling relations*

....



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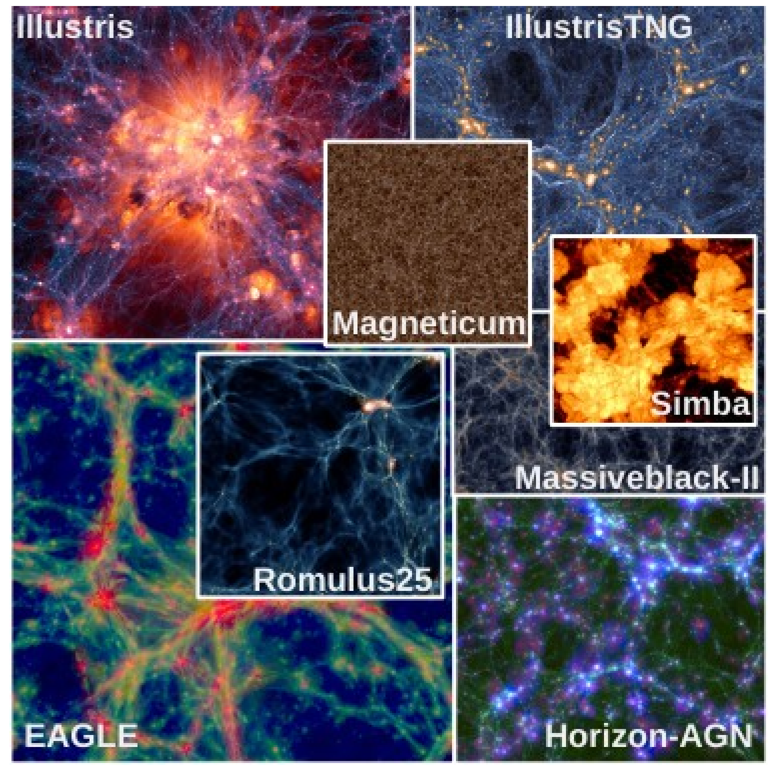


Zoom-in simulations of "Milky Way size objects"

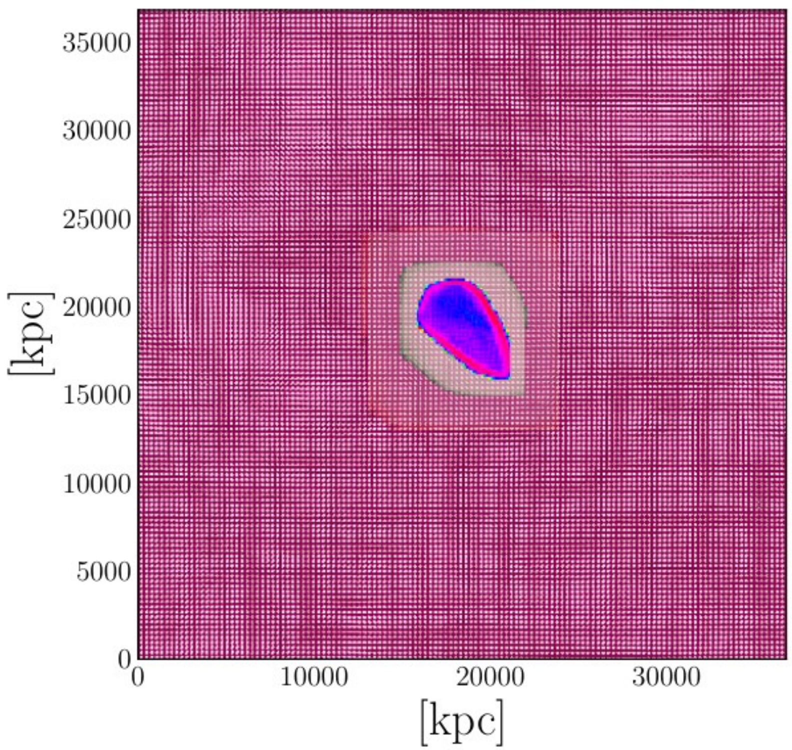
Zoom-in simulations of “Milky Way size objects”

Increase resolution around the initial Lagrangian volume of interest

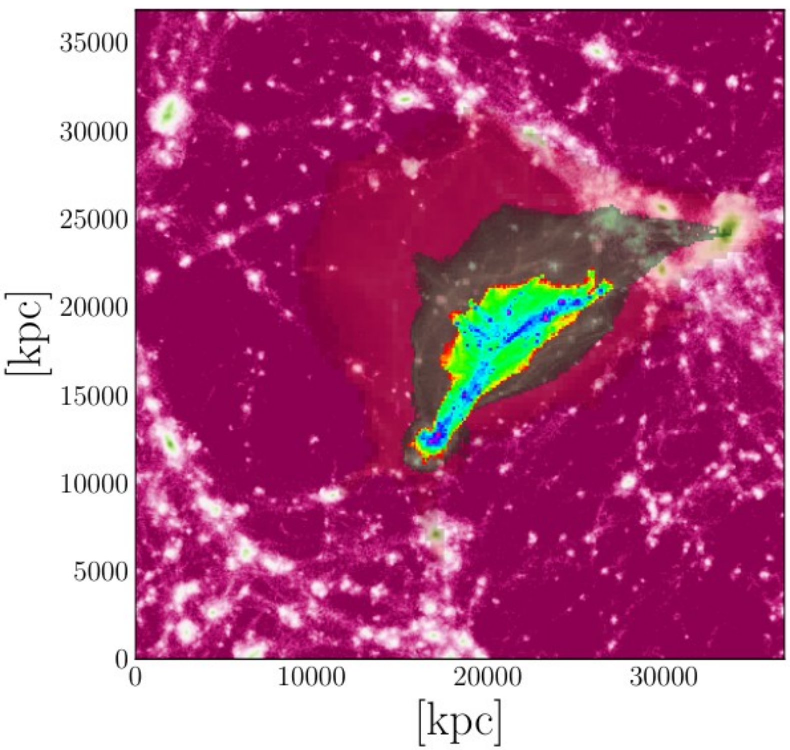
~ 10-50 Mpc boxes

(Gradual levels of zoom)

Begining of the simulation



End of the simulation



A. Nunez-Castineyra PhD

Density maps

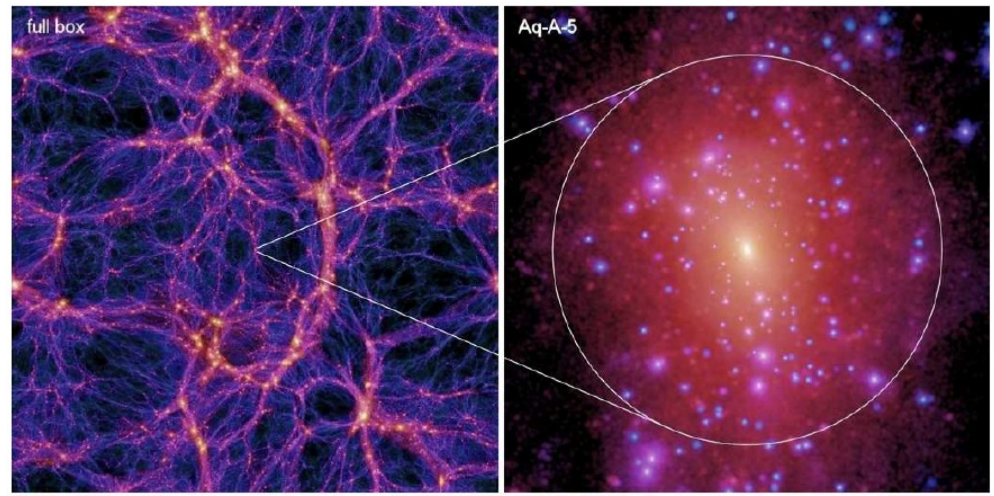
Zoom-in simulations of “Milky Way size objects”

Increase resolution around the initial
Lagrangian volume of interest

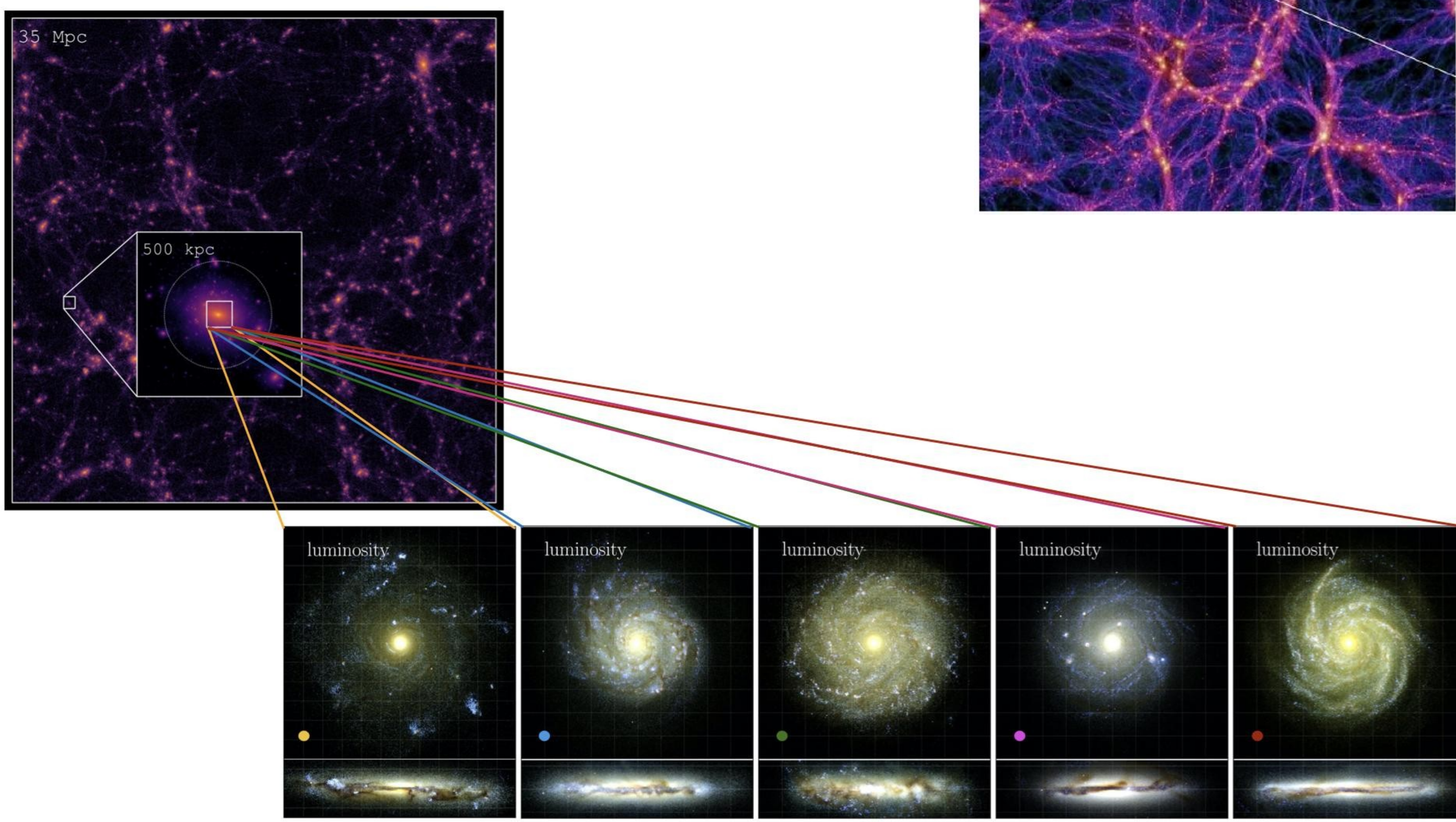
~10^{12} M⊙ halo

Aquarius

Springel et al 2009



arXiv:0809.0898



Mochima

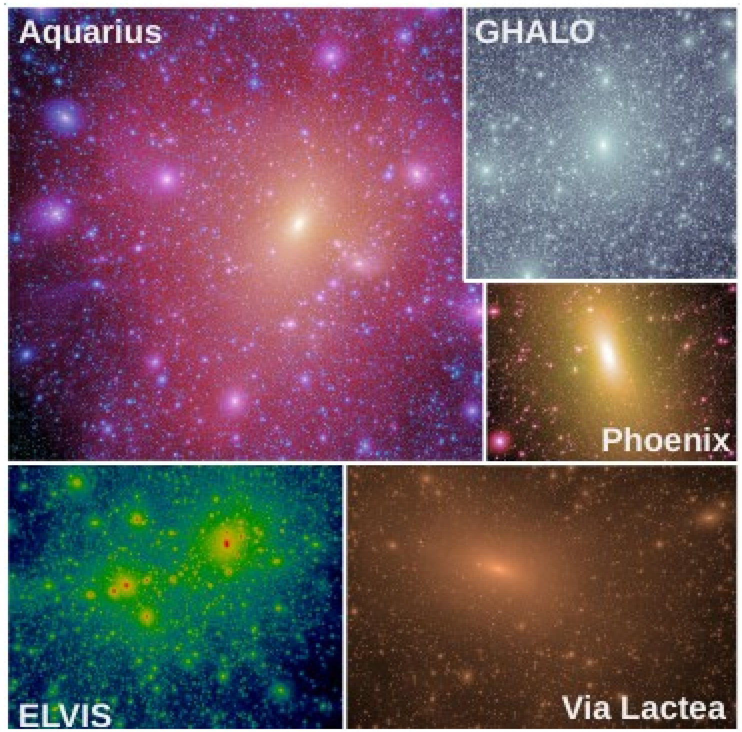
Nunez- Castineyra,

EN, Devriendt, Teyssier 2020

arXiv:2004.06008

Zoom-in simulations of “Milky Way size objects”

- Dark matter only (DMO): Zoom simulations of Milky Way size haloes



Zoom-in simulations of “Milky Way size objects”

- Dark matter only (DMO): Zoom simulations of Milky Way size haloes

Dark matter distribution ?

Substructures

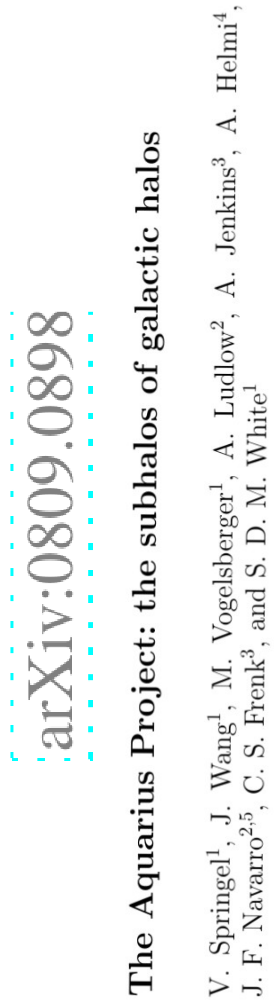
Subhalos
Mass spectrum
Concentration
Spatial distribution

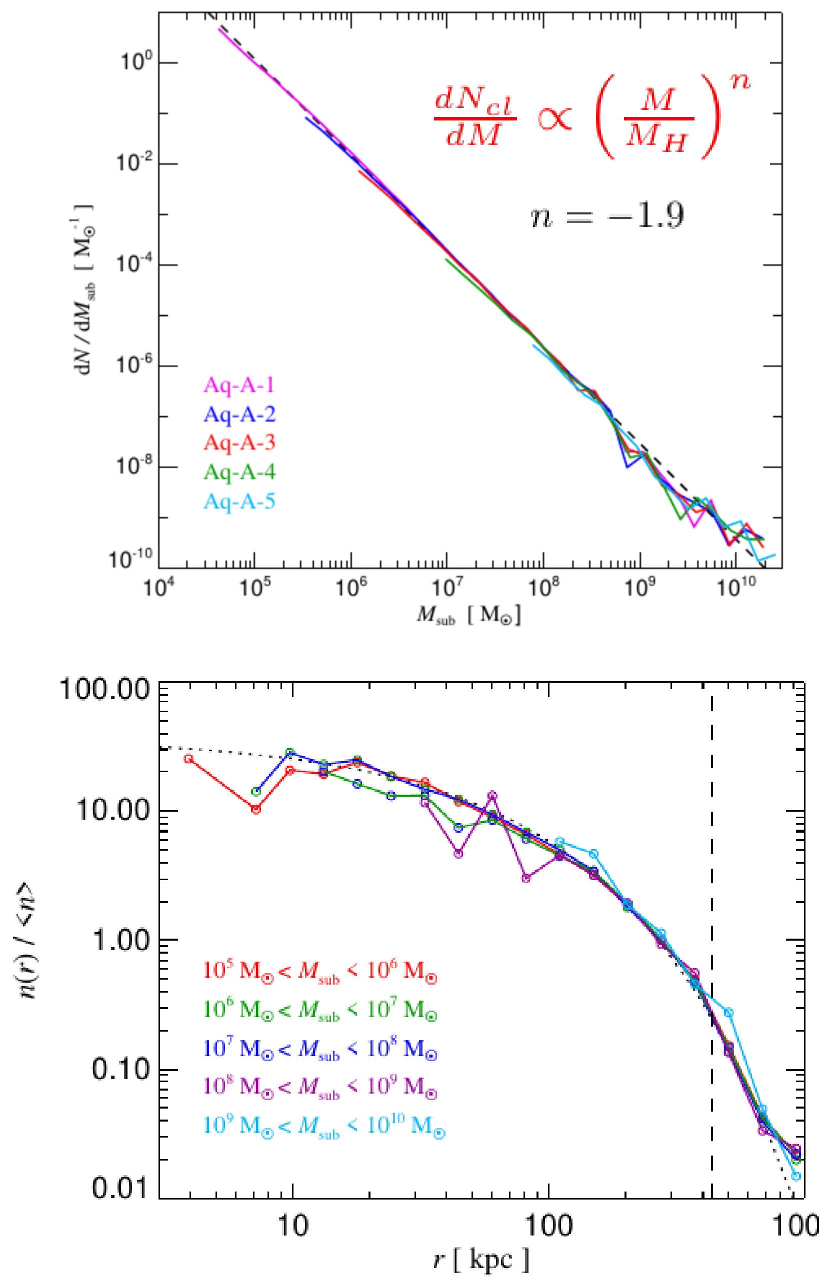
Streams

Main halo

Density profile
Cusp/NFW
Einasto

Velocity distribution





Zoom-in simulations of “Milky Way size objects”

- Dark matter only (DMO): Zoom simulations of Milky Way size haloes

NFW $\rho(r) = \rho_s r_s^3 / r(r + r_s)^2$

Einasto $\rho(r) = \rho_{-2} \exp[-2\alpha^{-1}((r/r_{-2})^\alpha - 1)]$

Dark matter distribution ?

arXiv:1911.09720

Universal structure of dark matter haloes over a mass range of 20 orders of magnitude

Wang, J.[1,5]*, Bose, S.[2], Frenk, C. S.[3†] Gao, L.[1,5], Jenkins, A.[3], Springel, V.[4] & White, S. D. M.[4‡]

Substructures

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Mass spectrum
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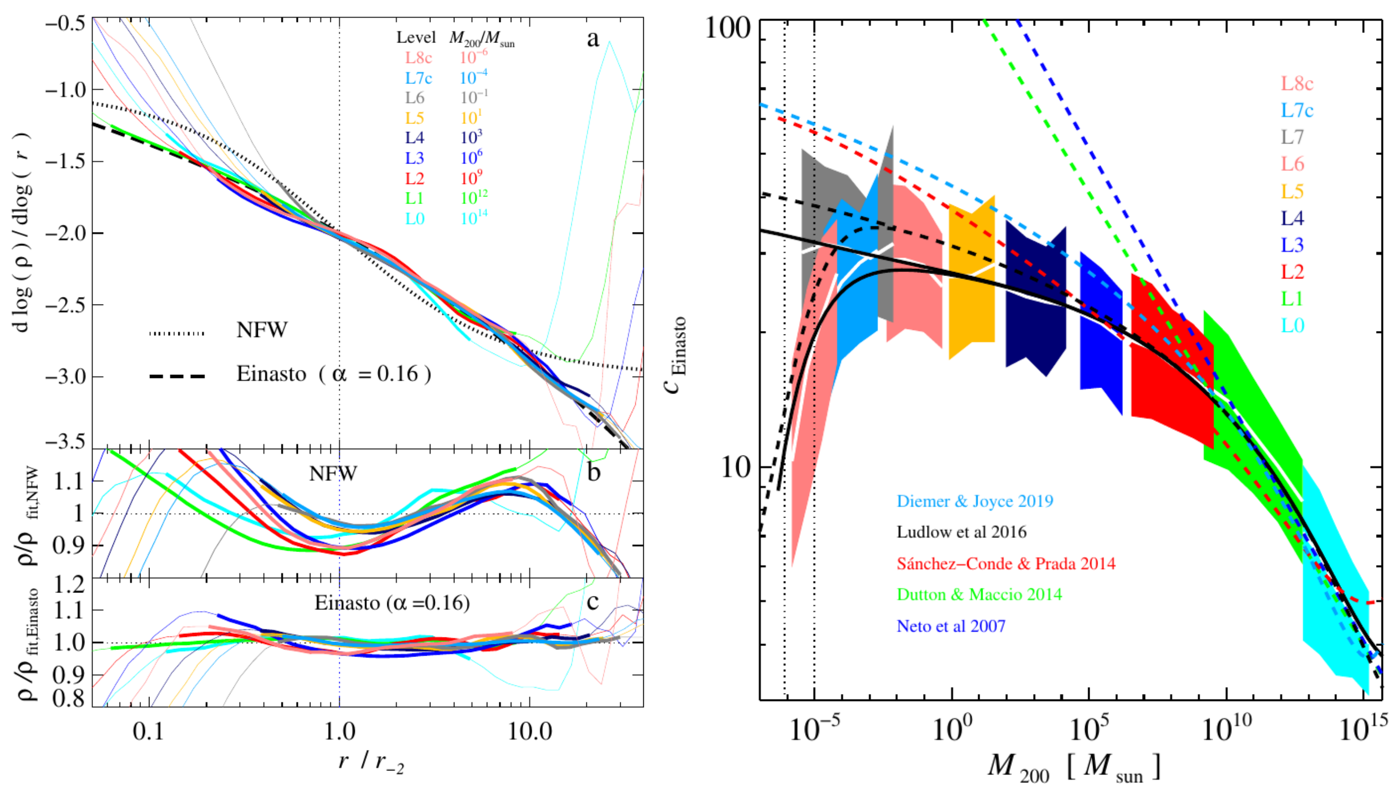

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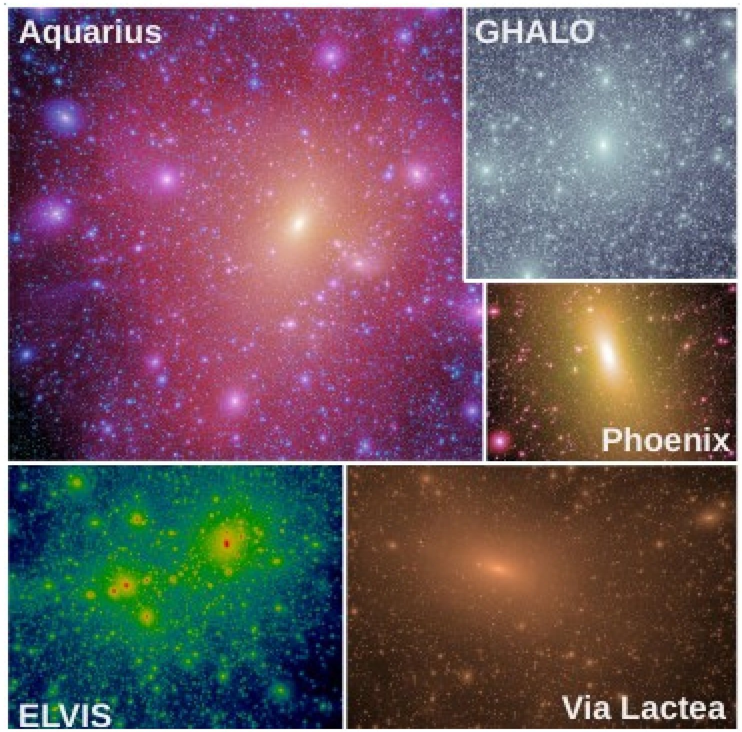

Zoom-in simulations of “Milky Way size objects”

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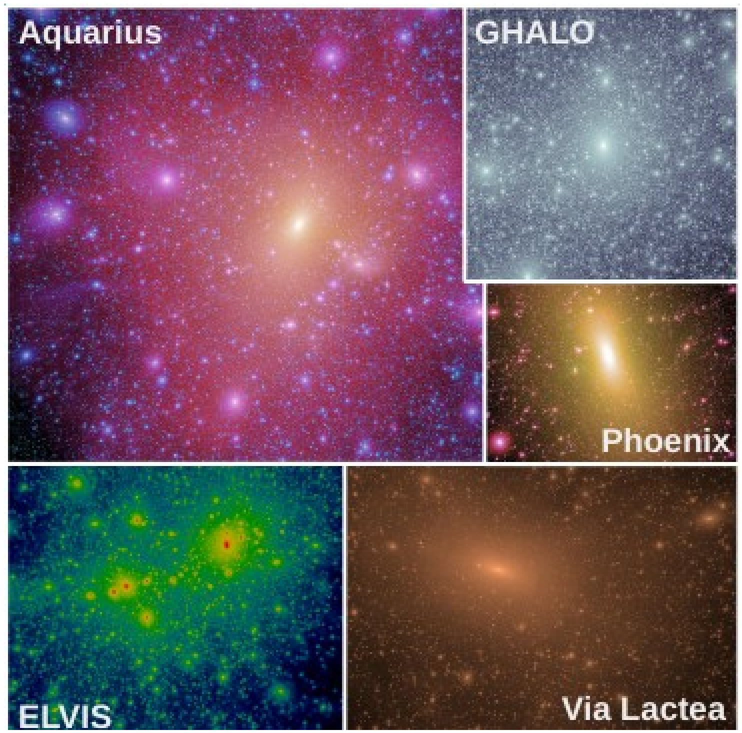
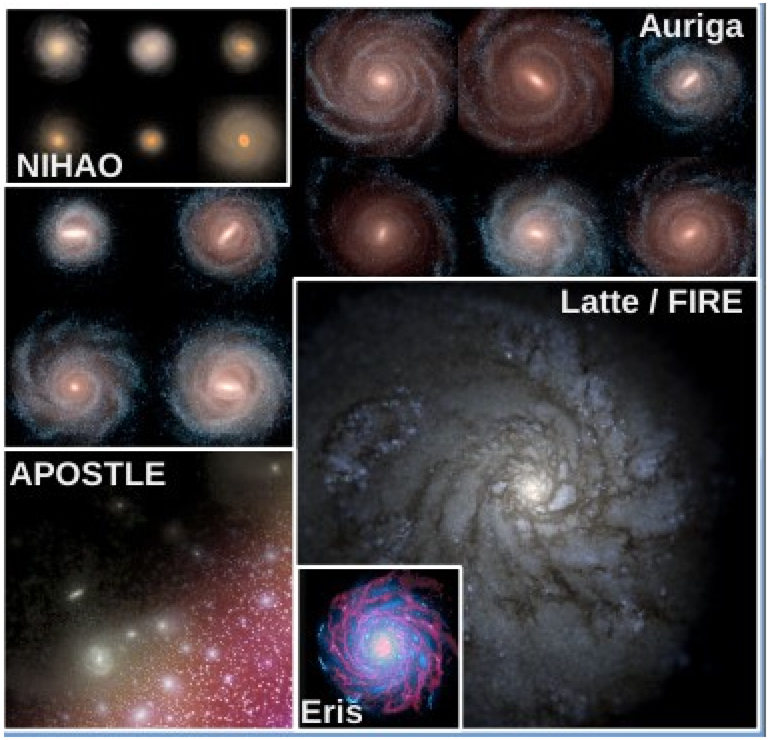


- Hydro: Zoom-in simulations of “Milky Way like” spiral galaxies

ERIS, NIHAO,EAGLE,FIRE, AURIGA,APOSTLE,GIMIC,ARTEMIS,VINTERGATAN, MOCHIMA, NEW HORIZON, ILLUSTRIS TNG …



Milky-Way “analog”: Spiral galaxie in ~ 10^{12} M⊙ halo

Zoom-in simulations of “Milky Way like” spiral galaxies

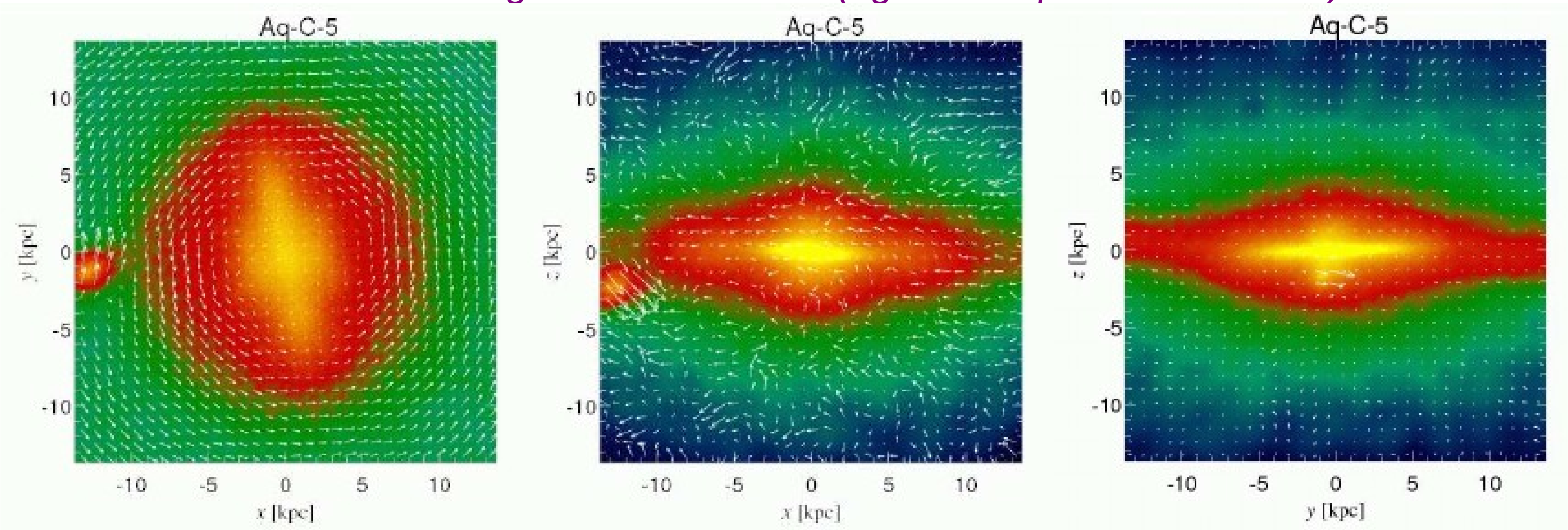
- Some time ago:

