

SURFACE ARRAY

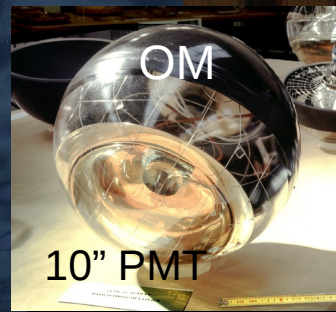
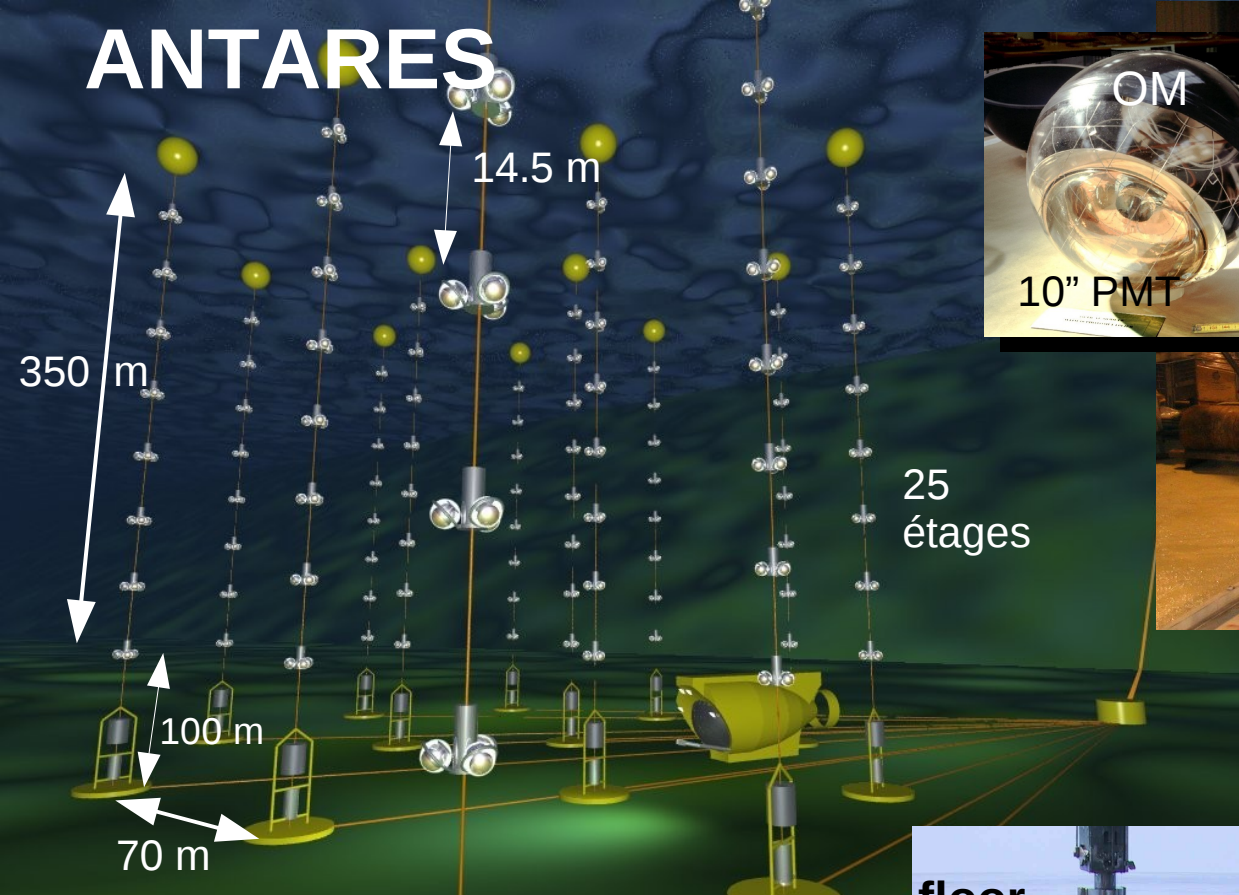
Détecteur de surface pour ANTARES



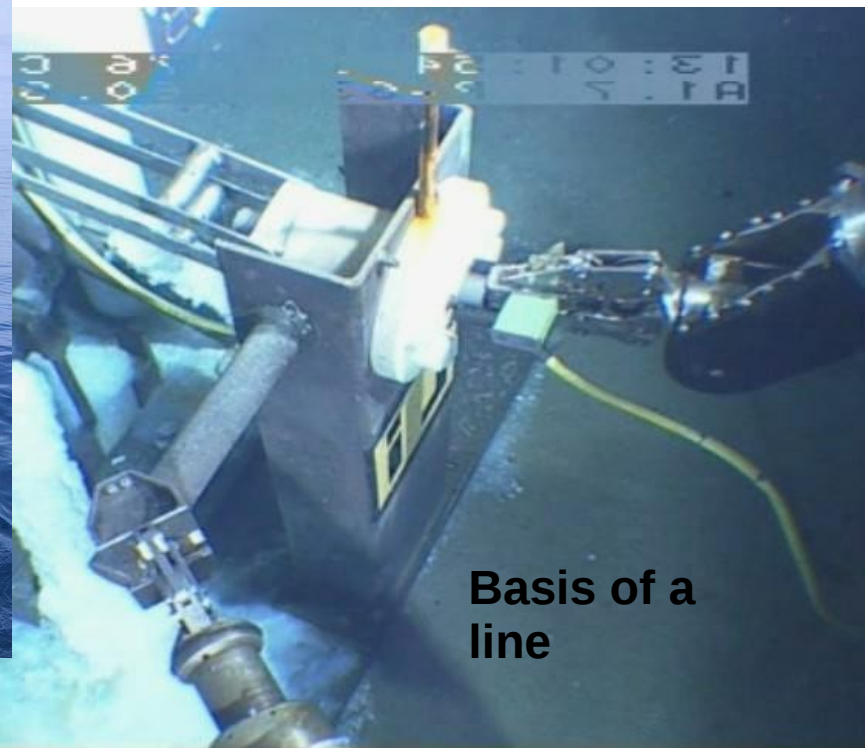
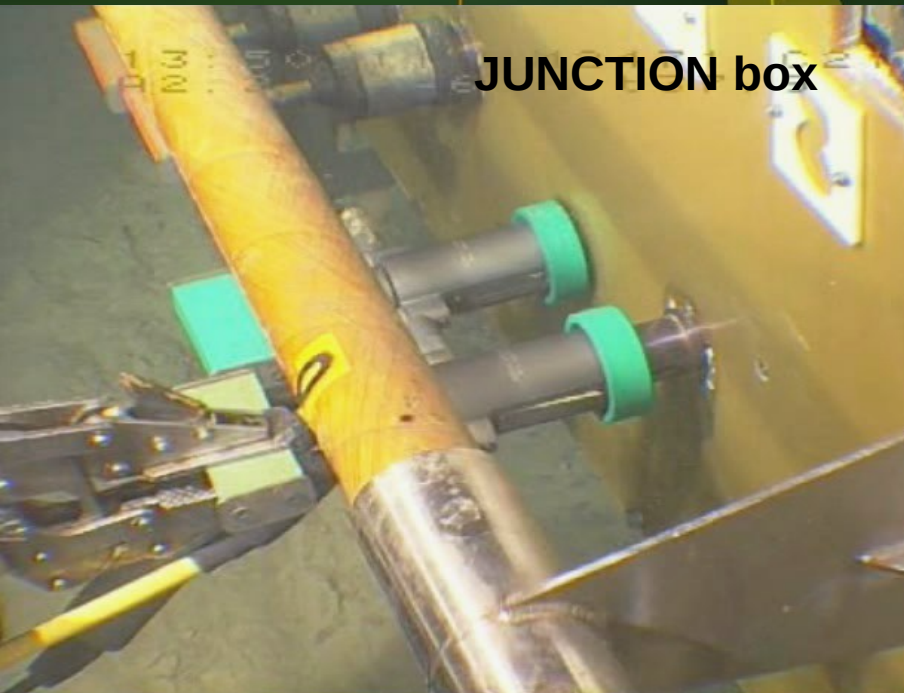
JP Ernenwein
CPPM

