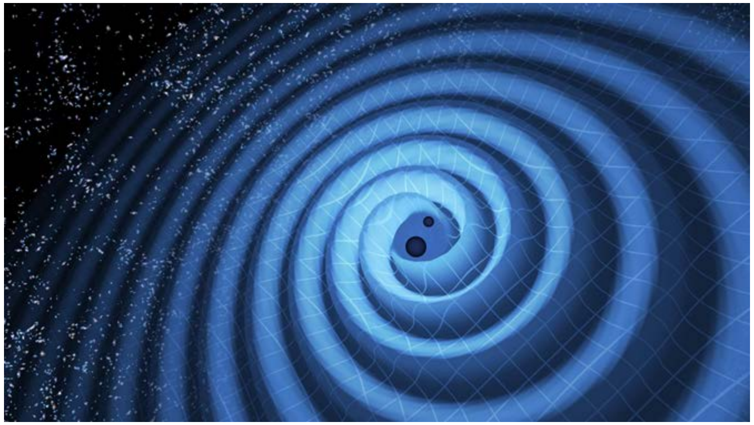


Gravitational Waves:

The challenges of the detection

Romain Gouaty
LAPP – Annecy
GraSPA summer school

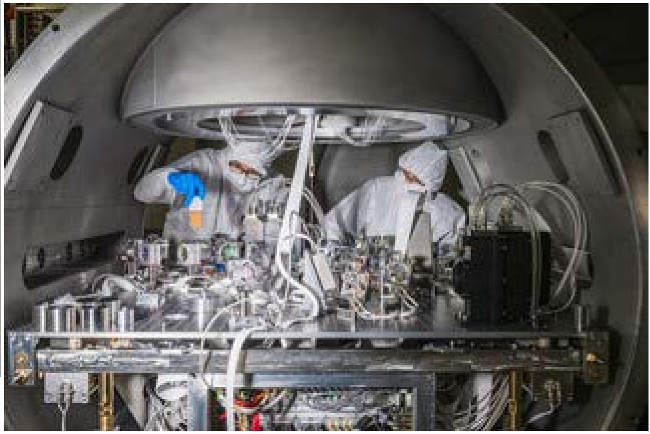


Table of Contents

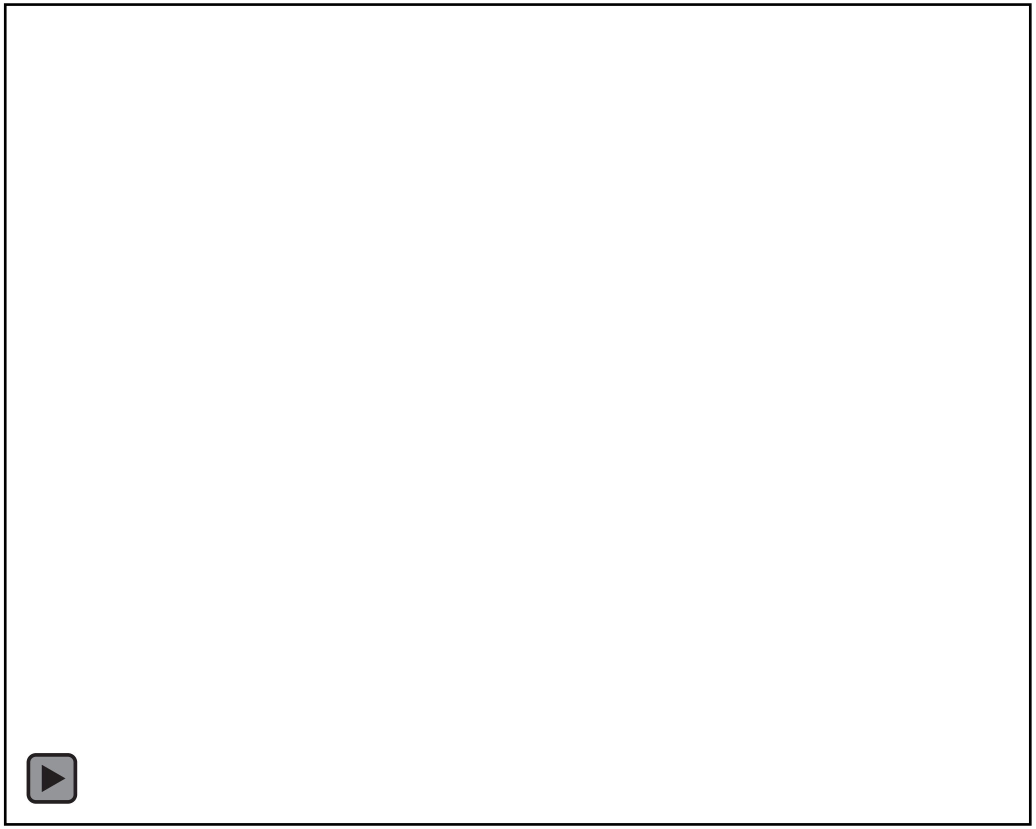
GW quest: a bit of history

- Joseph Weber invents the bar detector
  - The GW changes the resonance condition of a resonant bar of a few tons
  - First claim for detection in 1969… but contested
  - Triggered large interest, at least 18 bars in 8 countries
- Evolve to cryogenic resonant bars (80-90)
- Bar not enough sensitivity:
  - h : few 10^{-21} 1/sqrt(Hz) @ 900Hz

- ITF started in the 70's (Germany, Rai Weiss)
  - Broad band instrument
- Few ITF prototypes in the 80's
  - MIT, Glasgow, Garching, Caltech,...
  - ~10m long
  - Not made for detection
- Jump to km scale in early 90
  - LIGO, Virgo, GEO, TAMA
- **First detection in 2015 (GW150914)**