Gas cooling, star formation

Angular momentum catastrophe/overcooling problem (Balogh et al 2001, Brook et al 2011)

Too efficient SF and gas consumption at high redshift

→ Thick and not enough extended disks (eg Scannapieco et al 2009)



Zoom-in simulations of “Milky Way like” spiral galaxies

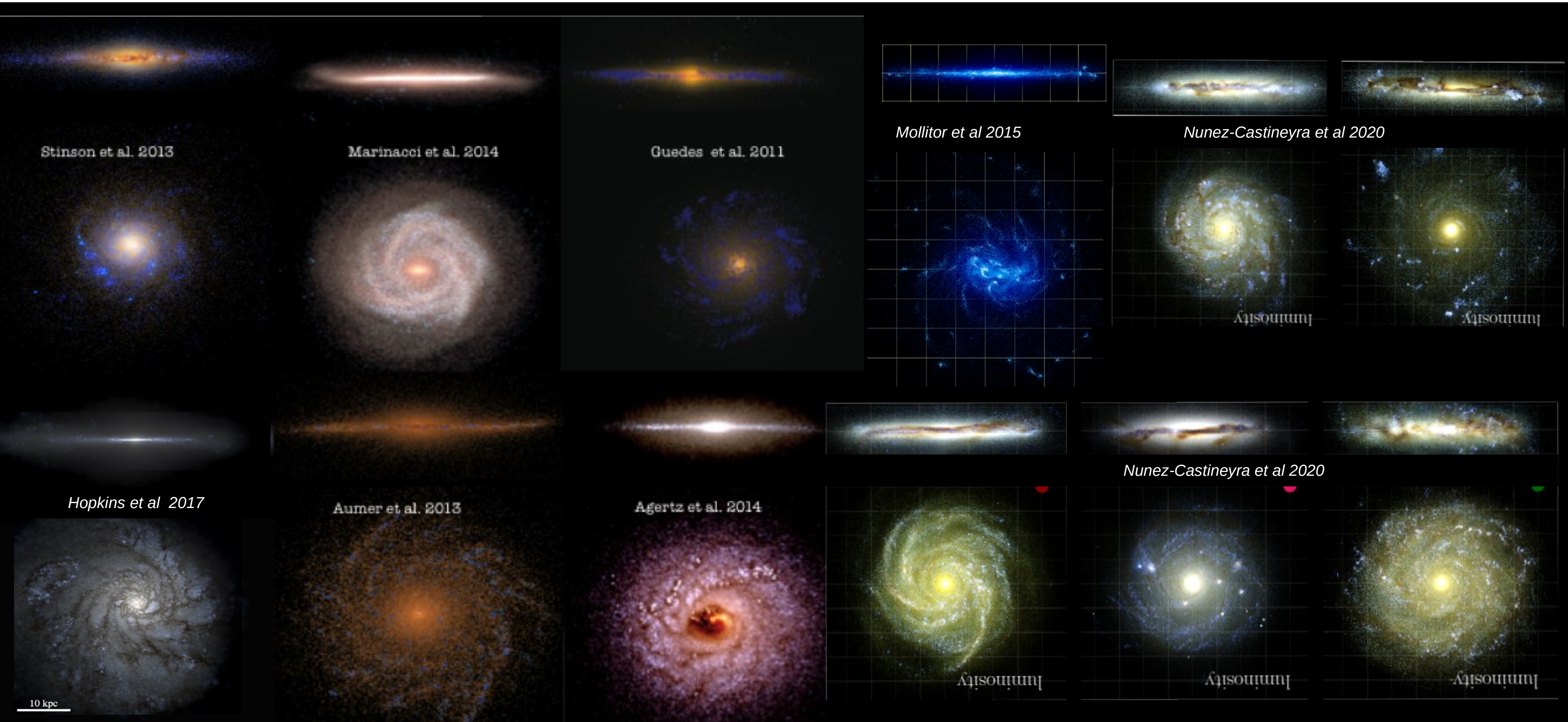
Improve star formation modeling + Including (strong enough) stellar feedback (+wind)

reduce early star formation

better stellar-to-halo mass ratio

close to 1977-78 predictions (Binney, Rees,Ostriker,Silk)

ERIS, NIHAO,EAGLE,FIRE,
AURIGA,APOSTLE
ARTEMIS,VINTERGATAN,MOCHIMA,
NEW HORIZON, ILLUSTRIS TNG ...



(non exhaustive) update of arXiv:1612.06891

Zoom-in simulations of “Milky Way like” spiral galaxies

Adapted from arXiv:1909.07976

Table 2: Recent structure and galaxy formation simulations

simulation	volume	method[a]	mass resolution[b]	spatial resolution[c]	primary reference
	[Mpc^3]		[$M_\odot$]	[kpc]	
Eris	zoom	Tree+SPH	$9.8\times10^4/2\times10^4$	0.12/0.12	Guedes et al. (2011)[349]
VELA	zoom	PM/ML + AMR	$8.3\times10^4/1.9\times10^5$	0.03/0.03[g]	Ceverino et al. (2014)[386]
NIHAO	zoom	Tree+SPH	$3.4\times10^3/6.2\times10^2$	0.12/0.05	Wang et al. (2015)[125]
APOSTLE	zoom	TreePM+SPH	$5.0\times10^4/1.0\times10^4$	0.13/0.13	Sawala et al. (2016)[387]
Latte/FIRE	zoom	TreePM+MLFM	$3.5\times10^4/7.1\times10^3$	0.02/0.001	Wetzel et al. (2016)[352]
Auriga	zoom	TreePM+MMFV	$4.0\times10^4/6.0\times10^3$	0.18/0.18[h]	Grand et al. (2017)[297]
Artemis	*zoom*	*SPH*	*2×10^4*	*0.125*	*Font et al 2020*
Vintergatan	*zoom*	*PM/ML+AMR*	*3.5×10^4 / 7.07×10^3*	*0.02*	*Agertz et al 2020*
Mochima	*zoom*	*PM/ML+AMR*	*1.9×10^5 / 5×10^4*	*0.035/0.035*	*Nunez-Castineyra et al 2020*

~ hundreds of thousands to millions CPU hours

Zoom-in simulations of “Milky Way like” spiral galaxies

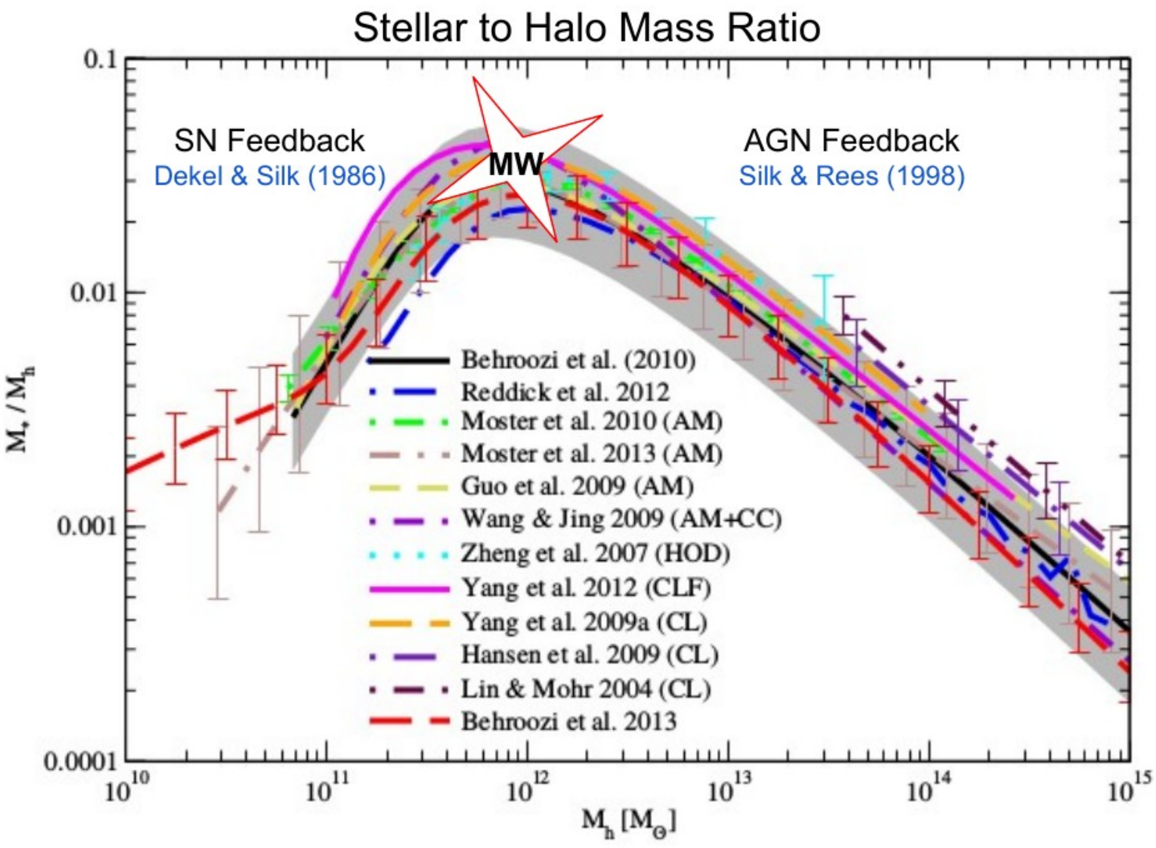
Selection of results:

- *Properties of simulated galaxies*

- *Dark matter distribution features of haloes*

Galaxies

- ***Stellar-to-halo mass ratio***
- *Star formation history*
- *Disk, bulge properties Surface density*
- *Chemistery*
- *Star forming gas region properties*

