Optimization of gravitational wave optical followup galaxy targeting

Ducoin Jean-Grégoire

In collaboration with :

Corre David Leroy Nicolas TS2020 2019







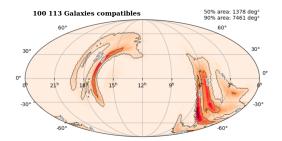


- Galaxies targeting Standard approach
- Quality Control of the Control of
- Conclusion

EM follow-up of GW events

Starting point : GW skymap

Example S190425z:



search of optical counterpart

- Poor 2D localization (up to some 1000 deg²)
- Imprecise distance information (~20% error)

Galaxies targeting - Standard approach

Galaxies targeting

Hypothesis : the source is located within a galaxy

- Choice of the catalog, what we need :
 - ▶ all sky
 - provide distance
 - completeness compatible with LIGO-Virgo range
 - ⇒ GLADE (http://aquarius.elte.hu/glade/)
 Constructed (combined and matched) from four existing galaxy catalogs:
 GWGC, 2MPZ, 2MASS XSC and HyperLEDA. GLADE contains 3,262,883
 objects.
- Selection in the catalogue of compatible galaxies for a certain 3D volume :
 RA, Dec, distance

Galaxies targeting - Standard approach

How do we use the galaxies?

We need to define a grade (weight) to put on each galaxy

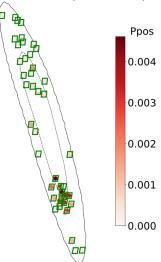
Standard definition of the grade

We use the 3D probability:

$$P_{pos} = P_{dV} = \frac{P_{pixel}}{Pixel\ area}\ N_{pixel}\ e^{-\frac{1}{2}\left(\frac{D_{galaxy} - \mu_{pixel}}{\sigma_{pixel}}\right)^2}$$

Where μ_{pixel} , σ_{pixel} and N_{pixel} are respectively the mean distance, the standard deviation and the normalization factor of the Gaussian distribution at the given pixel. D_{galaxy} is the galaxy distance fetch from the catalog.

Example of tiles obtained, GW170817, FOV = $20' \times 20'$



- Galaxies targeting Standard approach
- Q Galaxies targeting Adding galaxies properties
 - Previous works
 - Our works
 - Results
- 3 Conclusion

Upgrading the grade

More complete definition of the grade

Only information available on GLADE

- B,J,H,K Luminosity (not for all galaxies)
 - ⇒ sufficient to deduce interesting properties from it?

LCOGT grade \Rightarrow stellar mass

Use the B luminosity as an "indicator of mass"

We then calculate the B-band luminosity of the galaxy, L_B (based on the B-band magnitude and distance provided in the GLADE catalog), and assign it a score

$$S_{\text{lum}} = \frac{L_B}{\sum L_B},\tag{3}$$

(Arcavi et al. 2017)



The B band is highly sensitive to the galaxy dust attenuation

- ⇒ We should use near infrared band
- ⇒ K band is provided by GLADE but :
 - K band is still a bit affected by the dust extinction
 - only ~67% of the galaxies in the catalog (up to 400Mpc) have K band information
- \Rightarrow Utilization of the WISE1 band (3.4 μ m)

Our works

Cross-match AllWISE and GLADE (400Mpc):

After all treatment we have $\sim 90\%$ of the galaxies with WISE1 band

Determination of the stellar mass

From WISE1 band we can determine the stellar mass using a constant mass to light ratio : (Kettlety et al. 2017)

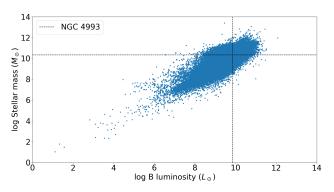
$$\Upsilon_*^{3.4\mu m} \sim 0.60 M_{\odot}/L_{\odot,3.4\mu m}$$

Adding a factor to the grade

As done previously we can now change the grade adding :

$$\mathsf{P}_{mass} = \frac{M_{*,galaxy}}{\sum M_{*,galaxy}} \Rightarrow P_{tot} = P_{pos} \times P_{mass}$$

Comparison with LCOGT method :



