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et de physique des particules

The Einstein Telescope Project

International Conference
on the Physics of the Two Infinities

Kyoto, March 27th, 2023

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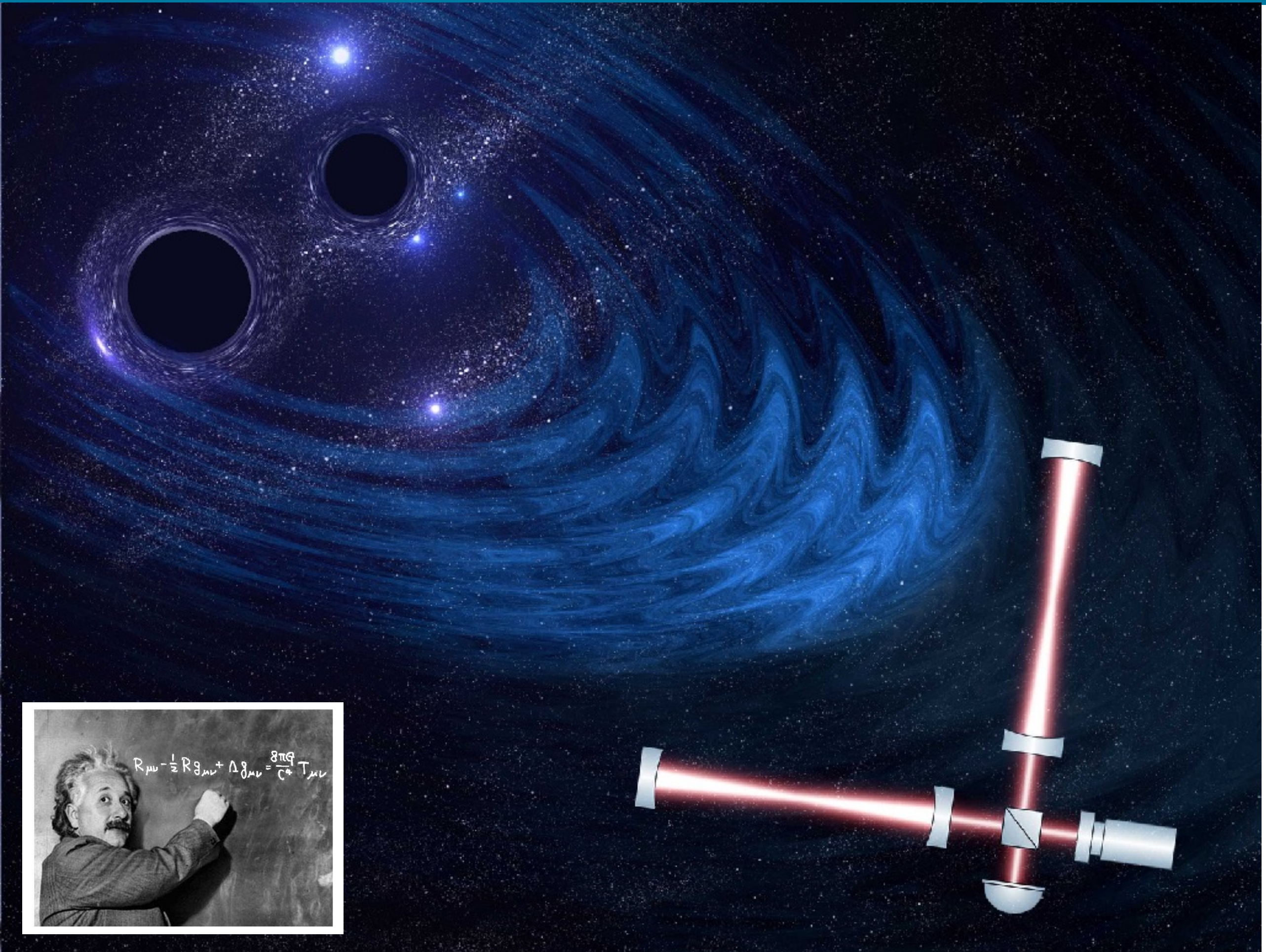
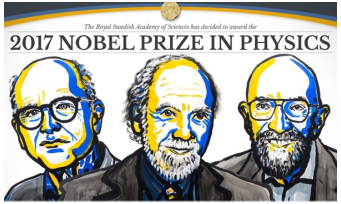
Introduction



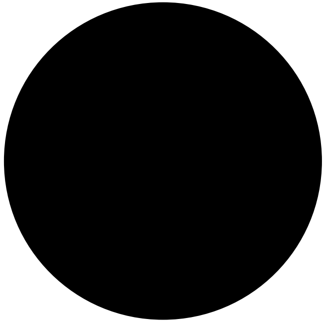
100 years after their prediction as part of Einstein's theory of general relativity, the discovery of gravitational waves opens a new way to explore and study the Universe!