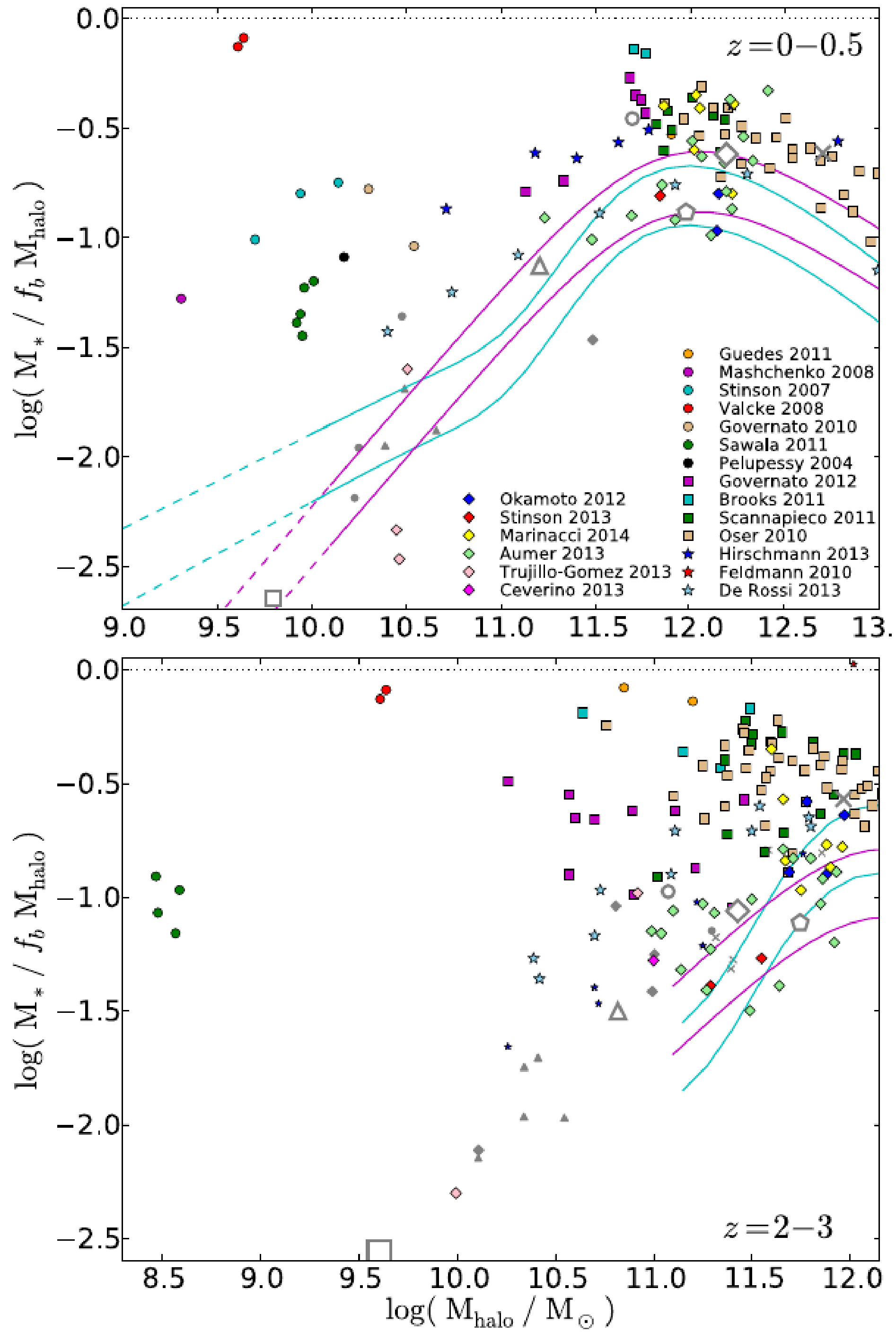


Stellar mass for 10^{12} M⊙ haloes

Specific scale

Galaxies

- **Stellar-to-halo mass ratio**

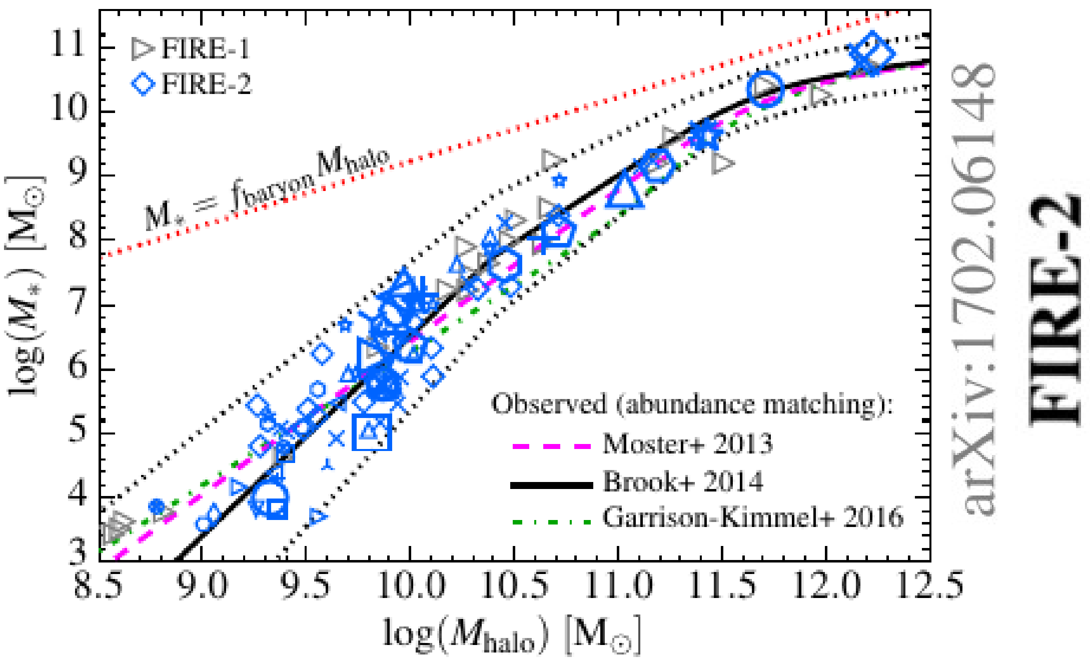


arXiv:1311.2073

Galaxies

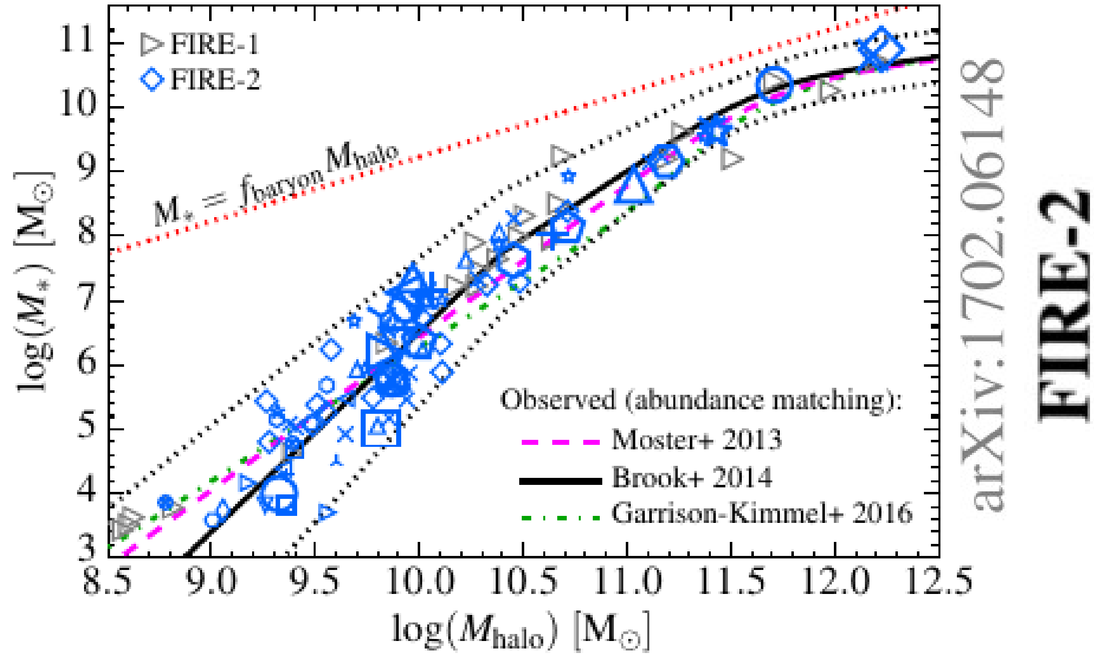
- ***Stellar-to-halo mass ratio***

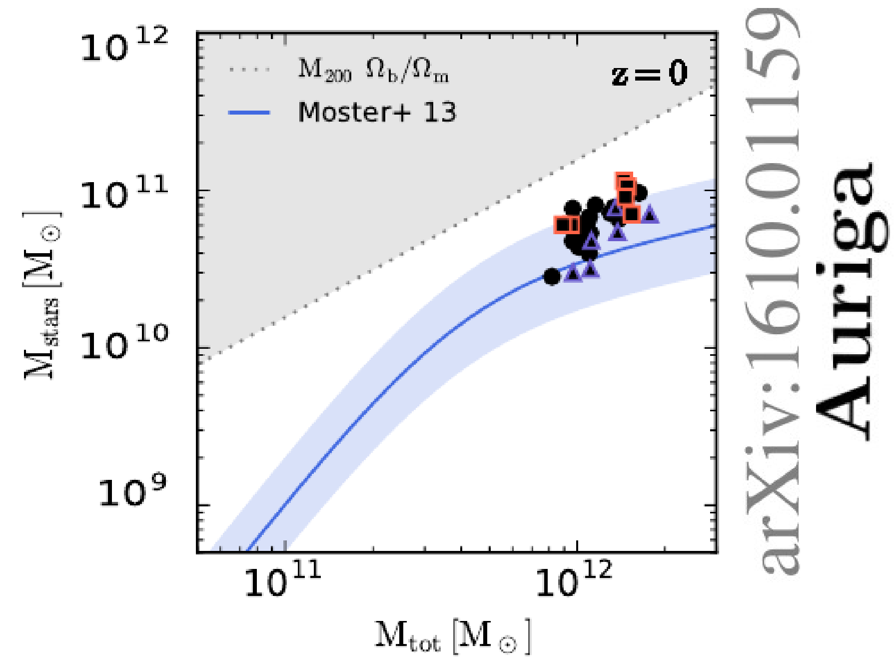
Fairly good agreement



- ***Stellar-to-halo mass ratio***

Fairly good agreement

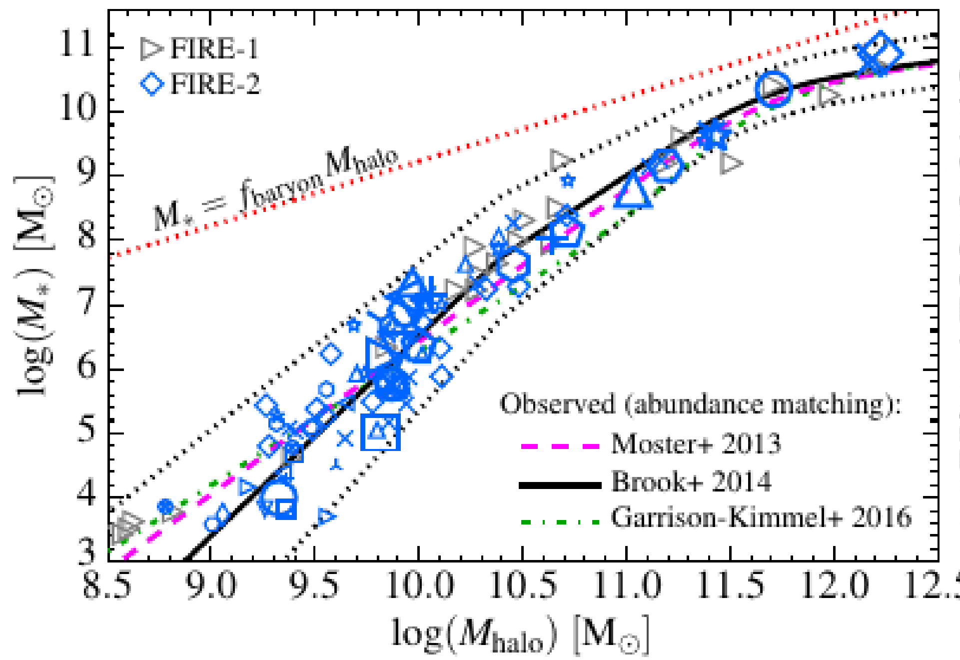




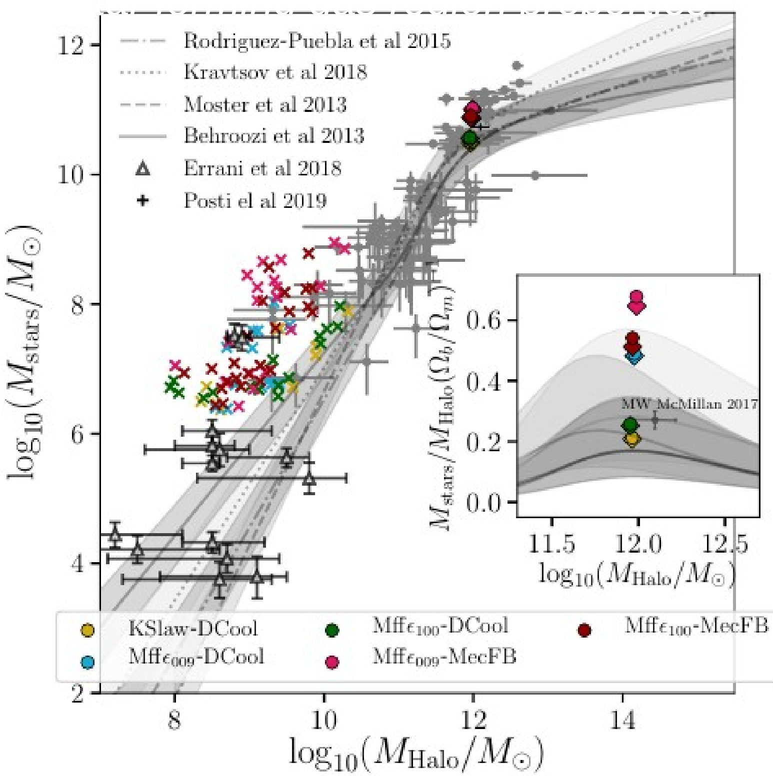
Galaxies

- ***Stellar-to-halo mass ratio***

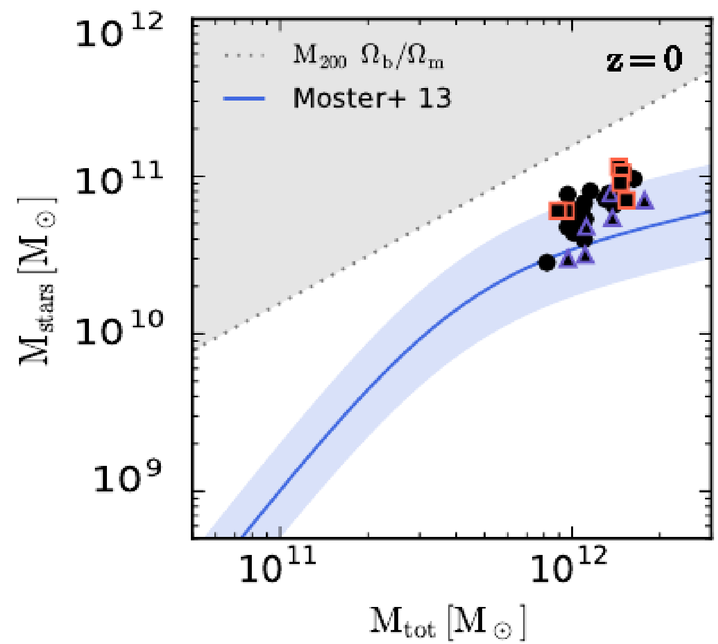
Fairly good agreement



arXiv:1702.06148 **FIRE-2**



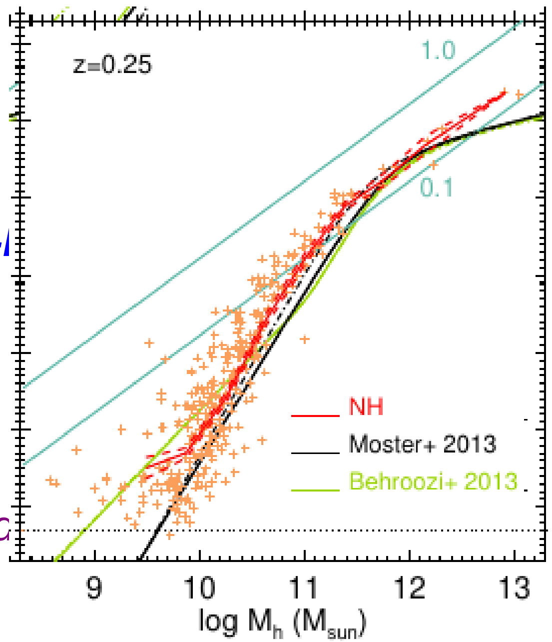
arXiv:2004.06008 Mochima



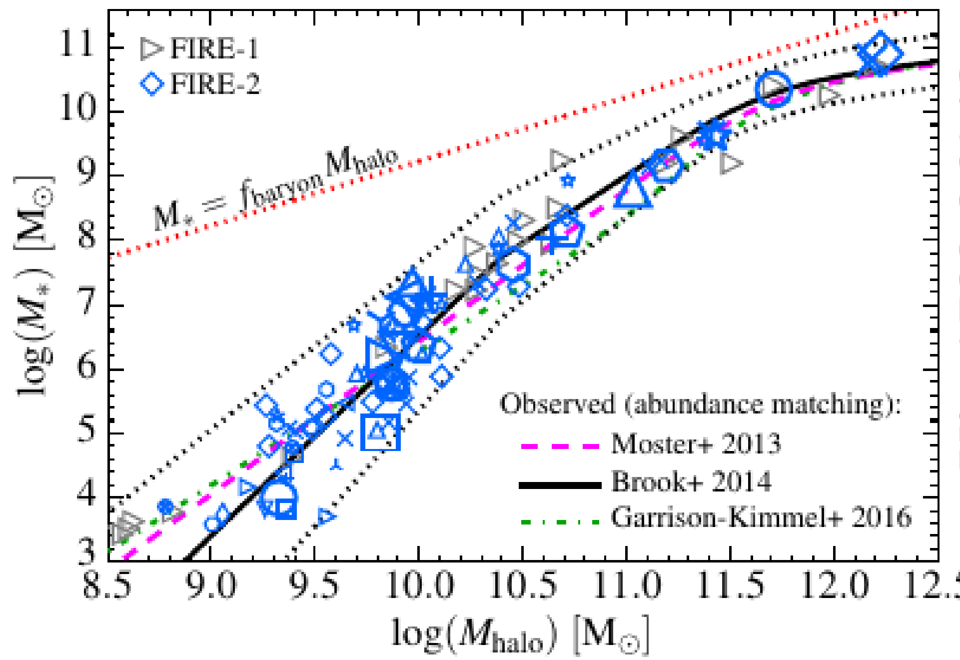
arXiv:1610.01159 Auriga

Galaxies

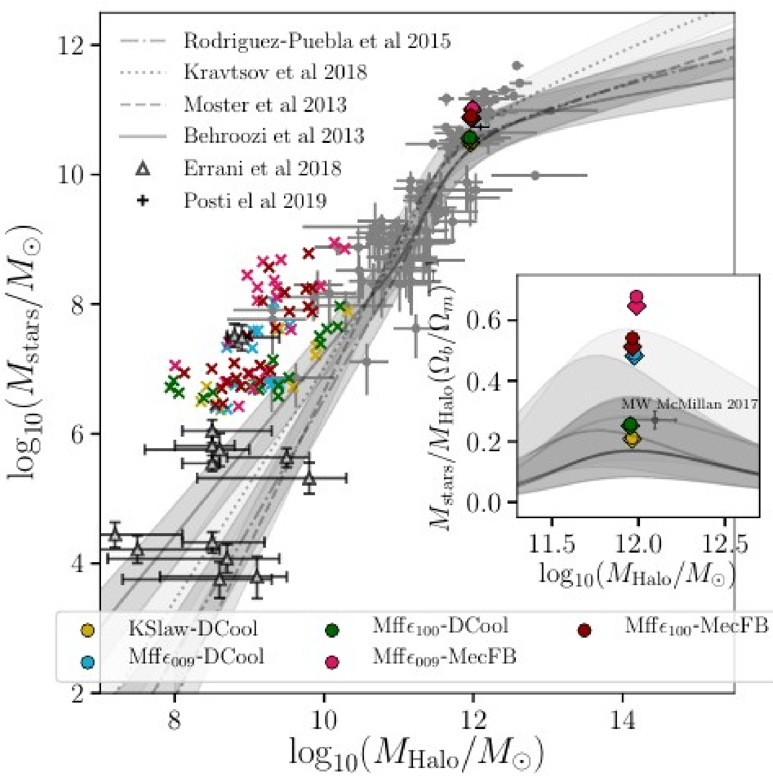
NewHorizon



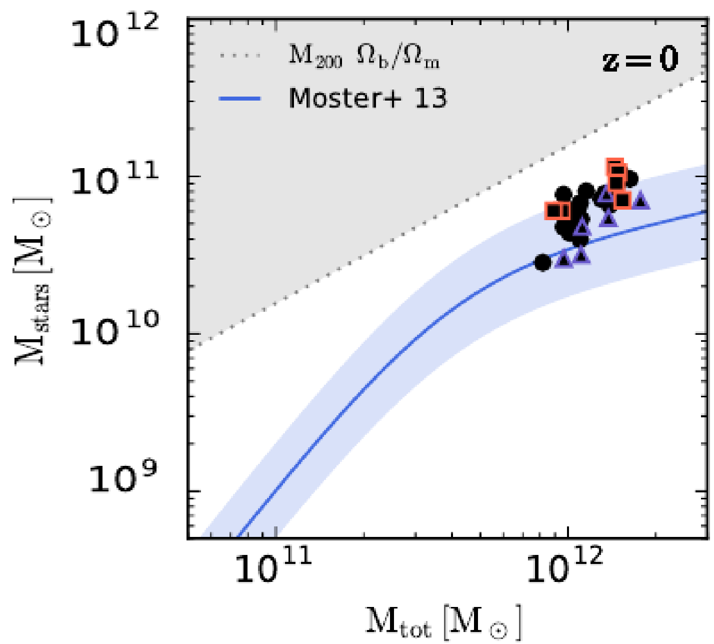
Fairly good



arXiv:1702.06148
FIRE-2

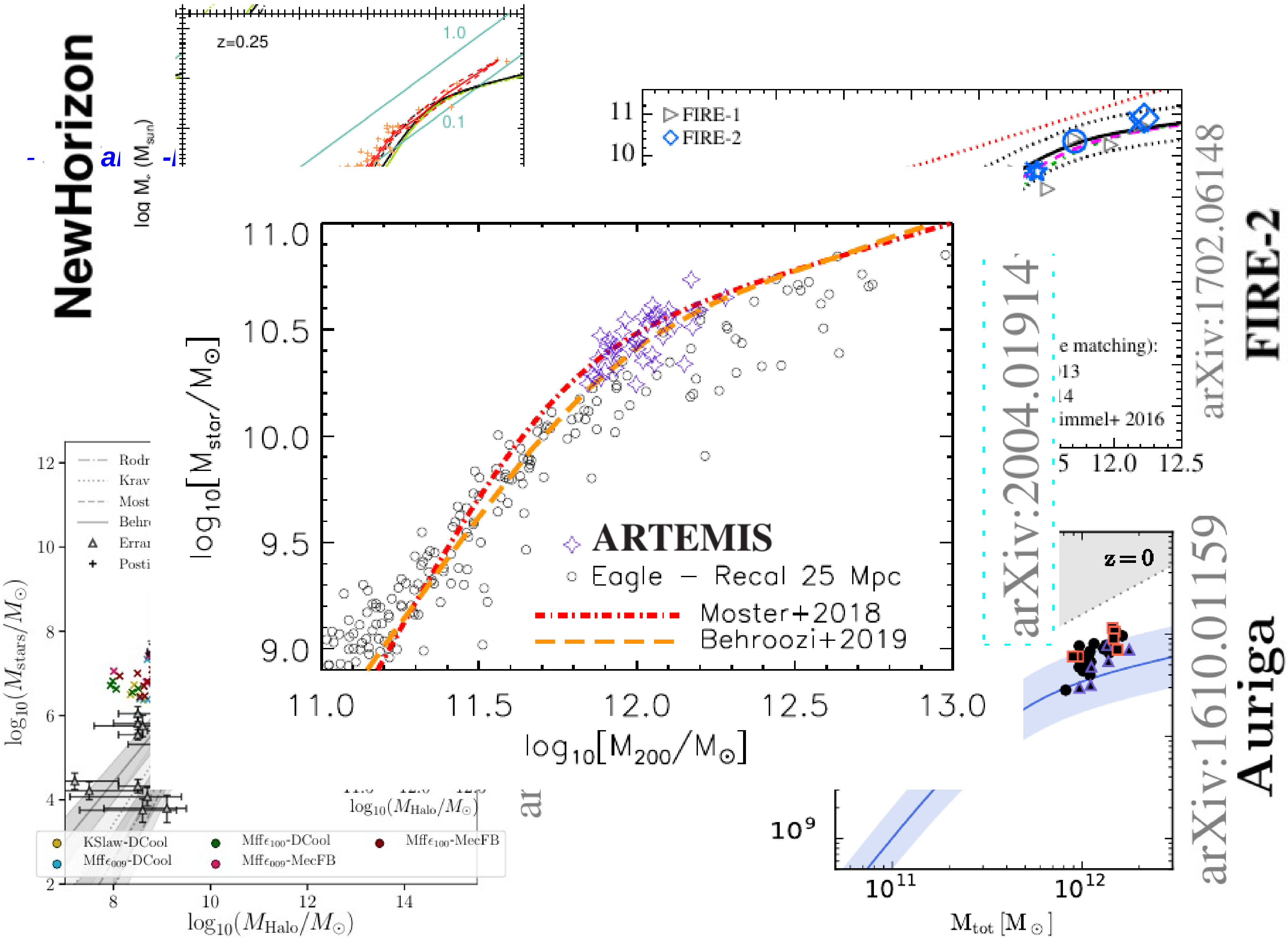


arXiv:2004.06008
Mochima



arXiv:1610.01159
Auriga

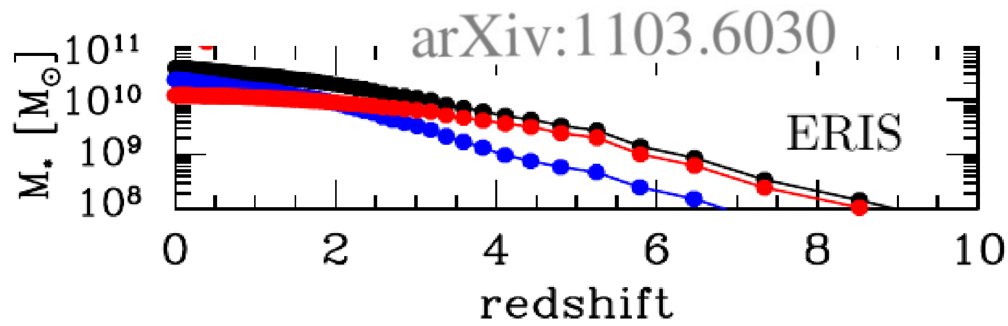
Galaxies



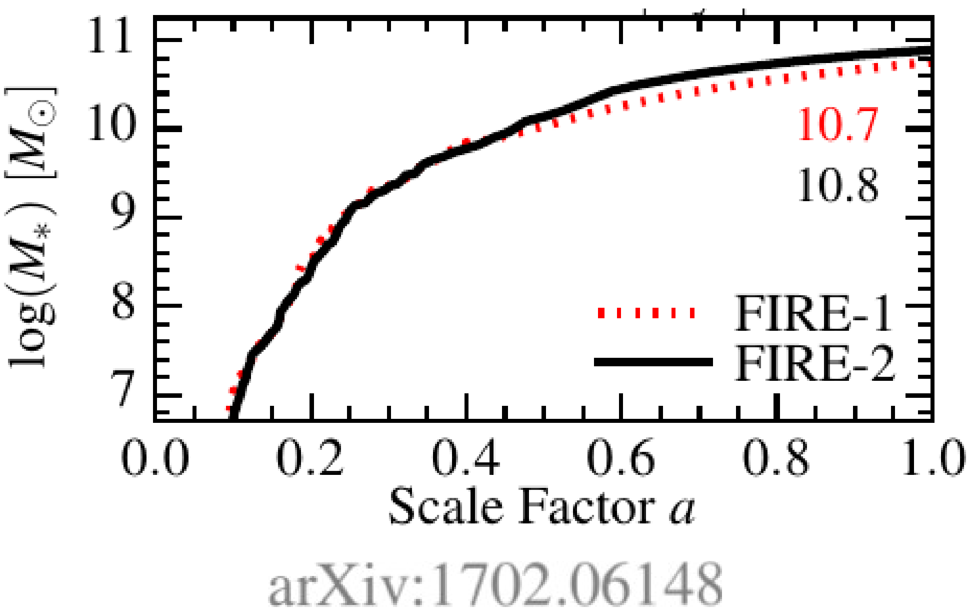
NewHorizon
z=0.25
1.0
0.1
log M* (Msun)
FIRE-1
FIRE-2
e matching):
immel+ 2016
12.0
12.5
arXiv:1702.06148
FIRE-2
ARTEMIS
Eagle – Recal 25 Mpc
Moster+2018
Behroozi+2019
log10[Mstar/M⊙]
log10[M200/M⊙]
11.0
10.5
10.0
9.5
9.0
11.0
11.5
12.0
12.5
13.0
arXiv:2004.01914
z = 0
arXiv:1610.01159
Auriga
10^9
10^11
10^12
Mtot [M⊙]
Rodr
Krav
Most
Behr
Errai
Posti
log10(Mstars/M⊙)
log10(MHalo/M⊙)
KSlaw-DCool
Mffϵ009-DCool
Mffϵ100-DCool
Mffϵ009-MecFB
Mffϵ100-MecFB

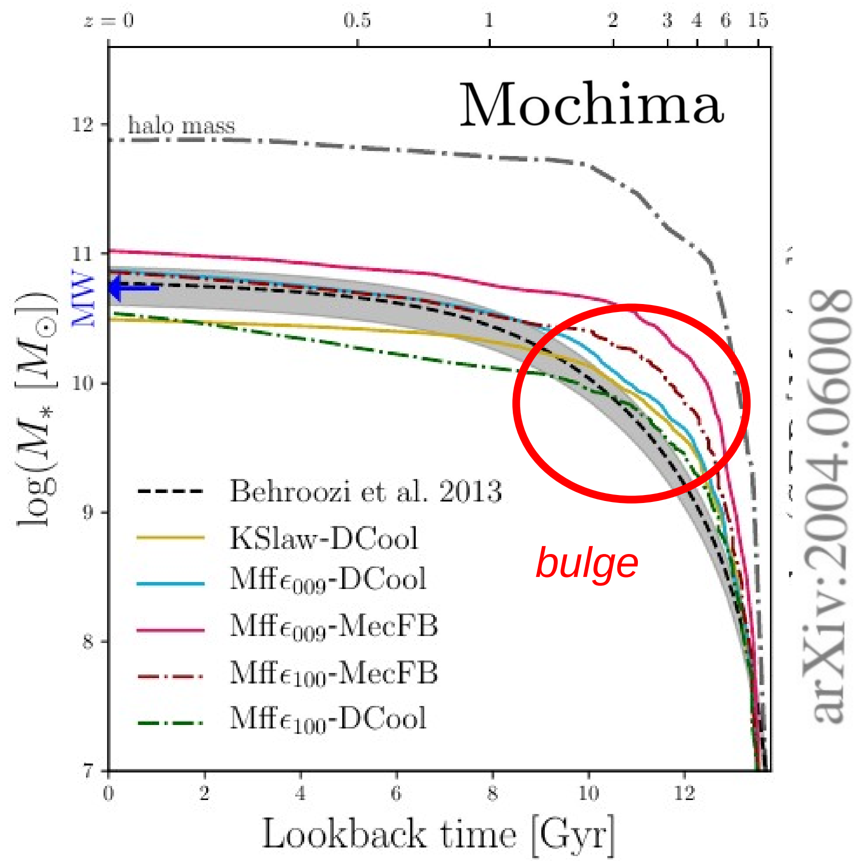
Galaxies

- *Stellar-to-halo mass ratio*
- ***Star formation history***

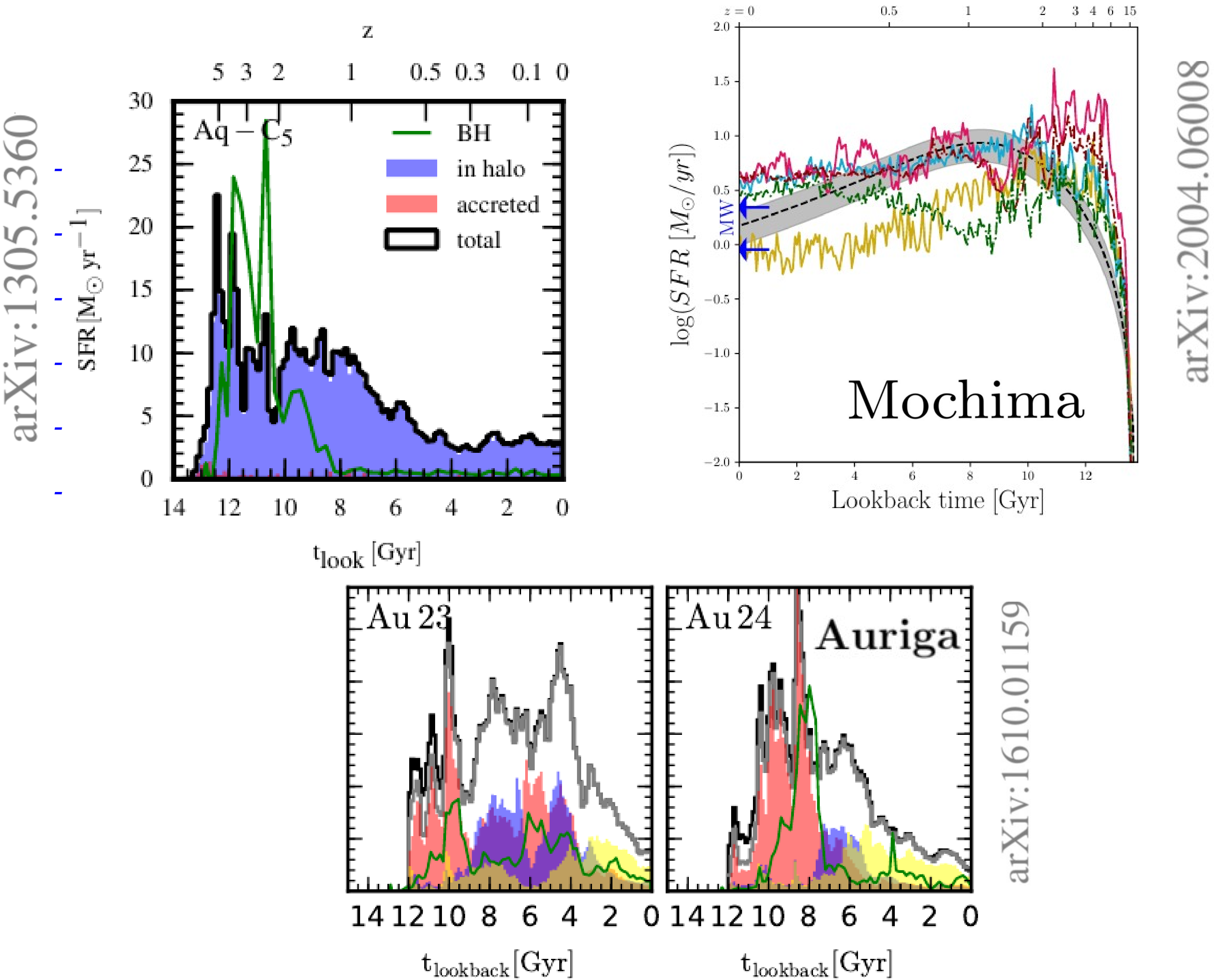


Reduce early star formation





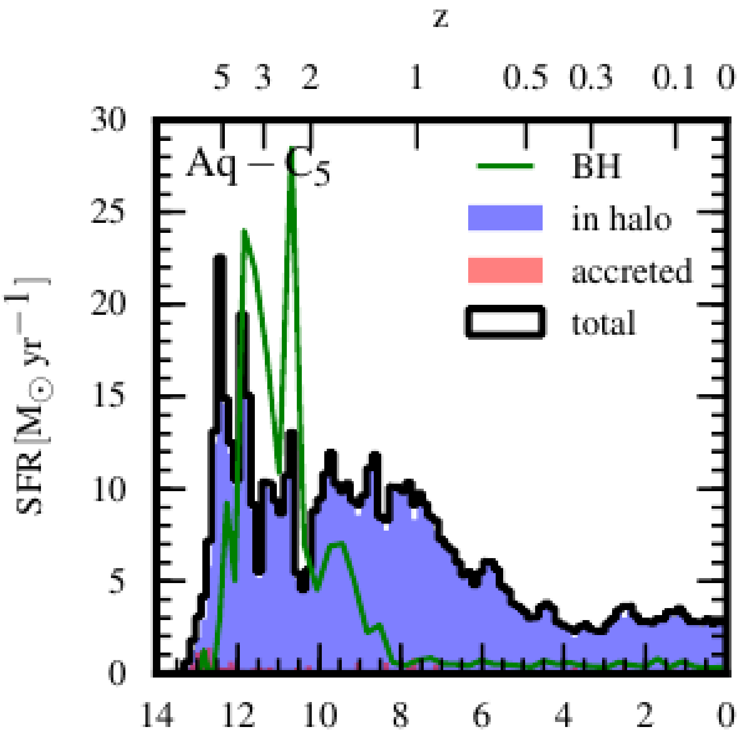
Galaxies

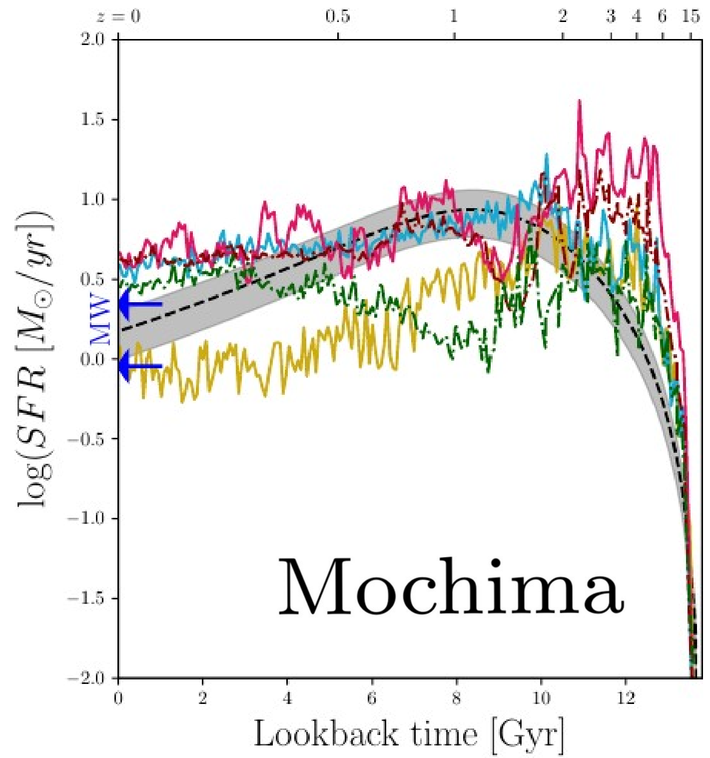


arXiv:1305.5360
Aq − C5
BH
in halo
accreted
total
SFR [M⊙ yr⁻¹]
z
t_look [Gyr]
arXiv:2004.06008
Mochima
log(SFR [M⊙/yr])
MW
Lookback time [Gyr]
arXiv:1610.01159
Au 23
Au 24
Auriga
t_lookback [Gyr]

