

Hybrid Quantum State in 2d Dilaton Gravity

Yohan Potaux, Institut Denis Poisson

June 1st, 2023

Joint work with Debajyoti Sarkar and Sergey Solodukhin

Based on 2112.03855 and 2212.13208

- Information loss problem: is black hole evolution unitary ?

- Information loss problem: is black hole evolution unitary ?
- We need to take quantum aspects into account:
semi-classical GR, quantized field propagate on a classical background, but can influence its geometry by back-reaction.

Introduction

- Information loss problem: is black hole evolution unitary ?
- We need to take quantum aspects into account: semi-classical GR, quantized field propagate on a classical background, but can influence its geometry by back-reaction.
- Why two dimensions ? Easier to study analytically solvable models, can help to gain some insight. If successful, one may try to generalize to four dimensions.

Classical CGHS model

$$\mathcal{S}_{CGHS} = \frac{1}{2\pi} \int_M d^2x \sqrt{-g} \left\{ e^{-2\phi} [R + 4(\nabla\phi)^2 + 4\lambda^2] - \frac{1}{2}(\nabla f)^2 \right\} \quad (1)$$

(Callan, Giddings, Harvey, Strominger, 1992)

Classical CGHS model

$$\mathcal{S}_{CGHS} = \frac{1}{2\pi} \int_M d^2x \sqrt{-g} \left\{ e^{-2\phi} [R + 4(\nabla\phi)^2 + 4\lambda^2] - \frac{1}{2}(\nabla f)^2 \right\} \quad (1)$$

(Callan, Giddings, Harvey, Strominger, 1992)

- $\phi \rightarrow$ dilaton scalar field, $e^{-\phi} \sim$ radius,

Classical CGHS model

$$\mathcal{S}_{CGHS} = \frac{1}{2\pi} \int_M d^2x \sqrt{-g} \left\{ e^{-2\phi} [R + 4(\nabla\phi)^2 + 4\lambda^2] - \frac{1}{2}(\nabla f)^2 \right\} \quad (1)$$

(Callan, Giddings, Harvey, Strominger, 1992)

- $\phi \rightarrow$ dilaton scalar field, $e^{-\phi} \sim$ radius,
- $g \rightarrow$ two-dimensional spacetime metric,
- $\lambda \rightarrow$ cosmological constant,
- $f \rightarrow$ classical matter.

Classical CGHS model

$$\mathcal{S}_{CGHS} = \frac{1}{2\pi} \int_M d^2x \sqrt{-g} \left\{ e^{-2\phi} [R + 4(\nabla\phi)^2 + 4\lambda^2] - \frac{1}{2}(\nabla f)^2 \right\} \quad (1)$$

(Callan, Giddings, Harvey, Strominger, 1992)

- $\phi \rightarrow$ dilaton scalar field, $e^{-\phi} \sim$ radius,
- $g \rightarrow$ two-dimensional spacetime metric,
- $\lambda \rightarrow$ cosmological constant,
- $f \rightarrow$ classical matter.

\rightarrow Classical black hole solution, now we need to include quantum matter.

Semi-classical RST model

$$\mathcal{S}_{RST} = \mathcal{S}_{CGHS} - \sum_{i=1}^2 \frac{\kappa_i}{2\pi} \int d^2x \sqrt{-g} \left\{ \frac{1}{2} (\nabla \psi_i)^2 + (\psi_i + \phi) R \right\} \quad (2)$$

(Russo, Susskind, Thorlacius, 1992)

Semi-classical RST model

$$\mathcal{S}_{RST} = \mathcal{S}_{CGHS} - \sum_{i=1}^2 \frac{\kappa_i}{2\pi} \int d^2x \sqrt{-g} \left\{ \frac{1}{2} (\nabla \psi_i)^2 + (\psi_i + \phi) R \right\} \quad (2)$$

(Russo, Susskind, Thorlacius, 1992)

- $\psi_1, \psi_2 \rightarrow$ auxiliary fields to take into account the conformal anomaly, each in a different quantum state, \rightarrow "hybrid" quantum state.

Semi-classical RST model

$$\mathcal{S}_{RST} = \mathcal{S}_{CGHS} - \sum_{i=1}^2 \frac{\kappa_i}{2\pi} \int d^2x \sqrt{-g} \left\{ \frac{1}{2} (\nabla \psi_i)^2 + (\psi_i + \phi) R \right\} \quad (2)$$

(Russo, Susskind, Thorlacius, 1992)

- $\psi_1, \psi_2 \rightarrow$ auxiliary fields to take into account the conformal anomaly, each in a different quantum state, \rightarrow "hybrid" quantum state.
- $\kappa_1, \kappa_2 \rightarrow$ central charges associated to ψ_1, ψ_2 .

