First neutrino interaction recorded in hydrogen bubble chamber

NEUTRINO PHYSICS: EXPERIMENTAL ASPECTS

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Overview

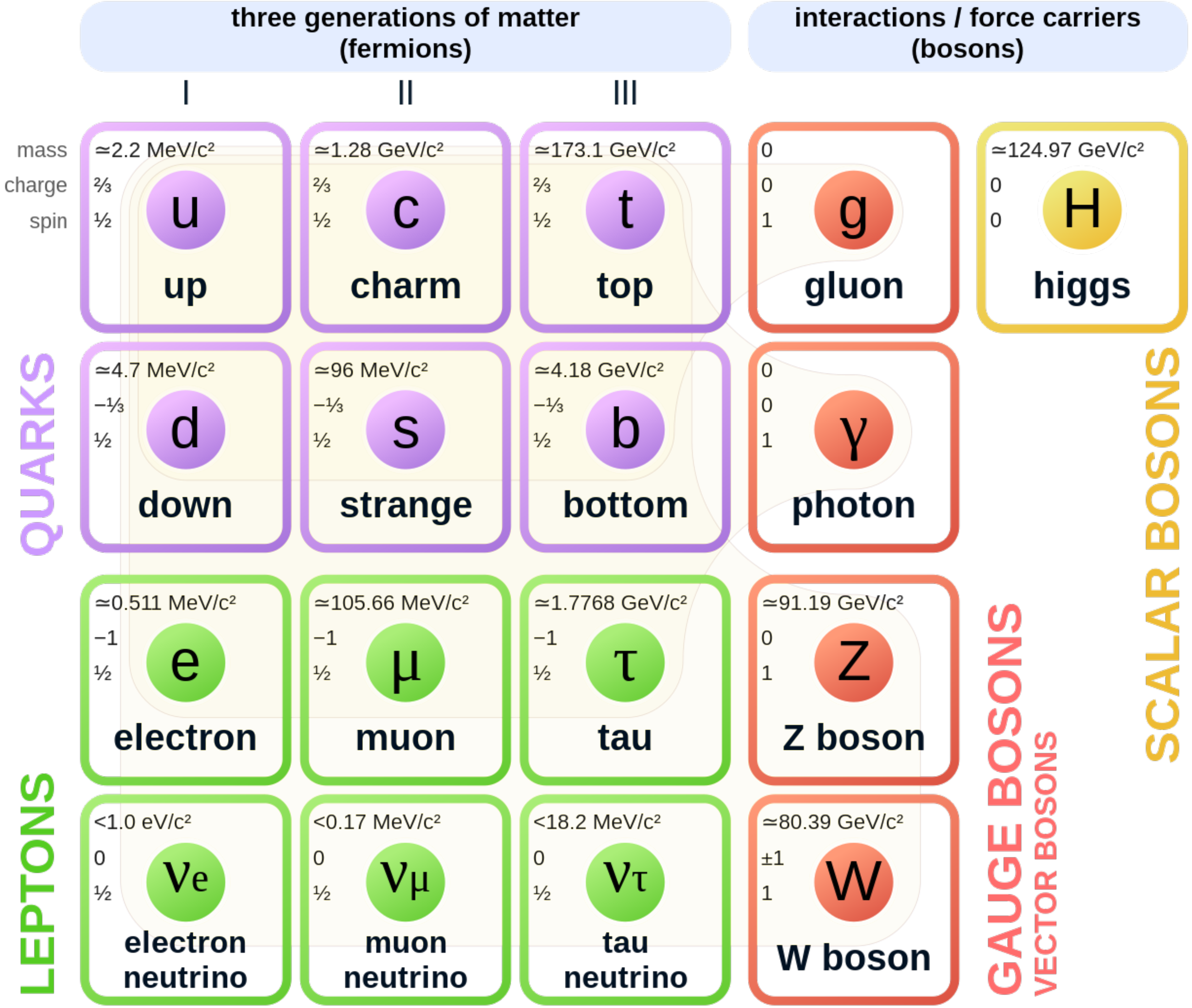
○ Recap on neutrino properties, sources, interactions

○ Neutrino Oscillation with a « historical » approach

○ Produce, Detect and neutrino oscillations today & tomorrow

In the Standard Model

- Neutrinos are leptons
- 3 flavors linked to their corresponding charged counterpart
- Can only interact through weak force (through W± and Z^0 bosons)

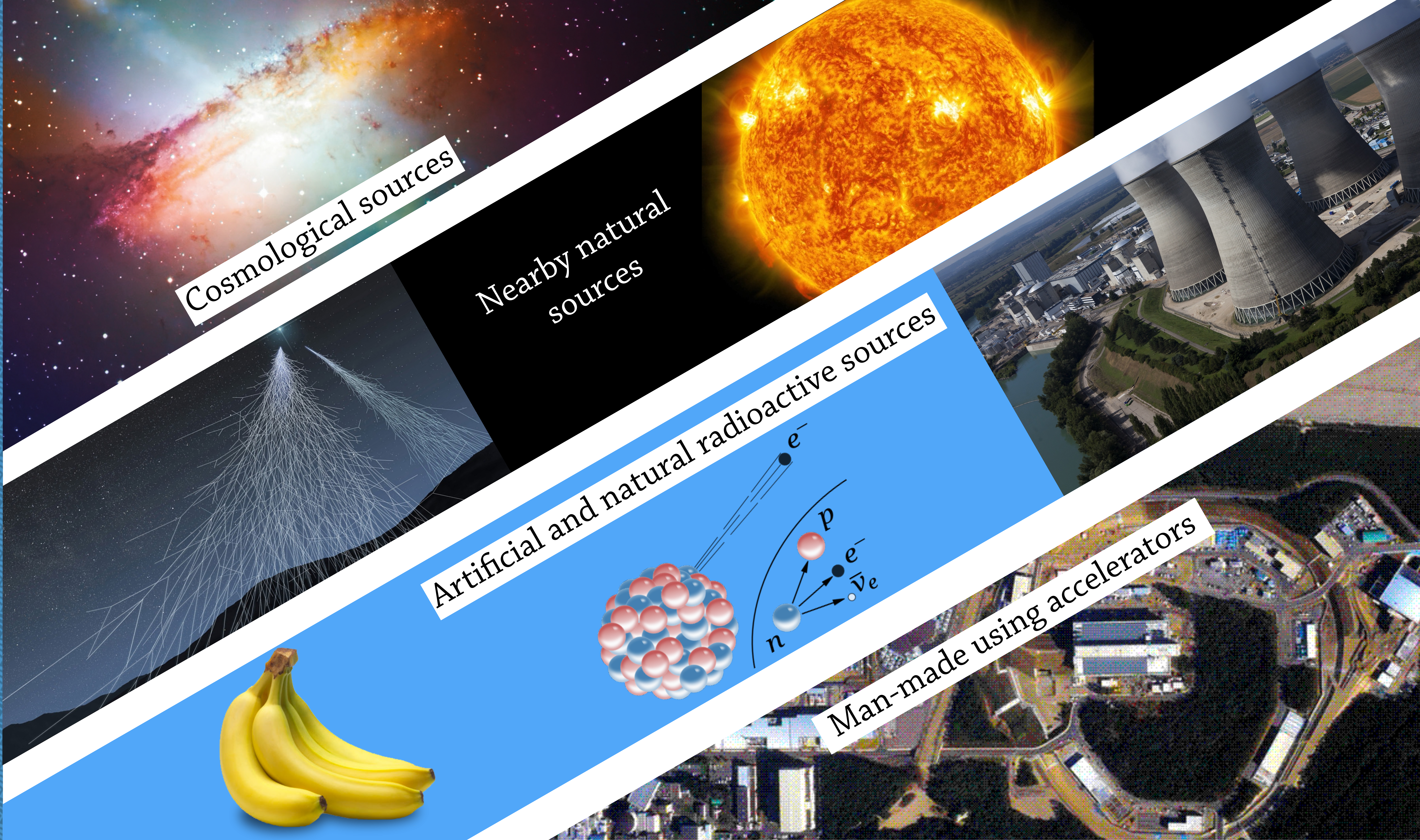


three generations of matter (fermions)

interactions / force carriers (bosons)

		I	II	III		
QUARKS	mass charge spin	≃2.2 MeV/c² ⅔ ½ u **up**	≃1.28 GeV/c² ⅔ ½ c **charm**	≃173.1 GeV/c² ⅔ ½ t **top**	0 0 1 g **gluon**	≃124.97 GeV/c² 0 0 H **higgs**
		≃4.7 MeV/c² −⅓ ½ d **down**	≃96 MeV/c² −⅓ ½ s **strange**	≃4.18 GeV/c² −⅓ ½ b **bottom**	0 0 1 γ **photon**	SCALAR BOSONS
LEPTONS		≃0.511 MeV/c² −1 ½ e **electron**	≃105.66 MeV/c² −1 ½ μ **muon**	≃1.7768 GeV/c² −1 ½ τ **tau**	≃91.19 GeV/c² 0 1 Z **Z boson**	
		<1.0 eV/c² 0 ½ ν_e **electron neutrino**	<0.17 MeV/c² 0 ½ ν_μ **muon neutrino**	<18.2 MeV/c² 0 ½ ν_τ **tau neutrino**	≃80.39 GeV/c² ±1 1 W **W boson**	GAUGE BOSONS VECTOR BOSONS

Neutrino Sources



Cosmological sources

Nearby natural sources

Artificial and natural radioactive sources

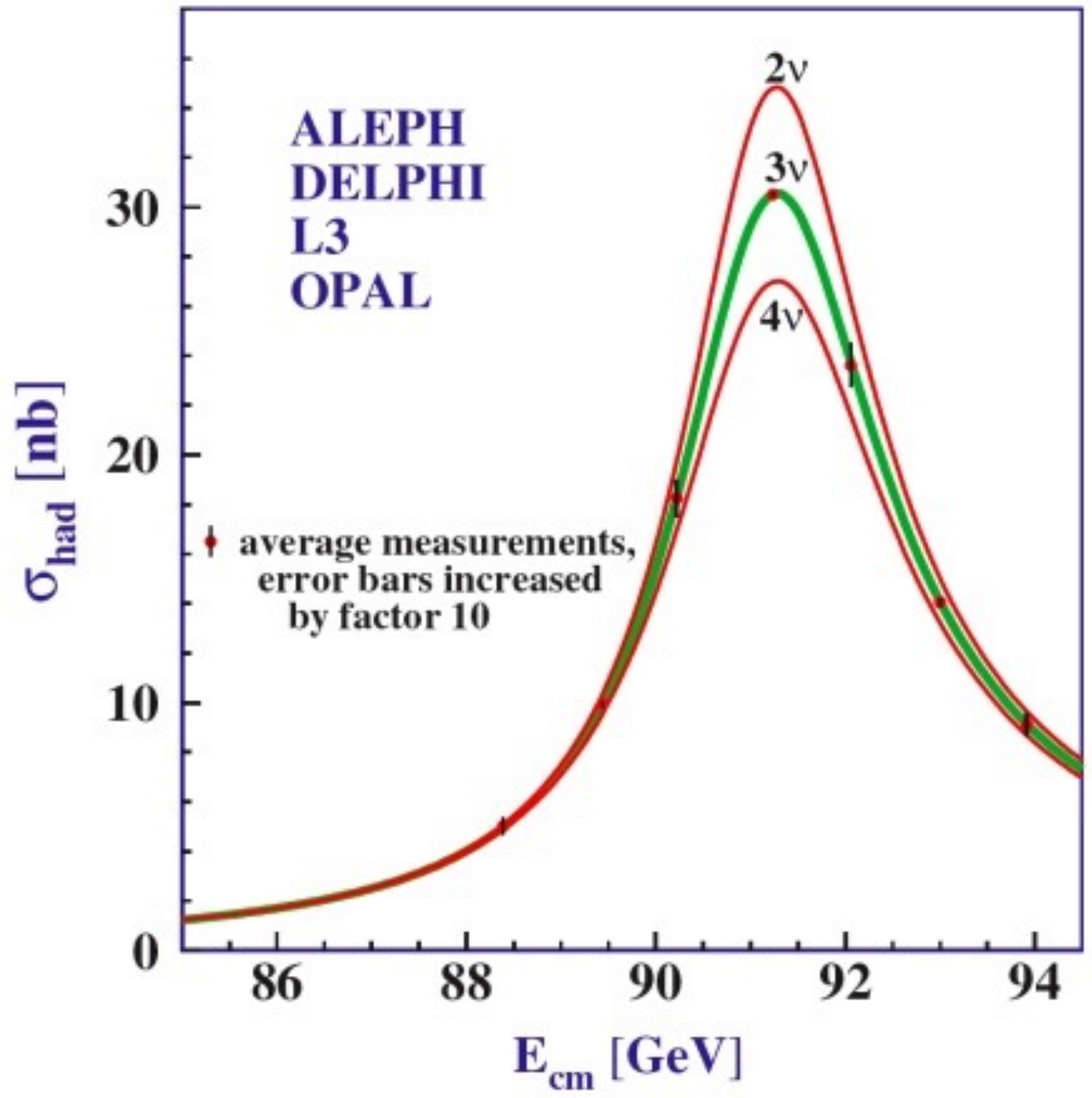
Man-made using accelerators

Key facts

○ Three flavors of light and active neutrinos named ν_e, ν_μ, ν_τ

→ In 1989, LEP mesures the Z invisible decay width :

$$N_\nu = 2.984 \pm 0.008$$



○ Neutrinos are only left-handed

→ Cannot couple to the Higgs field, therefore the neutrinos are considered massless in the Standard Model

→ But they in fact do have a mass:

$$m_\nu < 1 \text{ eV} \ ; \ \sum m_\nu > 0.06 \text{ eV}$$

○ Most abundant massive particule

$\Phi_{sun} = 65\times10^9\ \nu_e/cm^2/s$ on earth

$\Phi_{reactor} = 2\times10^{20}\ \bar{\nu}_e/s/GW_{th}$

$\Phi_{atmo} = 4\times10^2\ \nu_{e+\mu}/m^2/s/sr$

$\Phi_{accelerator} \sim 1\times10^{12}\ \nu_\mu/m^2$

○ Only interact through weak interaction

→ Small cross section :

$\sigma \sim 10^{-42}\ cm^2$ for IBD

$\sigma \sim 10^{-38}\ cm^2$ at 1 GeV

→ 50% chance a ν_e from the sun interact in you in your lifetime

HOW TO DETECT NEUTRINOS

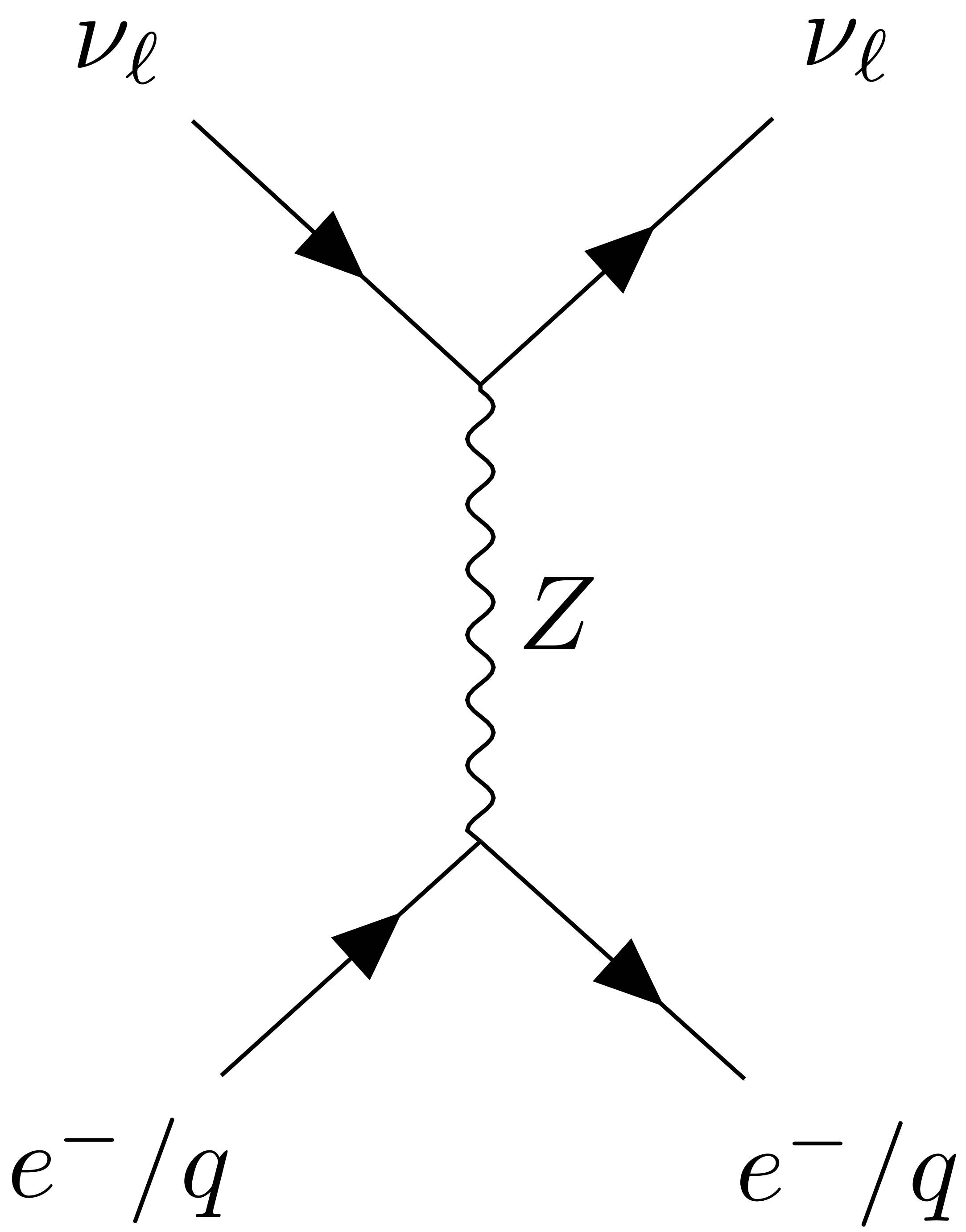
Charged and neutral currents

Neutrino have no electric charge -> We cannot detect them directly

We have to :
- Wait for a neutrino to interact
- Detect the products of the interaction
- Retrieve the original neutrino flavor/direction/energy/sign

Neutrinos can only interact by weak interaction
-> Through Z^0 exchange = Neutral Currents
-> Through $W^{\pm}$ exchange = Charged Currents

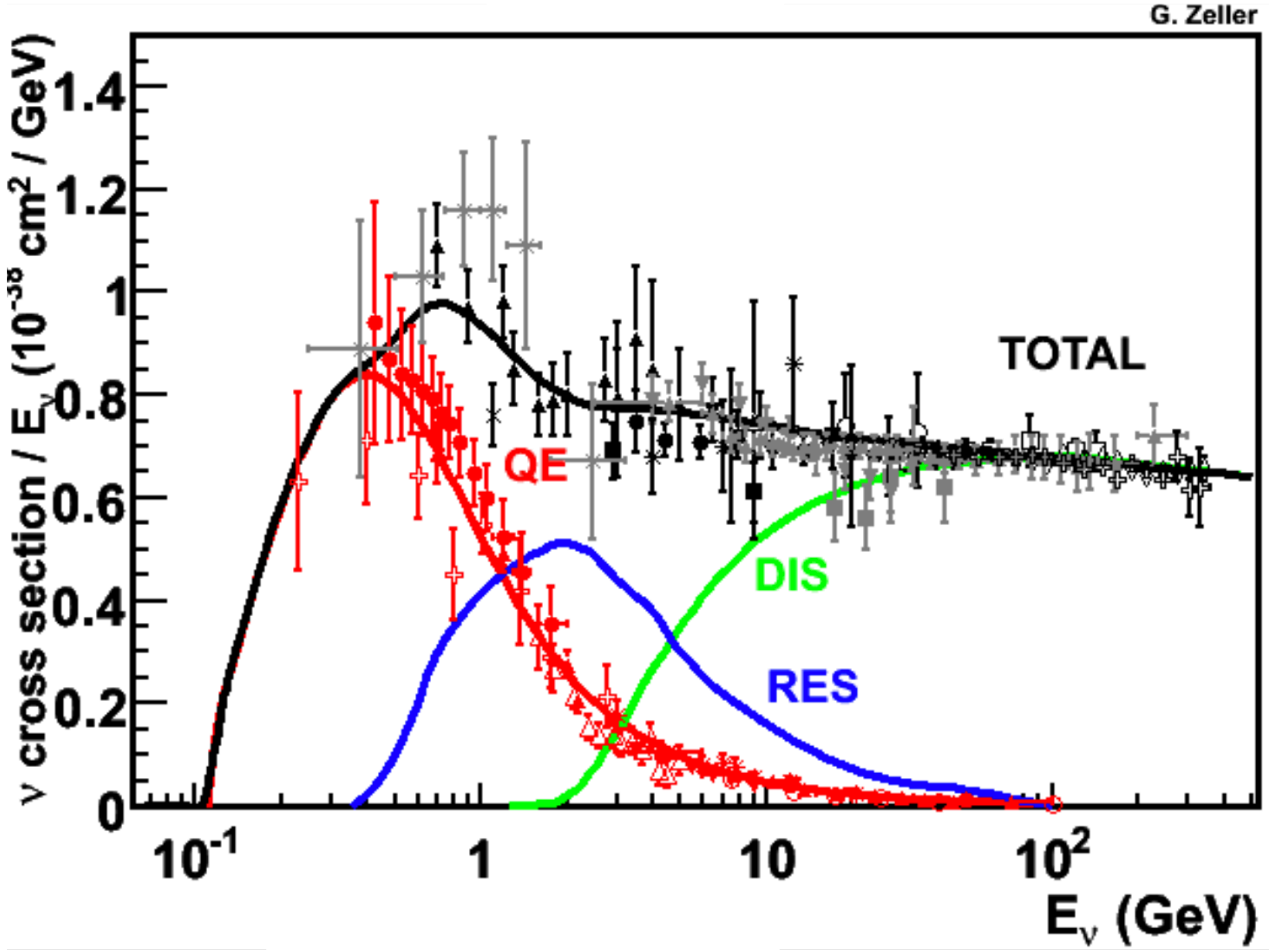
Neutral currents interactions



Elastic scattering

- ○ Cannot identify the incoming neutrino flavor
- ○ All neutrino interact with same potential

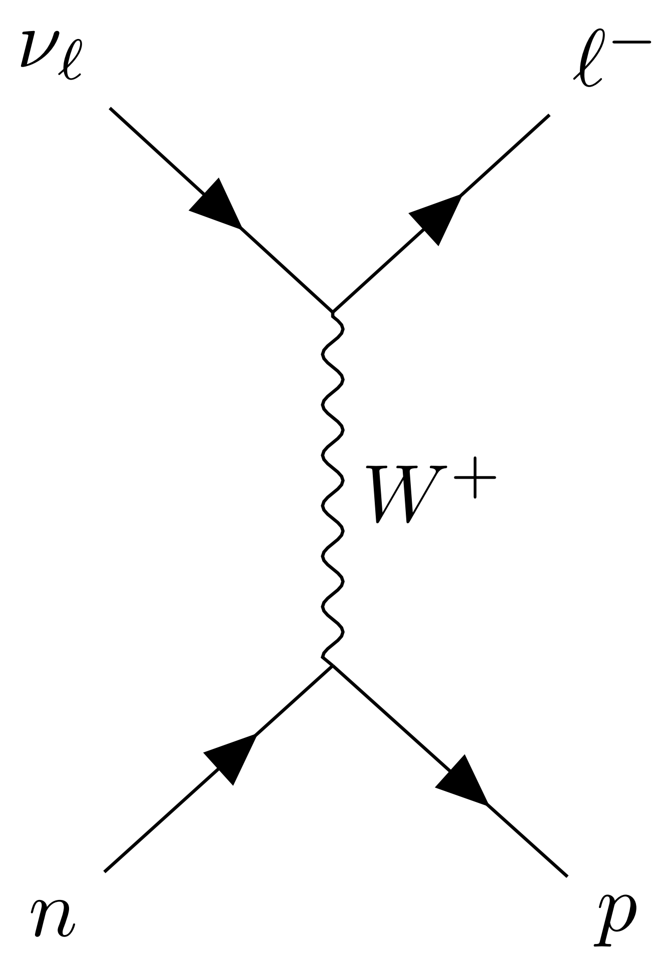
Charged currents interactions

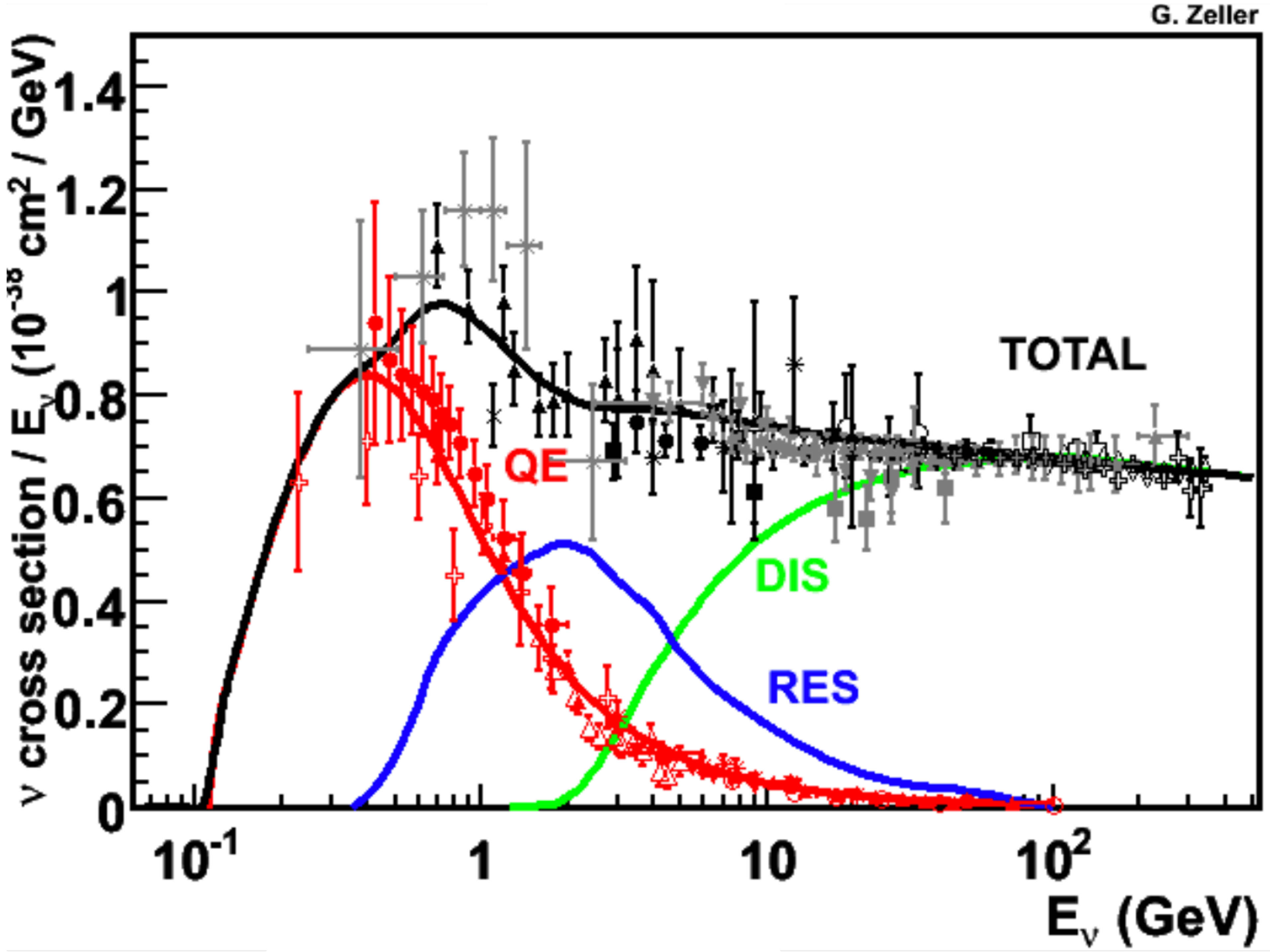


Charged currents interactions

The Quasi-Elastic interaction

The Golden Channel





- ○ ν flavor & sign tagged by the lepton
- ○ E_ν reconstructed with the lepton kinematics :

$$E_\nu = \frac{m_f^2 - (m_i - E_b)^2 - m_\mu^2 + 2(m_i - E_b)E_\mu}{2(m_i - E_b - E_\mu + p_\mu \cos\theta_\mu)}$$

m_i, m_f : initial, final nucleon masses;
E_b : nucleon binding energy in the nucleus

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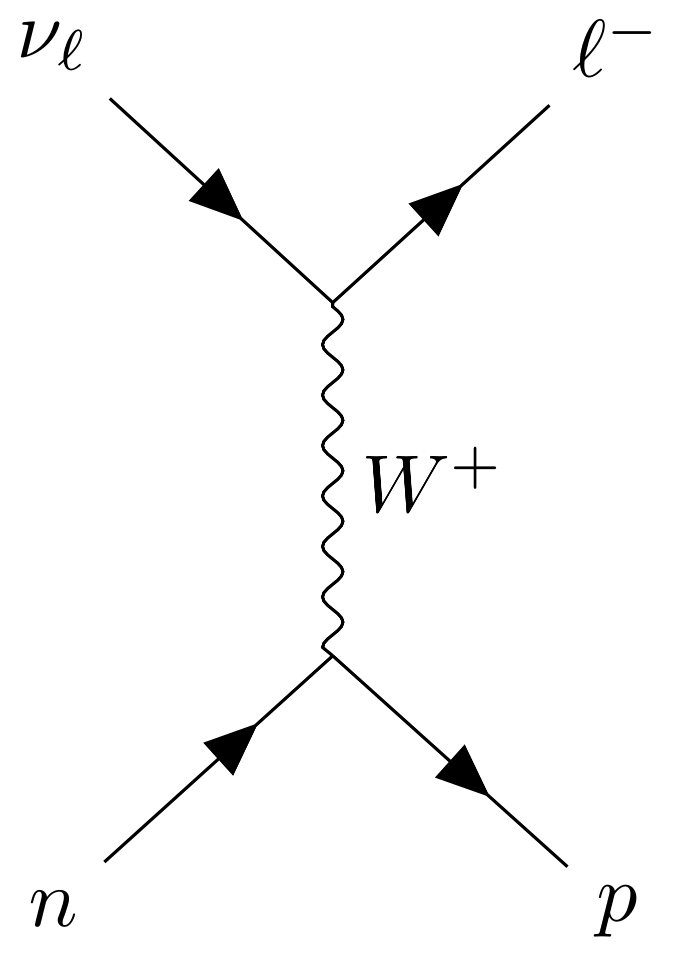

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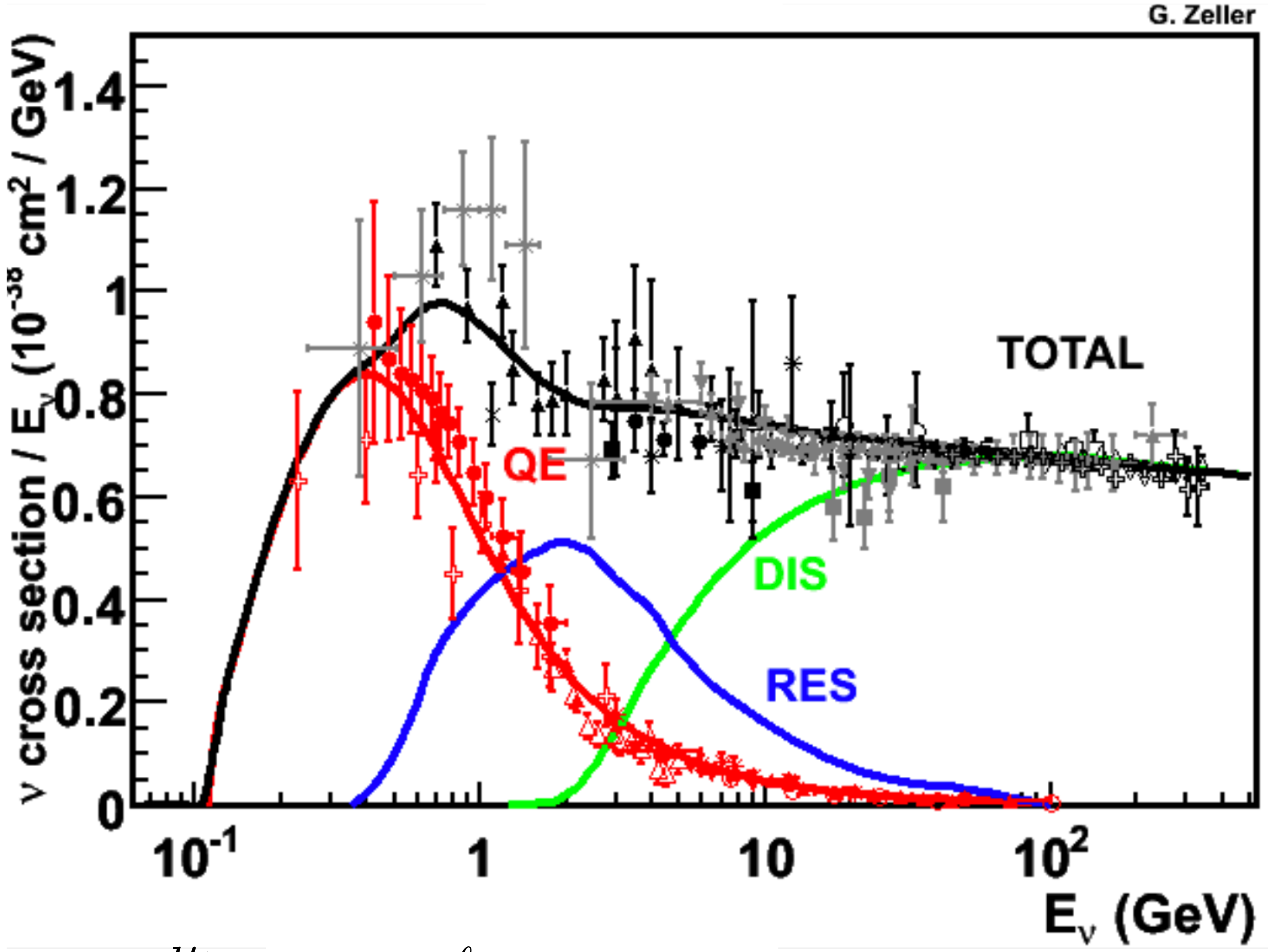

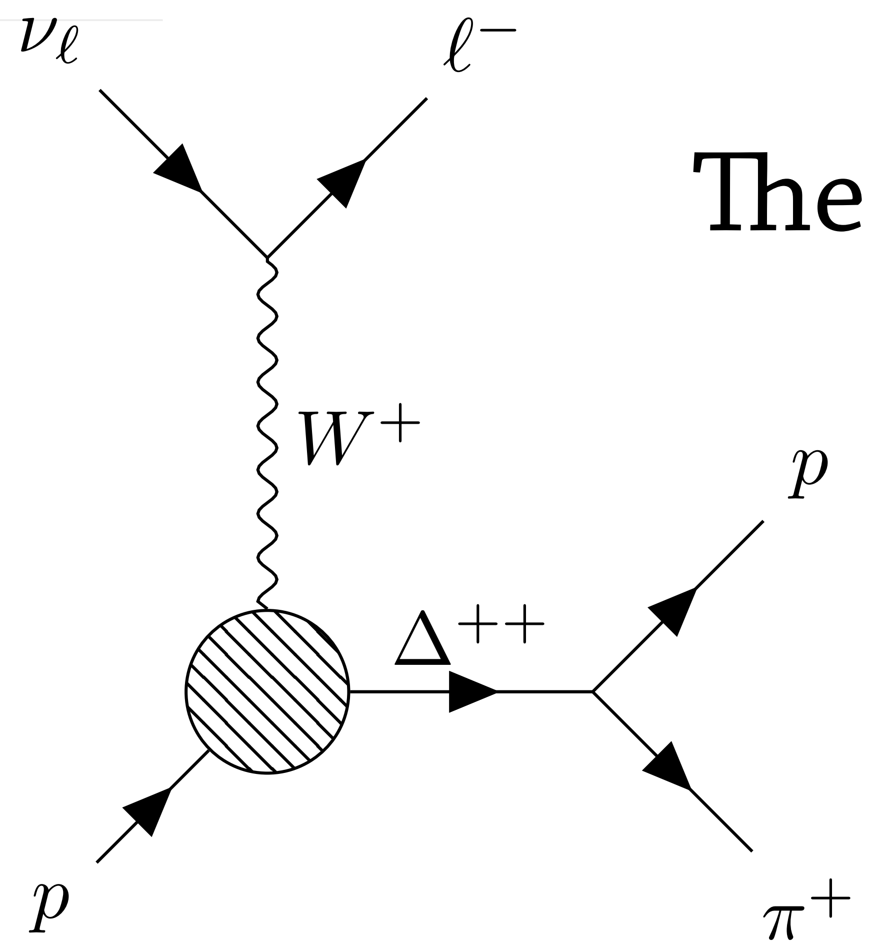

The Resonant interaction

Nucleon is excited
-> Many final states

Charged currents interactions

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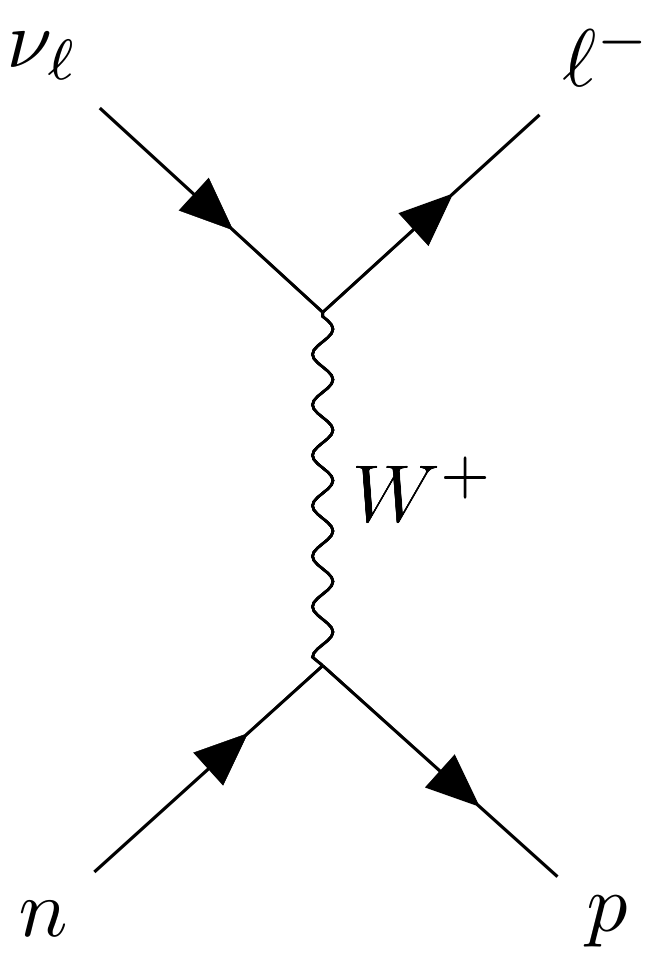

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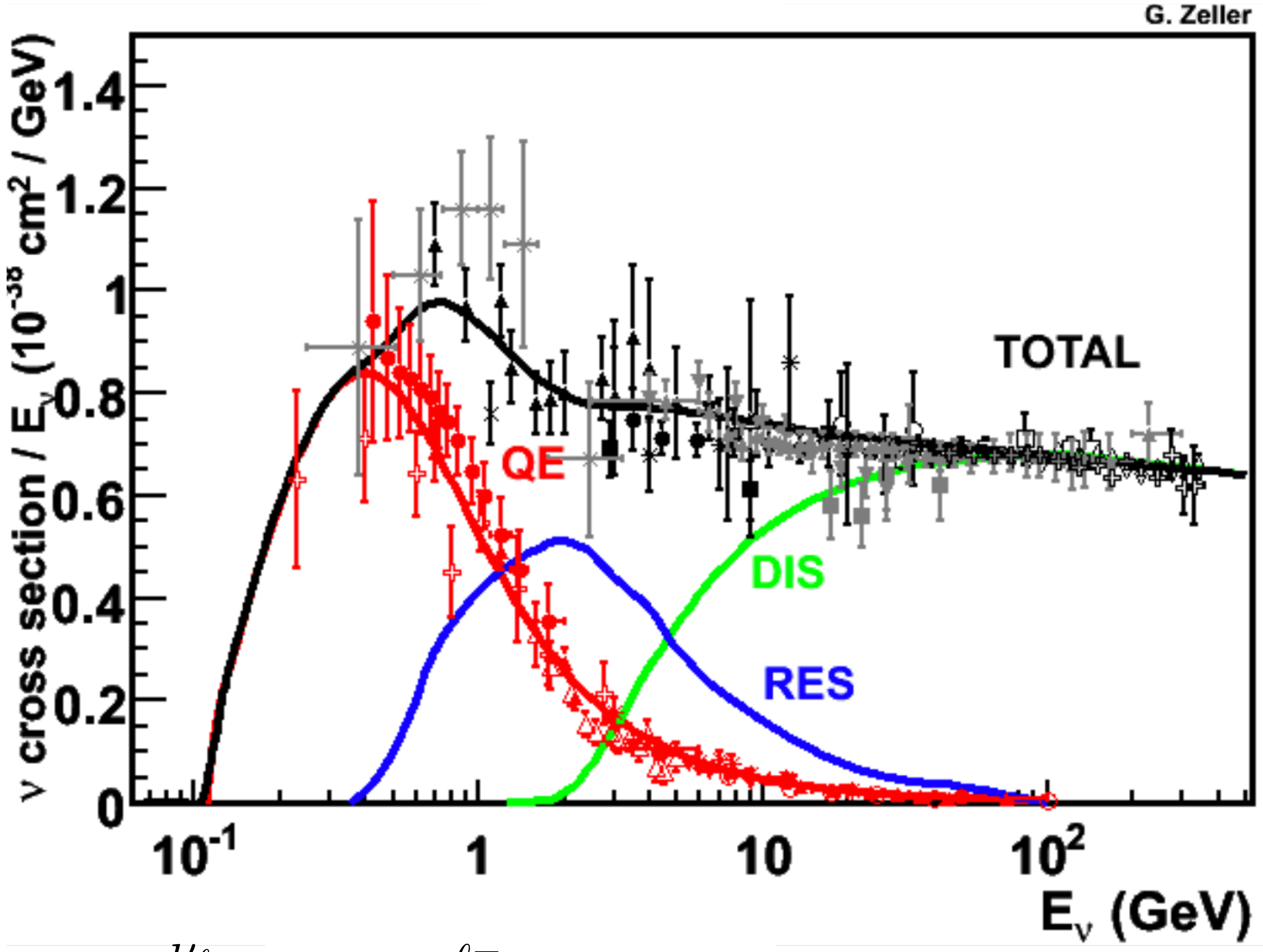
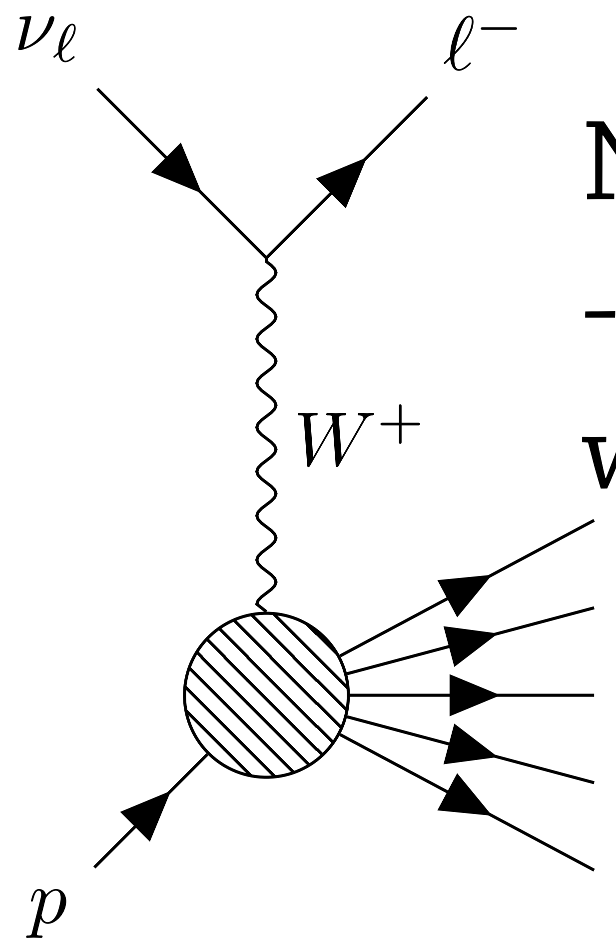
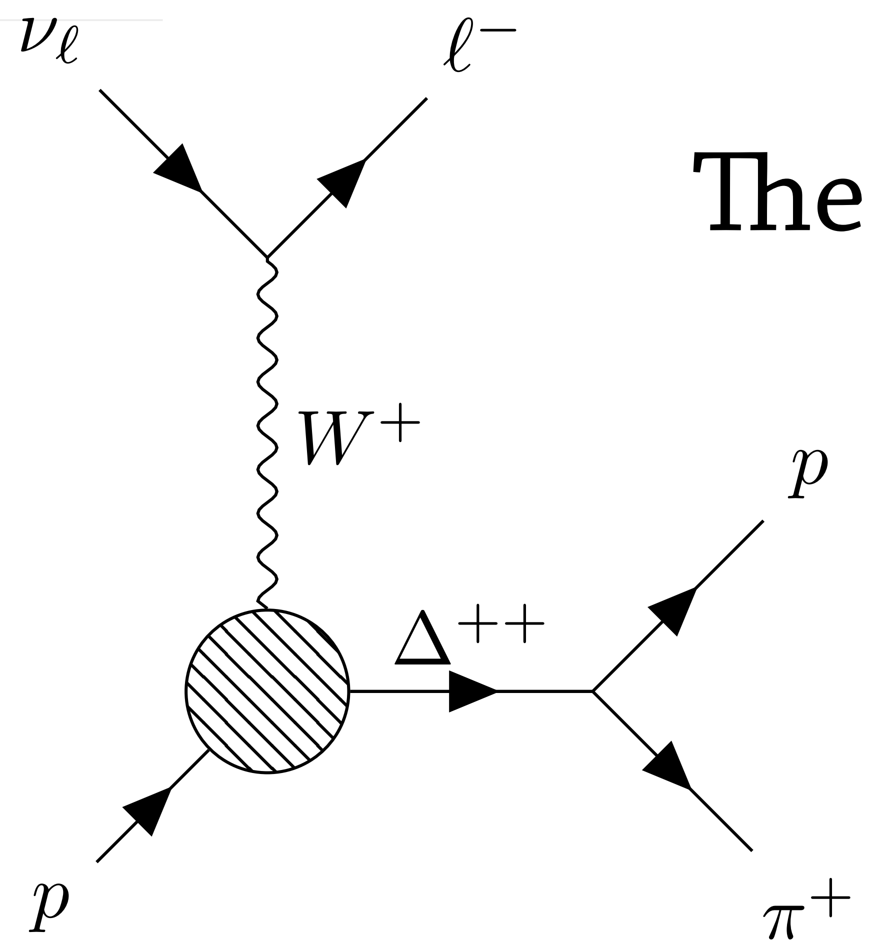


The Deep-Inelastic interaction



Nucleon breaks
-> Interactions with the quarks

The Resonant interaction



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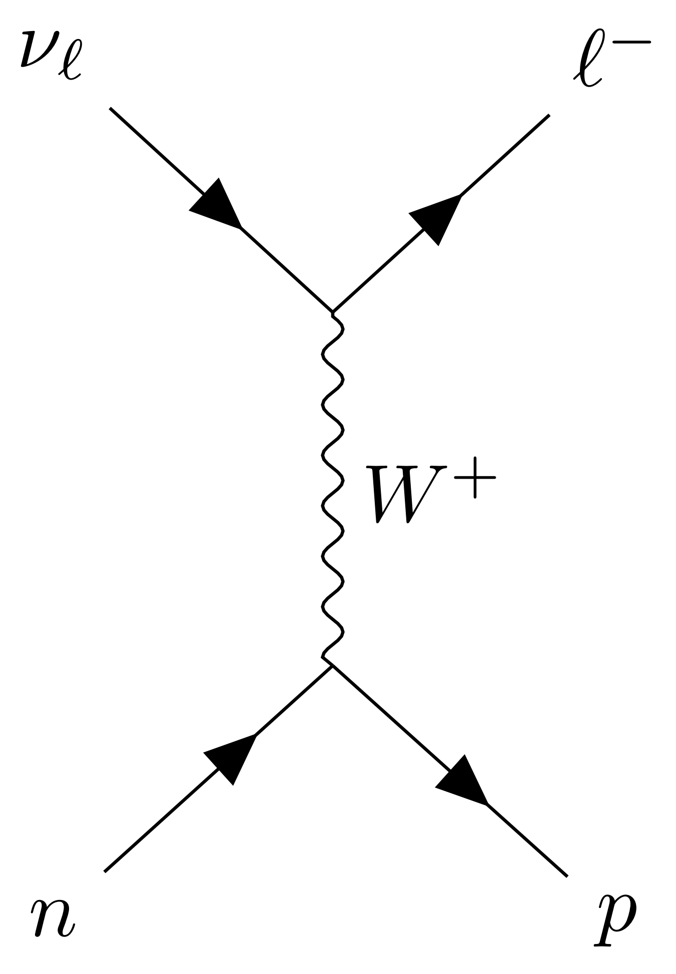

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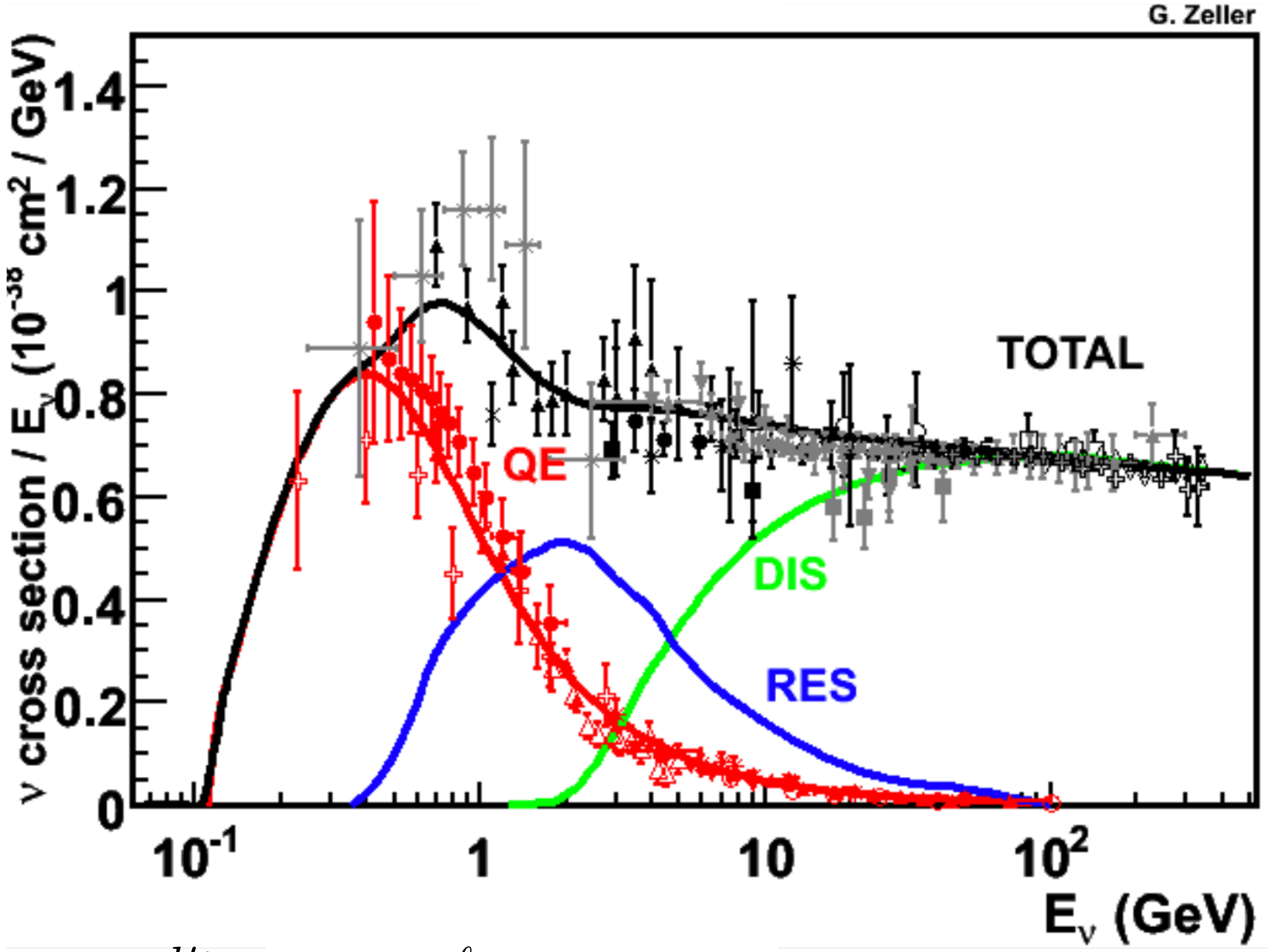

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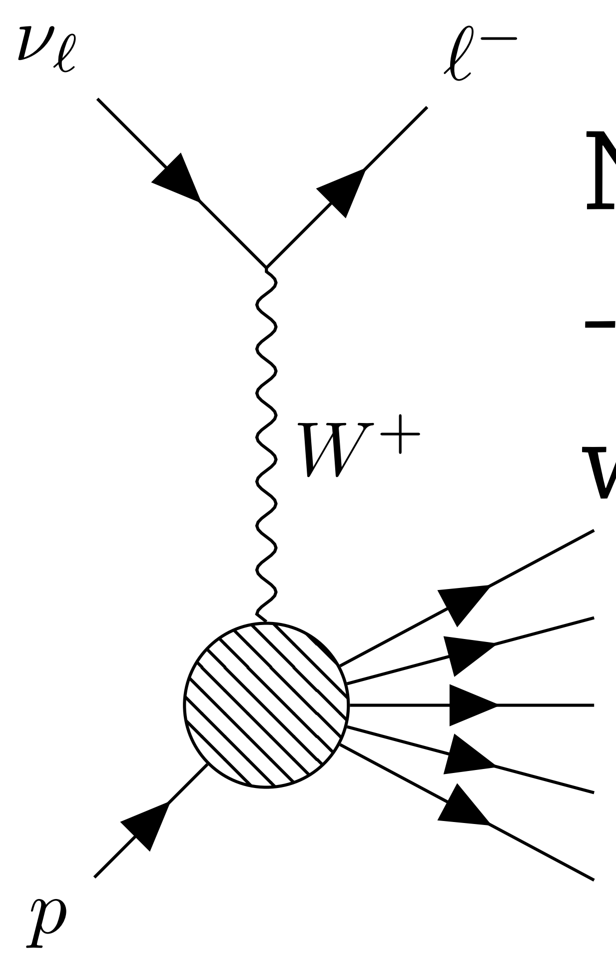