Galaxies

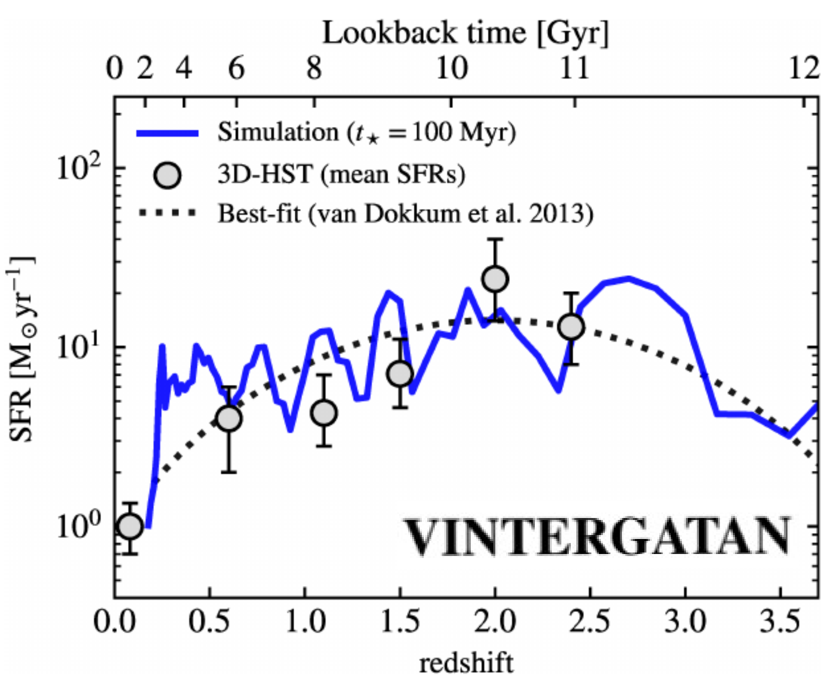
arXiv:1305.5360

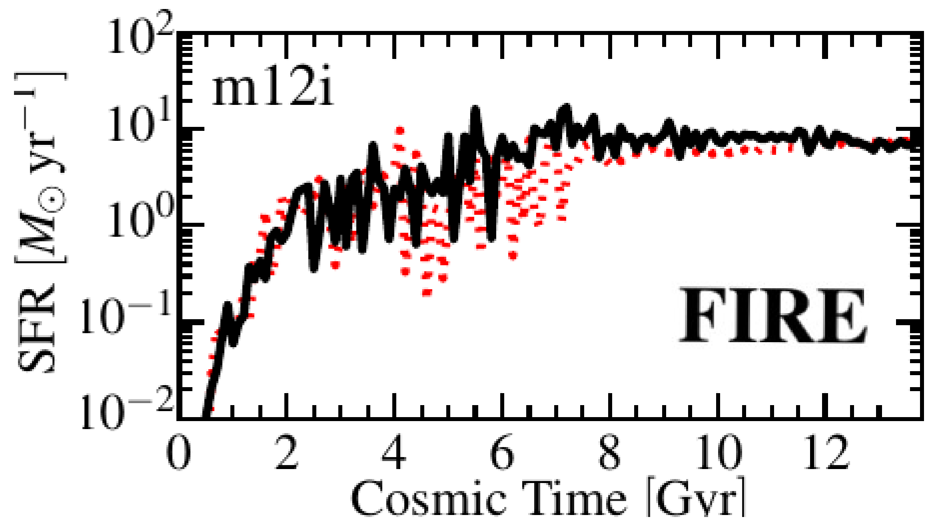




arXiv:2004.06008

arXiv:2006.06008



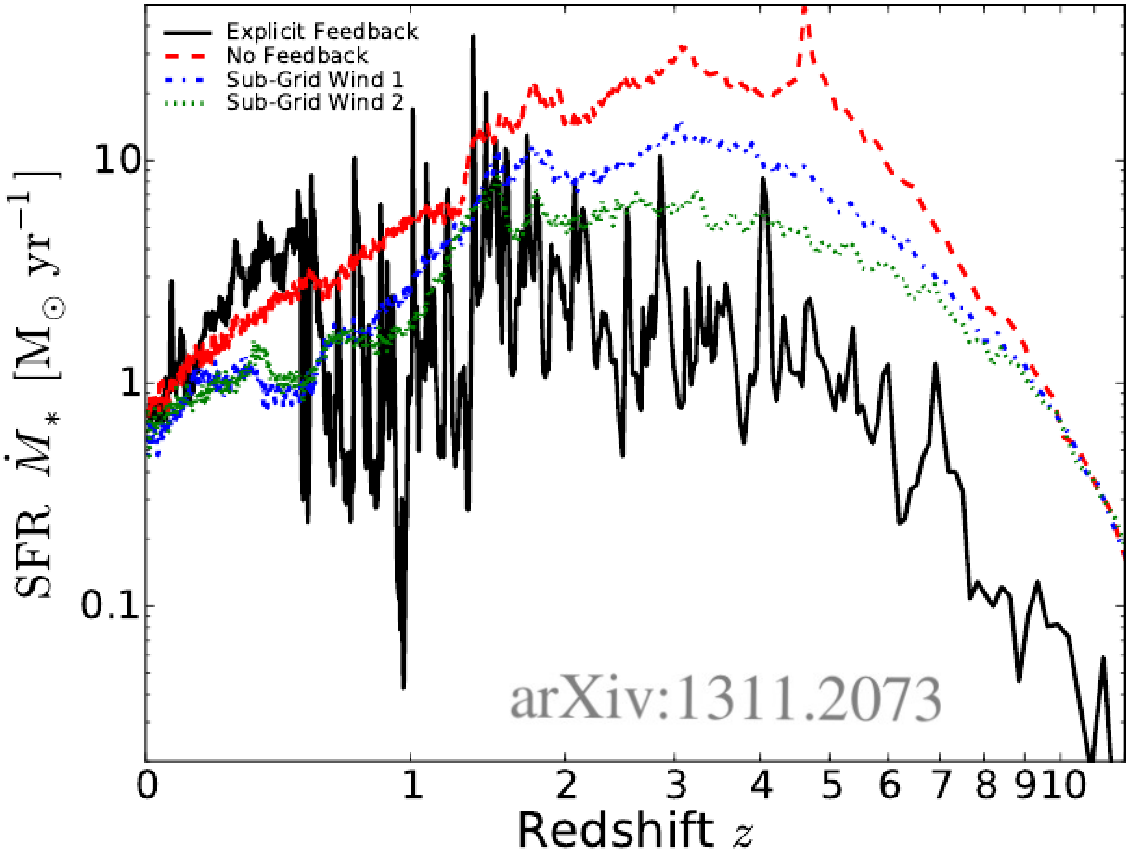


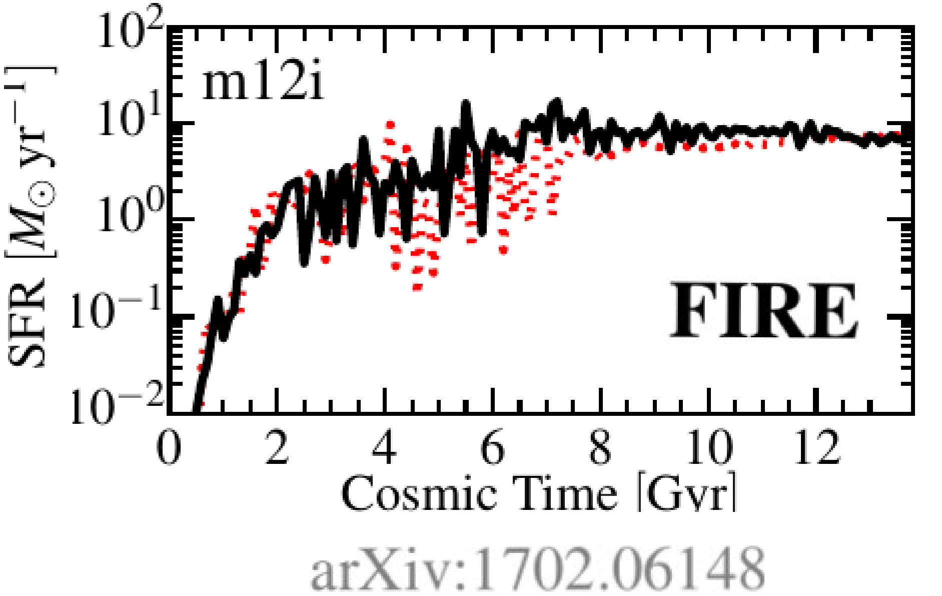
arXiv:1702.06148

Galaxies

- *Stellar-to-halo mass ratio*
- ***Star formation history***

Reduce early star formation





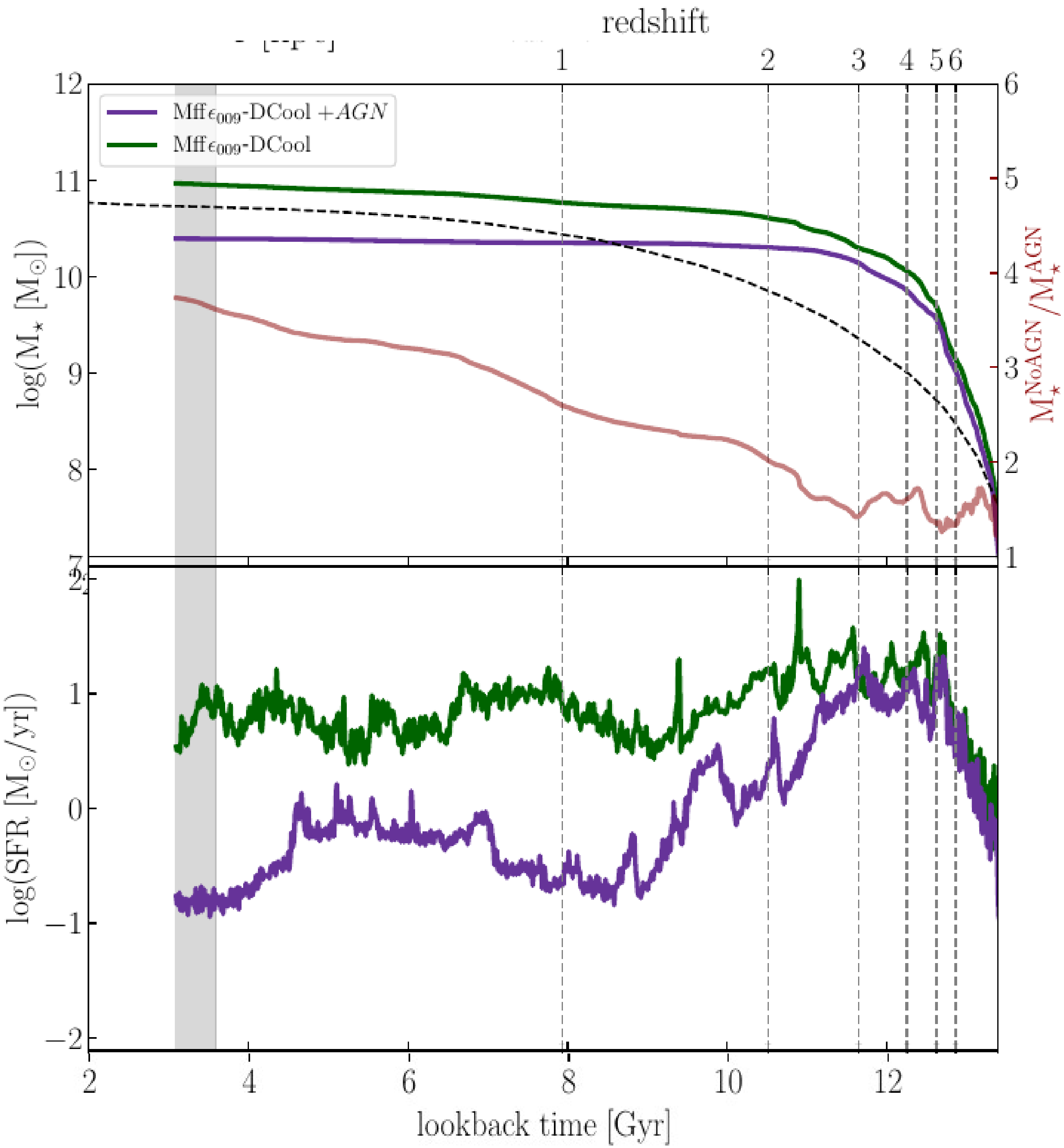
Galaxies

- *Stellar-to-halo mass ratio*
- ***Star formation history***

Reduce early star formation

Mochima

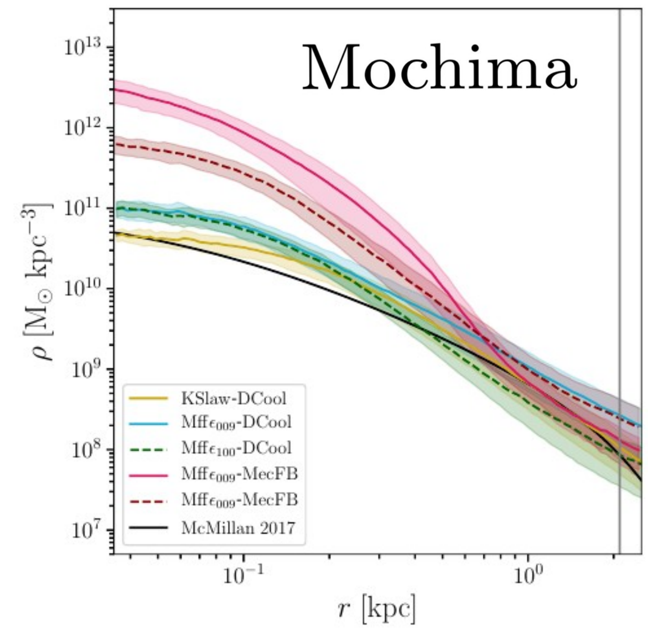
+ AGN (run in progress)
Z~0.4

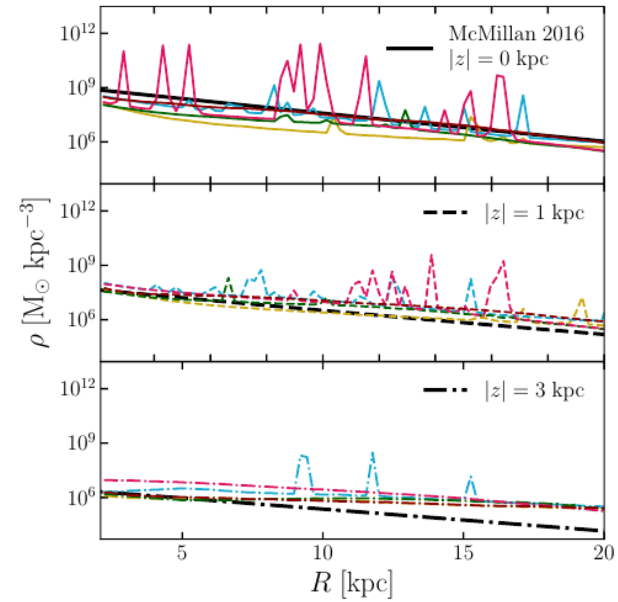


Galaxies

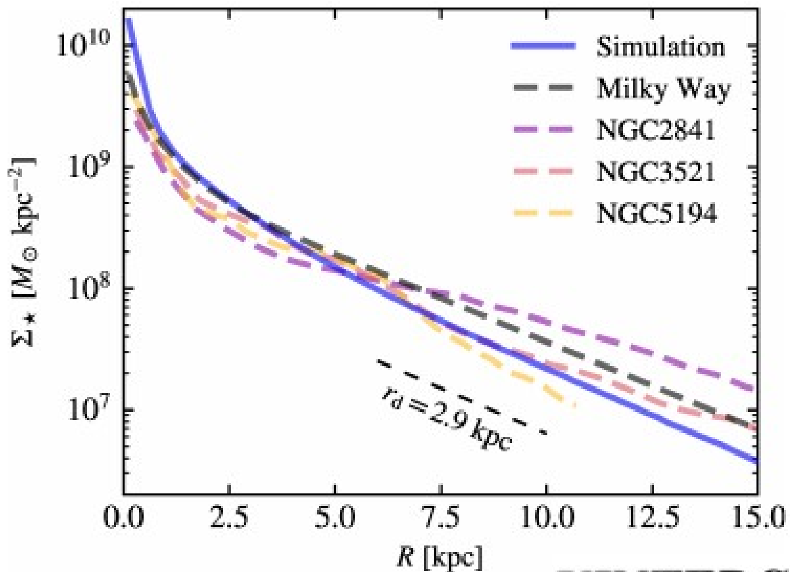
- *Stellar-to-halo mass ratio*
- *Star formation history*
- ***Disk, bulge properties Surface density, Rotation curve***

High central density





arXiv:2004.06008

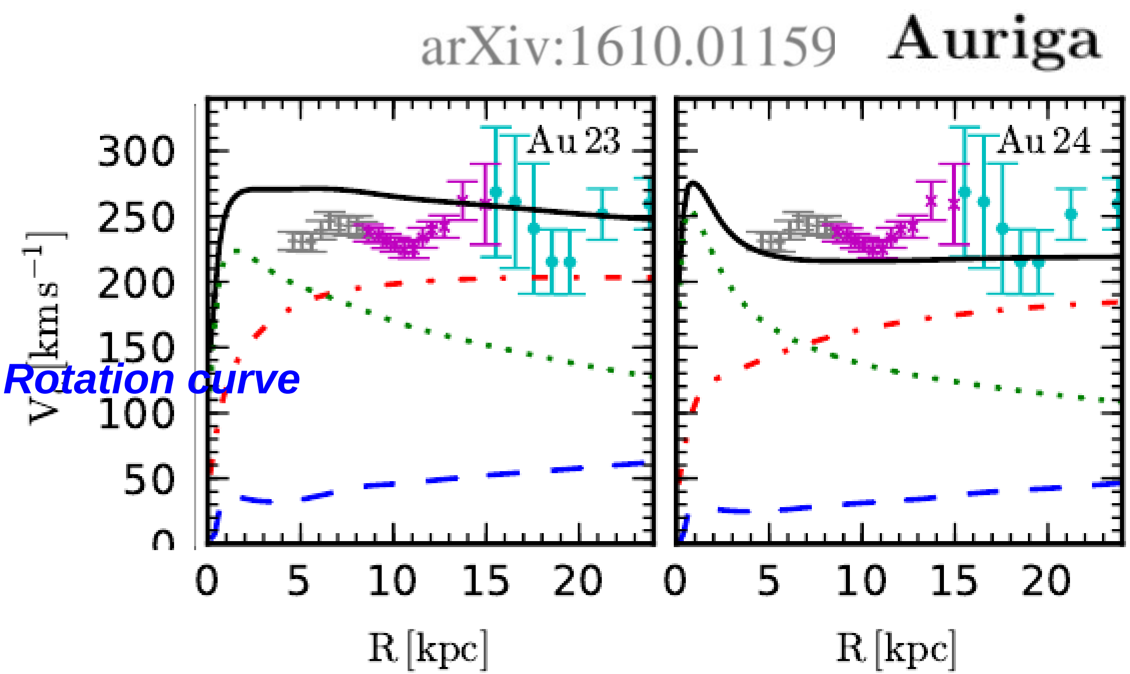


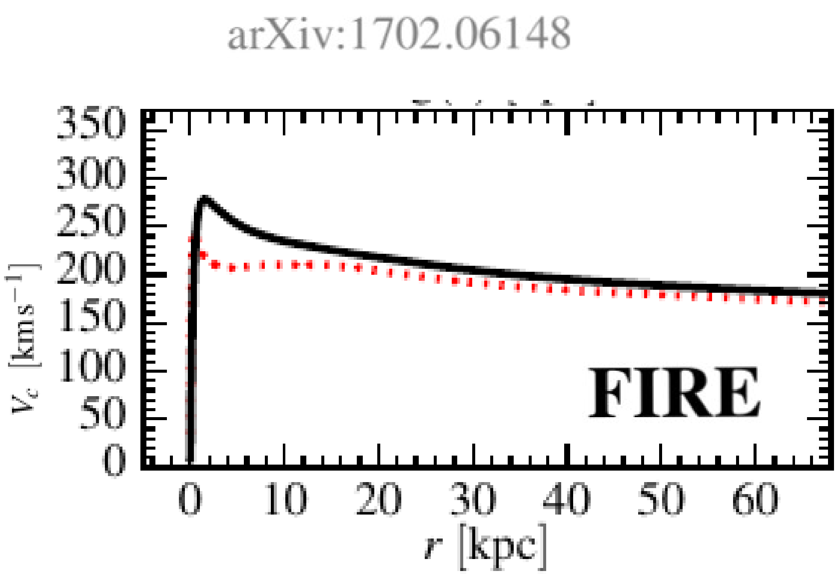
VINTERGATAN

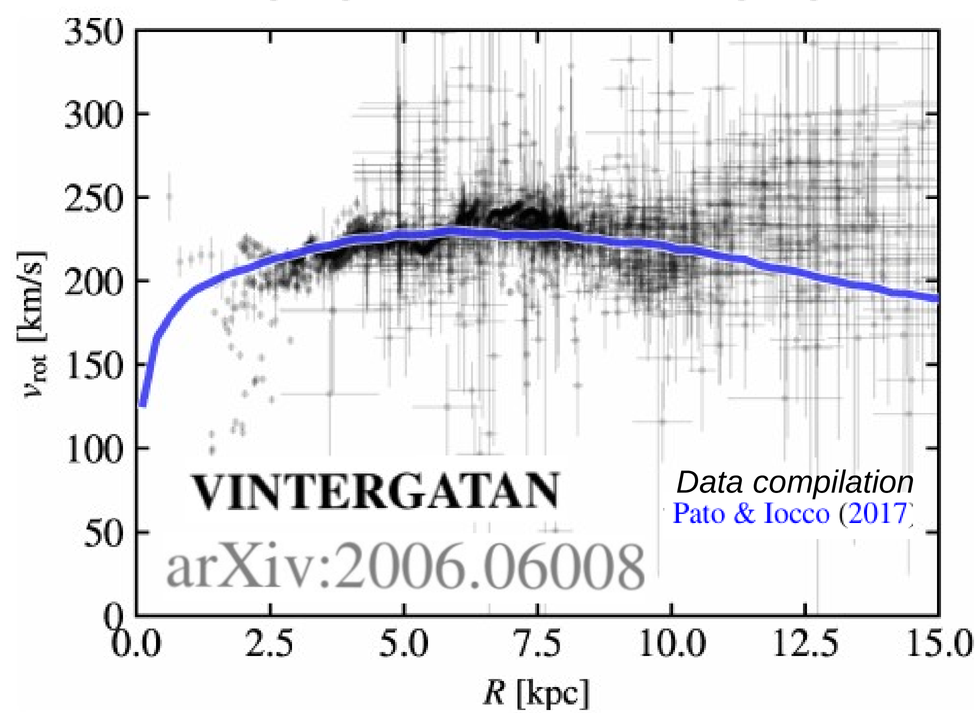
arXiv:2006.06008

Galaxies

- *Stellar-to-halo mass ratio*
- *Star formation history*
- ***Disk, bulge properties Surface density, Rotation curve***







Galaxies

- Stellar-to-halo mass ratio

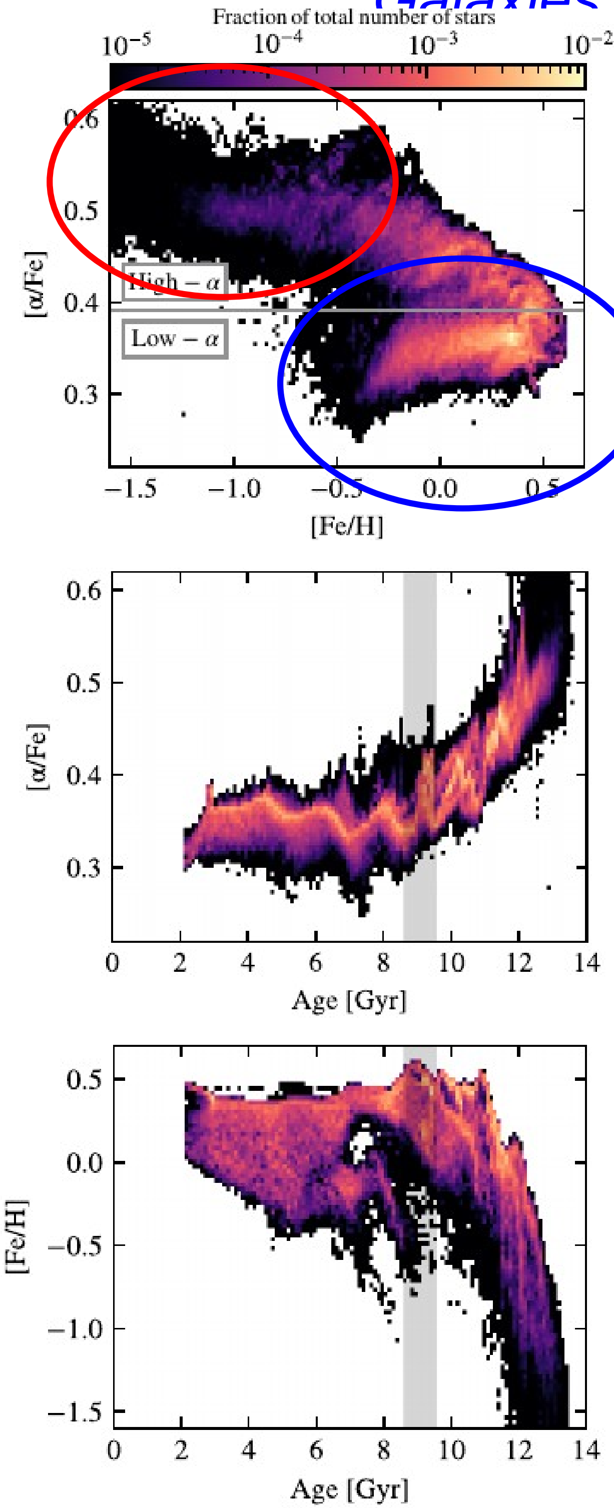
- Star formation history

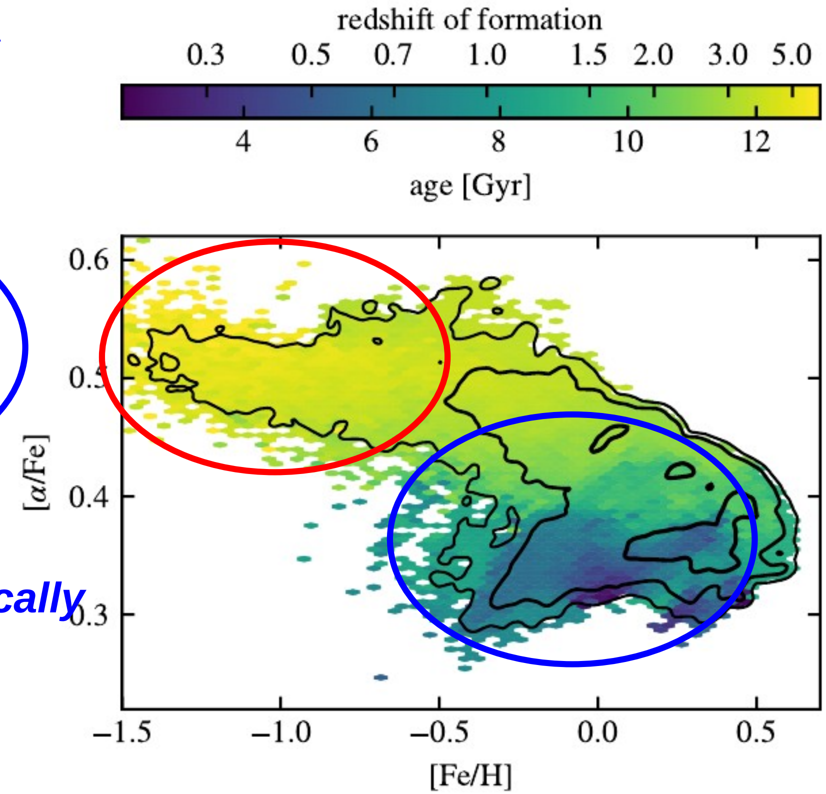
- Disk, bulge properties Surface density, Rotation curve

- Chemistery: identifying thin and thick discs chemically

Galaxies

- *Stellar-to-halo mass r*
- *Star formation history*
- *Disk, bulge propertie*
- ***Chemistery: identify*** ***cally***





Thick disc

Thin disc

Abundance elements (Fe, O, Mg, Si, Ca, and Ti)

Similar to the observed MW bimodality

VINTERGATAN

arXiv:2006.06008

arXiv:2006.06011

Galaxies

- Stellar-to-halo mass ratio

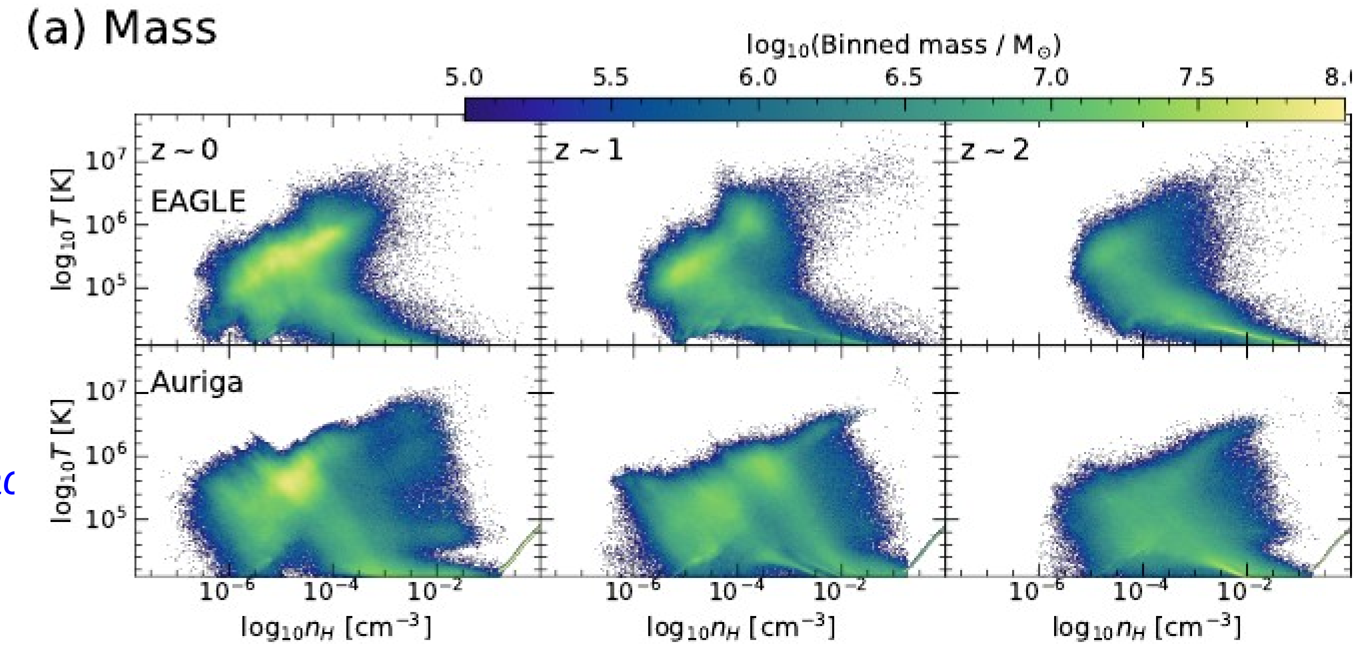
- Star formation history

- Disk, bulge properties Surface density, Rotation curve

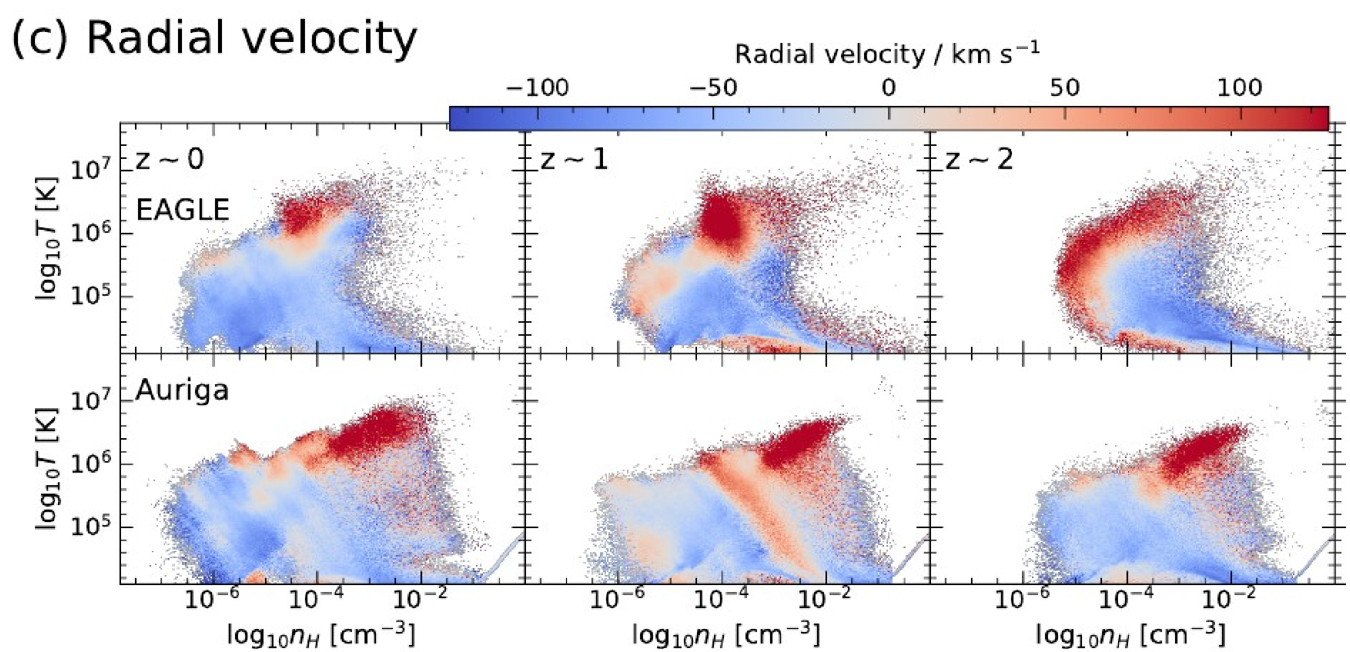
- Chemistery

- Gas cycle, Star forming gas region properties

Galaxies



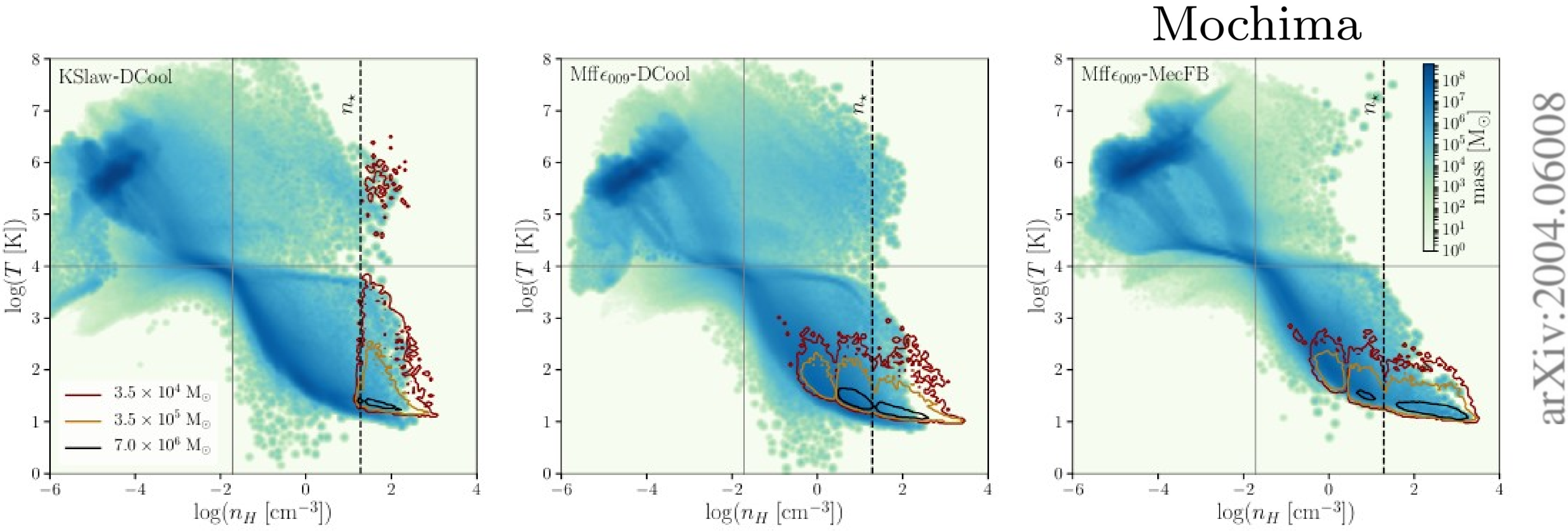
- *Stellar-to-halo mass ratio*
- *Star formation history*
- *Disk, bulge properties Surfac*
- *Chemistery*
- ***Gas cycle, Star forming gas region properties***



Same initial conditions and different baryonic physics → different gas properties

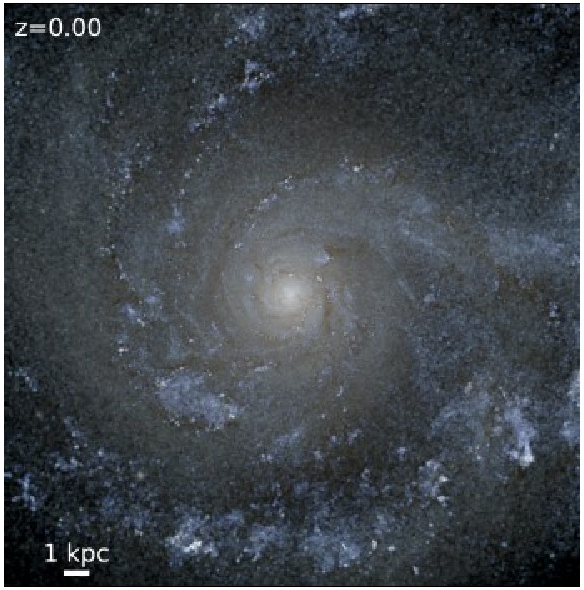
Galaxies

- *Stellar-to-halo mass ratio*
- *Star formation history*
- *Disk, bulge properties Surface density, Rotation curve*
- *Chemistery*
- ***Gas cycle, Star forming gas region properties***



Galaxies

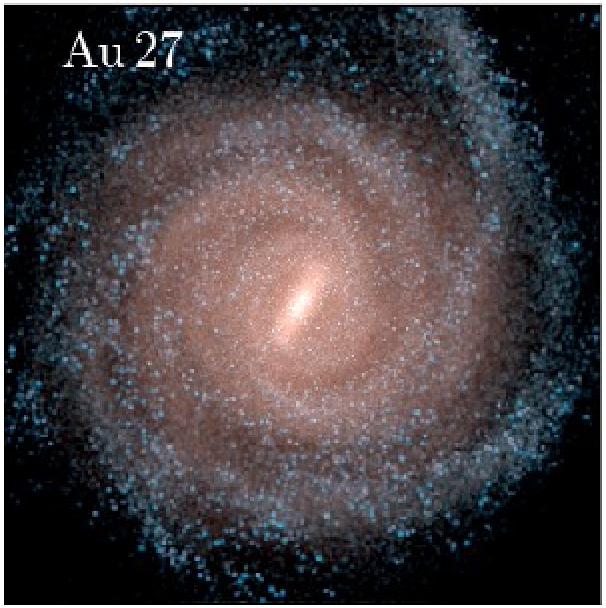
no Bar



FIRE

- Stellar-to-halo mass ratio
- Star formation history
- Disk, bulge properties Surface density, Rotation curve
- Chemistery
- Gas cycle, Star forming gas region properties
- **(no) Bars ?**

Bar ?