Reminder: effect of a GW on free fall masses

A gravitational wave (GW) modifies the distance between free-fall masses

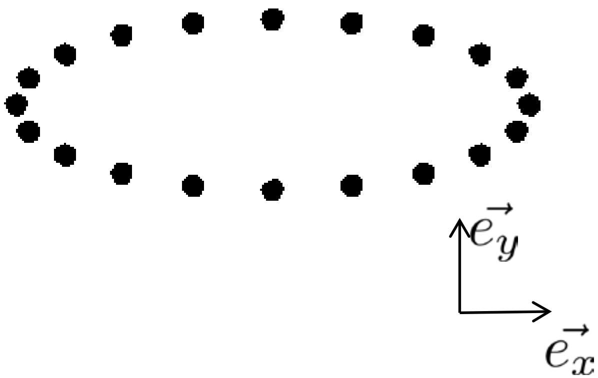


Reminder: effect of a GW on free fall masses

A gravitational wave (GW) modifies the distance between free-fall masses

$$\delta x(t) \ = \ -\delta y(t) \ = \ \frac{1}{2}\, h(t)\, L_0$$

h(t): amplitude of the GW (= strain)

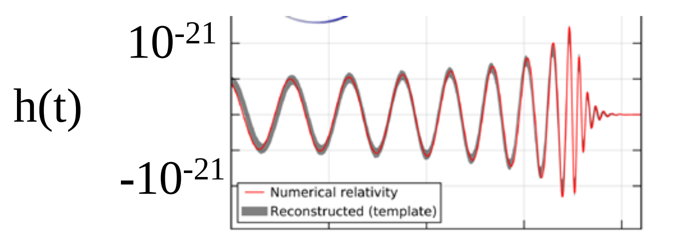


Typical amplitude of a GW crossing the Earth:
$h \sim 10^{-23}$

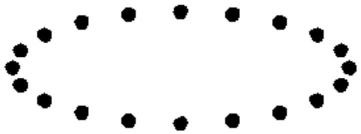
(h has no dimension/unit)

Case of a GW with polarisation + propagating along z

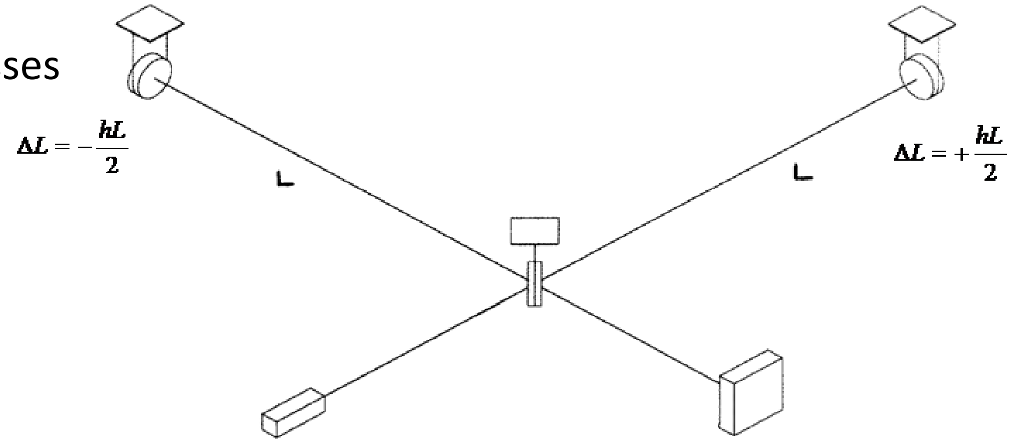
Reconstructed strain of GW150914



Terrestrial GW Interferometer: basic principle



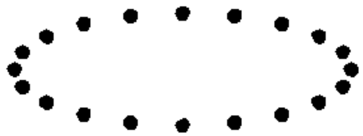
- Measure a variation of distance between masses
  - Measure the light travel time to propagate over this distance
  - Laser interferometry is an appropriate technique
    - Comparative measurement
    - Suspended mirrors = free fall test masses

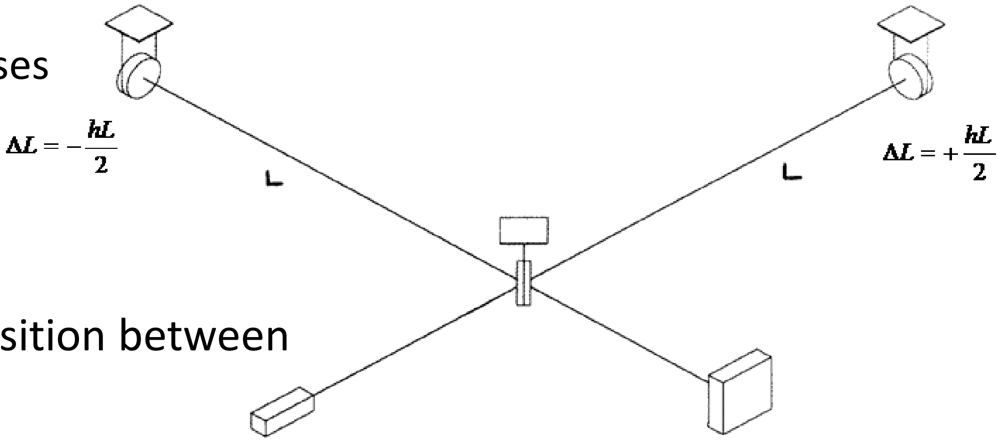


Terrestrial GW Interferometer: basic principle

Terrestrial GW Interferometer: basic principle

- Measure a variation of distance between masses
  - Measure the light travel time to propagate over this distance
  - Laser interferometry is an appropriate technique
    - Comparative measurement
    - Suspended mirrors = free fall test masses





- Michelson interferometer well suited:
  - Effect of a gravitational wave is in opposition between 2 perpendicular axes
  - **Light intensity of interfering beams is related to the difference of optical path length in the 2 arms**