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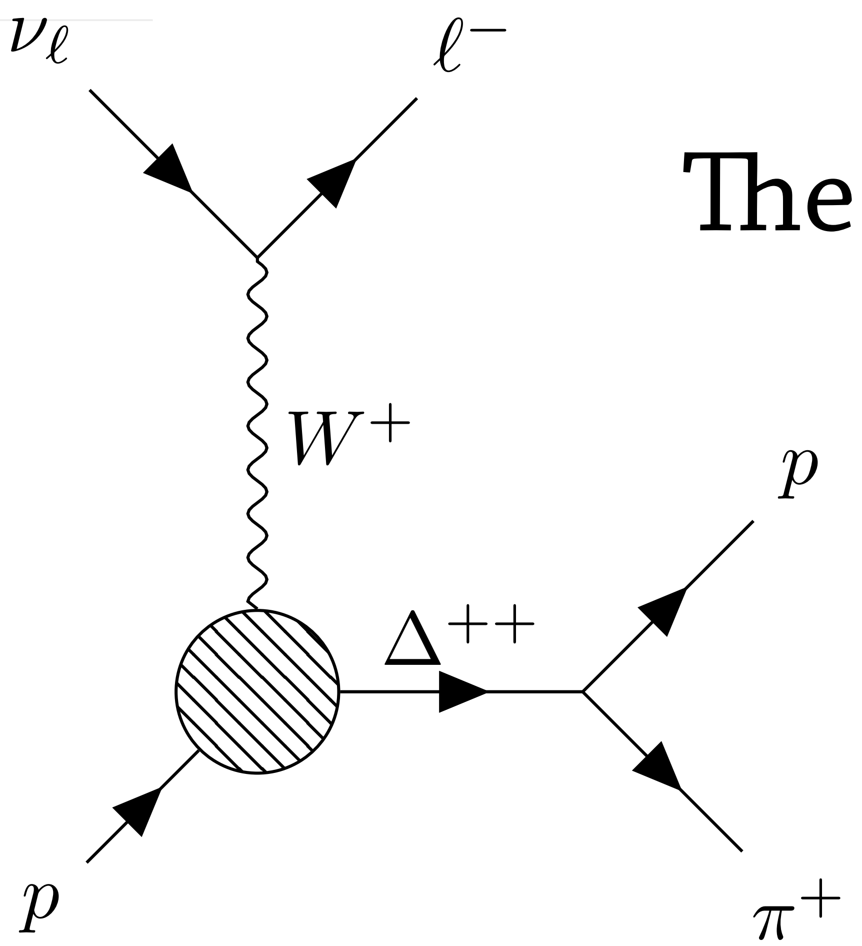


Nucleon is excited
-> Many final states

Cross section increase with energy but the final states are more complex

Charged currents interactions

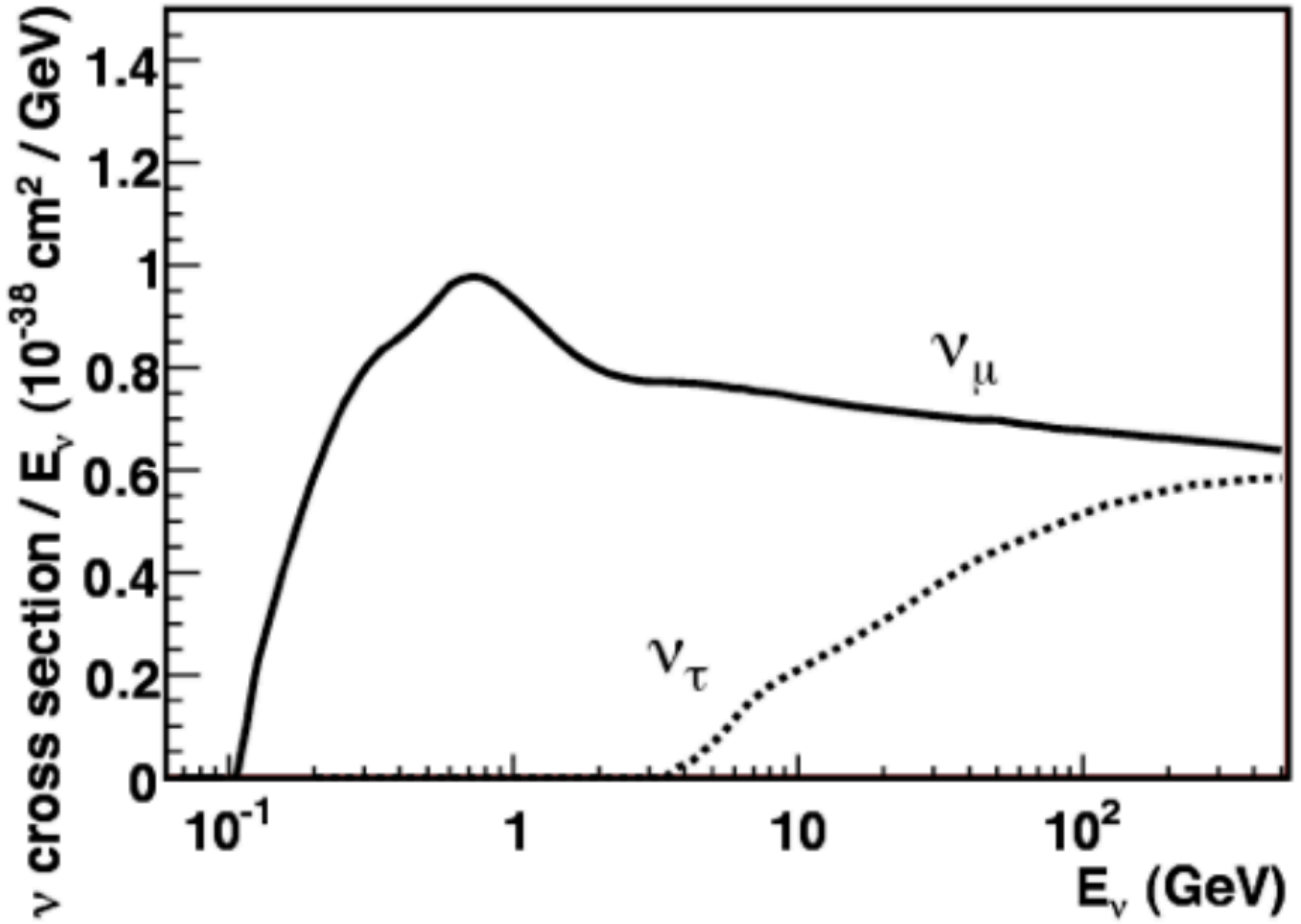
Interaction threshold

For CC interactions : $\nu_\ell + n \rightarrow \ell^- + p$

$$E_\nu \geq \frac{(m_\ell + m_p)^2 - m_n}{2m_n}$$

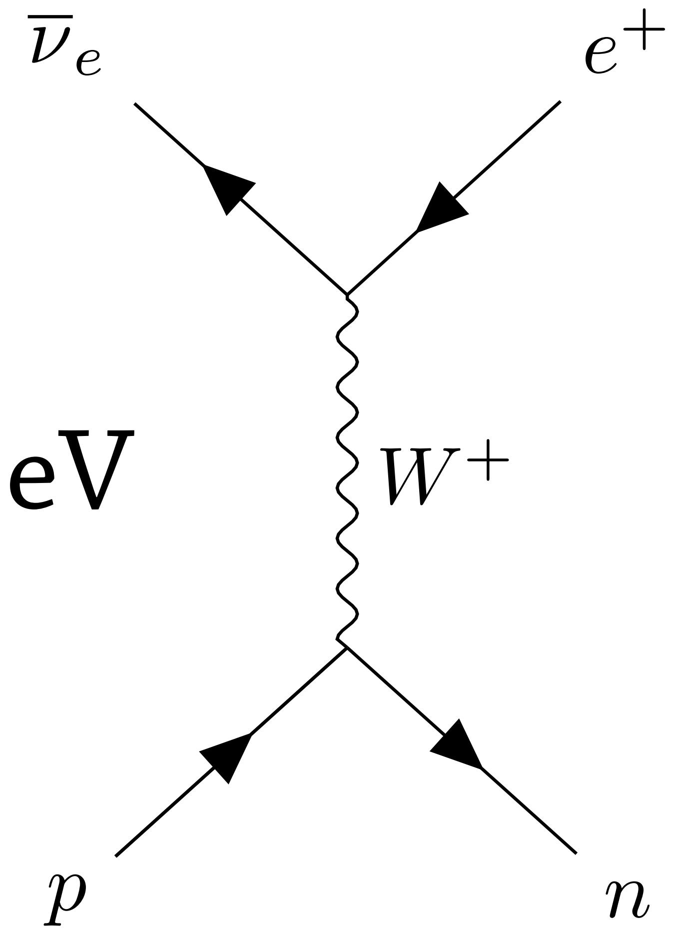
$E_{thr}(\nu_\mu)$ = 110 MeV

$E_{thr}(\nu_\tau)$ = 3.45 GeV

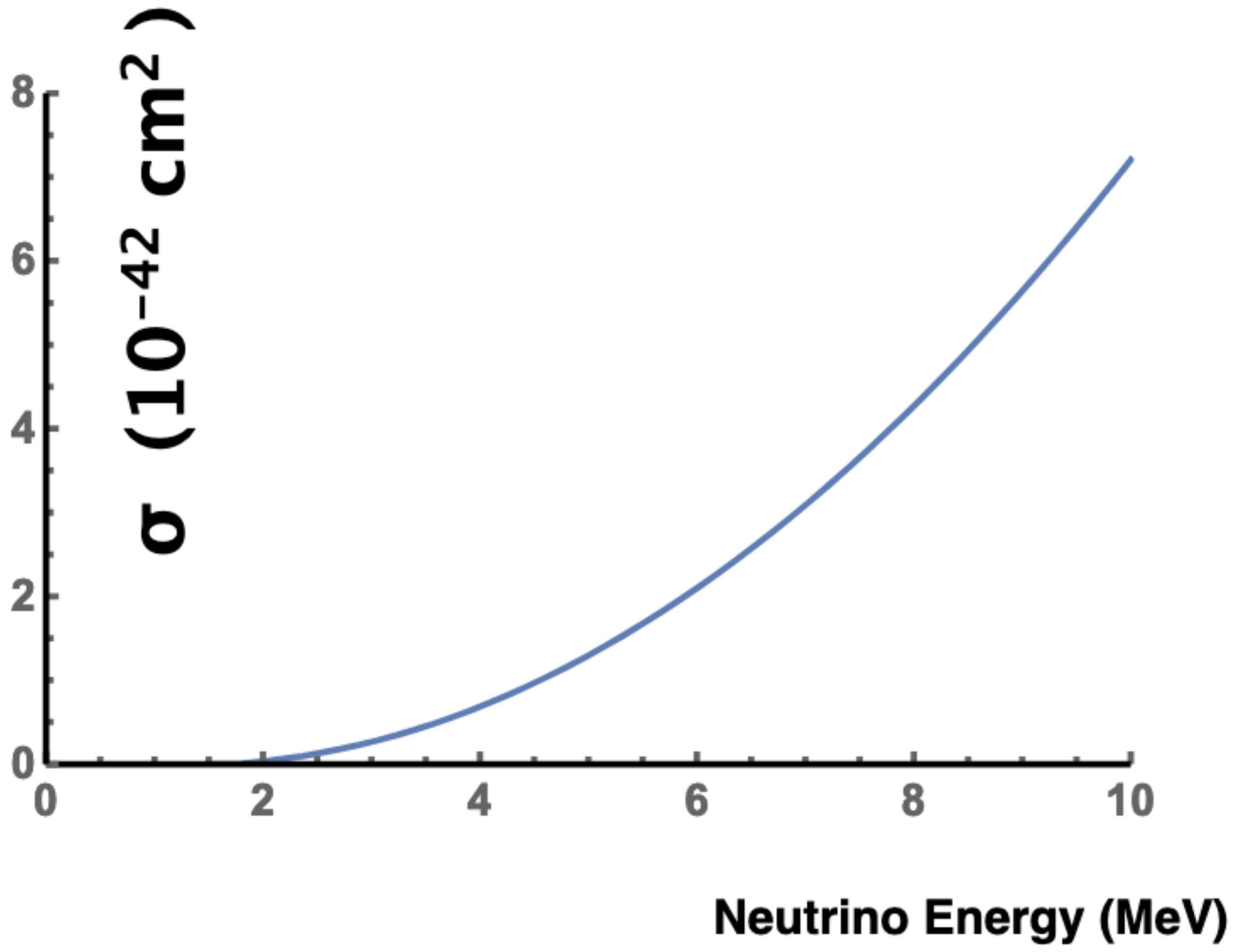


$\bar{\nu}_e$ low energy interaction

inverse β-decay :



$E_{thr}(\bar{\nu}_e)$ = 1.806 MeV

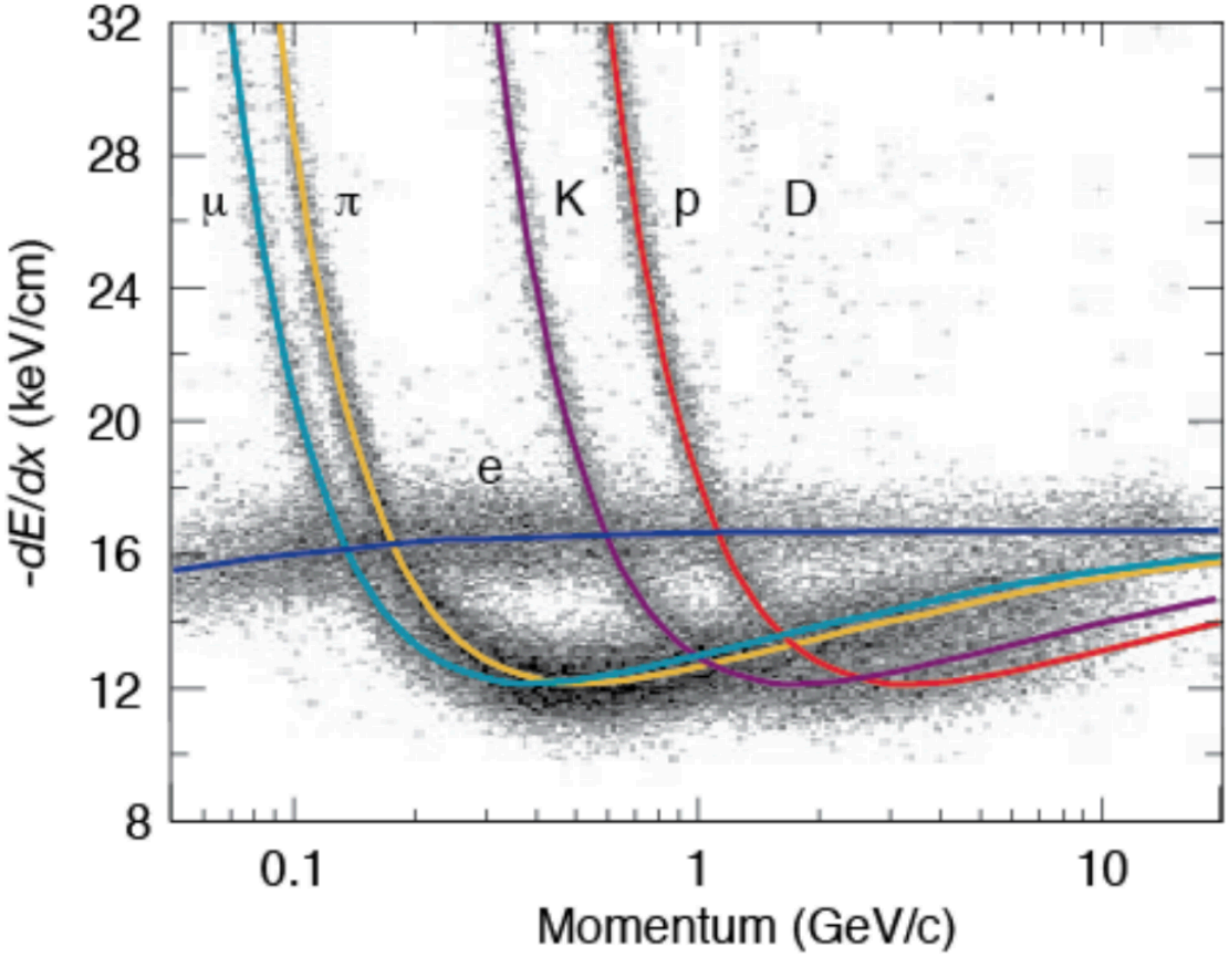


Using the Ionization potential

Principle : When a charged particle crosses a medium, it looses energy through ionization.
The mean amount of energy lost per cm through ionization is parametrized by the Bethe Bloch formula and depends on the particle energy (βγ)

$$\left\langle -\frac{dE}{dx} \right\rangle = Kz^2 \frac{Z}{A} \frac{1}{\beta^2} \left[\frac{1}{2} \ln \frac{2m_e c^2 \beta^2 \gamma^2 W_{max}}{I^2} - \beta^2 - \frac{\delta(\beta\gamma)}{2} \right]$$

PEP-4 detector, gas mixture 80:20 of Ar:CH_4



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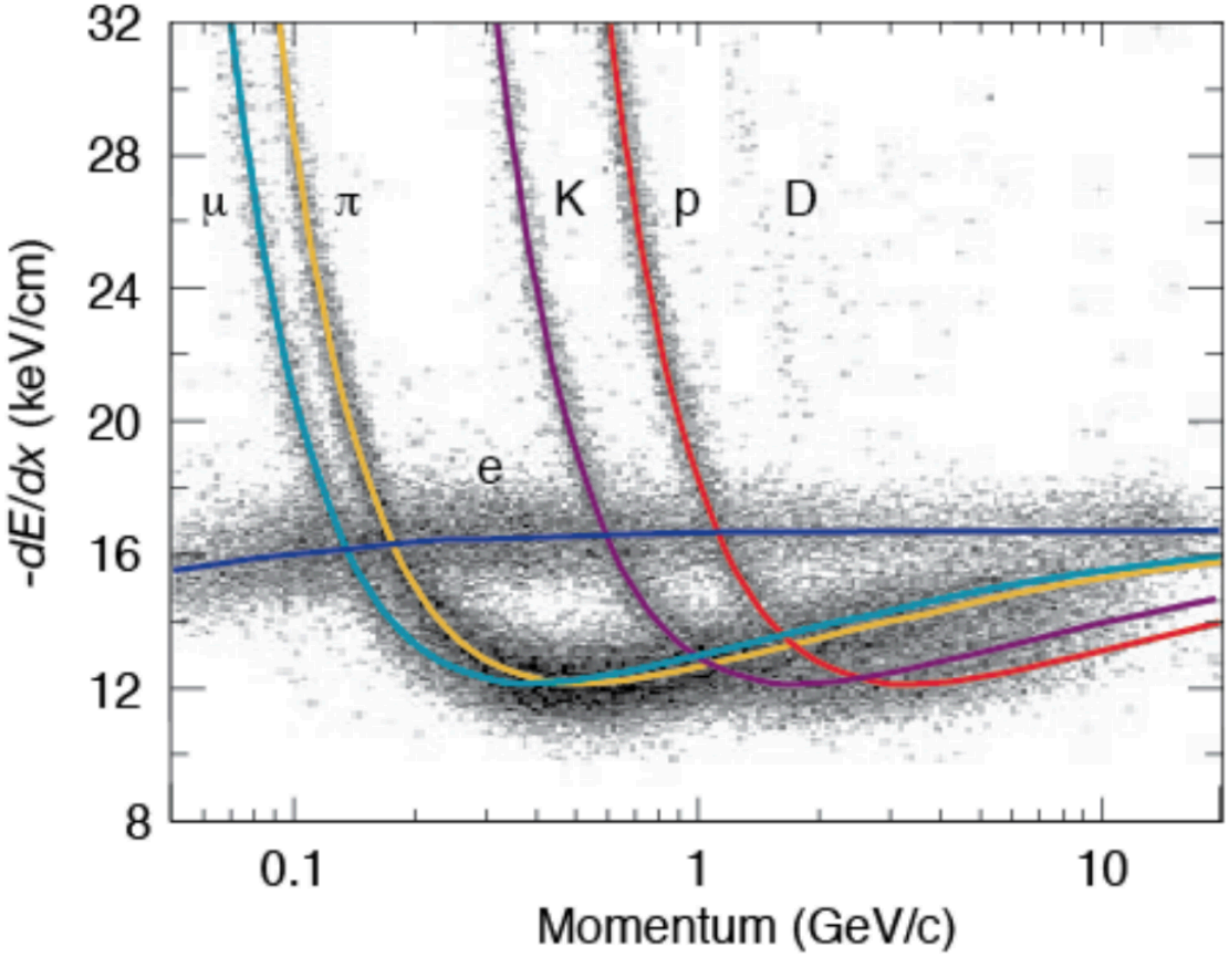


If this energy lost can be seen, one can have a 2D (or even 3D) image of the interaction. Through track topology, one can know the daughters identity.

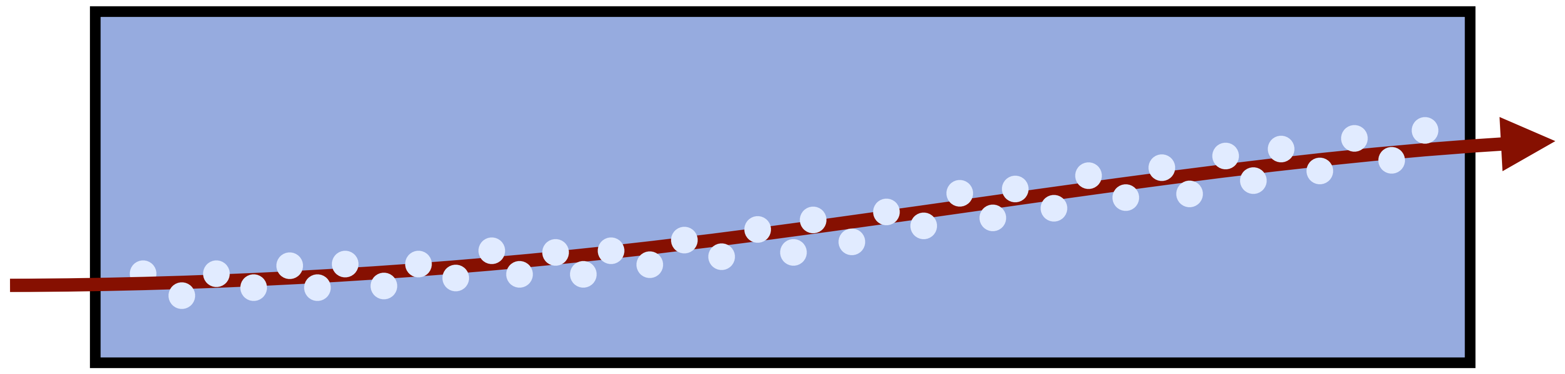
Moreover, if this energy can be collected, one can reconstruct the energy of the daughters, and hence fully reconstruct the interacting neutrino kinematics.

Using the Ionization potential

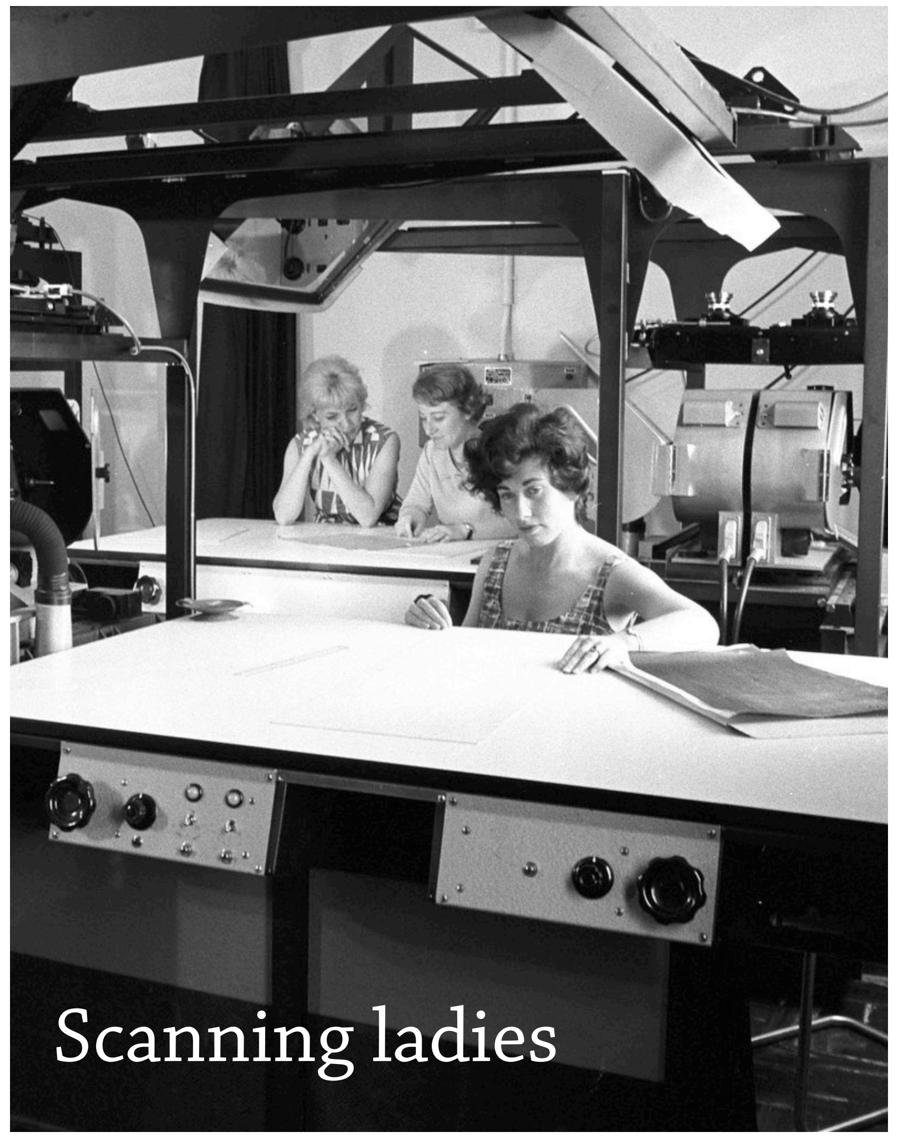

BEBC at CERN

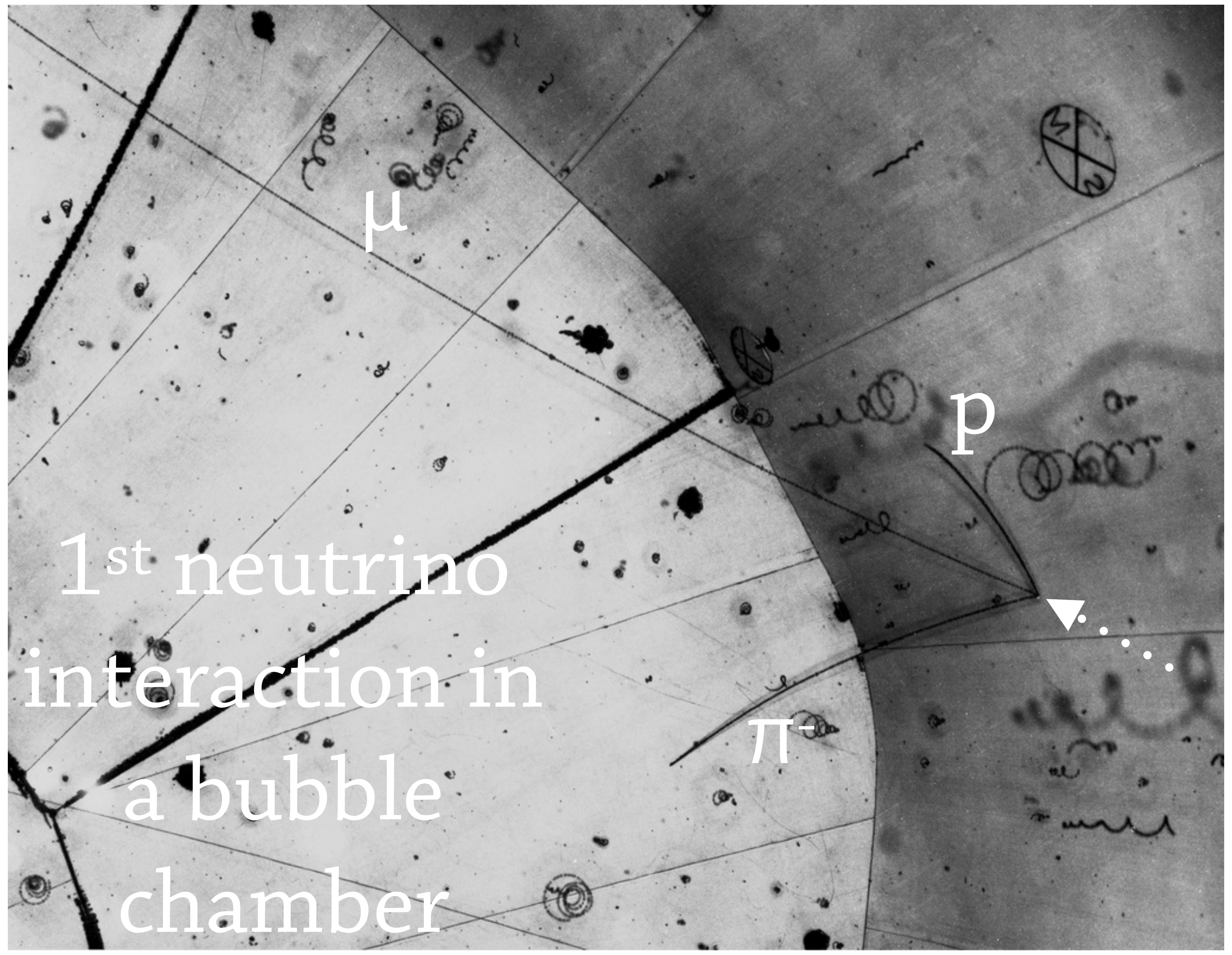
Bubble Chambers

Superheated fluid turns locally to gas (bubbles) when energy is deposited by a charged tracks :

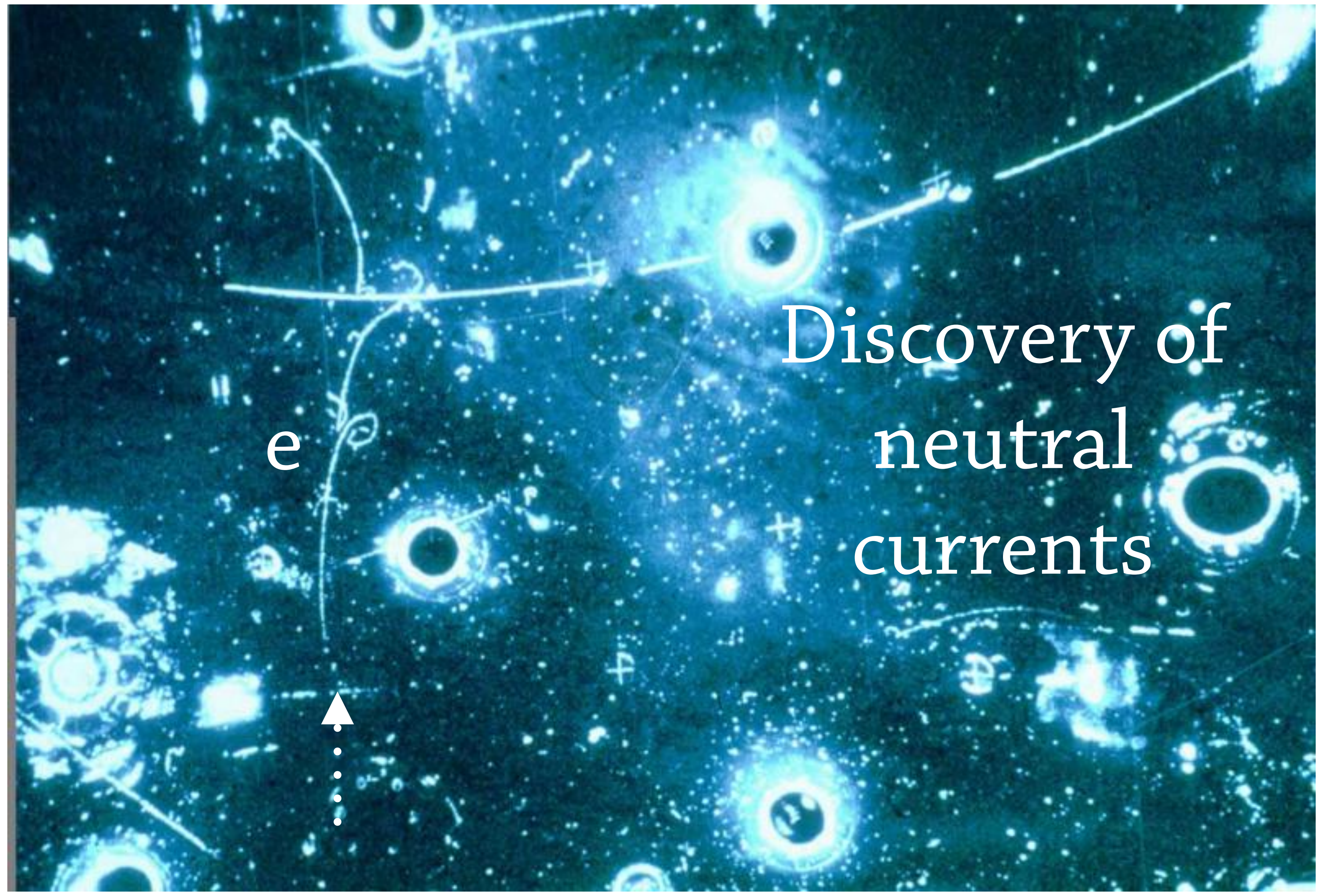


First bubble chambers where equipped with cameras, the pictures were scanned manually by the scanning ladies.


Scanning ladies



1st neutrino interaction in a bubble chamber



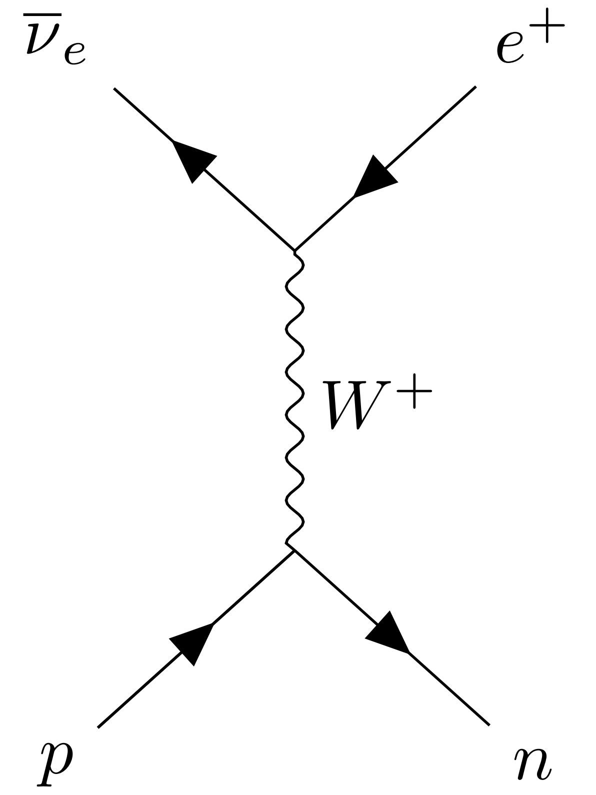
Discovery of neutral currents

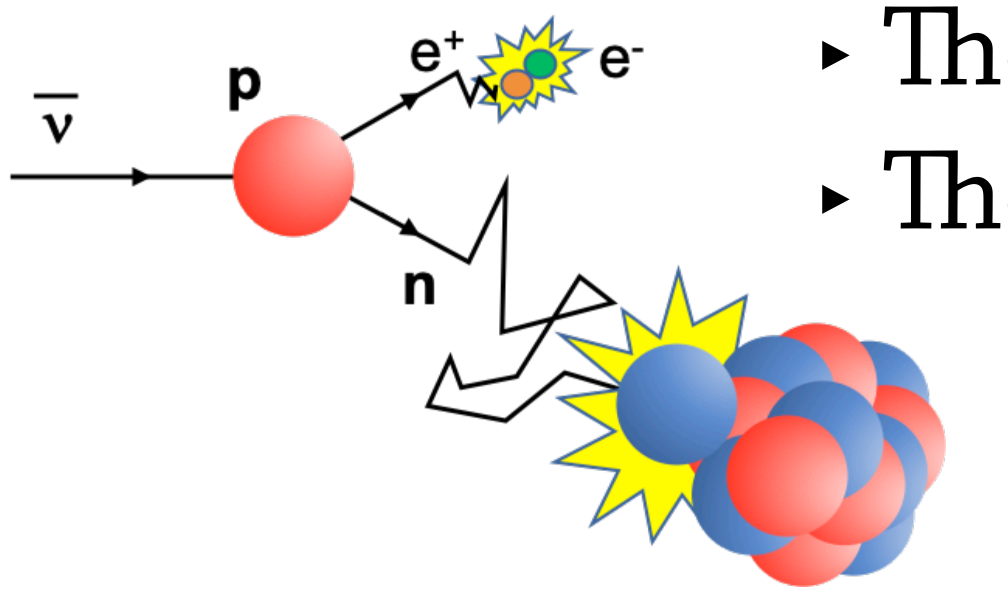
Using the Ionization potential

Liquid Scintillators

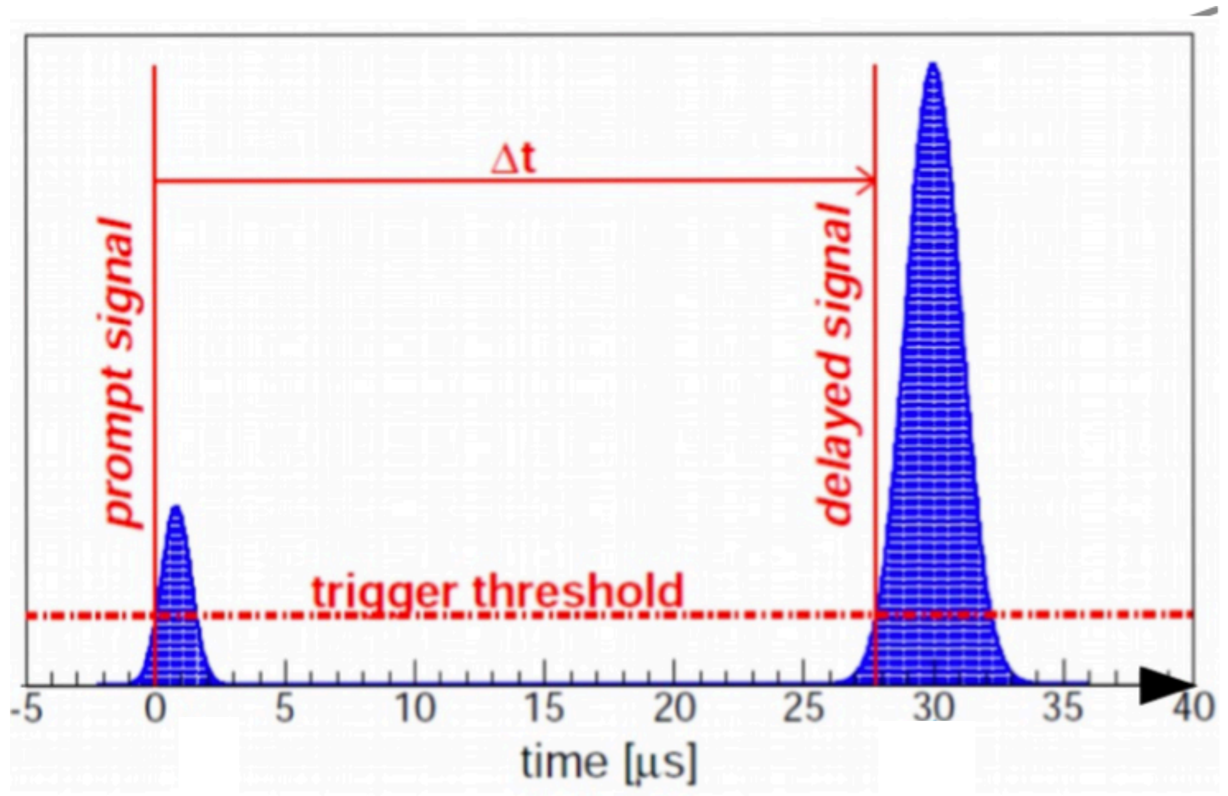
Organic liquid that scintillates when energy is deposited.

In neutrino physics, often used to tag inverse β-decay interactions

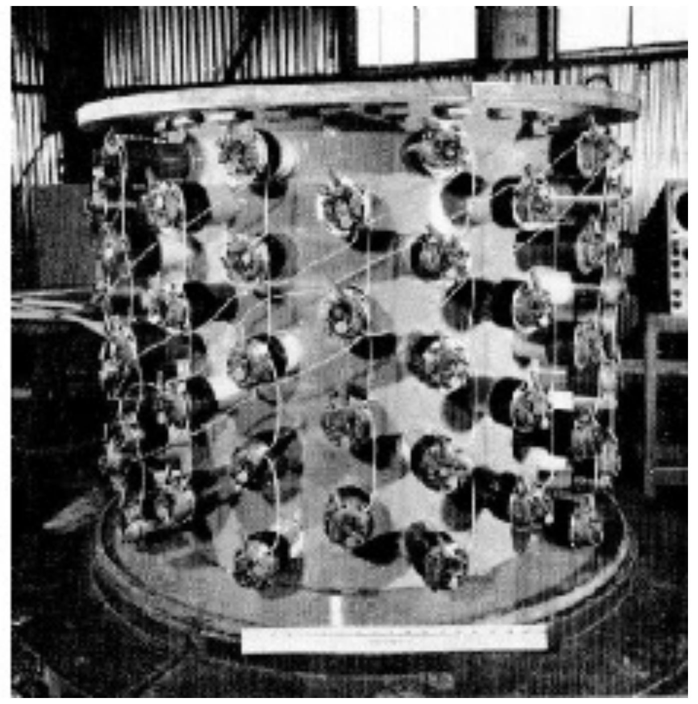


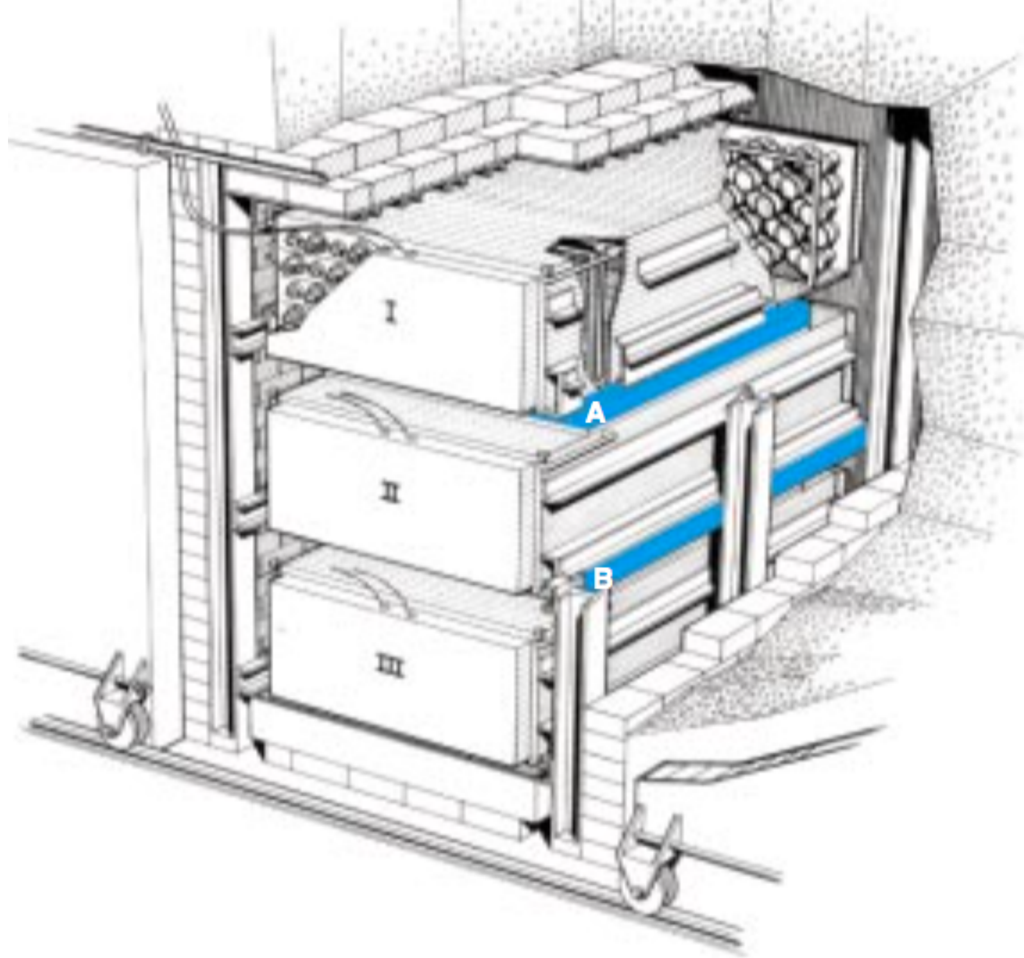


- The positron is quickly captured by an electron
- The neutron is captured later by a catcher-atom

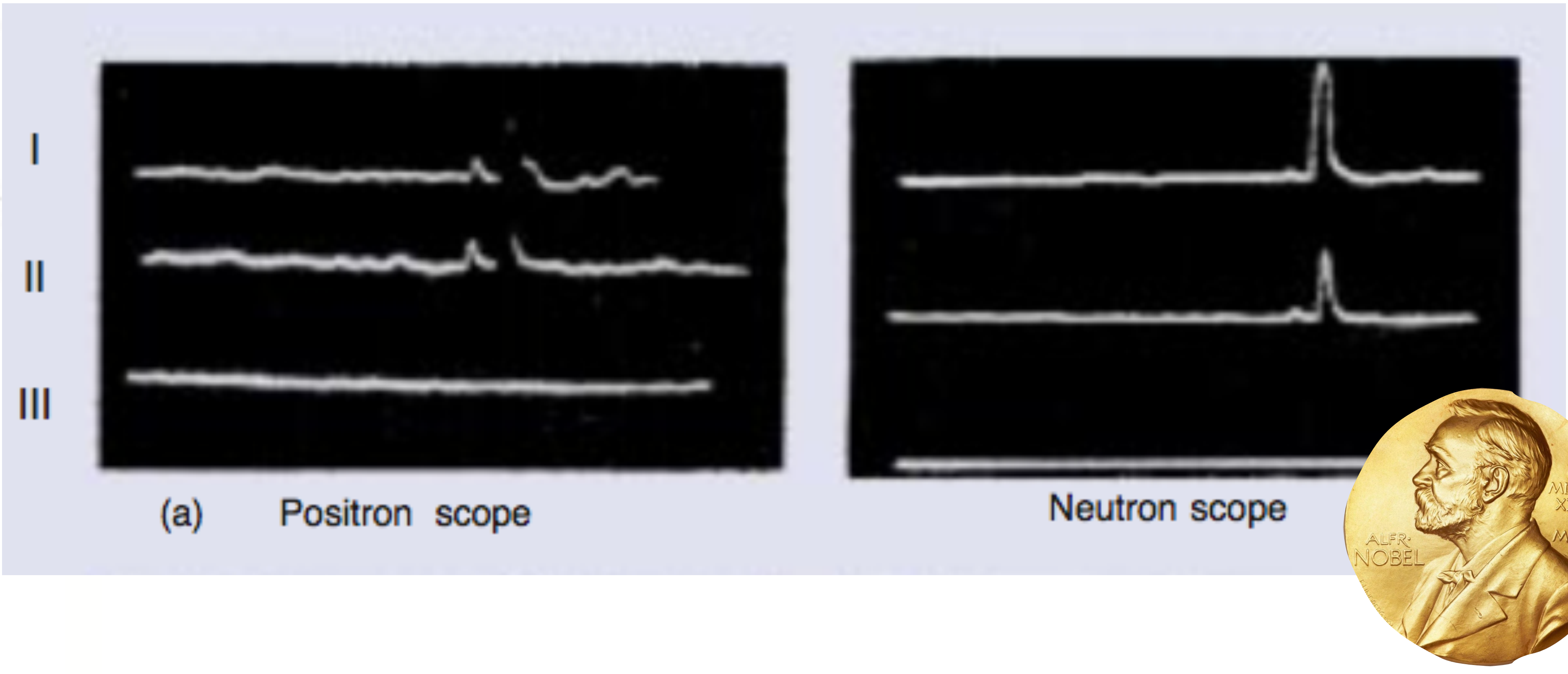


Savannah river experiment by Reines & Cowan in 1956





ν̄e discovery !

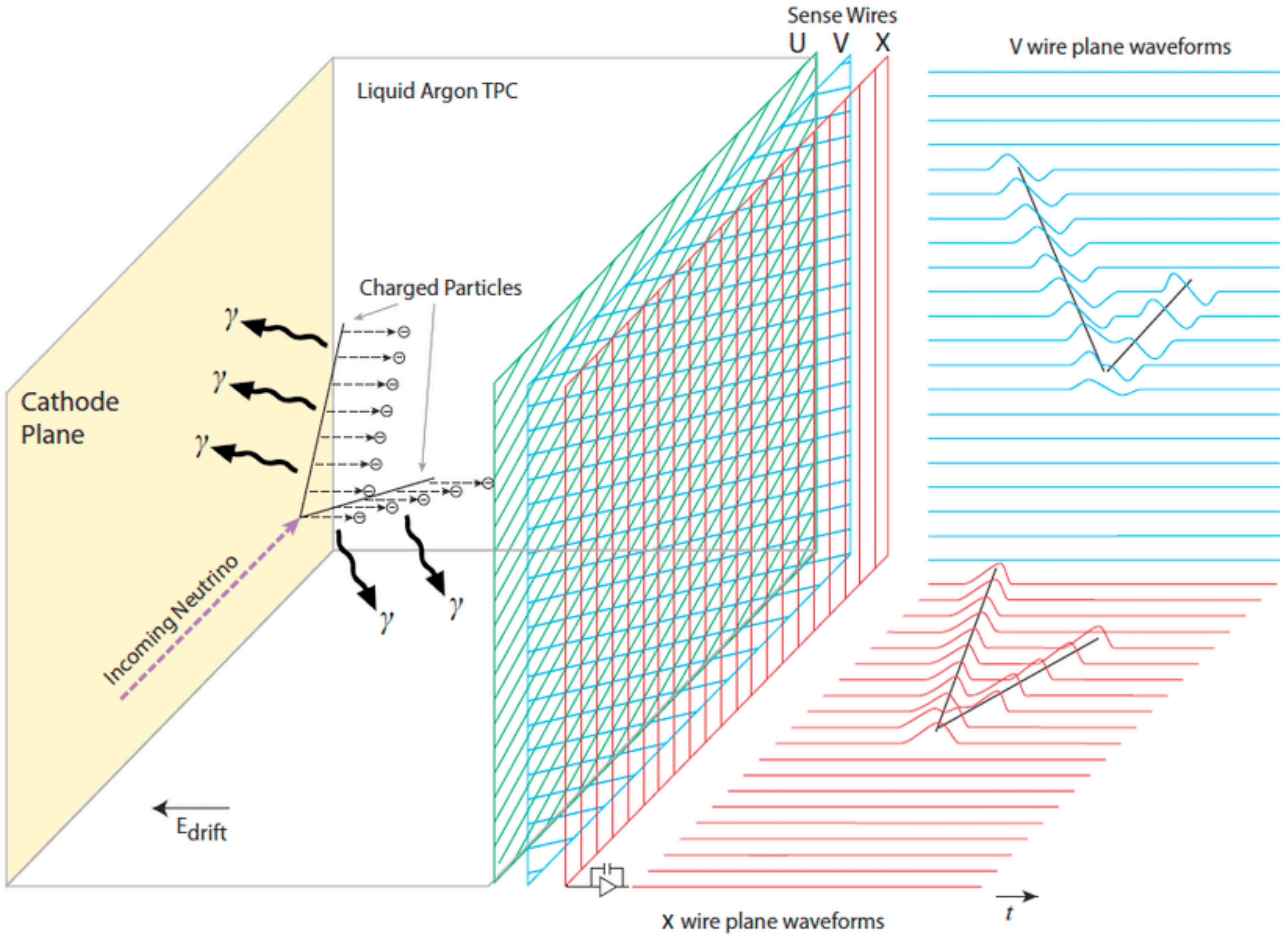


Using the Ionization potential

Time Projection Chamber [TPC]

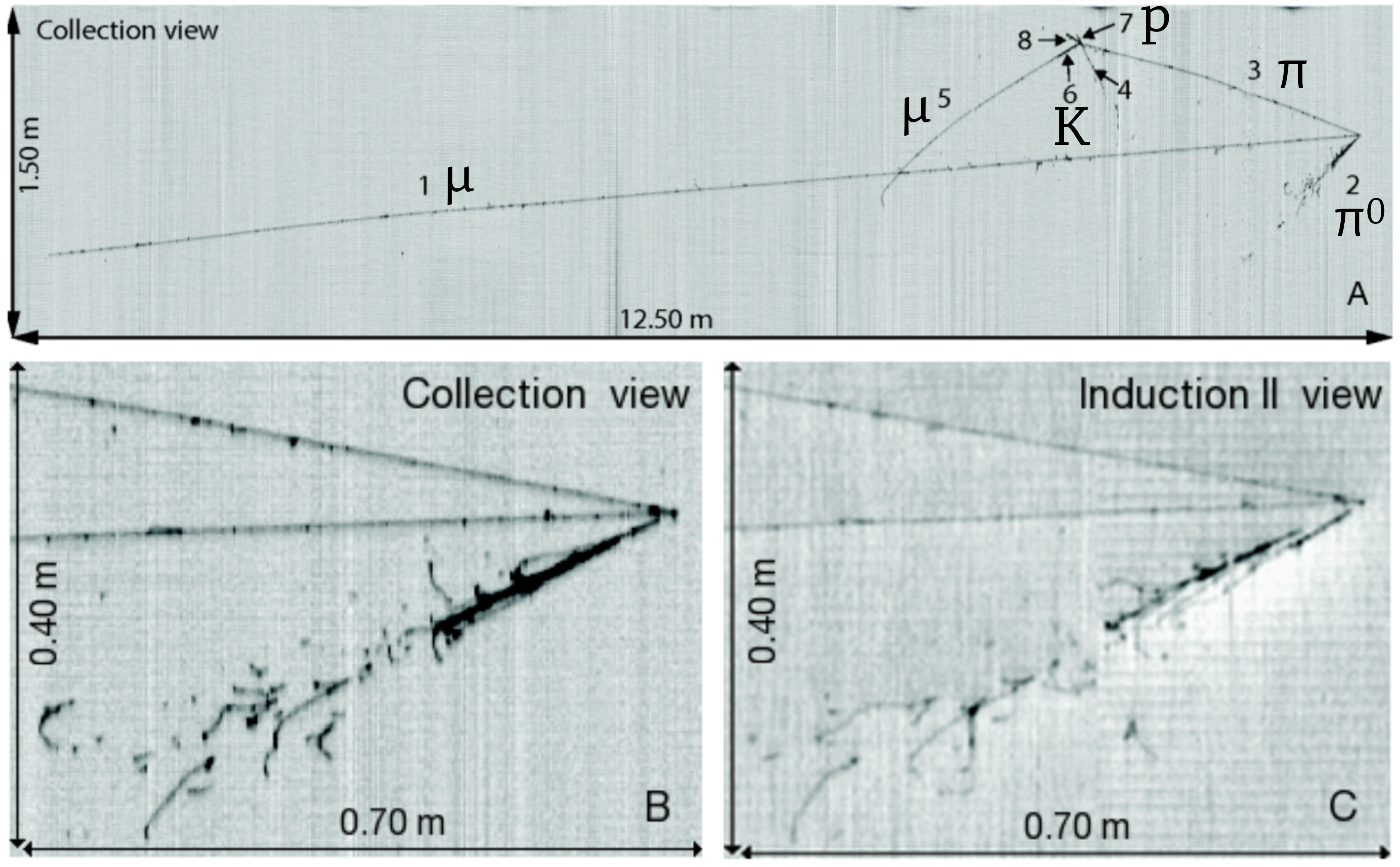
Uses a chamber filled with gas or liquid with an electric field applied across.