Auriga

Why ?
- Stabilization by the bulge or the halo (Debattista & Sellwood 2000; Kataria & Das 2017) ?
- Gas fraction/accretion (Kraljic et al. 2012) ?
...

Some bar effects:
Trigger star formation at its extremities (Renaud et al. 2015; Motte al. 2018),
Reduce star formation inside the bar (Longmore et al. 2013; Emsellem et al. 2015)
Fuel nuclear star formation in the very center where the gas accumulates
Affect the overall kinematics of the disk (resonances) Lynden-Bell & Kalnajs 1972).
...

Galaxies

- *Stellar-to-halo mass ratio*

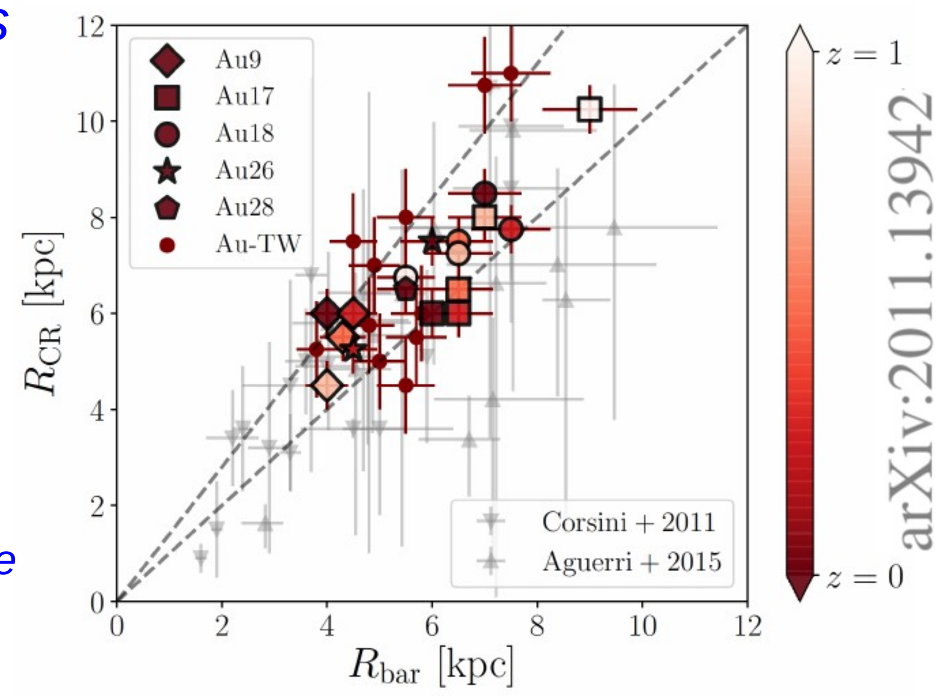
- *Star formation history*

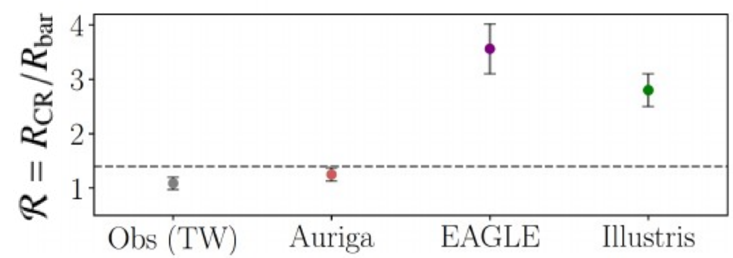
- *Disk, bulge properties Surface density, Rotation curve*

- *Chemistery*

- *Gas cycle, Star forming gas region properties*

- ***(no) Bars ?***



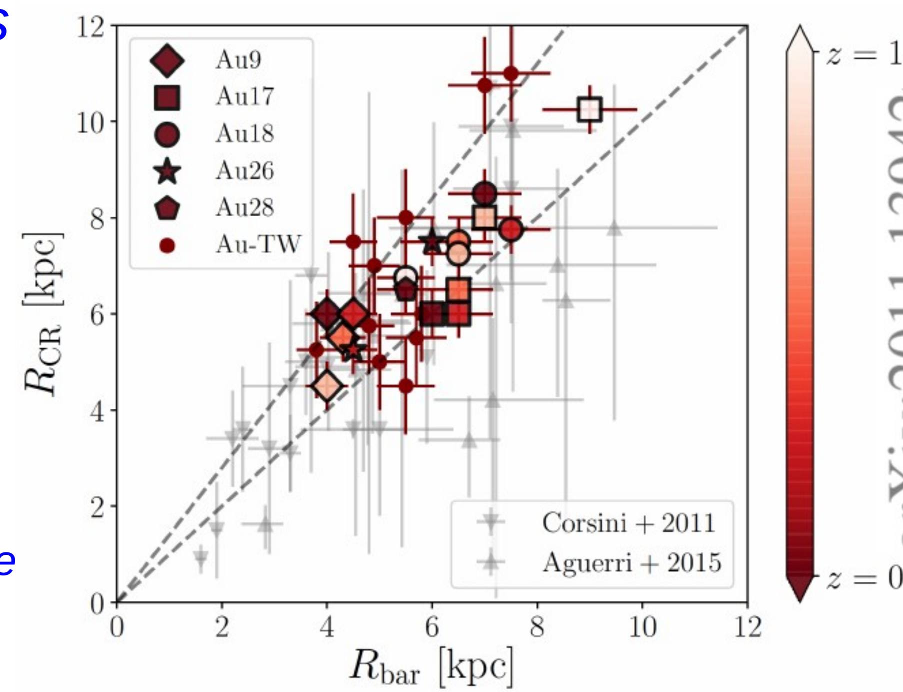


Slow

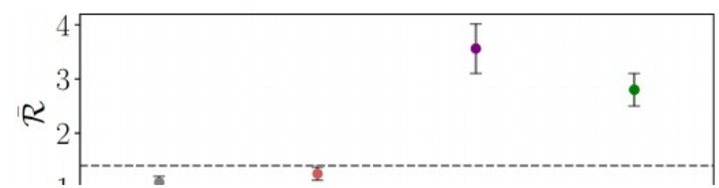
Fast

Auriga

Galaxies



- Stellar-to-halo mass ratio
- Star formation history
- Disk, bulge properties Surface density, Rotation curve
- Chemistery
- Gas cycle, Star forming gas region properties
- ***(no) Bars ?***

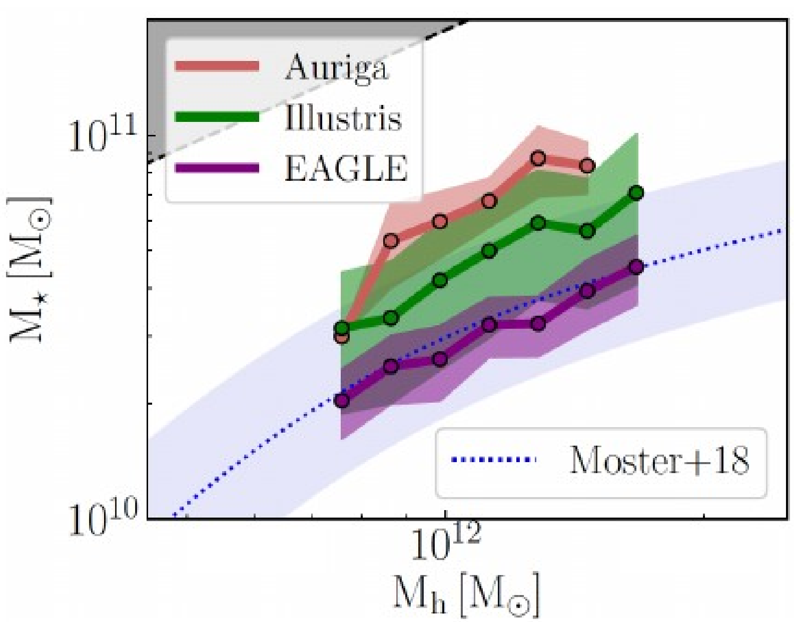


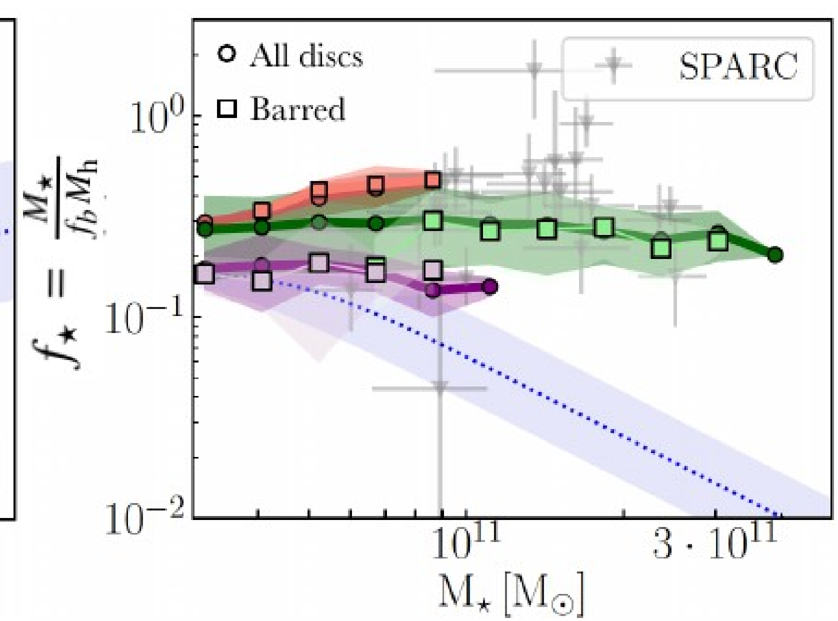
Fast bars at the price

of (too ?) high

stellar-to-halo

mass ratio





Dark matter

Distribution features feed DM detection calculations/prospects

Dark matter

Distribution features feed DM detection calculations/prospects

- *Mass density profiles*

- *Halo shape*

- *Phase-space/velocity distributions*

- *Substructures*

Dark matter

- ***Mass density profiles***
- *Halo shape*
- *Phase-space/velocity distributions*
- *Substructures*

Contraction with baryon ?
(Blumenthal 1986)
Angular momentum and mass conservation

$$M_i(r_i)r_i = [M_b(r_f) + M_{DM}(r_f)]r_f$$

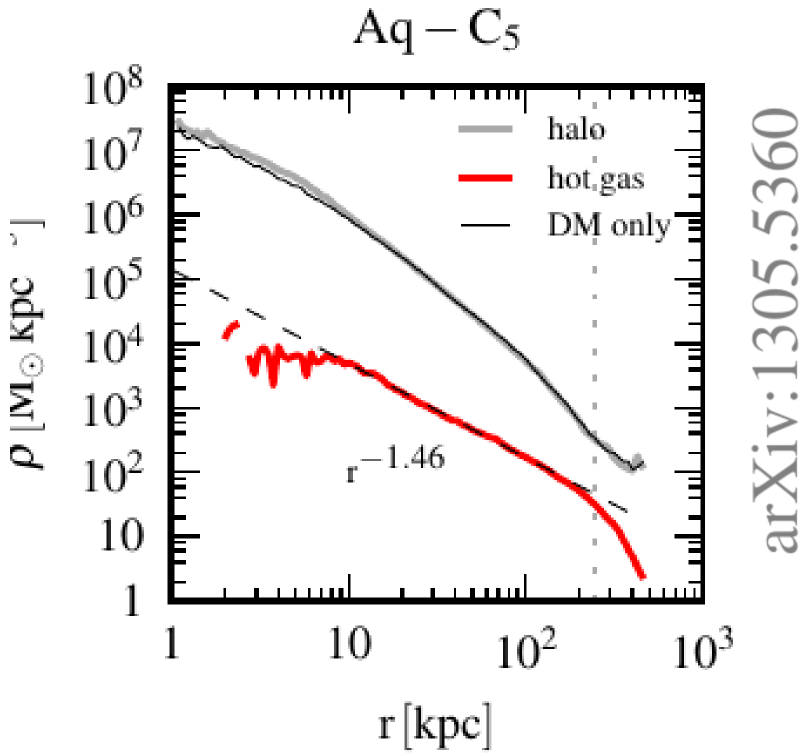
Steep cusp ?

Feedback induced core ? SN (AGN ?)
(Pontzen&Governato 2012-14)

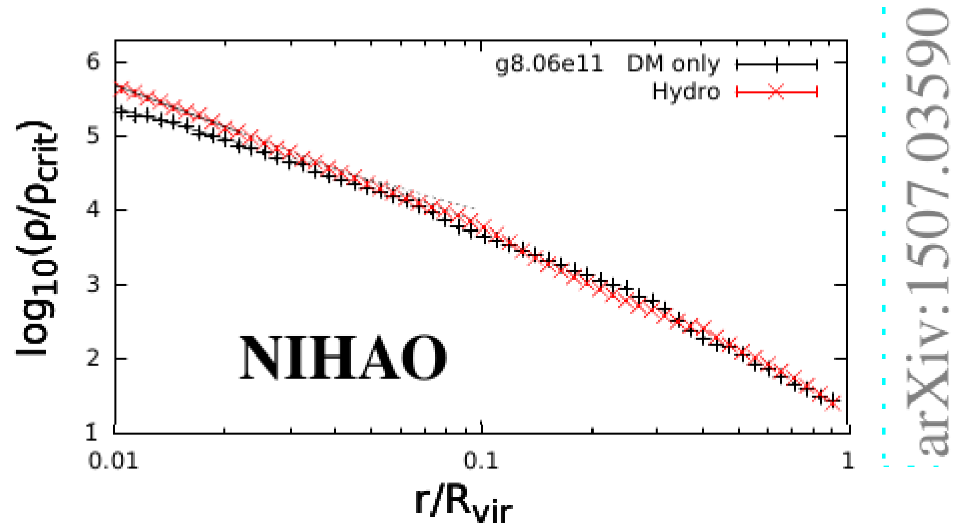
Relevant for indirect detection

$$\Phi_i \propto \int \rho^2_{DM}(r)dV$$

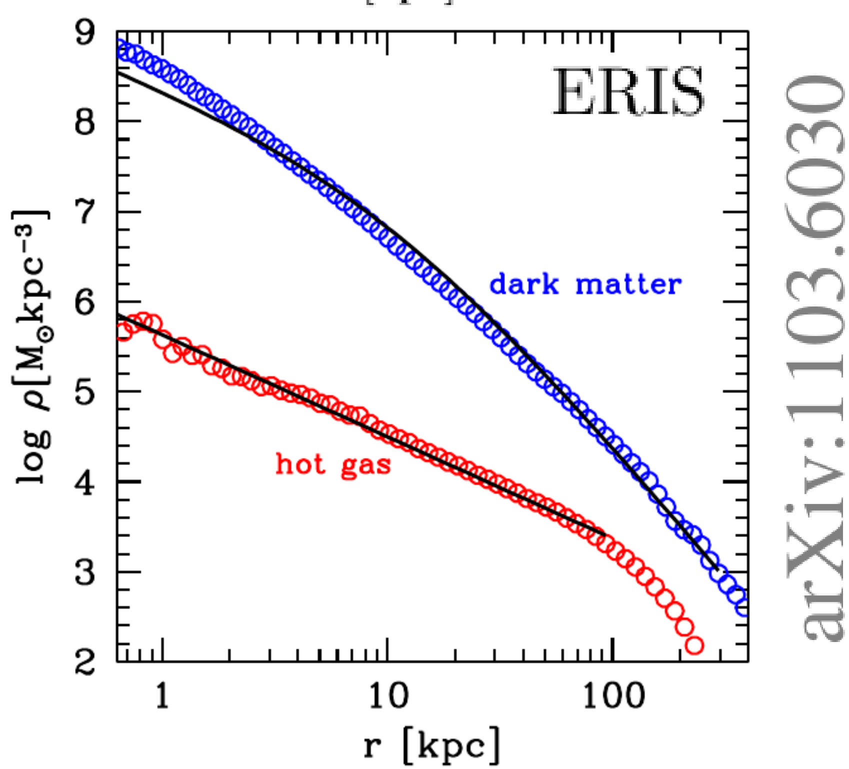
Dark matter



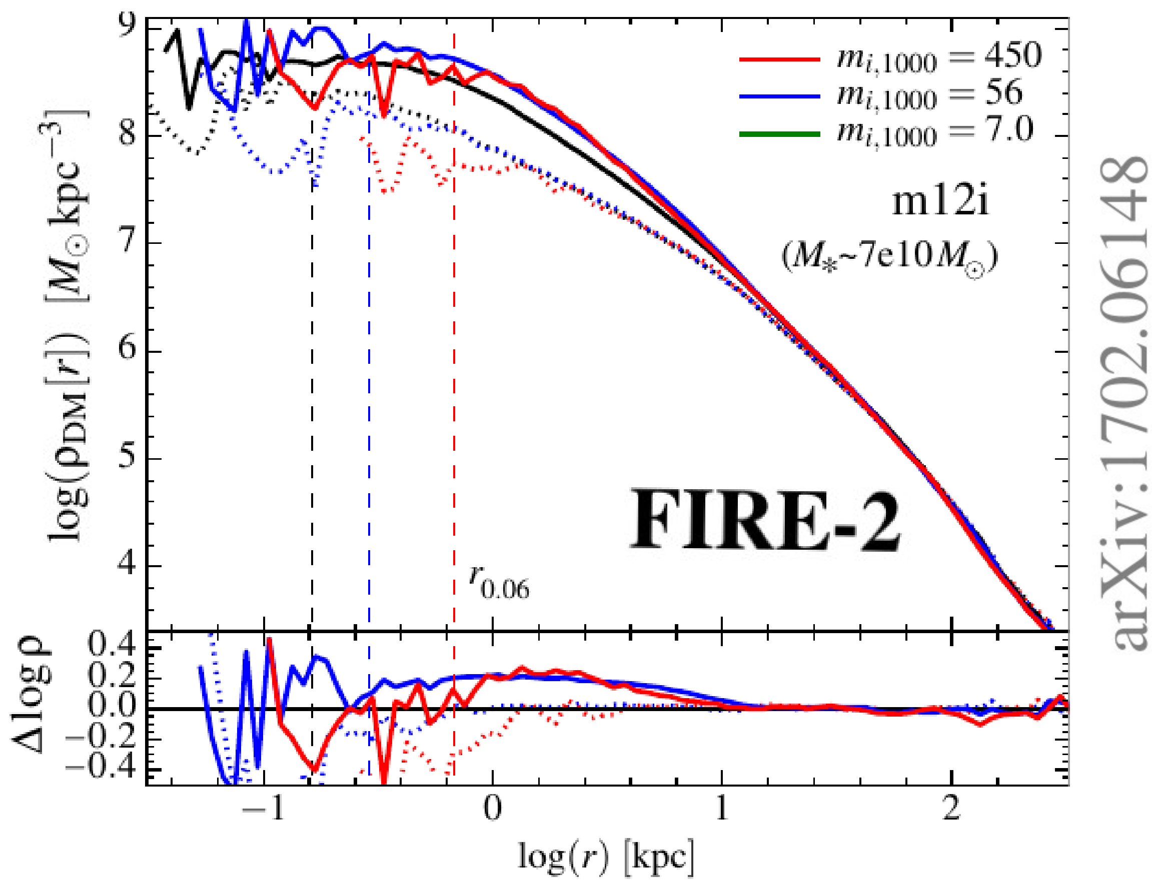
Aq – C5
halo
hot gas
DM only
$r^{-1.46}$
ρ [M⊙ kpc⁻³]
r [kpc]
arXiv:1305.5360



g8.06e11 DM only
Hydro
NIHAO
$\log_{10}(\rho/\rho_{crit})$
r/R_{vir}
arXiv:1507.03590



ERIS
dark matter
hot gas
log ρ[M⊙kpc⁻³]
r [kpc]
arXiv:1103.6030



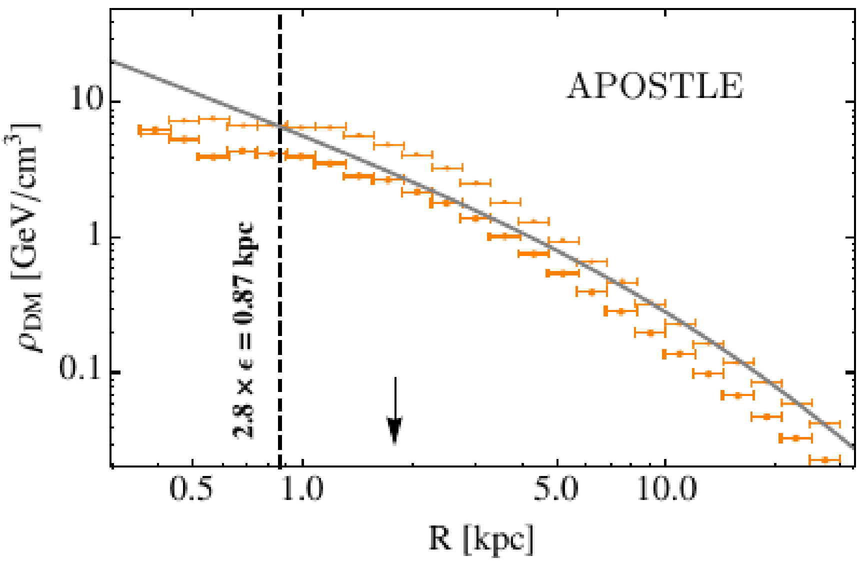
$m_{i,1000} = 450$
$m_{i,1000} = 56$
$m_{i,1000} = 7.0$
m12i
($M_* \sim 7e10 M_\odot$)
FIRE-2
$r_{0.06}$
$\log(\rho_{DM}[r])$ [M⊙ kpc⁻³]
Δlog ρ
log(r) [kpc]
arXiv:1702.06148