GDR V
29/10/2009

ANTARES



ANTARES, 2475m, 3 PMT/floor
12 lines of 25 floors 885 PMTs



2005**2006****2007****2008**

Mar:
MILOM

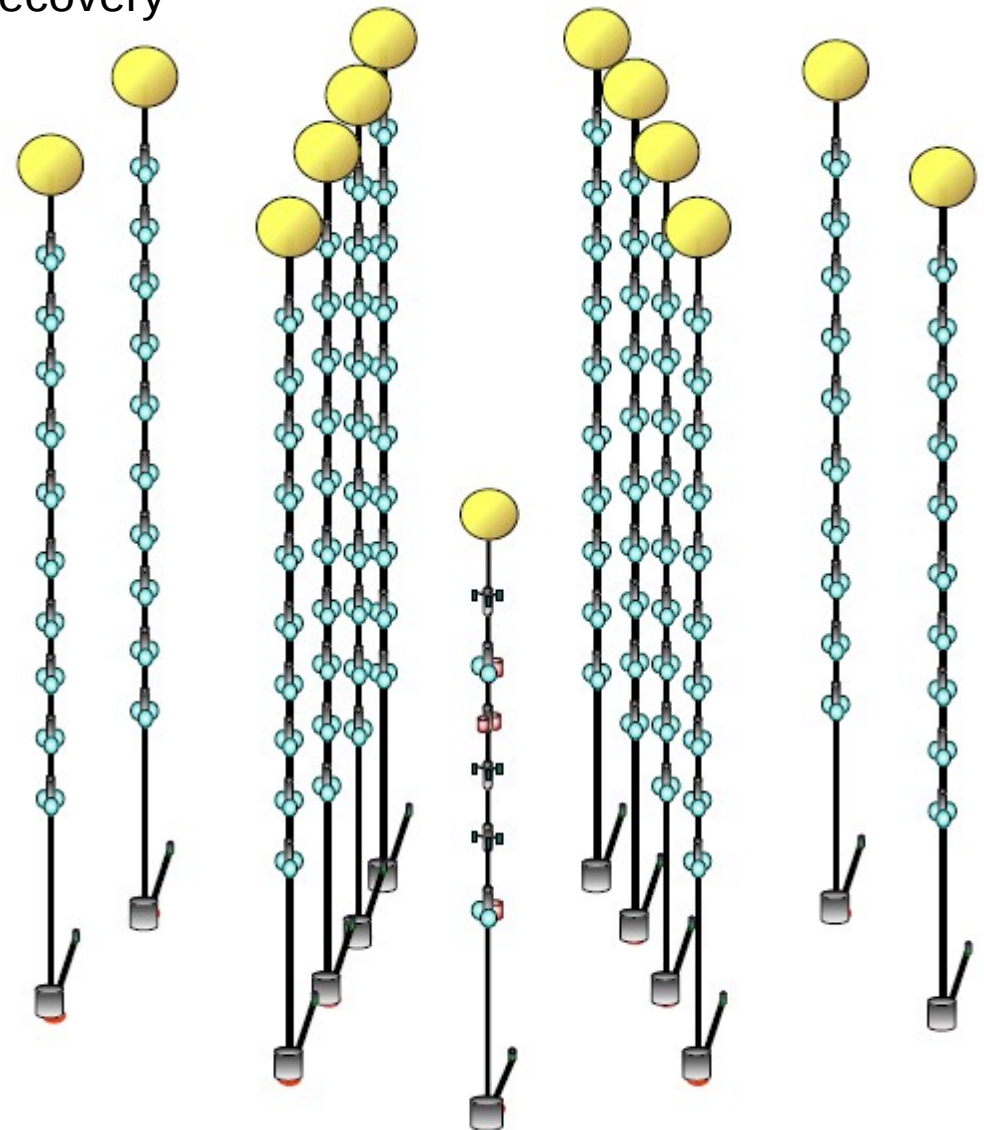
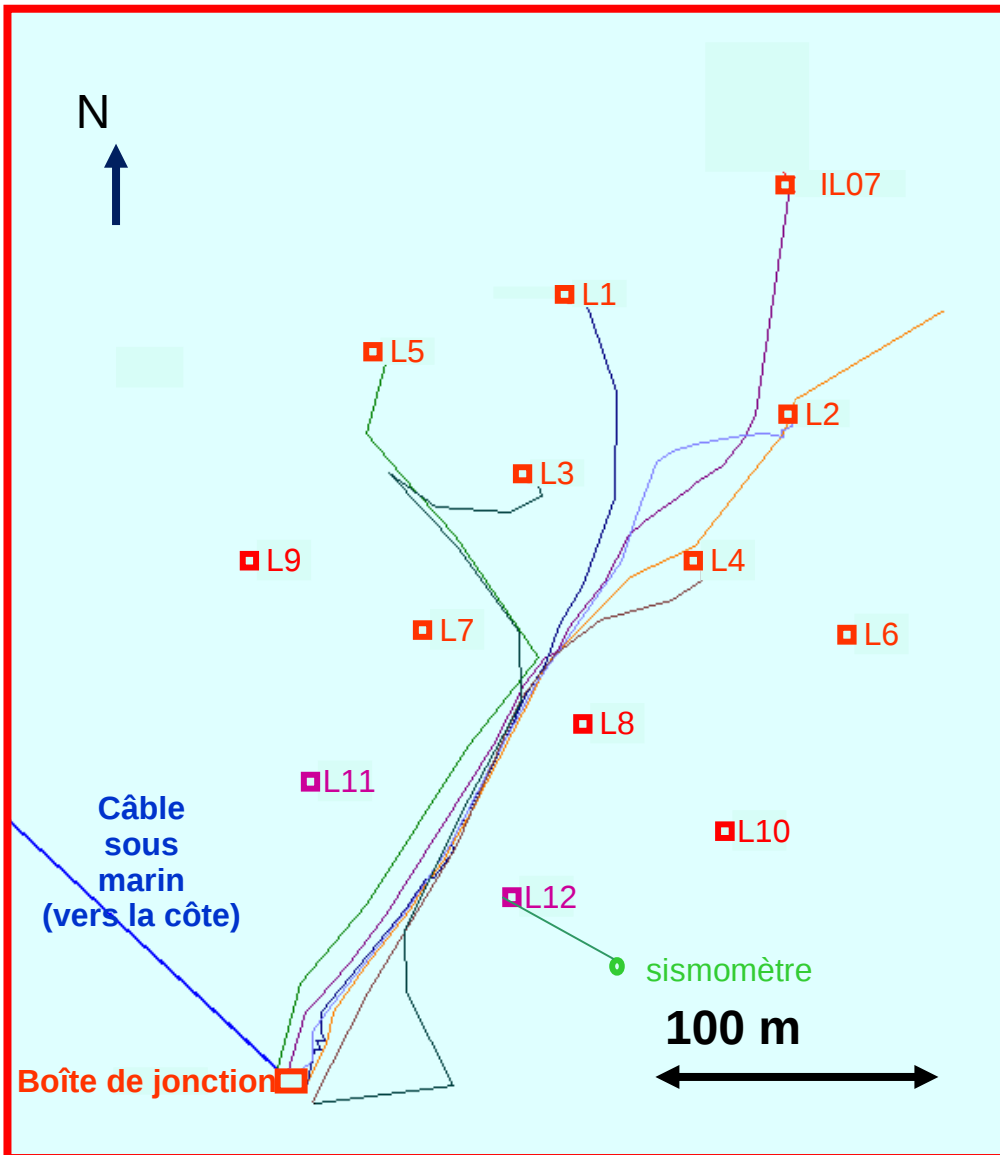
Mar:
L1

Sep:
L2

Jan :
L3-5
MILOM
recovery

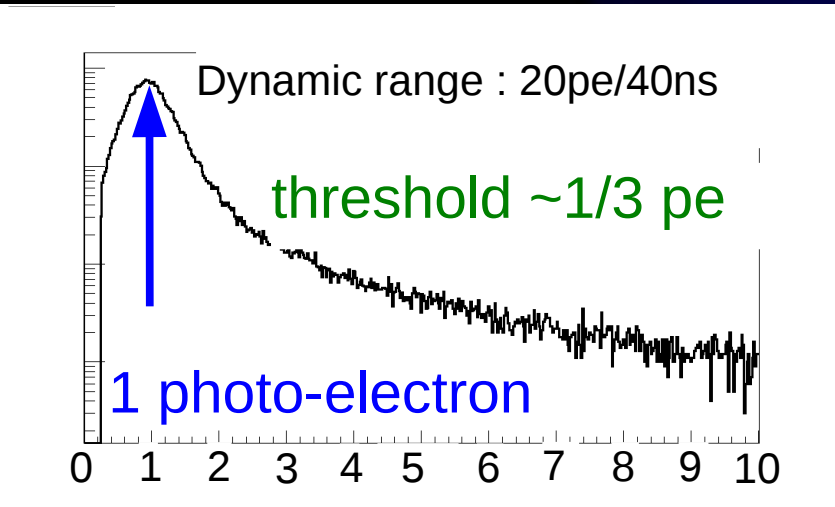
Dec:
L6-10
IL07

Jun:
L11,12

**Connection of lines**

Detection principle

3D matrix



Cherenkov light (μ)

required sensitivity: photo-electron

42°

shower

good energy resolution ($O(30\%)$),
poor angular resolution ($O(10^\circ)$)

ν_e

Charged current interaction (W)

track:

Good angular resolution ($O(\text{degree})$ @ $E > 10$ TeV),
Poor energy resolution (factor 2-3)

ν_μ

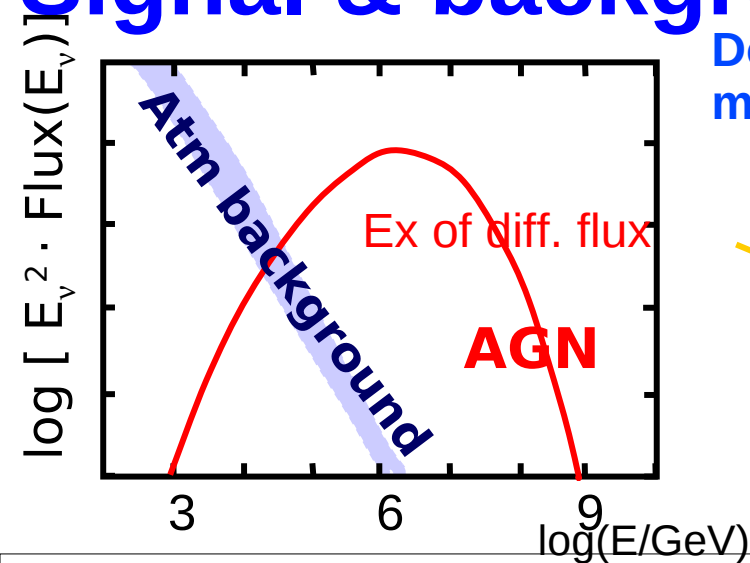
Measurements
time ($O(\text{ns})$), amplitude (30%)
& hit position ($O(10 \text{ cm})$)



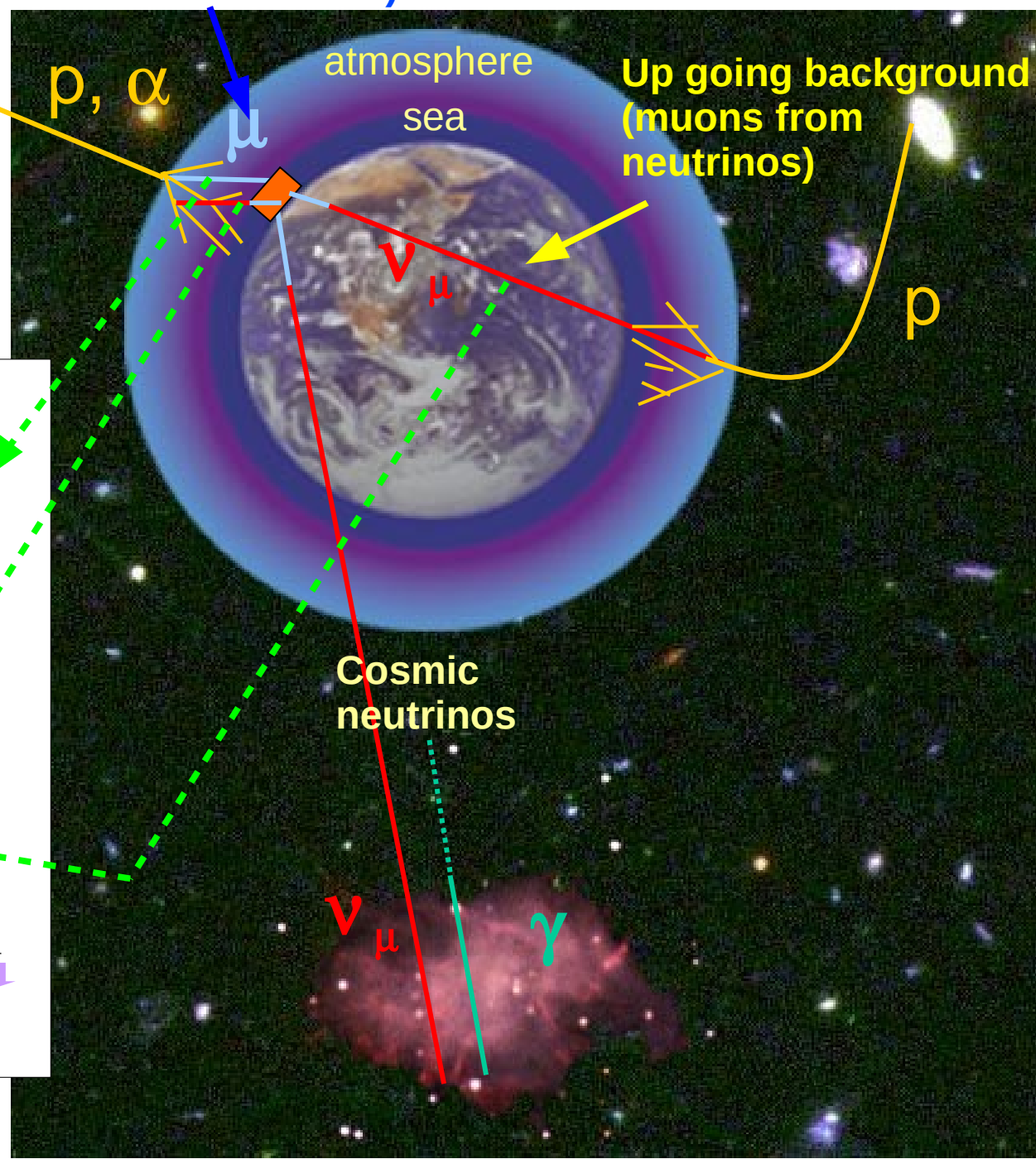
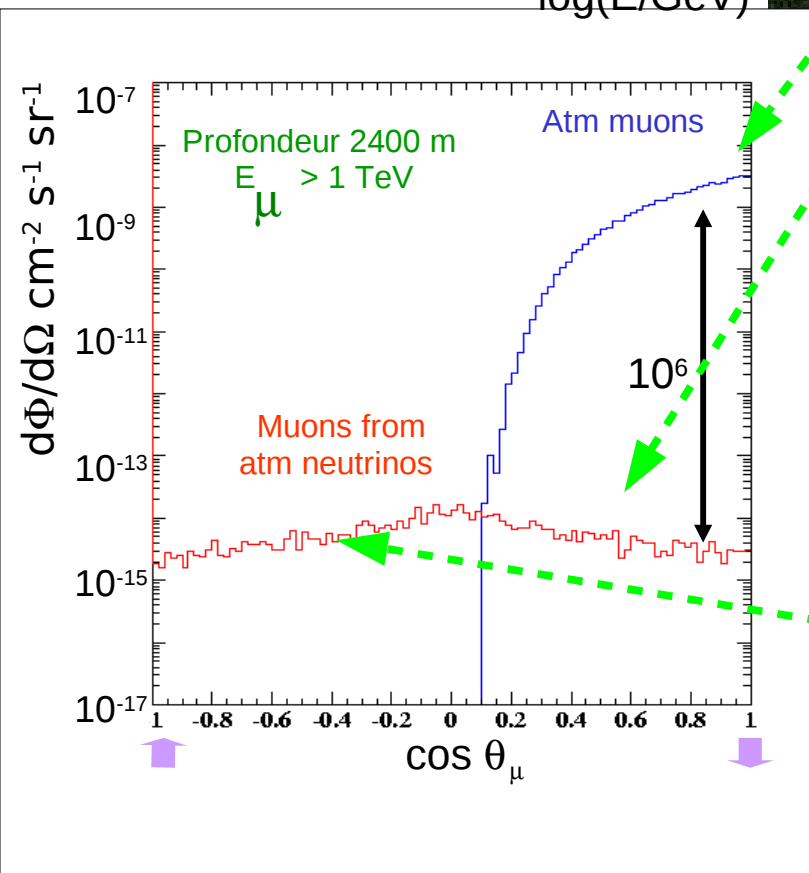
Muon track or shower measurement