Bandwidth: 10 Hz to few kHz

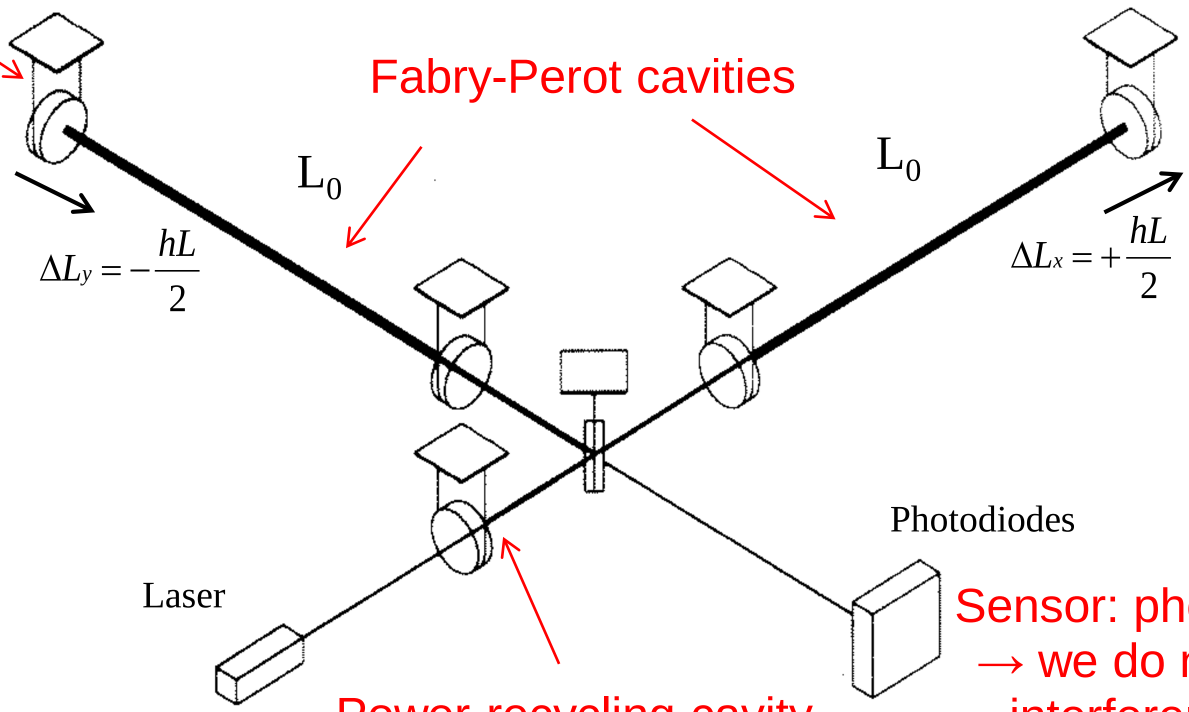
We need a big interferometer:

ΔL proportional to L

➔ need several km arms!

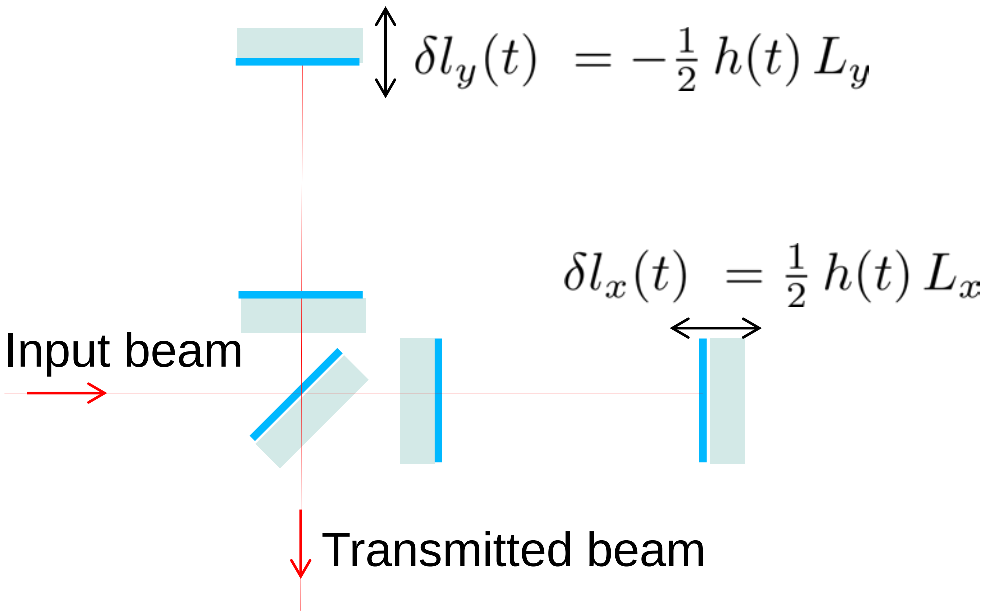
Virgo/LIGO: more complicated interferometers

Suspended mirrors → Mirrors can be considered as free-falling in the ITF plane for frequencies larger than ~10 Hz



WARNING: STILL VERY SIMPLIFIED SCHEME!