Nobel Prize in Physics 2017

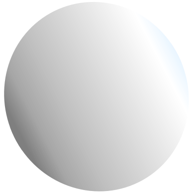


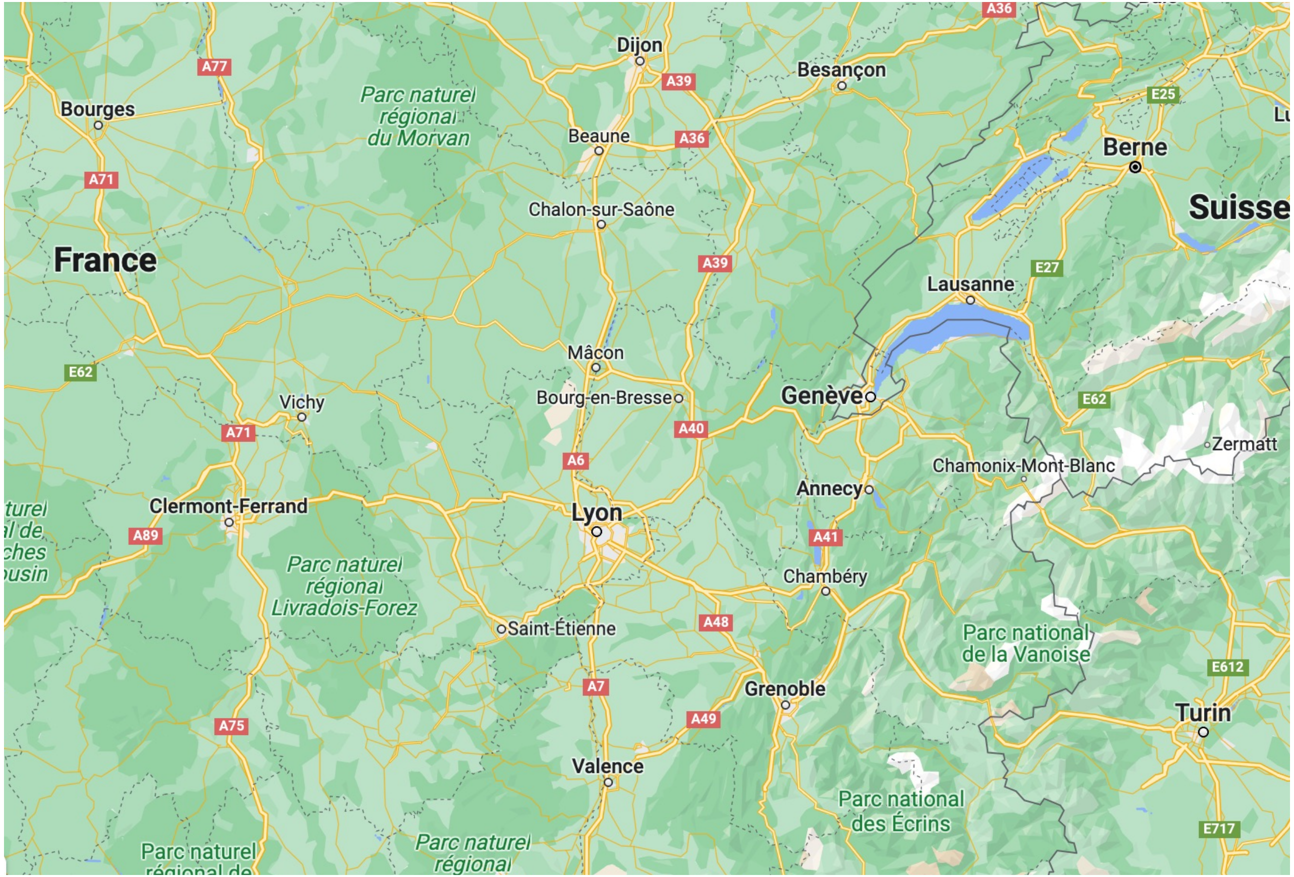
Introduction

Black hole 15$M_\odot$



Neutron star 1.4 $M_\odot$





Introduction

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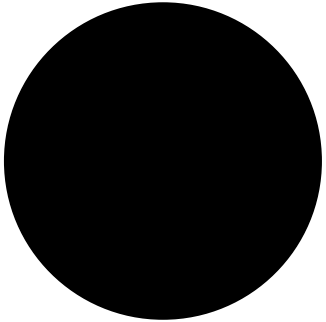
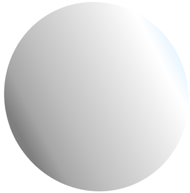
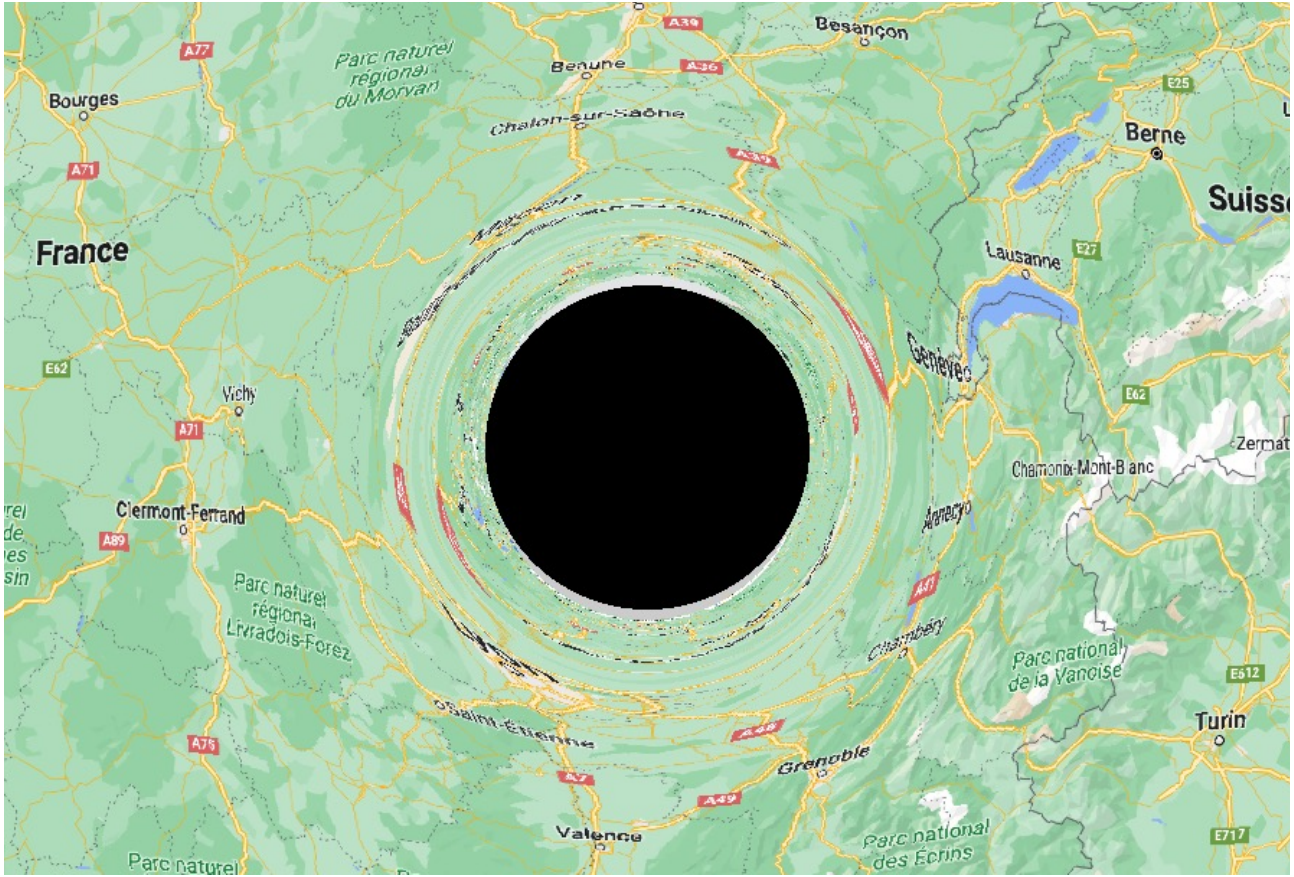
Black hole $15M_{\odot}$