© François Montanet

Signal & background



Down going background events (atm muons and muons from neutrinos)

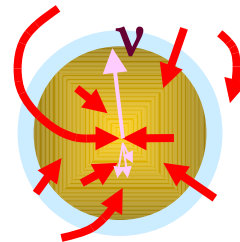
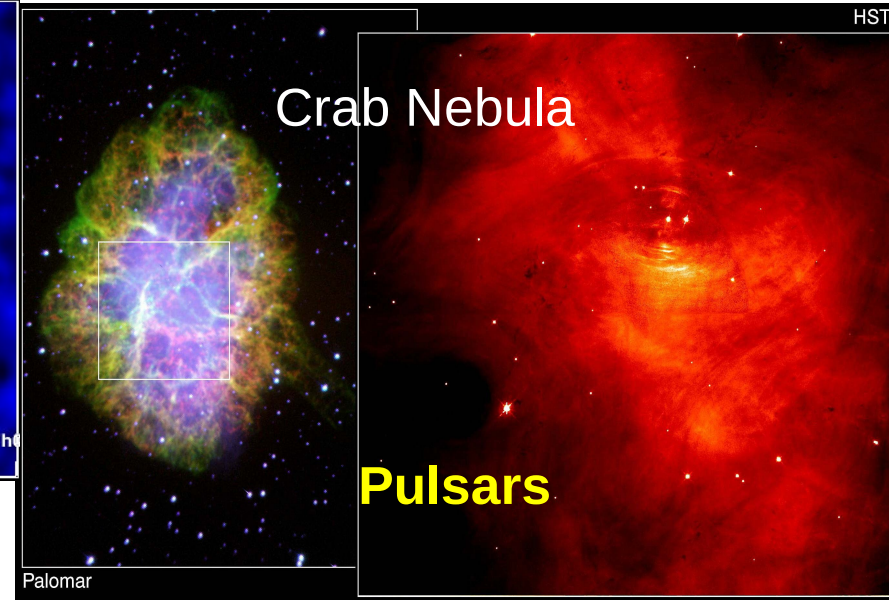
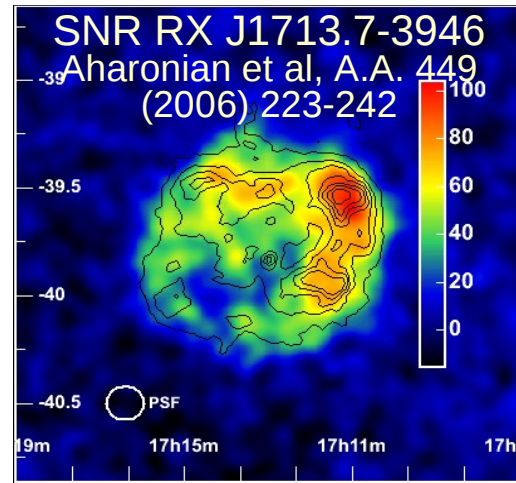
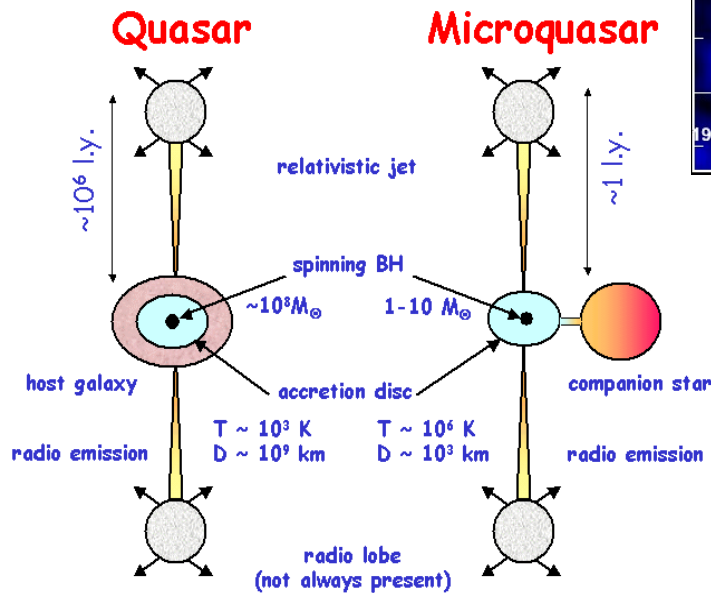


Expected neutrino sources

Galactic:

**Supernovae,
Supernovae remnants,**

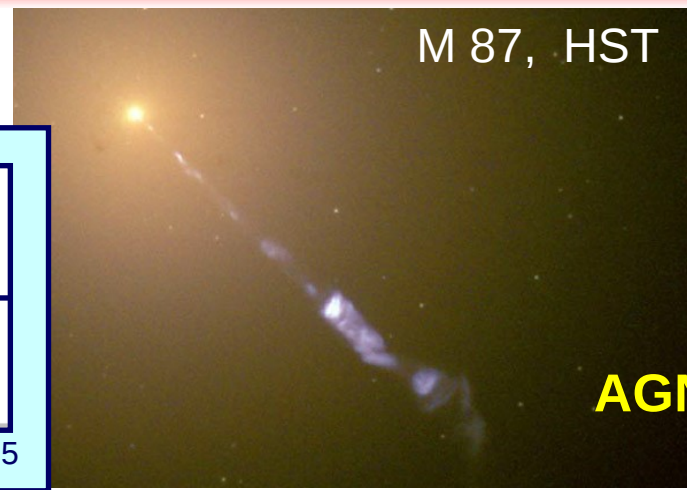
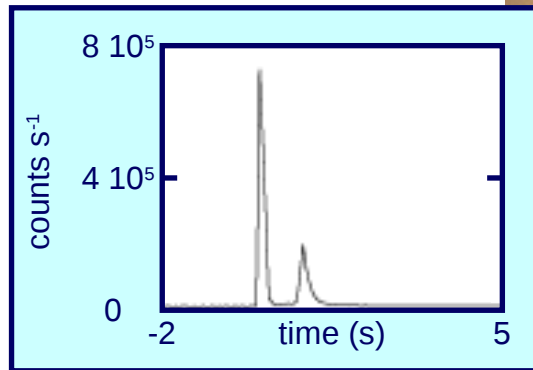
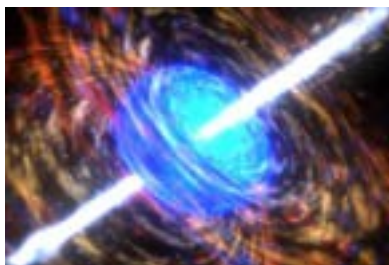
Micro Quasars:



Dark Matter : neutralino annihilation in massive objects (Sun, Earth, GC)

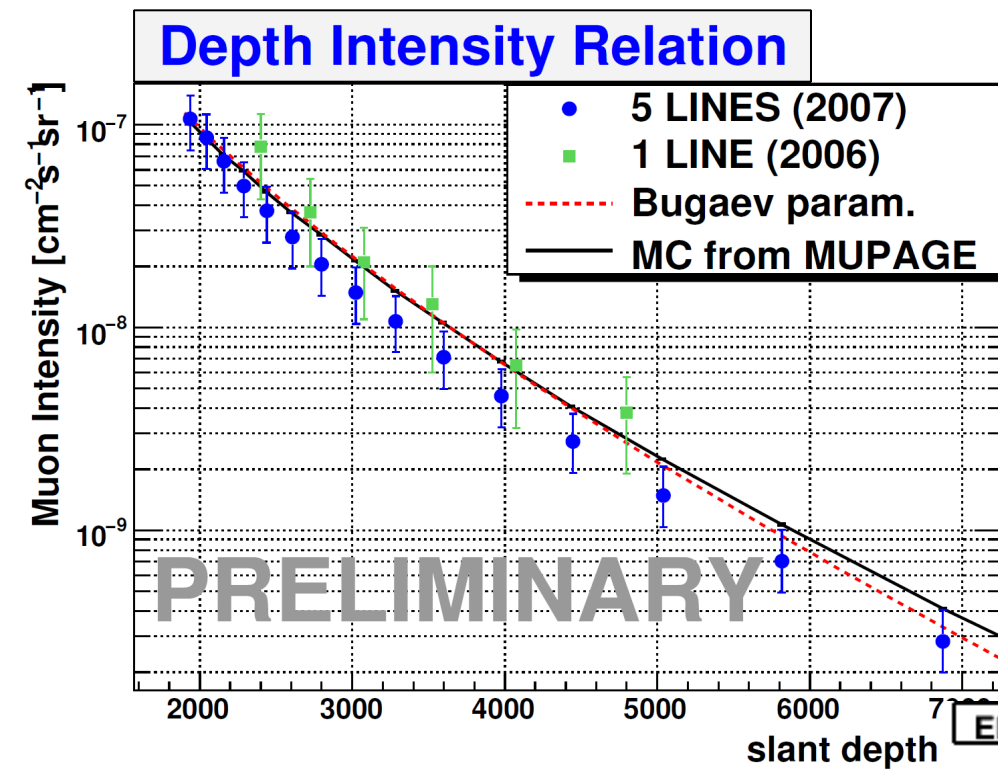
EXTRA GALACTIC:

GRB



Current results

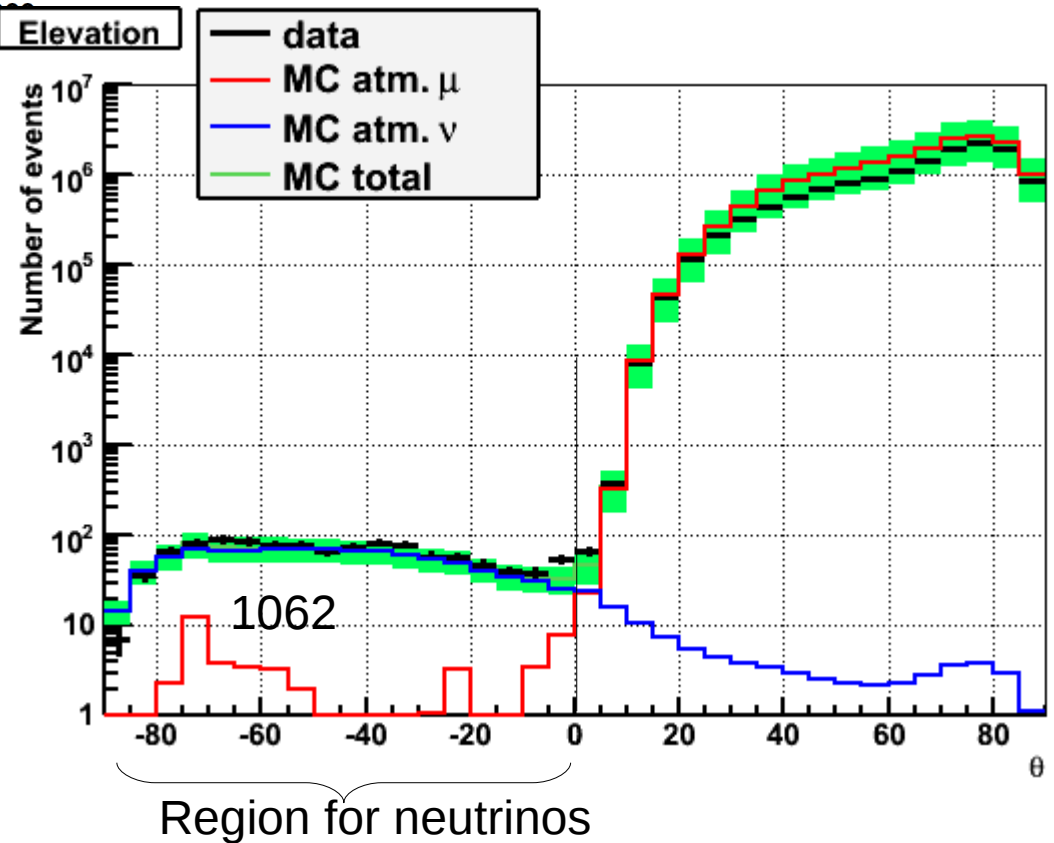
Atmospheric muons



Neutrinos & muons

341 days in 2007/2008

(5 lines, 9, 12 lines)

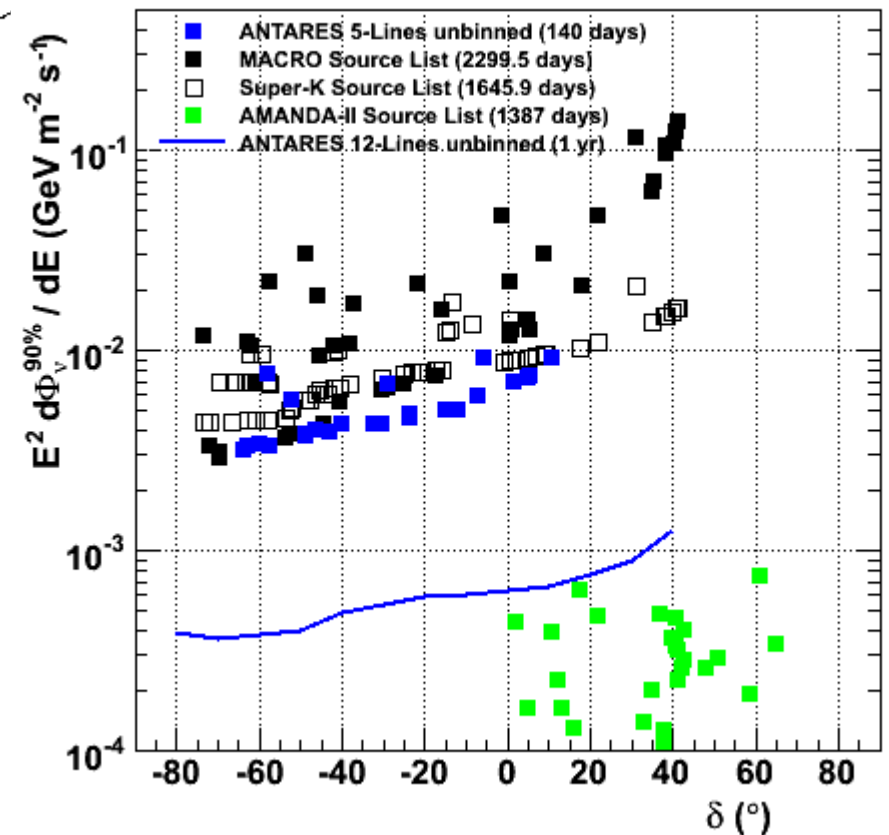


Current results

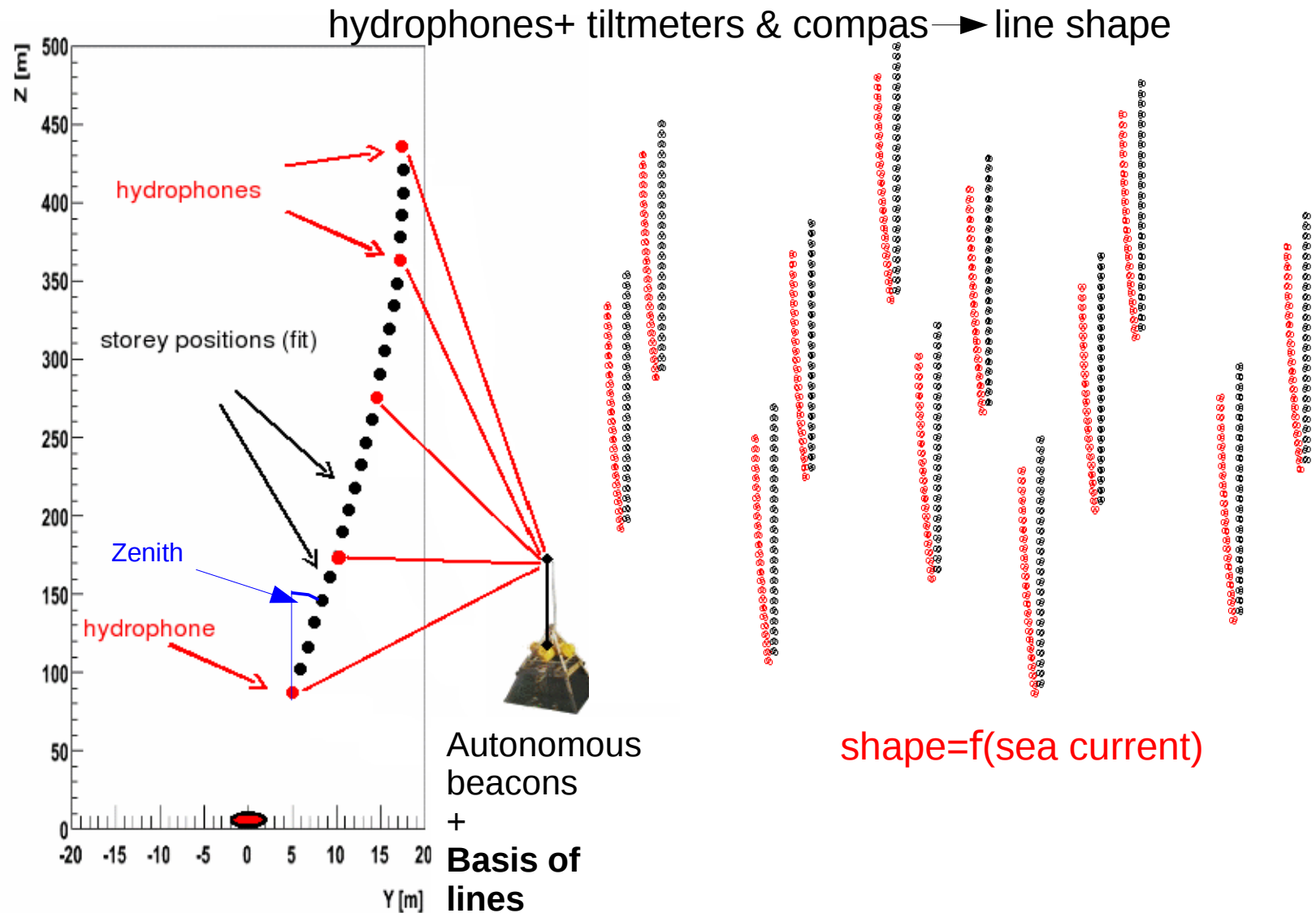
140 days

preliminary

Point source
search with ANTARES
in its 5 lines stage

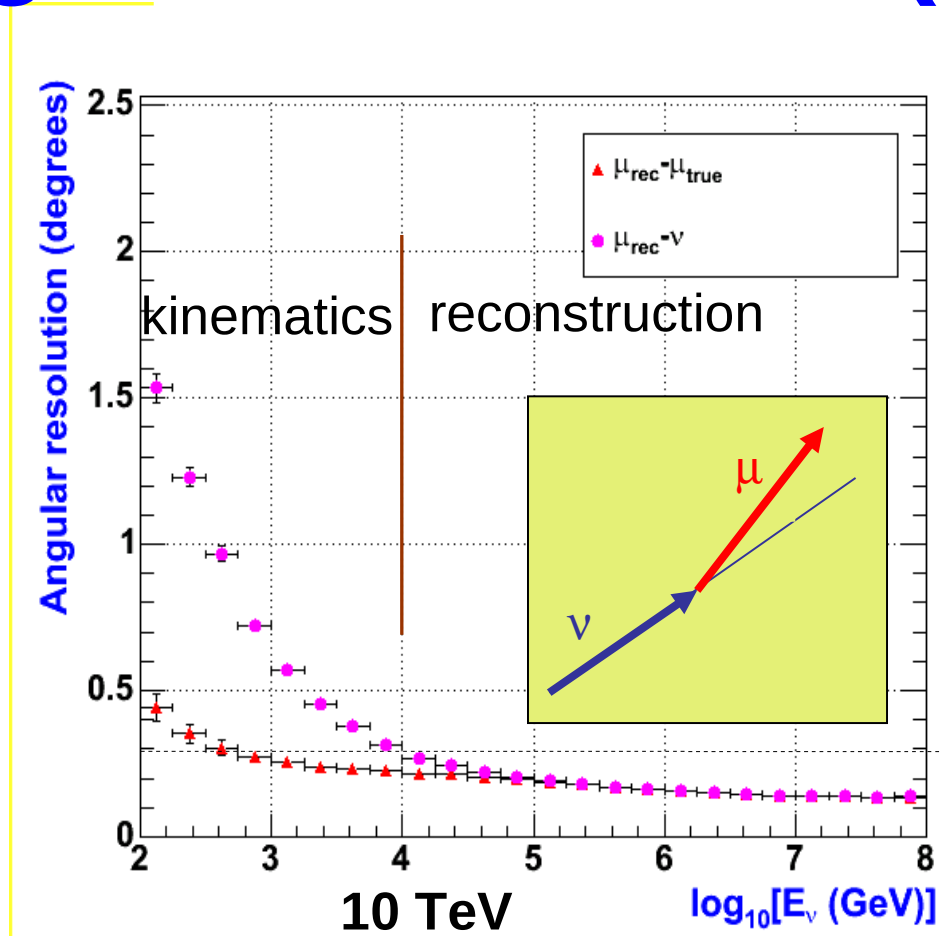


Measurement of the position of the lines



Position of line anchor : acoustic measurements from boat, itself positioned with a GPS.