Orders of magnitude



Typical amplitude of differential arm length variations when a GW crosses the Earth:

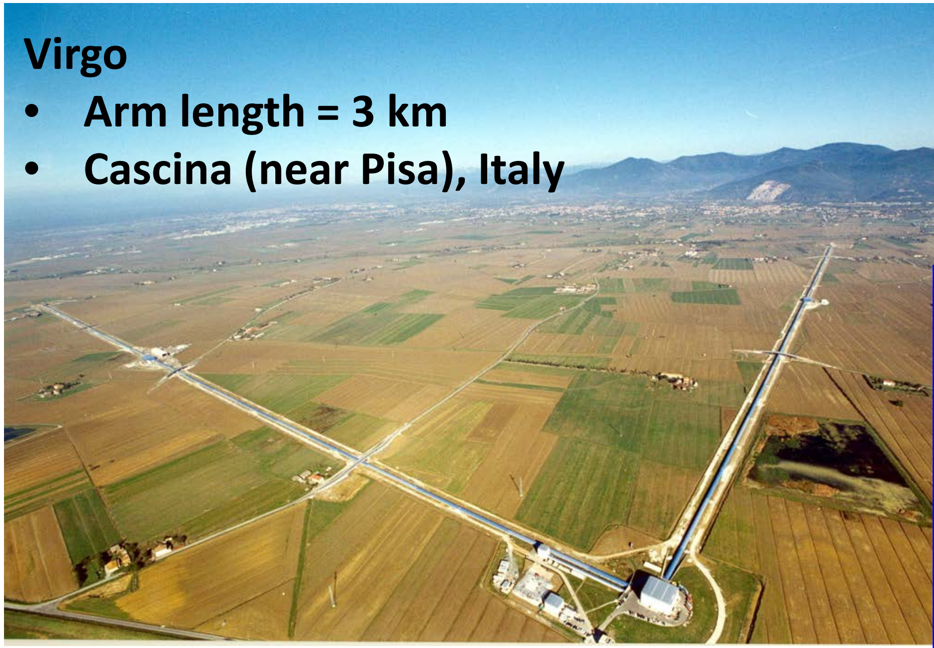
$$\delta \Delta L = \delta l_x(t) - \delta l_y(t)$$
$$= h(t)\, L_0$$

$$\mathrm{h} \sim 10^{-23} \qquad L_0 = 3\,\mathrm{km}$$

$$\rightarrow \delta \Delta L \sim 3 \times 10^{-20}\,\mathrm{m}$$

$$\sim \frac{\text{size of a proton}}{100000}$$

Km scale interferometers

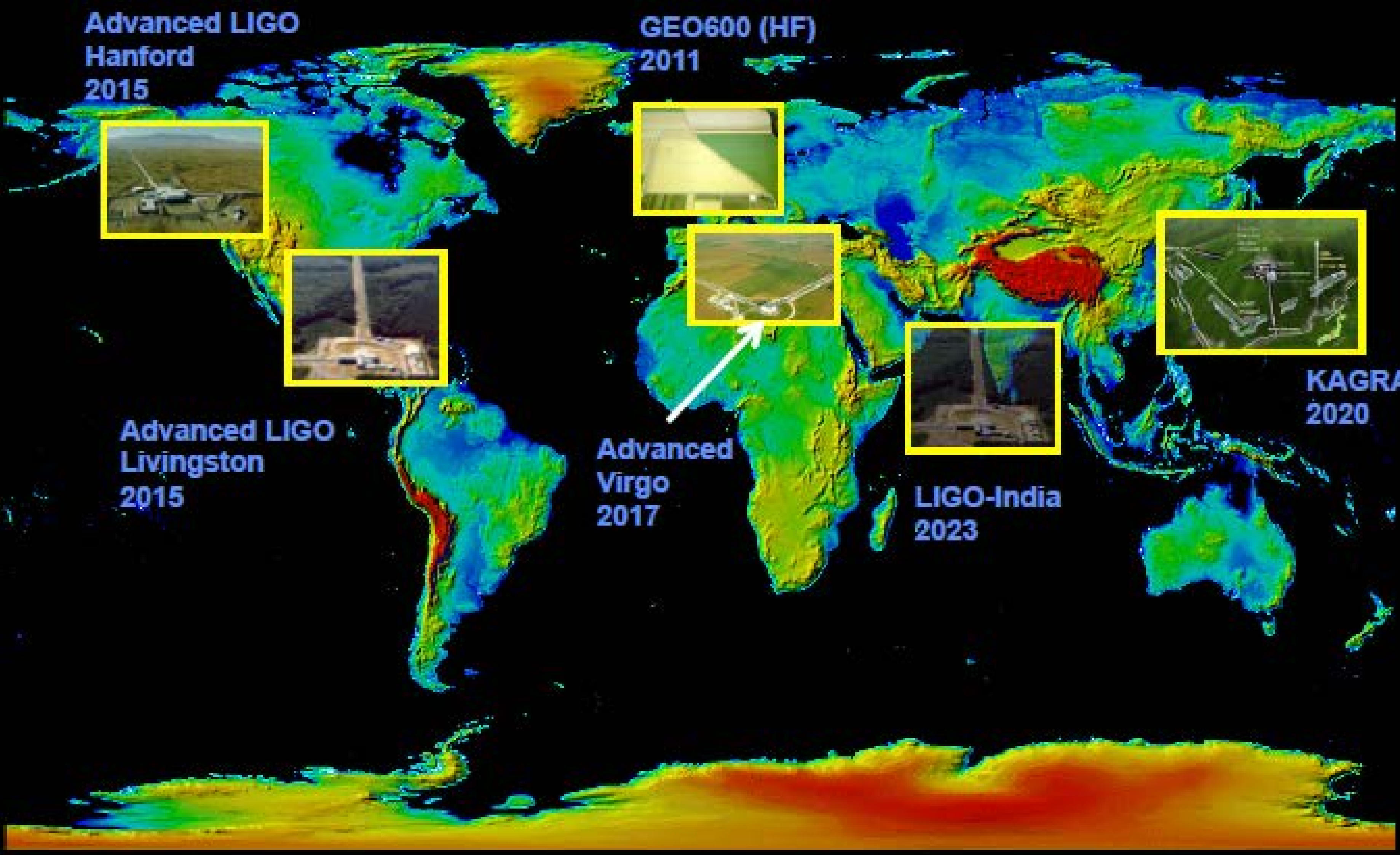


Virgo
- **Arm length = 3 km**
- **Cascina (near Pisa), Italy**

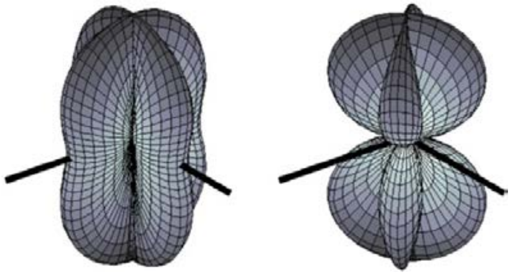
LIGO Livingston
- **Arm length = 4 km**
- **Louisiana**