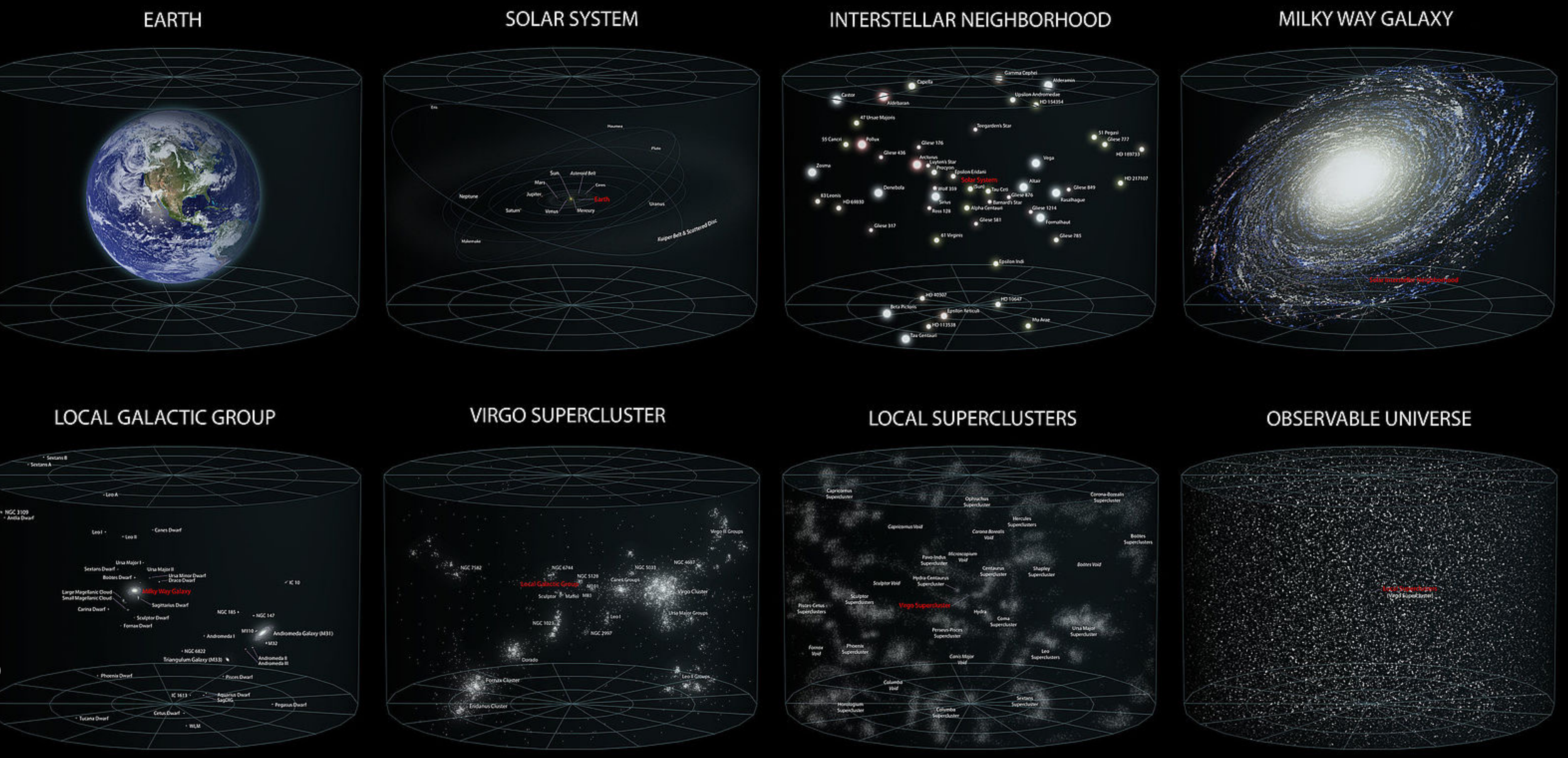
Neutron star $1.4\ M_{\odot}$





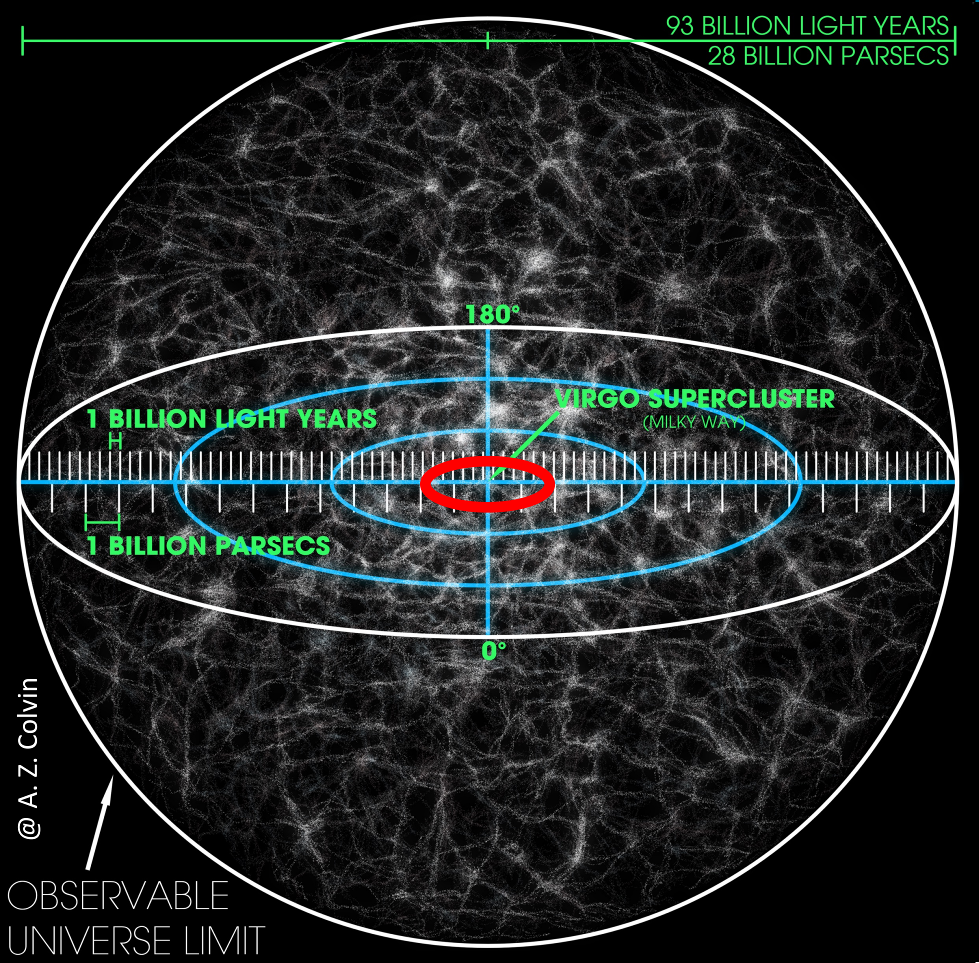
@ David Ball, University of Arizona: http://bhpire.arizona.edu/research/tools/black-hole-photo-filter/

