

Gravitational Wave Turbulence

- Multiple Time Scale Analysis -

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Introduction

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As [non-linear interactions](#) lead to turbulent cascades, we aim to study [how it occurs for GW](#). For that purpose, we base our approach on [wave turbulence](#) [Galtier & Nazarenko, PRL, 2017].



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- There is a reference frame (t, x, y, z) where the metric is diagonal and ∂_z is a Killing field [Hadad and Zakharov, 2014]:

$$g_{\mu\nu} = \begin{pmatrix} -e^{-2\varphi}(1 + \gamma)^2 & 0 & 0 & 0 \\ 0 & e^{-2\varphi}(1 + \beta)^2 & 0 & 0 \\ 0 & 0 & e^{-2\varphi}(1 + \alpha)^2 & 0 \\ 0 & 0 & 0 & e^{2\varphi} \end{pmatrix}.$$

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- The field φ has **small amplitude** oscillation:

$$|\varphi| \propto \varepsilon \ll 1.$$

Developpement of Einstein's equation

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Development of Einstein's equation

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- At **first non-linear order**, we have **dynamical equation**:

$$\square\varphi = -\partial_t [(\alpha + \beta - \gamma)\partial_t\varphi] + \partial_x [(\alpha - \beta + \gamma)\partial_x\varphi] + \partial_y [(-\alpha + \beta + \gamma)\partial_y\varphi],$$

and the **constraint equations**:

$$\partial_x\partial_t\alpha = -2\partial_t\varphi\partial_x\varphi, \quad \partial_y\partial_t\beta = -2\partial_t\varphi\partial_y\varphi, \quad \partial_x\partial_y\gamma = -2\partial_x\varphi\partial_y\varphi.$$

Wave Turbulence approach: a plunge into Fourier space

We introduce the **normal variables**:

$$a^s(\mathbf{k}, t) = \frac{1}{\varepsilon} \left(\sqrt{\frac{k}{2}} \hat{\varphi}(\mathbf{k}, t) + \frac{is}{\sqrt{2k}} \partial_t \hat{\varphi}(\mathbf{k}, t) \right) e^{is\omega_k t} \text{ with } s = \pm 1 \text{ and } \omega_k = k.$$

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We combine the expanded Einstein's equations to get their evolution. It can be written as the **standard form**:

$$\partial_t a^s(\mathbf{k}) = \varepsilon^2 \int_{\mathbb{R}^6} \sum_{s_1, s_2, s_3} \mathbf{L}_{\mathbf{k}\mathbf{k}_1\mathbf{k}_2\mathbf{k}_3}^{s s_1 s_2 s_3} a^{s_1}(\mathbf{k}_1) a^{s_2}(\mathbf{k}_2) a^{s_3}(\mathbf{k}_3) e^{i\Omega_{123}^0 t} \delta_{123}^0(\mathbf{k}) \prod_{i=1}^3 d^2\mathbf{k}_i,$$

where: $\mathbf{L}_{\mathbf{k}\mathbf{k}_1\mathbf{k}_2\mathbf{k}_3}^{s s_1 s_2 s_3}$ the **interaction coefficient**, $\Omega_{123}^0 = sk - s_1 k_1 - s_2 k_2 - s_3 k_3$ and $\delta_{123}^0(\mathbf{k}) = \delta(\mathbf{k} - \mathbf{k}_1 - \mathbf{k}_2 - \mathbf{k}_3)$. They define the **resonant manifold**.

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Main objective: to derive the evolution of the statistical mean of the wave action: $n(\mathbf{k}) = \langle a^s(\mathbf{k}) a^{-s}(-\mathbf{k}) \rangle$.

The multiple time scale method

We introduce a set of **time variables**: $T_0 = t$, $T_2 = \varepsilon^2 t$, $T_4 = \varepsilon^4 t, \dots$ so that:

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With these definitions, we obtain the following **expansion**:

$$\begin{aligned} & (\partial_{T_0} + \varepsilon^2 \partial_{T_2} + \varepsilon^4 \partial_{T_4} + \dots) (a_0^s(\mathbf{k}) + \varepsilon^2 a_2^s(\mathbf{k}) + \varepsilon^4 a_4^s(\mathbf{k}) + \dots) \\ &= \varepsilon^2 \int_{\mathbb{R}^6} \sum_{s_1, s_2, s_3} \mathbf{L}_{\mathbf{k}\mathbf{k}_1\mathbf{k}_2\mathbf{k}_3}^{s s_1 s_2 s_3} (a_0^{s_1}(\mathbf{k}_1) + \varepsilon^2 a_2^{s_1}(\mathbf{k}_1) + \dots) (a_0^{s_2}(\mathbf{k}_2) + \varepsilon^2 a_2^{s_2}(\mathbf{k}_2) + \dots) \\ & \quad \times (a_0^{s_3}(\mathbf{k}_3) + \varepsilon^2 a_2^{s_3}(\mathbf{k}_3) + \dots) \delta_{123}^0(\mathbf{k}) e^{i\Omega_{123}^0 T_0} \prod_{i=1}^3 d^2\mathbf{k}_i. \end{aligned}$$

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Their time integration introduce **seculars drifts** (terms $\propto T_0$ or $\propto T_0^2$).