Angular resolution (median)



► 0.3°

Can be cancelled by the uncertainty on the absolute pointing:

To be measured independently :

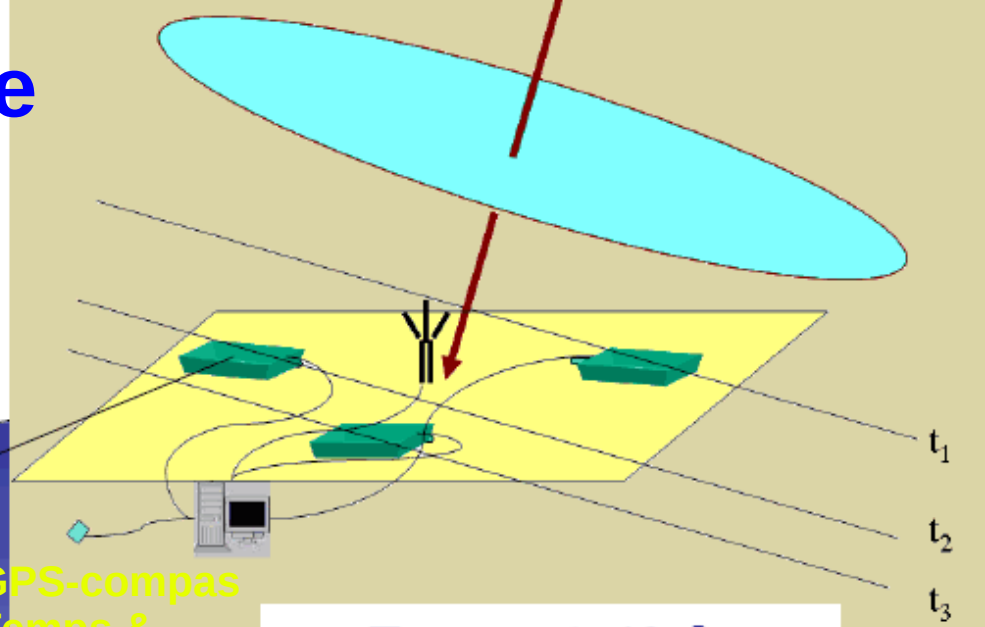
Moon shadow, cliff shadow : O(year) but free,

Surface Array : O(week) but not free!



Principle of the Surface Array

Gerbe atmosphérique

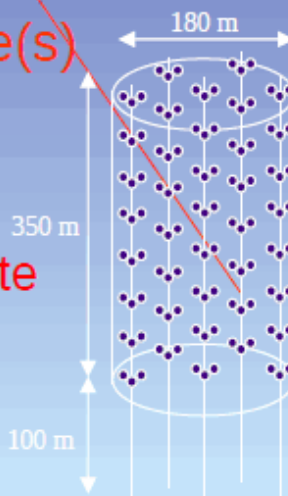


GPS-compas
Temps &
orientation
& inclinaison

Détecteur de surface (unités de 0.8 m^2)

Muon(s)
atmosphérique(s)

Trace reconstruite
par ANTARES



Temps relatifs →
direction de l'axe de
la gerbe



Événements
enregistrés
sur
disque dur

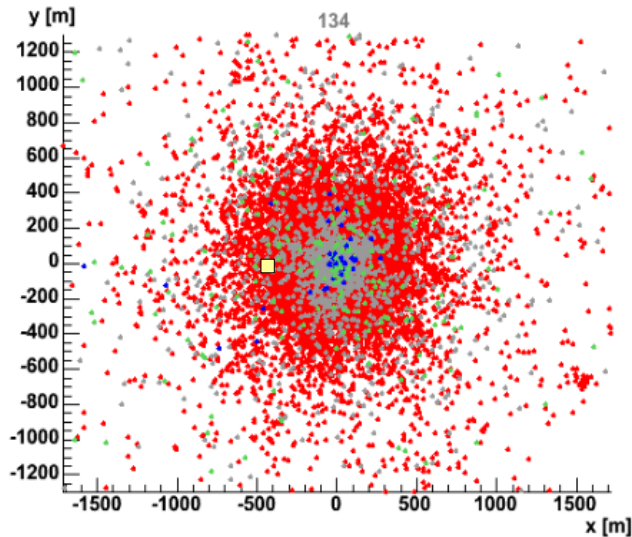
ANTARES
Événements
enregistrés
sur disque dur.



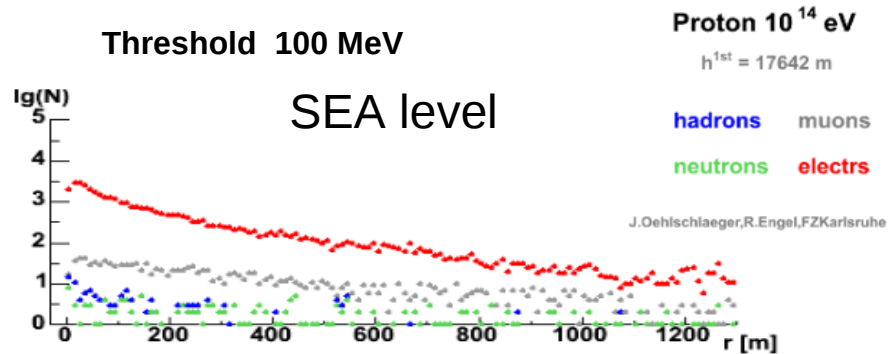
Association
en
Temps
(10 ms)
et
direction

Atmospheric showers

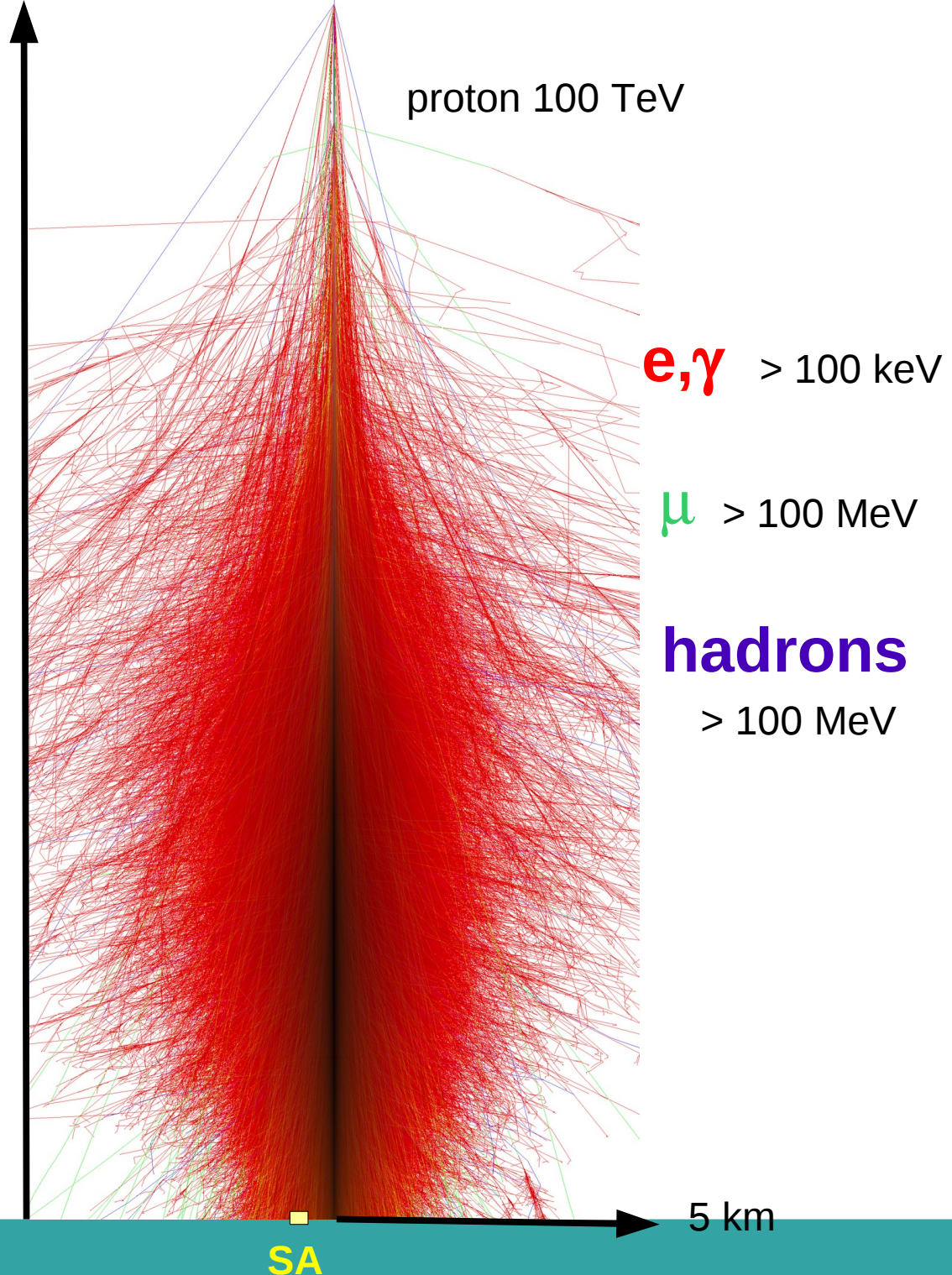
CORSIKA



30 km



F. Schmidt, "CORSIKA Shower Images", <http://www.ast.leeds.ac.uk/~fs/showerimages.html>



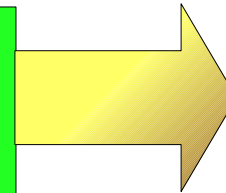
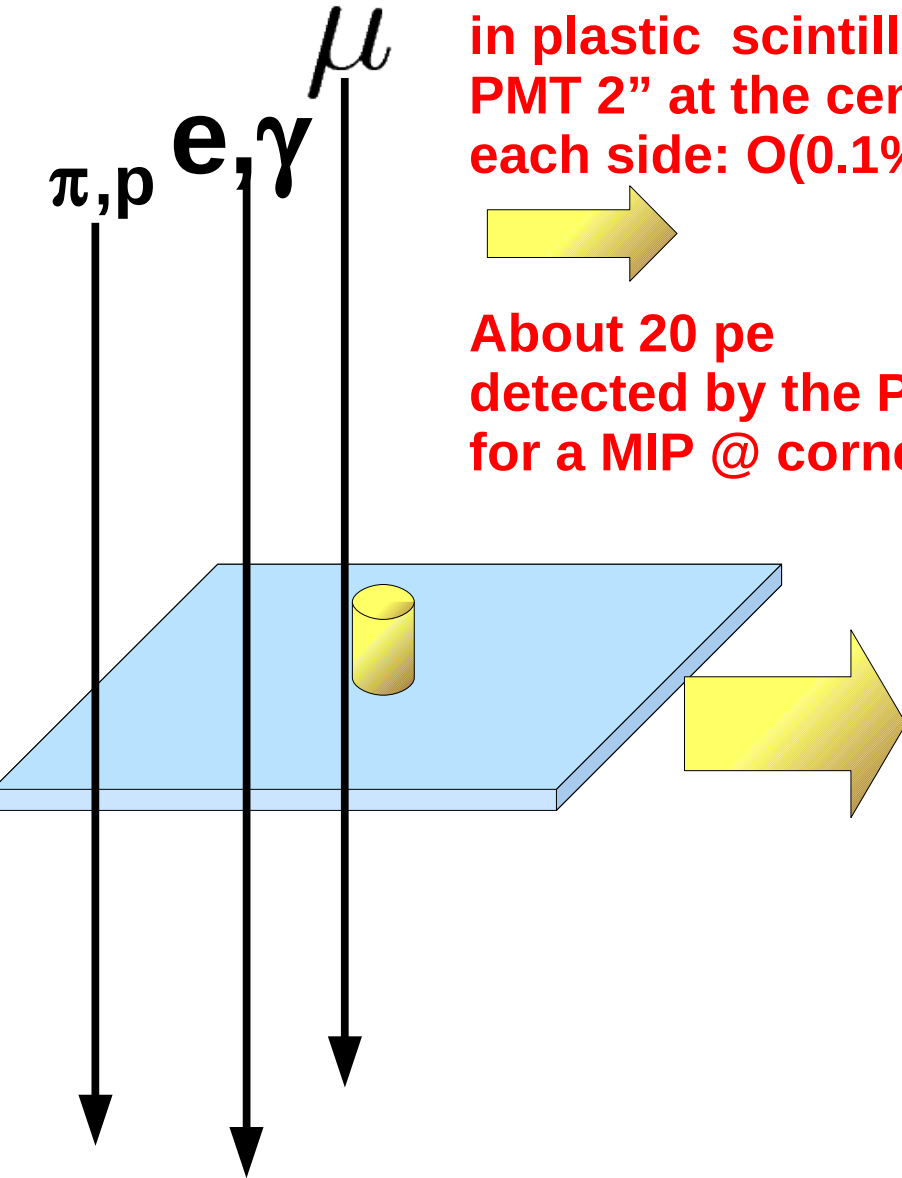
SEA