Free electrons from ionization are drifting towards the anode plane where they are collected : that gives a 2D image. The e$^-$ arrival time provides the 3rd coordinates. The amount of e$^-$ collected/cm is a handle to retrieve the particle identity/energy.



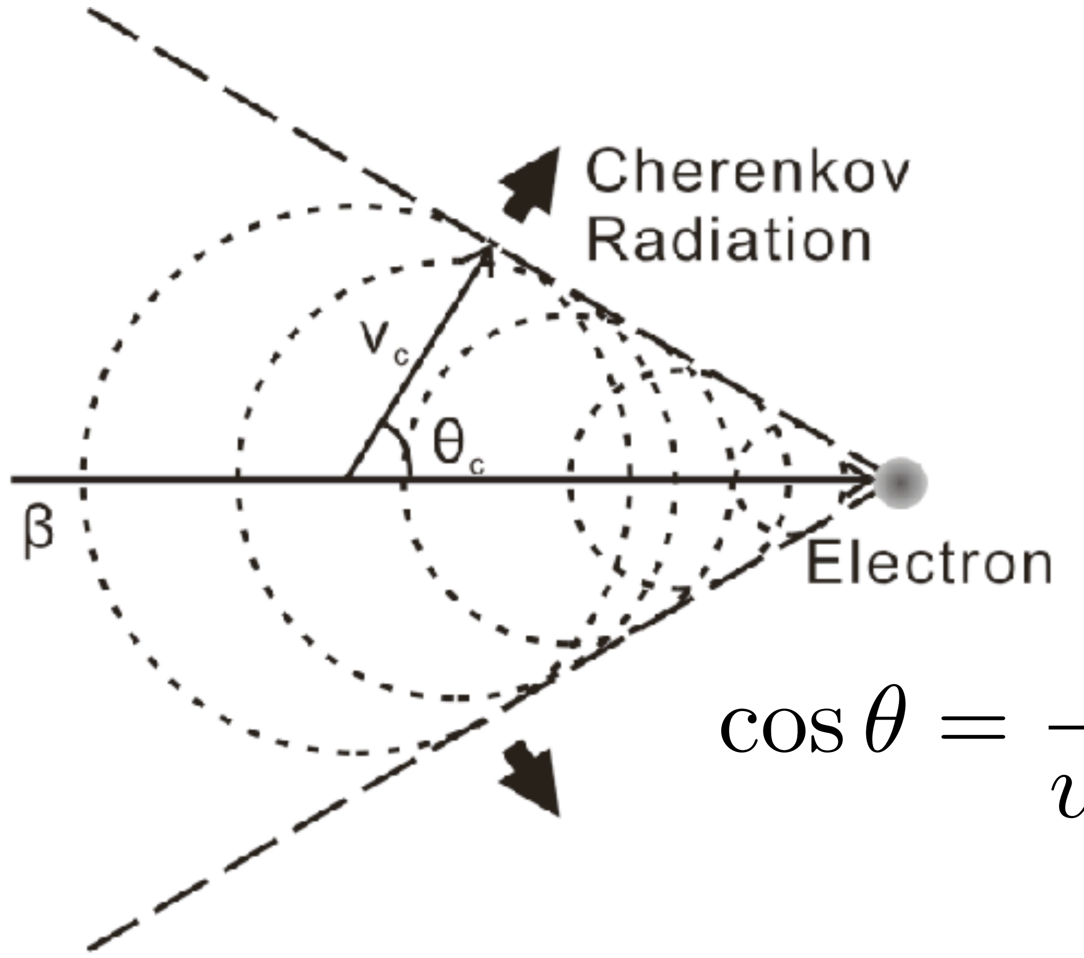
ICARUS experiment in Italy (now in USA)

- ν_μ interaction -



Using the Cherenkov effect

Principle : When a charged particle travels at a speed *v* higher than the speed of light in a medium *c/n* it radiates a cone of light :



$$\cos\theta = \frac{c}{vn}$$

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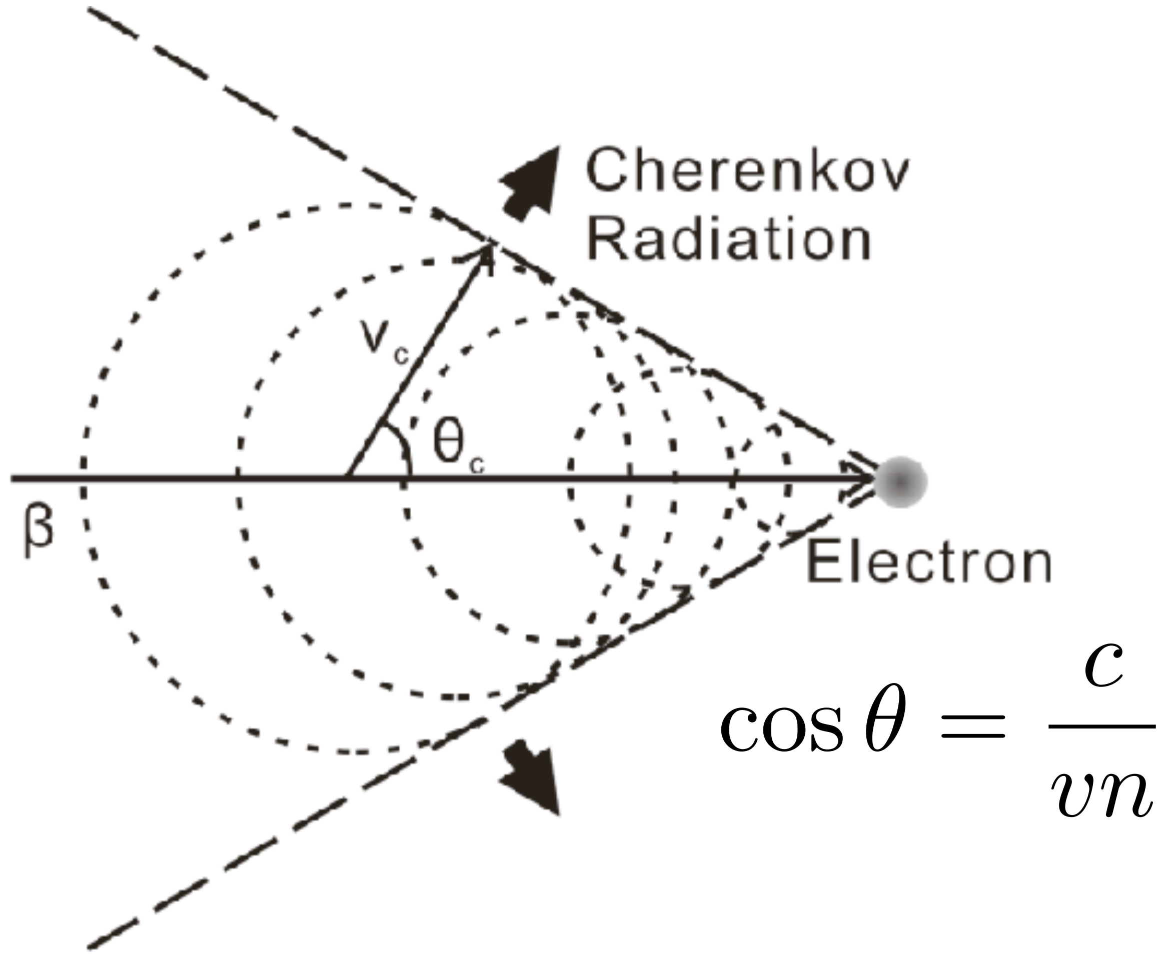

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Cherenkov detectors are widely used in neutrino physics :

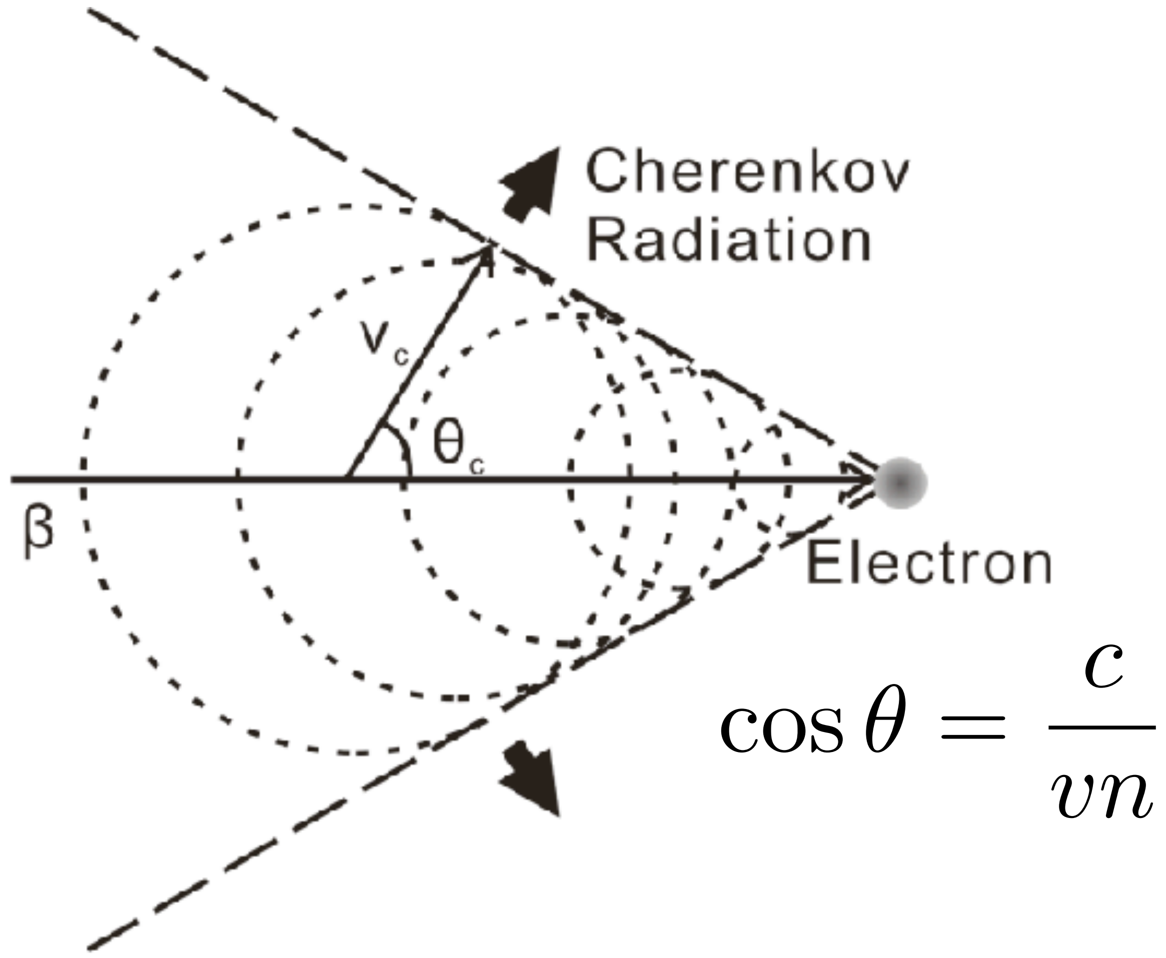
-> Can use cheap/free medium (ultra pure water, ice, sea)

-> Use photomultipliers to detect the light, very well known device

-> Can have large volume : bigger volume = more chances to catch a neutrino

-> Ring shape allows particle identification ; ring characteristics (diameter, nb of photons) is linked to the particle energy => Excellent e/μ separation

Using the Cherenkov effect



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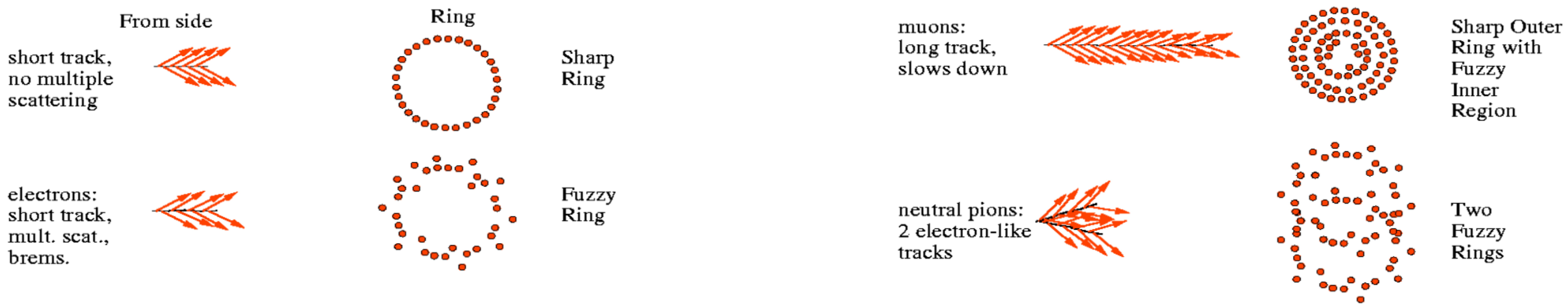
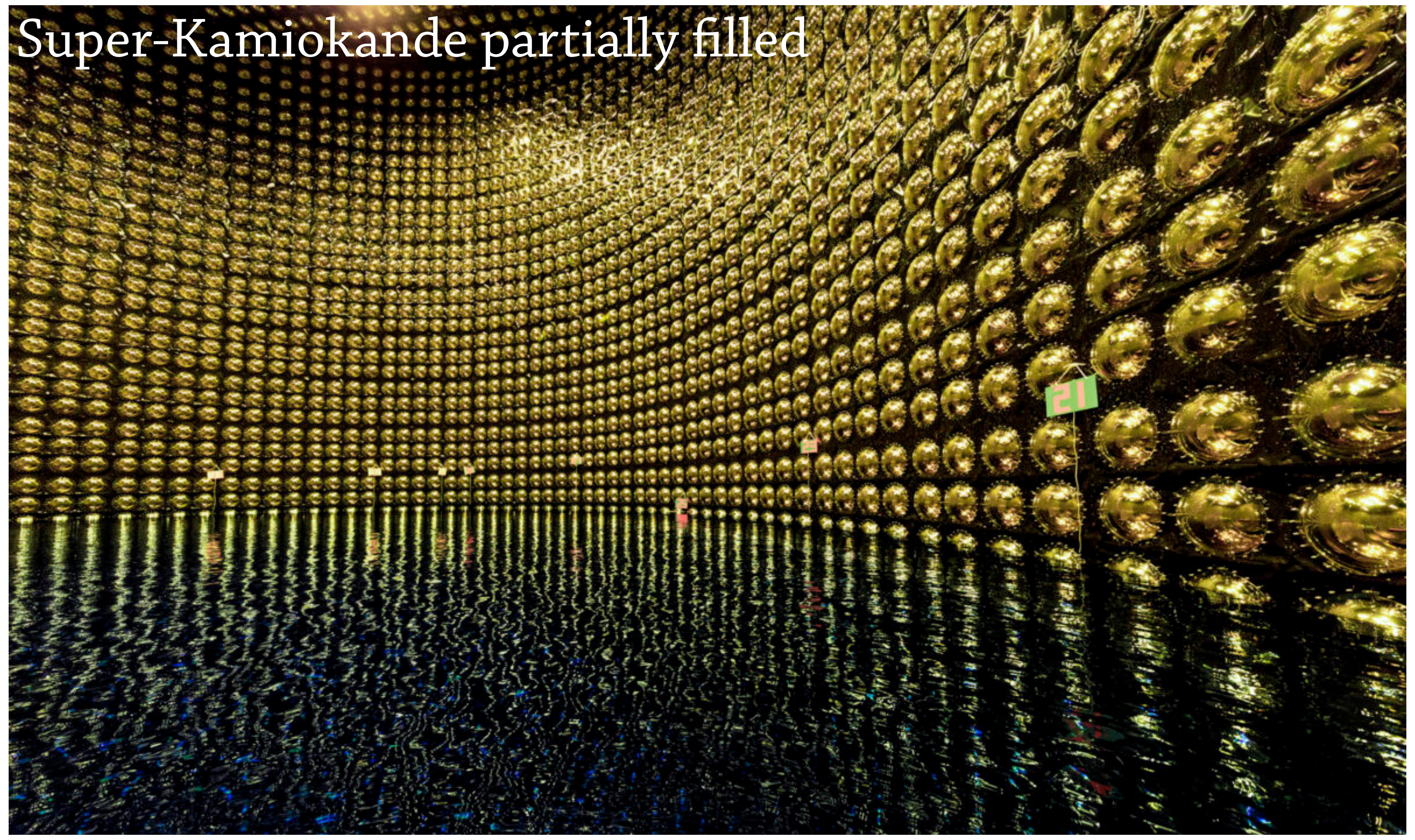
Particle identification using ring shape :

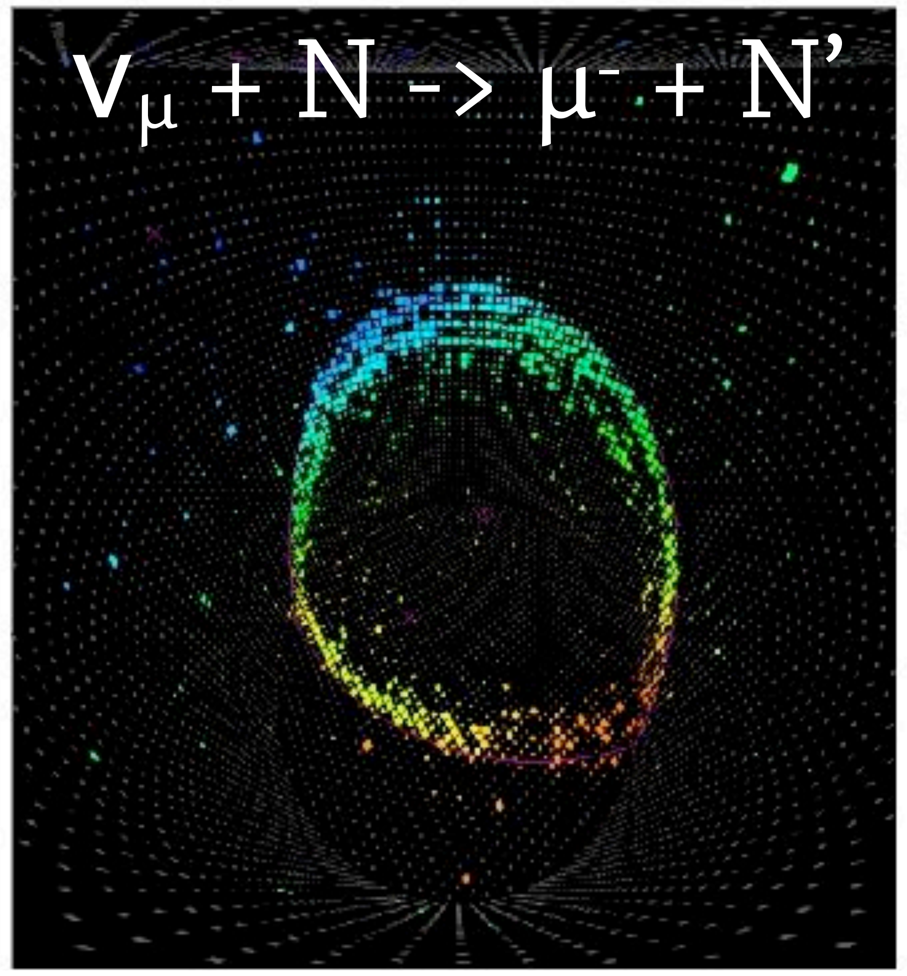


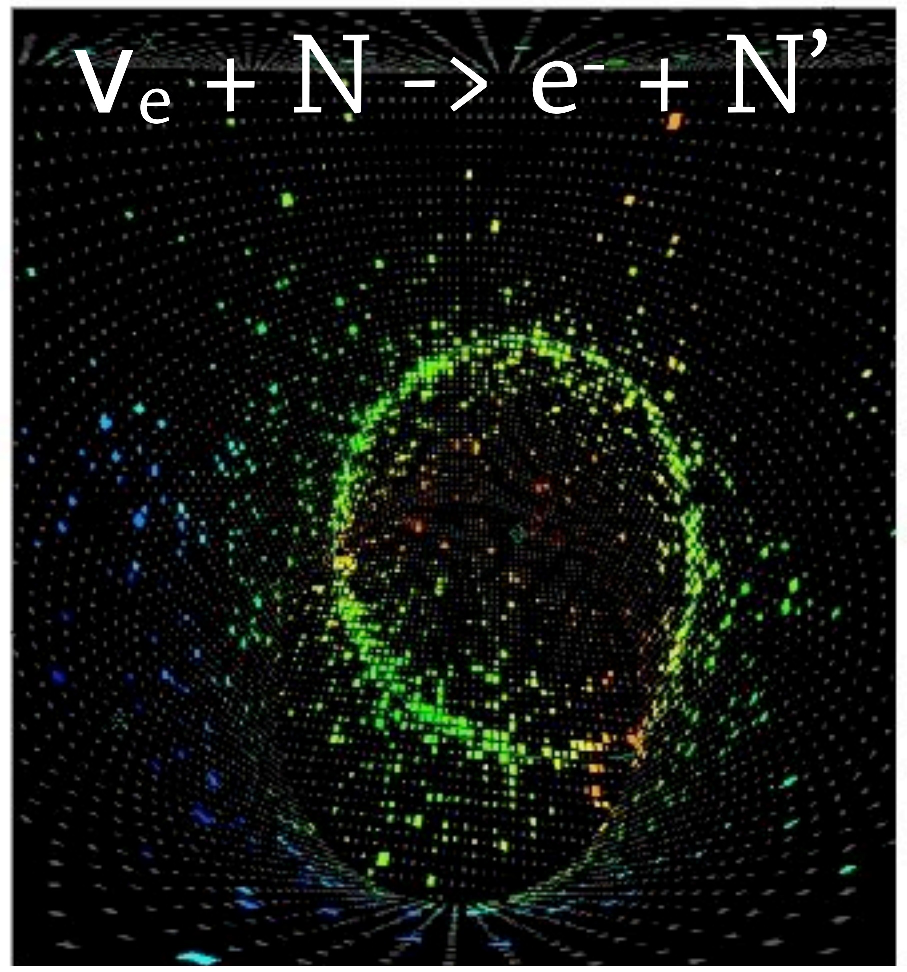
Using the Cherenkov effect

Super-Kamiokande in Japan

Tank of 50 kt of ultra pure water underneath a mountain, equipped with ~11k PMTs







Color = nb of photons collected

Using the Cherenkov effect

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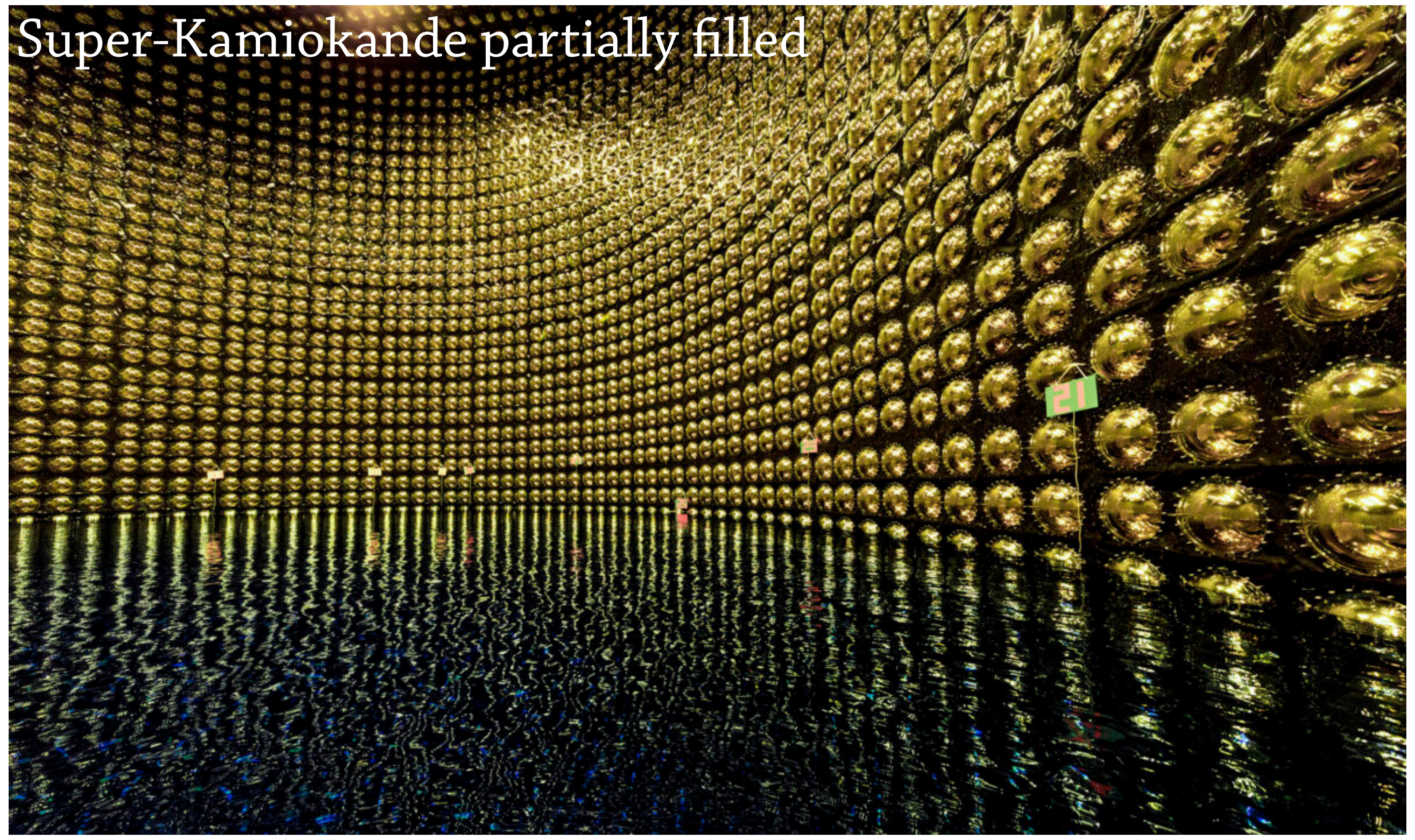
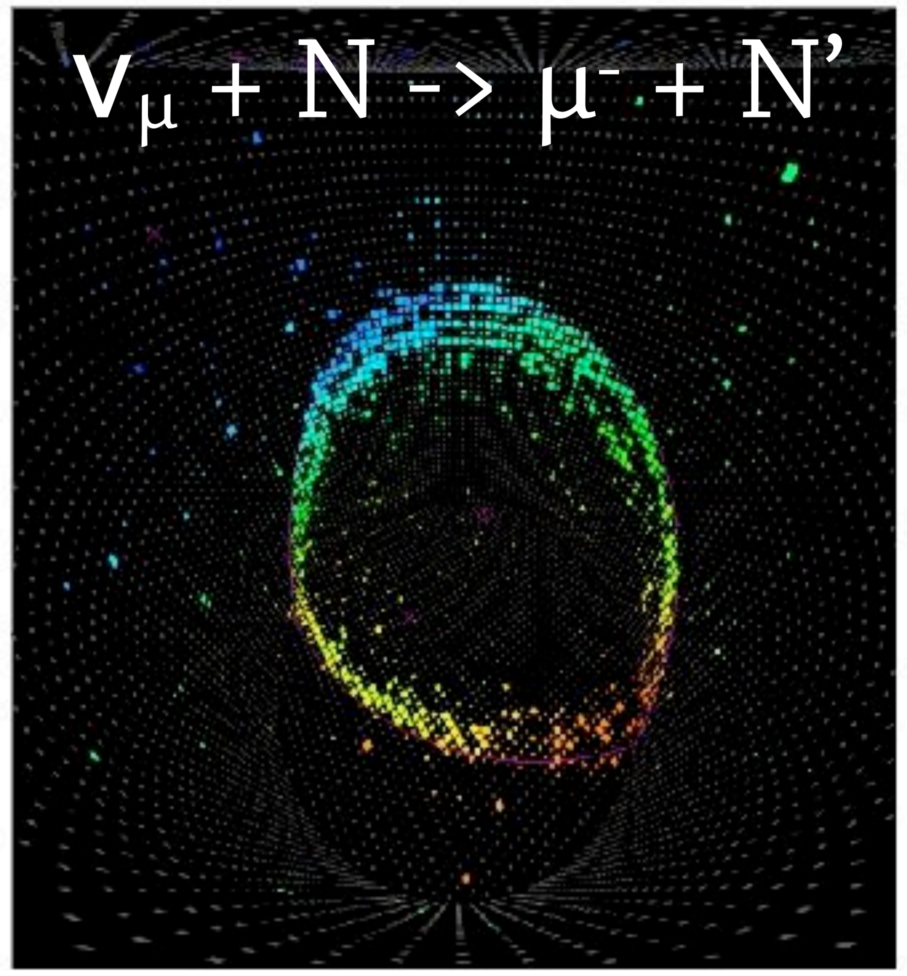
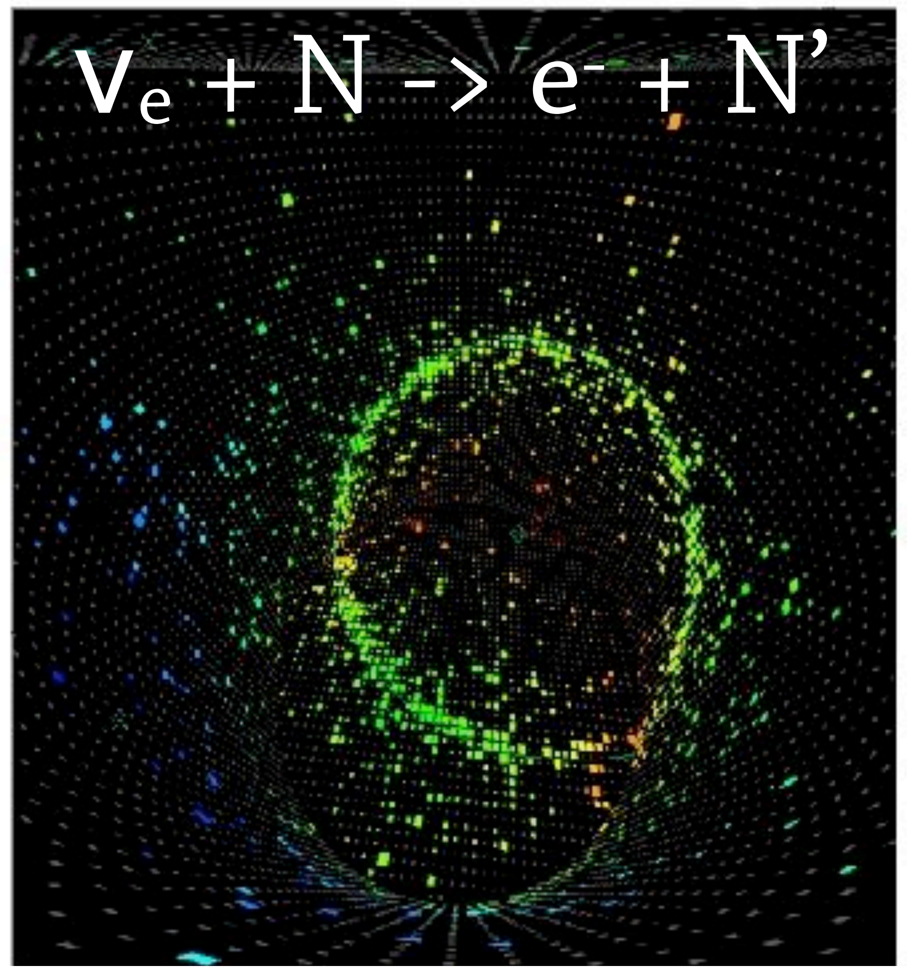

Super-Kamiokande partially filled





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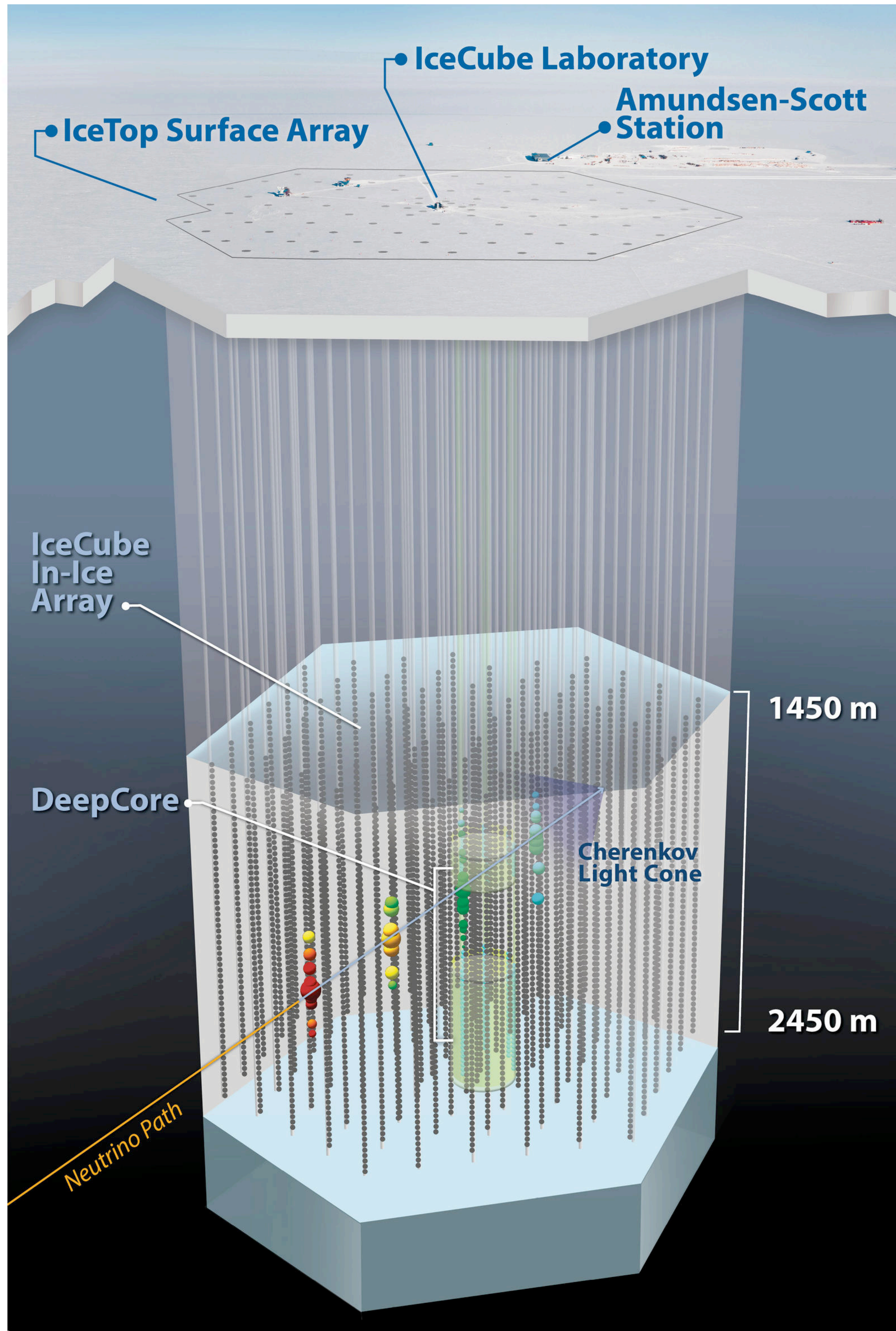
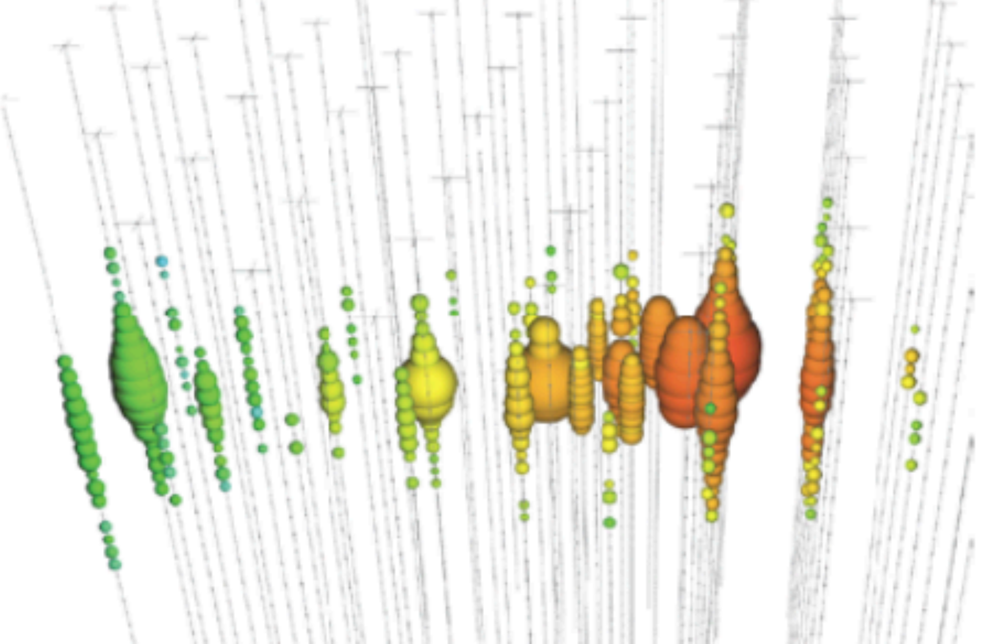
ICECUBE in Antartica

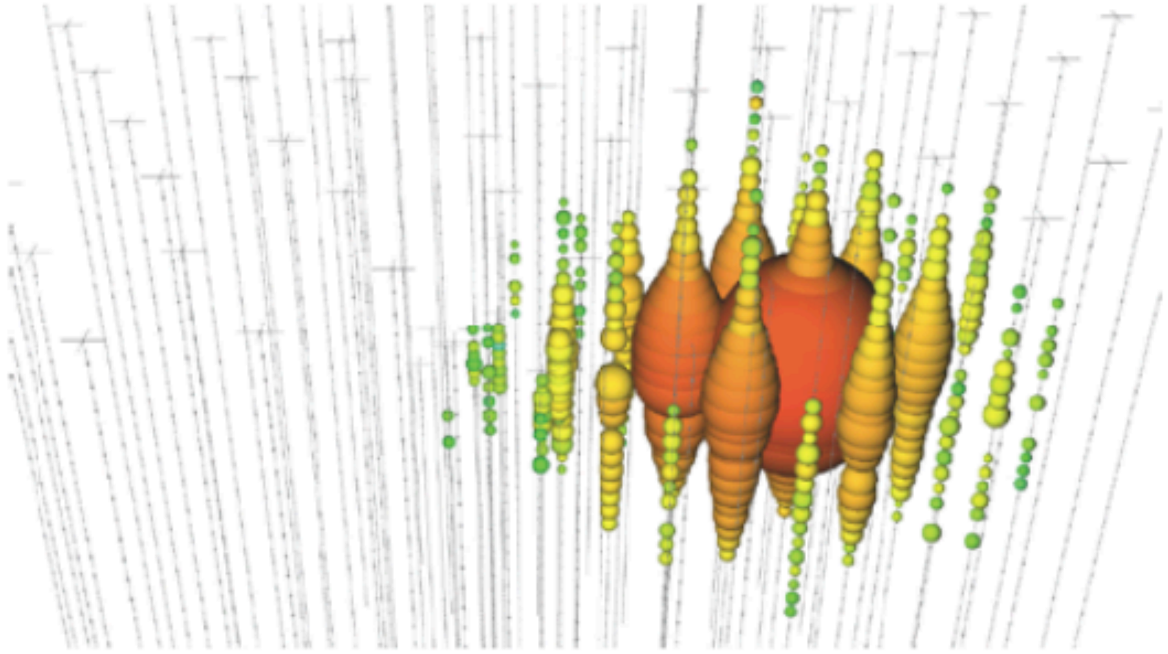
Giant detector in ice, equipped with 5k PTMs along 86 strings up to 2.4 km below the surface



$\nu_\mu + N -> \mu^- + N'$



$\nu_e + N -> e^- + N'$



Color = time
size = nb of photons collected

DISCOVERY OF NEUTRINO OSCILLATION

Solar neutrino flux

Most of 20th century research focused on nuclear reactions: radioactivity, fission and fusion.

Among all the consequences, it helped to understand stellar nucleosynthesis that powers stars.

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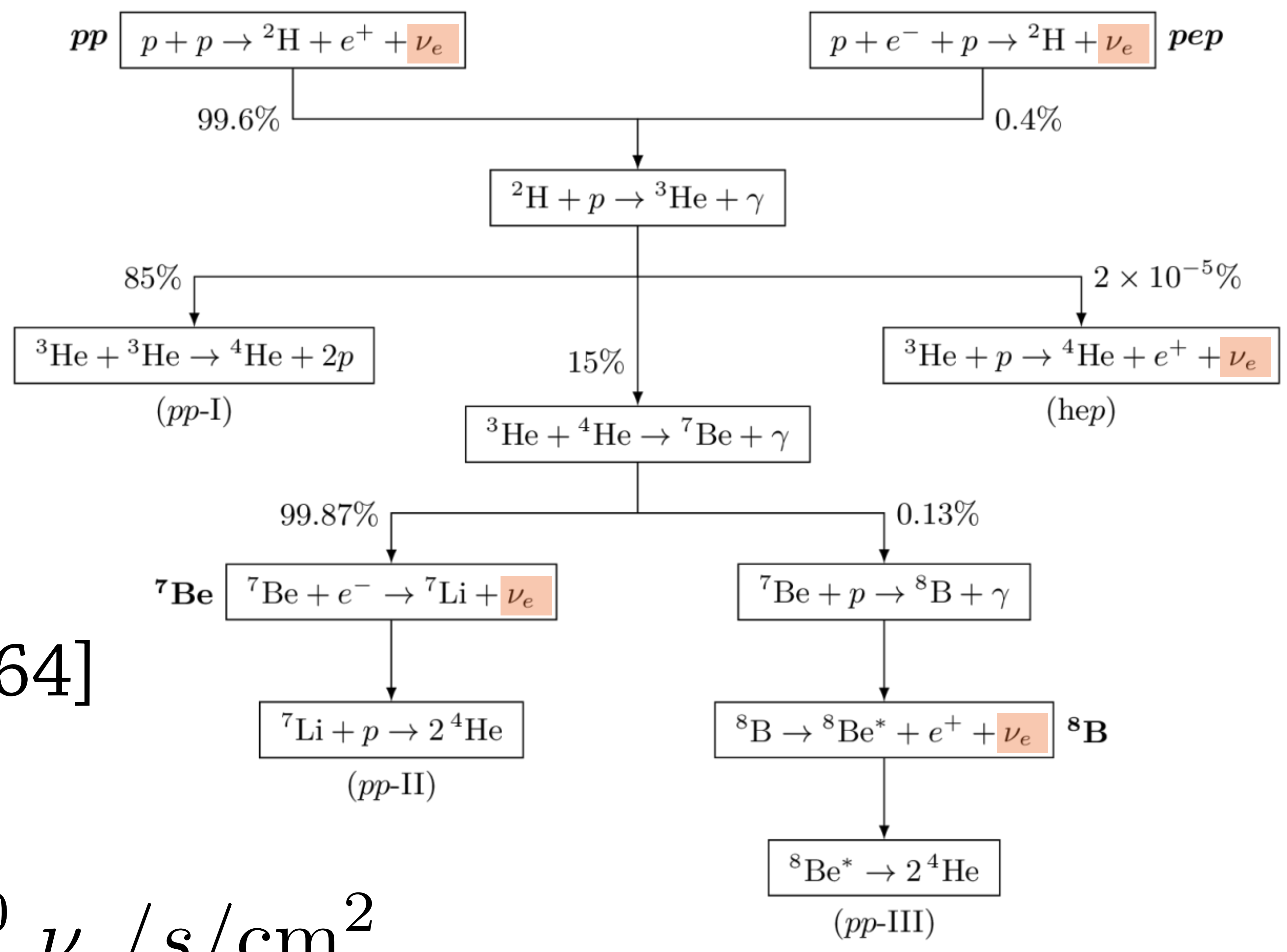
Proton-proton fusion chain in sun-like stars



pp $p + p \to {}^2\mathrm{H} + e^+ + \nu_e$

$p + e^- + p \to {}^2\mathrm{H} + \nu_e$ **pep**

99.6%

0.4%

${}^2\mathrm{H} + p \to {}^3\mathrm{He} + \gamma$

85%

15%

$2 \times 10^{-5}\%$

${}^3\mathrm{He} + {}^3\mathrm{He} \to {}^4\mathrm{He} + 2p$ (*pp*-I)

${}^3\mathrm{He} + {}^4\mathrm{He} \to {}^7\mathrm{Be} + \gamma$

${}^3\mathrm{He} + p \to {}^4\mathrm{He} + e^+ + \nu_e$ (he*p*)

99.87%

0.13%

⁷Be ${}^7\mathrm{Be} + e^- \to {}^7\mathrm{Li} + \nu_e$

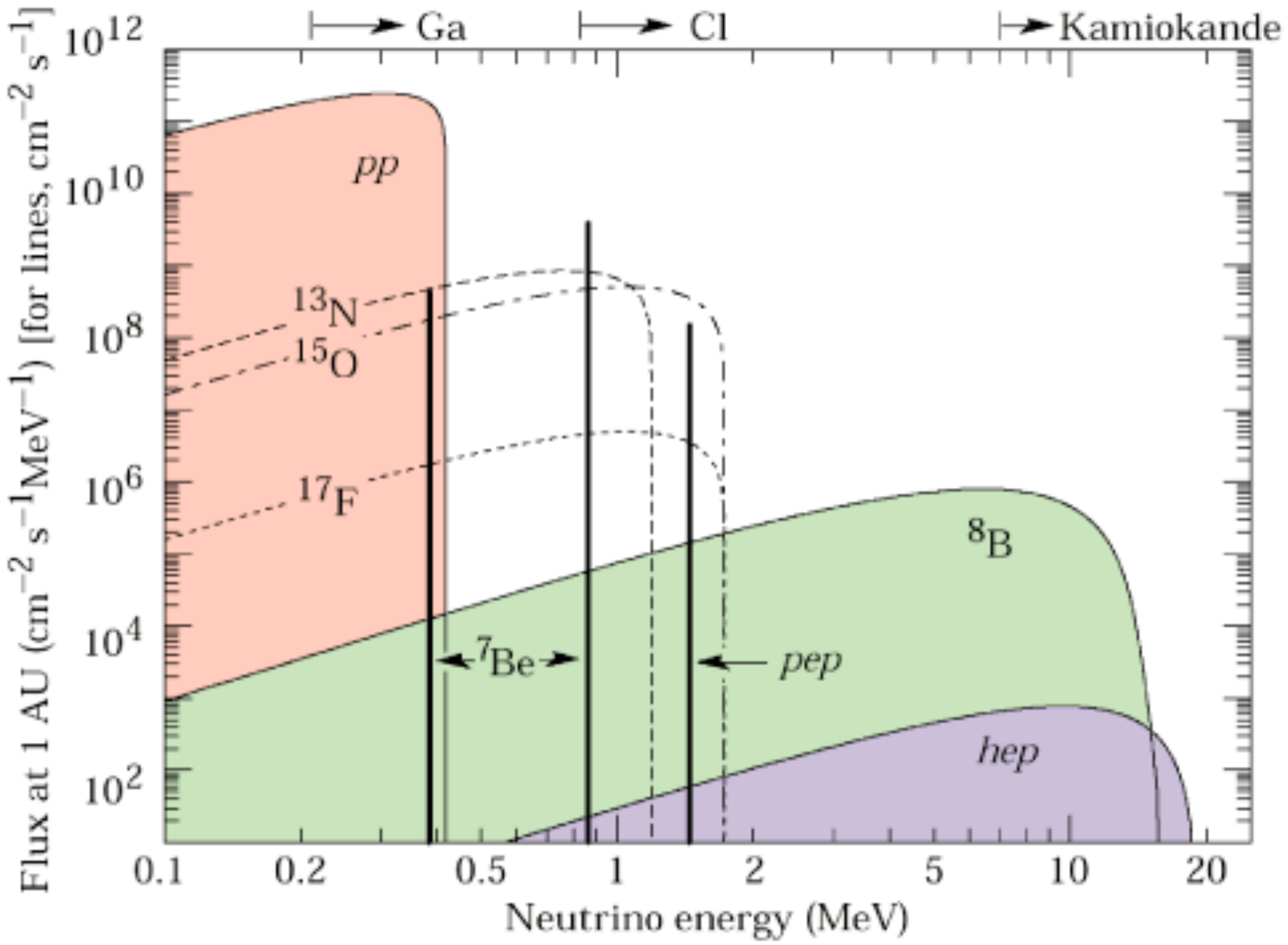
${}^7\mathrm{Be} + p \to {}^8\mathrm{B} + \gamma$

${}^7\mathrm{Li} + p \to 2\,{}^4\mathrm{He}$ (*pp*-II)

${}^8\mathrm{B} \to {}^8\mathrm{Be}^* + e^+ + \nu_e$ **⁸B**

${}^8\mathrm{Be}^* \to 2\,{}^4\mathrm{He}$ (*pp*-III)

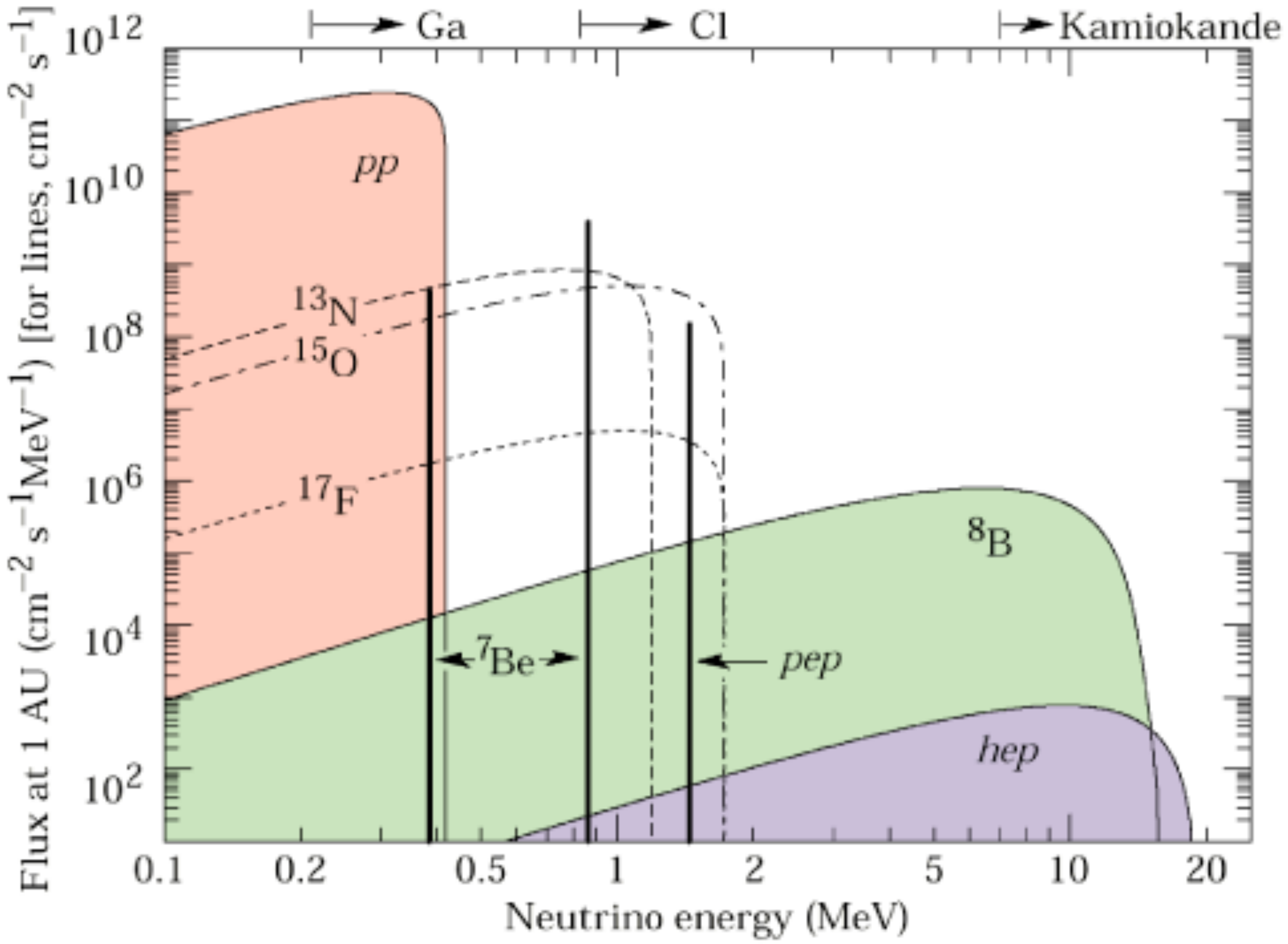
Bahcall made a prediction on the ν_e flux from the sun [1964]



$$\phi_{\nu_e}^{\mathrm{sun}} = 6.4 \times 10^{10}\ \nu_e/s/\mathrm{cm}^2$$

Neutrinos can escape the sun plasma unaffected. Detecting those neutrino would prove the fusion chain happening inside sun's core.