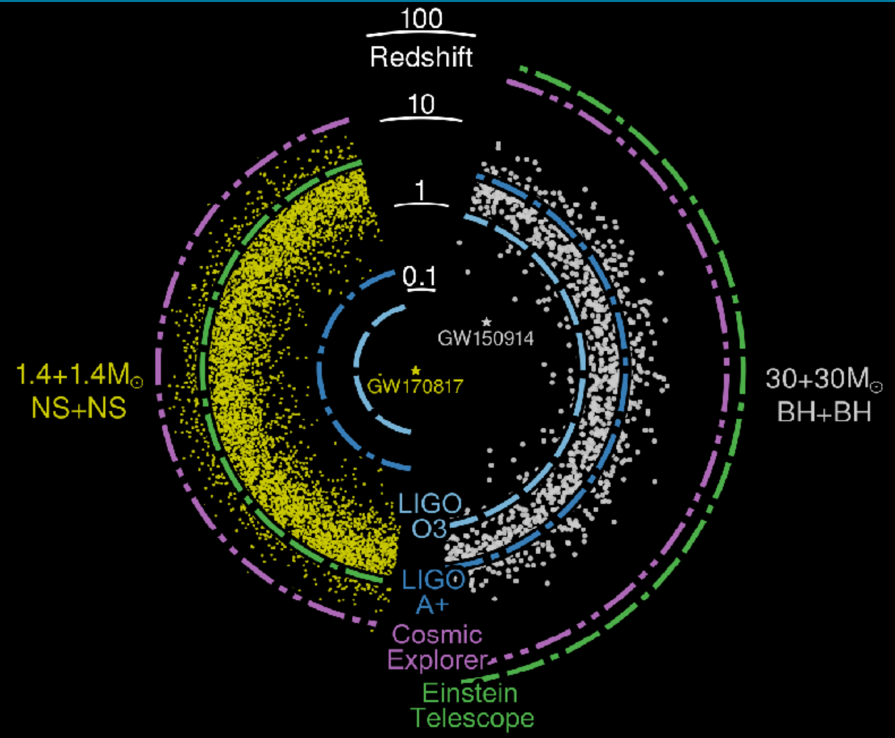
Introduction



@ A. Z. Colvin

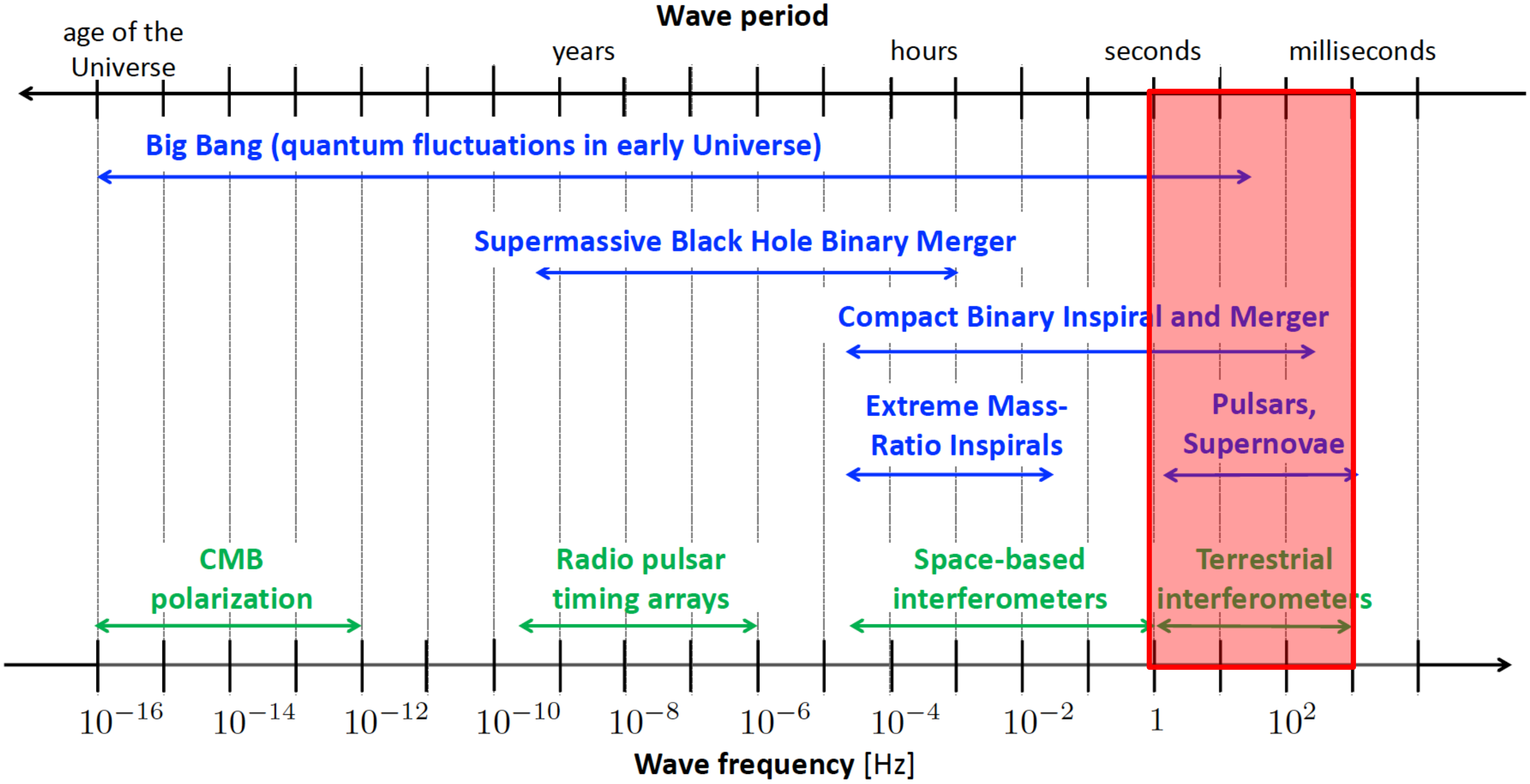
Introduction





Einstein Telescope aims to study most of the observable Universe

Physical phenomenon, Search techniques



The LIGO-Virgo-KAGRA (LVK) network

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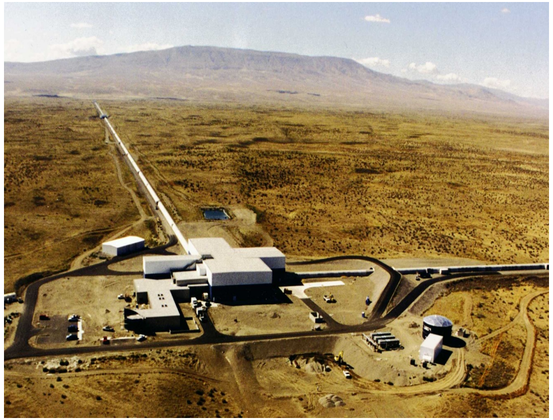
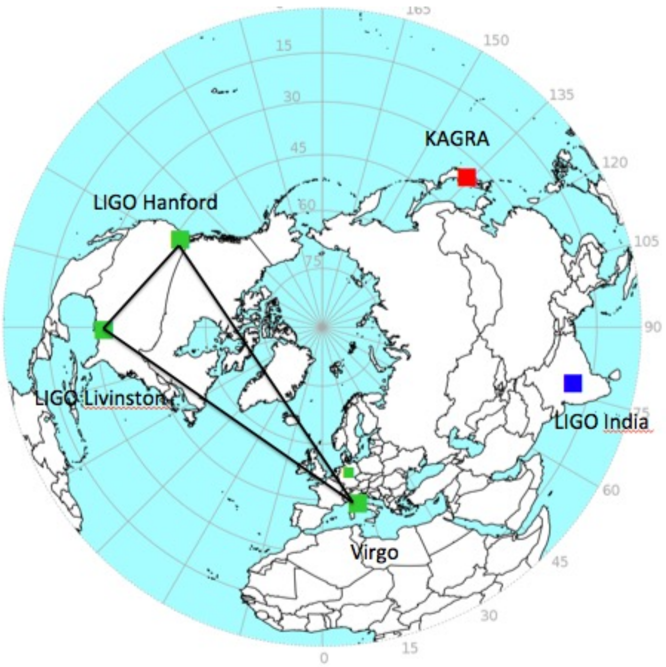
LIGO Hanford



KAGRA



LIGO Livingston

Virgo

Einstein Telescope

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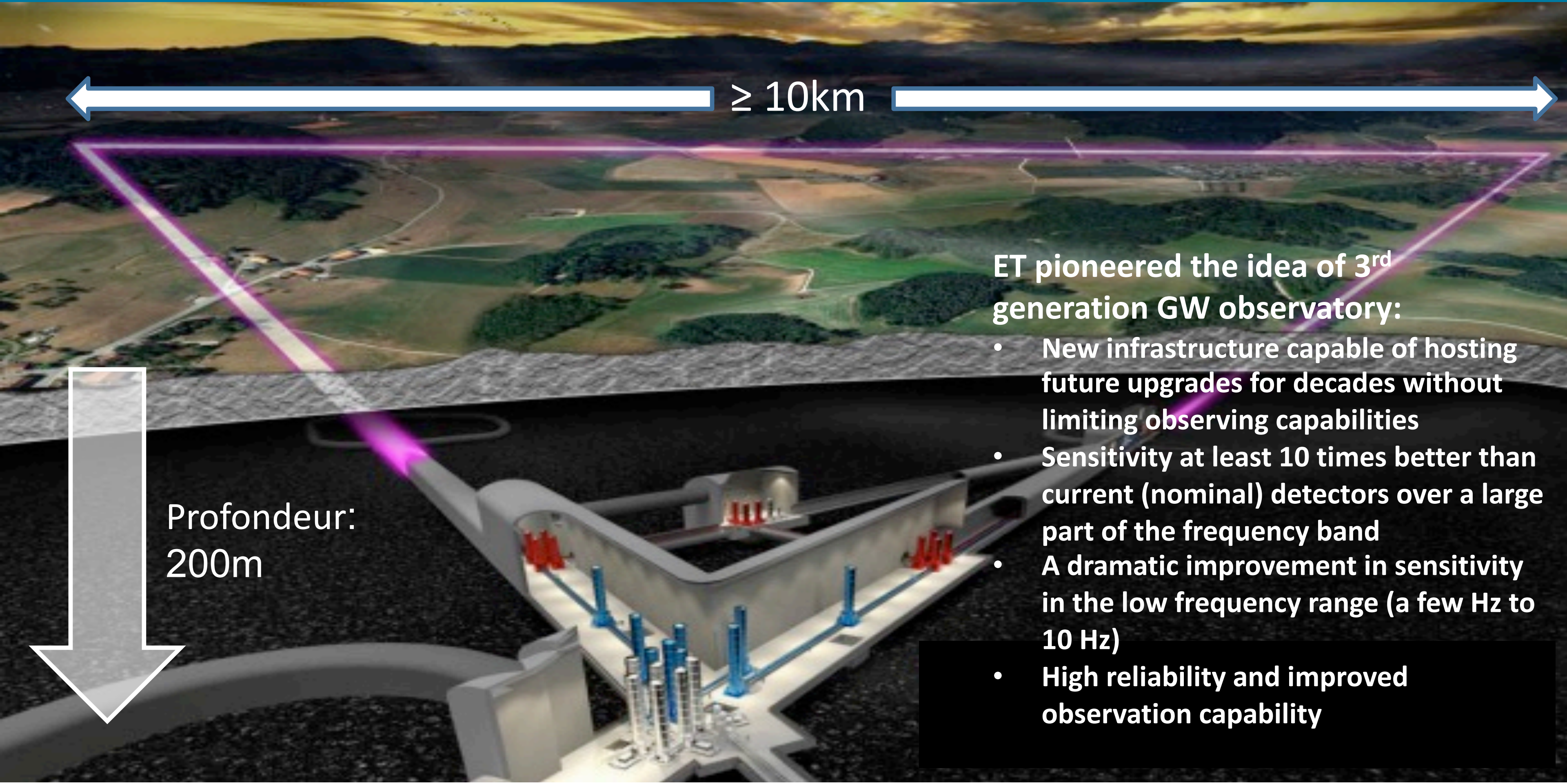

ET pioneered the idea of 3rd generation GW observatory:

- New infrastructure capable of hosting future upgrades for decades without limiting observing capabilities
- Sensitivity at least 10 times better than current (nominal) detectors over a large part of the frequency band
- A dramatic improvement in sensitivity in the low frequency range (a few Hz to 10 Hz)
- High reliability and improved observation capability

GW science programme

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ASTROPHYSICS

Black hole properties
origin (stellar vs. primordial)
evolution, demography

Neutron star properties
interior structure, equation of state & properties of dense matter, demography

Multi-band and –messenger astronomy
joint GW/EM observations (GRB, kilonova,...)
multiband GW detection (LISA)
neutrinos

Detection of new astrophysical sources
core collapse supernovae
isolated neutron stars (Radio, X, g and GW, FRB, …)
stochastic background of astrophysical origin

FUNDAMENTAL PHYSICS AND COSMOLOGY

The nature of compact objects
near-horizon physics, tests of no-hair theorem
exotic compact objects, phase transition in dense matter

Tests of General Relativity
post-Newtonian expansion, strong field regime

Dark matter
primordial BHs
axions, dark matter accreting on compact objects

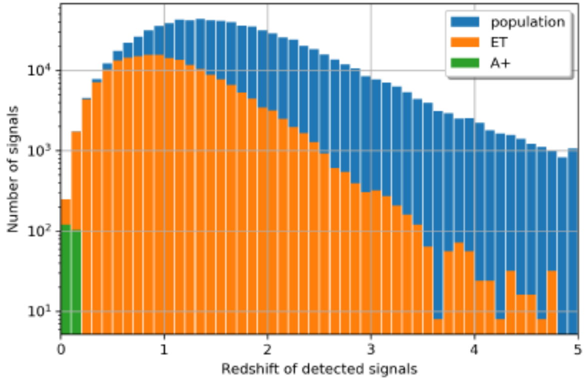
Dark energy and modifications of gravity on cosmological scales
dark energy equation of state,
modified GW propagation

Stochastic backgrounds of cosmological origin inflation, phase transitions, cosmic strings

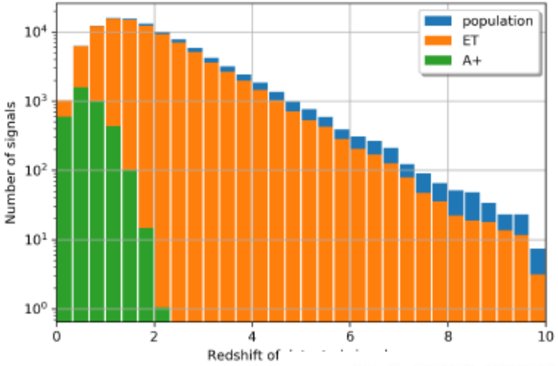
The “unexpected” ?