LIGO Hanford
- **Arm length = 4 km**
- **Washington State**

The detector network



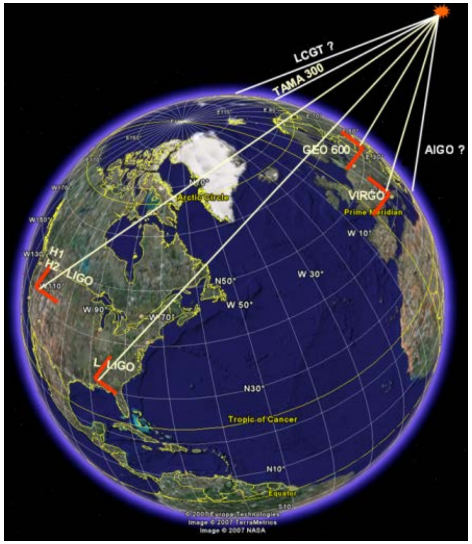
The benefits of the network



- A GW interferometer acts as a wide beam antenna
  - A single detector cannot localize the source
  - Need to compare the signals found in coincidence between several detectors (triangulation):

    → **allow to point towards the source position in the sky**

    → **the telescope is obtained by a network of interferometers**



- Looking for rare and transient signals: can be hidden in detector noise

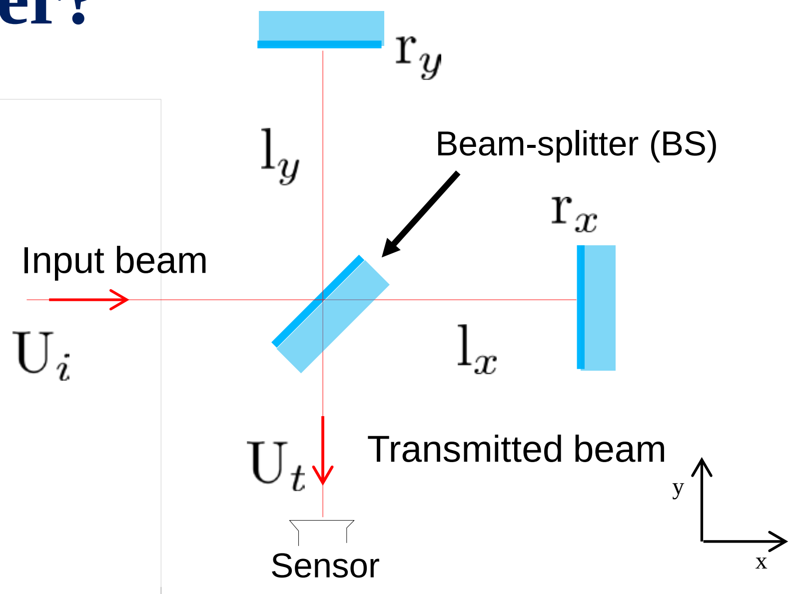
  → requires observation in coincidence between at least 2 detectors
- Since 2007, Virgo and LIGO share their data and analyze them jointly

Table of Contents

How do we « observe » ΔL with a Michelson interferometer?

- Input wave $U_i(x,t) = \underline{A}_i e^{jkx}$
  $= \underline{A}_i \quad \text{on BS}$
- BS located at (0,0)
- Sensor located at $(0,-y_s)$
- Amplitude reflection and transmission coefficients: r and t

→ We are interested in the beam transmitted by the interferometer: it is the sum of the two beams (fields) that have propagated along each arm



Around the mirrors:

- Radius of curvature of the beam ~ 1400 m
- Size of the beam ~ few cm

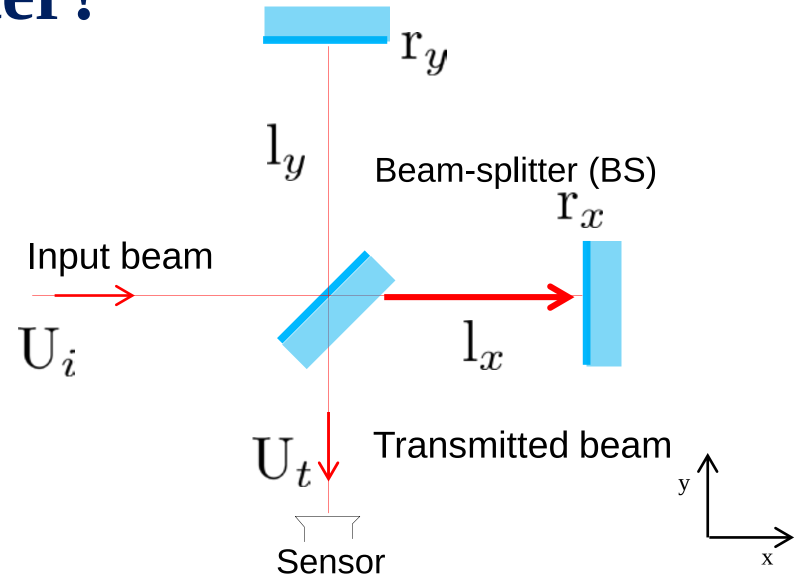
→ The beam can be approximated by plane waves

How do we « observe » ΔL with a Michelson interferometer?