Solar neutrino deficit



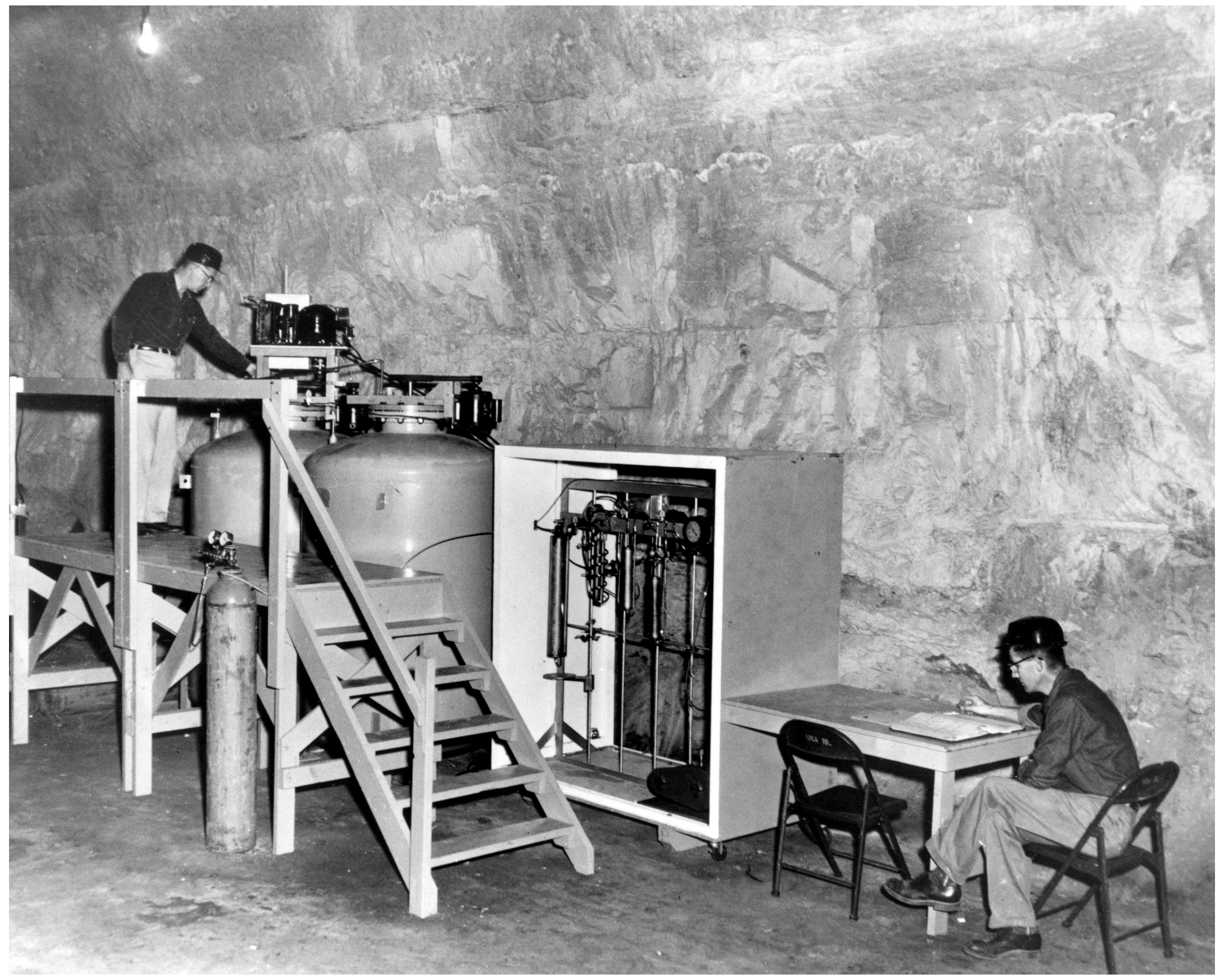
Solar Neutrino Unit = 10^{-36} interaction/s/atom

Solar neutrino deficit

Homestake Experiment designed to detect solar neutrinos

Underground detector observing

Cl to Ar conversion by: $\nu_e + {}^{37}\mathrm{Cl} \rightarrow {}^{37}\mathrm{Ar}^{+} + e^{-}$

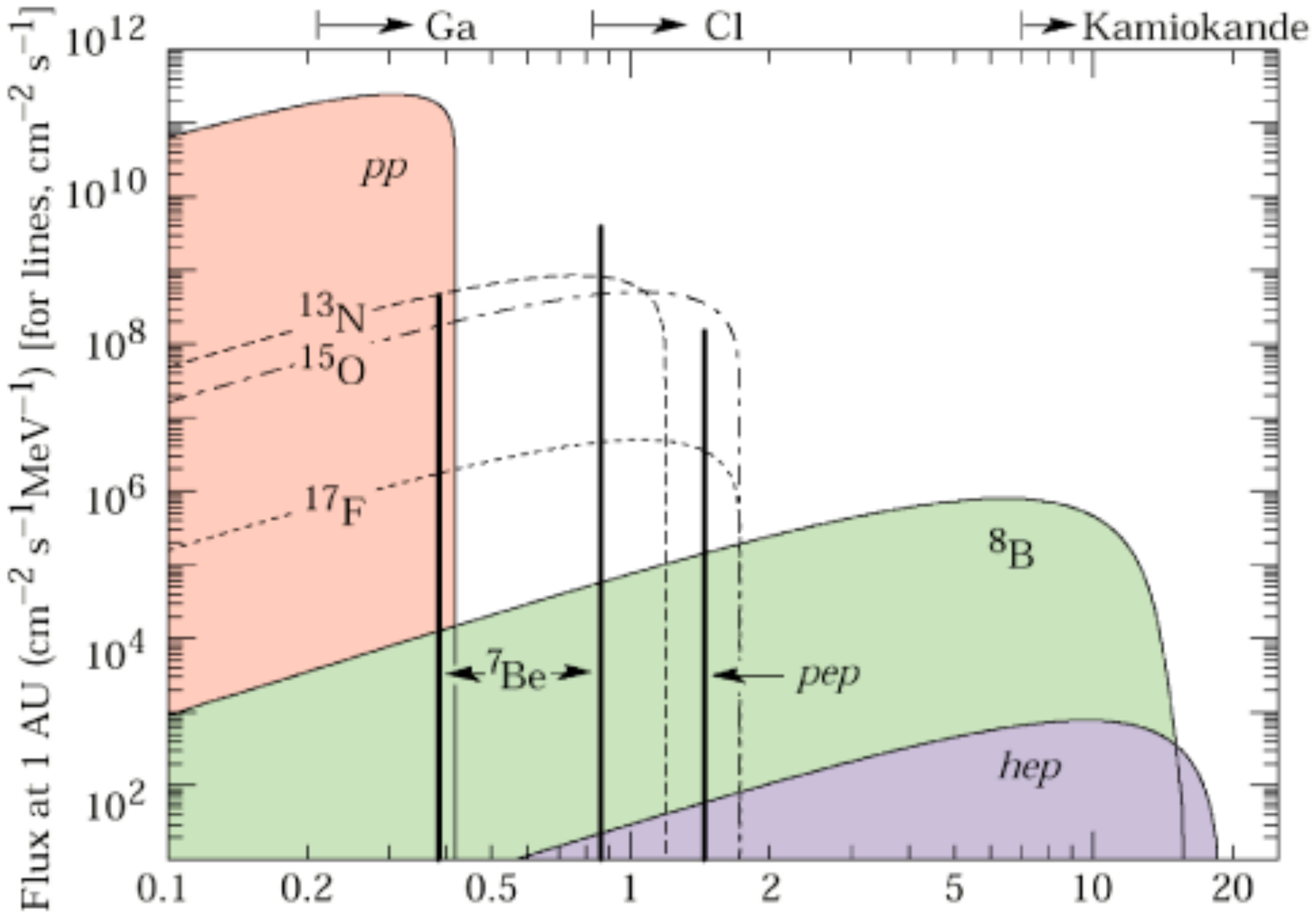


Number of Ar atom in the chamber was counted every few weeks with filters

Expected : 8.2 ± 1.8 SNU

Observed : 2.56 ± 0.23 SNU

60% ν_e missing



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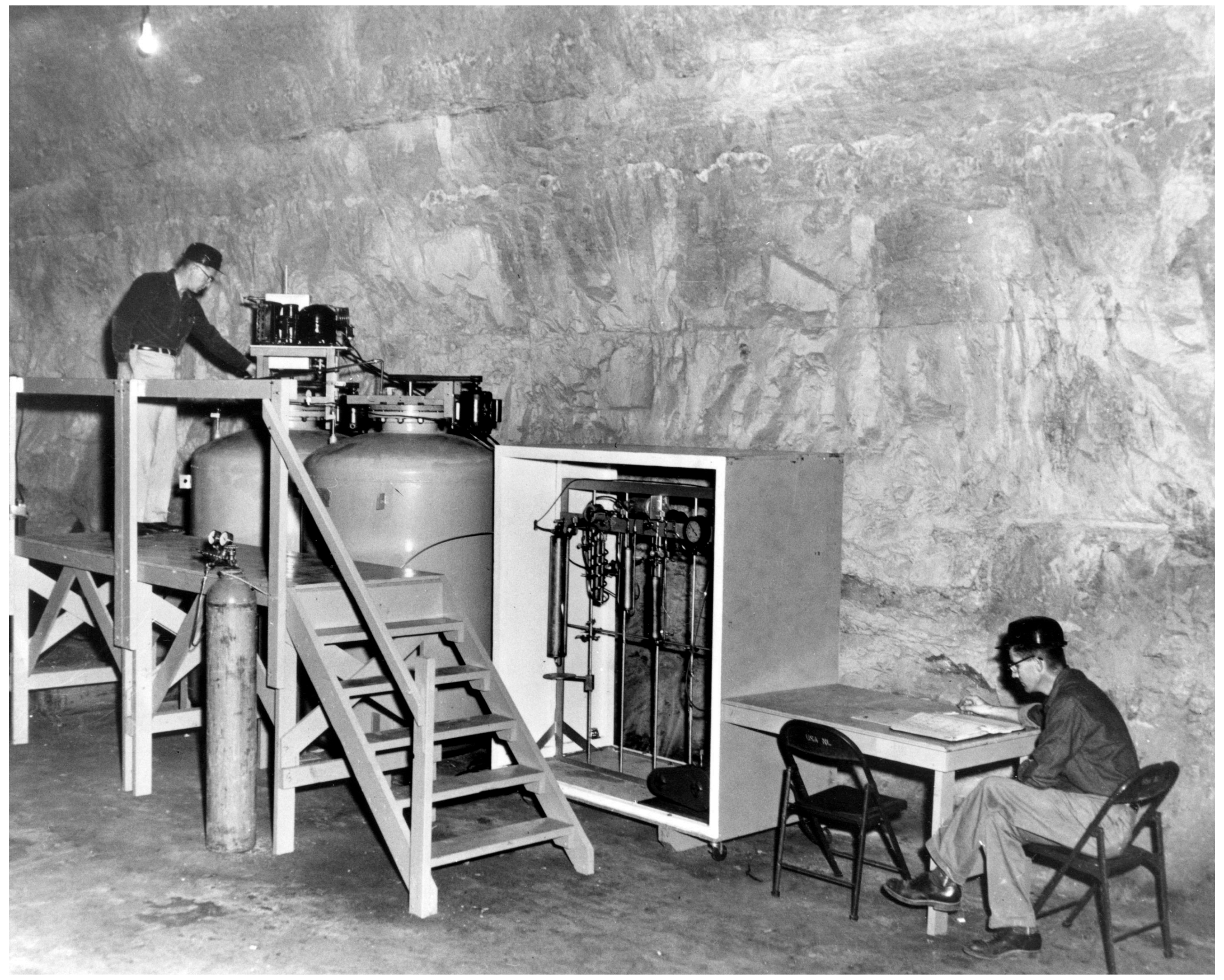

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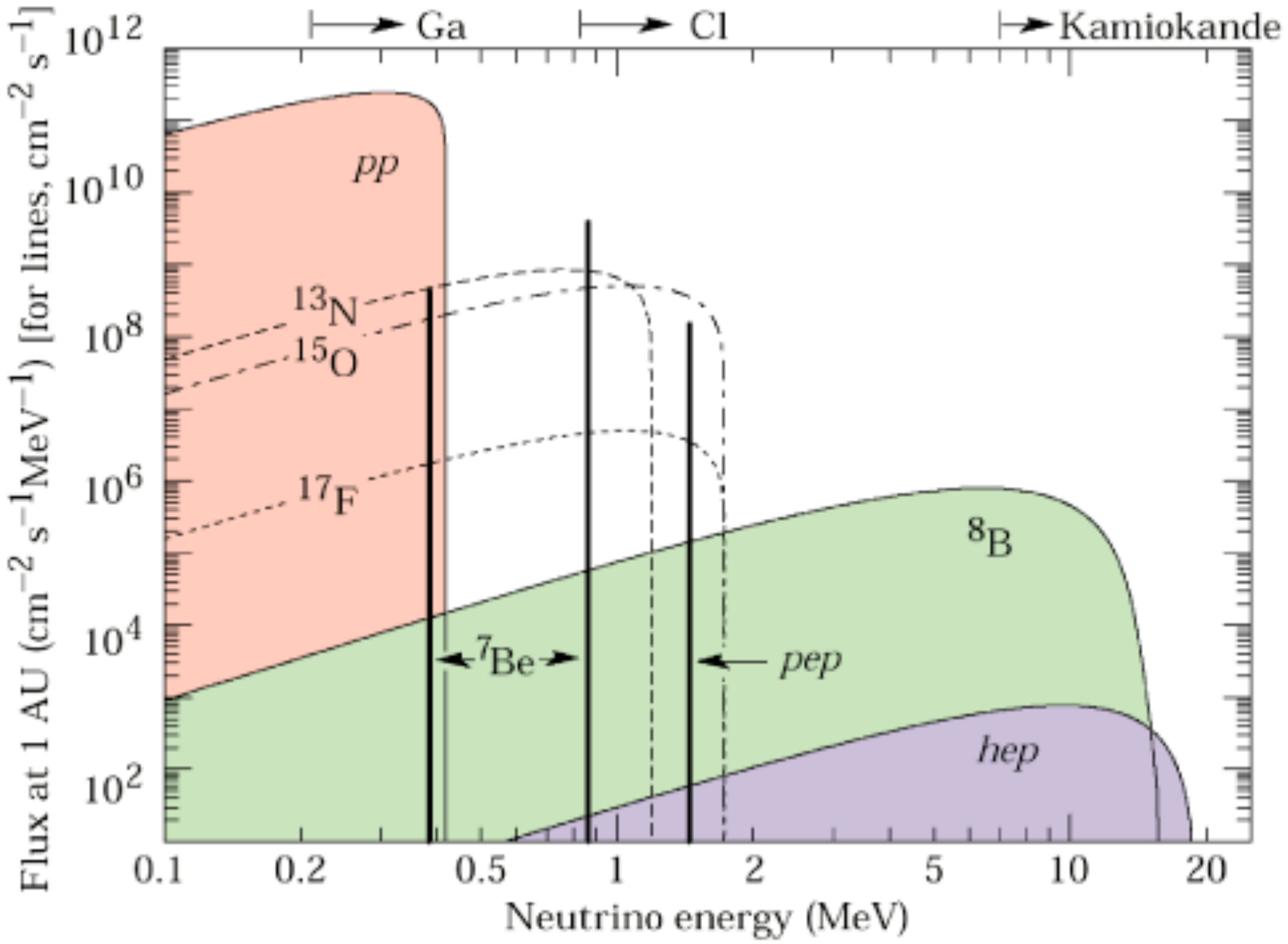

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The ***GALLEX*** and ***SAGE*** experiments used
Ga to Ge conversion : $\nu_e + {}^{71}\mathrm{Ga} \to {}^{71}\mathrm{Ge} + e^{-}$

Expected : 127 ± 12 SNU
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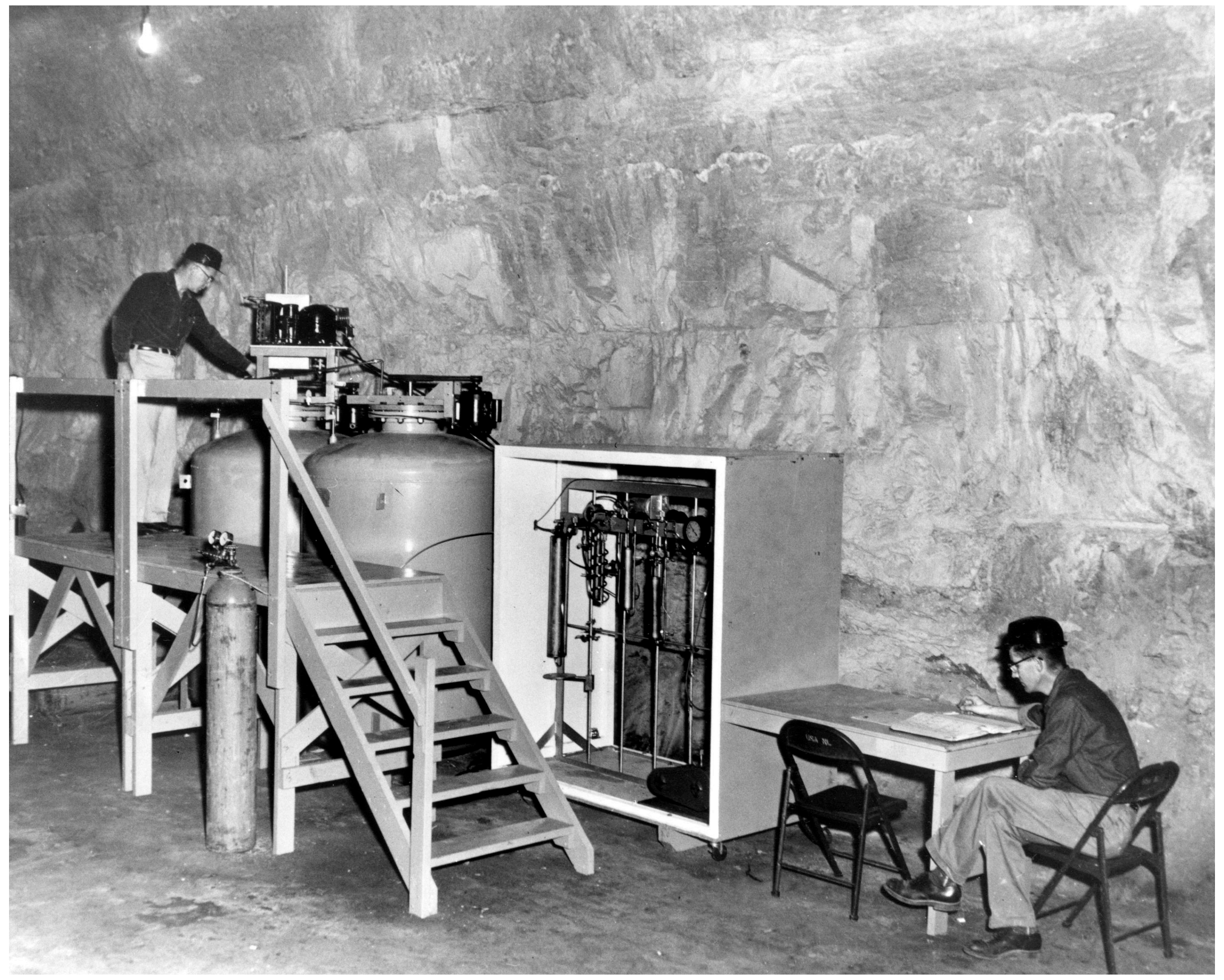

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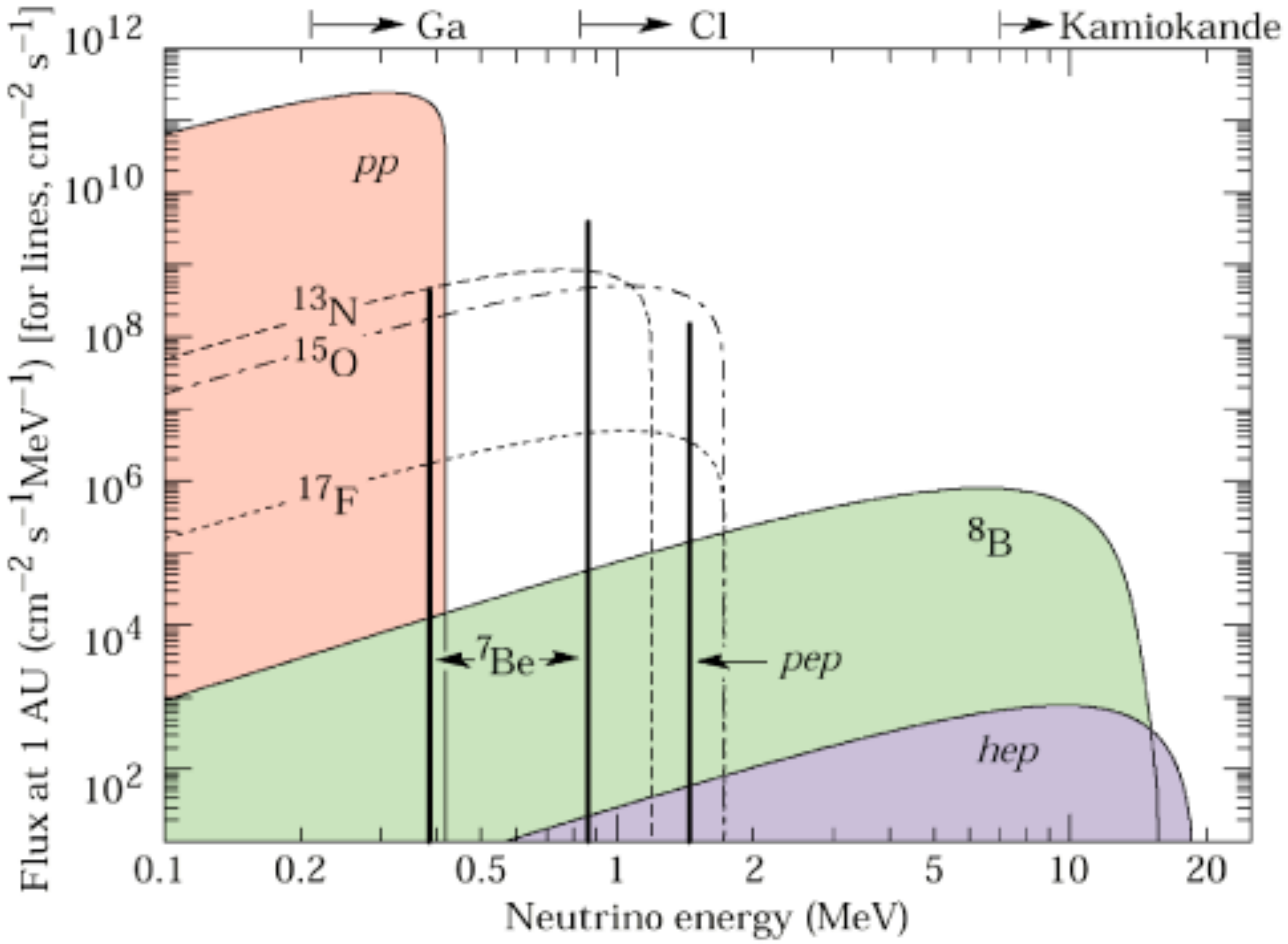

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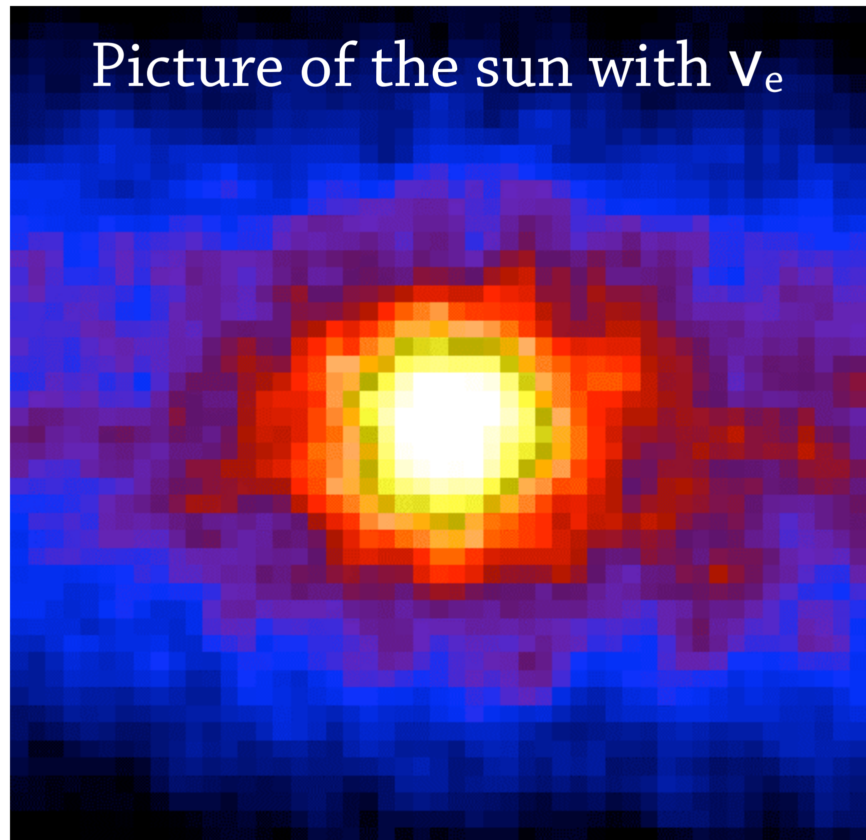
The ***Kamiokande*** experiment

observed solar neutrino through elastic scattering :

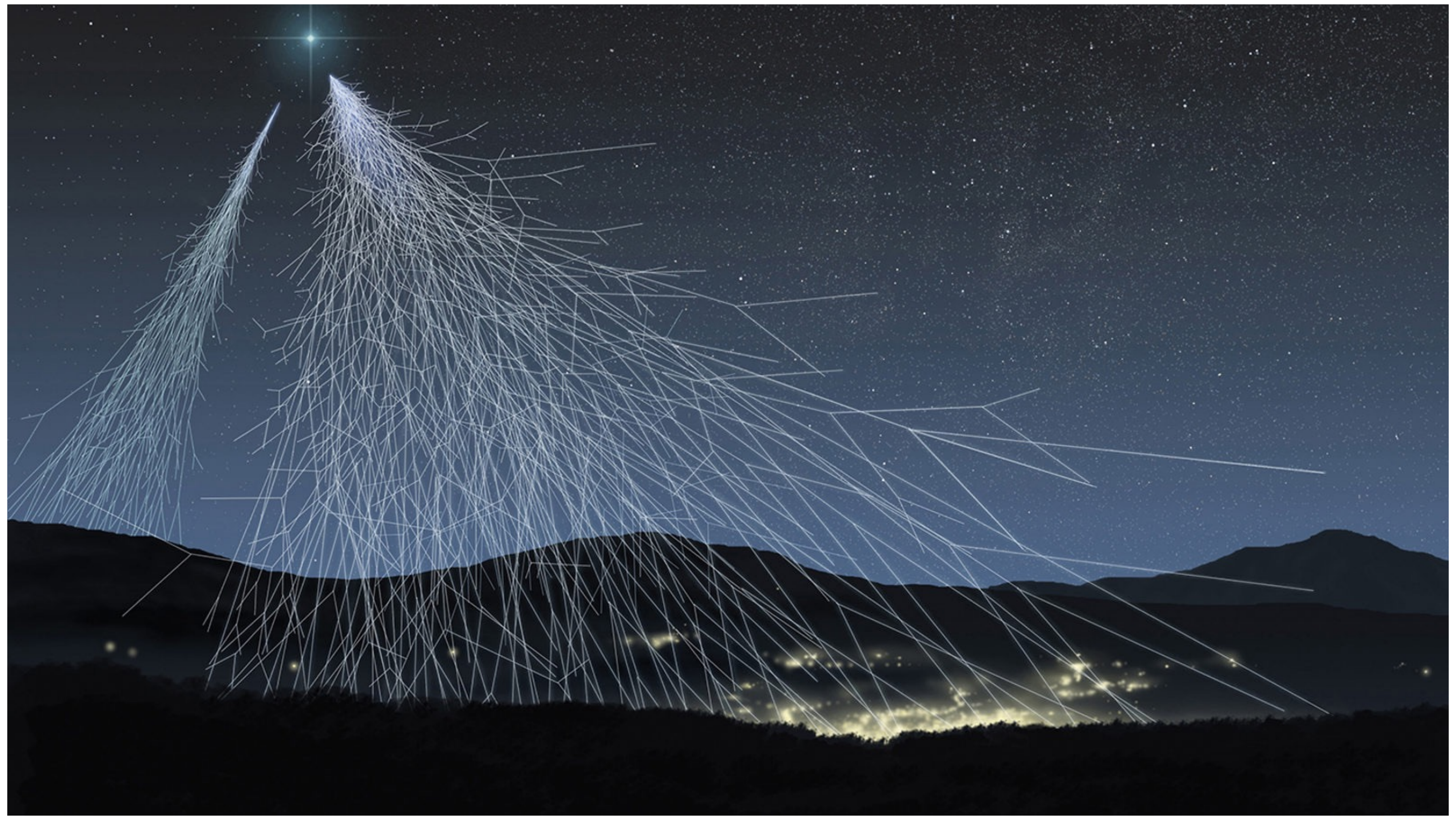
$$\nu_e + e^{-} \rightarrow \nu_e + e^{-}$$

45% ν_e missing



Picture of the sun with ν_e

Atmospheric neutrino deficit



In parallel, interest in neutrinos from cosmic rays

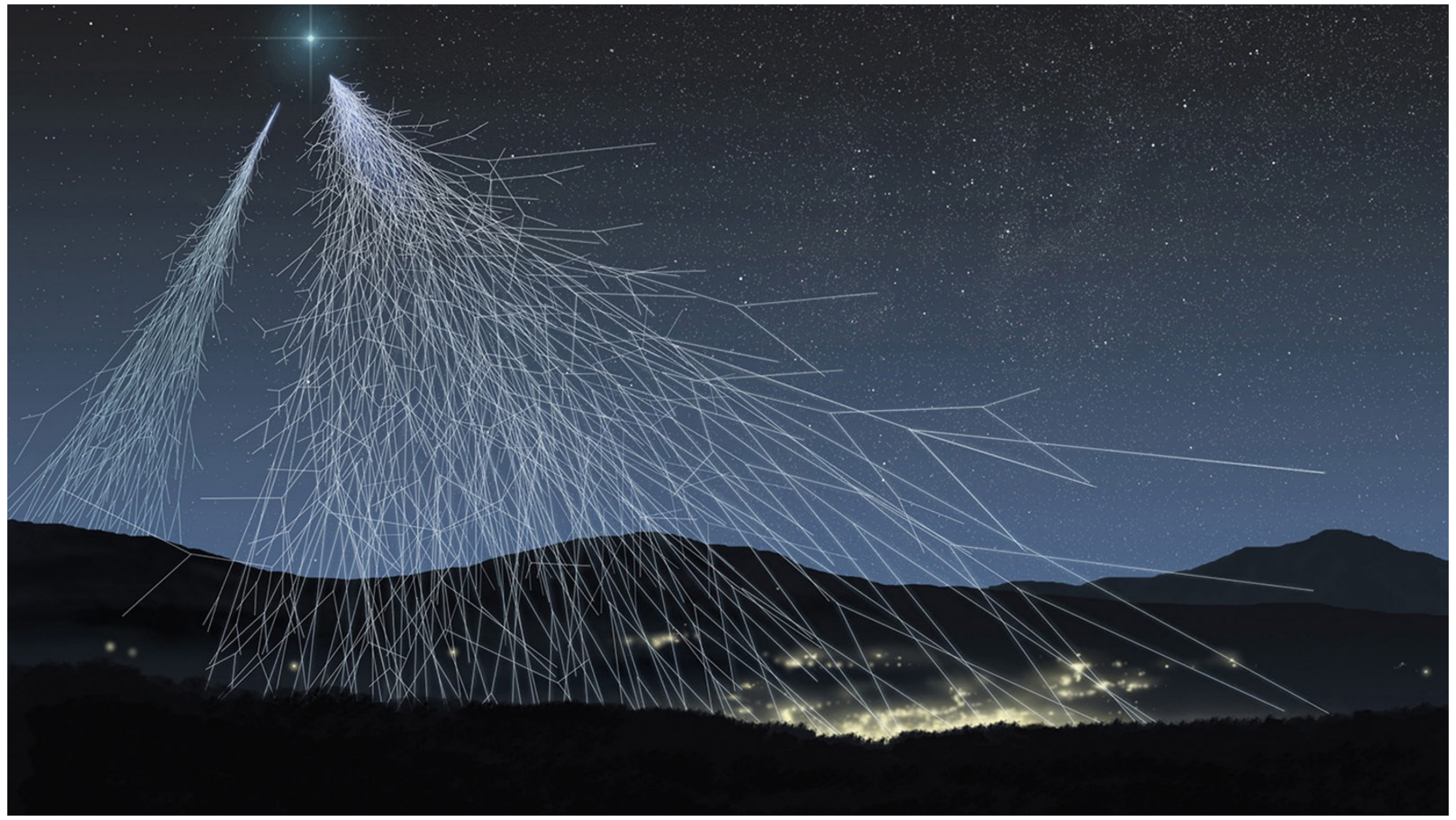
→ When cosmic rays hit earth, they interact with the atmosphere and produce pions and muons

$$p + atm \rightarrow \pi^+ + \dots$$
$$\pi^+ \rightarrow \mu^+ + \nu_\mu$$
$$\mu^+ \rightarrow e^+ + \nu_e + \bar{\nu}_\mu$$

At ground, we expect:

$\nu_\mu : \nu_e = 2 : 1$

Atmospheric neutrino deficit



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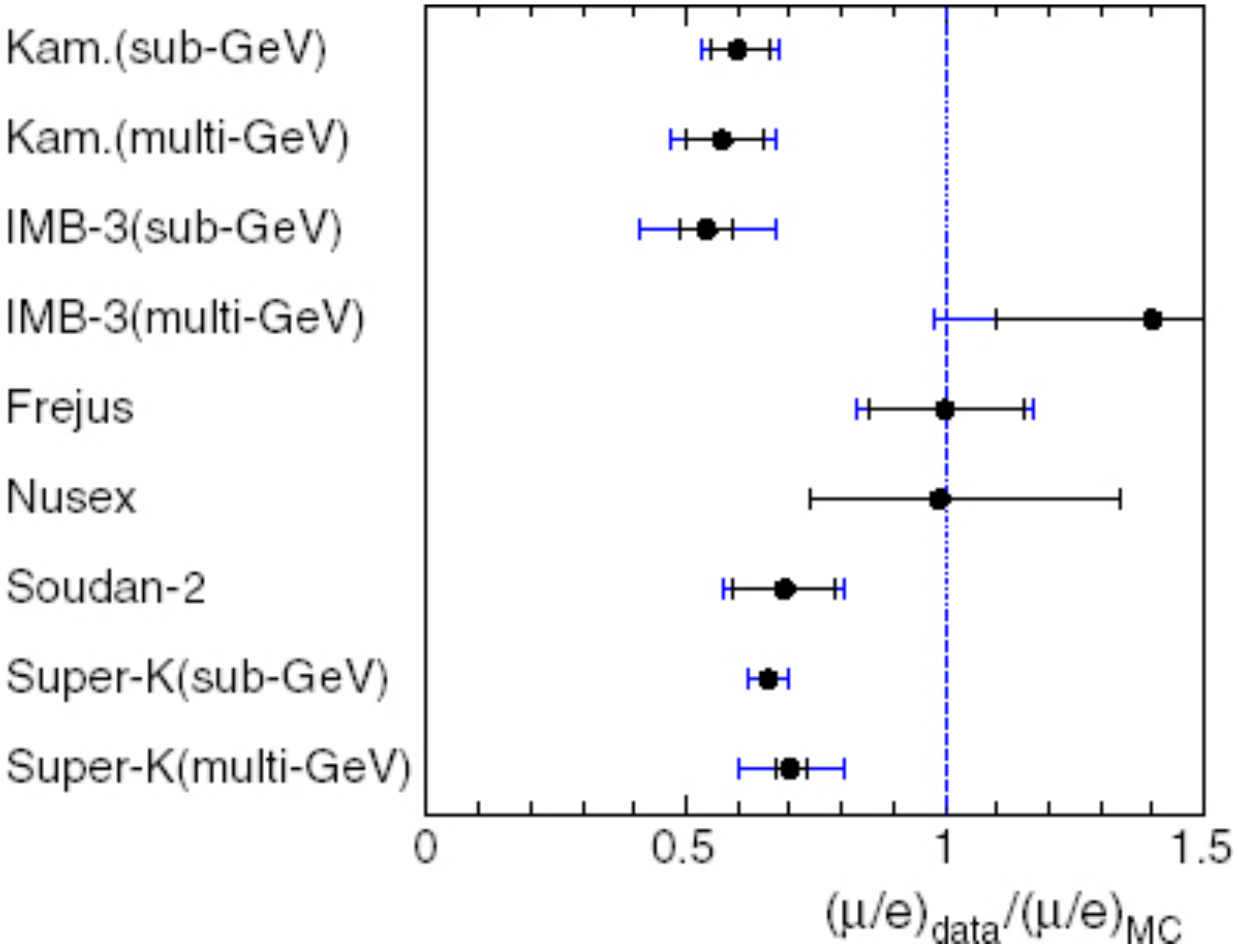

About 50% ν missing

Understanding the anomalies

- Several hypothesis to explain the anomalies:
  - Problems with fluxes computations, experiments
  - Neutrino behavior: ν-decay, ν-decoherence, flavor changing neutral currents, oscillations, …

- In 1957, Pontecorvo suggested the $\nu \rightarrow \bar{\nu}$ oscillations, in analogy with $K^0 \rightarrow \bar{K}^0$ mixing

- Principle : Neutrino flavor and mass eigenstates are **not superimposed** but **linked** by a 3×3 unitary mixing matrix (the PMNS matrix) :

$$|\nu_\alpha\rangle = \sum_{i=1}^{3} U^*_{\alpha i}|\nu_i\rangle$$

α = (e, μ, τ) := Flavor states
i = (1, 2, 3) := Mass states
U = PMNS matrix

Where PMNS stands for Pontecorvo-Maki-Nakagawa-Sakata

→ ***Neutrinos would be massive !***

Understanding the anomalies

Simplified 2 flavors case

The mixing matrix is written as : $\begin{pmatrix}\nu_\alpha \\ \nu_\beta\end{pmatrix} = \begin{pmatrix}\cos\theta & \sin\theta \\ -\sin\theta & \cos\theta\end{pmatrix}\begin{pmatrix}\nu_1 \\ \nu_2\end{pmatrix}$

- θ is the mixing angle
- Δm^2 is the mass splitting :

$$\Delta m^2 = m_1^2 - m_2^2$$

Understanding the anomalies

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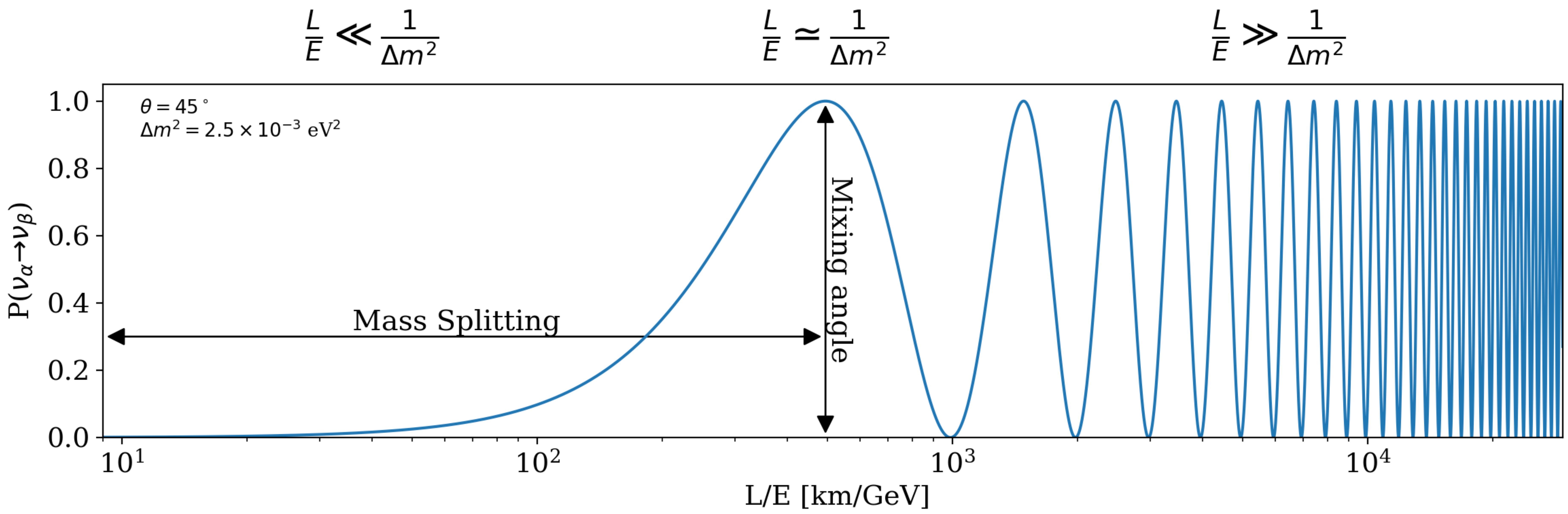
- θ is the mixing angle
- Δm² is the mass splitting :

$$\Delta m^2 = m_1^2 - m_2^2$$

With a source ν_α at an energy E, the probability to detect a ν_β at a distance L is :

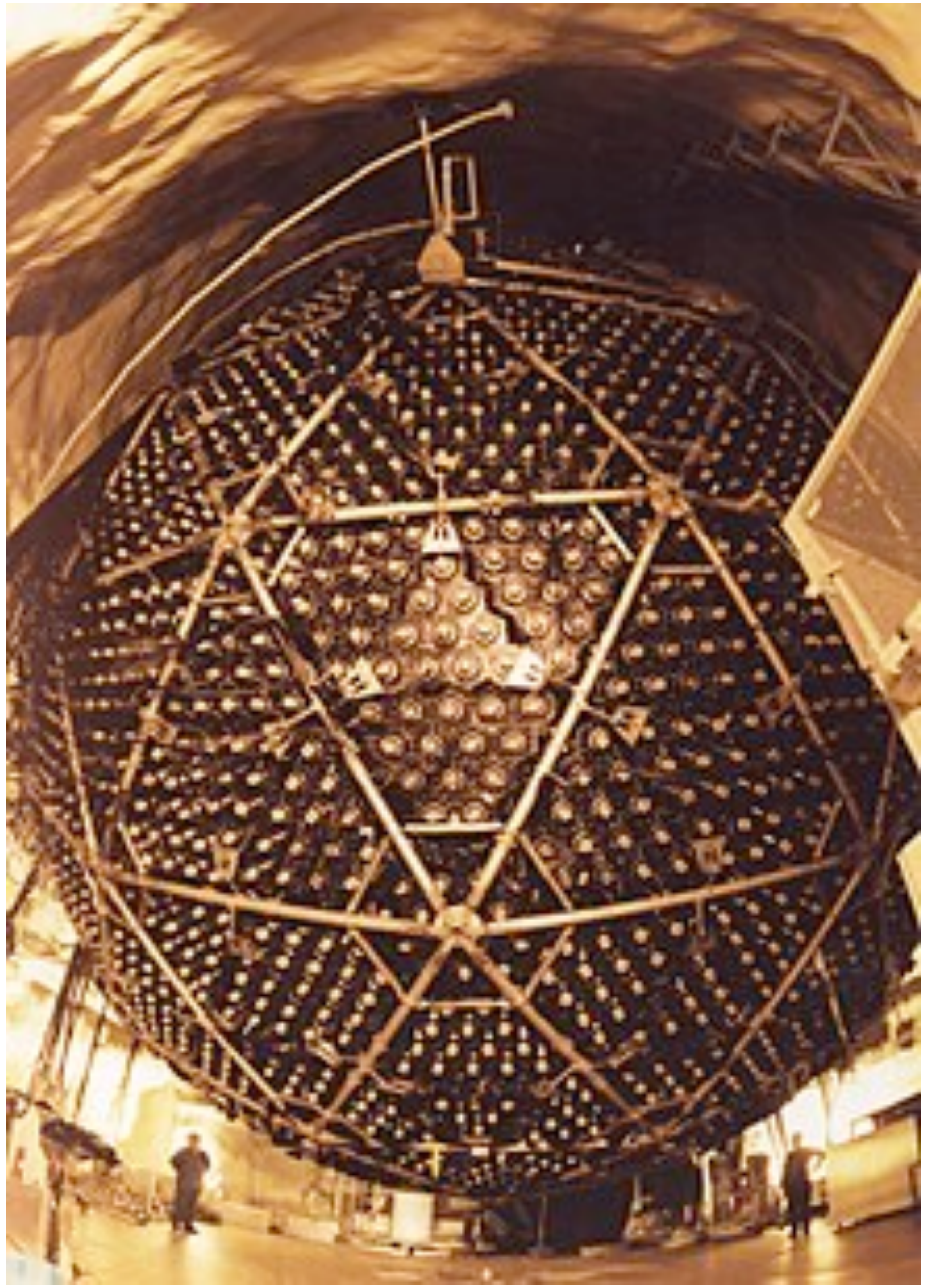
$$P(\nu_\alpha \to \nu_\beta) = \sin^2(2\theta)\sin^2\left(\frac{\Delta m^2 L}{4E}\right)$$

- L : source → detector distance
- E : neutrino energy



→L, E can be tuned by experiments to measure θ, Δm²

Proofs of neutrino oscillations - Solar



SNO (1kton of heavy water) was designed to detect solar neutrinos through:

- ○ **CC** interactions $\nu_e + d \rightarrow p + p + e^-$
  ν_e only (ν_μ & ν_τ don't have enough energy)
- ○ **ES** interactions $\nu_x + e^- \rightarrow \nu_x + e^-$
  all flavors
- ○ **NC** interactions $\nu_x + d \rightarrow p + n + \nu_x$
  all flavors

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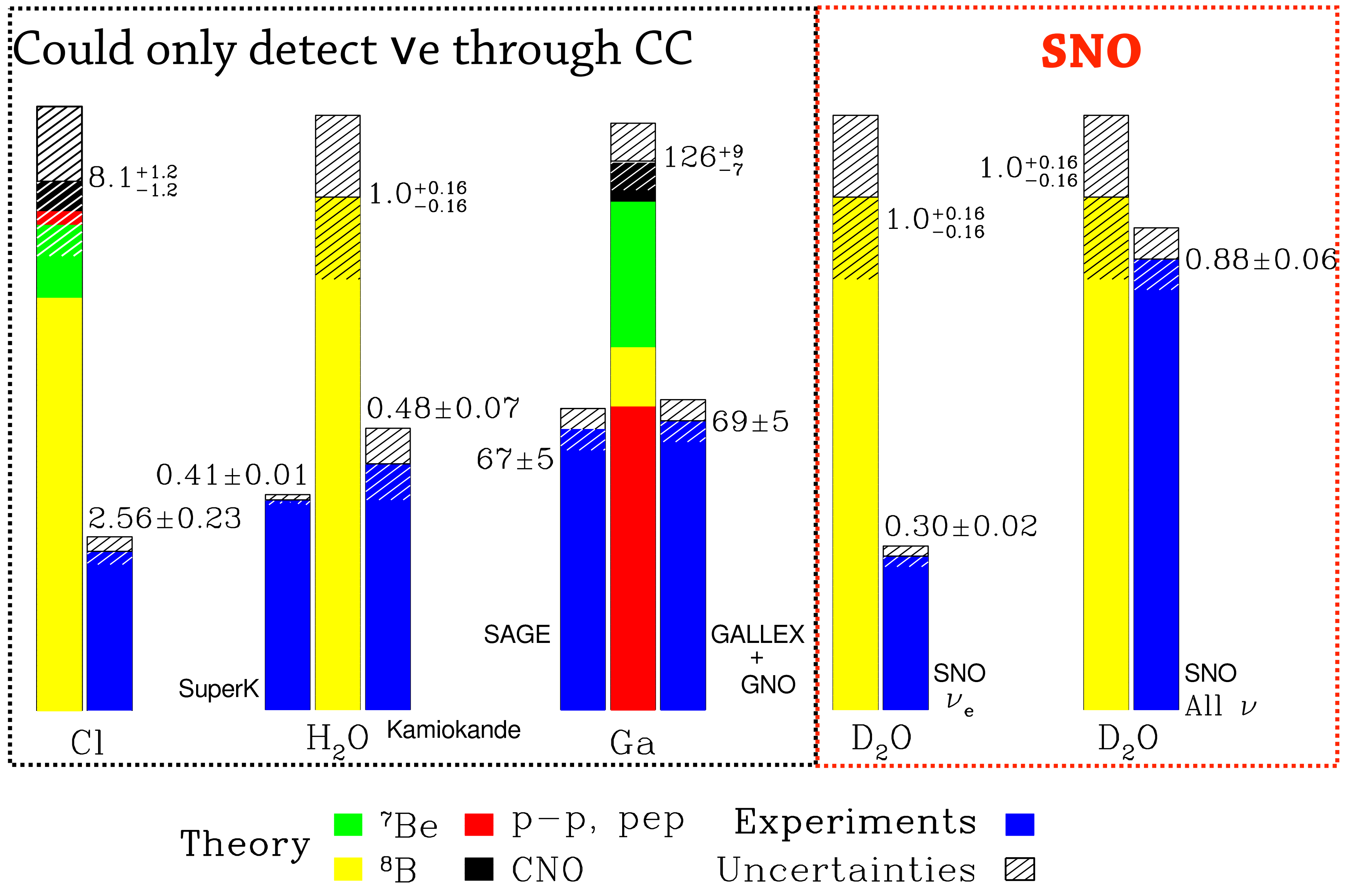