Required properties of a detection unit :

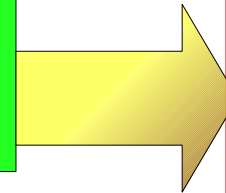
1 MIP = 2.1 MeV ~ 20 000 photons
in plastic scintillator 1cm thick (REXON RP200).
PMT 2" at the center + reflectors on
each side: O(0.1%) detection efficiency



About 20 pe
detected by the PMT
for a MIP @ corner

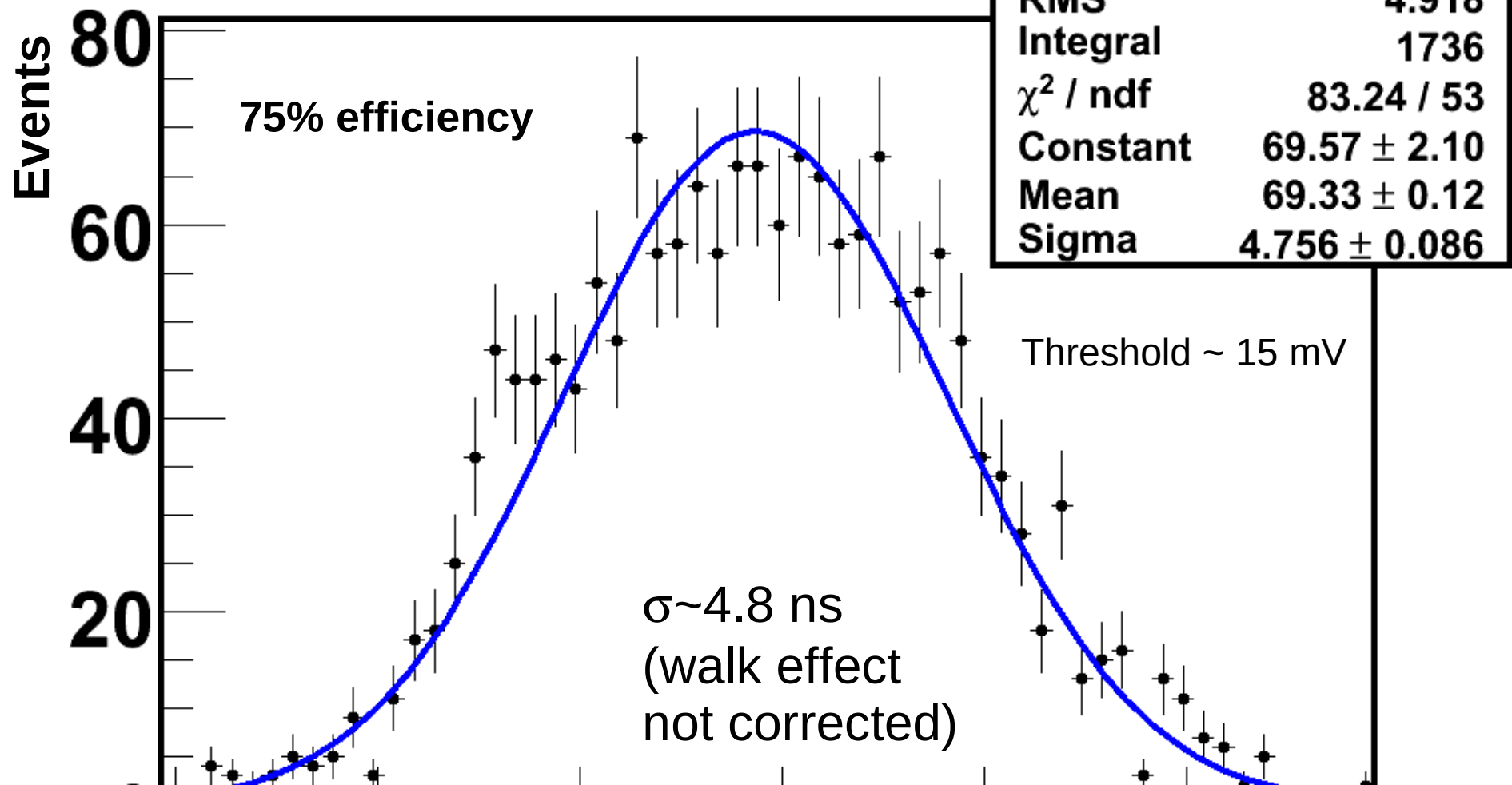
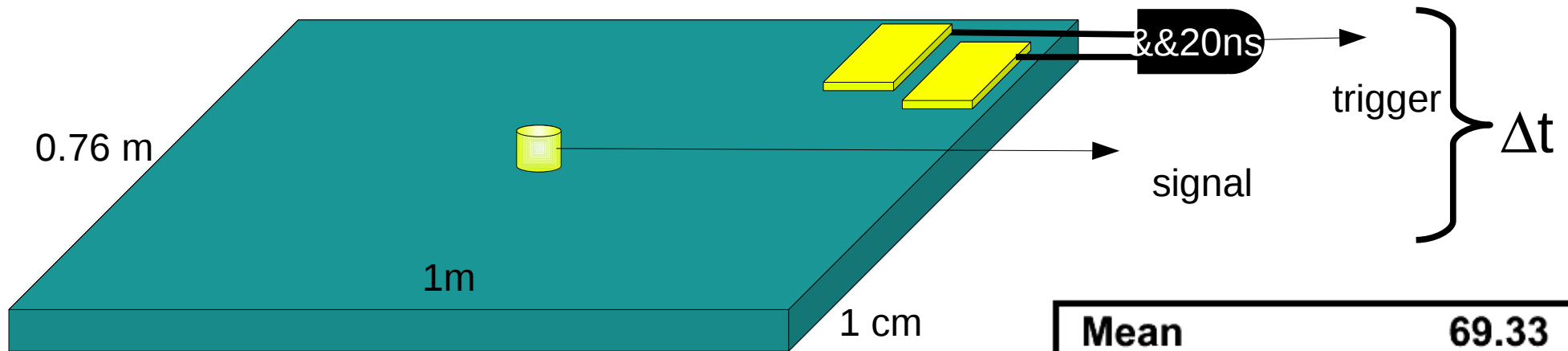


Gain $3 \cdot 10^5$
(Hamamatsu R6231)
threshold ~ 15 mV
for one MIP



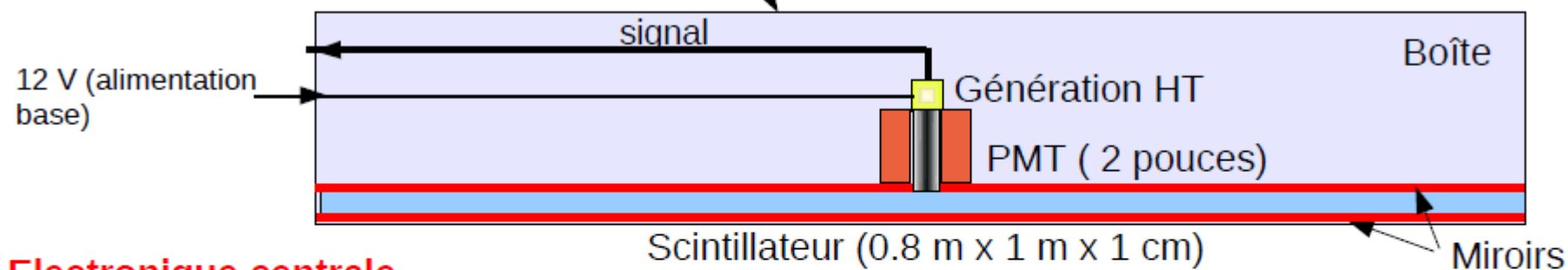
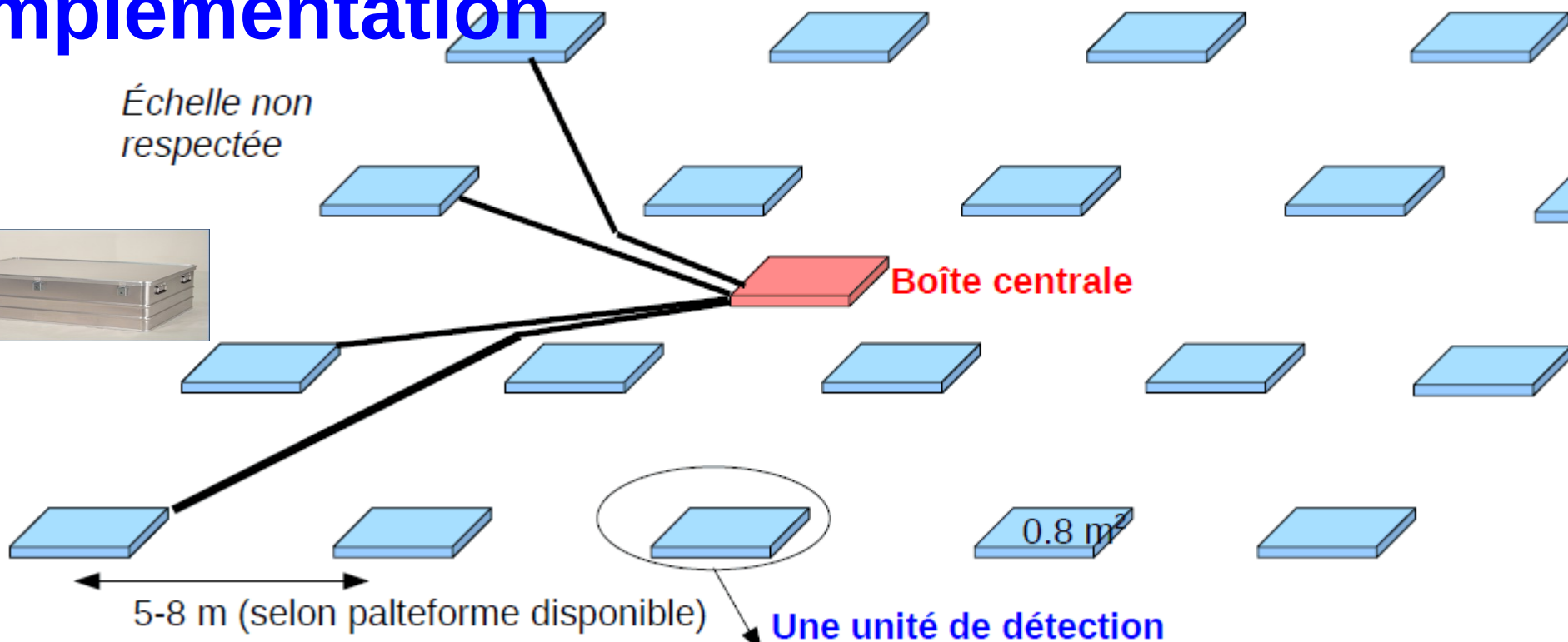
Time resolution
Crossing of the
threshold:
 σ of a few ns