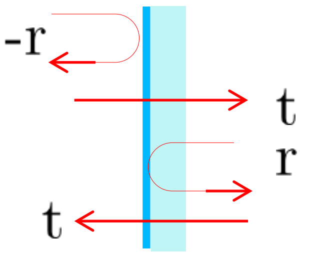
- Input wave $U_i(x,t) = \underline{A}_i e^{\jmath kx}$
  $= \underline{A}_i \quad \text{on BS}$

- Beam propagating along x-arm:

$$U_{tx} = \underline{A}_i \, t_{BS} \, e^{\jmath k l_x} \; ......$$



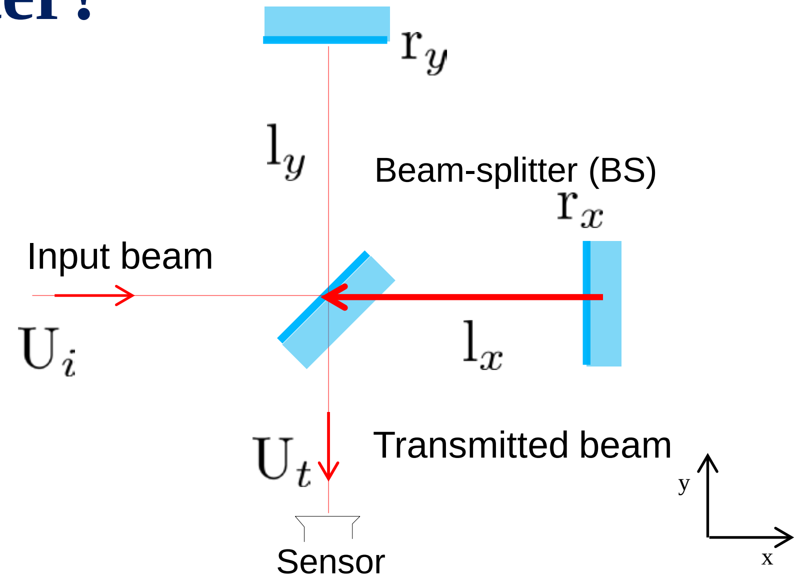
Sign convention for amplitude reflection and transmission coefficients



Without losses:
$t^2 + r^2 = 1$

How do we « observe » ΔL with a Michelson interferometer?

- Input wave $U_i(x,t) = \underline{A_i} e^{jkx}$
  $= \underline{A_i}$ on BS

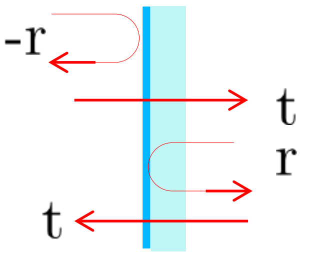


- Beam propagating along x-arm:

$$U_{tx} = \underline{A_i}\, t_{BS}\, e^{jkl_x} \quad (-r_x) e^{jkl_x} \ldots\ldots$$

Sign convention for amplitude reflection and transmission coefficients

Without losses:
$t^2 + r^2 = 1$



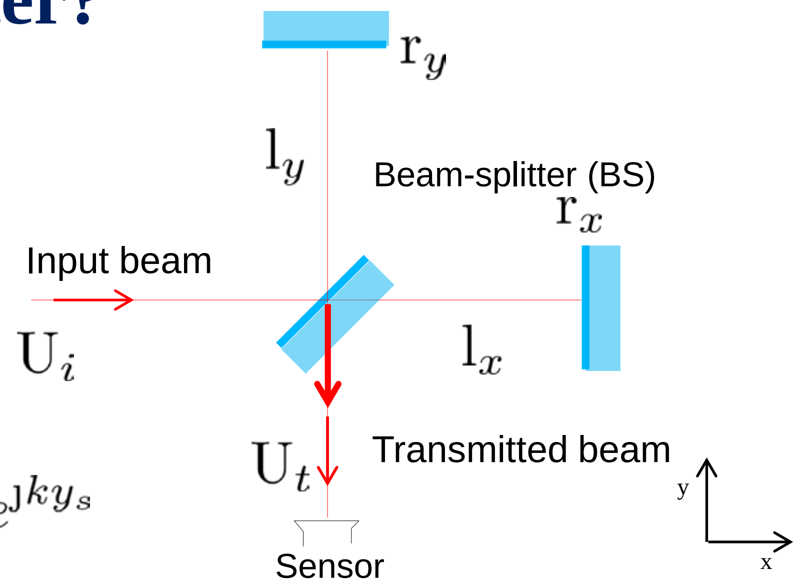
How do we « observe » ΔL with a Michelson interferometer?

- Input wave $U_i(x,t) = \underline{\mathcal{A}}_i e^{\jmath kx}$

  $= \underline{\mathcal{A}}_i \quad \text{on BS}$

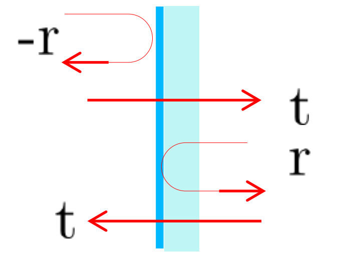
- Beam propagating along x-arm:

$$U_{tx} = \underline{\mathcal{A}}_i \, t_{BS} \, e^{\jmath k l_x} \quad (-r_x) e^{\jmath k l_x} \quad r_{BS} \, e^{\jmath k y_s}$$