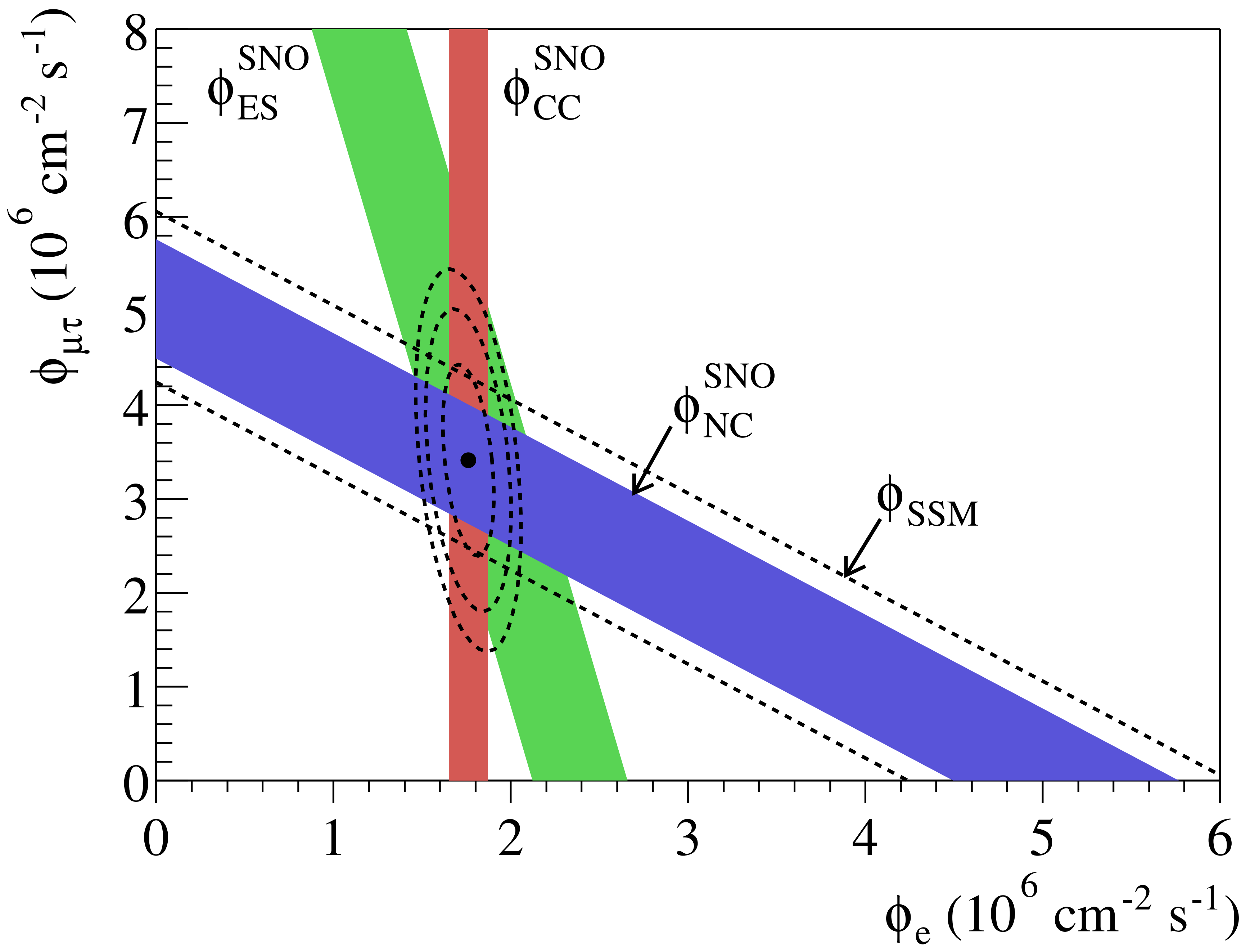

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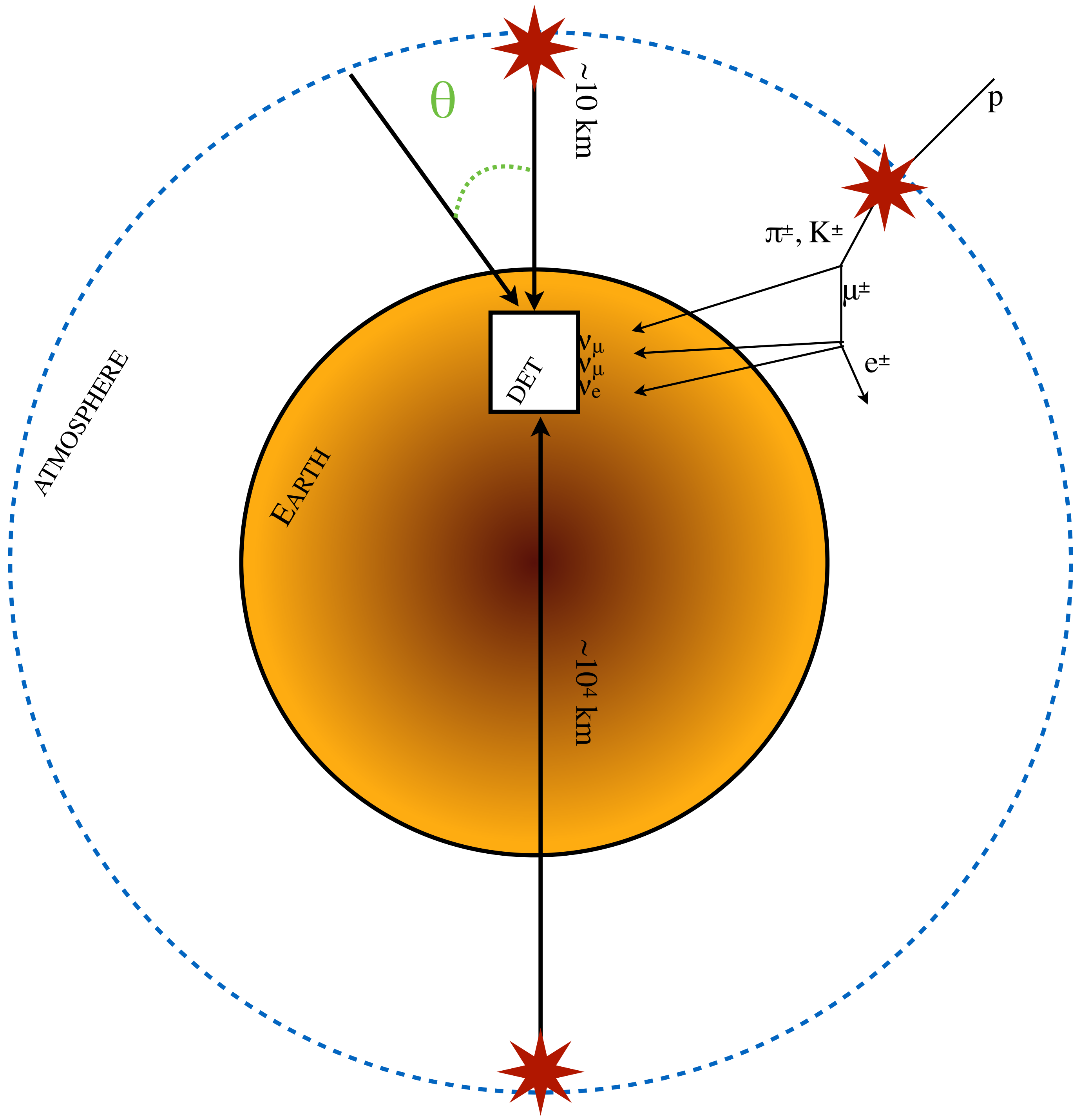
SNO measured the ratio : $\dfrac{\Phi_{CC}}{\Phi_{NC}} = 0.34 \pm 0.023(\text{stat.})^{+0.029}_{-0.031}$

And showed that the **total** flux of solar neutrino is **compatible** with the solar standard model

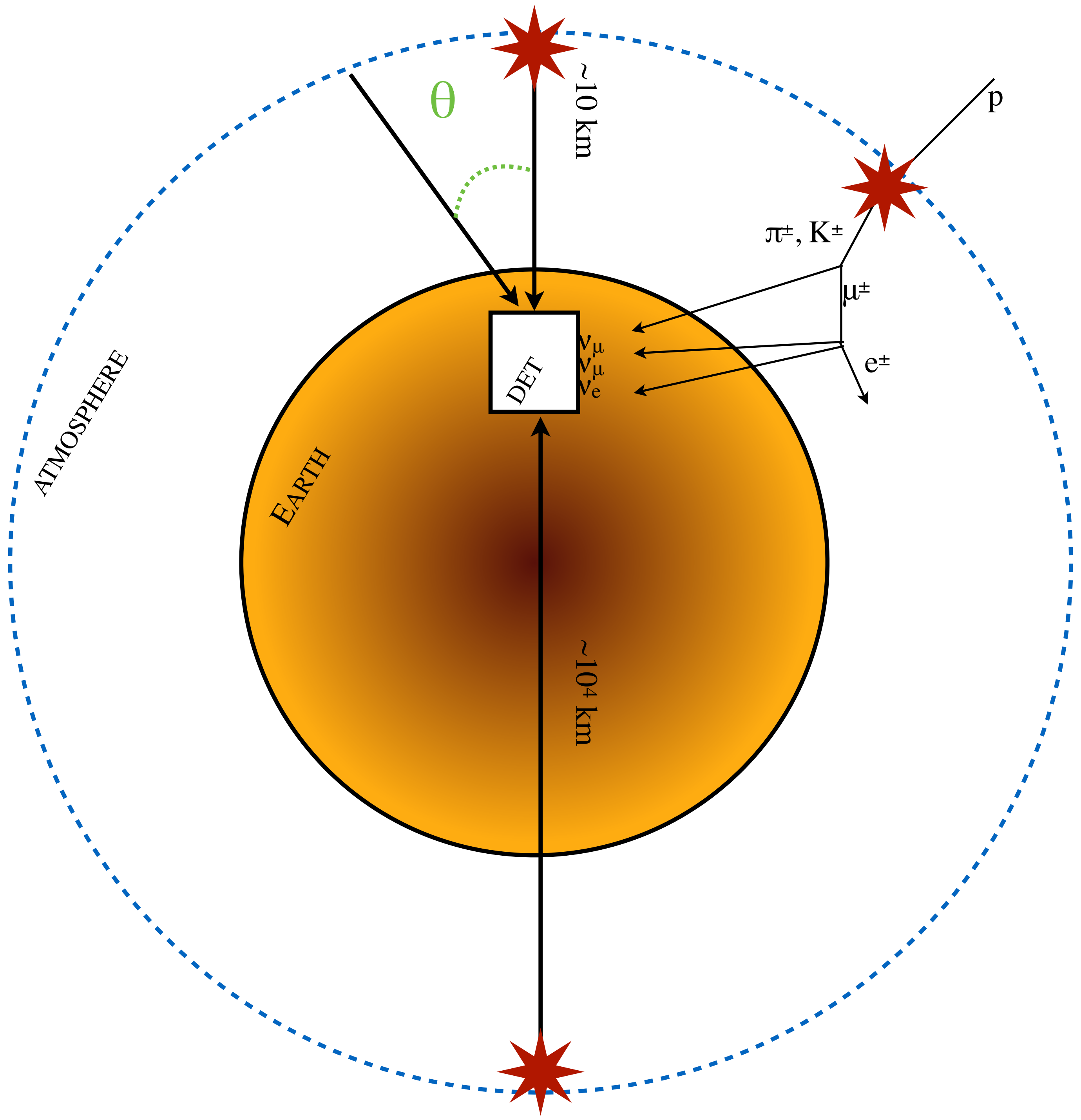
SNO proved that neutrino change flavors

Proofs of neutrino oscillations - Atmospheric



Super-Kamiokande measured the atmospheric ν_e and ν_μ energy as a function of cosθ ↔ L

Proofs of neutrino oscillations - Atmospheric



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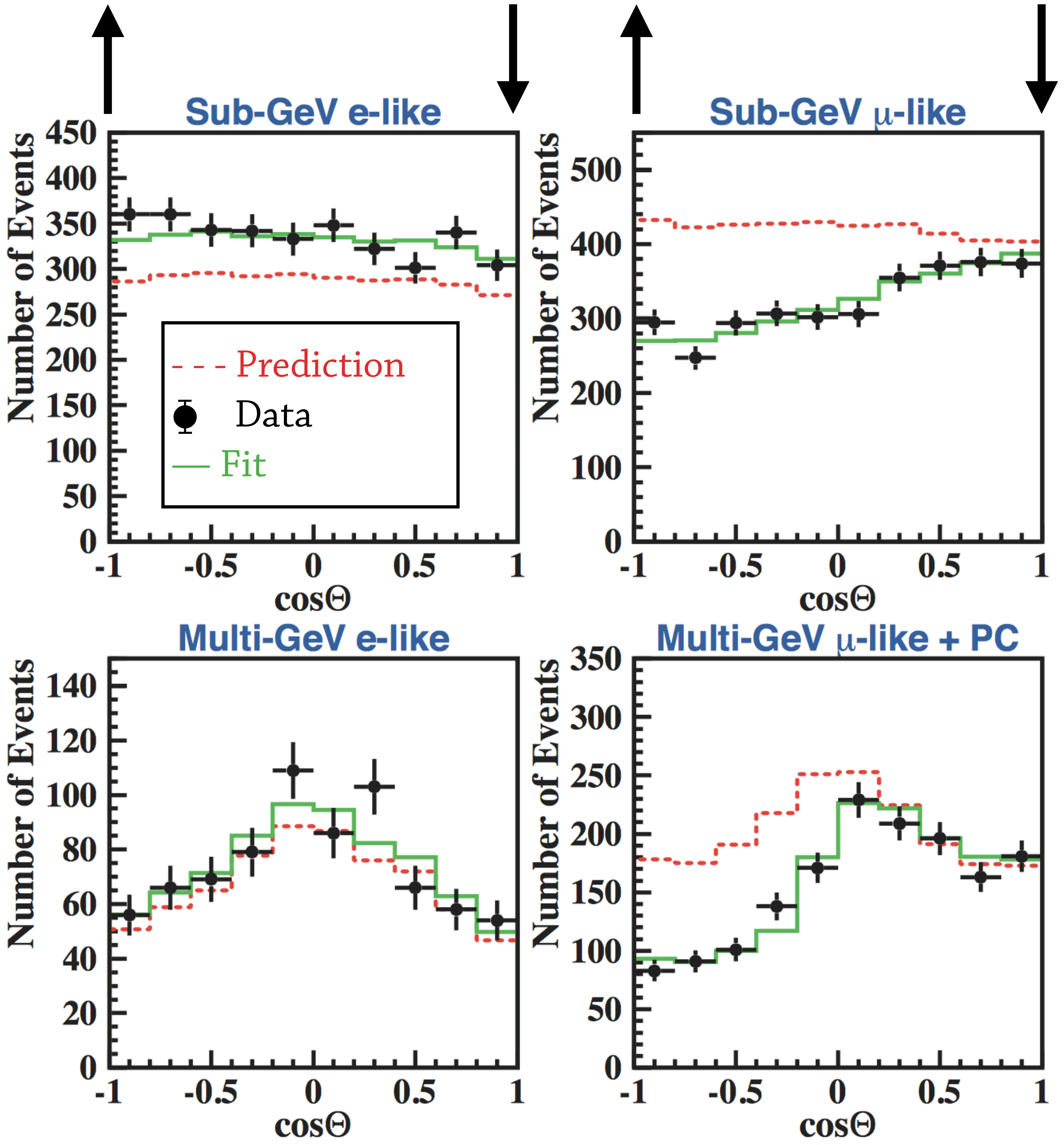


For ν_e :

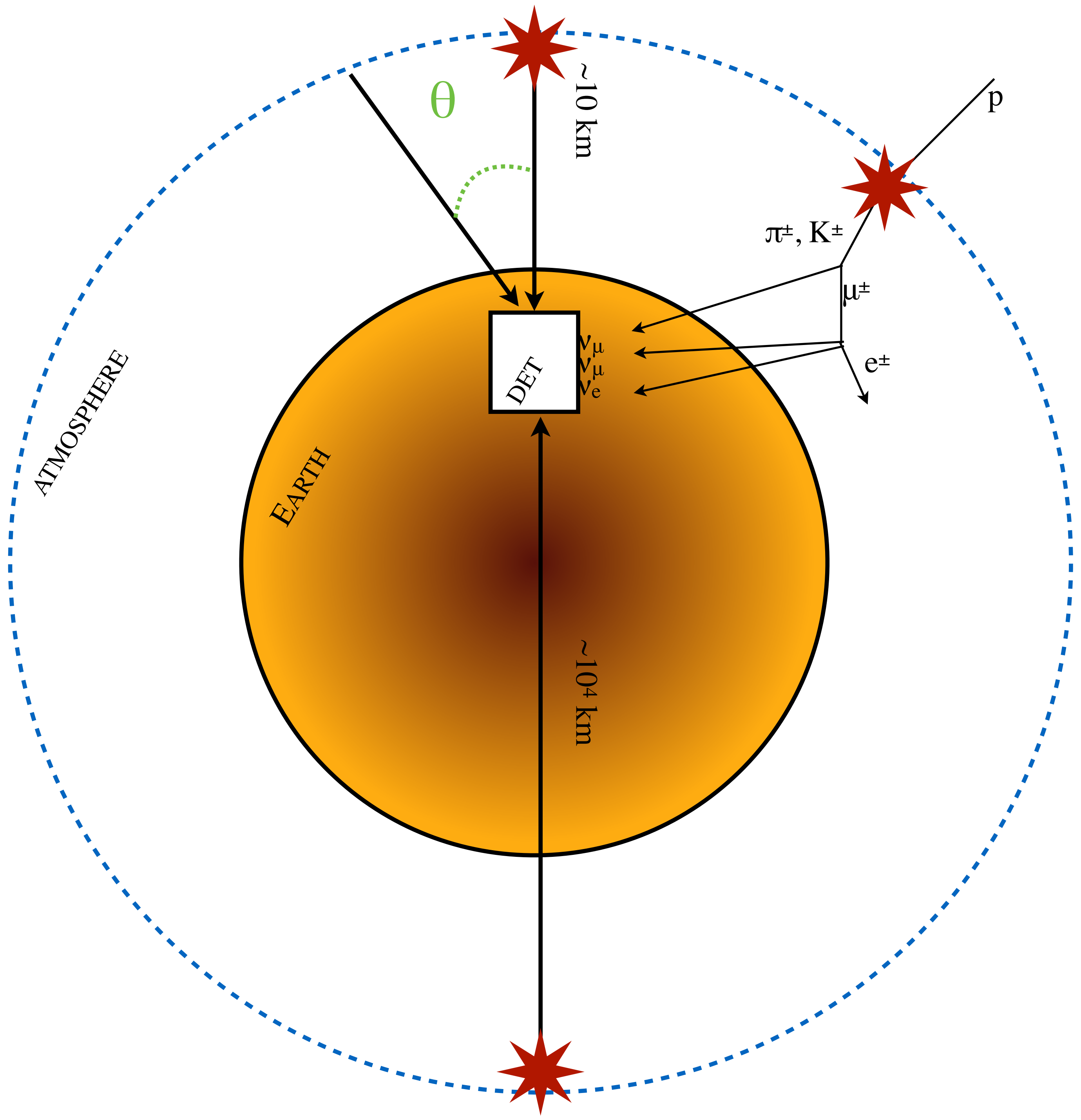
As predicted for all directions and energy

For ν_μ :

Loss of upwards ν_μ (L ~ 10^4 km)

As expected for downwards ν_μ (L ~ 10 km)

Proofs of neutrino oscillations - Atmospheric



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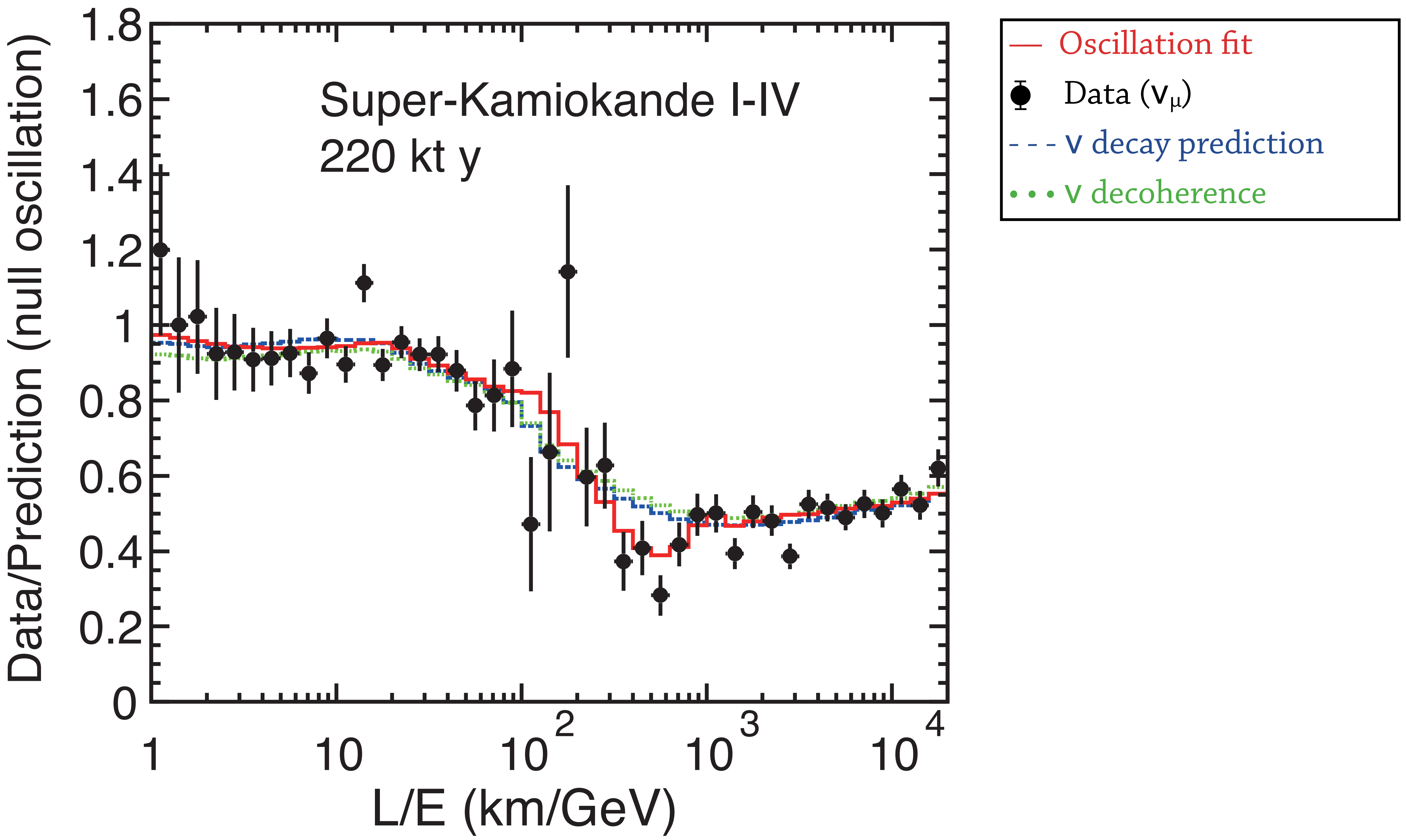
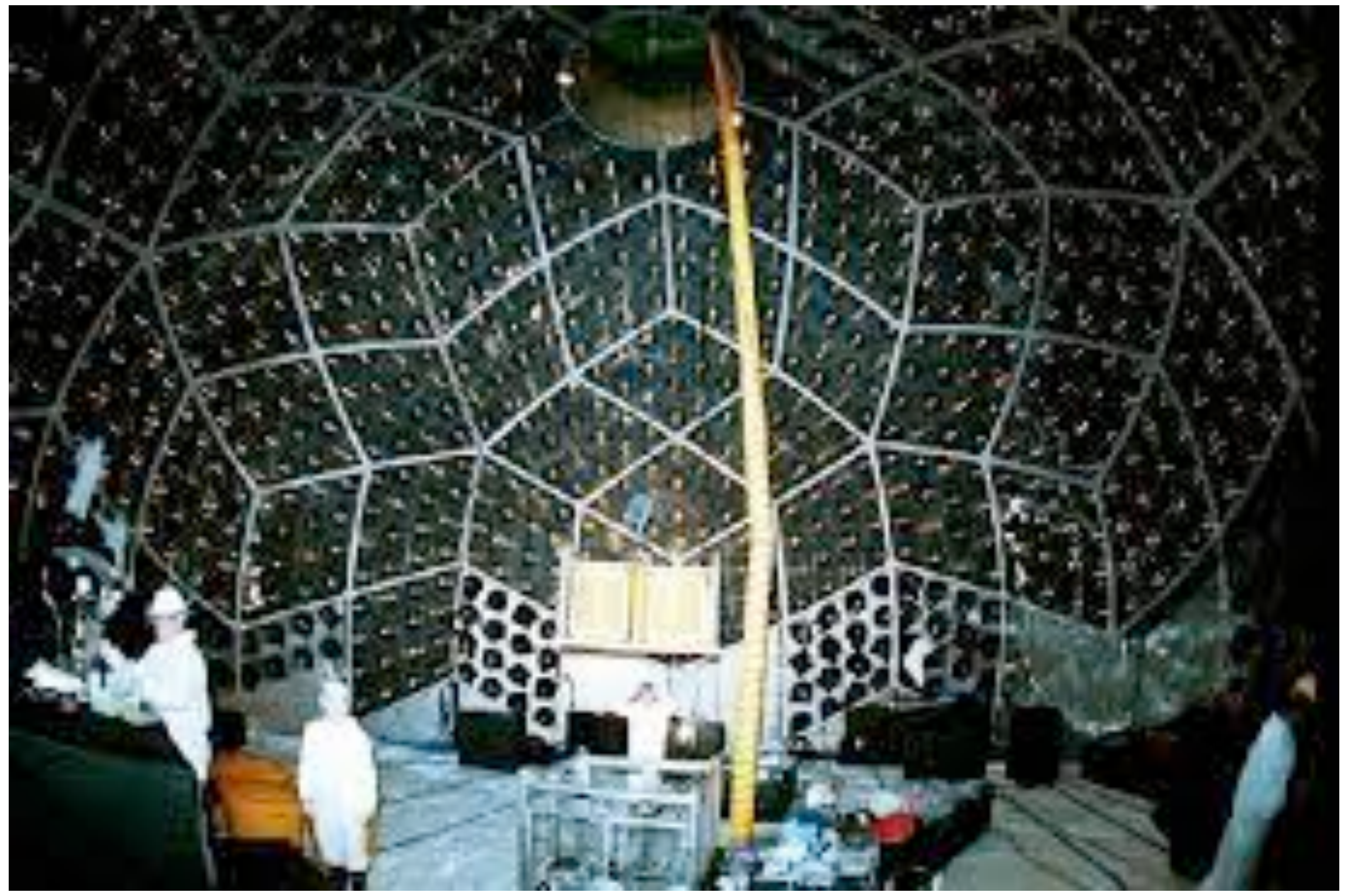
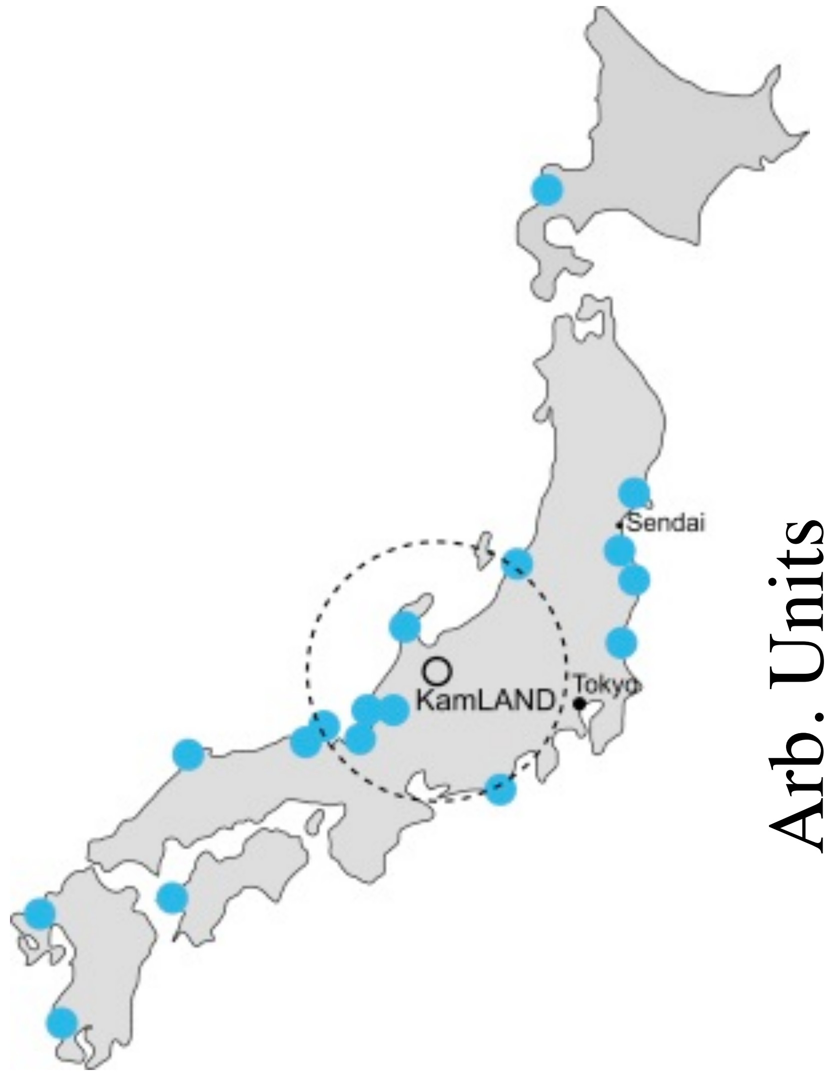


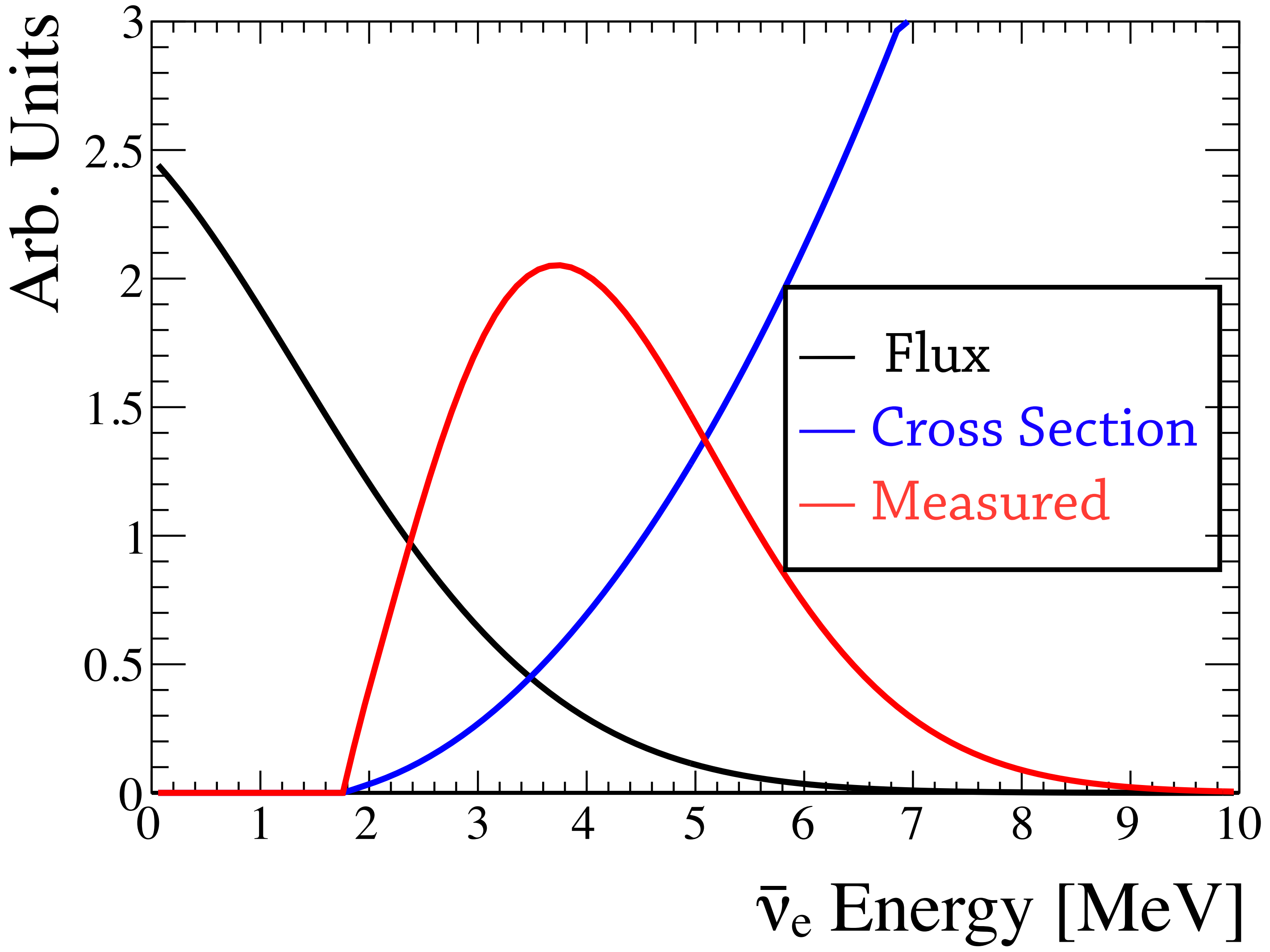
Super-Kamiokande proved that ν_μ disappear as a function of L/E (possibly into ν_τ)

Proofs of neutrino oscillations - Reactors



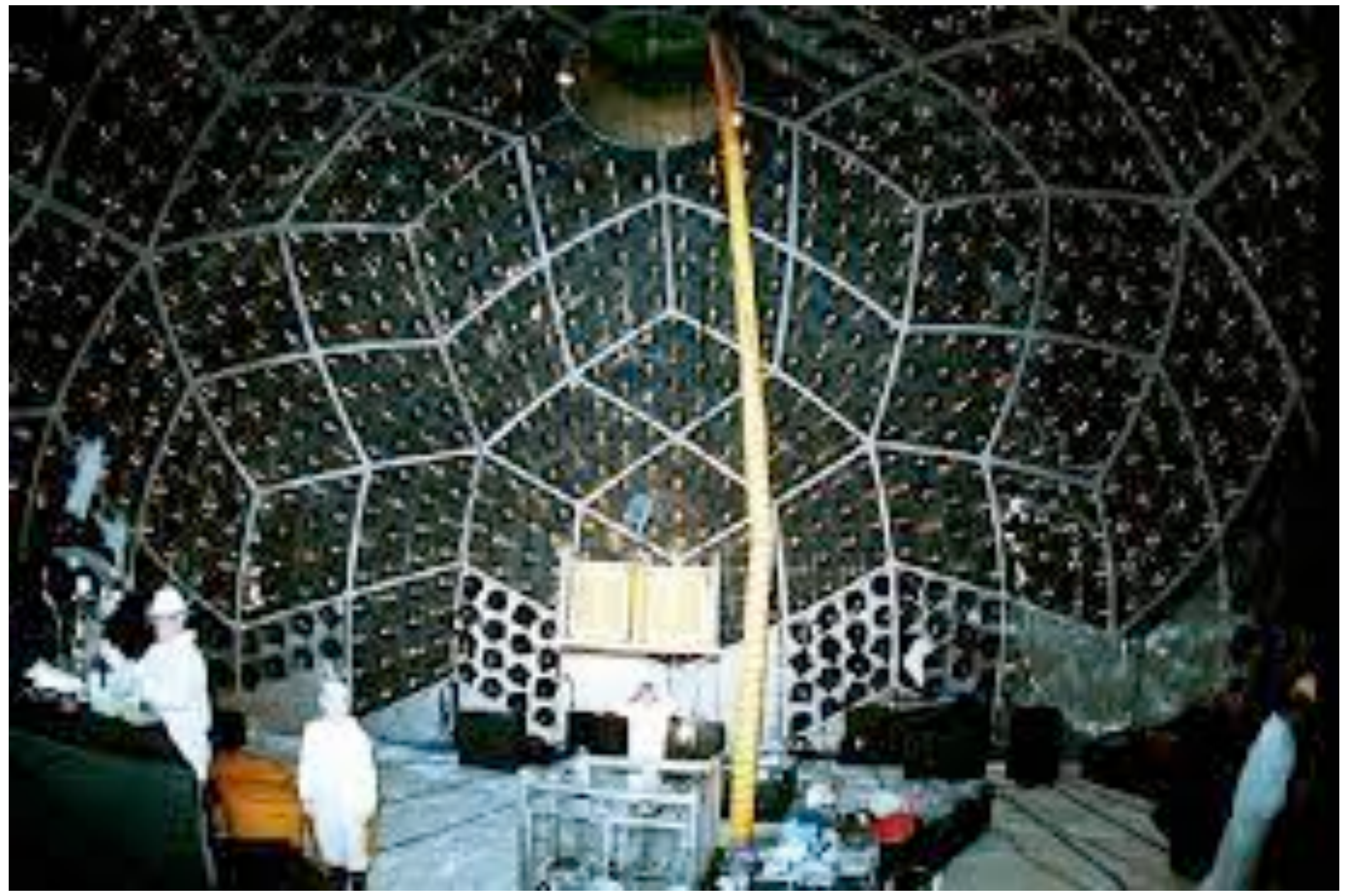


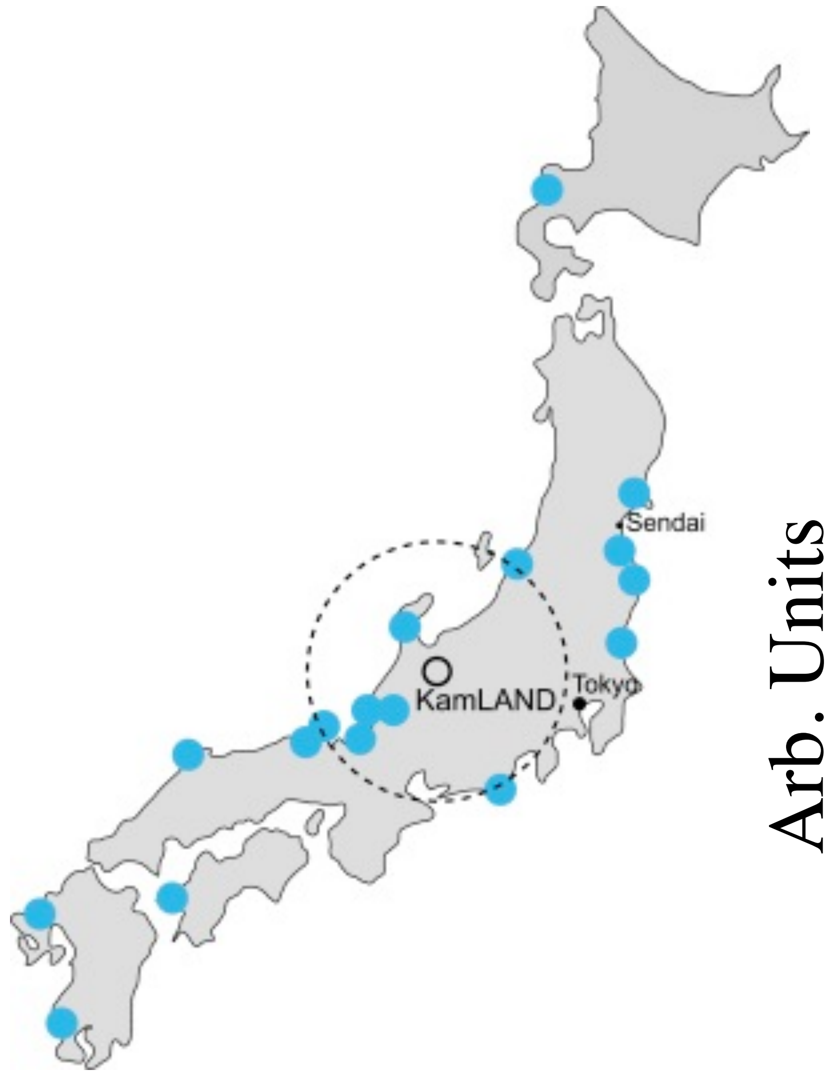
Kamland experiment in Japan measured the $\bar{\nu}_e$ flux from 53 nuclear reactors (L_{mean} ~ 180 km)



Reactor $\bar{\nu}_e$ spectrum up to ~10 MeV
-> Cannot measure appearance of new flavors

Proofs of neutrino oscillations - Reactors





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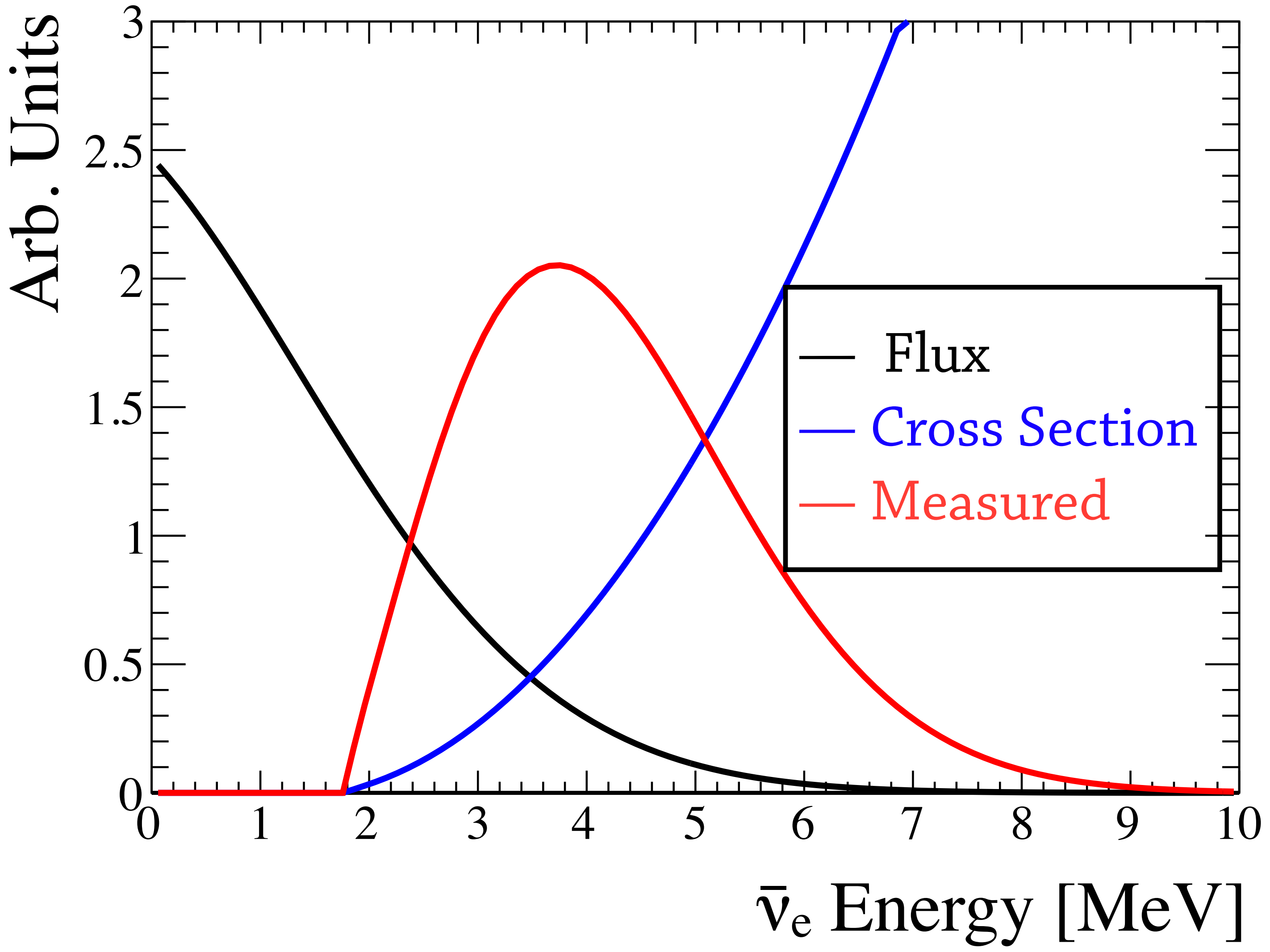

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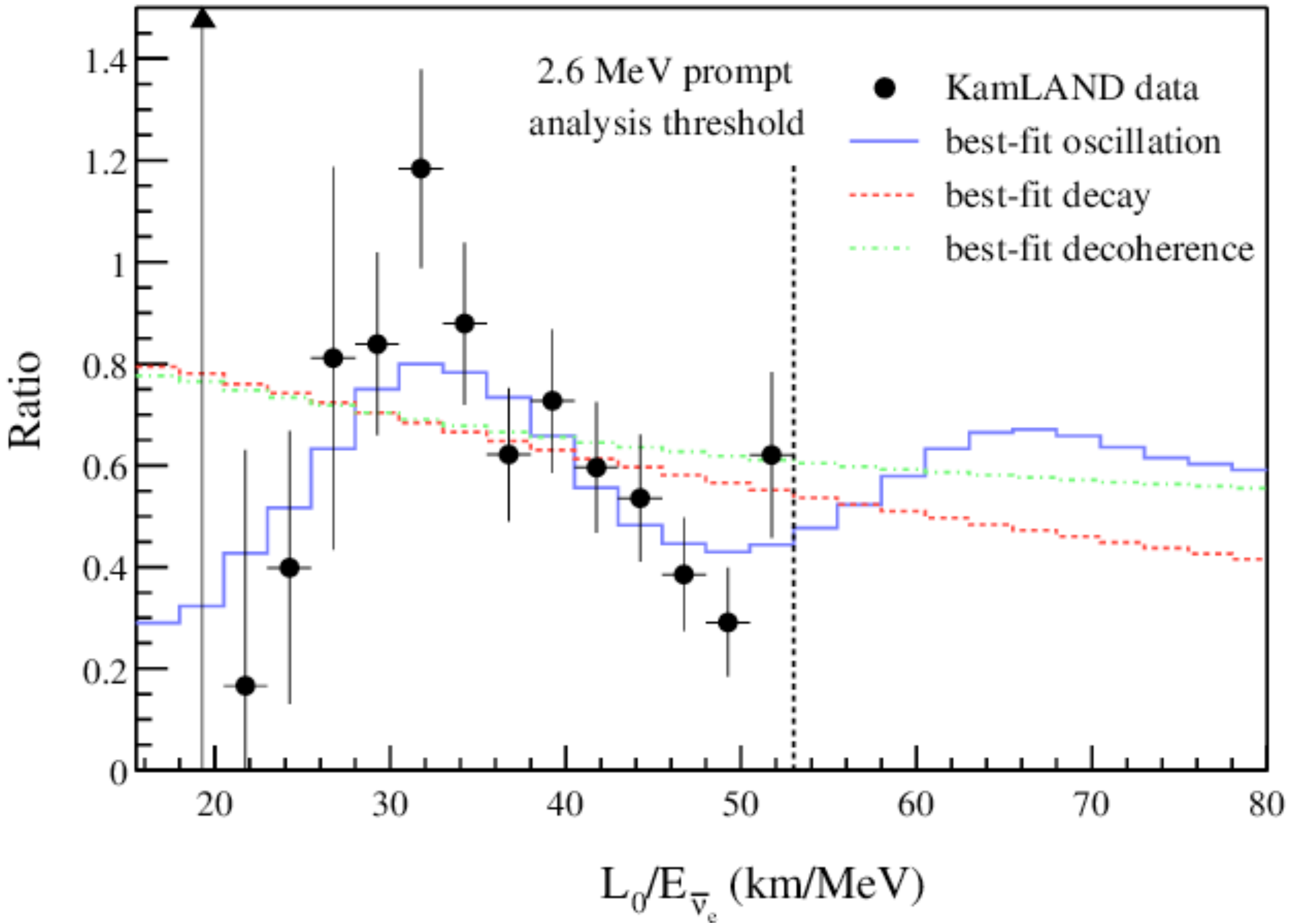
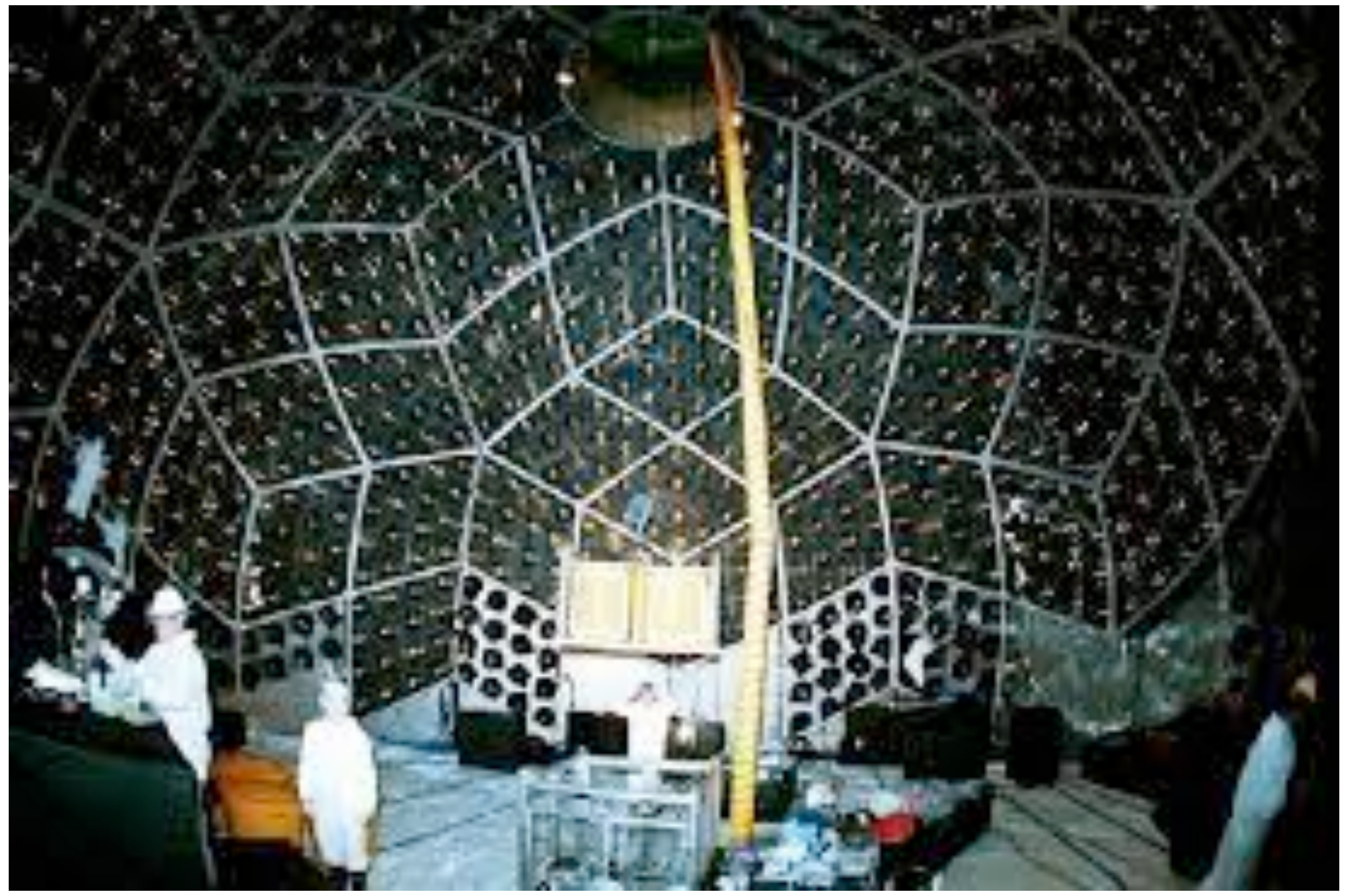


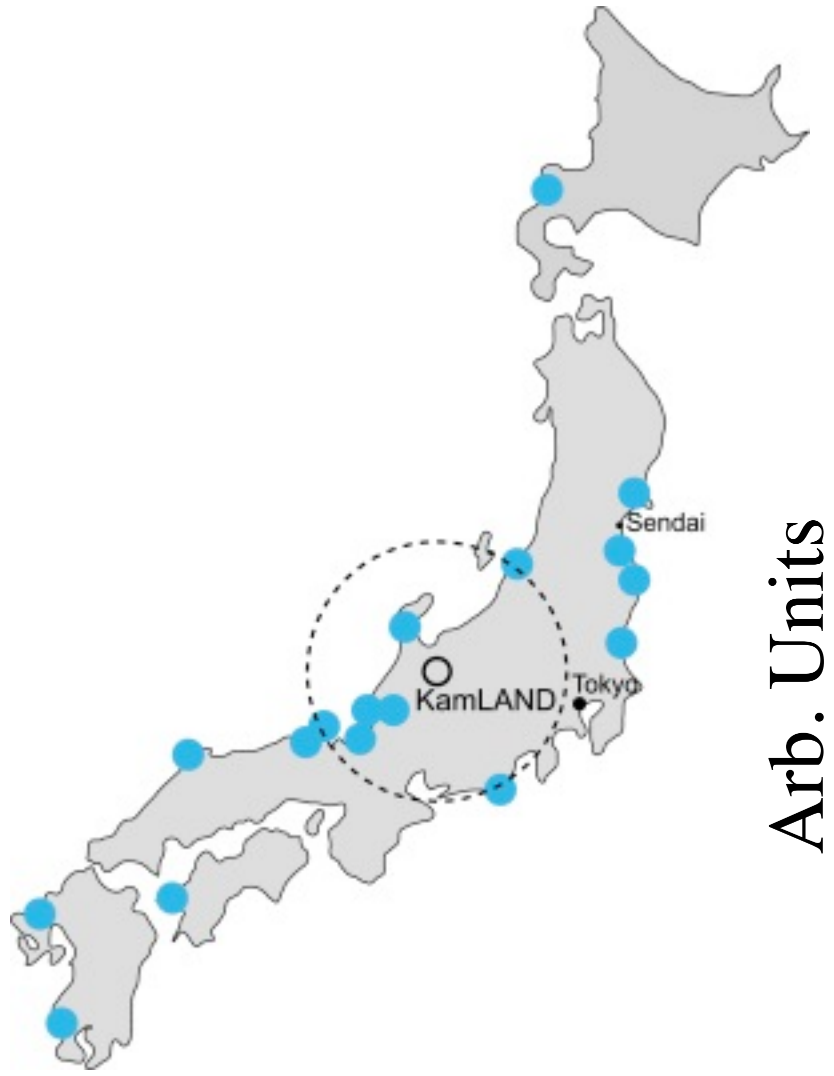
- First results -

○ Rejection of the ν-decay and ν-decoherence hypotheses

○ ν-oscillation preferred

Proofs of neutrino oscillations





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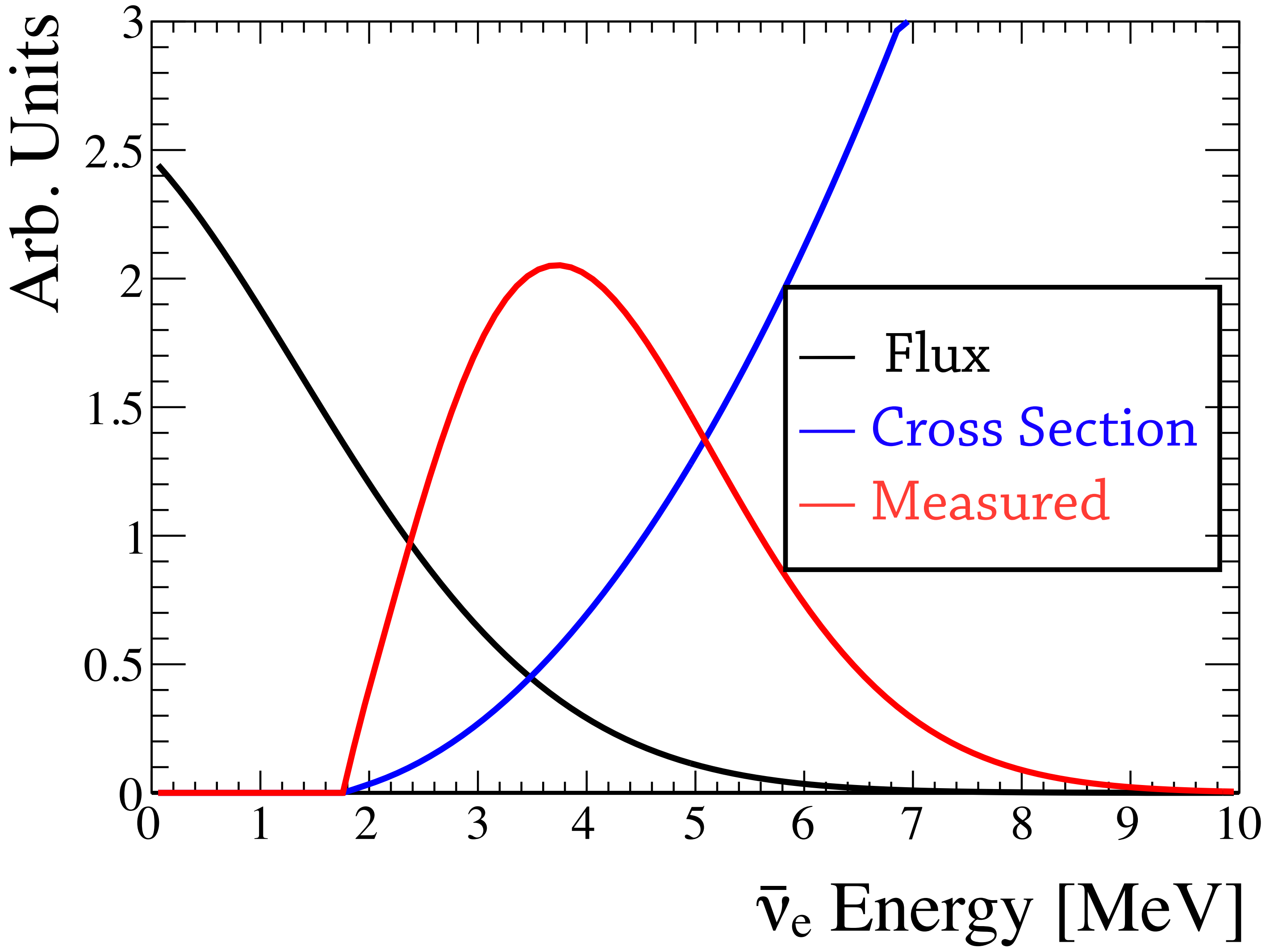

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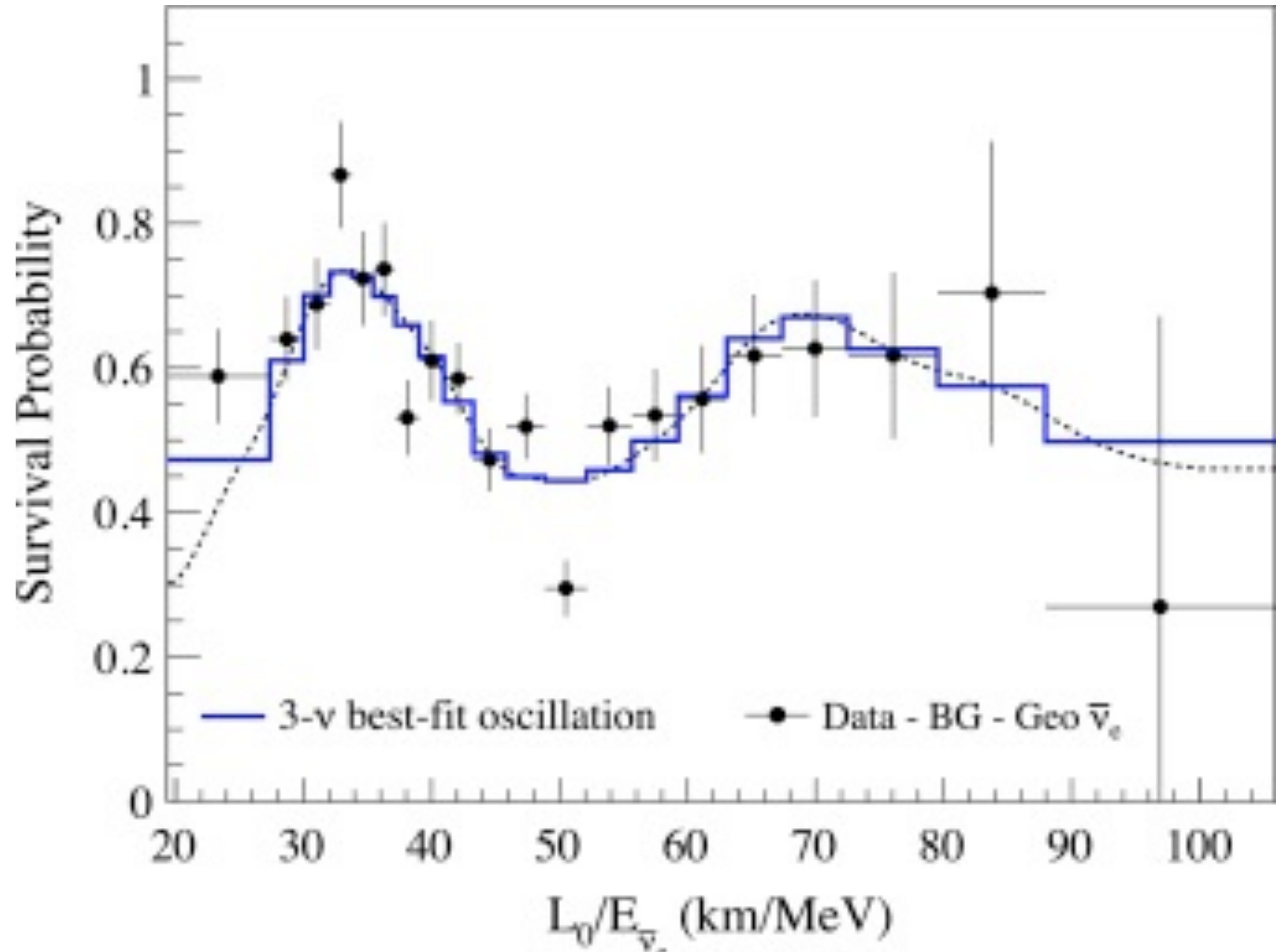


- Final results -

○ Very clear L/E pattern

○ Can see the disappearance dip, and re-appearance of $\bar{\nu}_e$!