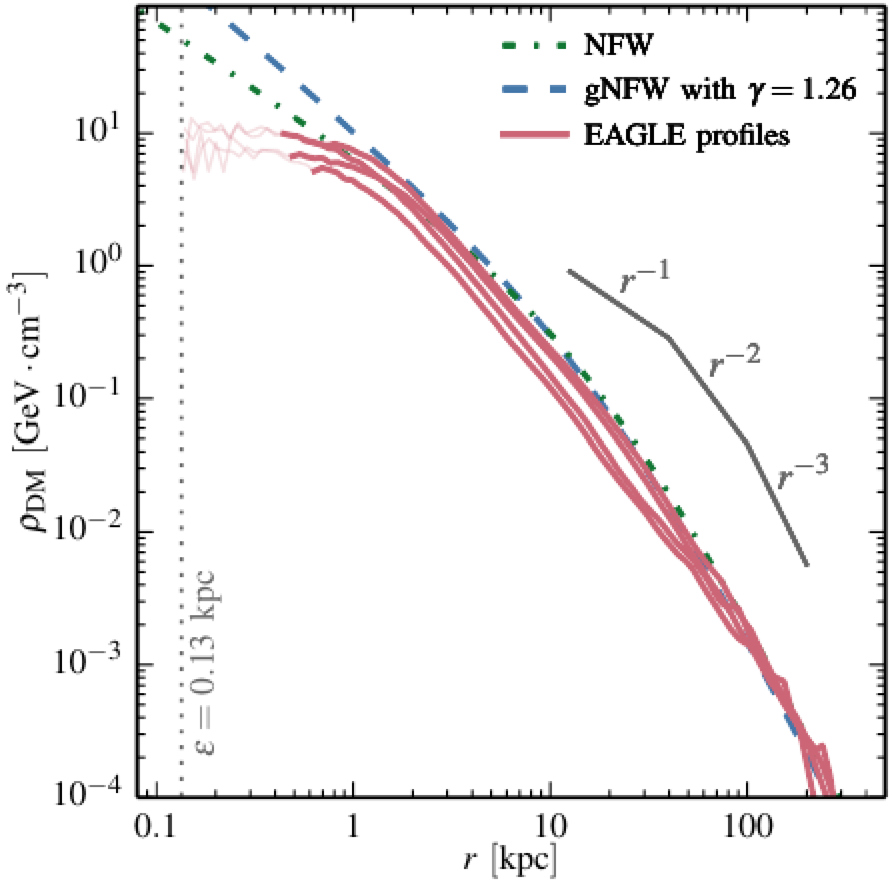
Dark matter

- ***Mass density profiles***

- *Contraction (+ flattening ?)*

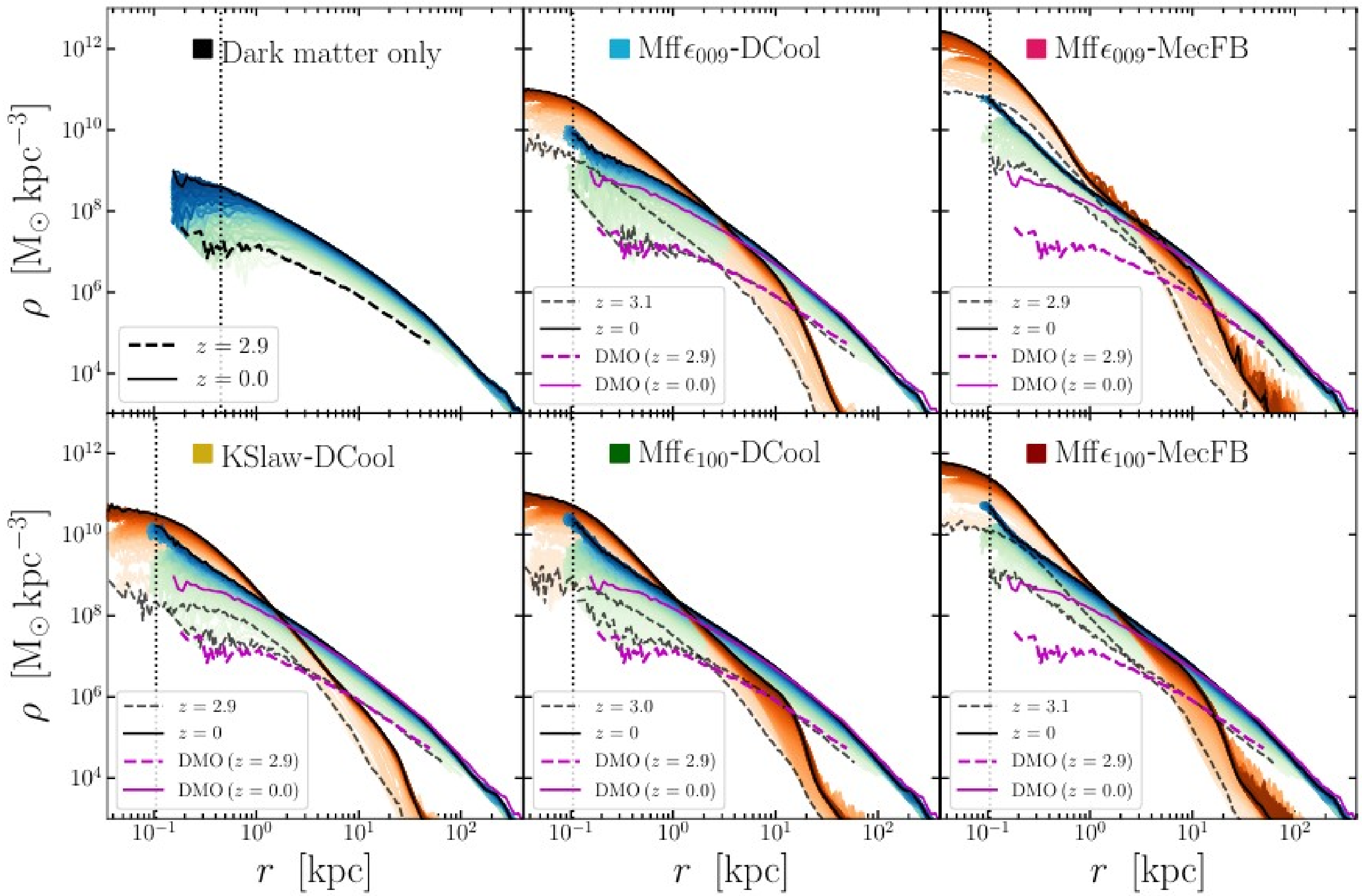
arXiv:1509.02164





arXiv:1509.02166

Dark matter

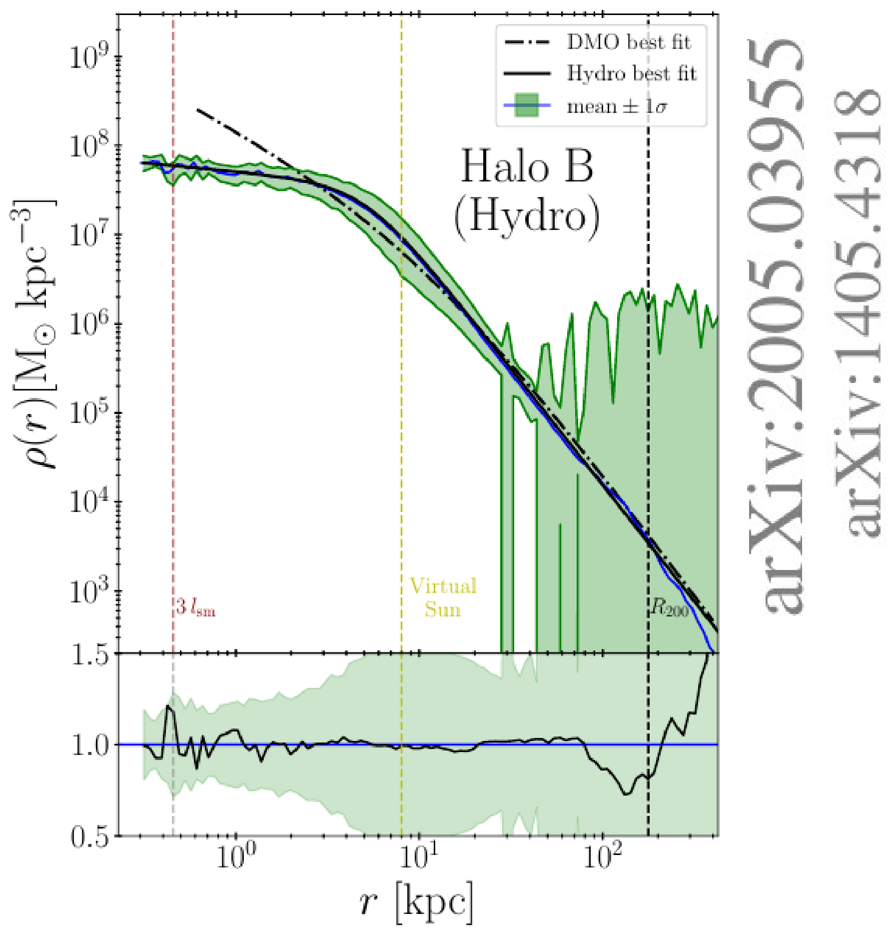


Mochima arXiv:2301.06189

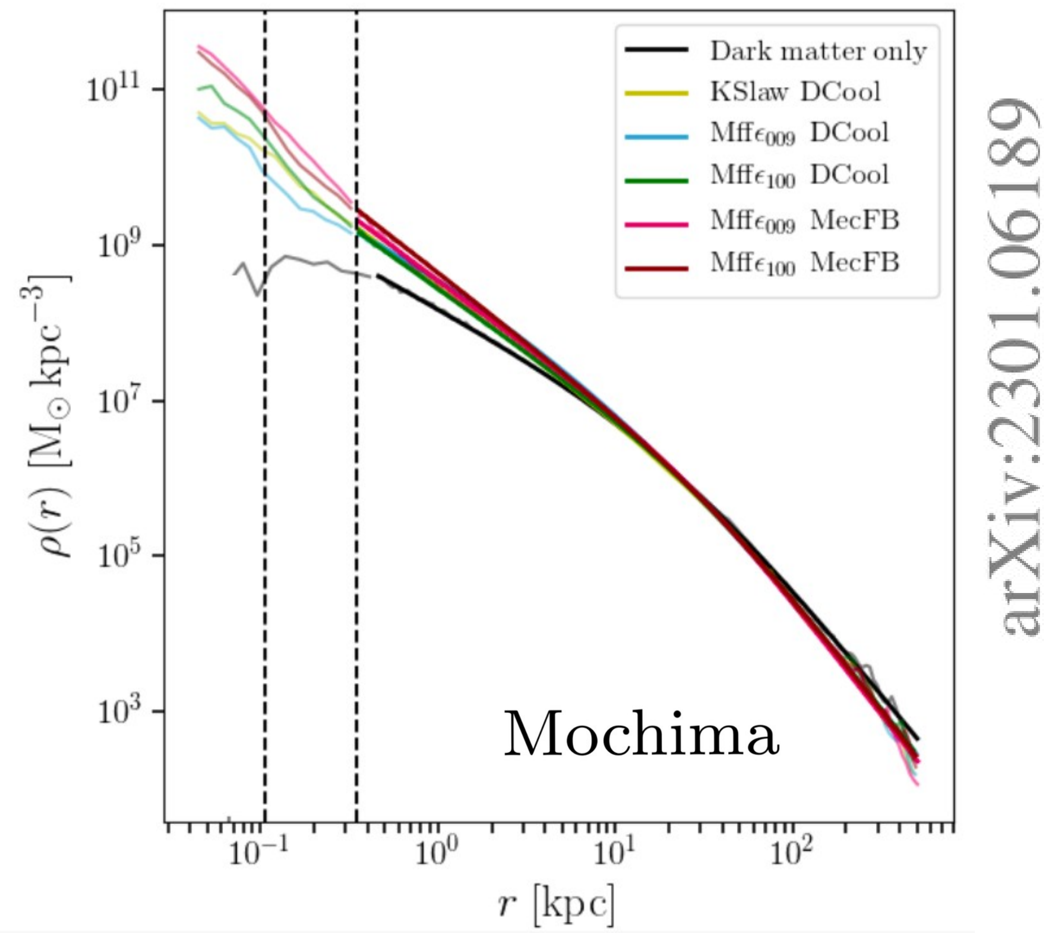
Response of DM halo driven by the history of assembly of baryons (e.g Pedrosa et al 2009)

→ DM profile depends on baryonic physics. SF and feedback reciepes (model, parameters, resolution ...)

Dark matter



Strong SN feedback

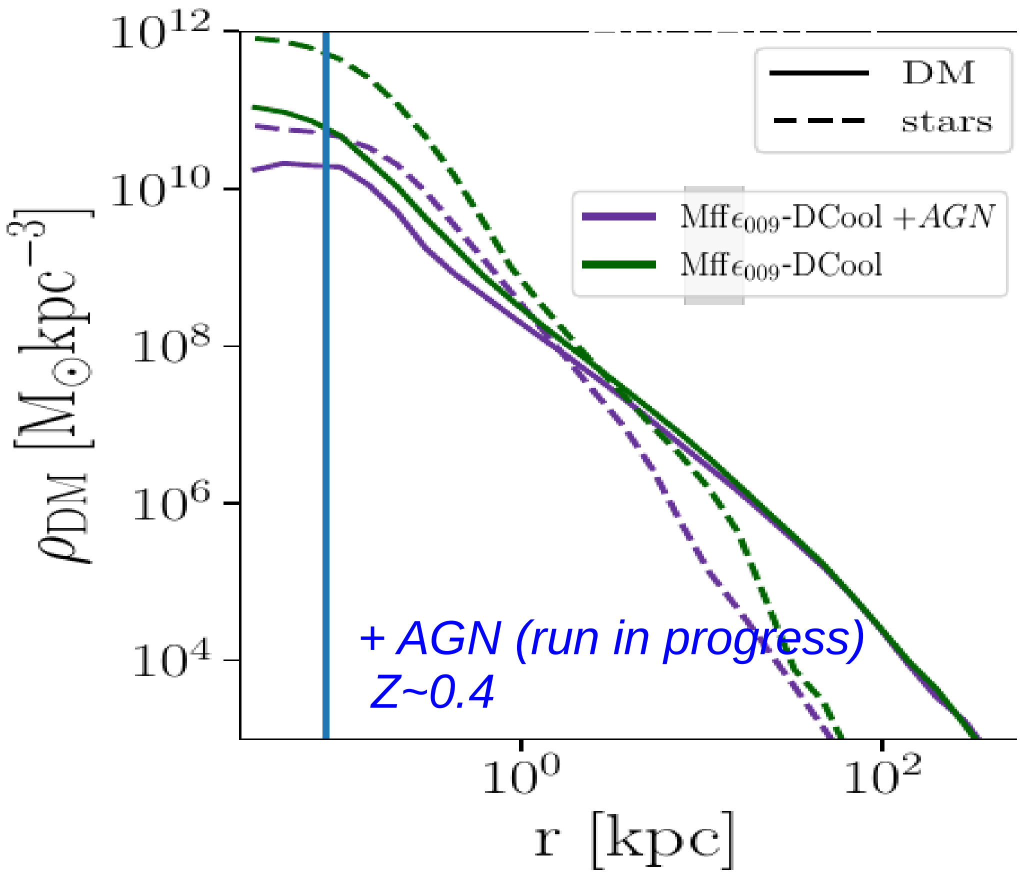


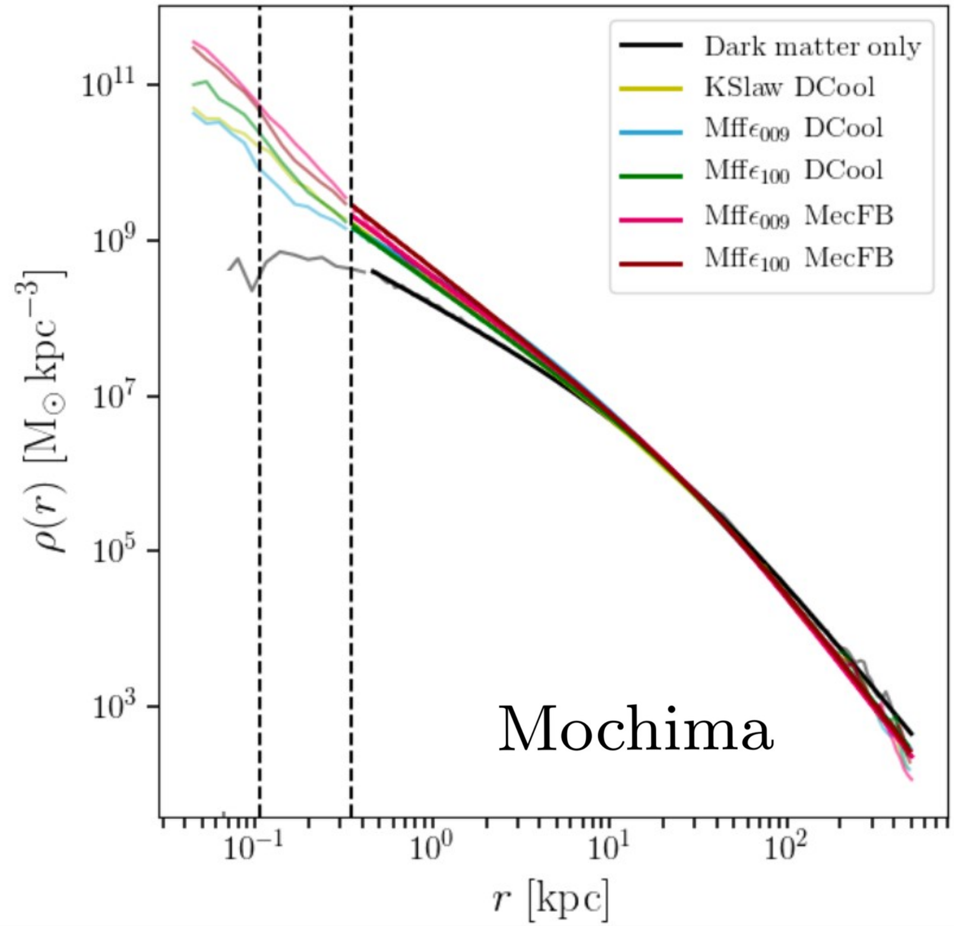
Response of DM halo driven by the history of assembly of baryons
(e.g Pedrosa et al 2009)

→ DM profile depends on baryonic physics. SF and feedback reciepes (model, parameters, resolution ...)

NFW ? Einasto ?

Dark matter





arXiv:2301.06189

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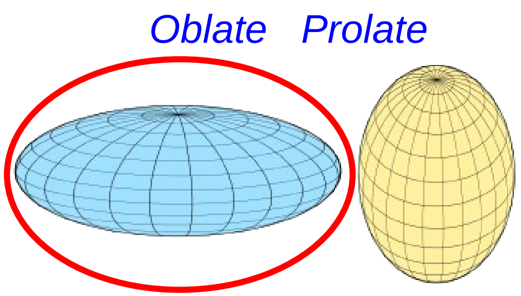
NFW ? Einasto ?

Dark matter

- Mass density profiles

- Halo shape

Dark matter

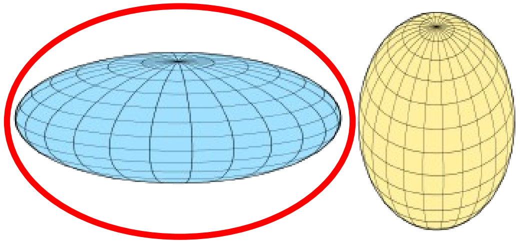


- Mass density profiles

- Halo shape

(Uncertain) observations suggest slightly oblate halo in the center and become triaxial at large distances (Law and Majewski 2010, Ibata et al 2013, Vera-Ciro and Helmi 2013)

Dark matter



- *Mass density profiles*

- ***Halo shape: rounder halo than DMO***

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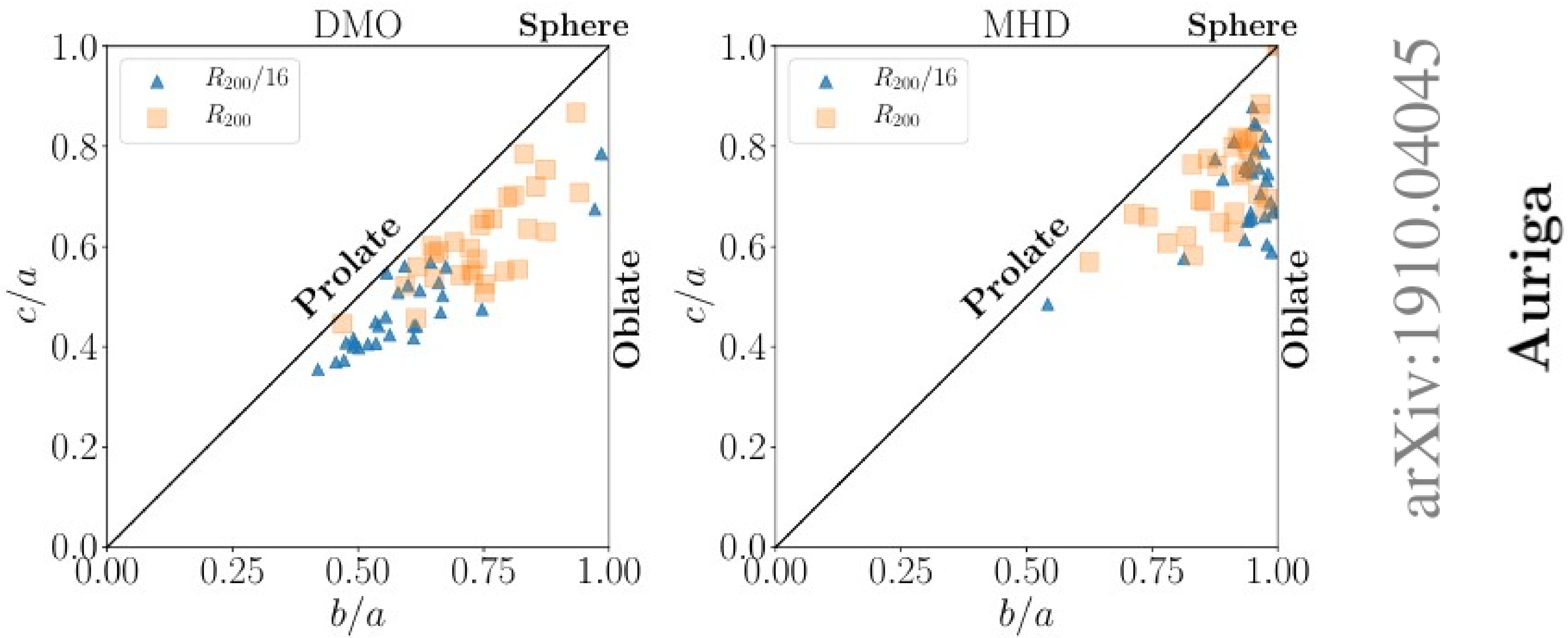
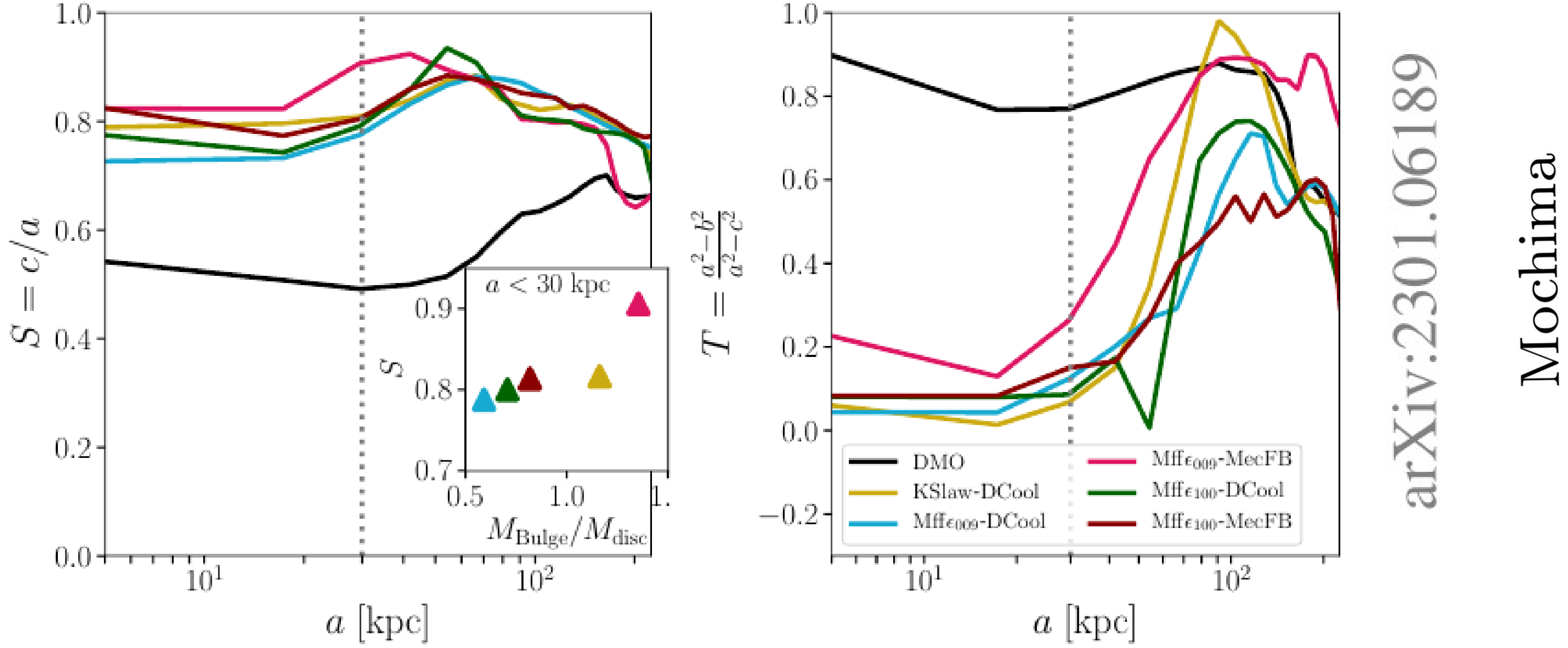


Dark matter

- *Mass density profiles*

- ***Halo shape: rounder halo than DMO***



Same halo, varying baryonic physics. Results might change with weaker bulge, bar ...

Dark matter

- Mass density profiles

- Halo shape

- Phase-space/velocity distributions (complex/realistic ?)

Accretion history → Distribution features
beyond analytical functions ?
Dark disc ?

Fit ? Maxwellian, Tsallis … ? SHM ?

Agreement with analytical predictions ?
(e.g Eddington inversion)

Relevant for (in)direct detection,
capture in celestial bodies

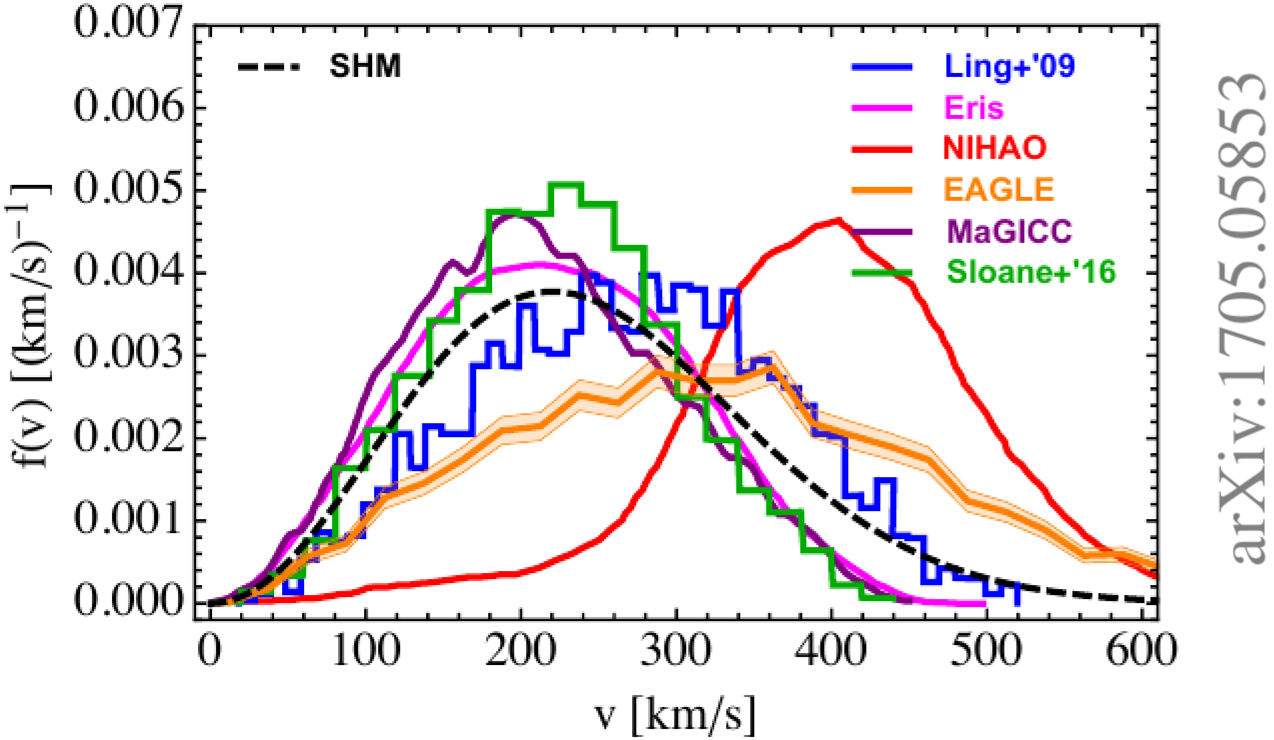
$$\frac{d\mathcal{R}}{dE_R} \propto \int_{v_{min}}^{v_{esc}} d^3\vec{v}\, \frac{\mathrm{f}(\vec{v}(t))}{v}$$

Dark matter

- *Mass density profiles*

- *Halo shape*

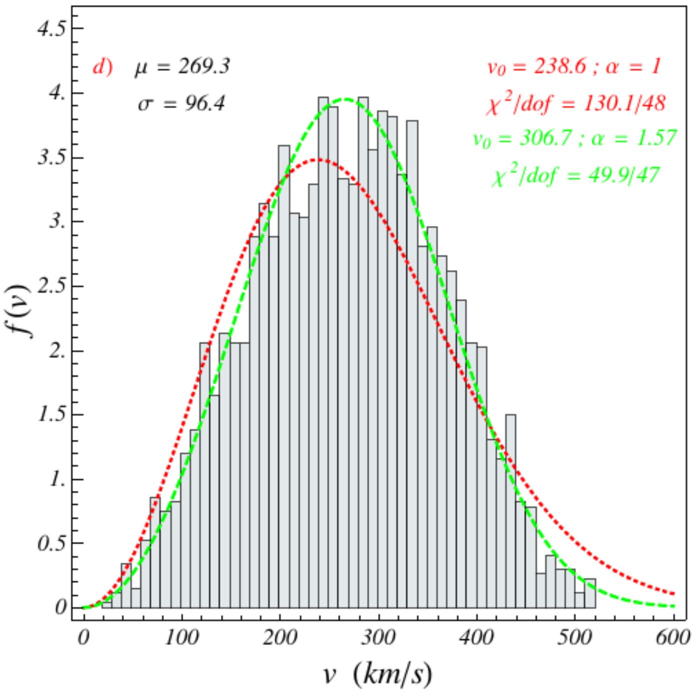
- ***Phase-space/velocity distributions (complex/realistic ?)***



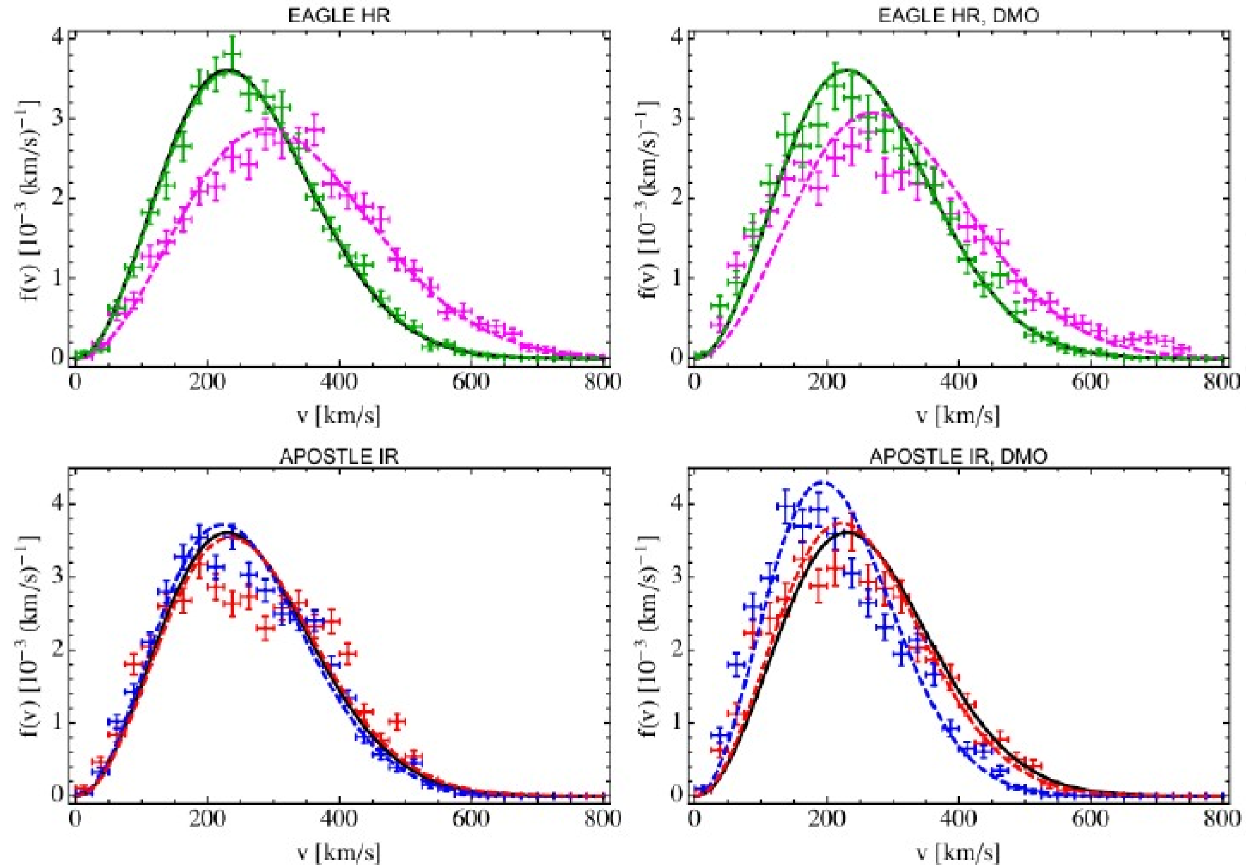
Methods (meaning !) of particle selections ?