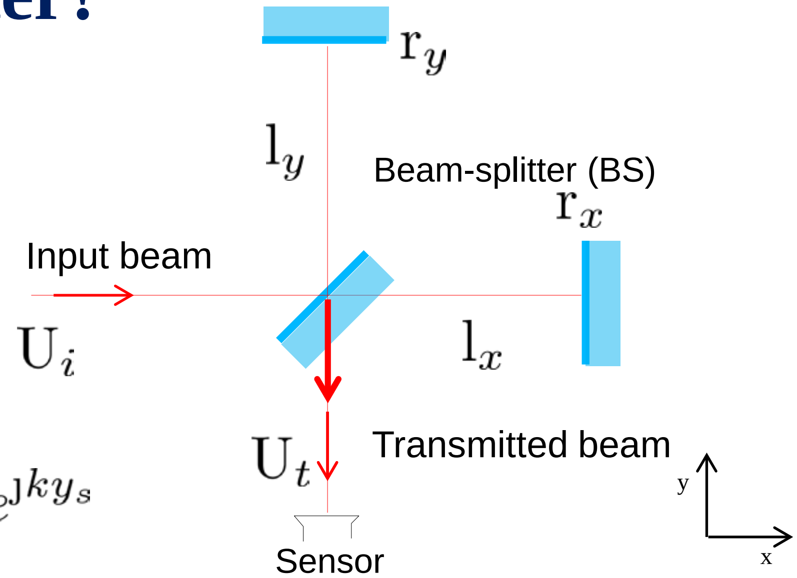


Sign convention for amplitude reflection and transmission coefficients

Without losses:
$t^2 + r^2 = 1$



How do we « observe » ΔL with a Michelson interferometer?



- Input wave $U_i(x,t) = \underline{A}_i e^{\jmath kx}$

$$= \underline{A}_i \quad \text{on BS}$$

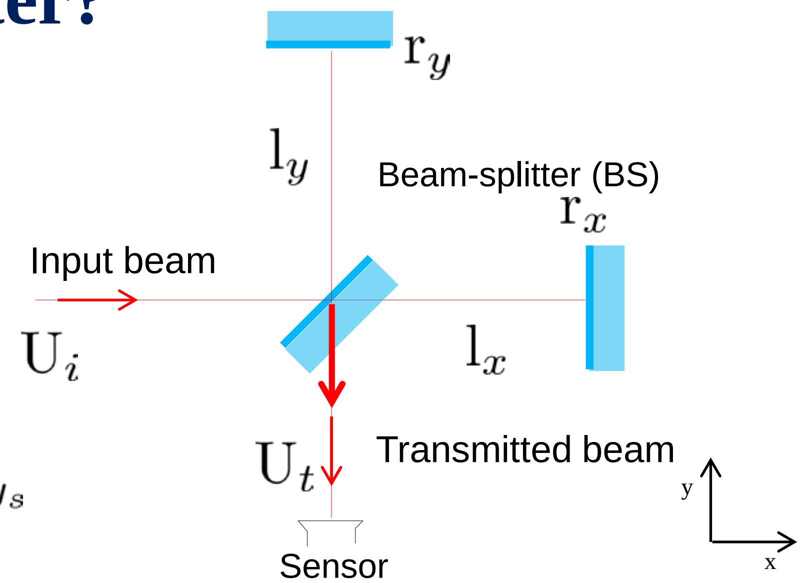
- Beam propagating along x-arm:

$$U_{tx} = \underline{A}_i \, t_{BS} \, e^{\jmath k l_x} \quad (-r_x) e^{\jmath k l_x} \quad r_{BS} \, e^{\jmath k y_s}$$

$$= \underline{A}_i \, t_{BS} \, r_{BS} \, (-r_x) \, e^{2\jmath k l_x} \, e^{\jmath k y_s}$$

$$= \frac{\underline{A}_i}{2} \times \underbrace{\left(-\, r_x \, e^{2\jmath k l_x} \right)}_{\text{Complex reflection of the x-arm}} e^{\jmath k y_s} \qquad \text{with } t_{BS} = r_{BS} = \frac{1}{\sqrt{2}}$$

How do we « observe » ΔL with a Michelson interferometer?



- Input wave

$$U_i(x,t) = \underline{A}_i e^{\jmath kx}$$
$$= \underline{A}_i \quad \text{on BS}$$

- Beam propagating along x-arm:

$$U_{tx} = \underline{A}_i\, t_{BS}\, e^{\jmath k l_x} \quad (-r_x) e^{\jmath k l_x} \quad r_{BS}\, e^{\jmath k y_s}$$
$$= \underline{A}_i\, t_{BS}\, r_{BS}\, (-r_x)\, e^{2\jmath k l_x}\, e^{\jmath k y_s}$$
$$= \frac{\underline{\underline{A}}_i}{2} \times \underbrace{\left(-r_x\, e^{2\jmath k l_x}\right)}_{\text{Complex reflection of the x-arm}} e^{\jmath k y_s}$$

- Beam propagating along y-arm:

$$U_{ty} = -\frac{\underline{\underline{A}}_i}{2} \times \underbrace{\left(-r_y\, e^{2\jmath k l_y}\right)}_{\text{Complex reflection of the y-arm}} e^{\jmath k y_s}$$

Transmitted field:

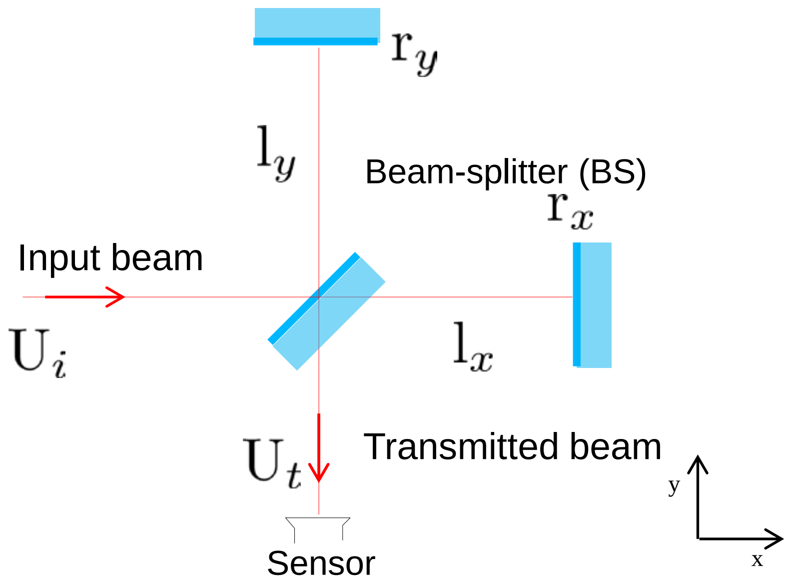
$$U_t = U_{tx} + U_{ty}$$
$$= \frac{\underline{\underline{A}}_i}{2} e^{\jmath k y_s} \left(r_y\, e^{2\jmath k l_y} - r_x\, e^{2\jmath k l_x}\right)$$