KAMLAND proved that $\bar{\nu}_e$ oscillates !

NEUTRINO OSCILLATIONS

Oscillations with 3 flavors

The PMNS 3×3 unitary mixing matrix can be written as :

$$\begin{pmatrix}\nu_e\\ \nu_\mu\\ \nu_\tau\end{pmatrix} = \overset{\text{Atmospheric}}{\begin{pmatrix}1 & 0 & 0\\ 0 & c_{23} & s_{23}\\ 0 & -s_{23} & c_{23}\end{pmatrix}} \overset{\text{Reactor/Accelerator}}{\begin{pmatrix}c_{13} & 0 & s_{13}e^{-i\delta}\\ 0 & 1 & 0\\ -s_{13}e^{i\delta} & 0 & c_{13}\end{pmatrix}} \overset{\text{Solar}}{\begin{pmatrix}c_{12} & s_{12} & 0\\ -s_{12} & c_{12} & 0\\ 0 & 0 & 1\end{pmatrix}} \begin{pmatrix}\nu_1\\ \nu_2\\ \nu_3\end{pmatrix}$$

$$s_{ij} = \sin\theta_{ij}$$
$$c_{ij} = \cos\theta_{ij}$$
$$\Delta m^2_{ij} = m^2_i - m^2_j$$

There are 3 mass splittings : $\Delta m^2_{21},\ \Delta m^2_{31},\ \Delta m^2_{32}$

But only two are relevant, since : $\Delta m^2_{12} + \Delta m^2_{23} + \Delta m^2_{31} = 0$

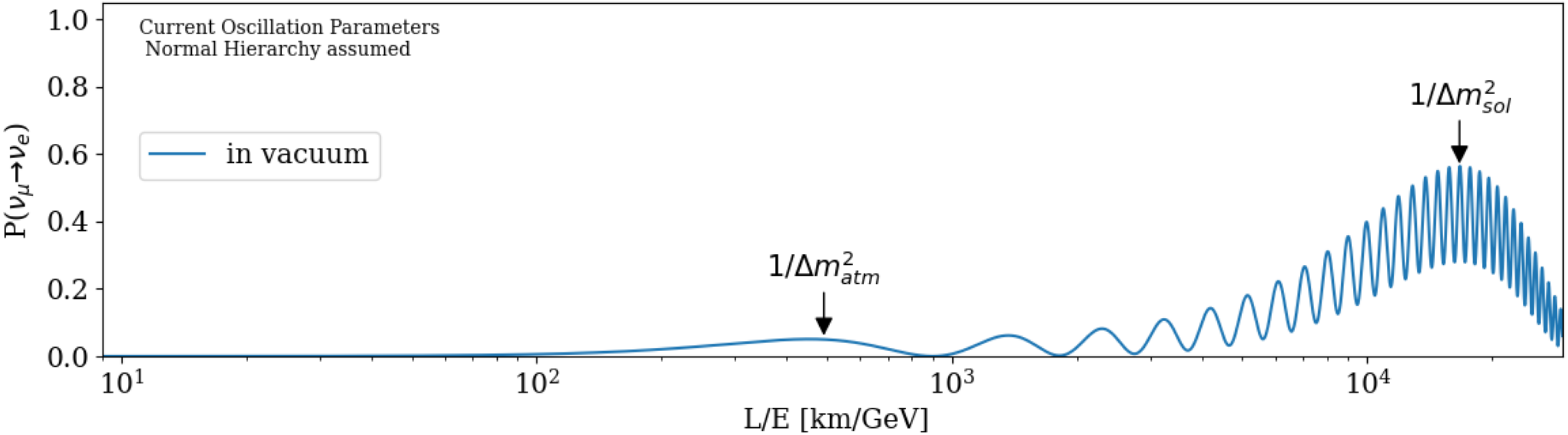
In the 3 ν-flavor case, the oscillation phenomena is described by:

- **3** mixing angles: θ_{12}, θ_{23} and θ_{13}
- **2** mass splittings: $\Delta m^2_{21} = \Delta m^2_{sol}$ and $\Delta m^2_{31} = \Delta m^2_{atm}$
- **1** CP violation phase δ

Oscillations with 3 flavors in vacuum

The oscillation probability is written as : $P(\nu_\alpha \to \nu_\beta) = |\langle \nu_\beta(L)|\nu_\alpha\rangle|^2 = \left| \sum_j U^*_{\alpha j} U_{\beta j} e^{-i\frac{m_j^2 L}{2E}} \right|^2$

As we have 2 mass splittings, we have 2 oscillation frequencies interfering :

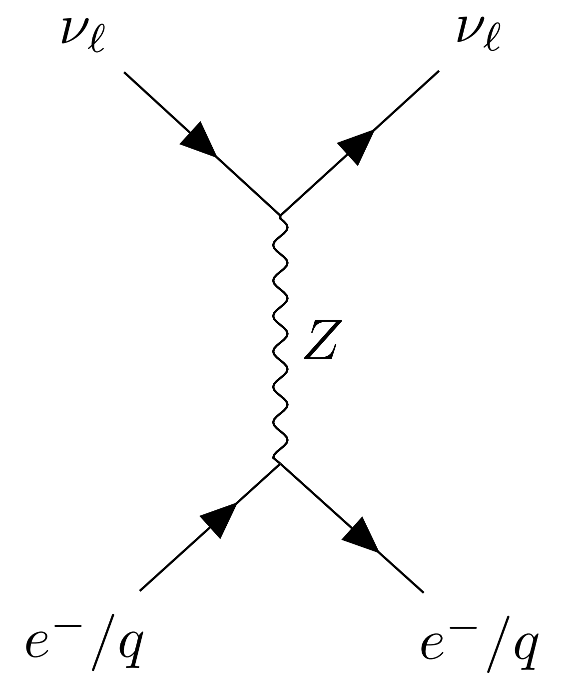


Oscillations with 3 flavors in matter

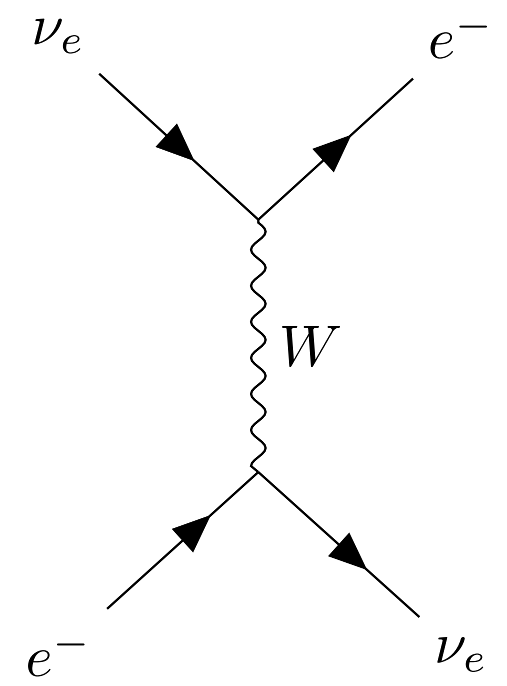
When neutrinos propagates in matter, the oscillations probabilities are modified

The e-type neutrinos have an extra interaction potential with electrons through charged current

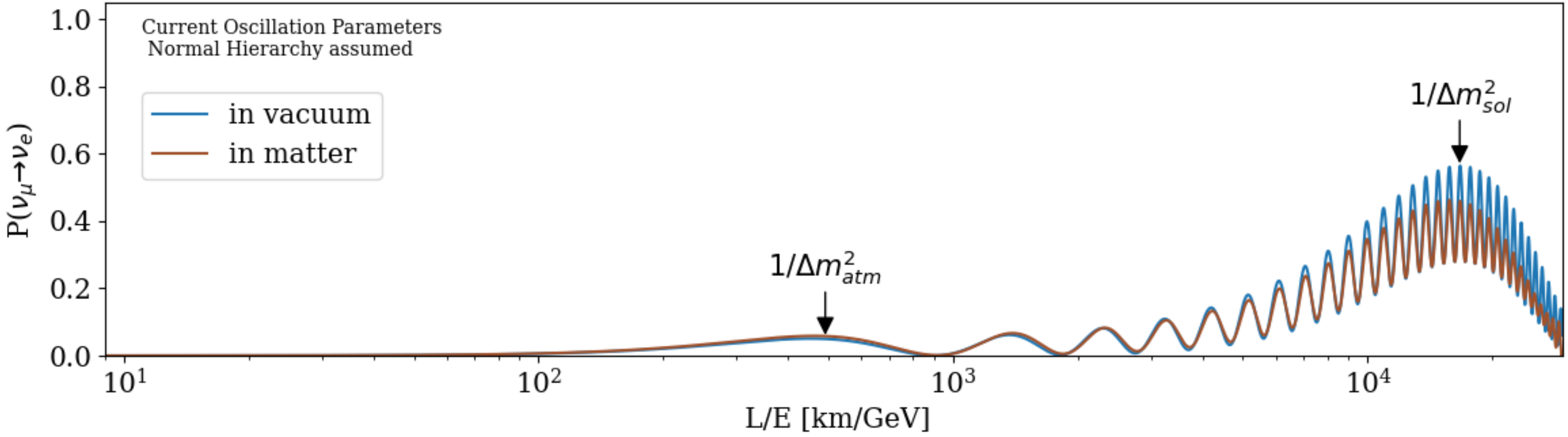
All neutrinos



ν_e only



Even earth crust has an impact on neutrino oscillation (denser matter -> stronger effect)



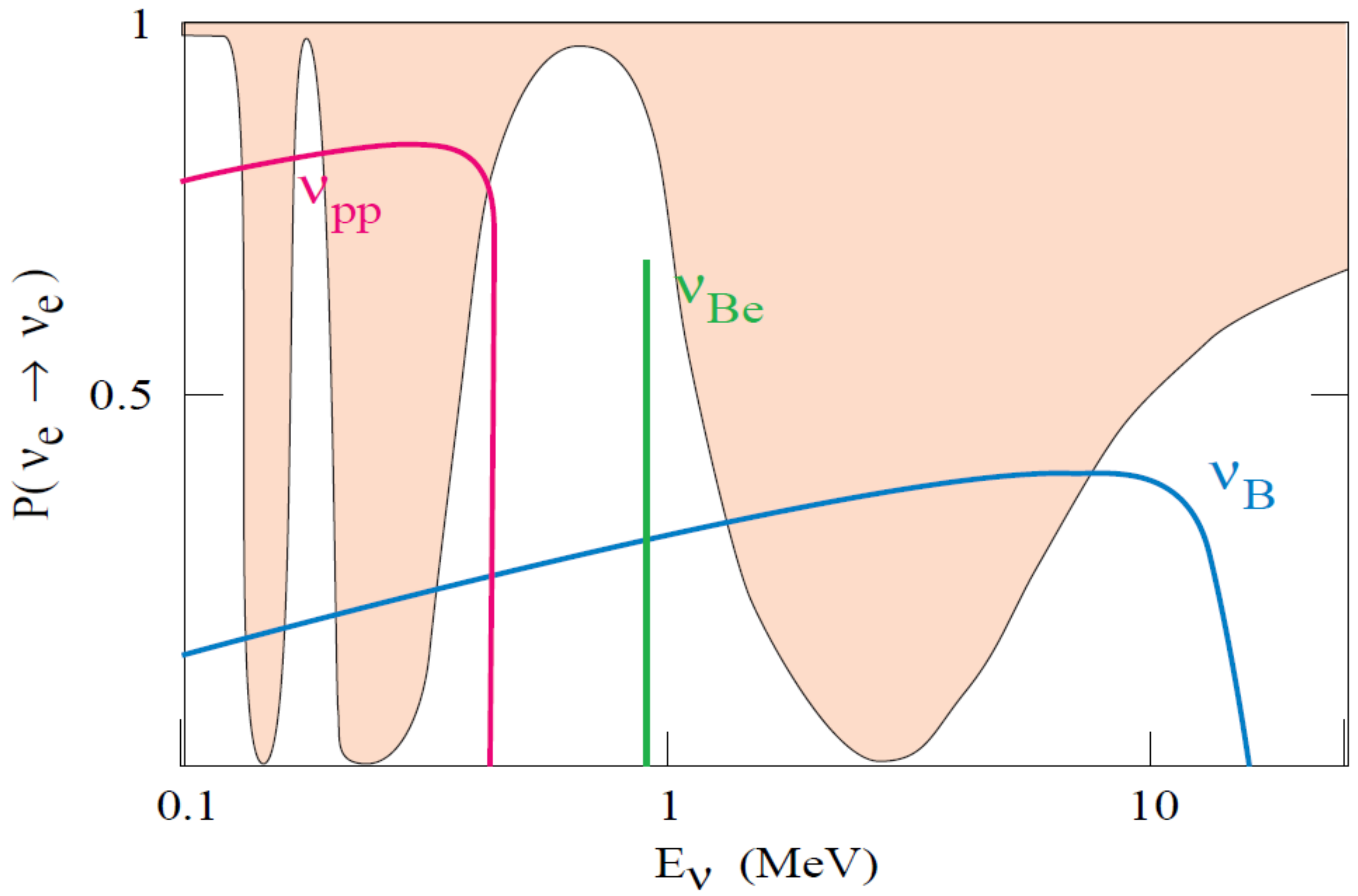
Oscillations parameters

Values of the oscillation parameters as of July 2022 :

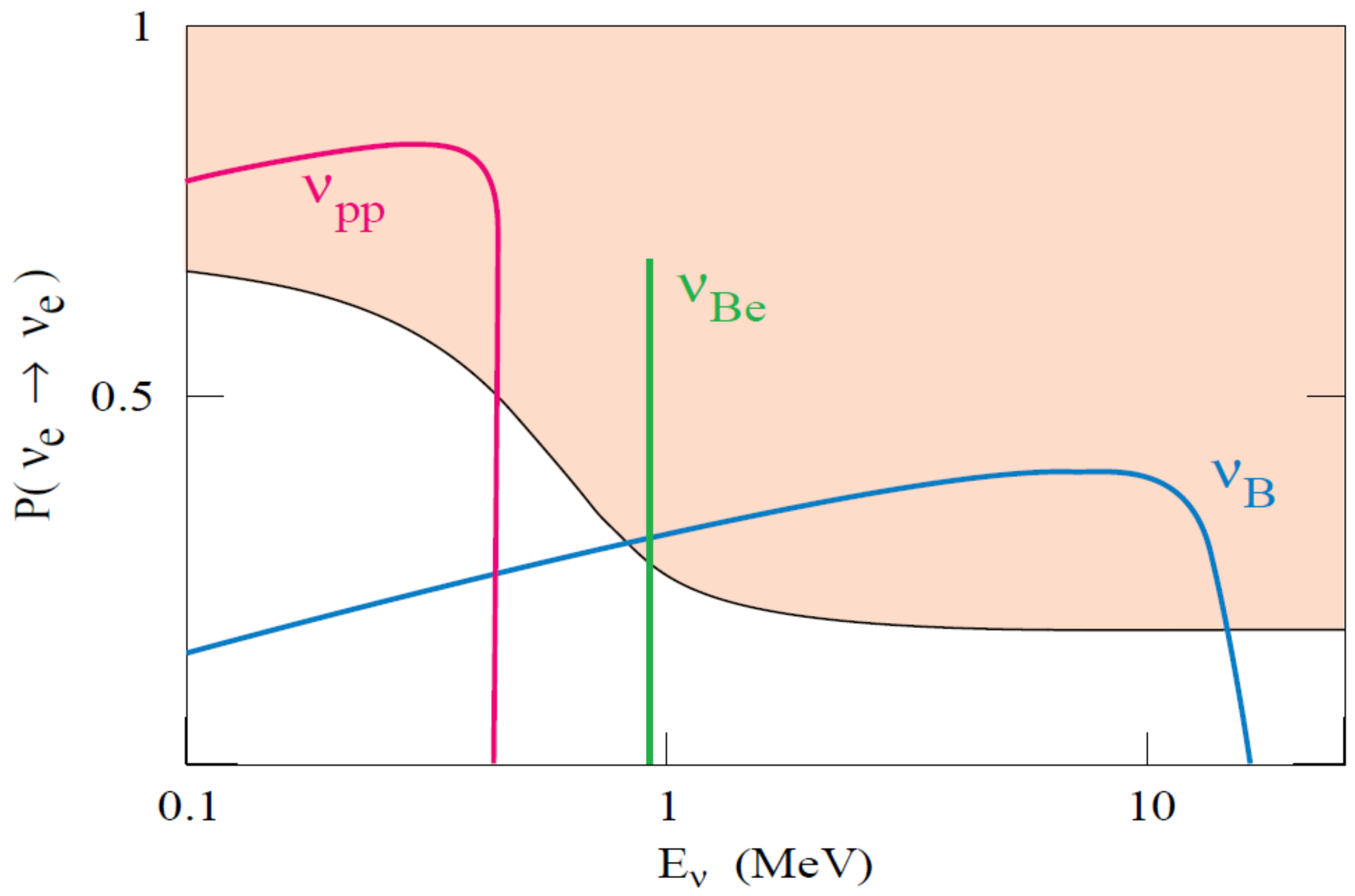
$$
\begin{aligned}
\theta_{12} &= 33.45^{+0.77}_{-0.75} \\
\theta_{23} &= 42.1^{+1.1}_{-0.9} \\
\theta_{13} &= 8.62^{+0.12}_{-0.12} \\
\Delta m^2_{\text{sol}} = \Delta m^2_{12} &= 7.42^{+0.21}_{-0.20} \times 10^{-5}\ \text{eV}^2 \\
|\Delta m^2_{\text{atm}}| = |\Delta m^2_{3\ell}| &= 2.510^{+0.027}_{-0.027} \times 10^{-3}\ \text{eV}^2
\end{aligned}
$$

- Oscillations in vacuum are not sensitive to the sign of Δm^2
- Matter effects helps to determine Δm^2 sign:
  - $m_2 > m_1$ from solar ν_e
  - Not yet resolved for m_3

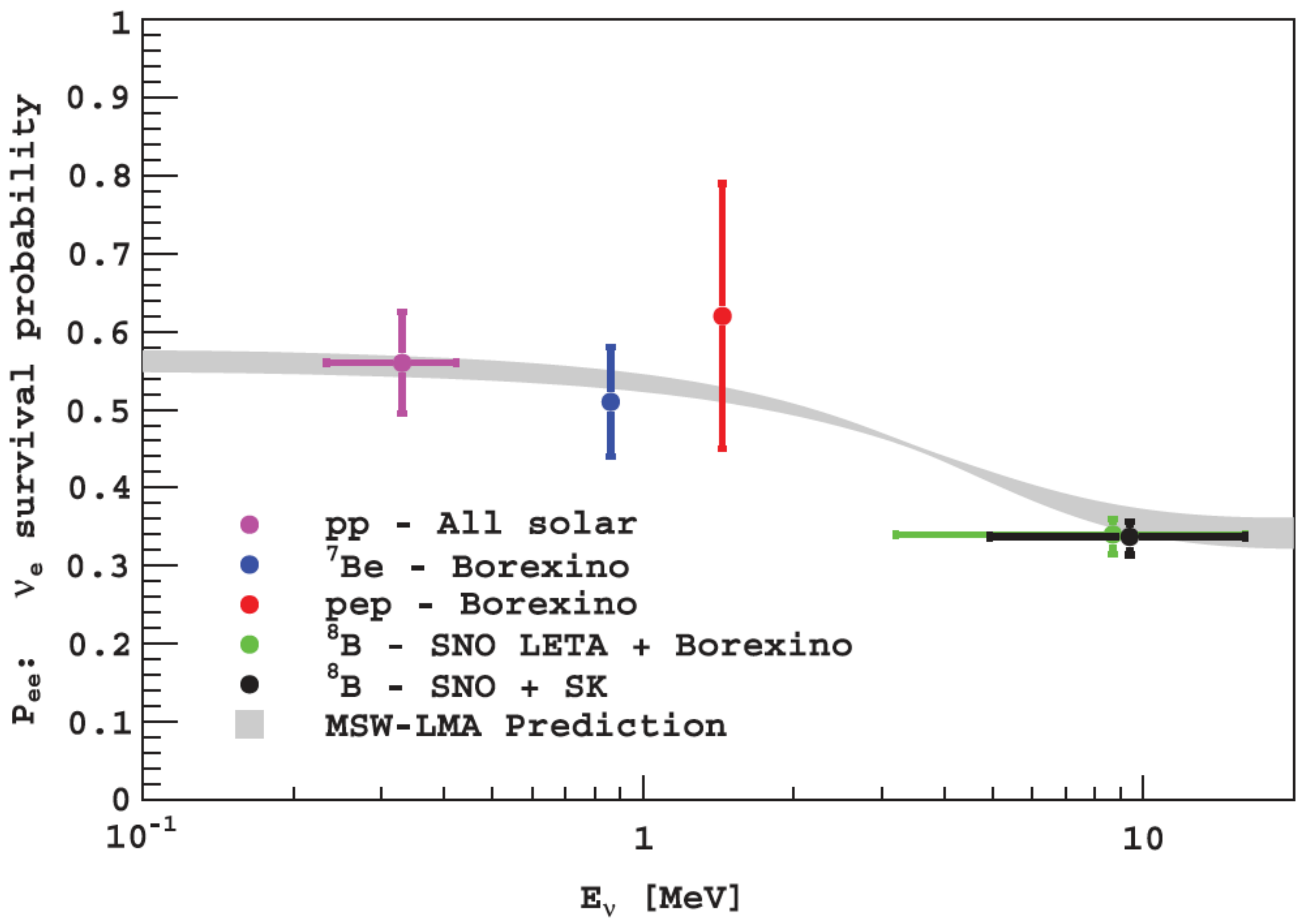
ν_e survival in vacuum



ν_e survival in matter

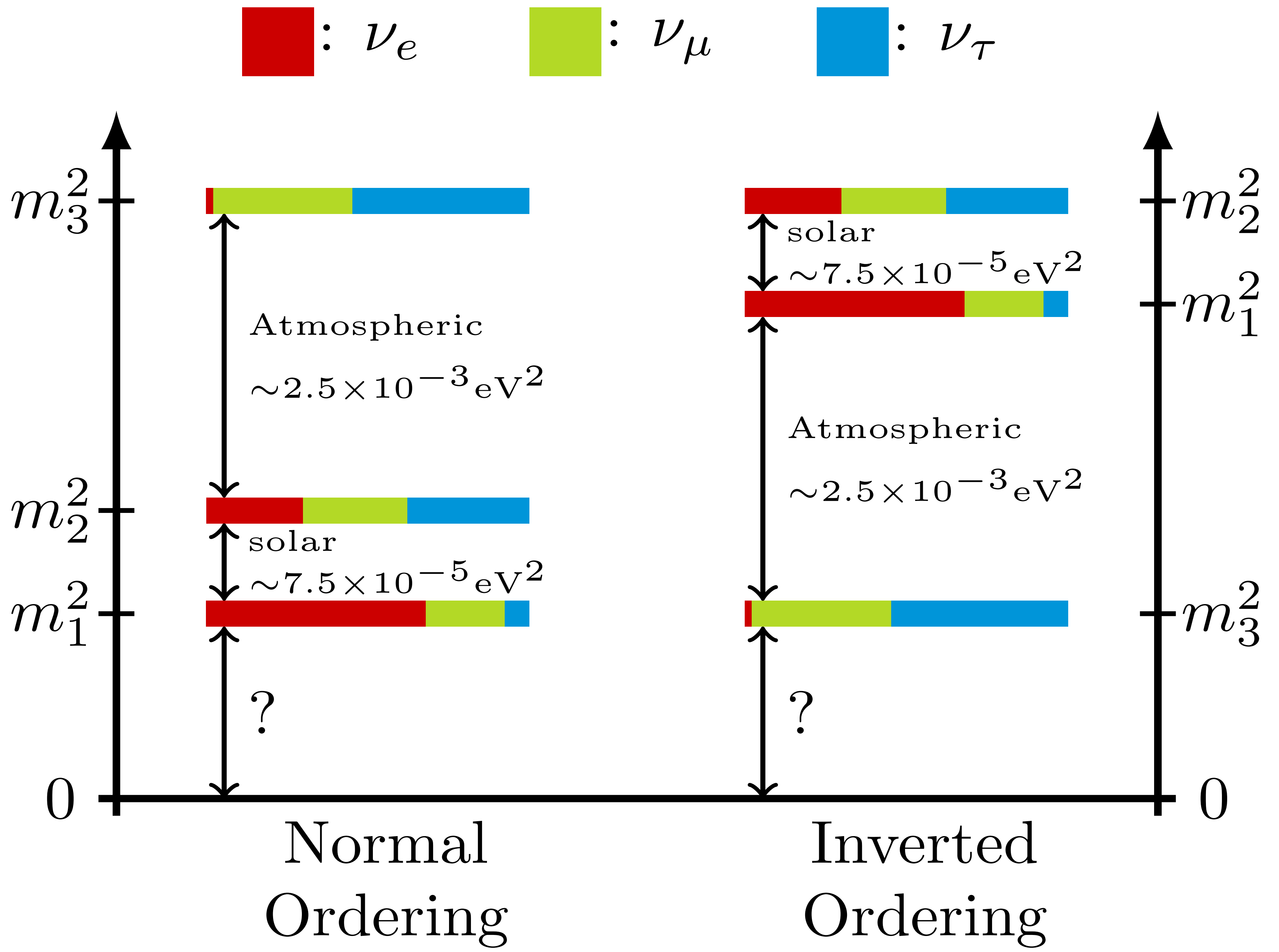


Solar ν_e flux measurement



Oscillations parameters

	Normal Ordering	Inverted Ordering
$\theta_{12} =$	$33.45^{+0.77}_{-0.75}$	$33.45^{+0.78}_{-0.75}$
$\theta_{23} =$	$42.1^{+1.1}_{-0.9}$	$49.0^{+0.9}_{-1.3}$
$\theta_{13} =$	$8.62^{+0.12}_{-0.12}$	$8.61^{+0.14}_{-0.12}$
$\Delta m^2_{\text{sol}} = \Delta m^2_{12} =$	$7.42^{+0.21}_{-0.20} \times 10^{-5}\ \text{eV}^2$	$7.42^{+0.21}_{-0.20} \times 10^{-5}\ \text{eV}^2$
$\Delta m^2_{\text{atm}} = \Delta m^2_{3\ell} =$	$+2.510^{+0.027}_{-0.027} \times 10^{-3}\ \text{eV}^2$	$-2.490^{+0.026}_{-0.028} \times 10^{-3}\ \text{eV}^2$

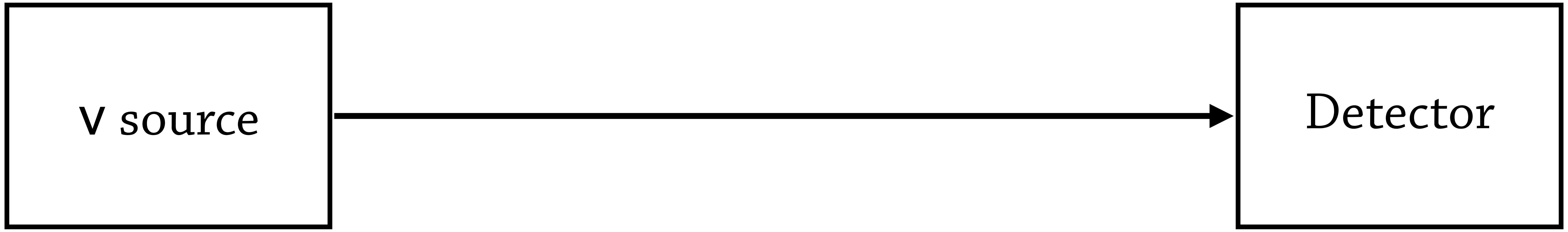


Three unknowns of neutrino oscillations :

- **Mass Hierarchy** : Normal or inverted ?
- **θ_{23} octant** : $\theta_{23} < 45°$ or $\theta_{23} > 45°$?
- **δ_{CP}** : Do ν behaves as $\bar{\nu}$?

CURRENT AND FUTURE EXPERIMENTS

Principle for precision measurement



ν source

Detector

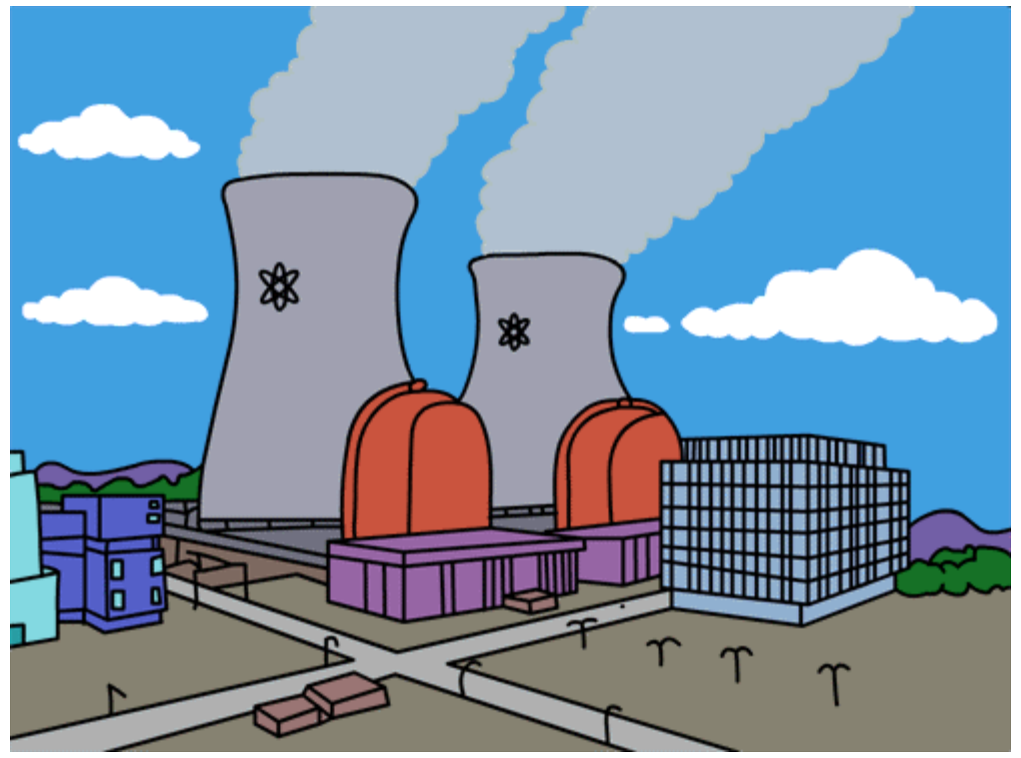
Requirements :

- Powerful source
- Initial location known
- Initial flavor content known
- Initial energy spectrum known

Requirements :

- At L/E for oscillation
- Able to distinguish e/μ/τ
- Energy reconstruction
- Big and/or dense

Experiments Using Reactors



Continuous powerful emission of $\bar{\nu}_e$ through the fission of ^{235}U, ^{239}Pu and ^{239}Pu.

Energy spectrum from 2 to 8 MeV

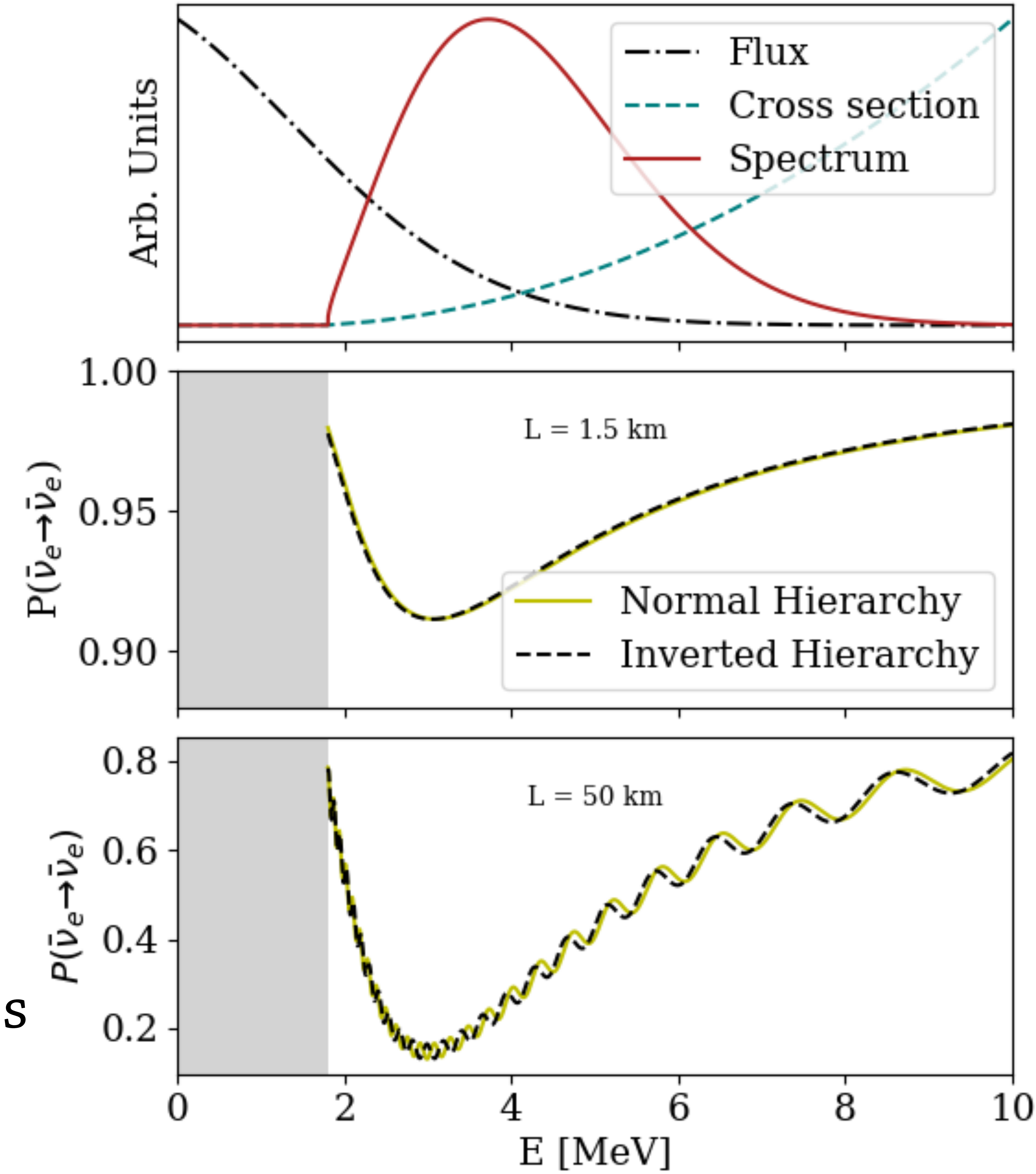
- Cannot tag new flavor appearance ($\bar{\nu}_\mu$ or $\bar{\nu}_\tau$)
- Only the disappearance measurement is possible

L ~ 1 km to be at atmospheric oscillation

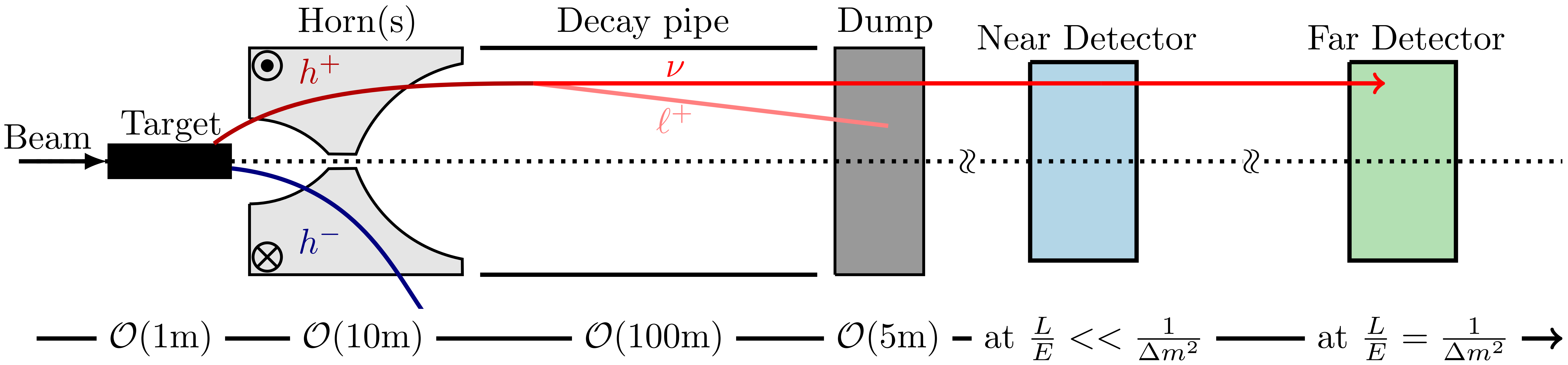
- Considered to be in vacuum, the 2 flavor approximation is valid. No sensitivity to δ_{CP} or MH

L ~ 50 km to be at solar oscillation

- Study of the interference between the 2 oscillations gives sensitivity to MH [***JUNO*** experiment]

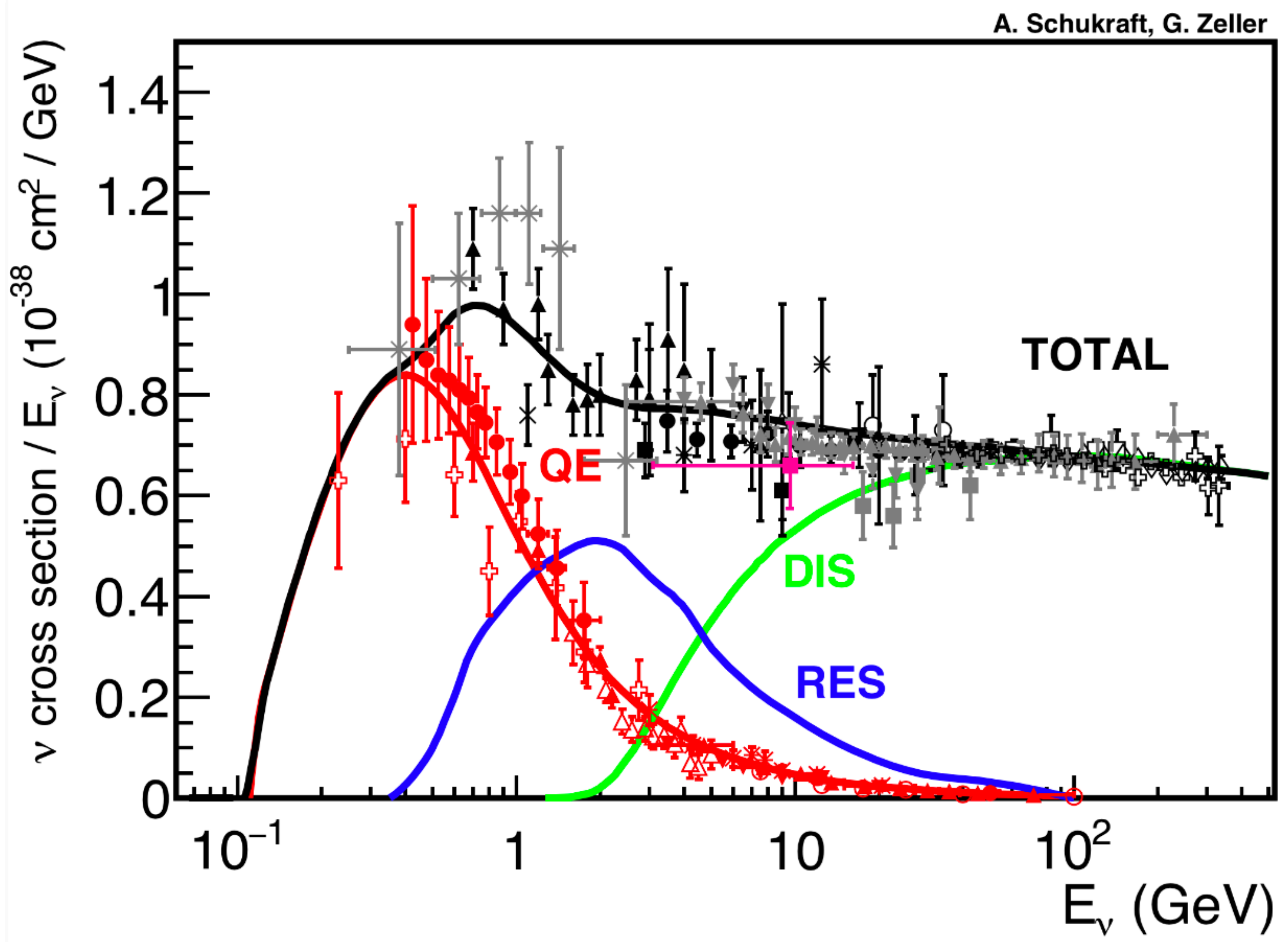


Experiments Using Accelerators



Principle :

- Accelerated proton collides into a target, produces mostly $\pi^\pm$.
- Pions main decay channel (99%) : $\pi \rightarrow \mu + \nu_\mu$
- Focussing horns to select π^+ (ν_μ flux) or π^- ($\bar{\nu}_\mu$ flux)
- A near detector to measure the flux *before* oscillations
- A far detector at the L/E to observe oscillations
  - ν beamline parameters tuned for optimal E



Experiments Using Accelerators

Three Channels possible (same for $\bar{\nu}_\mu$):

○ **$\nu_\mu \rightarrow \nu_\mu$:**

- No CP violation : $P(\nu_\mu \rightarrow \nu_\mu) \equiv P(\bar{\nu}_\mu \rightarrow \bar{\nu}_\mu)$
- Negligible matter effects

○ **$\nu_\mu \rightarrow \nu_e$:** The Golden Channel

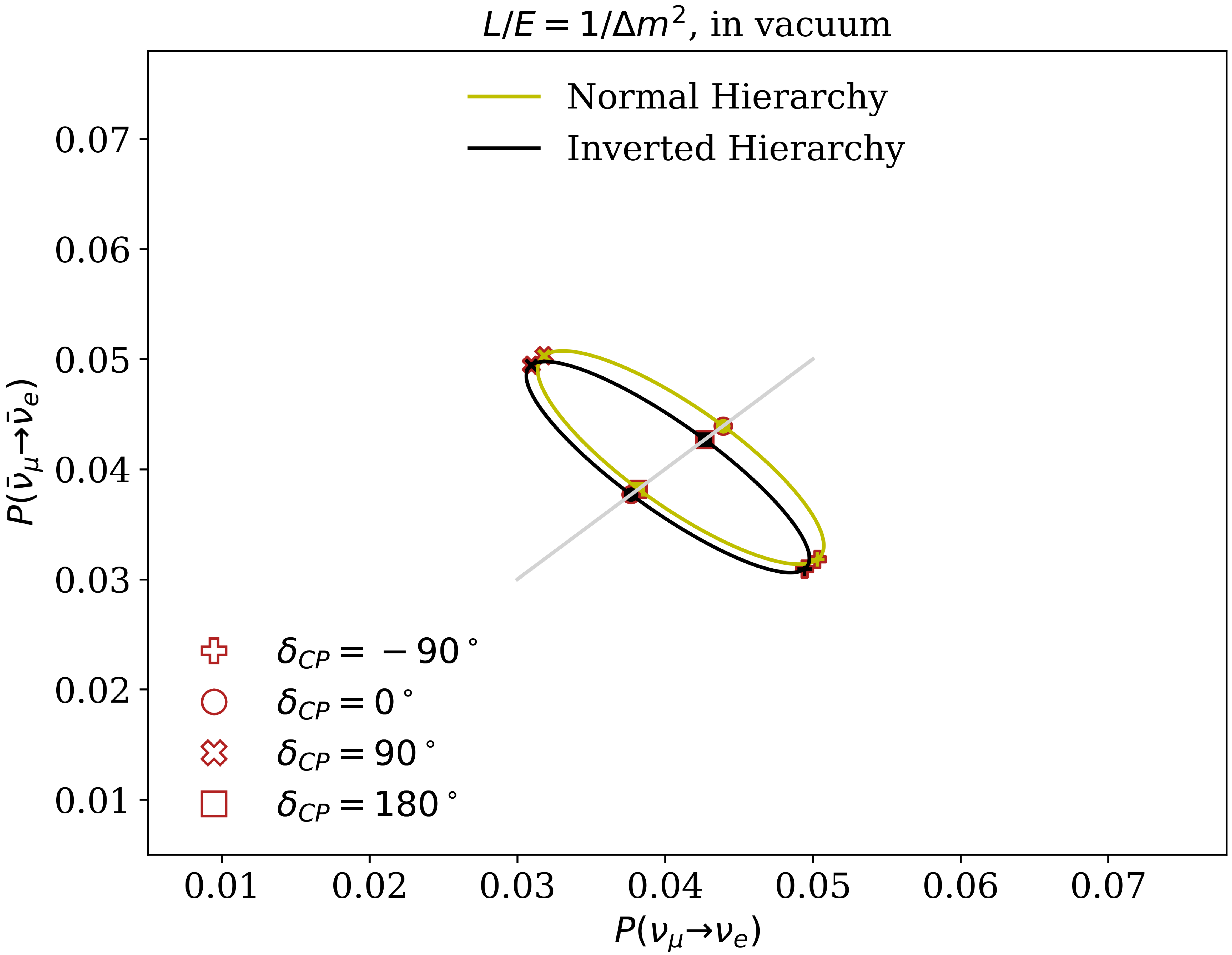
- Very sensitive to CP
- Very sensitive to MH with matter
- Very sensitive to θ_{23} octant

○ **$\nu_\mu \rightarrow \nu_\tau$:**

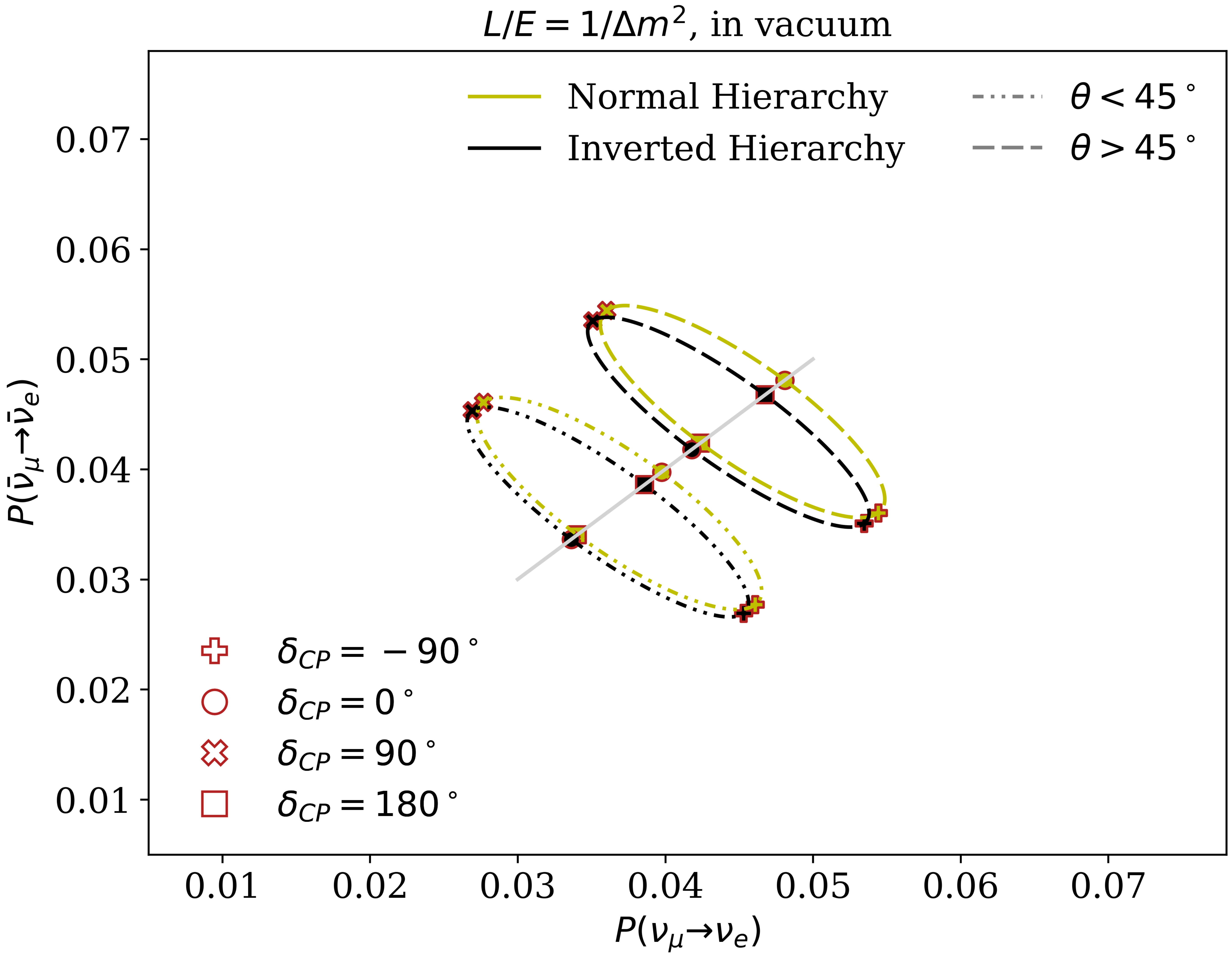
- Similar discovery potential as ν_e appearance but:

$m_\tau = 1.7$ GeV, $c\tau_\tau = 87$ μm and $\tau^\pm$ have hundreds of complicated decay channels