Measurement of time resolution

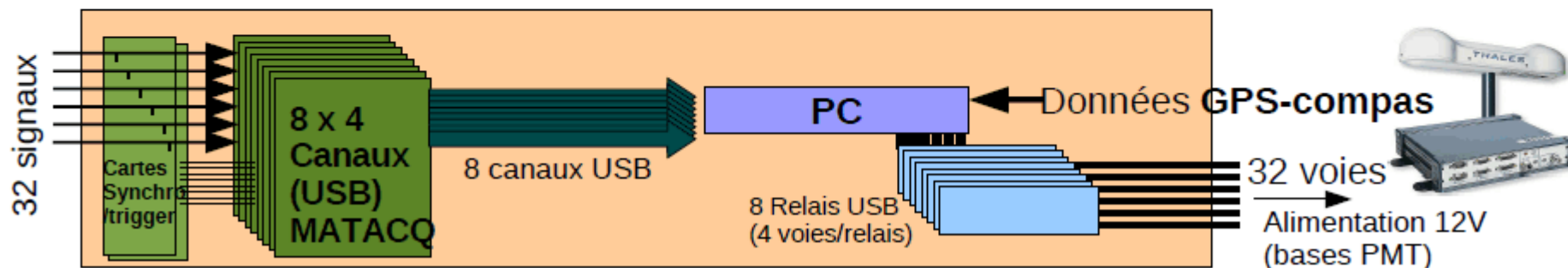


Implementation

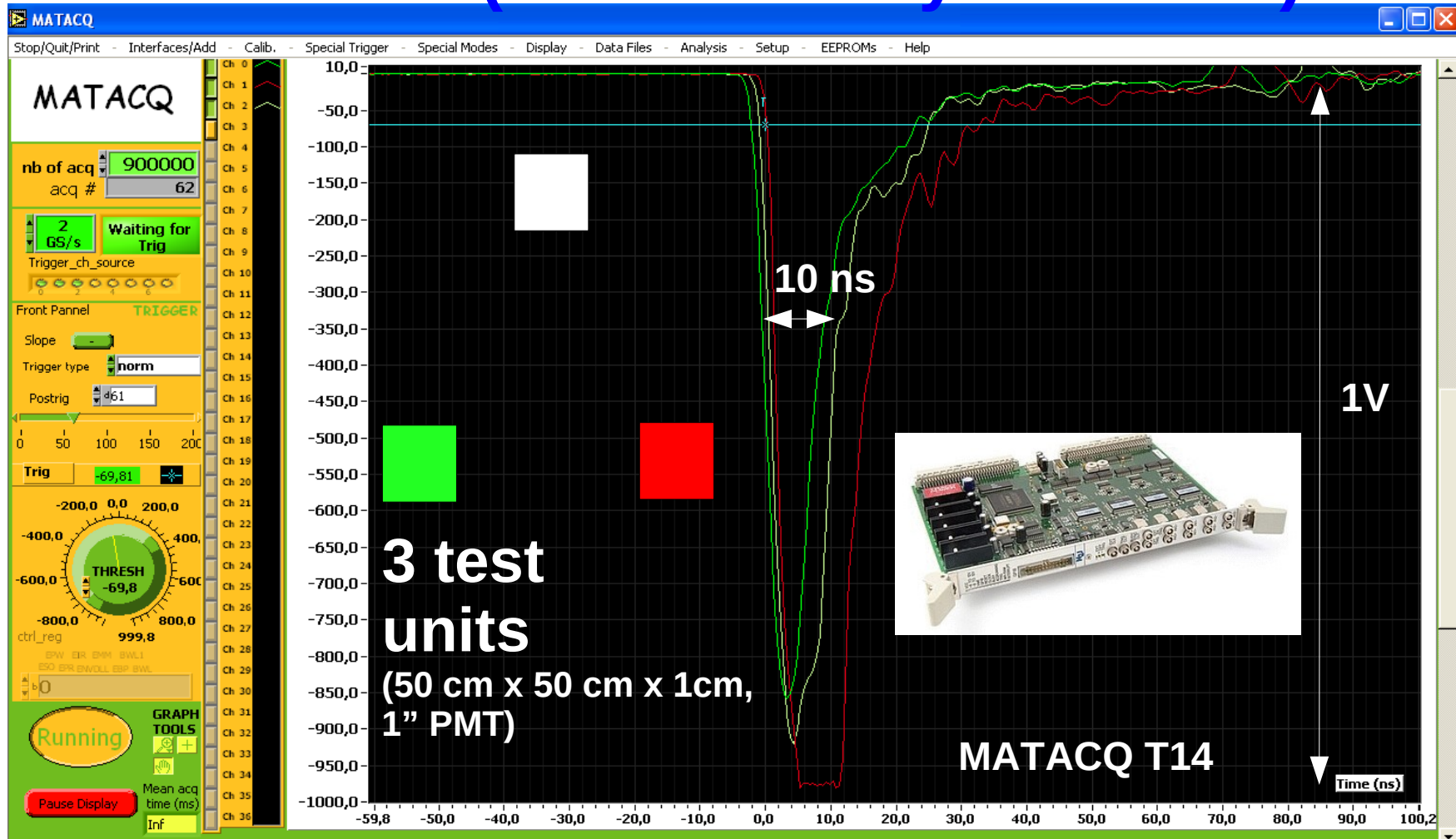
Échelle non respectée



Electronique centrale

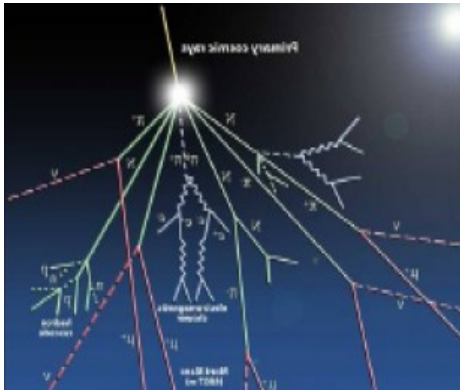


Electronics (commercially available)



We search for some offset in the mean of $\Delta\theta$ and $\Delta\phi$: **simulations of performances**

event by event comparison of :
zenith, azimuth of reconstructed tracks (SA vs ANTARES).
Estimate of the event rate.



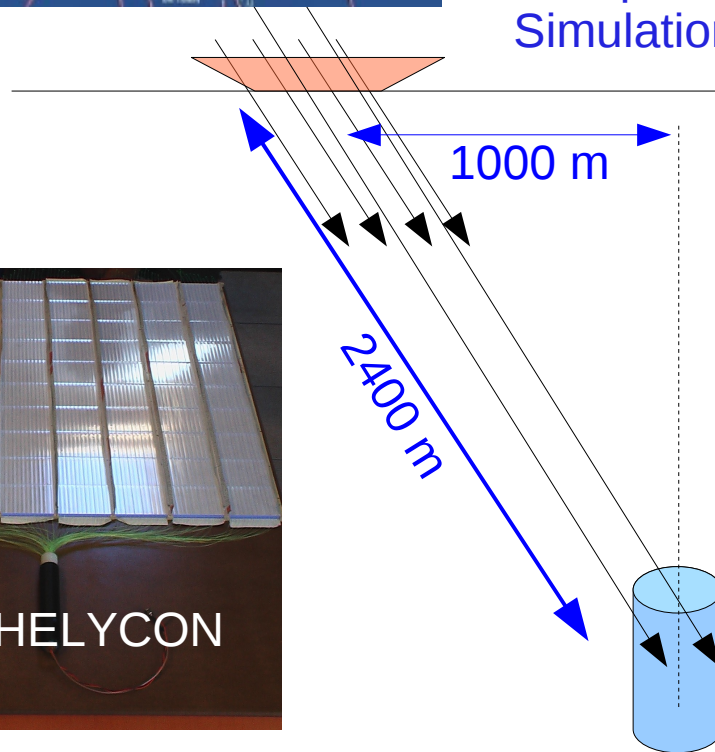
Corsika shower simulation :

protons with E in $[10^5, 5 \cdot 10^6]$ GeV

5 or 8 days equivalent ($55 \cdot 10^4$ showers / day).

Reconstruction based on Δt from a least 3 scintillators requiring 4 MIPs of threshold for each (resolution $\sim 3^\circ$ rms for setups with ~ 10 detectors) : ~ 3 rec evts / min.

Simulation made in coll with HOU (HELYCON detectors)



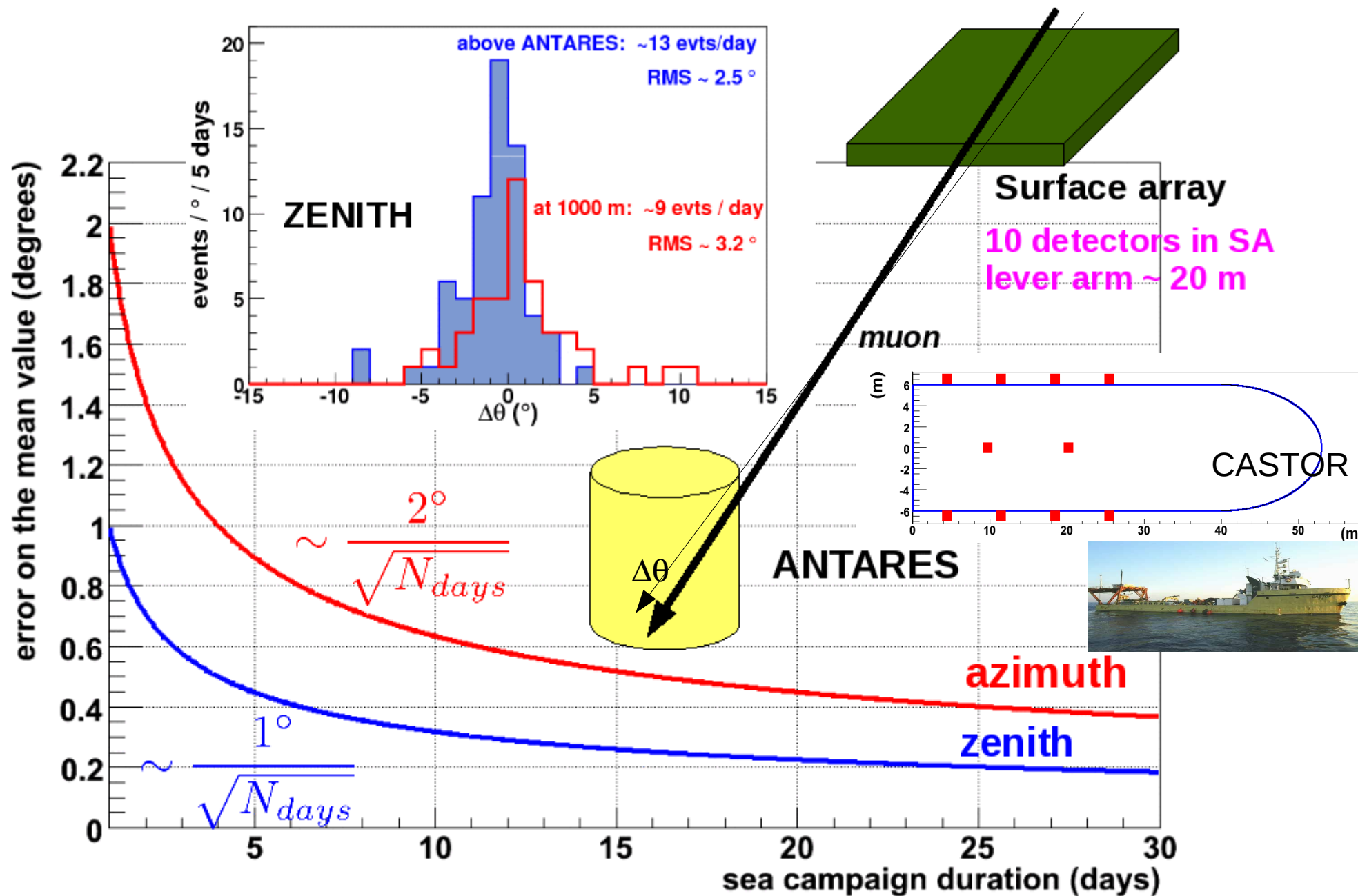
2 tested distances :

just above ANTARES (0m)

1 km apart (1000m, zenith $\sim 24^\circ$)

HELYCON

Results of simulations for 10 detection units



COSTS

3.6 k€ per detection unit including electronics for digitization (MATACQ)

90 k€ for 5 effective days of sea campaign

FULL APPARATUS

32 Detection Units on a platform 35m x 15m or a boat
(limitation: size of the vessel, electronics channels)

VESSELS ...