- ψ_1, ψ_2 not completely determined by the equations

- ψ_1, ψ_2 not completely determined by the equations
→ we fix them by choosing boundary conditions
(= quantum state),

- ψ_1, ψ_2 not completely determined by the equations
→ we fix them by choosing boundary conditions
(= quantum state),
→ impose value of quantum energy density $T^{(q)}$ at flat
infinity.

- ψ_1, ψ_2 not completely determined by the equations
→ we fix them by choosing boundary conditions
(= quantum state),
→ impose value of quantum energy density $T^{(q)}$ at flat infinity.
- **Hartle-Hawking** state: $T^{(q)} \xrightarrow{\infty}$ thermality

Quantum states

- ψ_1, ψ_2 not completely determined by the equations
→ we fix them by choosing boundary conditions
(= quantum state),
→ impose value of quantum energy density $T^{(q)}$ at flat infinity.
- **Hartle-Hawking** state: $T^{(q)} \xrightarrow{\infty} \text{thermality}$
- **Boulware** state: $T^{(q)} \xrightarrow{\infty} 0$

Hybrid quantum state

- For Hartle-Hawking and Boulware states: solutions are known (see *e.g.* Sarkar, Solodukhin, Potaux, 2022).

Hybrid quantum state

- For Hartle-Hawking and Boulware states: solutions are known (see *e.g.* Sarkar, Solodukhin, Potaux, 2022).
- Hartle-Hawking: radiation observed at $\infty \Rightarrow$ well suited to physical particles (central charge $\kappa > 0$).

Hybrid quantum state

- For Hartle-Hawking and Boulware states: solutions are known (see *e.g.* Sarkar, Solodukhin, Potaux, 2022).
- Hartle-Hawking: radiation observed at $\infty \Rightarrow$ well suited to physical particles (central charge $\kappa > 0$).
- Boulware: no radiation observed at $\infty \Rightarrow$ well suited to non-physical particles (*e.g.* ghosts, central charge $\kappa < 0$).

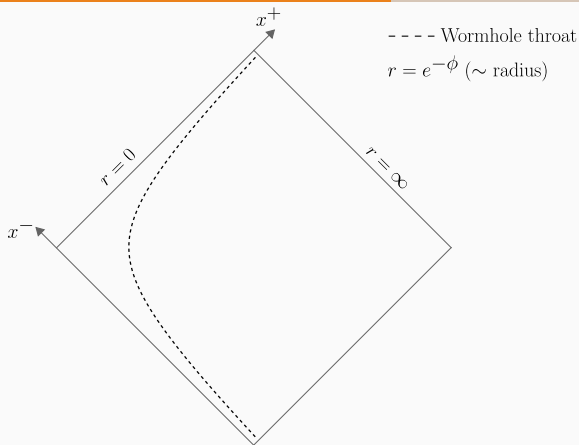
Hybrid quantum state

- For Hartle-Hawking and Boulware states: solutions are known (see *e.g.* Sarkar, Solodukhin, Potaux, 2022).
- Hartle-Hawking: radiation observed at $\infty \Rightarrow$ well suited to physical particles (central charge $\kappa > 0$).
- Boulware: no radiation observed at $\infty \Rightarrow$ well suited to non-physical particles (*e.g.* ghosts, central charge $\kappa < 0$).
- **Question:** what happens when both physical and non-physical particles are present ?

Hybrid quantum state

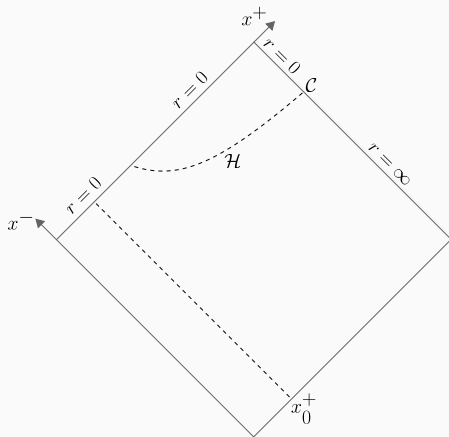
- For Hartle-Hawking and Boulware states: solutions are known (see *e.g.* Sarkar, Solodukhin, Potaux, 2022).
- Hartle-Hawking: radiation observed at $\infty \Rightarrow$ well suited to physical particles (central charge $\kappa > 0$).
- Boulware: no radiation observed at $\infty \Rightarrow$ well suited to non-physical particles (*e.g.* ghosts, central charge $\kappa < 0$).
- **Question:** what happens when both physical and non-physical particles are present ?
 - \rightarrow hybrid quantum state with $\kappa_1 > 0$ and $\kappa_2 < 0$,
 - \rightarrow we focus on the $\kappa = \kappa_1 + \kappa_2 < 0$ case.