ET sensitivity

BINARY NEUTRON-STAR MERGERS

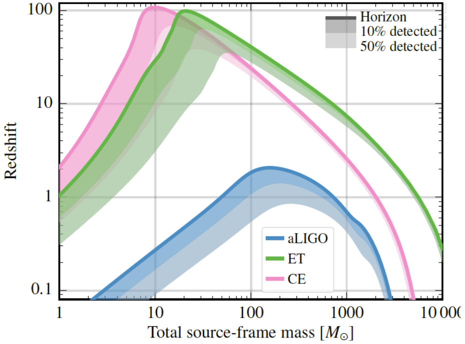


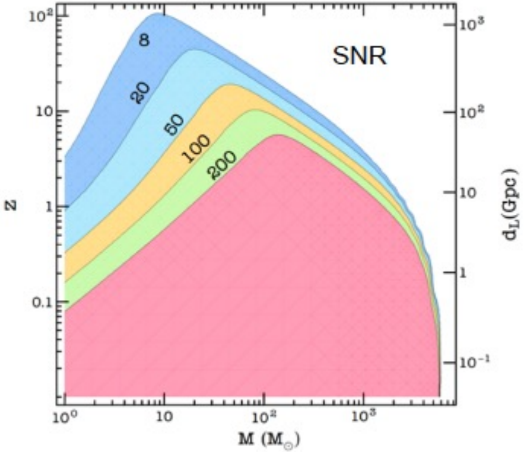
BINARY BLACK-HOLE MERGERS



- 10^5-10^6 BBH detections per year
- 10^4-10^5 BNS detections per year among which ~10-100 with EM counterparts
- High SNR events
- Overlapping events

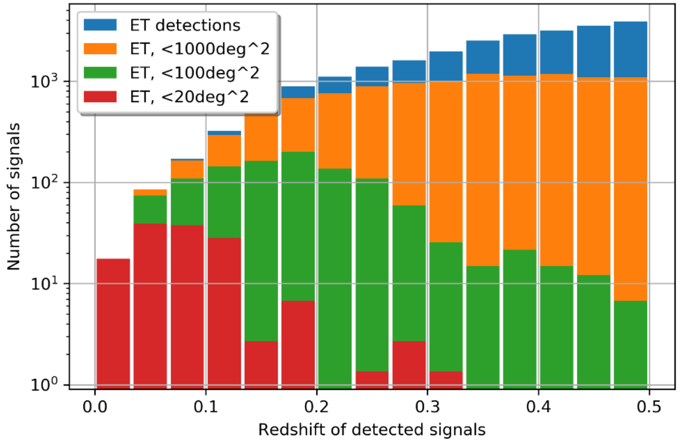
~1 detection every 30s



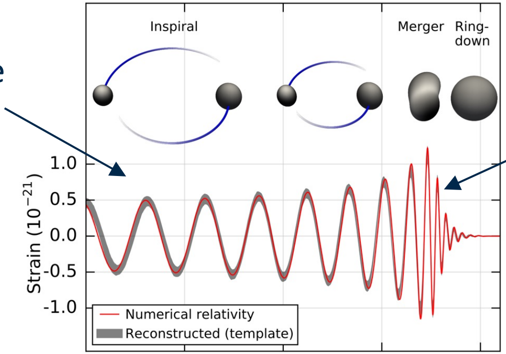


ET sensitivity

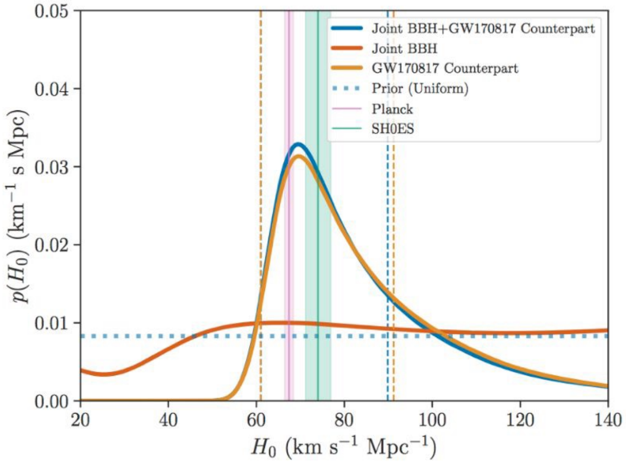
- **BNS detection with EM counterparts and localization precision < 20 deg²** : o(10-100) per year
- Overlap with many BBH signals
- Potentially, very long signals
- ET will be able to provide alerts few hours before the merger



Identify early the inspiral …

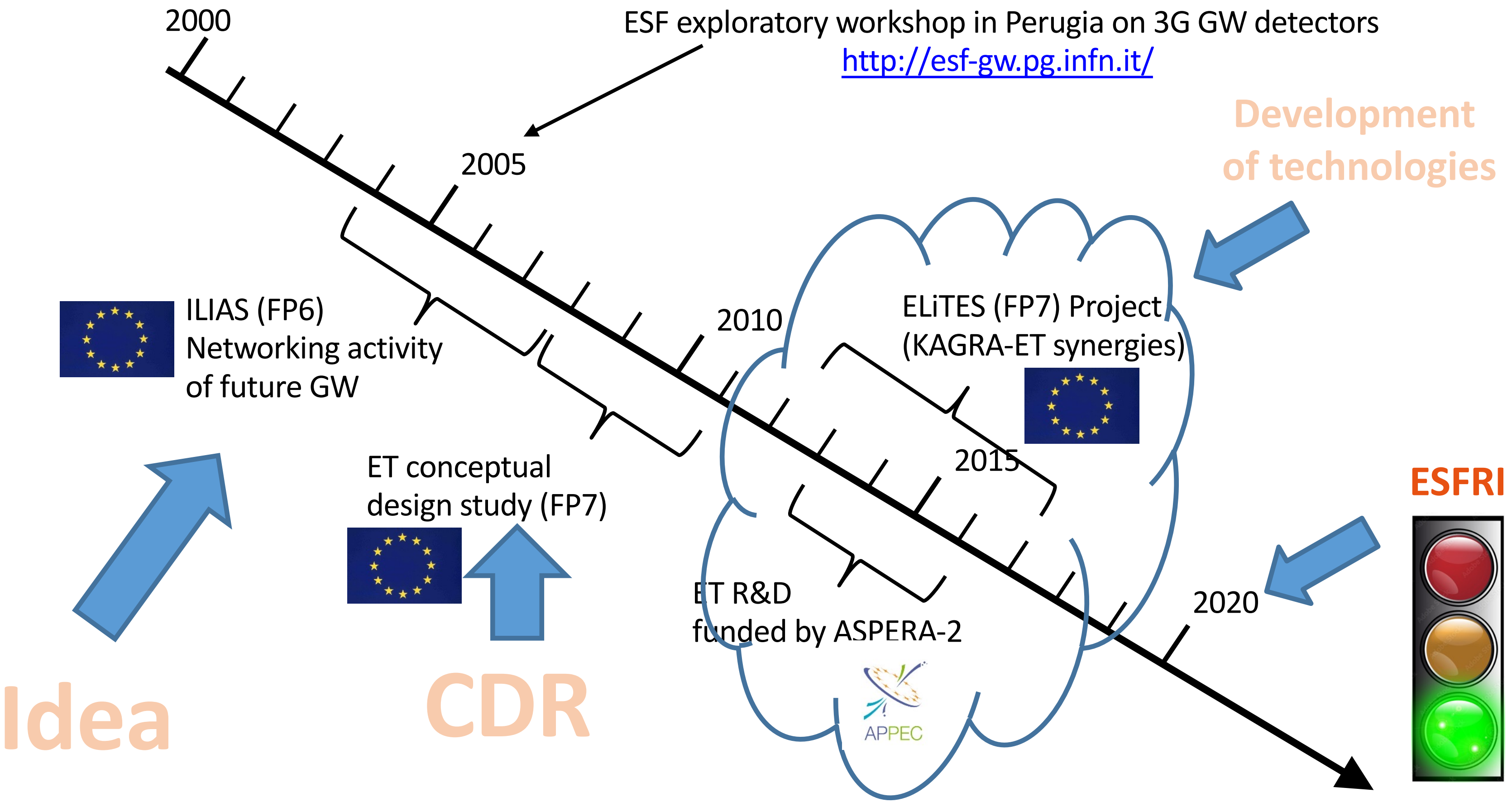


… and provide alert before the merger phase



- And with ~500 BNS-EM detection, we can reach Planck resolution on H_0 measurement