A NGC4993 like galaxy ($\sim 7.2 \times 10^9 L_\odot$ B band luminosity) can have a stellar mass which can span from $\sim 3.8 \times 10^7 M_\odot$ to $\sim 1.0 \times 10^{12} M_\odot$

- ⇒ Our grade is going to behave very differently from one using B band luminosity
 ⇒ B band luminosity is a very poor indicator of the stellar mass (assuming our
- ⇒ B band luminosity is a very poor indicator of the stellar mass (assuming our determination of stellar mass is ok)

Reformulation

$$P_{tot} = P_{pos} \times P_{mass}$$

Huge drawback of the product expression :

Can't define P_{mass} when you don't have the stellar mass info (= the W1 mag)

 \Rightarrow forced to throw away $\sim 10\%$ of the catalog

We chose to reformulate the grade :

$$P_{tot} = P_{pos} \left(1 + \alpha P_{mass} \right)$$

Reformulation

$$P_{tot} = P_{pos} \left(1 + \alpha P_{mass} \right)$$

whit α that ensure the two factor in the addition are, in mean, contributing as much :

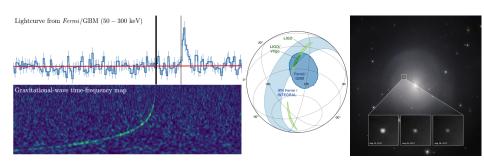
$$\frac{\sum P_{pos}}{N} = \frac{\sum P_{pos} \,\alpha \, P_{mass}}{N}$$

$$\Rightarrow \alpha = \frac{\sum P_{pos}}{\sum P_{pos} P_{mass}}$$

 \Rightarrow Put $P_{mass} = 0$ to fall back on P_{pos}

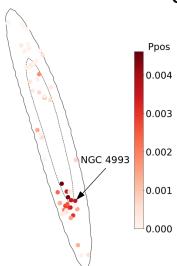
GW170817

Only counterpart for a GW found at the moment



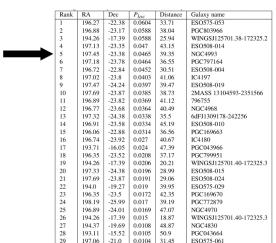
⇒ Mandatory to test our grade on it

GW170817



- 90% skymap $\sim 30 deg^2$
- distance 40 ± 8 Mpc
- 65 galaxies compatibles

With the standard 3D localization \Rightarrow NGC 4993 ranked 5



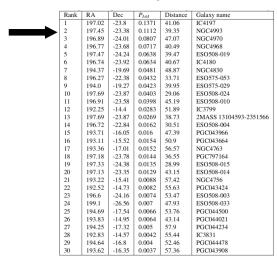
30

193.84 -17.1

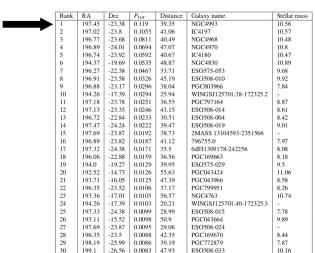
PGC044023

0.0092 53.88

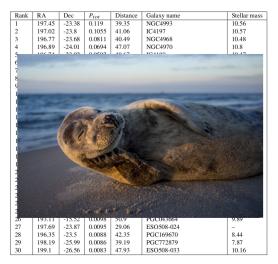
With the B luminosity \Rightarrow NGC 4993 ranked 2



With the stellar mass addition \Rightarrow NGC 4993 ranked 1



With the stellar mass addition \Rightarrow NGC 4993 ranked 1



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Conclusion

A brand new galaxy selection!



- This method is already used by Kilonova-Catcher and will be then by GRANDMA and SVOM
- Paper about this development is on writing process
- All the tools will be publicly available via dedicated web-site

MERCI!

Backup slide

$$P_{tot} = P_{pos} (1 + \alpha \beta P_{mass})$$

- $\Rightarrow \beta$ which will determine at which point the mass factor will count in the grade
- $\Rightarrow \beta$ is skymap independent
- \Rightarrow β should be to fit with a statistically significant sample of gravitational wave host galaxies, but as we don't have such sample yet \Rightarrow β = 1