Dark matter

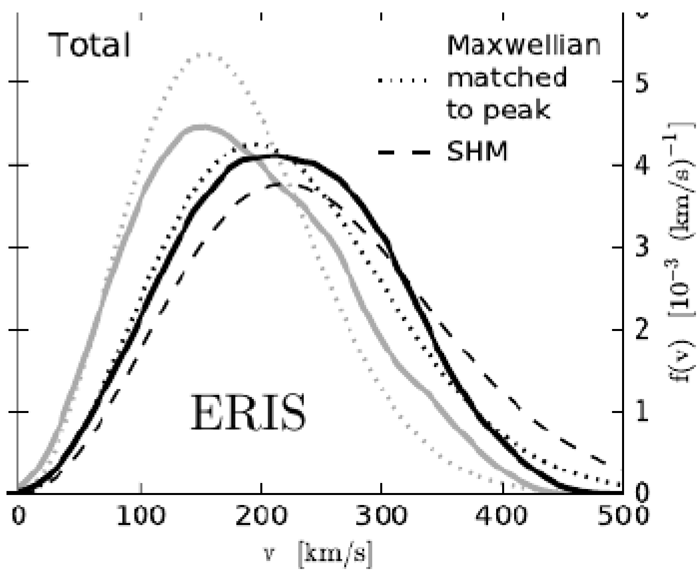
arXiv:0909.2028



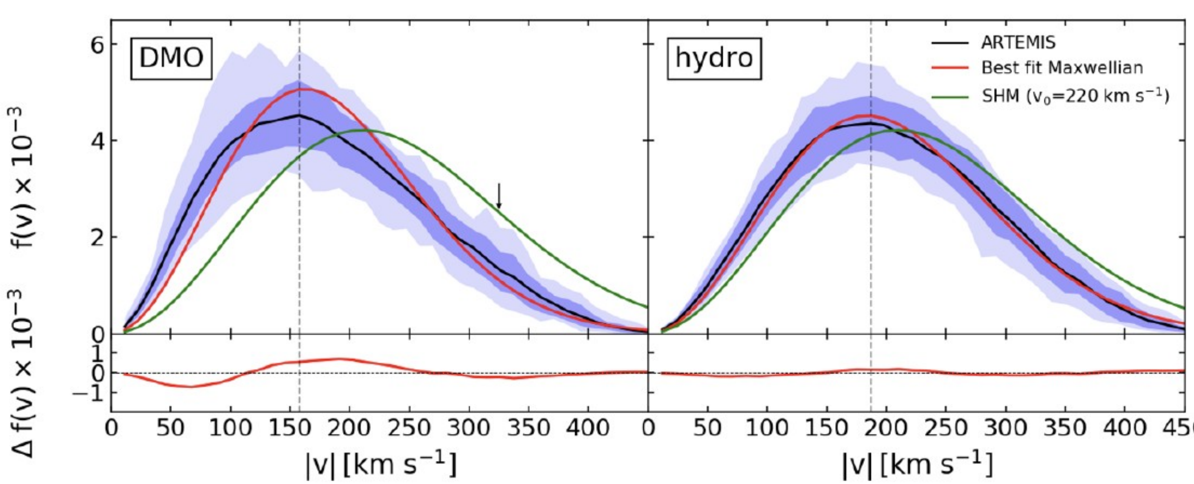
arXiv:1601.04707



arXiv:1308.1703

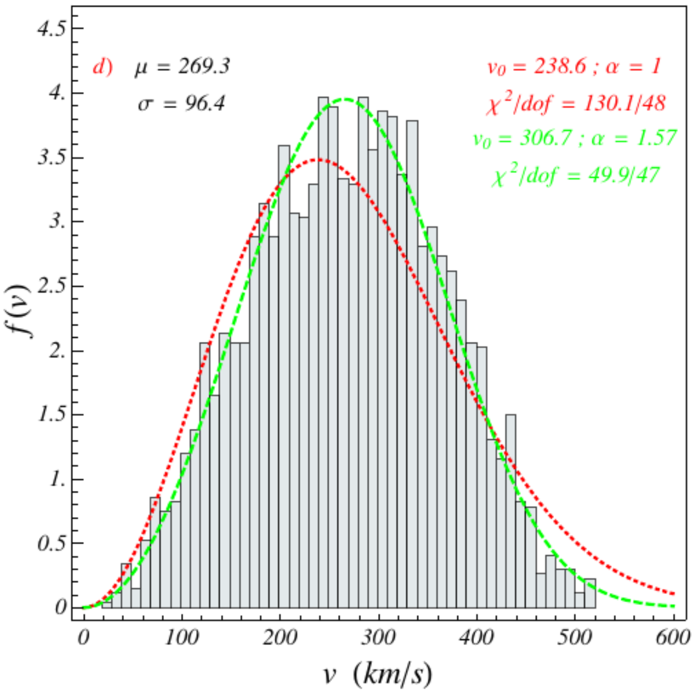


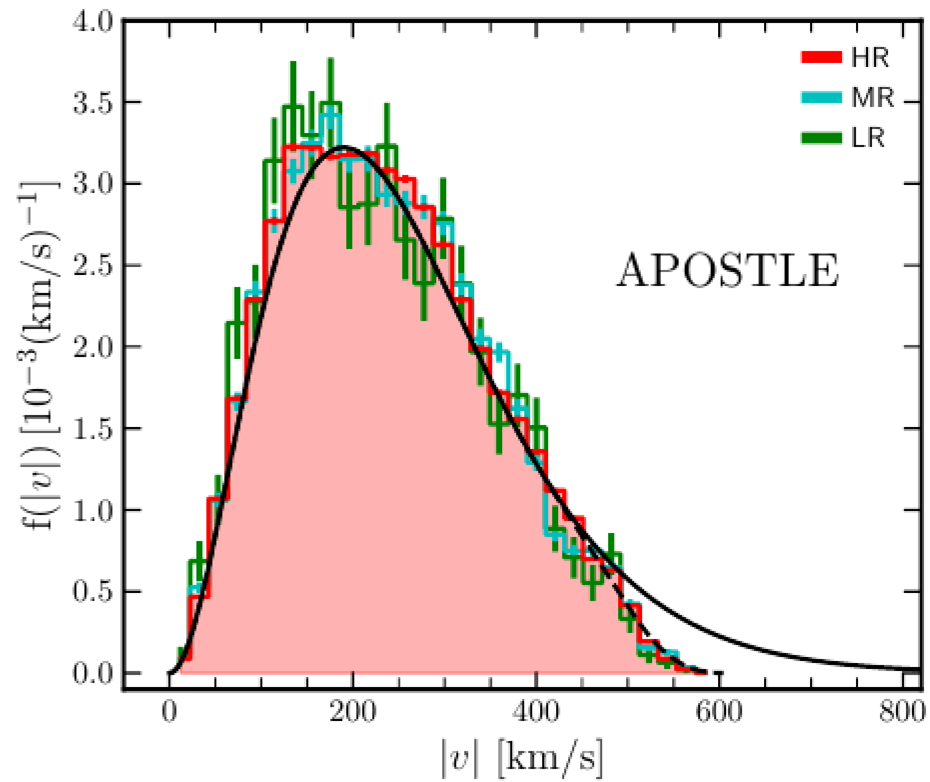
arXiv:2006.15159



Dark matter

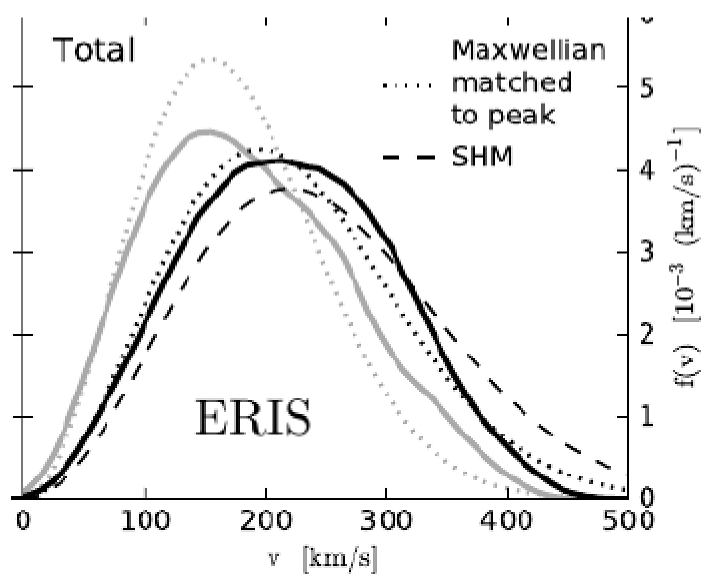
arXiv:0909.2028

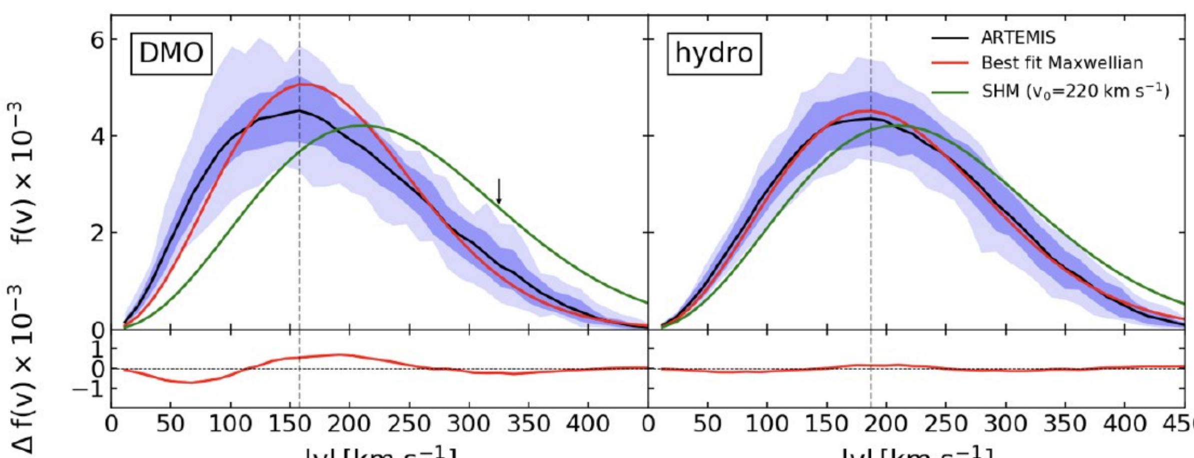




arXiv:2308.15388

arXiv:1308.1703





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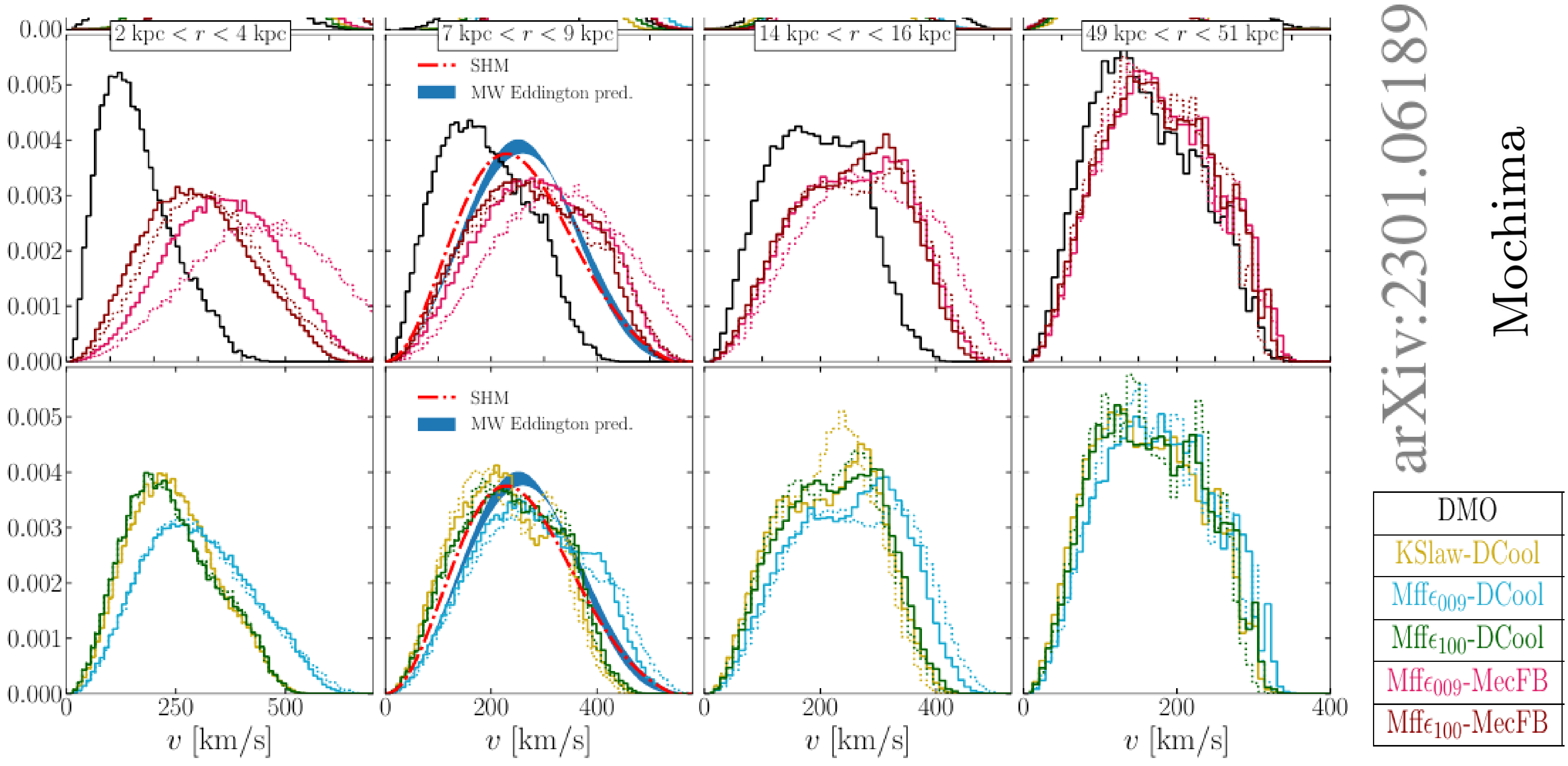
Dark matter



Features from formation history

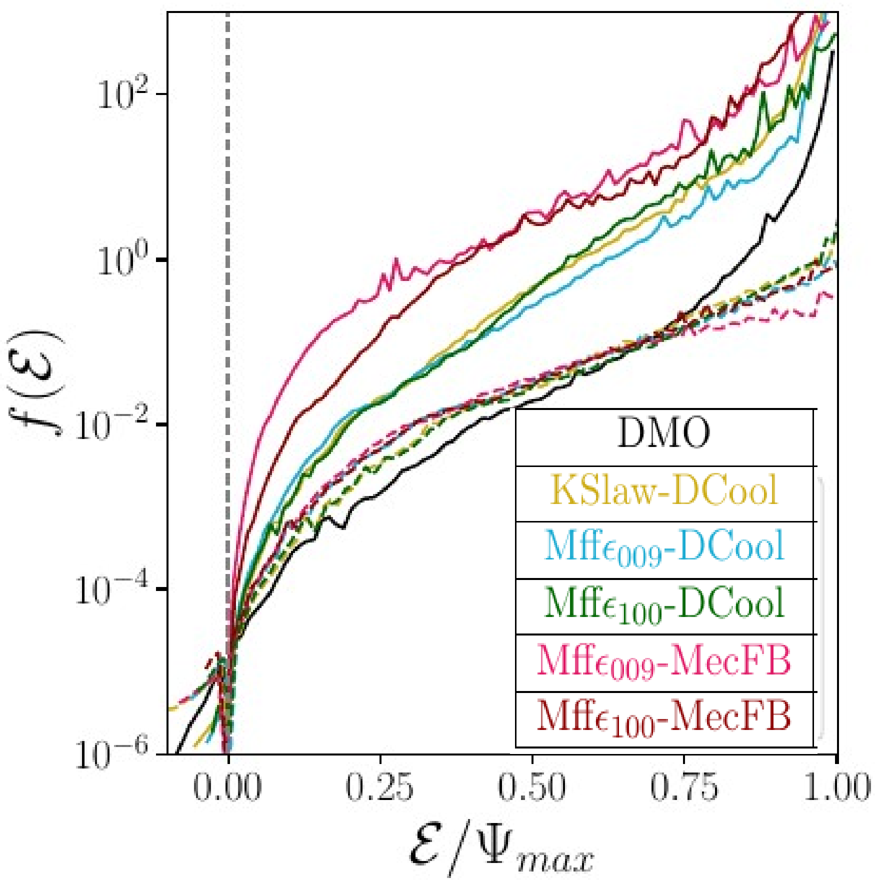
Baryons : shift central value and broader distributions in central part.

Dark matter

- *Mass density profiles*

- *Halo shape*

- ***Phase-space/velocity distributions (complex/realistic ?)***



arXiv:2301.06189

Mochima

Pseudo-phase space distribution

Higher energy particles due to baryonic potential in hydro runs.

all particles (solid line)

excluding central particles R<3 kpc (dashed line)

Dark matter

- *Mass density profiles*
- *Halo shape*
- *Phase-space/velocity distributions*
- ***Substructures***

Mass spectrum modified by baryons (tidal effects, disc, concentration …) ?

$$\frac{dN_{cl}}{dM} \propto \left(\frac{M}{M_H}\right)^n$$

$$n \sim -1.8 - 2$$

Relevant for detection
(Boost factor, local DM distribution)

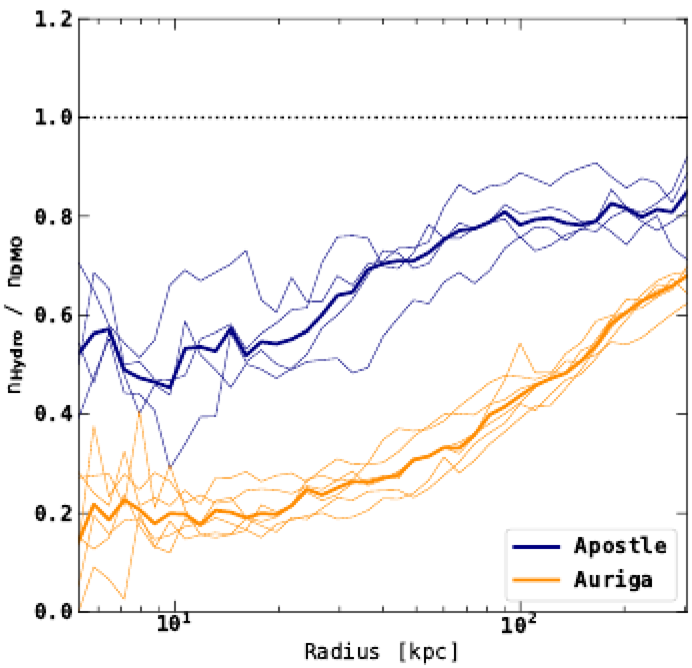
Dark matter

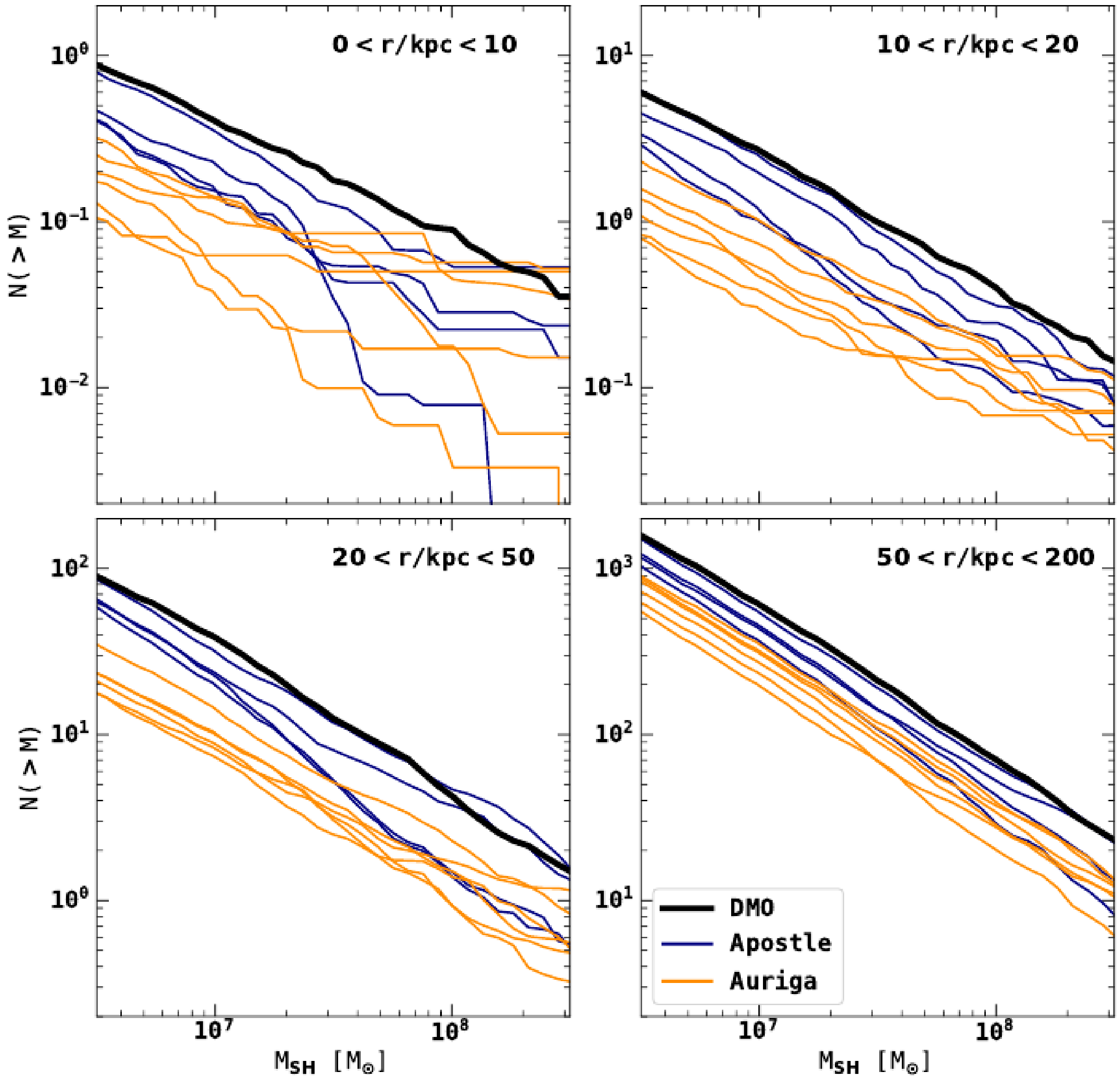
- *Mass density profiles*

- *Halo shape*

- *Phase-space/velocity distribu*

- ***Substructures***





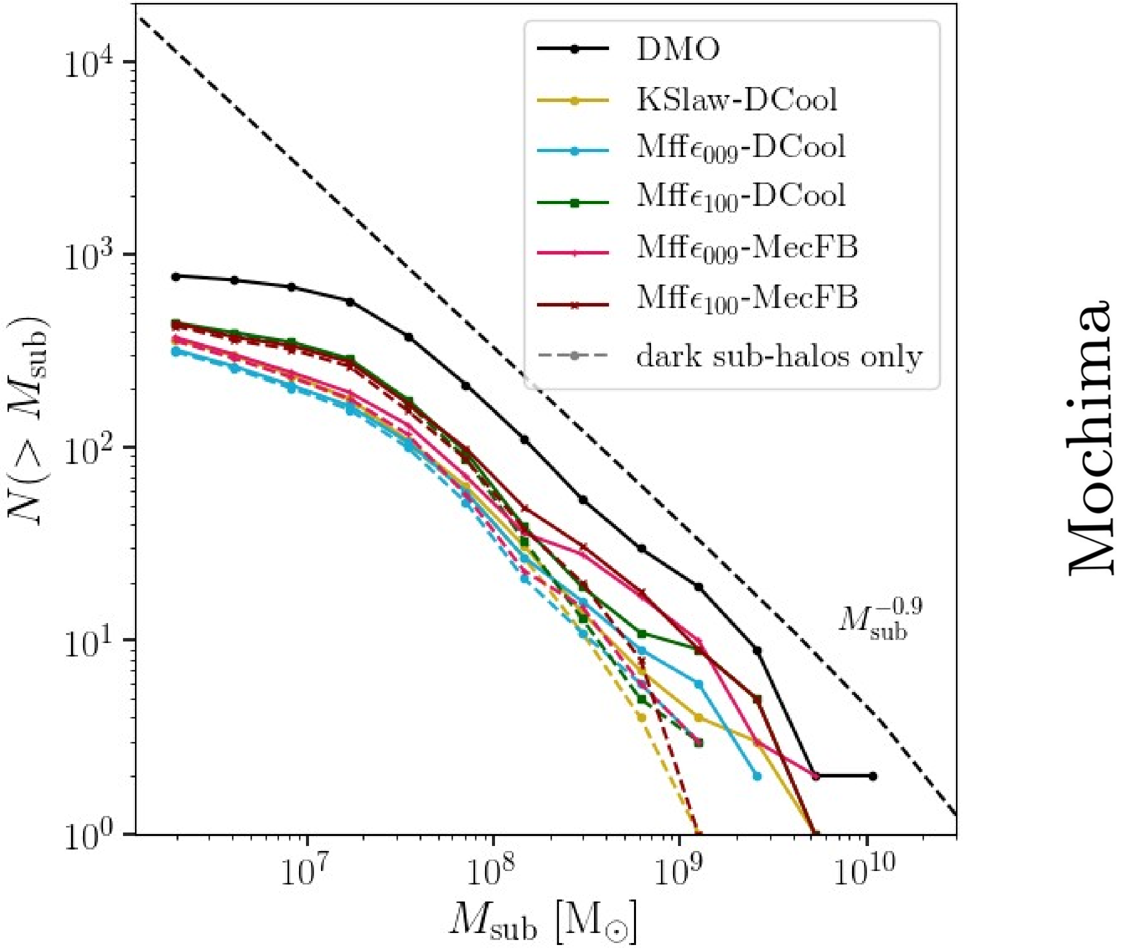
APOSTLE and AURIGA

Relative to DMO simulation, the abundance of subhalos is reduced.
APOSTLE: by 50% near the centre and by 10% within r200 .
AURIGA: 80% and 40%

Dark matter

- *Mass density profiles*
- *Halo shape*
- *Phase-space/velocity distributions*
- ***Substructures***

Reduced abundance of subhaloes

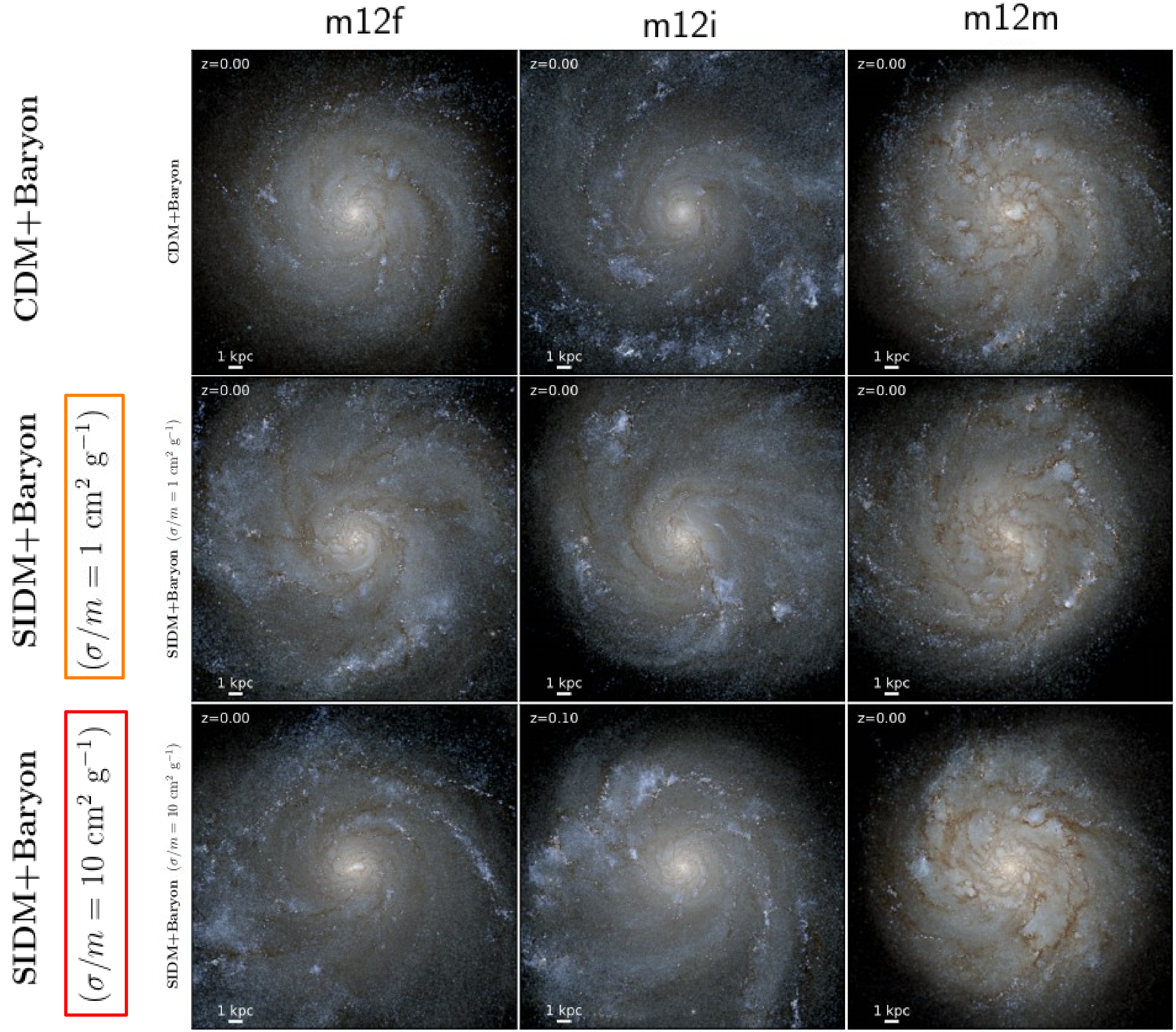


Nunez-Castineyra, EN et al in prep

Beyond CDM ?