The genesis of the Einstein Telescope



The genesis of the Einstein Telescope

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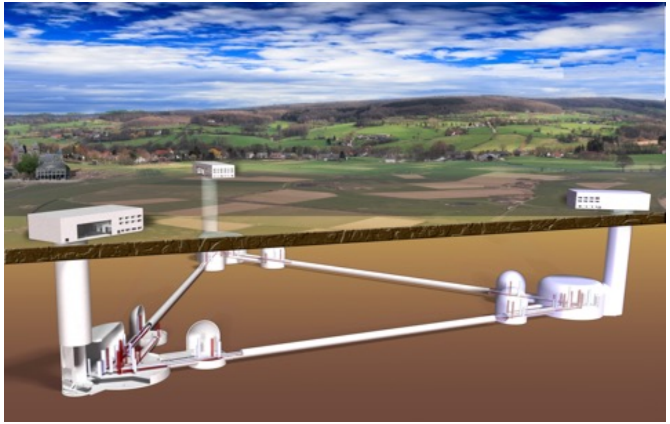
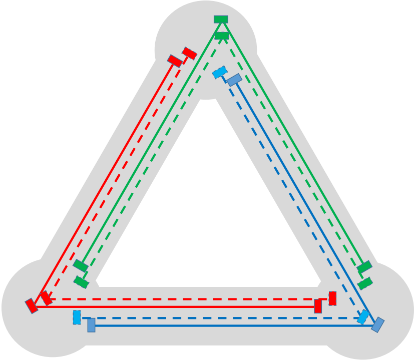
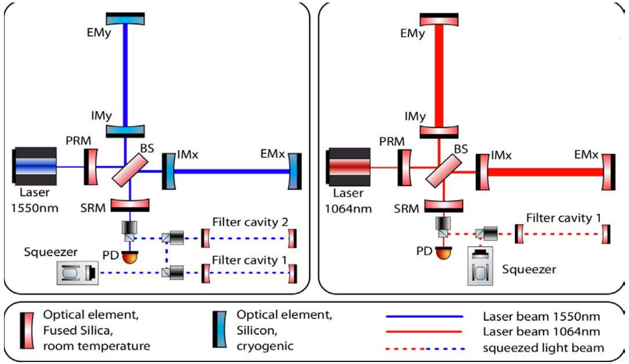
ESFRI

Underground infrastructure



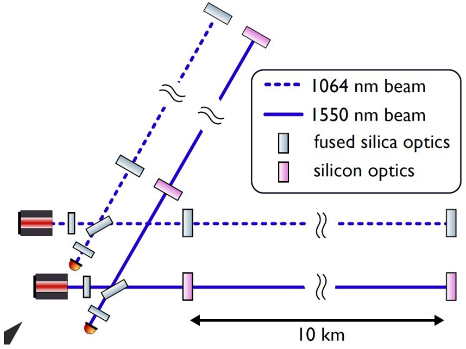




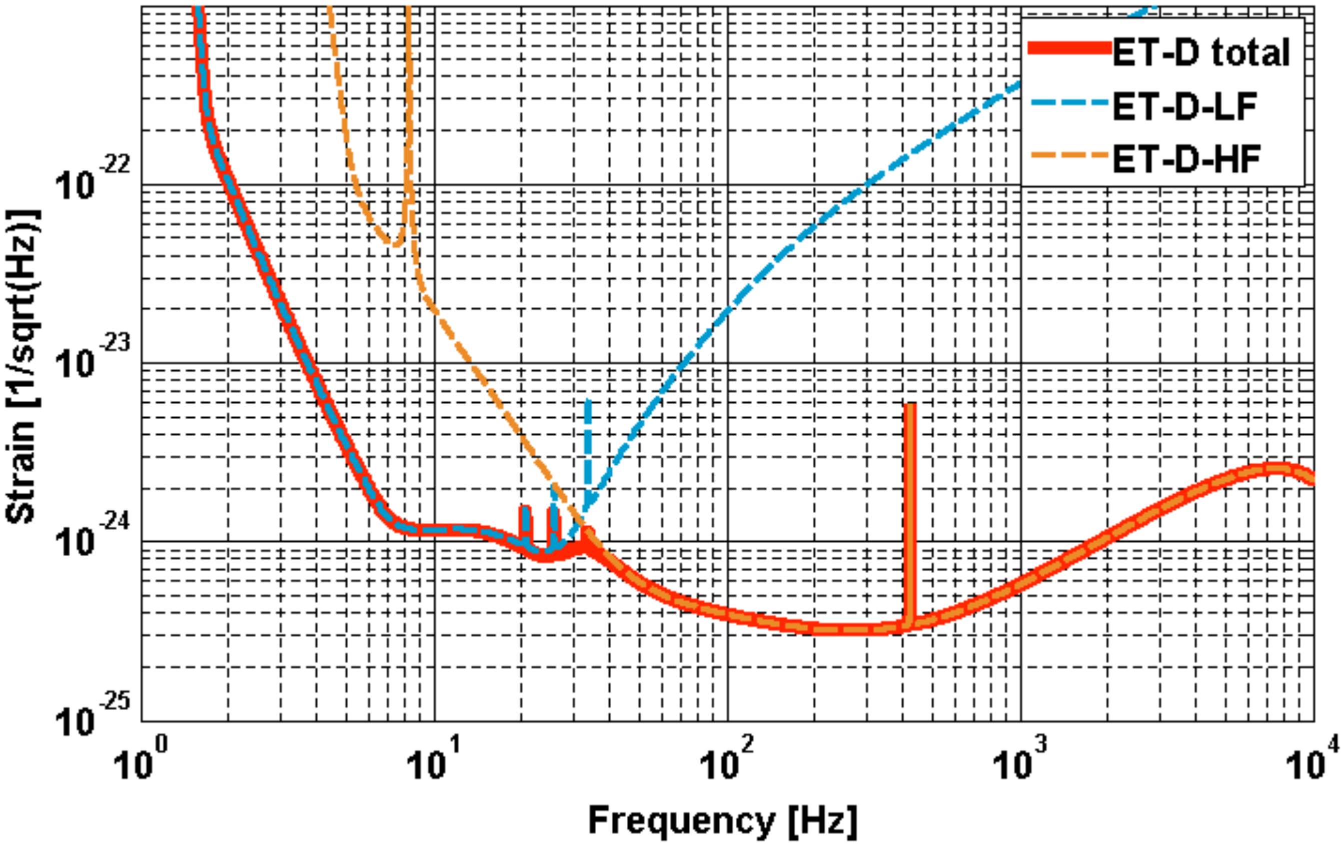
Parameter	ET-HF	ET-LF
Arm length	10 km	10 km
Input power (after IMC)	500 W	3 W
Arm power	3 MW	18 kW
Temperature	290 K	10-20 K
Mirror material	fused silica	silicon
Mirror diameter / thickness	62 cm / 30 cm	45 cm/ 57 cm
Mirror masses	200 kg	211 kg
Laser wavelength	1064 nm	1550 nm
SR-phase (rad)	tuned (0.0)	detuned (0.6)
SR transmittance	10 %	20 %
Quantum noise suppression	freq. dep. squeez.	freq. dep. squeez.
Filter cavities	1×300 m	2×1.0 km
Squeezing level	10 dB (effective)	10 dB (effective)
Beam shape	TEM_{00}	TEM_{00}
Beam radius	12.0 cm	9 cm
Scatter loss per surface	37 ppm	37 ppm
Seismic isolation	SA, 8 m tall	mod SA, 17 m tall
Seismic (for $f > 1$ Hz)	$5 \cdot 10^{-10}$ m/f^2	$5 \cdot 10^{-10}$ m/f^2
Gravity gradient subtraction	none	factor of a few

Xylophone: 2 sensitive interferometers at different frequencies

Triangle configuration to have 3 detectors in the same infrastructure



ET Sensitivity



ET technologies

The multi-interferometer approach asks for two parallel technology developments:

Challenging engineering

New technology in cryo-cooling

New technology in optics

New laser technology

High precision mechanics and low noise controls

High quality opto-electronics and new controls

ET-LF:

- Underground
- Cryogenics
- Silicon (Sapphire) test masses
- Large test masses
- New coatings
- New laser wavelength
- Seismic suspensions
- Frequency dependent squeezing

ET-HF:

- High power laser
- Large test masses
- New coatings
- Thermal compensation
- Frequency dependent squeezing