Static solution

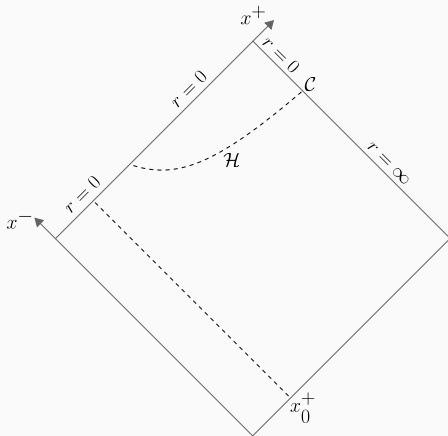


Structure of a Damour-Solodukhin wormhole (no singularity,
no horizon),
→ black hole mimicker.

Dynamical solution

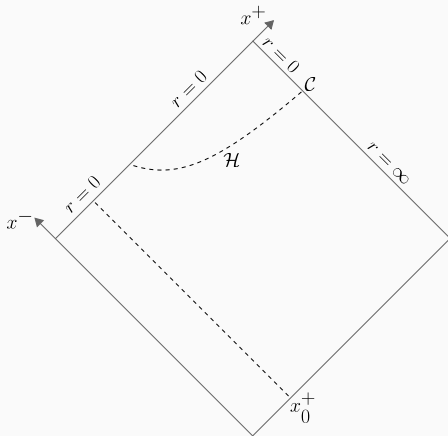


Dynamical solution



- Creation of an apparent horizon \mathcal{H} but still no singularity.

Dynamical solution



- Creation of an apparent horizon \mathcal{H} but still no singularity.
- What about the radiation entropy ?

Define radiation entropy at infinity by

$$\partial_- S = 2\pi(-x^-)T_{--}^{(q)} \quad (3)$$

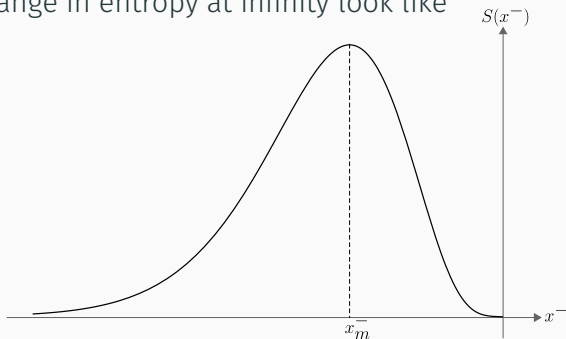
($dS = T^{-1}dE$ in asymptotically flat coordinates).

Define radiation entropy at infinity by

$$\partial_- S = 2\pi(-x^-)T_{--}^{(q)} \quad (3)$$

($dS = T^{-1}dE$ in asymptotically flat coordinates).

→ The change in entropy at infinity look like



→ Page Curve, sign of unitary evolution (Page, 2013).

Conclusion

- Boulware particles drastically change the static spacetime geometry: from black hole with Hawking radiation (Hartle-Hawking state) to completely regular spacetime with a wormhole throat.

Conclusion

- Boulware particles drastically change the static spacetime geometry: from black hole with Hawking radiation (Hartle-Hawking state) to completely regular spacetime with a wormhole throat.
- They also help in recovering all the information of the system, as the Page curve shows.

Conclusion

- Boulware particles drastically change the static spacetime geometry: from black hole with Hawking radiation (Hartle-Hawking state) to completely regular spacetime with a wormhole throat.
- They also help in recovering all the information of the system, as the Page curve shows.
- Is this reproducible in four dimensions ?

Thank you!

- Static solution:

$$\Omega = (\kappa_1 + \kappa_2)\phi + e^{-2\phi} = -\lambda^2 x^+ x^- - \frac{\kappa_2}{2} \ln(-\lambda^2 x^+ x^-) + \frac{M}{\lambda} \quad (4)$$

- Dynamical solution:

$$\Omega = -\lambda^2 x^+ \left(x^- + \frac{m}{\lambda^3 x_0^+} \right) - \frac{\kappa_2}{2} \ln(-\lambda^2 x^+ x^-) + \frac{M + m}{\lambda} \quad (5)$$

where m is the energy of the shock wave.