Self-Interacting DM

Self-Interacting DM



arXiv:2104.14069

FIRE

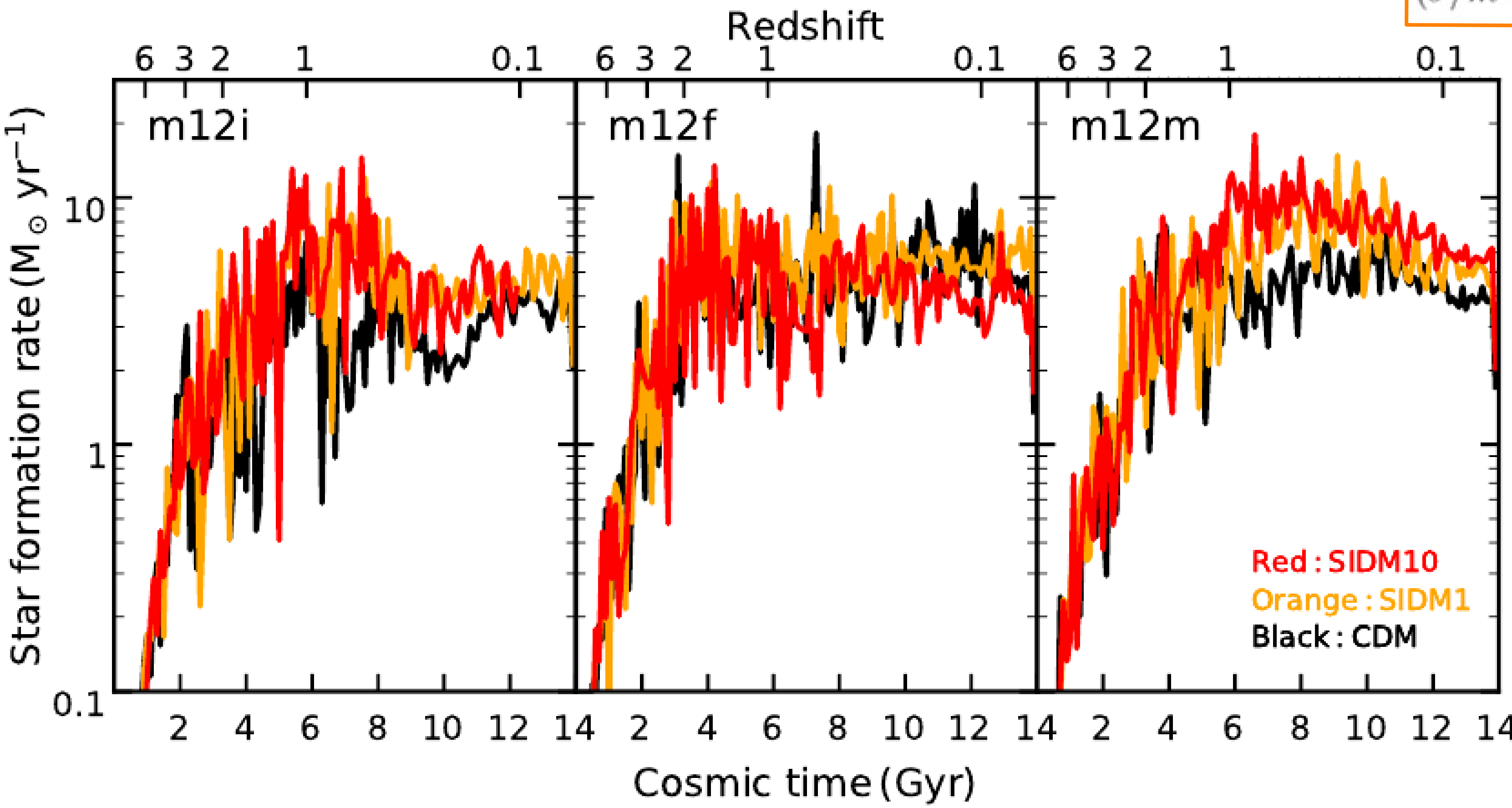
Similar galaxies

Self-Interacting DM

$(\sigma/m = 10\ \mathrm{cm}^2\ \mathrm{g}^{-1})$

$(\sigma/m = 1\ \mathrm{cm}^2\ \mathrm{g}^{-1})$

- *Star formation rate*



arXiv:2102.12480

FIRE

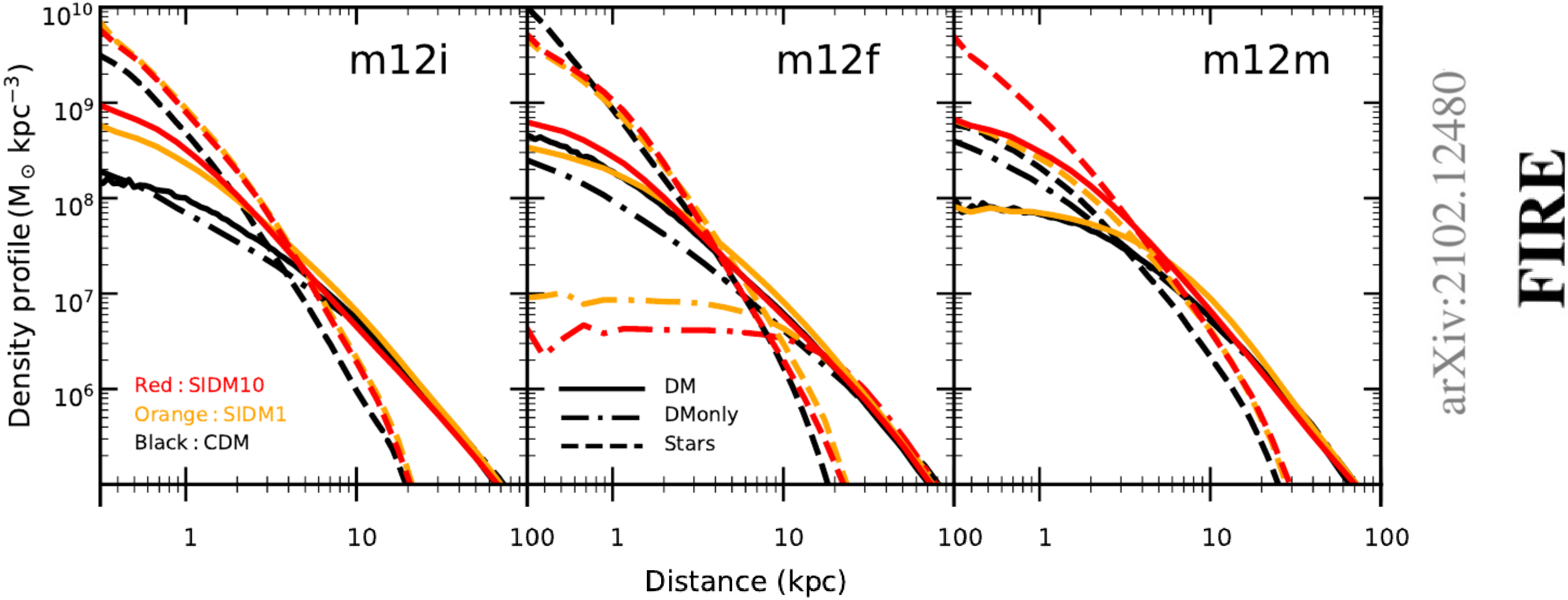
Higher SFR

SIDM

- *Mass density profiles*

$(\sigma/m = 10\ \mathrm{cm^2\ g^{-1}})$

$(\sigma/m = 1\ \mathrm{cm^2\ g^{-1}})$



Strong(er) stellar cusp than CDM

SIDM profile responds more significantly to presence/contraction by baryons than CDM

SIDM	$\bar{V}_{\mathrm{2kpc,\,DMO}}/\bar{V}_{\mathrm{2kpc,\,Hydro.}} \sim 0.10$	*Strong cusp*
CDM	$\bar{V}_{\mathrm{2kpc,\,DMO}}/\bar{V}_{\mathrm{2kpc,\,Hydro.}} \sim 0.25\text{–}0.35$	*Contraction + flattening*

FuzzyDM
(No Hydro)

FuzzyDM
(No Hydro)

Name	$m_{\rm p}$ [$\rm M_\odot$]	ϵ [pc]	$N_{\rm hr}$	$N_{\rm lr}$	M_{200} [$\rm M_\odot$]
Aq-A-1	1.712×10^3	20.5	4,252,607,000	144,979,154	1.839×10^{12}
Aq-A-2	1.370×10^4	65.8	531,570,000	75,296,170	1.842×10^{12}
Aq-A-3	4.911×10^4	120.5	148,285,000	20,035,279	1.836×10^{12}
Aq-A-4	3.929×10^5	342.5	18,535,972	634,793	1.838×10^{12}
Aq-A-5	3.143×10^6	684.9	2,316,893	634,793	1.853×10^{12}

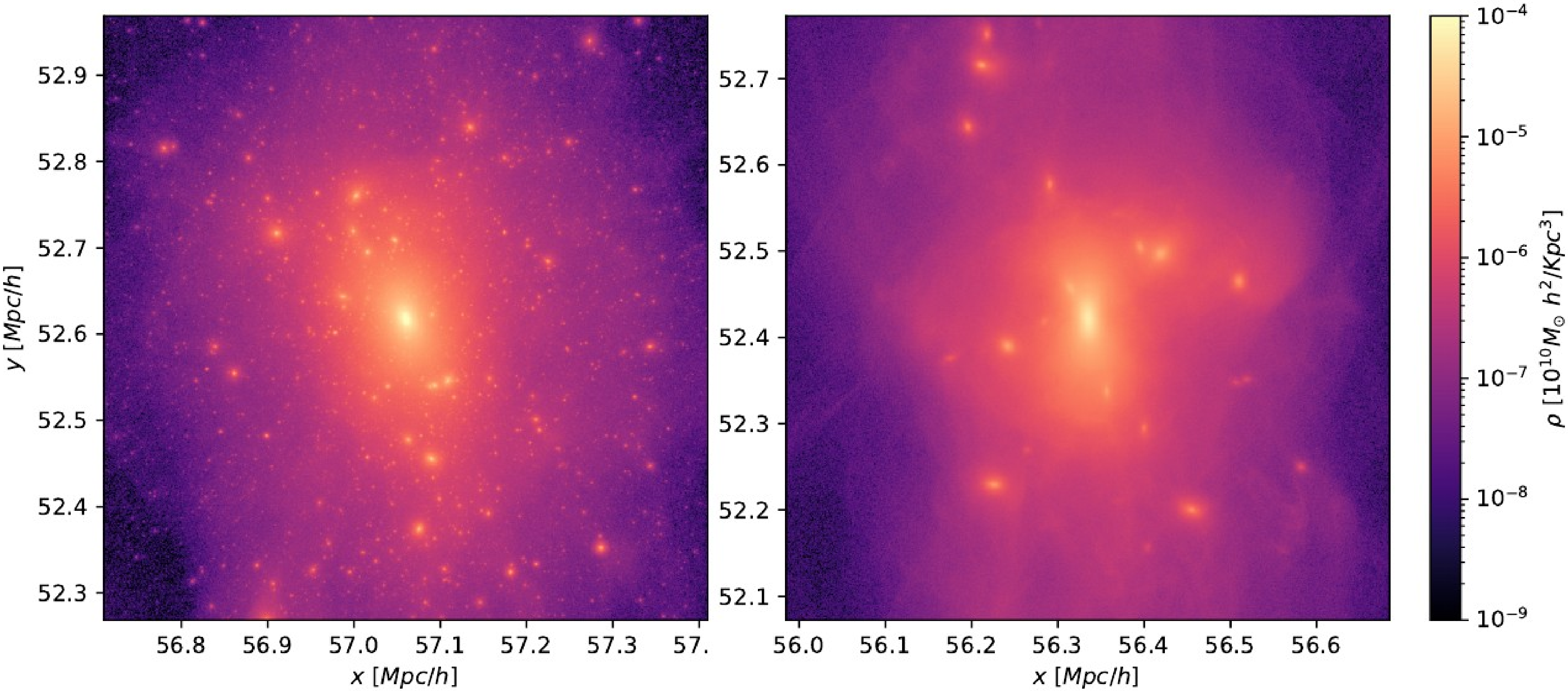
arXiv:2210.08022

Fuzzy Aquarius

FuzzyDM
(No Hydro)

- Density maps

CDM | FuzzyDM



Suppress small mass objects

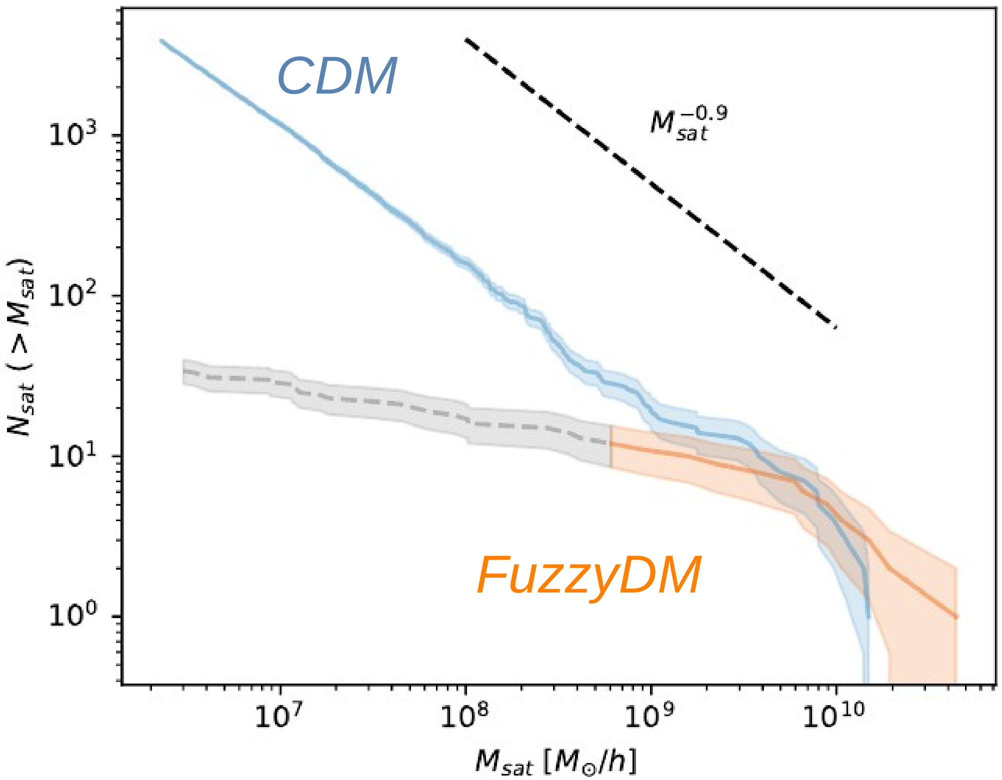
Cored profiles

arXiv:2210.08022

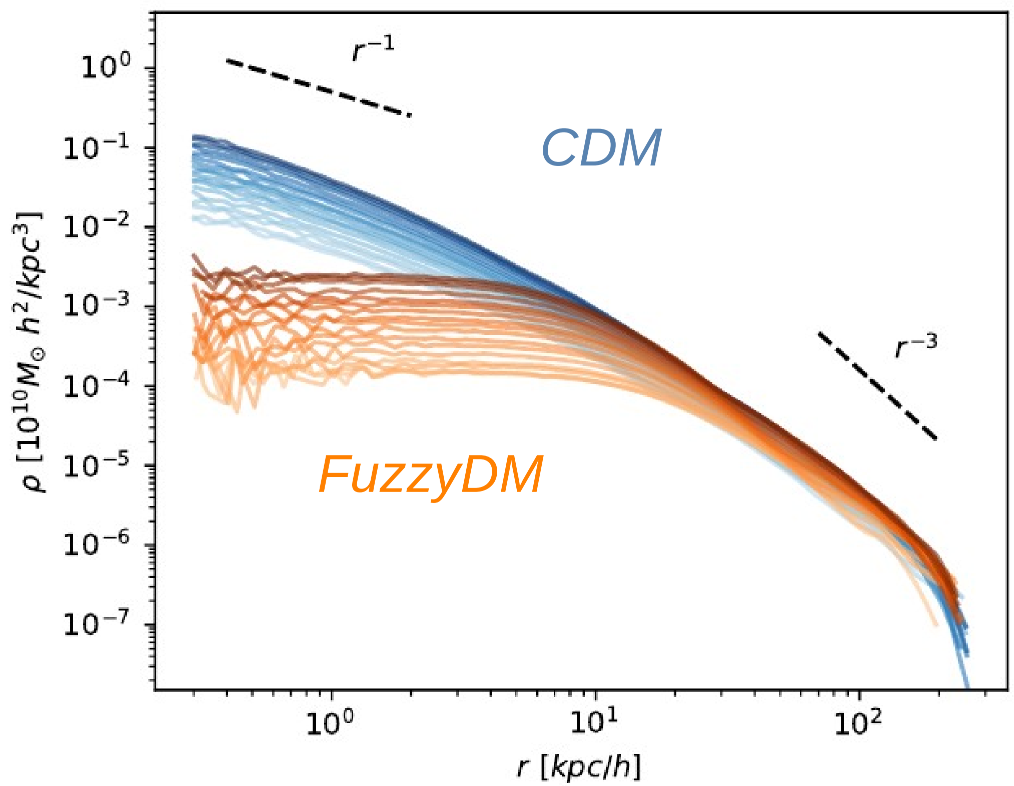
Fuzzy Aquarius

FuzzyDM
(No Hydro)

Subhalo mass functions



Mass density profiles



cosmology. Full hydro simulations will be needed to probe the effects of the different dynamical evolution of the stellar content of the satellites, since dark matter and stars react differently to stellar stripping (Peñarrubia et al. 2008; Macciò et al. 2021).

arXiv:2210.08022

Fuzzy Aquarius

Status

- Different baryonic physics change the resulting galaxy and the DM distribution in the halo

- Add more physics not necessarily give better agreement with observations (!)

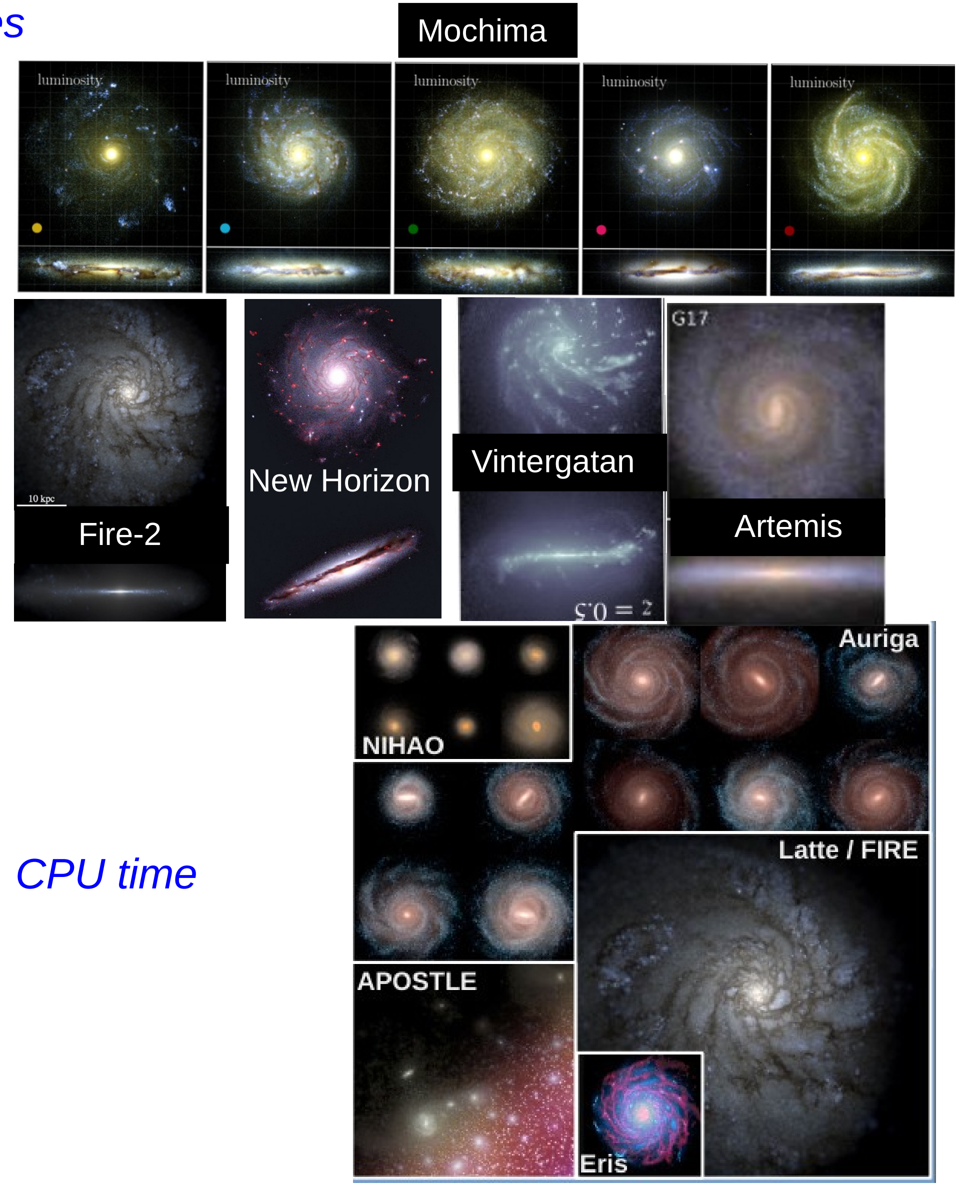
(recipes/models, parameters, calibration, resolution ...)

And even if baryonic physics under control → formation history changes the galaxy morphology and DM distribution

Summary-Conclusion

Successes

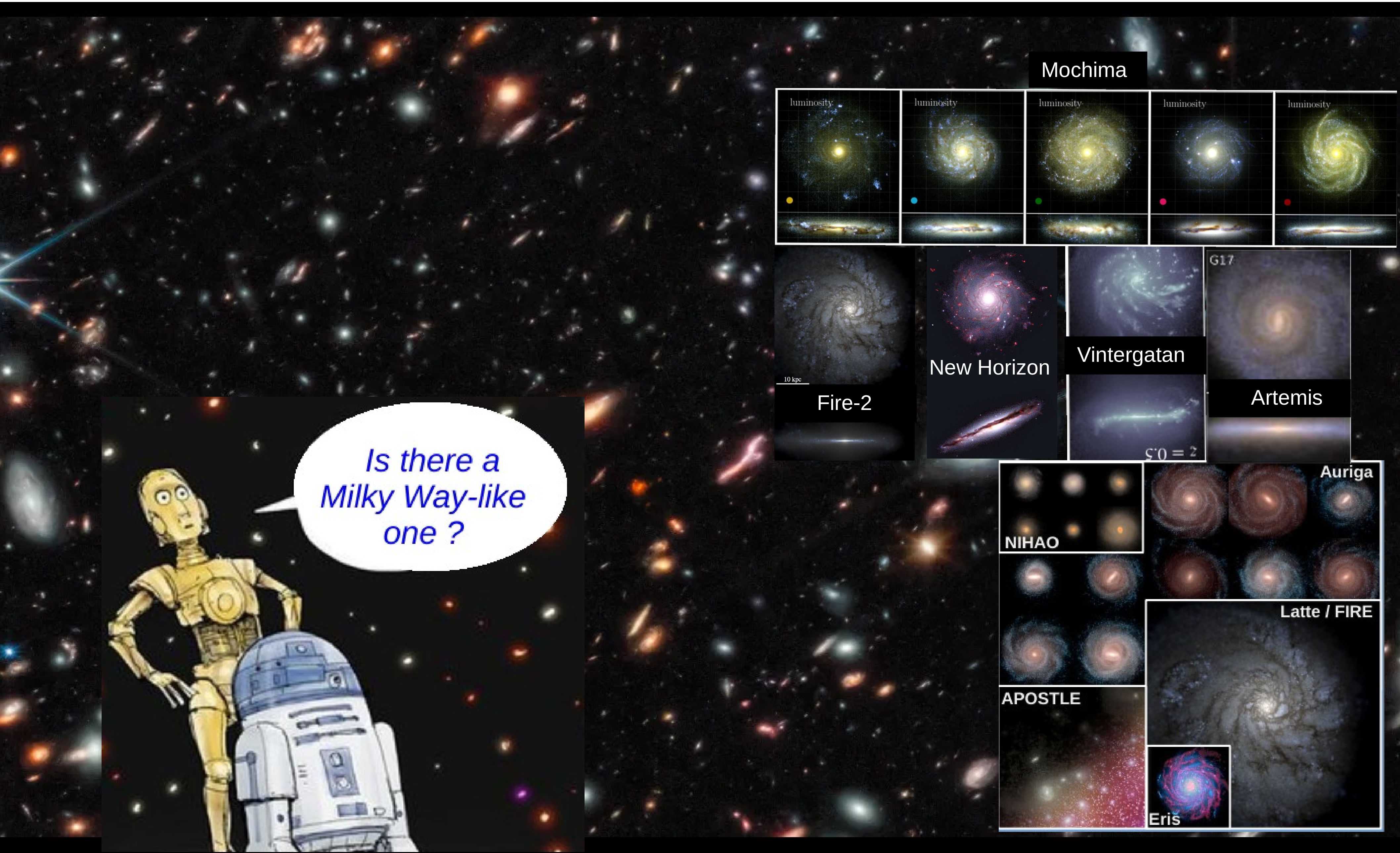
- *Consistent (realistic !?) galaxies from first principles*
- *Numerical experiment to understand physical processes*
- *Comparisons with observations*
- *Test against theoretical models and calibration of semi-analytical models*
- *Dynamical studies*
- *Useful for DM detection rate predictions/uncertainties*



Challenges

- *Improve baryonic physics modeling*
- *Increase resolution, reach individual star formation ?*
- *Additionnal relevant processes (MHD ? cosmic rays ?). CPU time*
- *Early star formation, Bar ? Bulge ? core/cusp ?*
- *GPU ?*
- *Scaling (inhomogeneous volume/resolution)*

Thank you for your attention



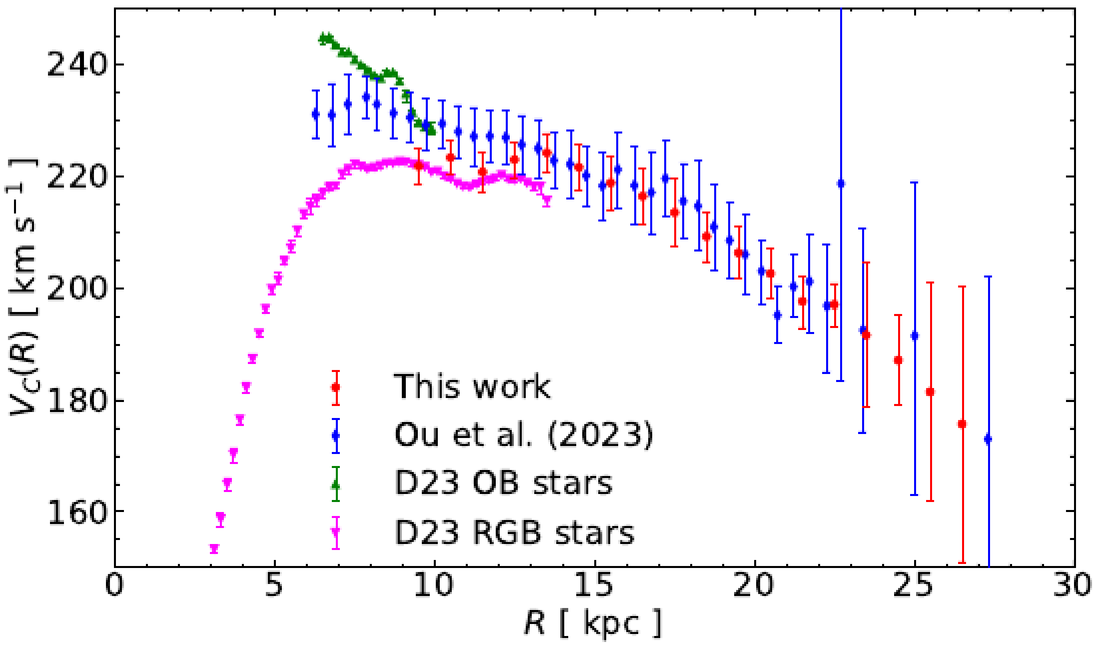
Back-up

Decreasing RC from GAIA data ?

Detection of the Keplerian decline in the Milky Way rotation curve

Yongjun Jiao[1], François Hammer[1], Haifeng Wang[2], Jianling Wang[1,3], Philippe Amram[4], Laurent Chemin[5], and Yanbin Yang[1]

arXiv:2309.00048



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