Evolved laser technology

Evolved technology in optics

Highly innovative adaptive optics

High quality opto-electronics and new controls

Technologies and challenges for ET

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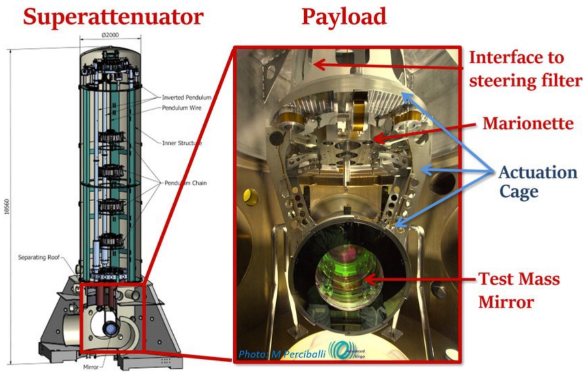
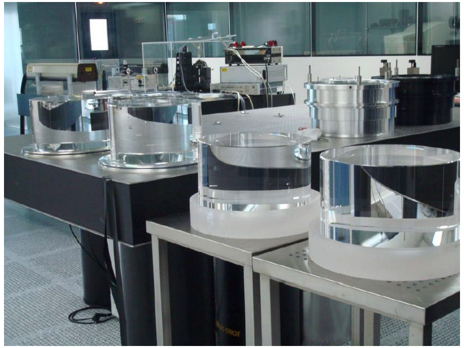
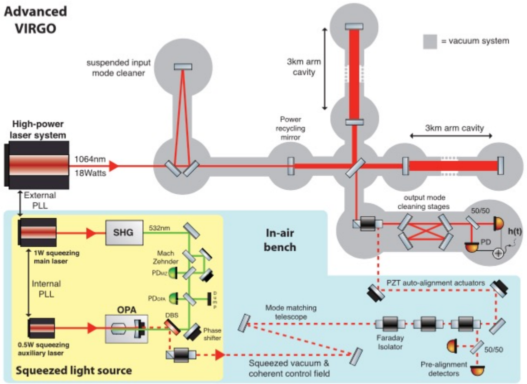
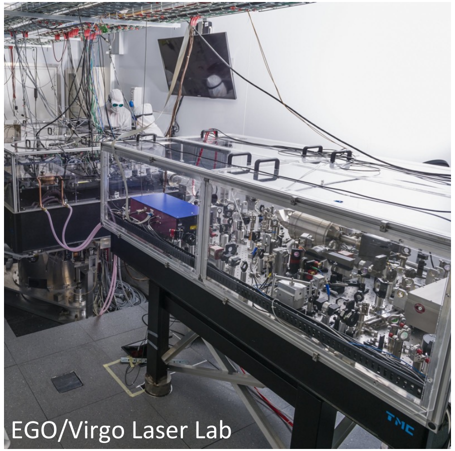
- Extrapolation of current or planned technologies for Virgo and LIGO
  – Squeezing (quantum states of light)
  – High power lasers
  – Bigger mirrors
  – New thin films for mirrors
  – Thermal compensation techniques
  – Seismic suspension systems
- Technologies not tested in Virgo and LIGO (prototypes and/or R&D in progress)
  – Cryogenics (=> KAGRA)
  – New cryogenic materials
  – New laser wavelengths

⇒ Implementation of R&D programs







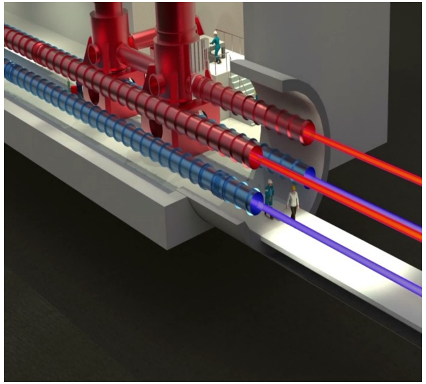

EGO/Virgo Laser Lab

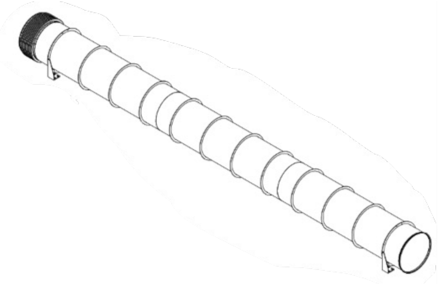
Vacuum pipe system

Vacuum systems planned for 3G detectors are likely to be the largest UHV systems built

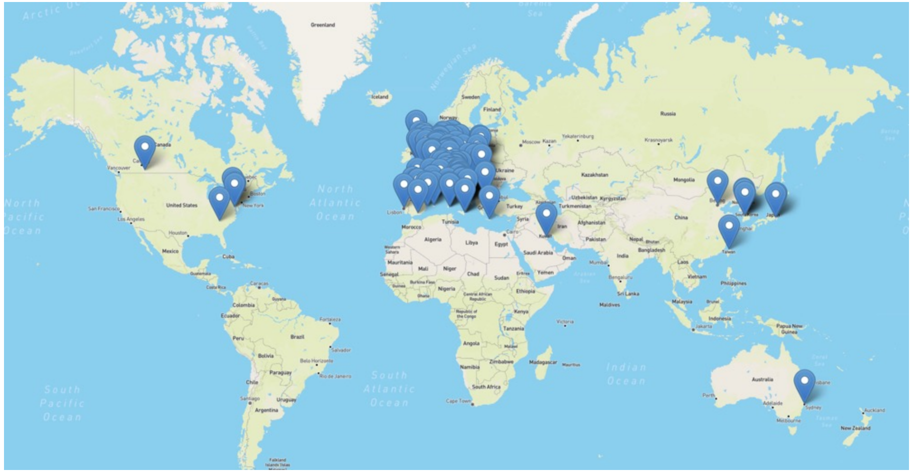
- The beam tube is its most important component (~1/2 of the cost of the system)
- 120 km of UHV tubes of 1 m diameter, total volume ~10^5 m3
- Vacuum requirements: factor > 5 stricter than Virgo:
- 10^{-10} mbar for H2
- 10^{-11} mbar for N2
- < 10^{-14} mbar for hydrocarbons
- Lifespan: 50 years
- Preliminary estimated cost ~560 M€

Joint development with CERN involving Einstein Telescope and Cosmic Explorer (US Project)





ET Collaboration



1370 collaborators from 202 institutes in 23 countries worldwide

Birth of the ET Collaboration
in June 2022 in Budapest

ET Infrastructure and ET Collaboration

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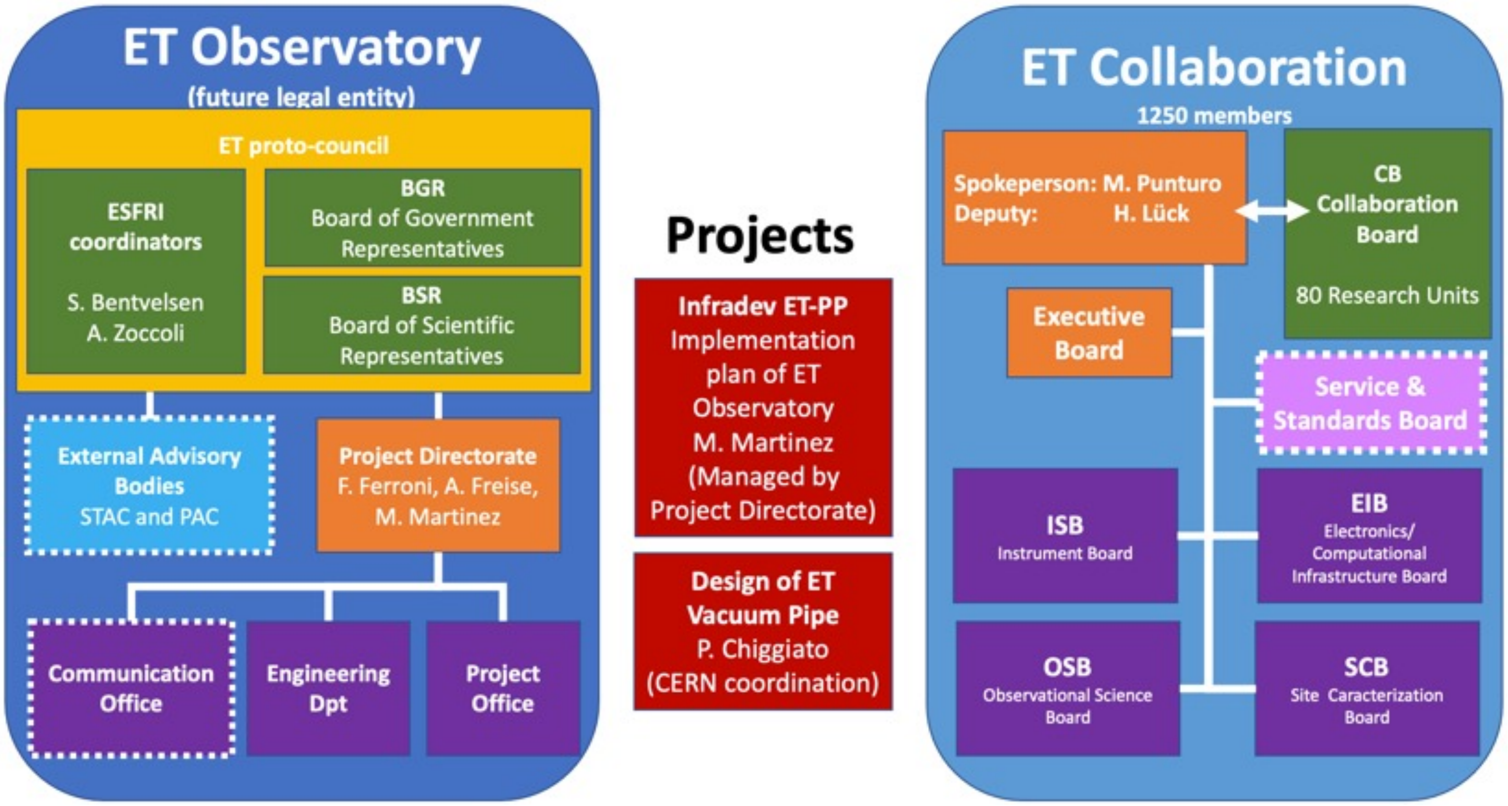
The EU supports the creation of the ET infrastructure (ETO) through the financing of an Infradev project:

Einstein Telescope Preparatory Phase (ET-PP)



Since the summer of 2022, the ET structures are being put in place

Schedule

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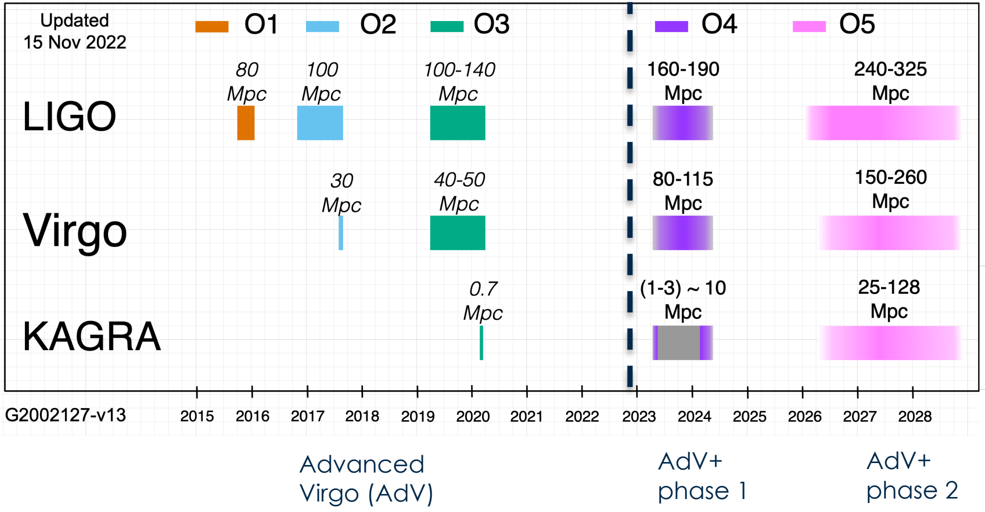

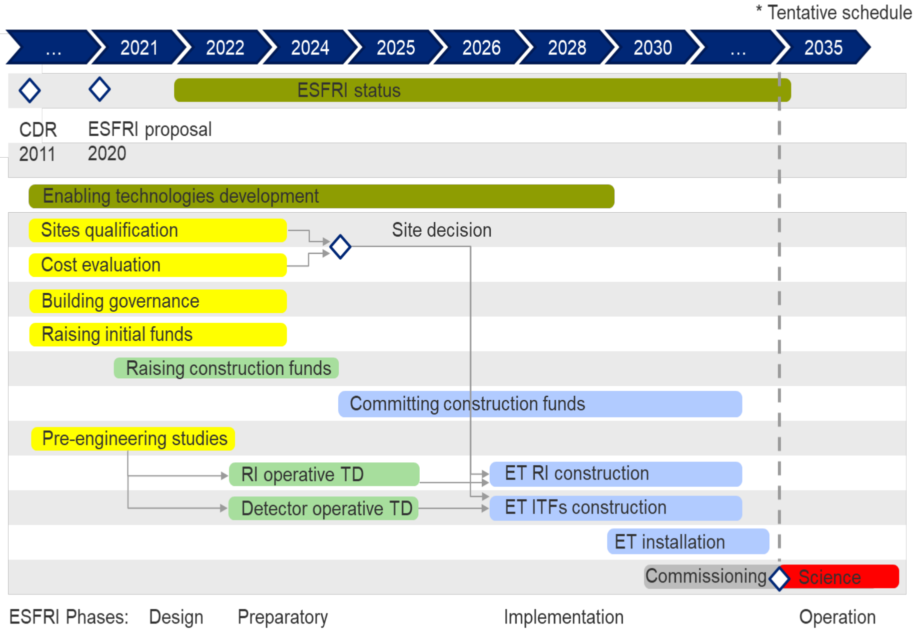

Virgo and LIGO are currently studying post-O5 upgrades, Virgo_nEXT and A#

Site candidates in Europe

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There are currently two candidate sites in Europe to host ET:

- The Sardinia site, near the Sos Enattos mine
- The EU Regio Rhine-Meusse site, close to the NL-B-D border
- A third option in Saxony (Germany) is under discussion

Sites are studied through

- seismic noise measurements at surface, borehole and mine (Sardinia)
- Magnetic and ambient noise measurements
- Geophysical and geotechnical characterizations
- …

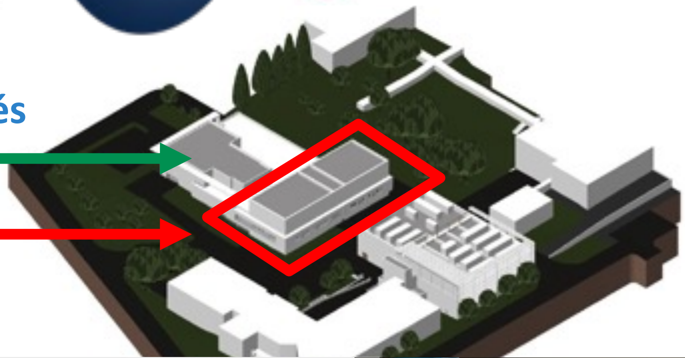
Significant funds are required to develop a site application

Extension of the LMA in Lyon

Laboratoire des Matériaux Avancés

Current

Extension



Contrat de Plan Etat-Région 2021-2027

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RÉPUBLIQUE FRANÇAISE

La Région Auvergne-Rhône-Alpes

RHÔNE LE DÉPARTEMENT

MÉTROPOLE GRAND LYON

An investment for:

- the extension of the LMA building
- the construction of a new coater allowing the deposition of thin films on very large substrates: Ø 1.6 m, 600 kg
- associated optical and metrology tools

LMA is a research infrastructure unique in the world

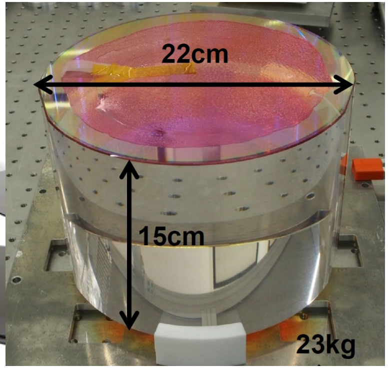
Developing technologies for future experiments (e.g. Einstein Telescope)

Mirror substrates

At cryogenic temperature, silica is no longer compatible

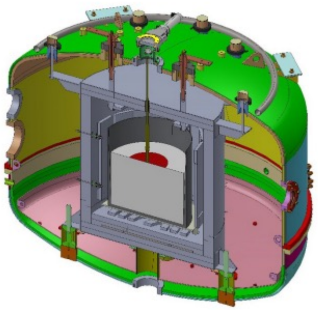
- Crystalline substrates:
  - silicon (λ > 1300 nm)
  - Sapphire (λ > 400 nm)
- R&D on low T thin films
- Must be available in Ø 450 mm, 200 kg and with good optical properties

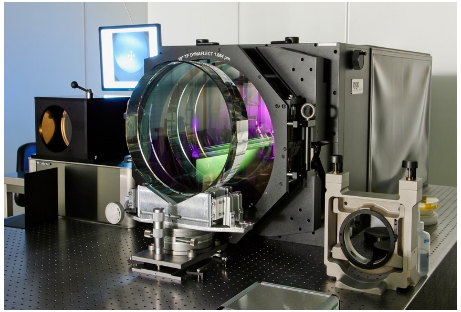
Today, no solution for such substrates!



Project in Lyon: *Sapphire Optics for Gravitational Astronomy*

- Funded by Université de Lyon
- Sized to make substrates of Ø 450 mm, 200 kg
- Production on Ø 200 mm in progress
- Combined with LMA : measurement of optical absorption birefringence measurement of surface quality





Work in progress: stay tuned !

Conclusion

Einstein Telescope is a very ambitious project which requires:

- An intensive R&D program to develop new technologies
- A large, structured and organized scientific collaboration
- Strong partnerships between academic research and industries

There are important synergies between Einstein Telescope, existing 2G detectors (LIGO-Virgo-KAGRA) and future detectors (Cosmic Explorer)