Boat or platform + boat

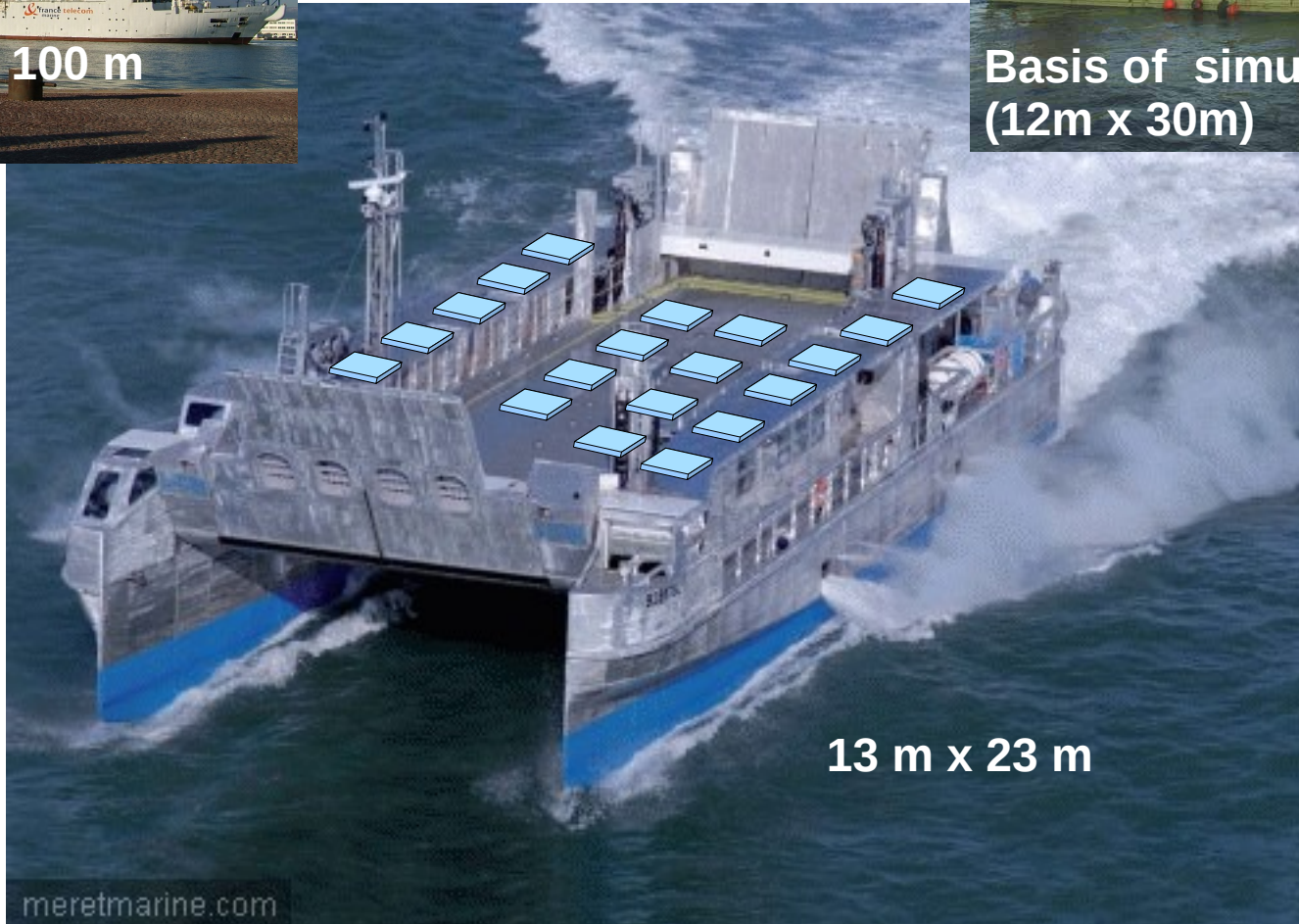
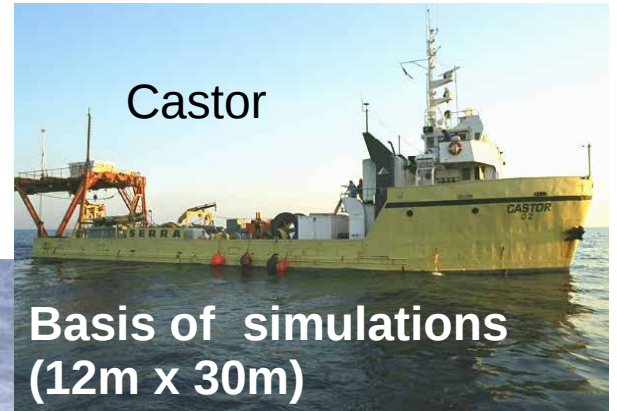
Raymond Croze

Length up to 100 m
Width 20 m



Castor

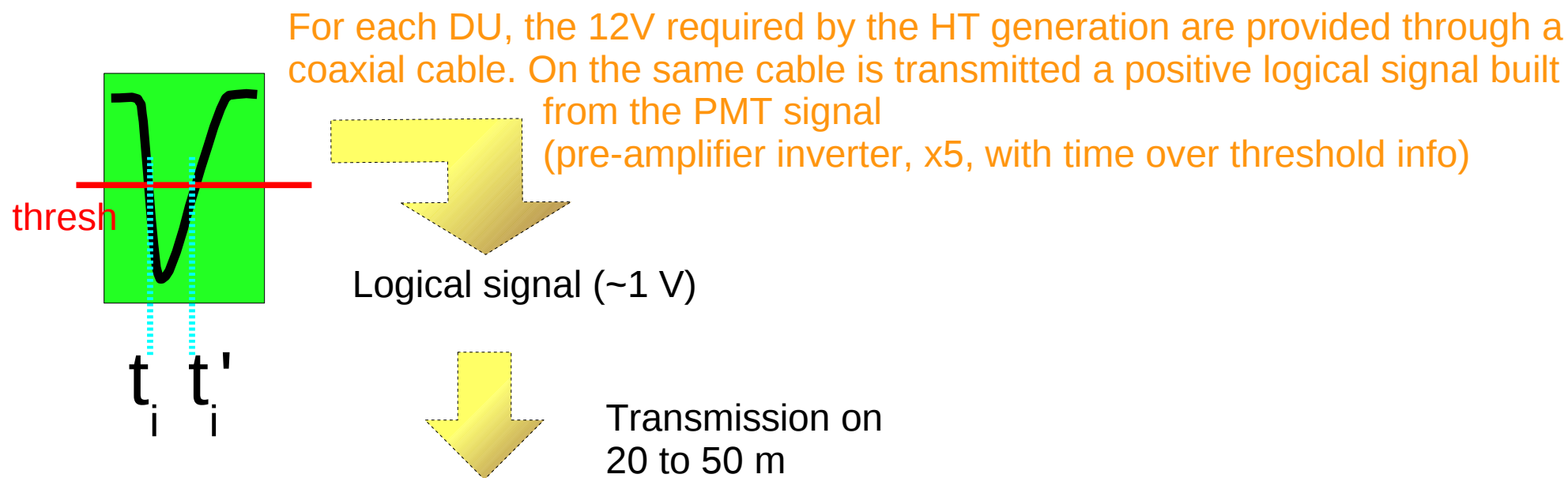
Basis of simulations
(12m x 30m)



13 m x 23 m

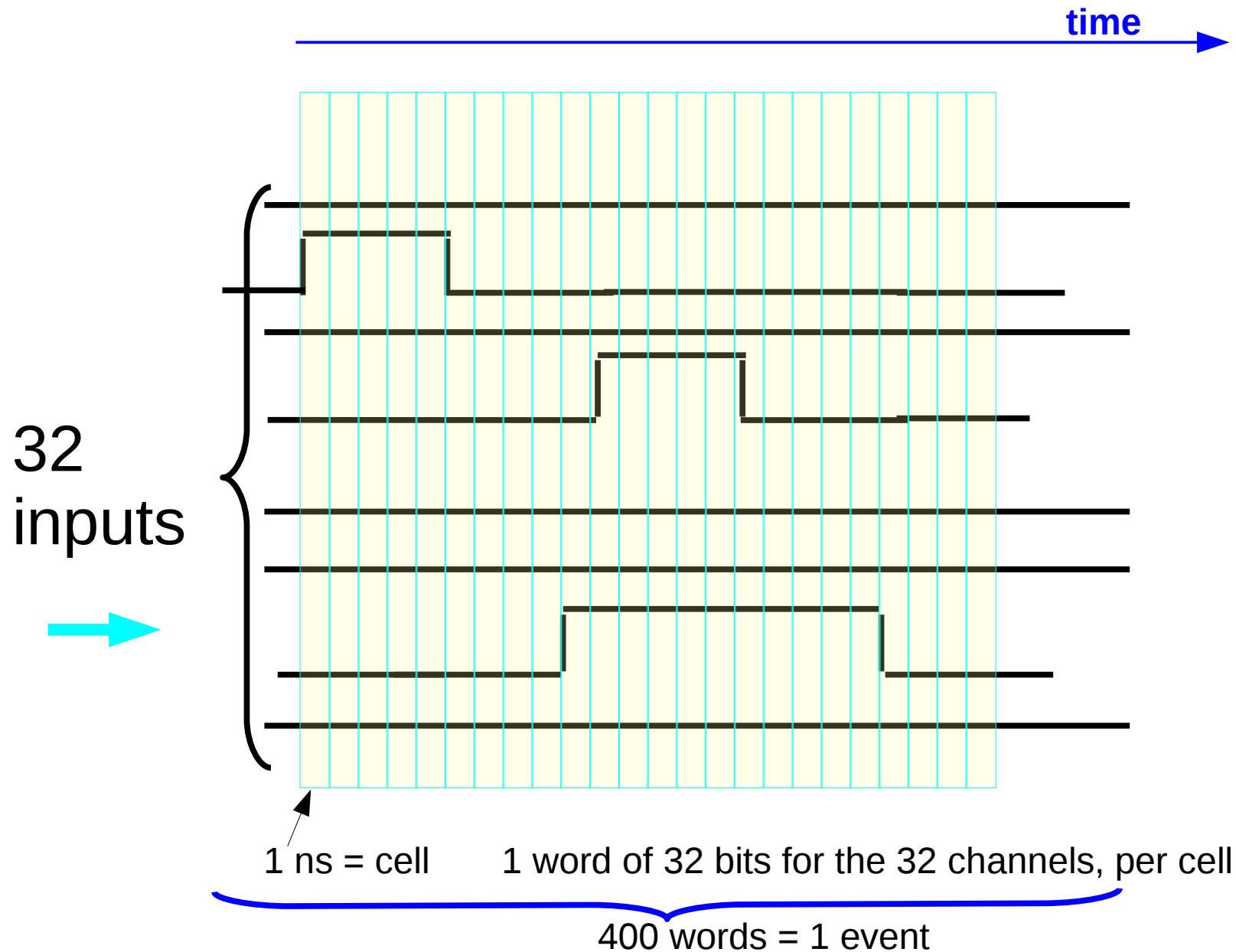
How to make it cheaper: electronics

CPPM home made solution for electronics (under study) :
from 1200€ / channel to 200€/channel :
re-use and adaptation of an existing solution.



All 32 channels connected to a single central card sampling the signals with a 1-2 ns cell.

sampling **How to make it cheaper: electronics**



Size of 1 event :
 $400 * 32 \text{ bits} =$
1600 bytes.

Repetition :
~ 3/minute

Communication
with PC:
TCPIP

Steps of the project :

November 2009: 4 units under test (1 MATACQ card)

January 2010: 8 units (2 MATACQ card,
measurement of angular resolution by splitting detector)

March 2010: decision on the choice *commercial electronics/home made*

Summer 2010: N units ready (depending of the funding) for sea operation

**Achievable constraint in 10 days, 32 units :
0.2° in zenith and 0.4° in azimuth.**