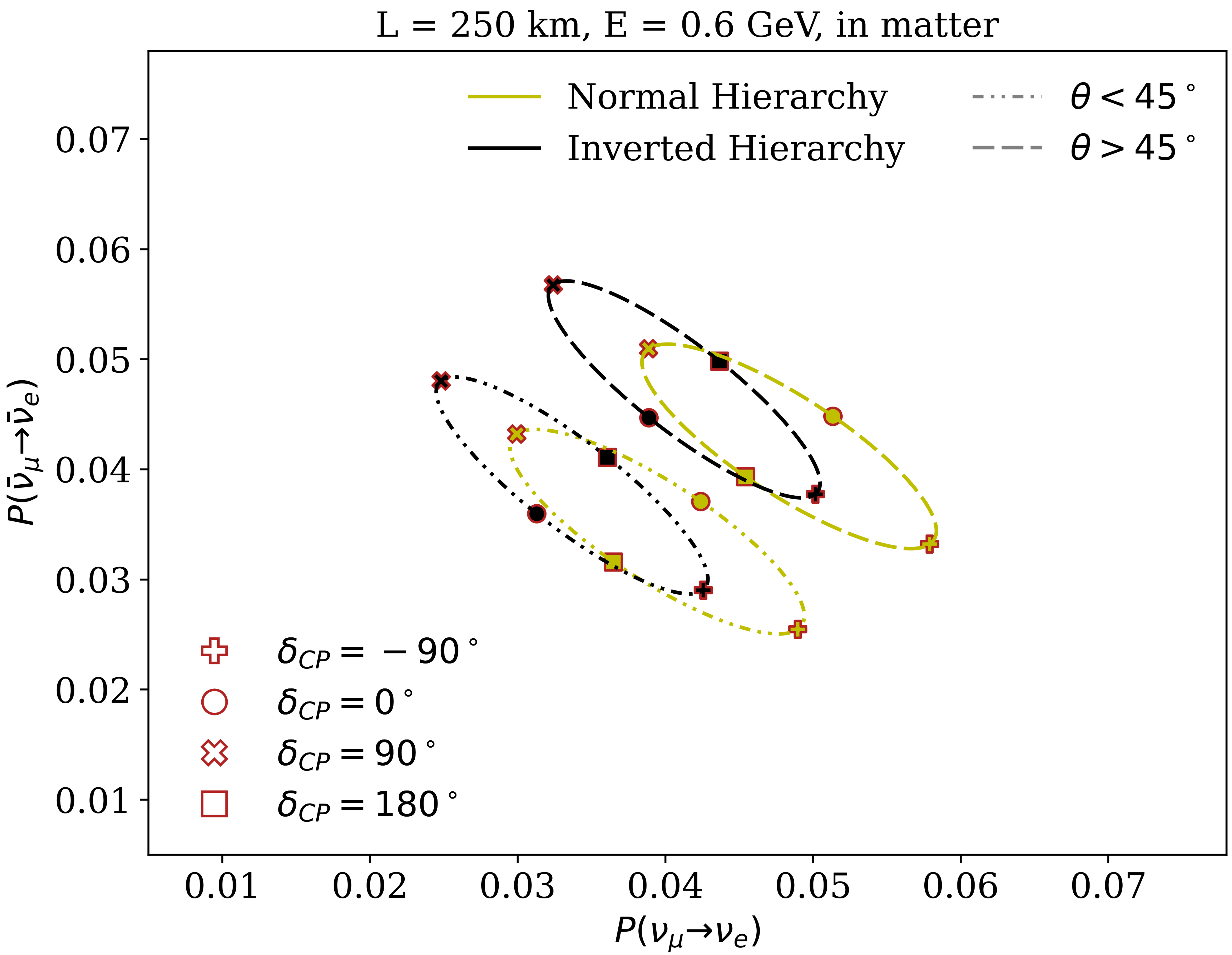
Experiments Using Accelerators



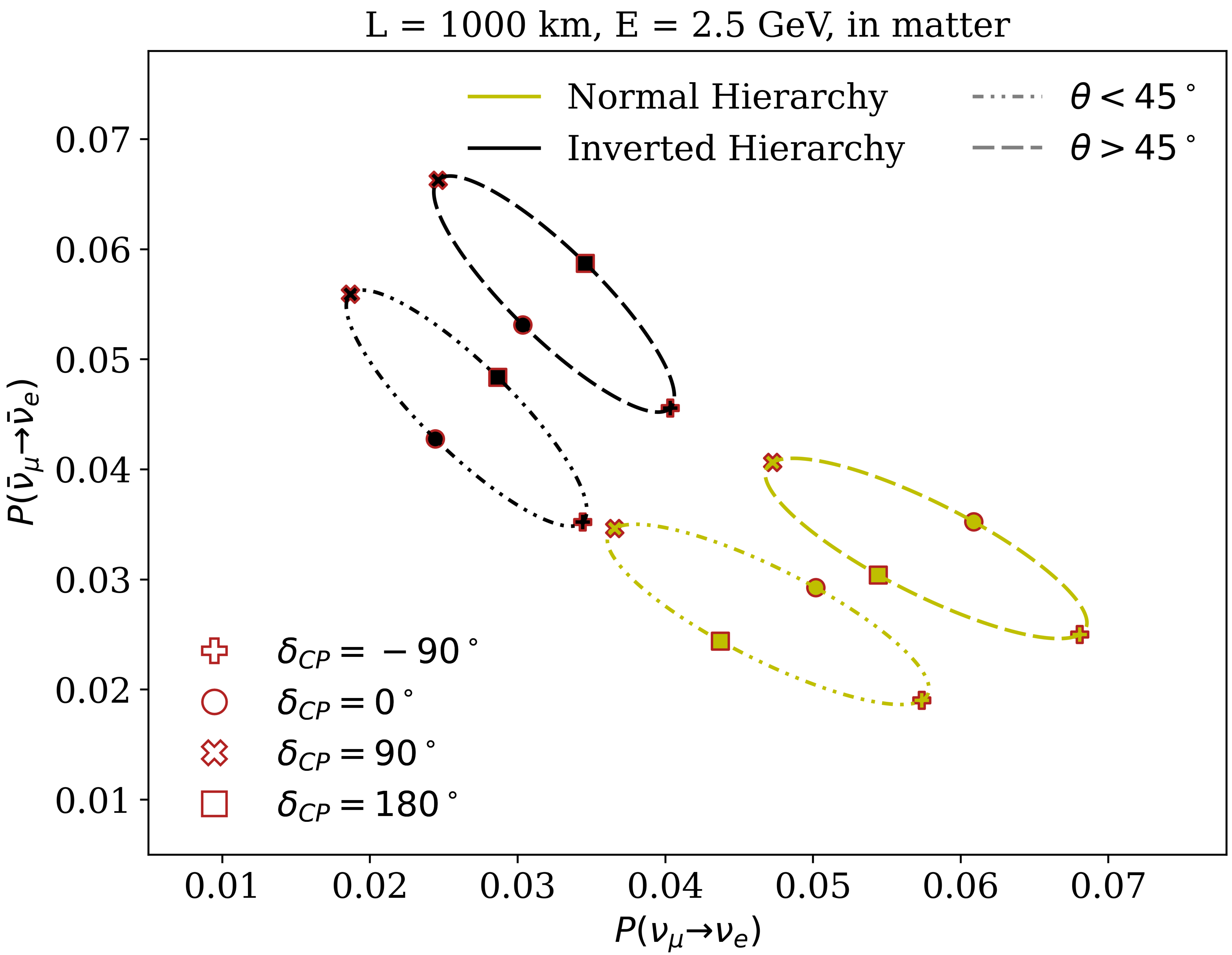
Experiments Using Accelerators



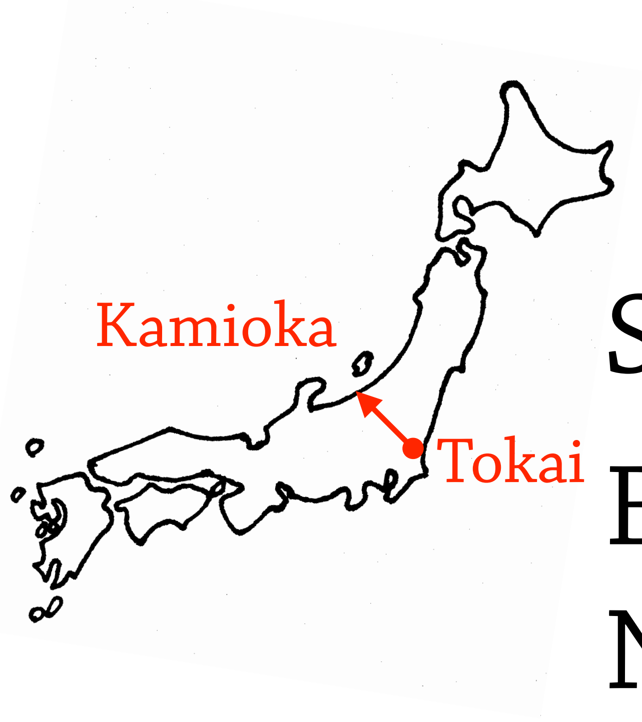
Experiments Using Accelerators



Experiments Using Accelerators



Current ν accelerator experiments



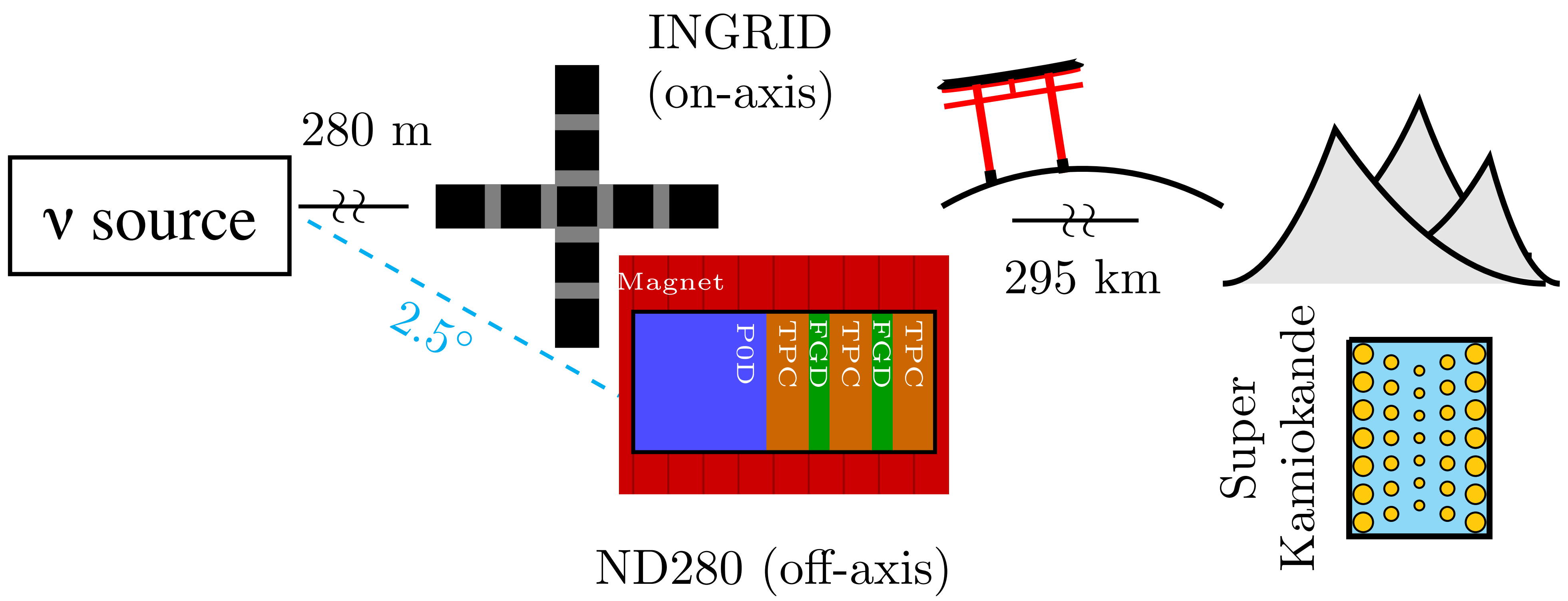
T2K in Japan

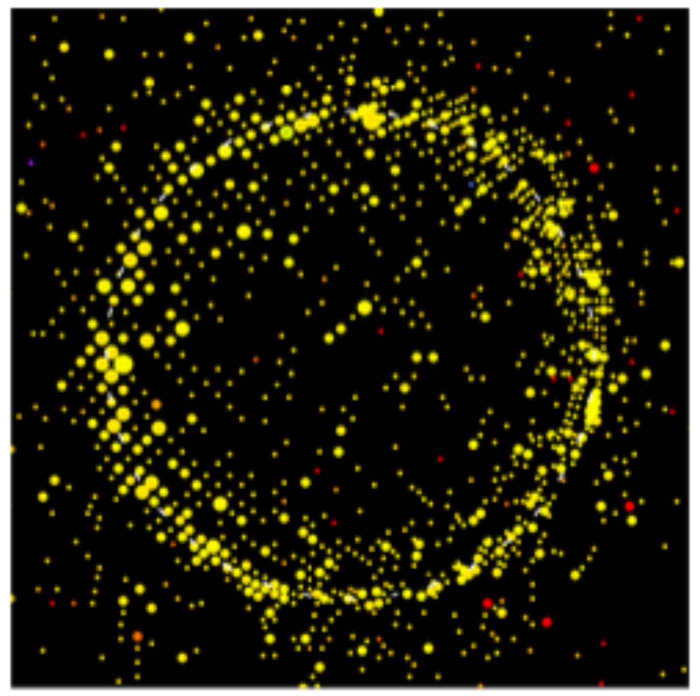
Since 2010, L = 295 km, E = 0.6 GeV

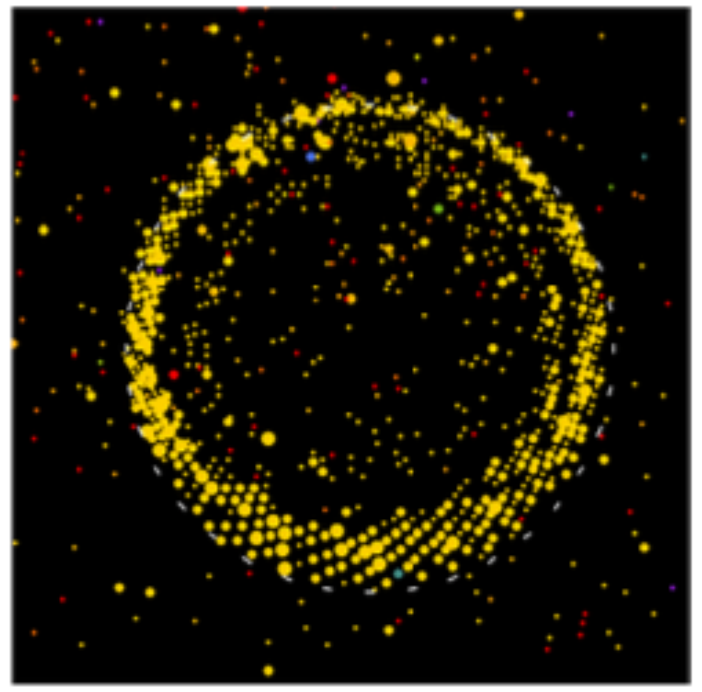
Equal ν and ν̄ runs

Near detector is a gaseous TPC

Far detector is Super-Kamiokande




ν_e-like

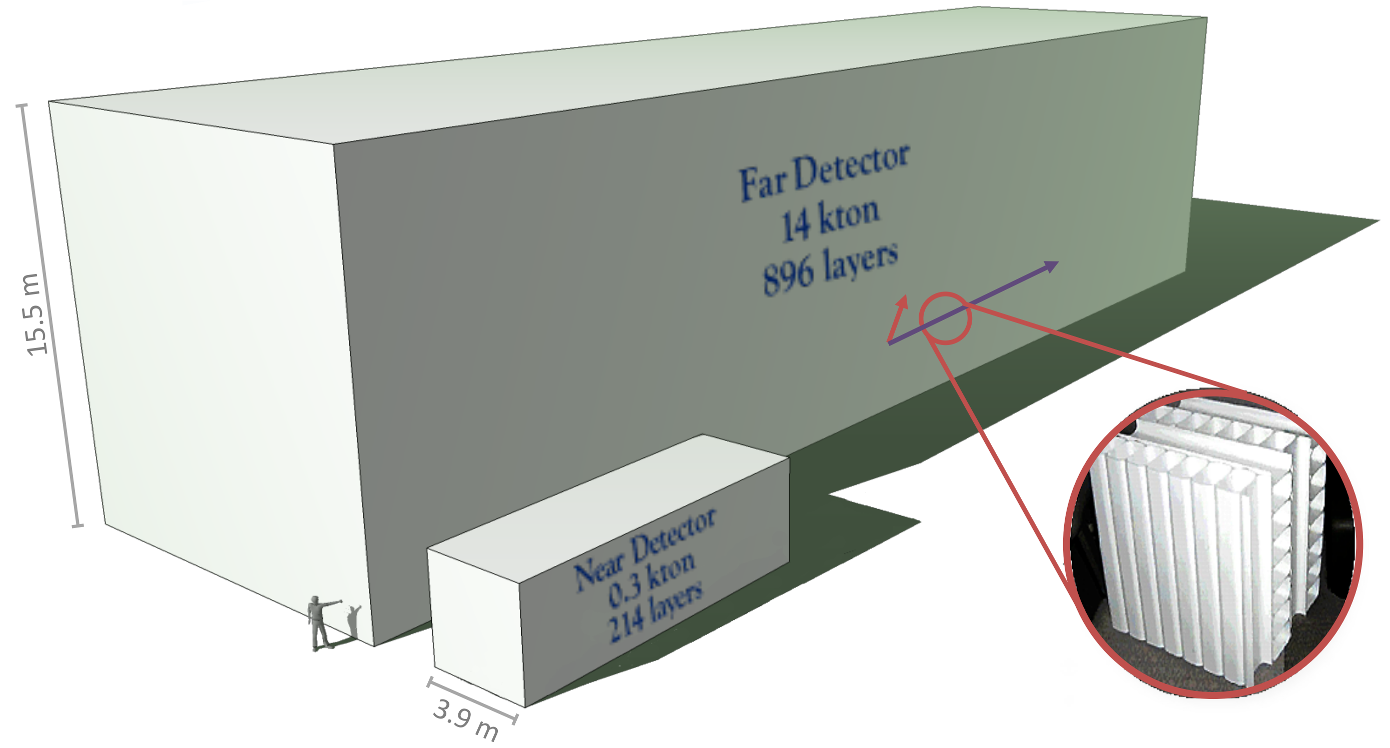

ν_μ-like

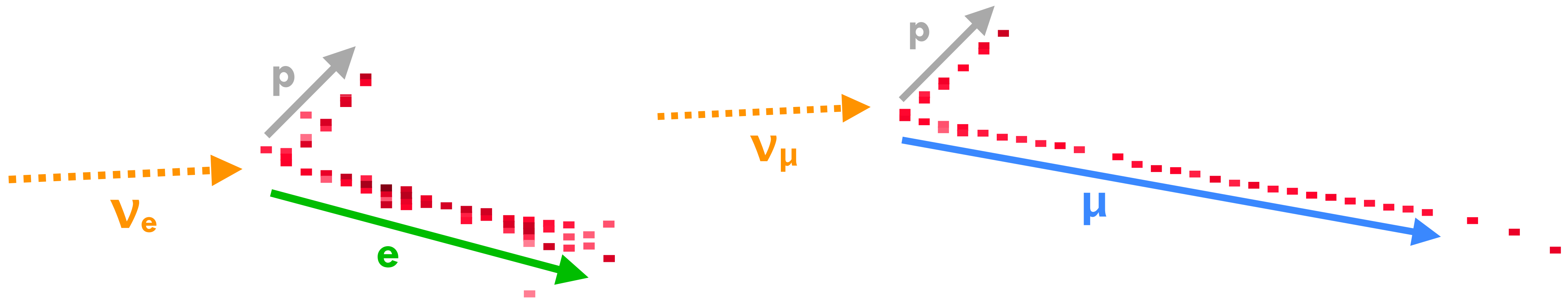
NOνA in the US

Since 2013, L = 810 km, E = 2 GeV

Equal ν and ν̄ runs

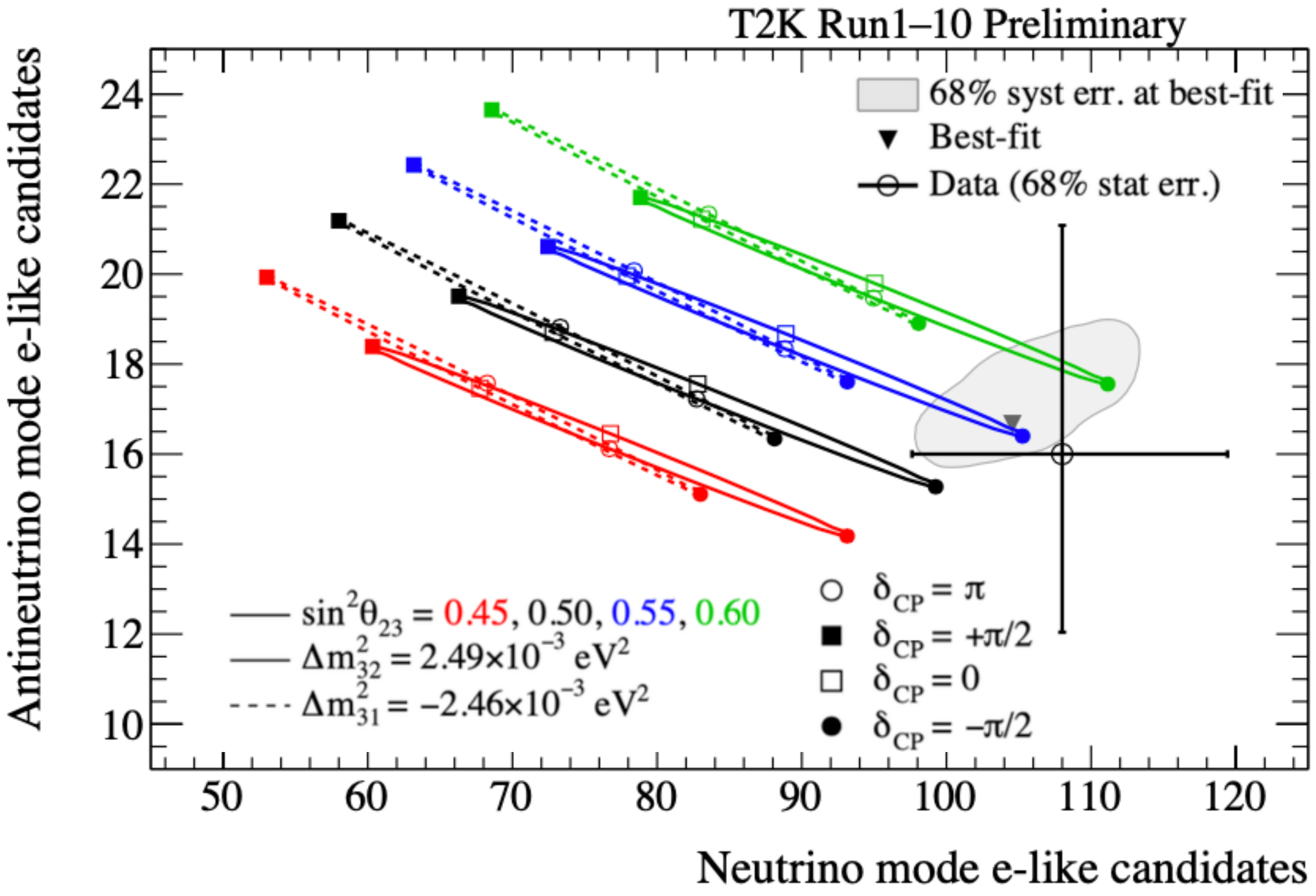
Near and Far detectors are plastic scintillators



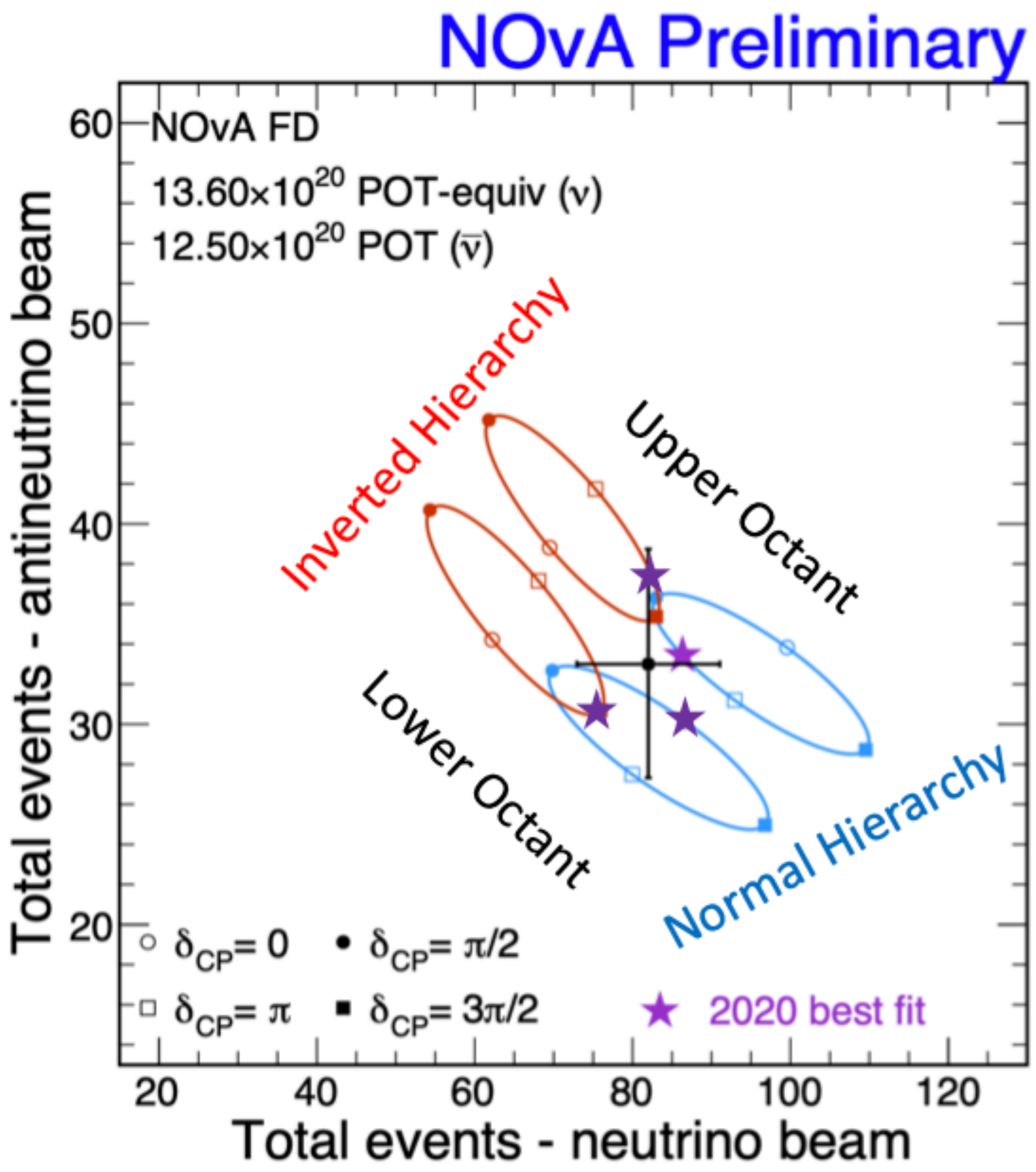


Current ν accelerator experiments

T2K in Japan

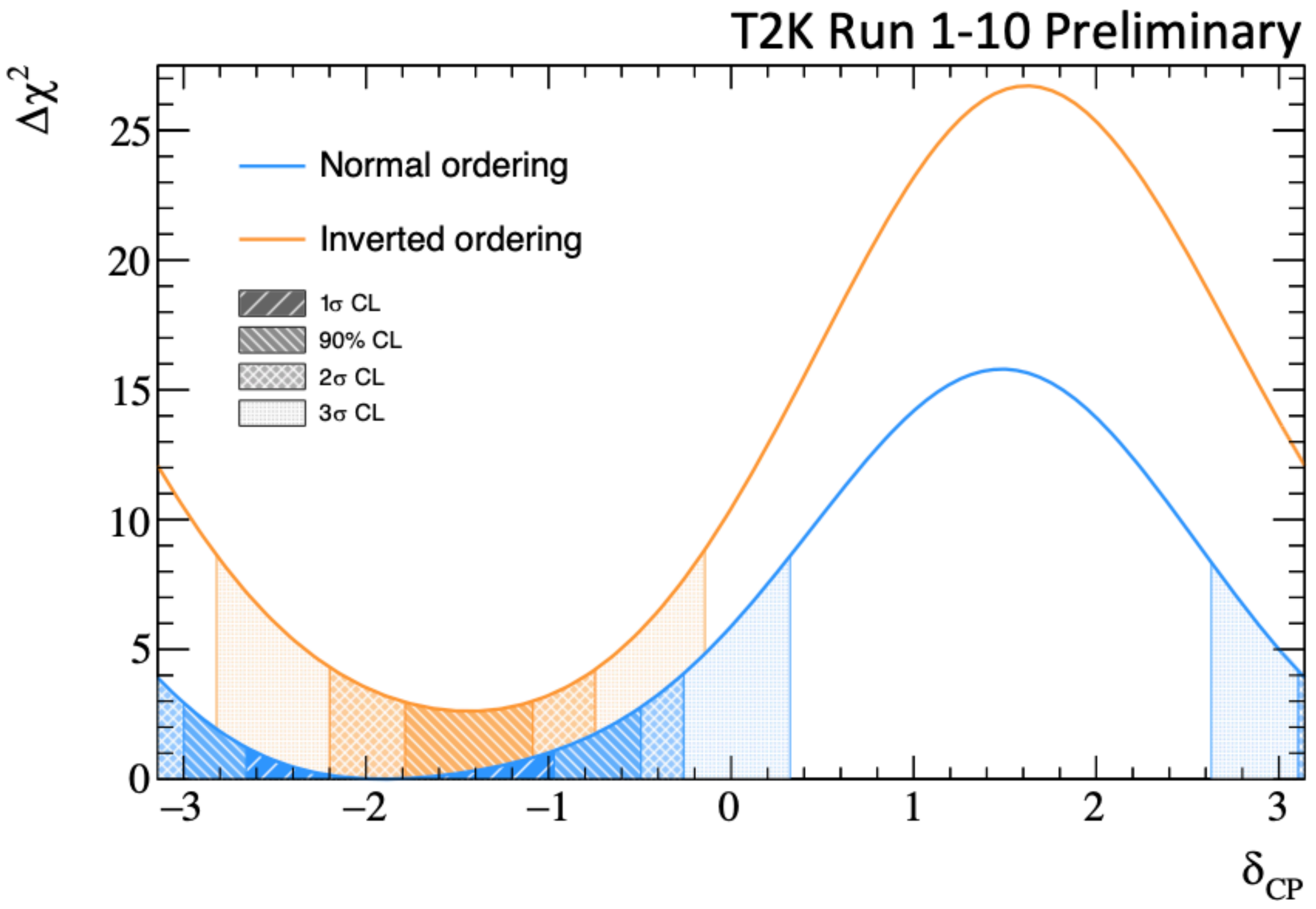


NOvA in the US



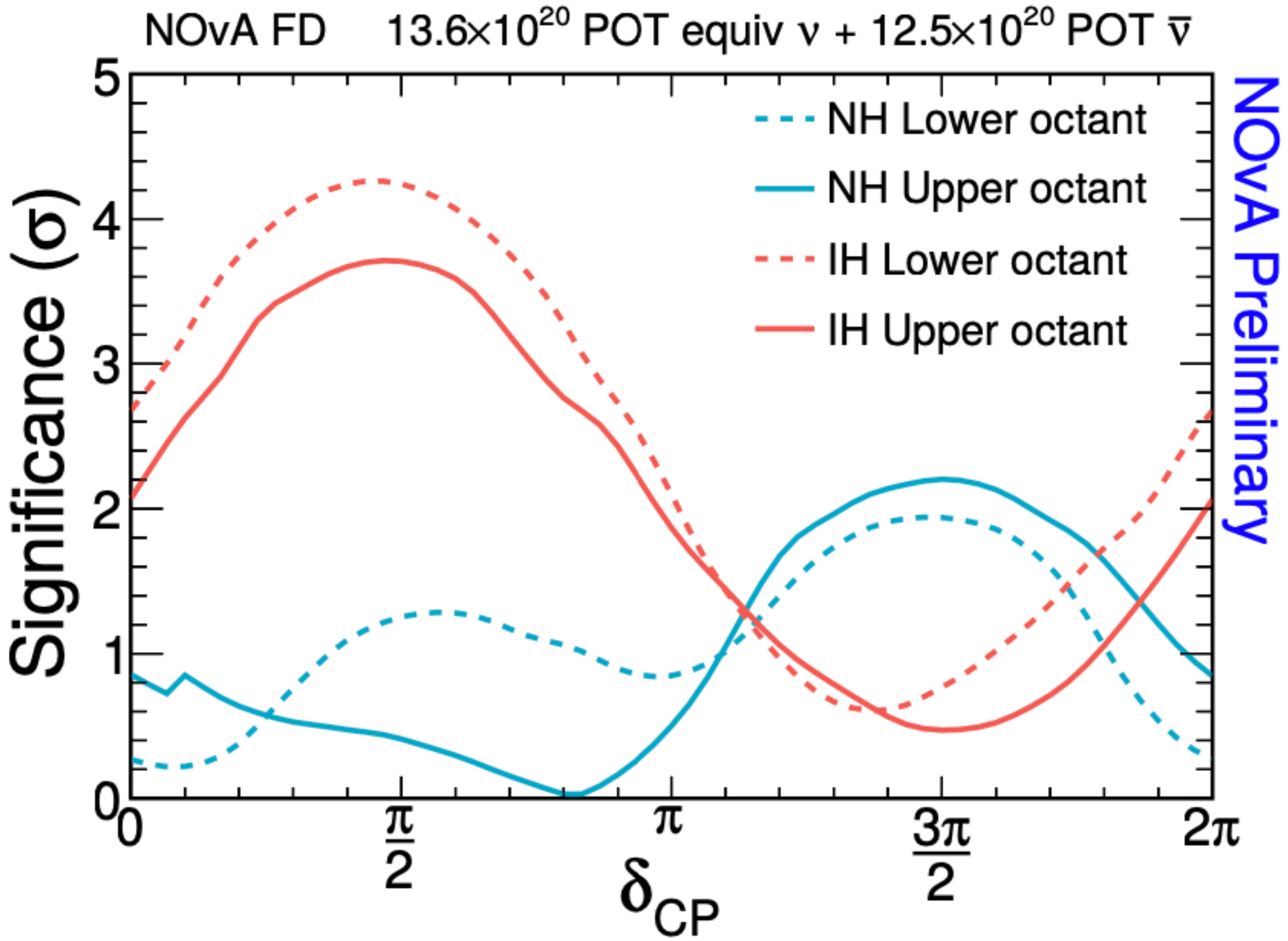
Current ν accelerator experiments

T2K in Japan



- ○Slight preference for Normal Hierarchy
- ○δ_{CP} = (0, π) excluded at 95% C.L. for both MH
- ○Large range around δ_{CP} = +π/2 excluded at 3σ

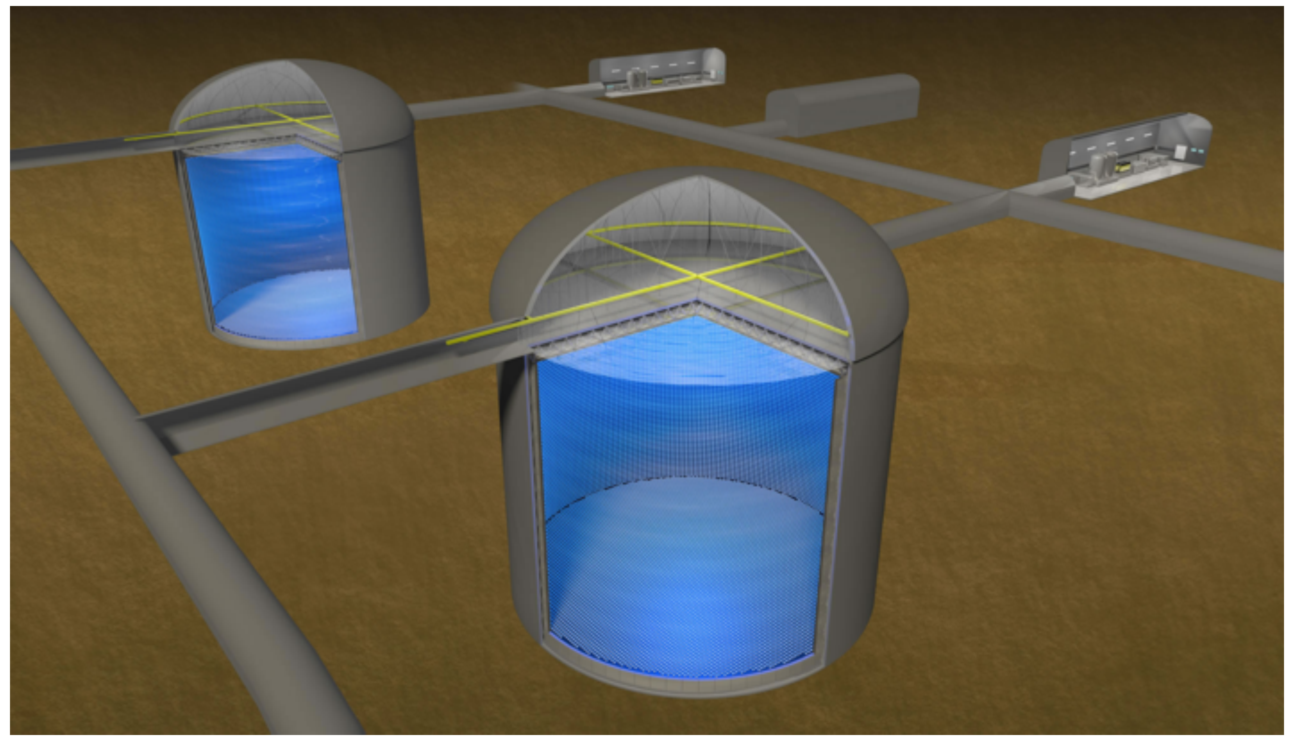
NOvA in the US



- ○Prefers Normal Hierarchy at 1.0σ
- ○Exclude δ_{CP} = π/2 + IH at >3σ
- ○Exclude δ_{CP} = 3π/2 + NH at 2σ

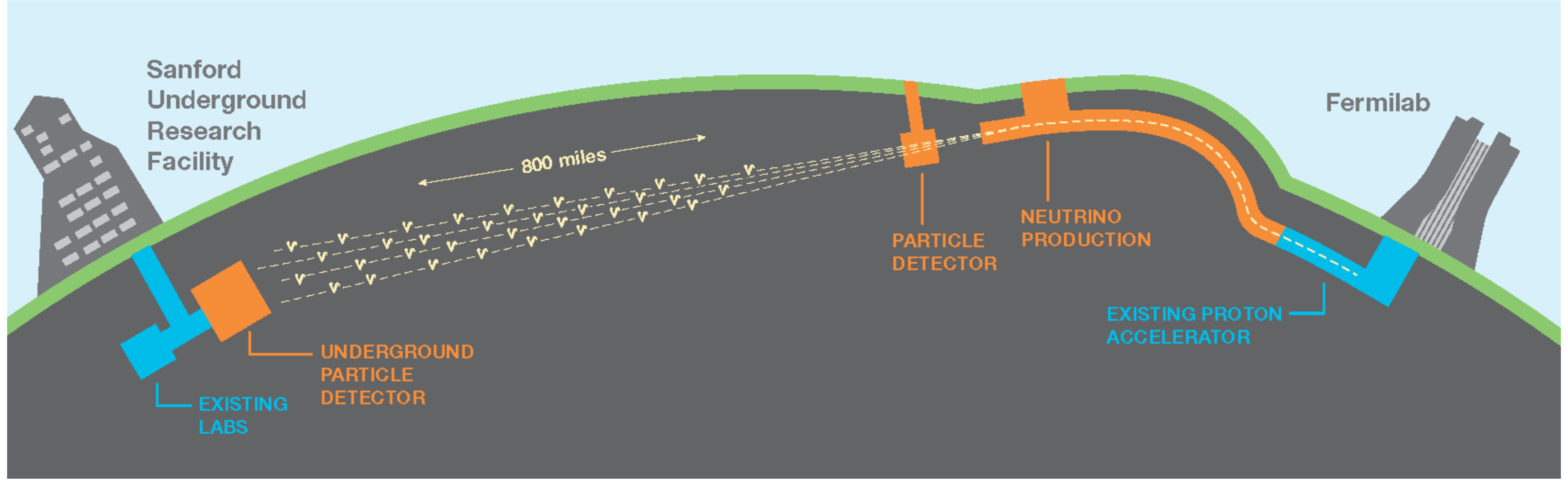
Future ν accelerator experiments

T2HK in Japan



- L=300 km, E ~ 0.6 GeV
- 260 kt water Cherenkov detector
- Proven and scalable technology
- Excellent e-μ ring separation
- Little R&D foreseen
- Only low energy beam possible (< 1 GeV)

DUNE in the US

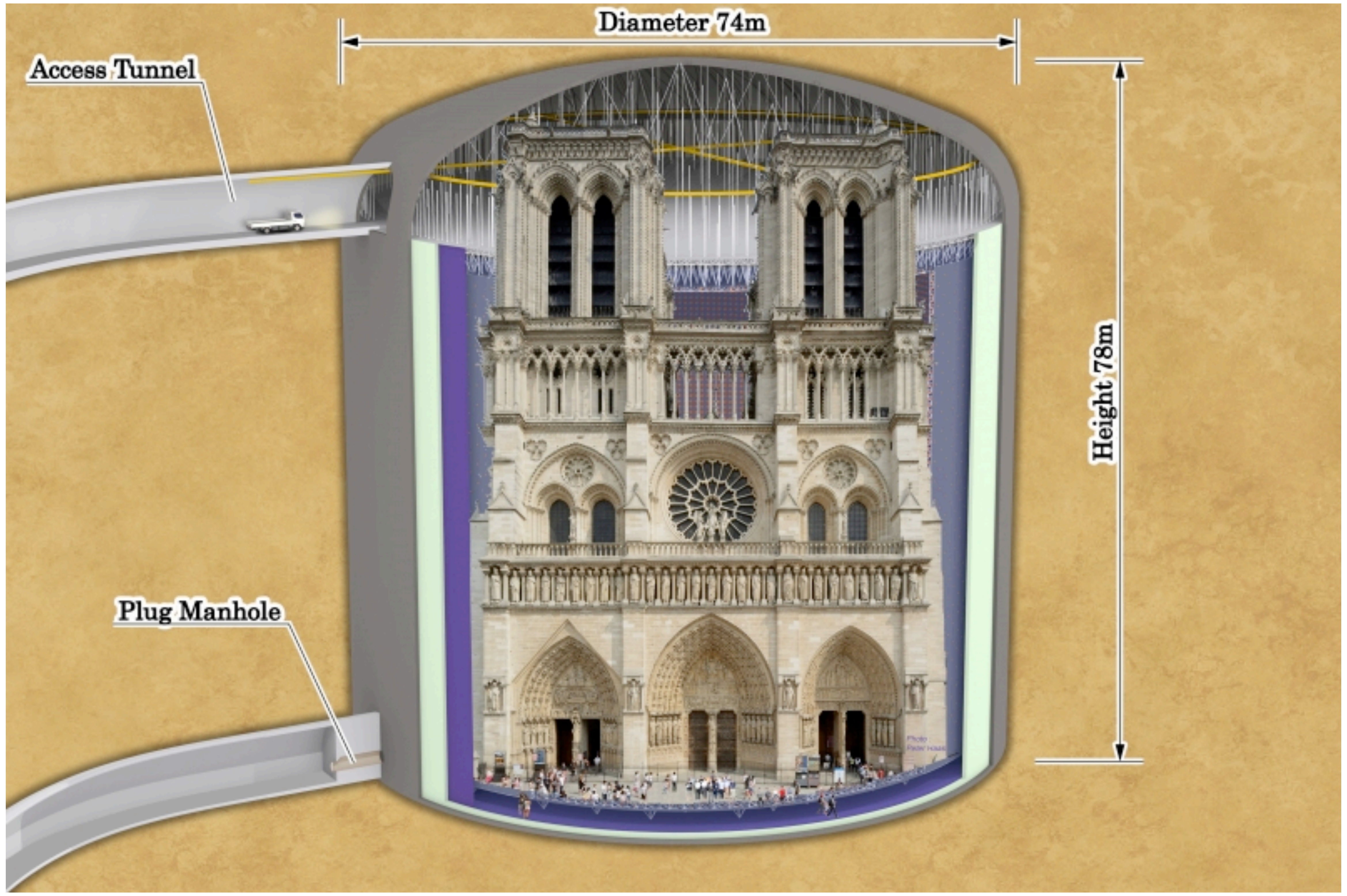


- L=1300 km, E ~ 1-3 GeV
- 40 kt liquid argon TPC detector
- 3D imaging with high granularity for precise tracking
- Low energy threshold (~10s MeV)
- Important R&D efforts ongoing : Scalability, Engineering

Both planning of starting data taking in ~2026

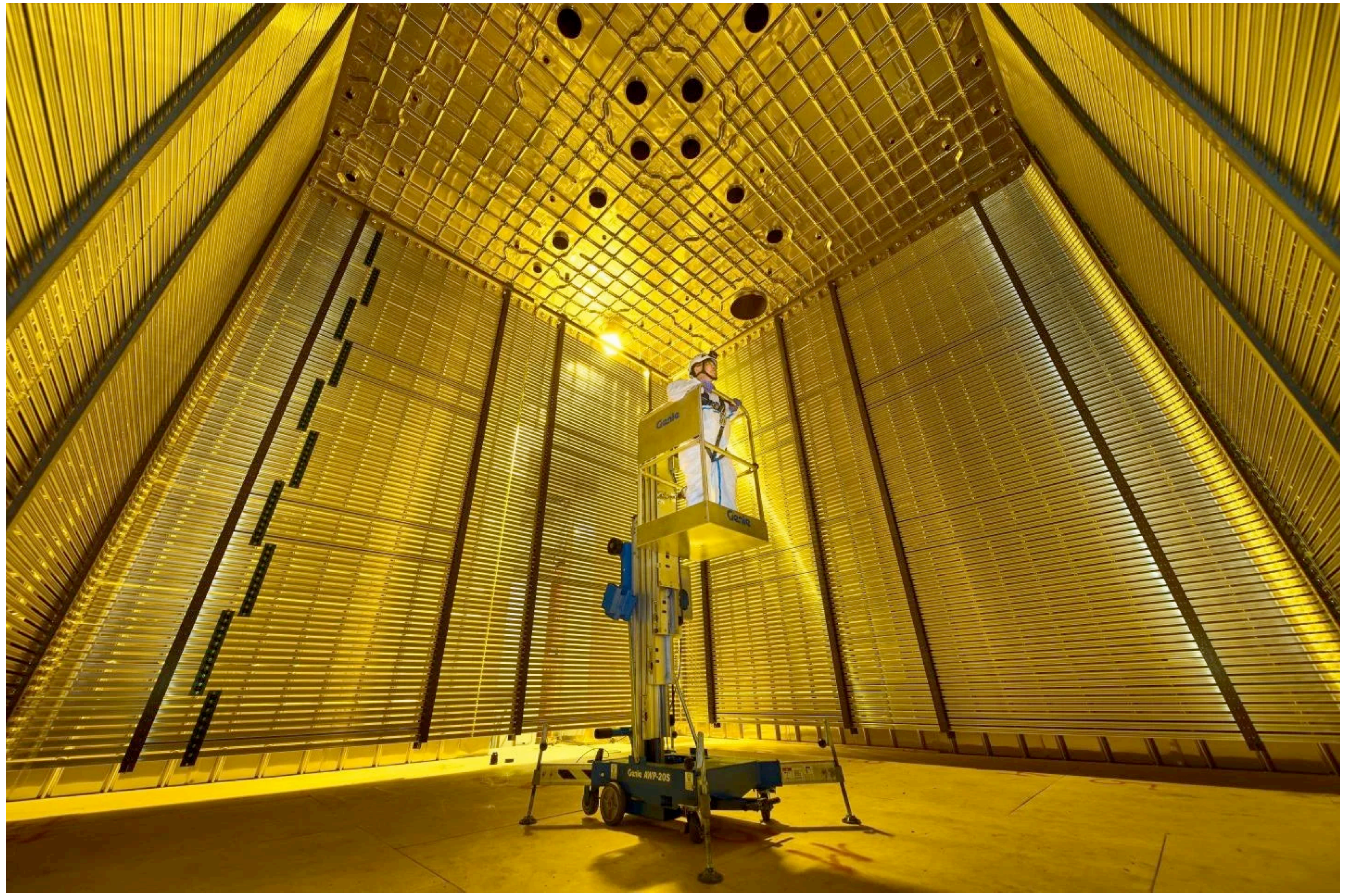
Future ν accelerator experiments

T2HK in Japan



Notre-Dame will fit inside Hyper-Kamiokande !