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and:

$$\begin{aligned} \partial_{T_4} n(\mathbf{k}) = & 36\pi \int_{\mathbb{R}^6} \sum_{s_1, s_2, s_3} s \left| \mathbf{L}_{\mathbf{k}\mathbf{k}_1\mathbf{k}_2\mathbf{k}_3}^{ss_1s_2s_3} \right|^2 \left(\frac{s}{n(\mathbf{k})} - \frac{s_1}{n(\mathbf{k}_1)} - \frac{s_2}{n(\mathbf{k}_2)} - \frac{s_3}{n(\mathbf{k}_3)} \right) \\ & \times n(\mathbf{k})n(\mathbf{k}_1)n(\mathbf{k}_2)n(\mathbf{k}_3)\delta_{123}^0(\omega)\delta_{123}^0(\mathbf{k}) \prod_{i=1}^3 d^2\mathbf{k}_i. \end{aligned}$$

The kinetic equation

Using [all the symmetries](#) of the interaction coefficient, we find:

$$\begin{aligned} \partial_t n(\mathbf{k}) = & 36\pi\epsilon^4 \int_{\mathbb{R}^6} \left| \mathbf{L}_{\mathbf{k}-\mathbf{k}_1\mathbf{k}_2\mathbf{k}_3}^{s-sss} \right|^2 \left(\frac{1}{n(\mathbf{k})} + \frac{1}{n(\mathbf{k}_1)} - \frac{1}{n(\mathbf{k}_2)} - \frac{1}{n(\mathbf{k}_3)} \right) \\ & \times n(\mathbf{k})n(\mathbf{k}_1)n(\mathbf{k}_2)n(\mathbf{k}_3)\delta_{23}^{01}(\omega)\delta_{23}^{01}(\mathbf{k}) \prod_{i=1}^3 d^2\mathbf{k}_i. \end{aligned}$$

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inverse cascade in wave action

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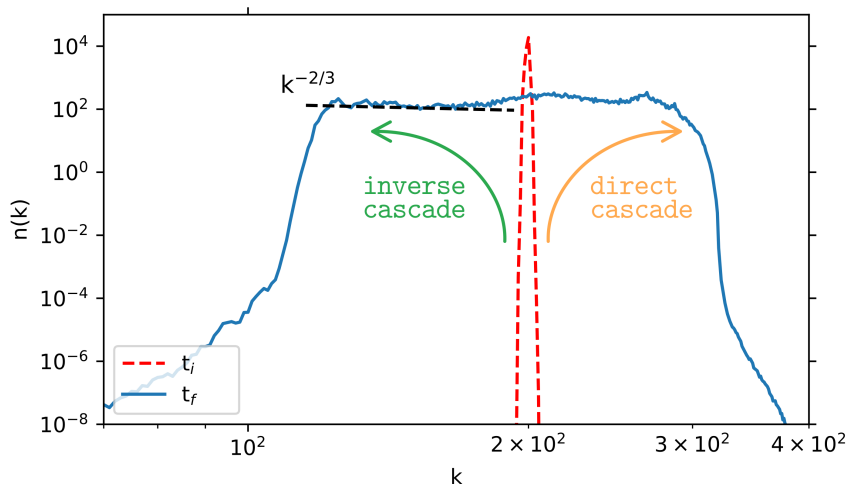
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$$\underbrace{n(k) \propto (-\zeta)^{1/3} k^{-2/3}}_{\text{inverse cascade in wave action}} \quad \text{and} \quad \underbrace{n(k) \propto \epsilon^{1/3} k^{-1}}_{\text{direct cascade in energy}}$$

Numerical results

First numerical results on GPU (DNS, 1024×1024):



Conclusion

We aim to describe the **weakly non linear regime** of gravitational waves using **statistical and analytical** tools [Gay *et al.*, in prep.], in order to predict the existence of **a dual cascade** numerically observed.

Further works need to be performed:

- How to generalize this method to a more general model? What about the other polarization?
- What about the strong turbulent regime? Is there a link with inflation ?