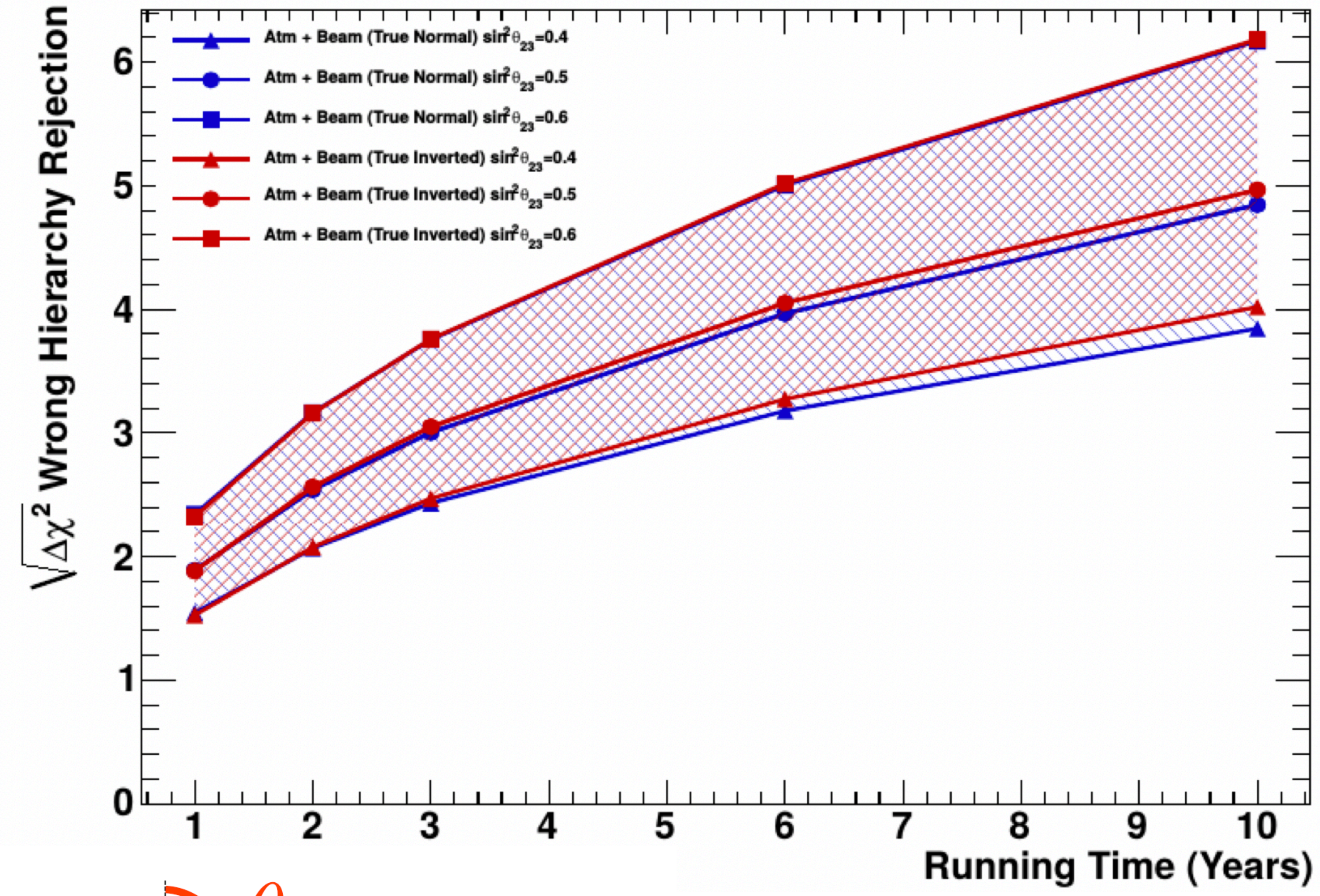
DUNE in the US



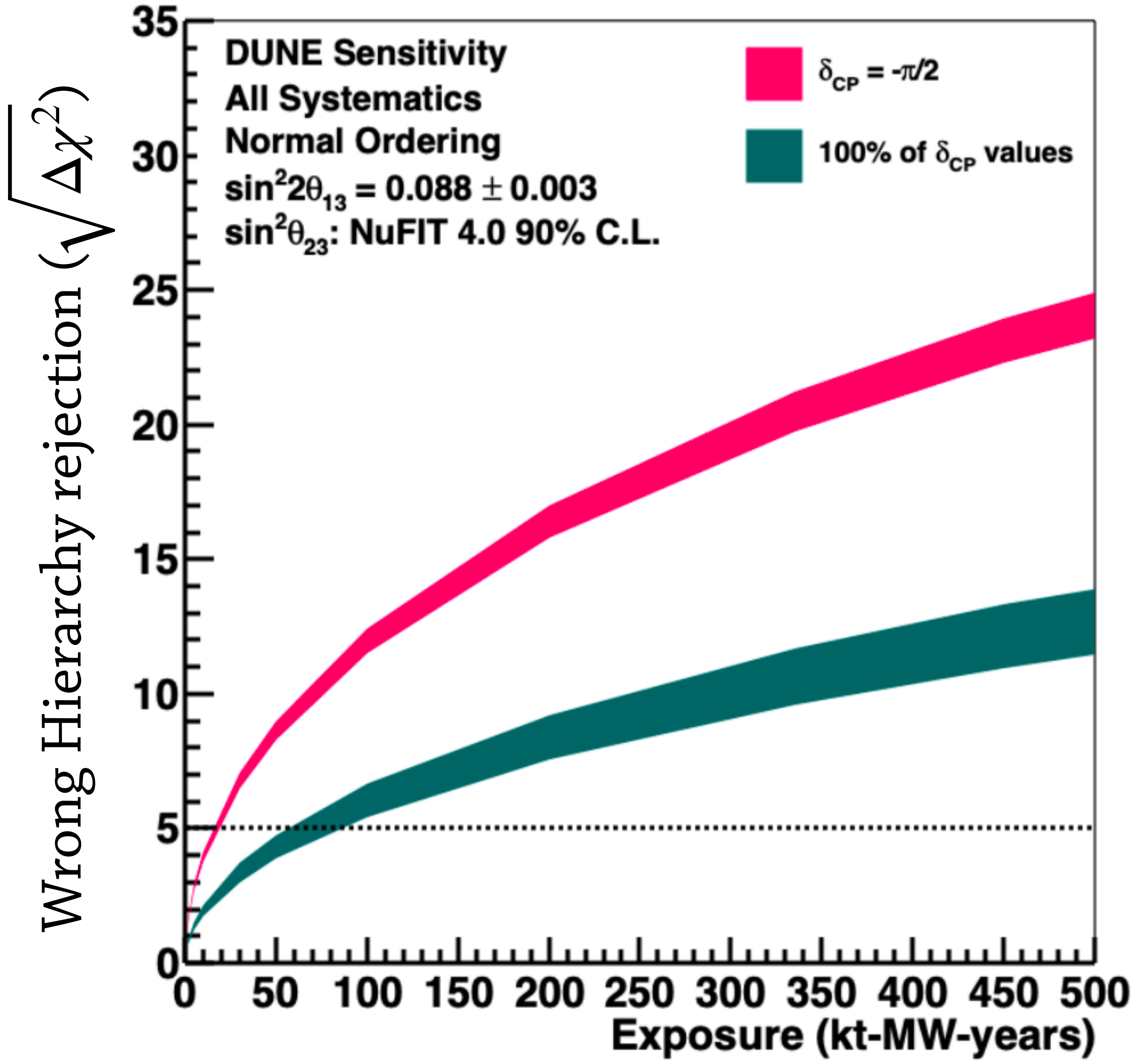
Inside DUNE prototype (6×6×6 m^3) at CERN
-> Future : 4 modules of 60×12×12 m^3 each

Future ν accelerator experiments

T2HK in Japan

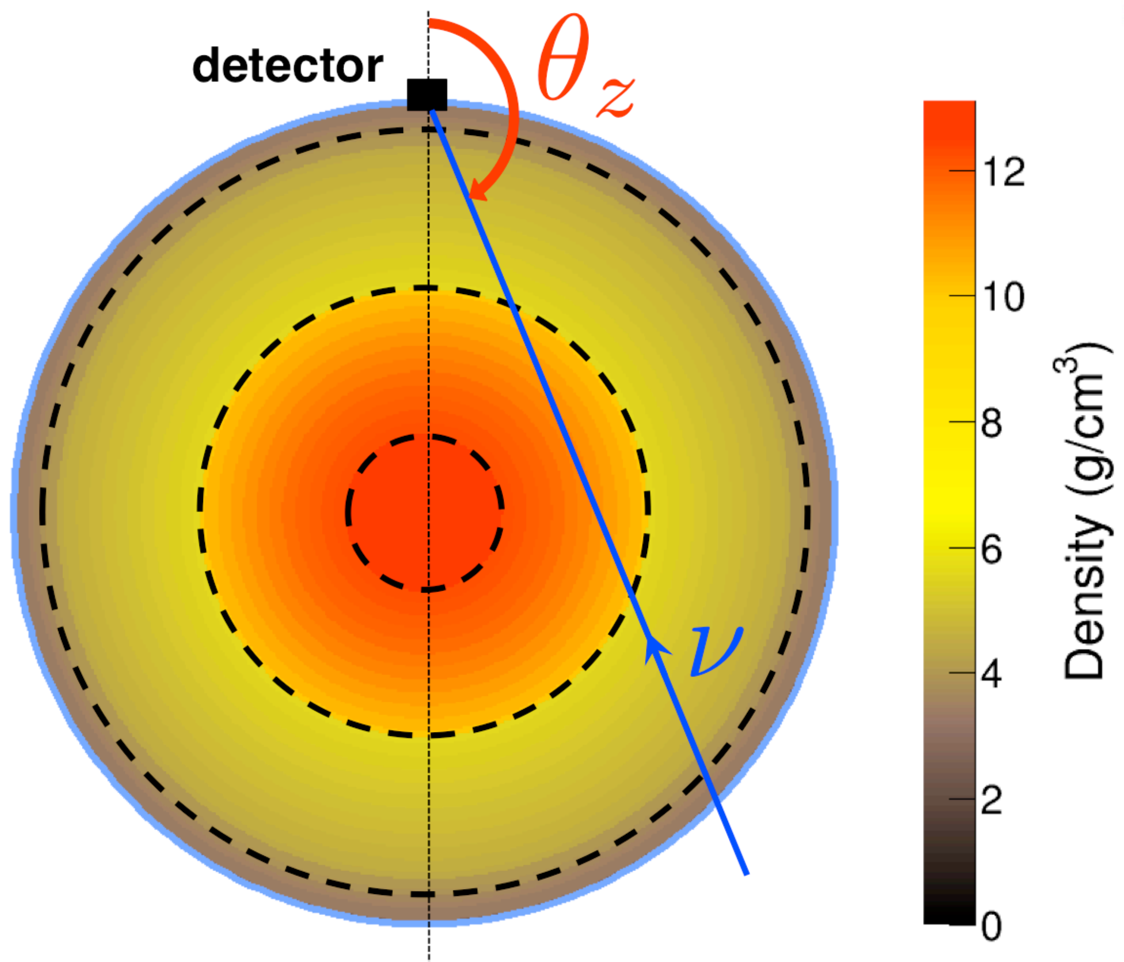


DUNE in the US



DUNE default operation :
40 kton of LAr staged
Beam power at 1.2 ~ 2.4 MW

kt•MW•yr	Staged years
30	1.2
100	3.1
200	5.2
336	7
624	10
1104	15



In the ideal case of δ_{CP} = -π/2

○ ***DUNE*** will resolve the MH at 5σ in ~1.5y

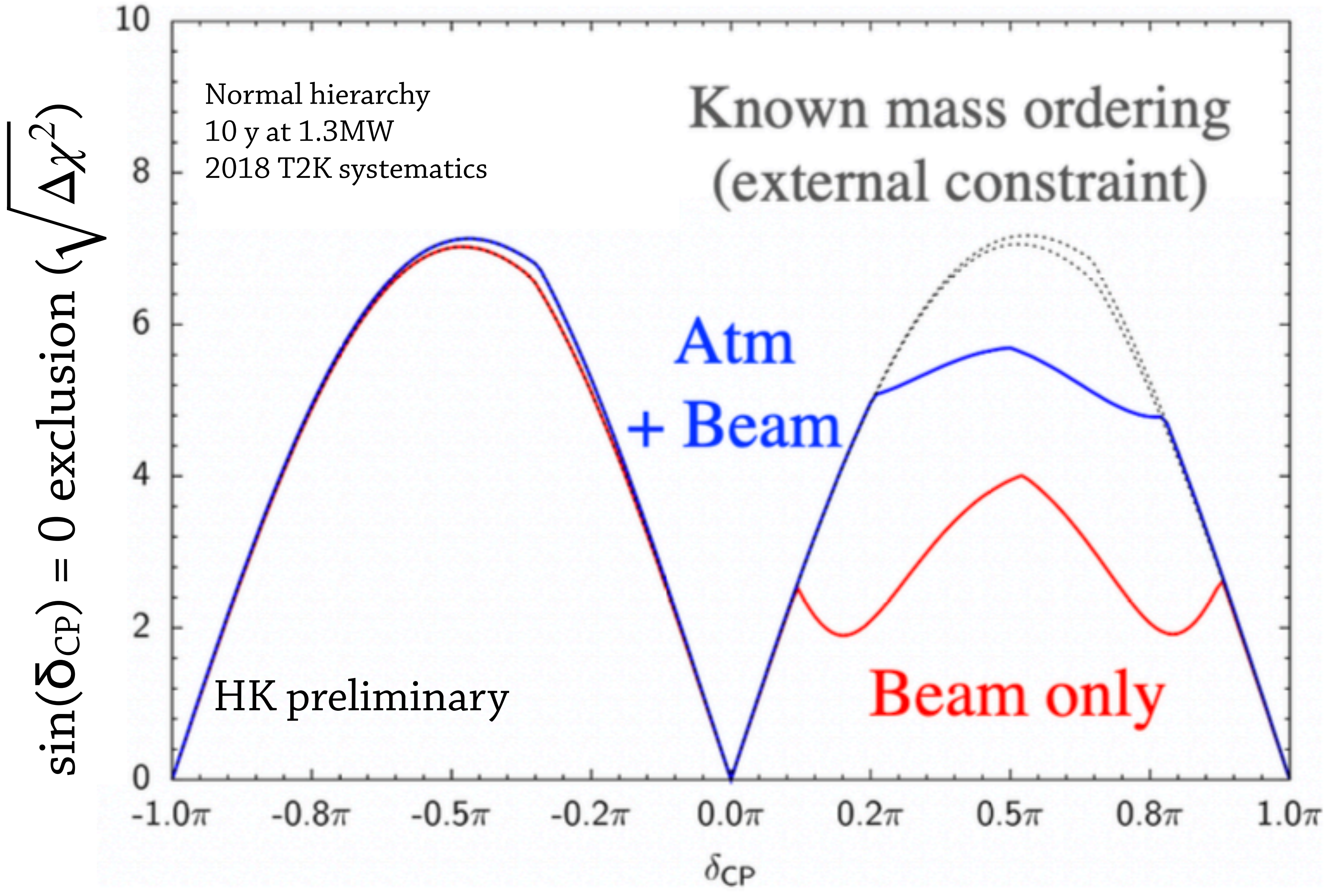
[3y to exclude the wrong MH for any δ_{CP} value]

○ ***T2HK*** itself do not have a lot of sensitivity

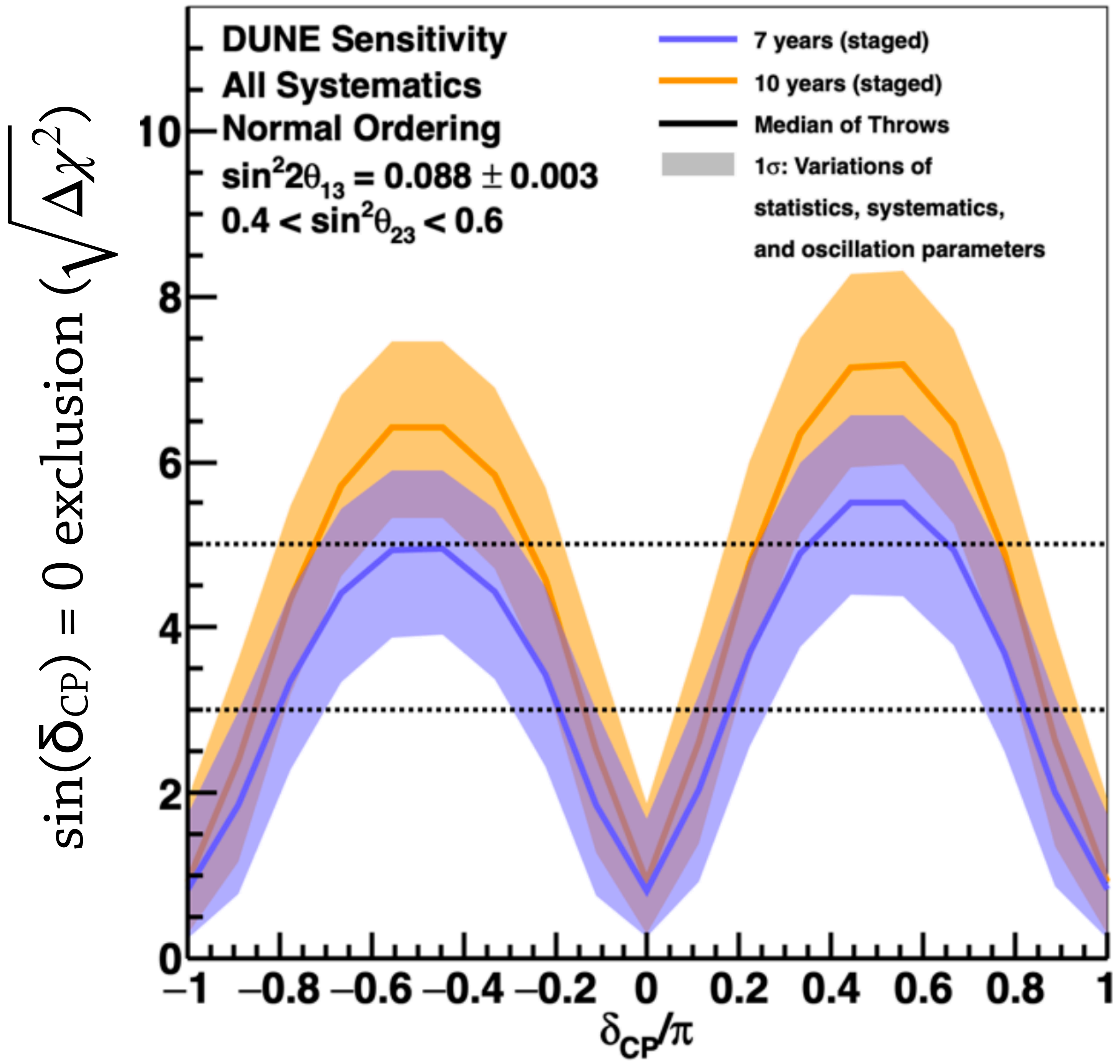
[can reach 5σ in 10y with beam + atmospheric ν]

Future ν accelerator experiments

T2HK in Japan



DUNE in the US



In 10 years of operation, if the MH is known:

- ○ ***DUNE*** can exclude $\delta_{CP} = (0, \pi)$ for 50% of δ_{CP} values
- ○ ***T2HK*** can reach 5σ for 60% of δ_{CP} values

CONCLUSIONS

Conclusions

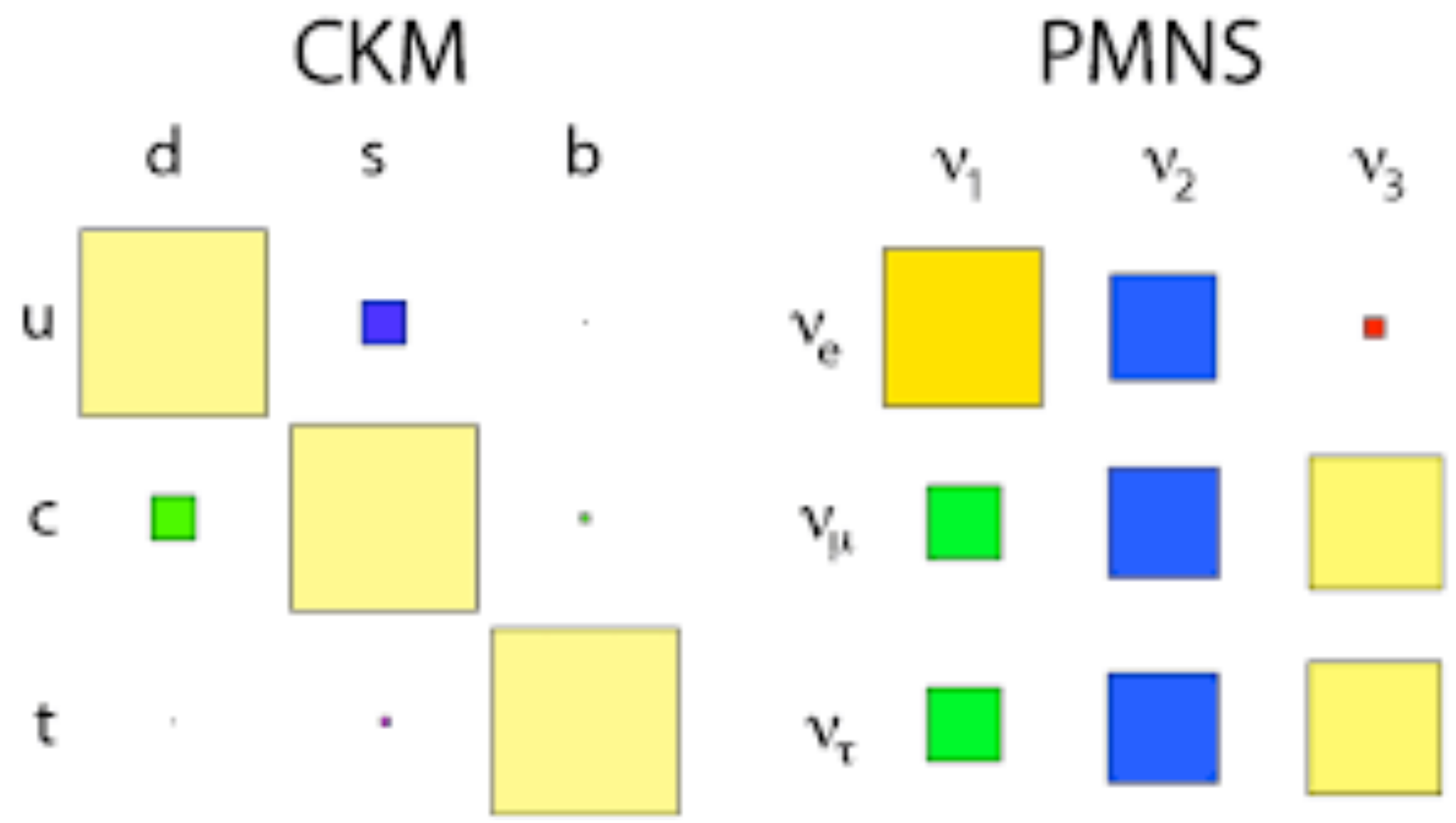
Neutrinos **oscillates** :

- ○ $\nu_e, \nu_\mu, \nu_\tau \neq \nu_1, \nu_2, \nu_3$
- ○ Two oscillation frequencies:

  fast (solar) and slow (atmospheric)
- ○ Neutrinos **mix** a lot more than quarks
- ○ In the next decade(s), all parameters measured:

  - matter/anti-matter asymmetry in the leptonic sector

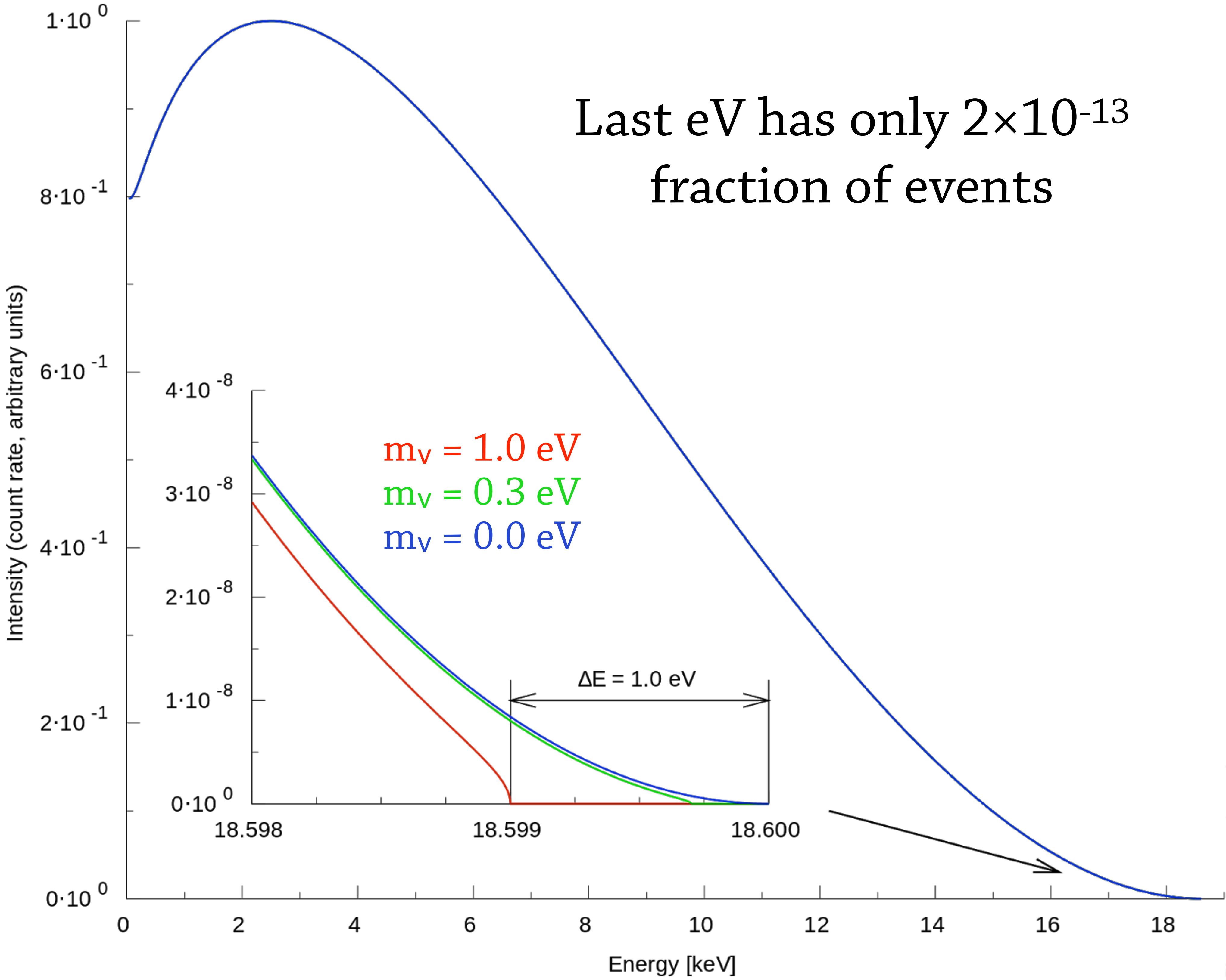
  - neutrino mass ordering
- ○ Neutrinos are **massive** - and it raises many other questions !
  - ‣ What mass ?
  - ‣ Mass mechanism ?
  - ‣ Could there be other neutrinos ?



Neutrino absolute mass ? ***KATRIN experiment in Germany***

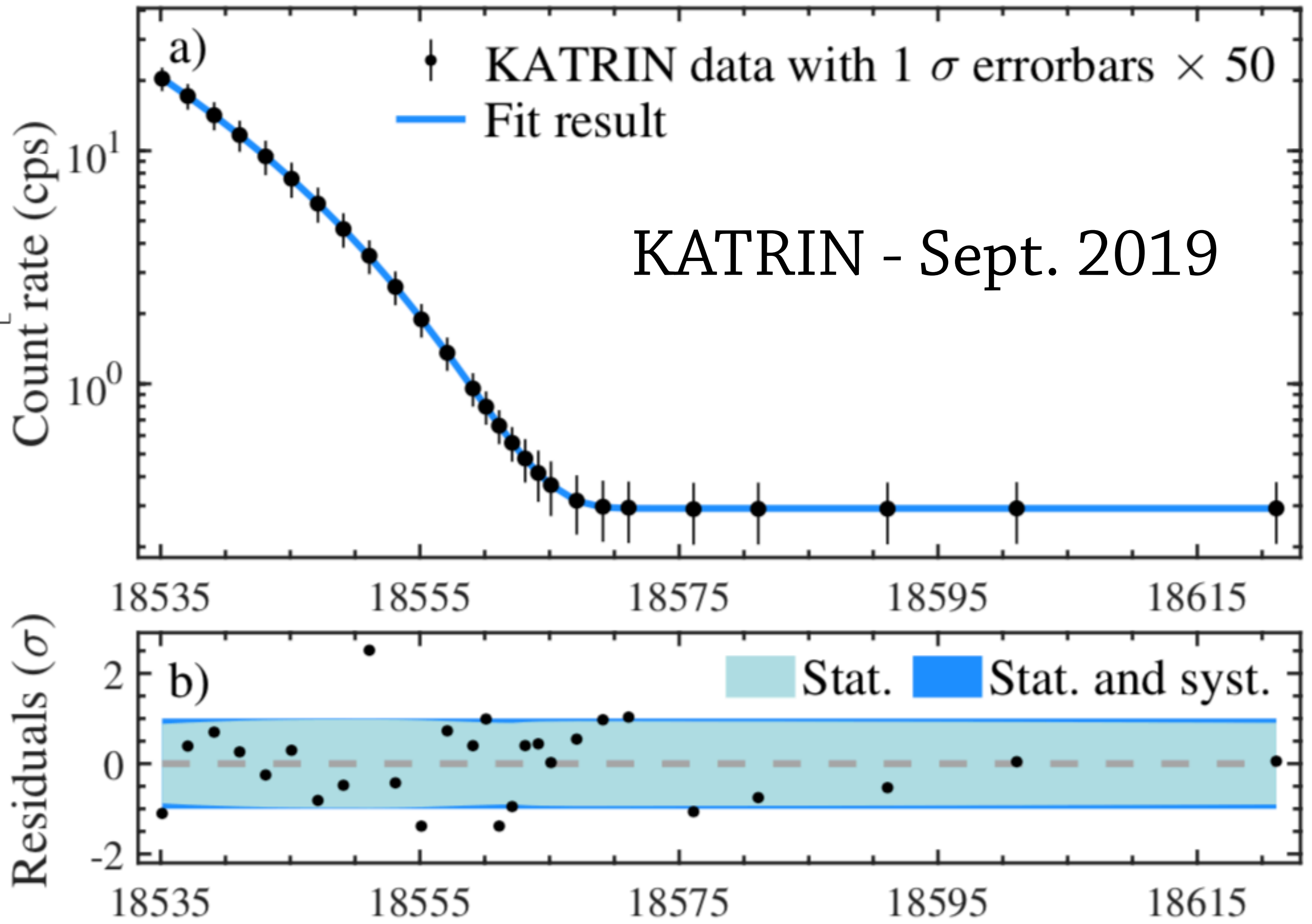
Look at the **end-point** of the β spectrum

↳ rare cases were the e^- takes most of the available energy



Last eV has only 2×10^{-13} fraction of events

Installation of KATRIN experiment



Current limit :

$m_{\bar{\nu}_e} \leq 1.1$ eV at 90% CL

How neutrinos get massive ? ***ββ0ν experiments (SuperNEMO, CUORE, SNO+)***

- The **Dirac** way

Through Higgs coupling

Need a sterile right handed ν

$$\mathcal{L}^{D}_{mass} = -m_D(\bar{\nu}_R\nu_L + \bar{\nu}_L\nu_R)$$

$$m_D = \frac{v}{\sqrt{2}}Y_v \leftarrow \sim 10^{-12} \text{ (why?)}$$

- The **Majorana** way

No distinction between ν and ν̄

Mass given by seesaw mechanism

Need massive neutrinos

$$\nu_R = C\bar{\nu}_L^T = \nu_L^C$$

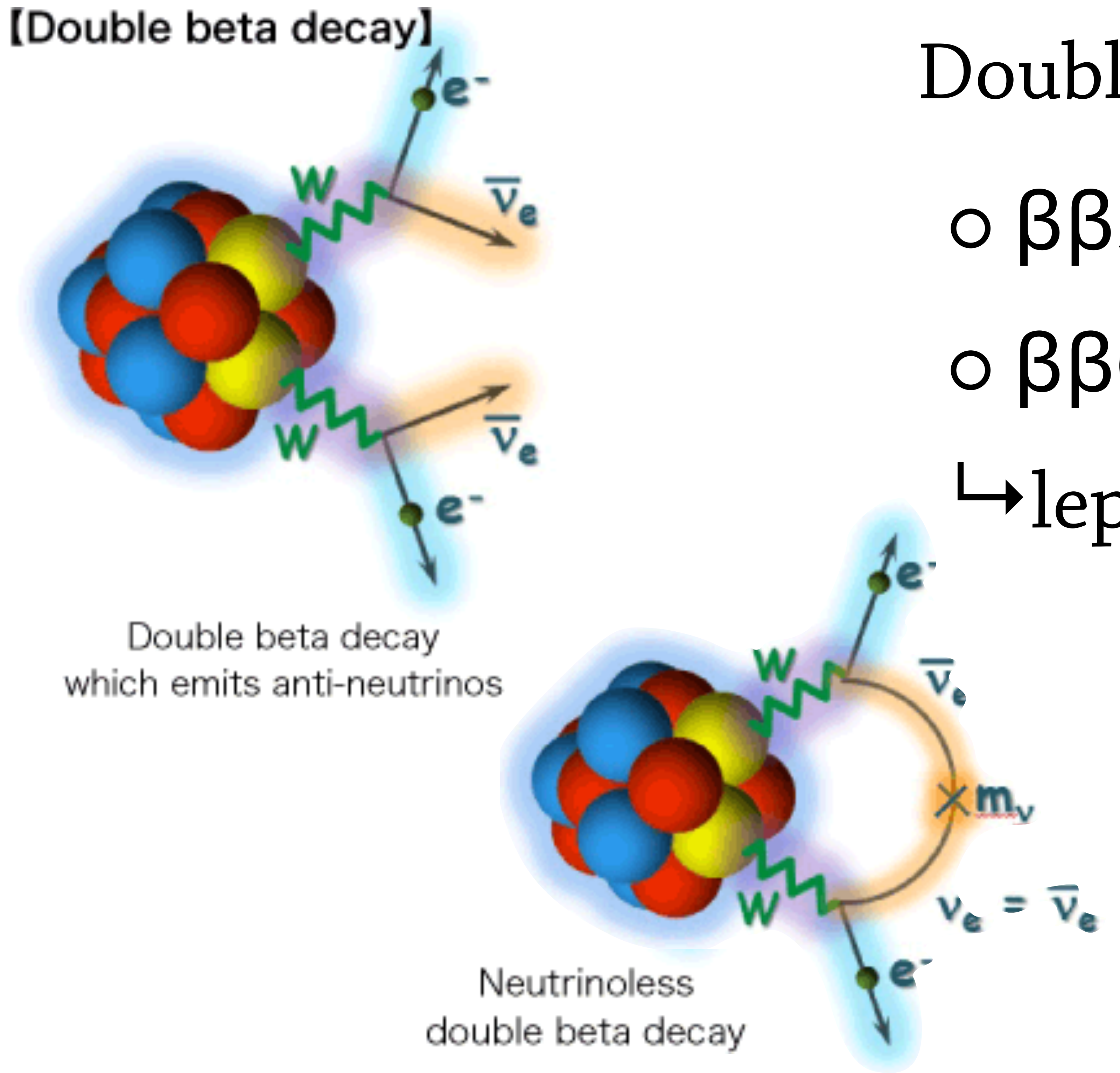
$$m = \frac{m_D^2 \leftarrow \text{Dirac term}}{m_R \leftarrow \text{Very big}}$$

→ ***Only one way to prove that neutrino are Majorana particles :***

Double β decay with **no** neutrino emission

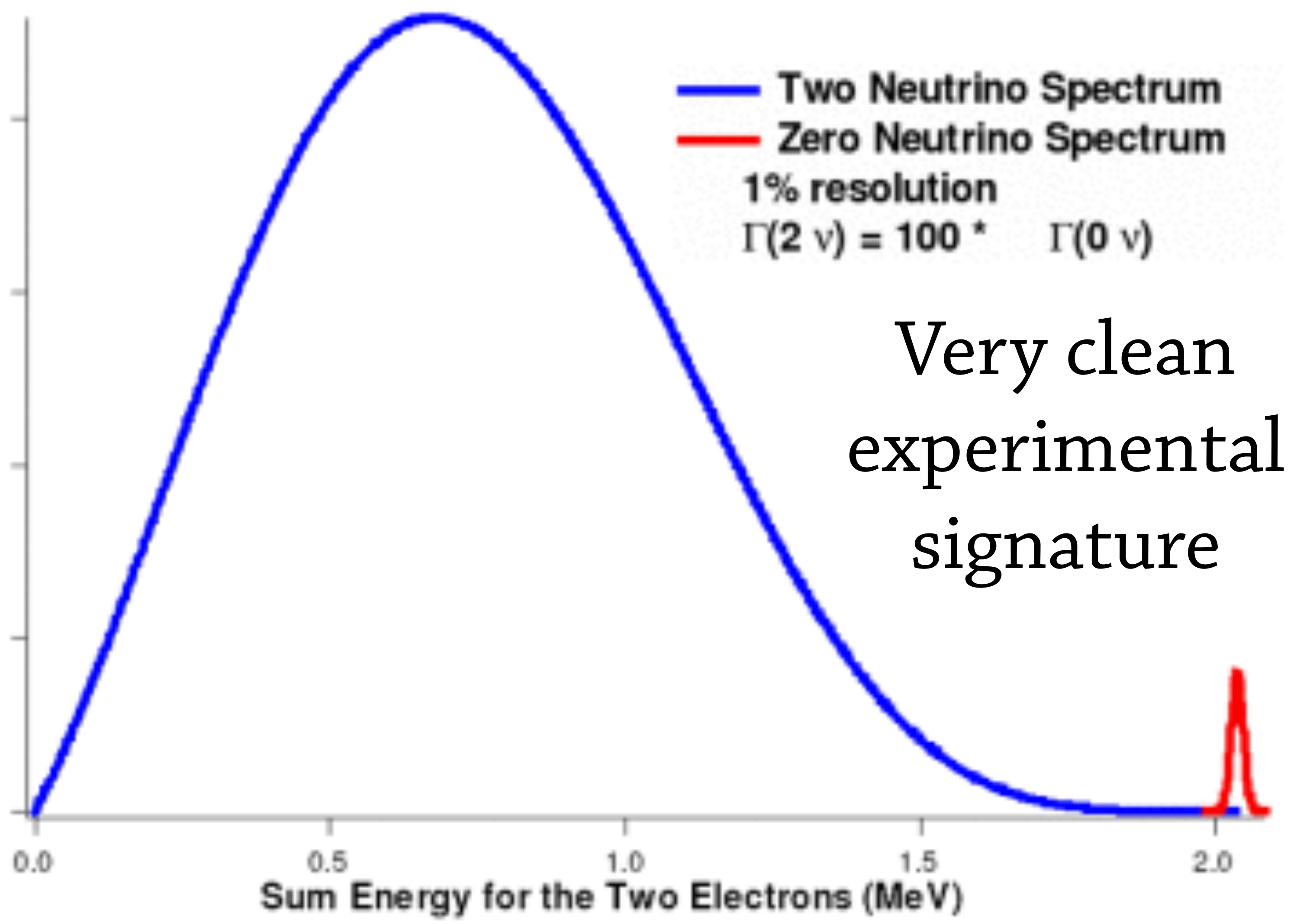
- ββ2ν is very rare (half life ~10^{18} - 10^{24} y)
- ββ0ν is **forbidden** in SM

↳lepton number violated by 2 units



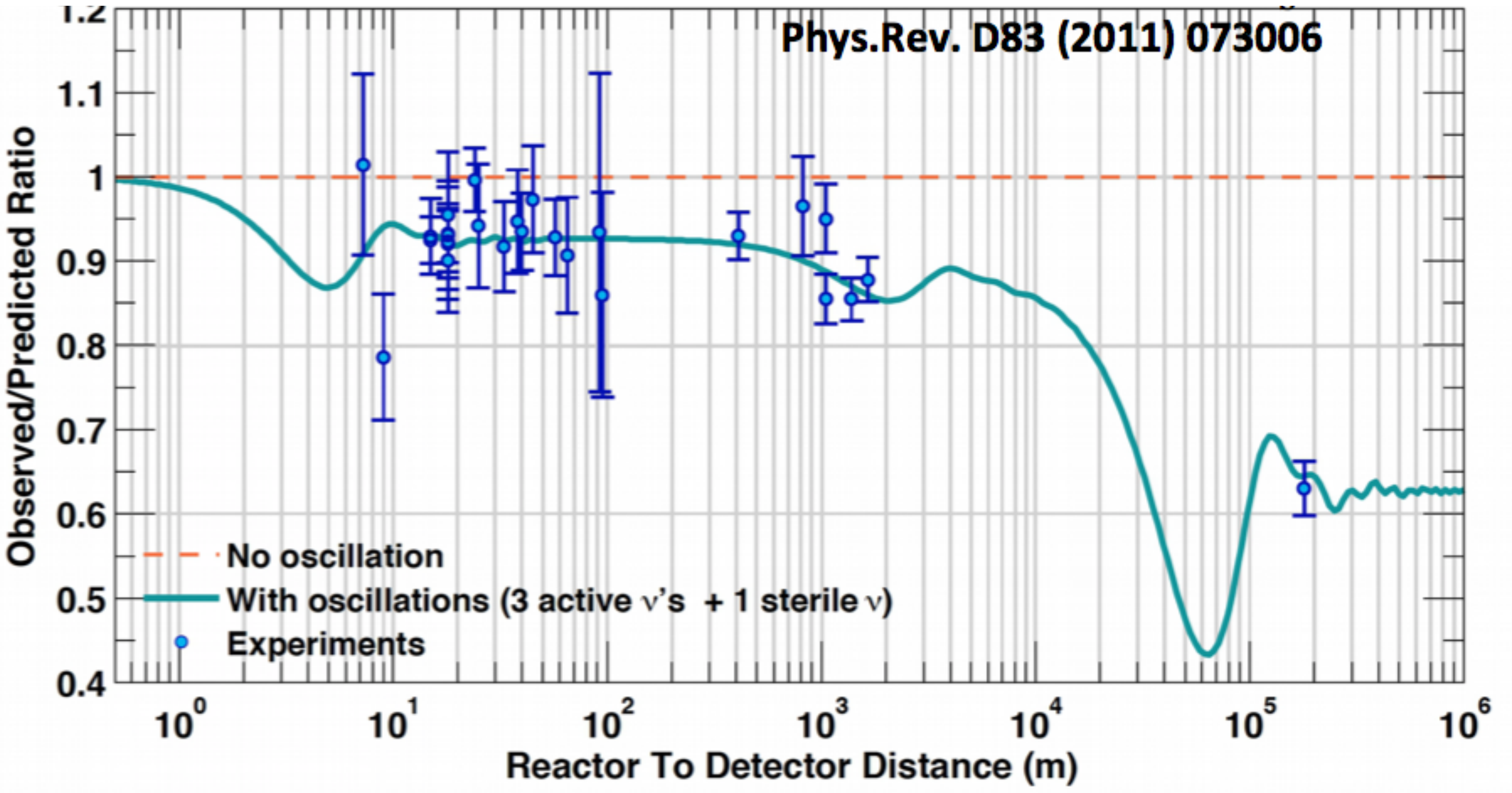
Double beta decay which emits anti-neutrinos

Neutrinoless double beta decay



Very clean experimental signature

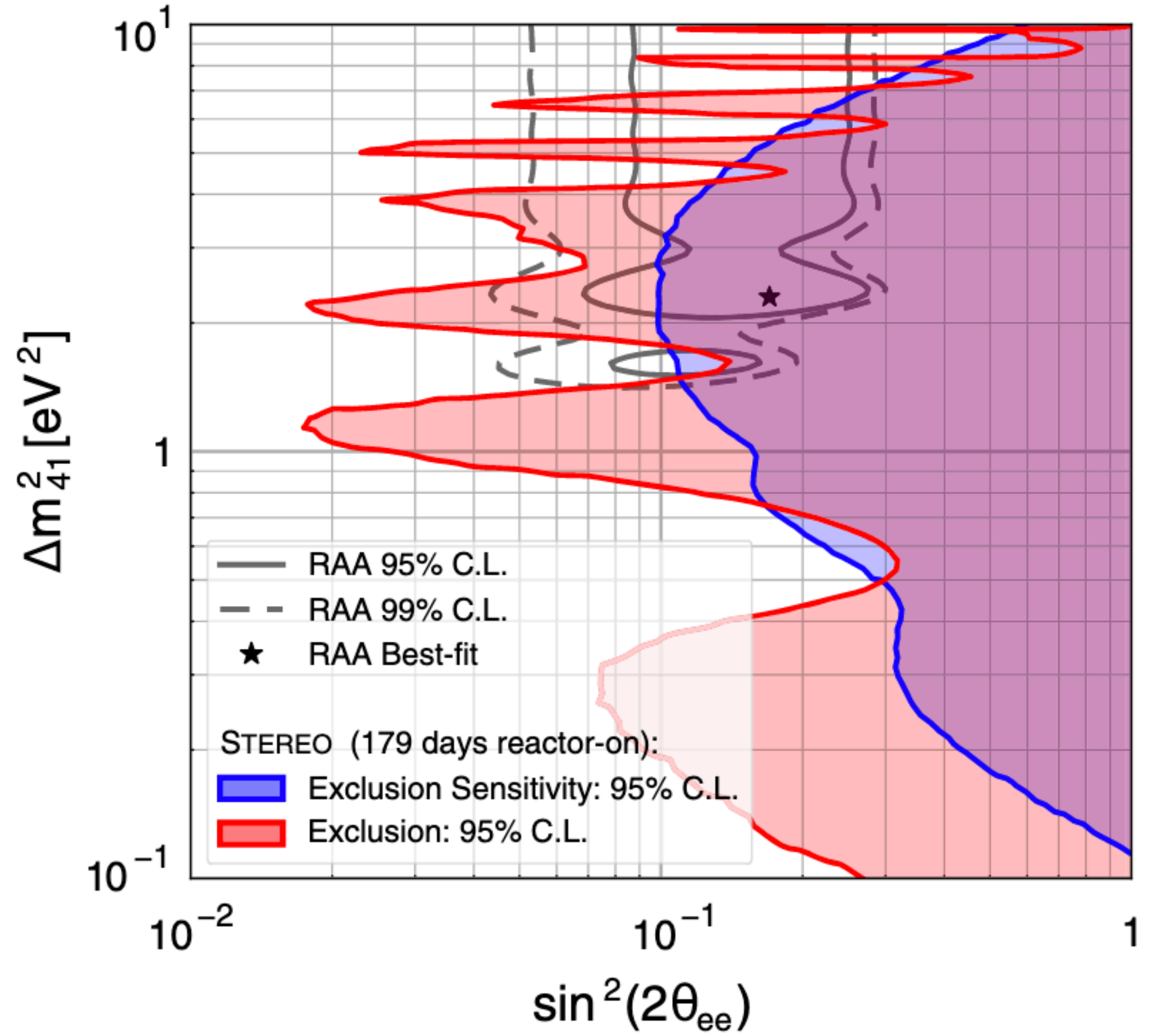
Only 3 Neutrinos ? ***STEREO, SOLID, PROSPECT,…***



A revised reactor $\bar{\nu}e$ flux analysis showed that all past ν experiments had a **~6% deficit** at small distances (3σ)

-> Problem with reactor flux ?
-> Existence of a sterile neutrinos ?

Latest results from STEREO



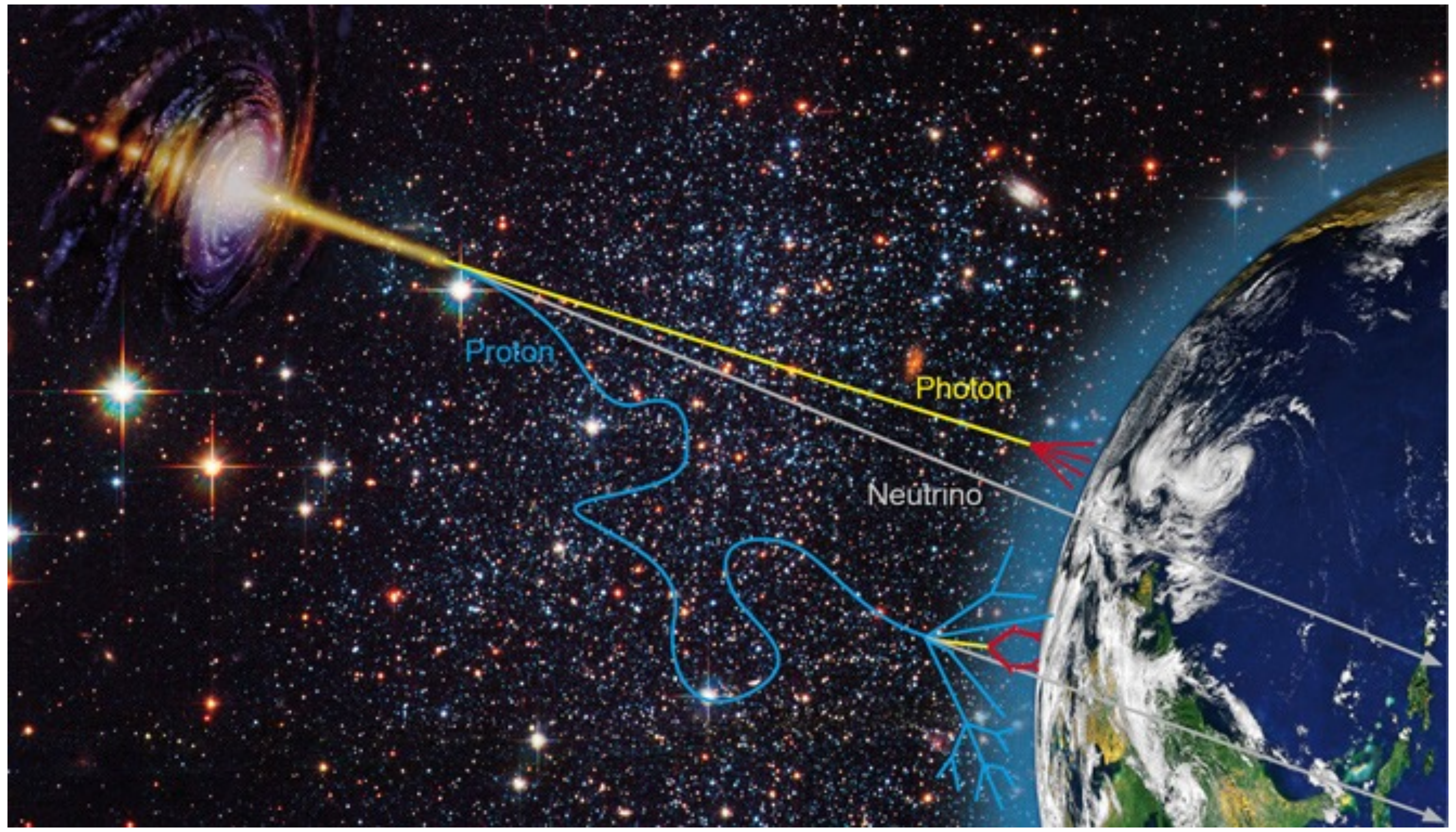
- Sterile because this neutrino cannot interact with weak force: it would be invisible
- But all 4 neutrinos could oscillate within each others

→ ***New mass splitting and new mixing angle***

Best fit parameters of reactor anomaly:

$\Delta m^2 \sim 2$ eV^2
$\sin^2(2\theta) \sim 0.15$
$L_{osc} \sim$ few m

Neutrino Astronomy: ***ICECUBE, KM3NET***

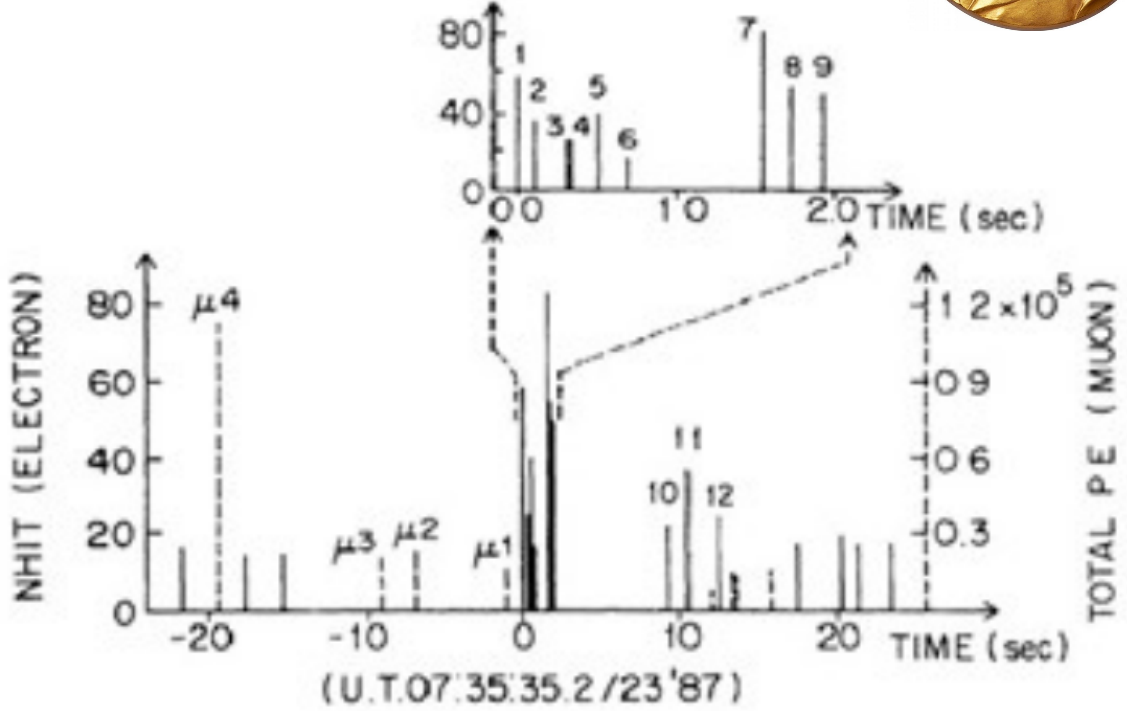


- Unlike protons & gammas, neutrinos **points to their sources**
- Can probe the inside of the structure
- **No GZK threshold** : can probe far away objects

On February 23rd 1987, a supernova exploded in the large magellan cloud (170 000 l.y.)

3h before the light signal, three neutrino detectors observed a large number of events in a very short time (**24 events in 13s**)

9@Kamiokande

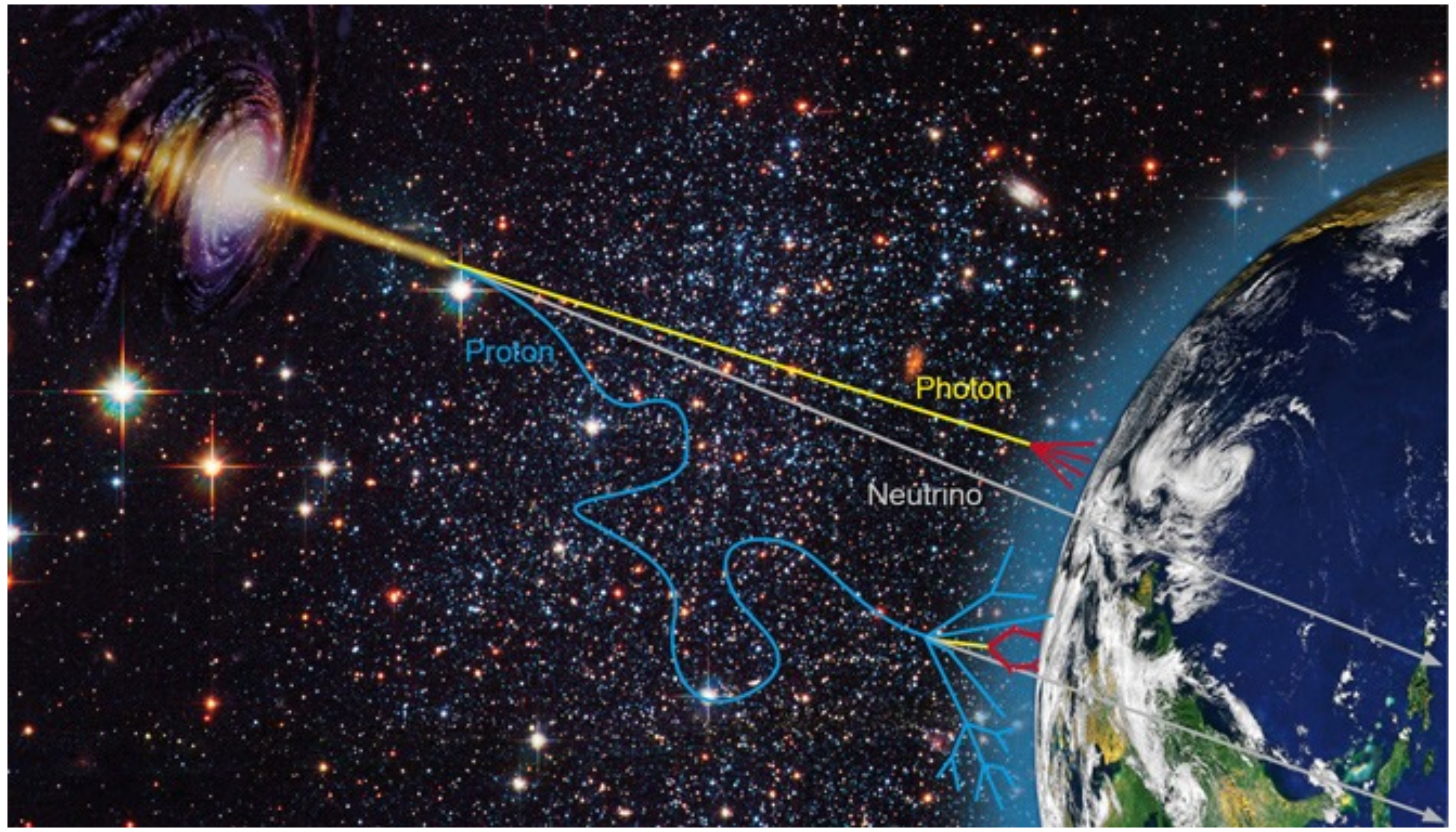


- 99% of the SN energy is released as neutrinos
- 1st case of neutrino astronomy and multi-messenger
- all ν experiments waiting for next nearby SN explosion

SN1987A

© Anglo-Australian Observatory/David Malin Images

Neutrino Astronomy: ***ICECUBE, KM3NET***



- Unlike protons & gammas, neutrinos **points to their sources**
- Can probe the inside of the structure
- **No GZK threshold** : can probe far away objects

On September 22nd 2017 : **Simultaneous** light & neutrino detection from the TXS 0506+056 blazar (3σ, $E_\nu = 290$ TeV)

(blazar = Active Galactic Nucleus with one jet pointing to earth)

- 1st case of planned **multi-messenger** astronomy
- Confirmed that blazar emits neutrinos