Simple Michelson interferometer: transmitted power

Field transmitted by the interferometer

$$U_t = \frac{\mathcal{A}_i}{2}\left(r_y\, e^{2\jmath k l_y} - r_x\, e^{2\jmath k l_x}\right)$$

k is the wave number, k = 2π/λ
λ is the laser wavelength (λ=1064 nm)



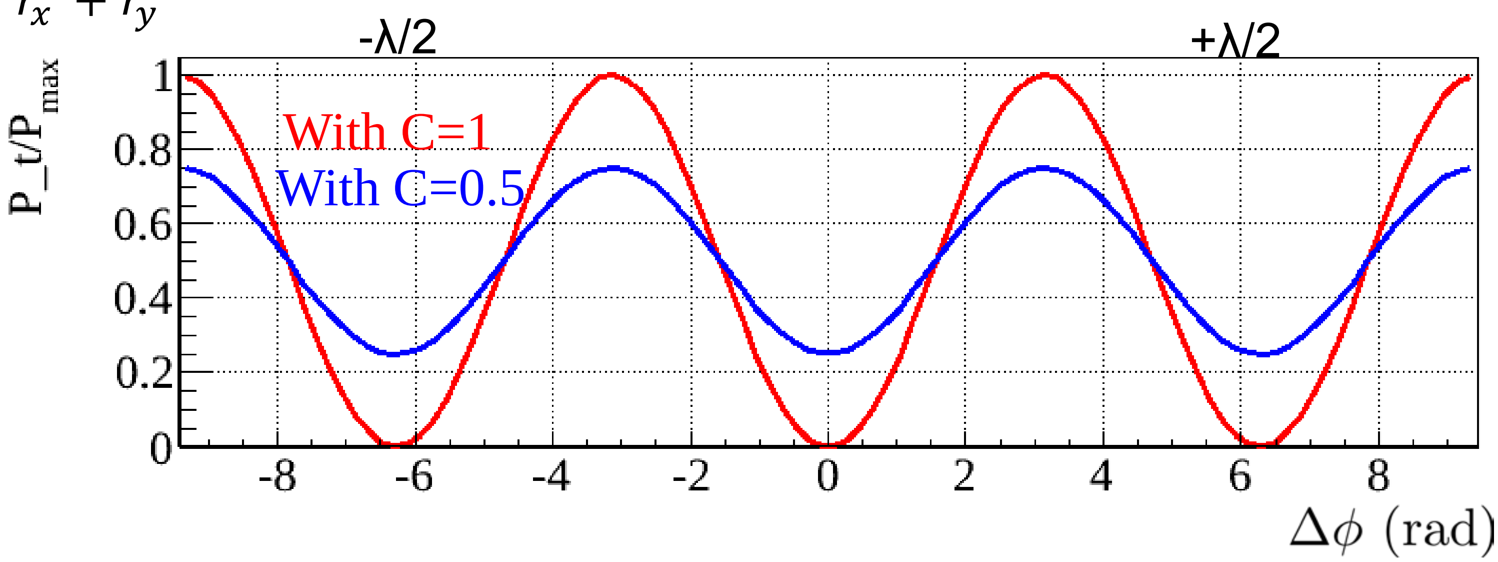
Transmitted power

$$P_t \propto |U_t|^2 = \frac{P_{max}}{2}\left(1 - C\cos(\Delta\phi)\right)$$

where $\Delta\phi = 2k(l_y - l_x)$

ITF contrast: $C = \dfrac{2 r_x r_y}{r_x^2 + r_y^2}$

$$P_{max} = \frac{P_i}{2}(r_x^2 + r_y^2)$$



What power does Virgo measure?

- In general, the beam is not a plane wave but a spherical wave

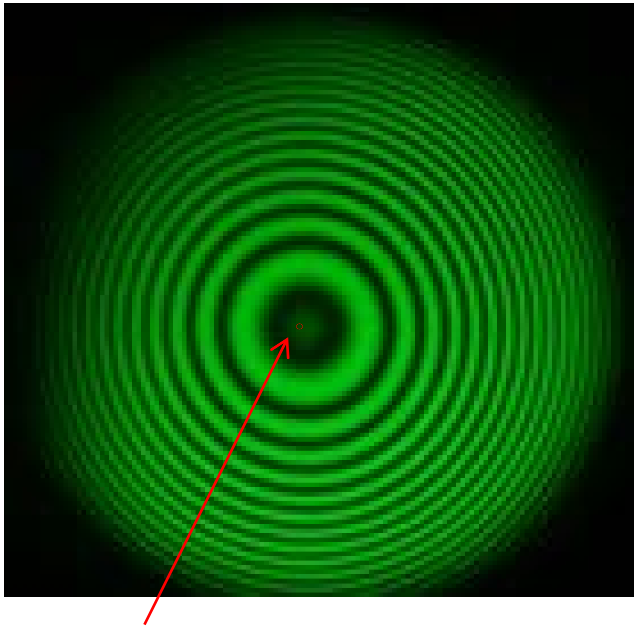
  → interference pattern

  (and the complementary pattern in reflection)

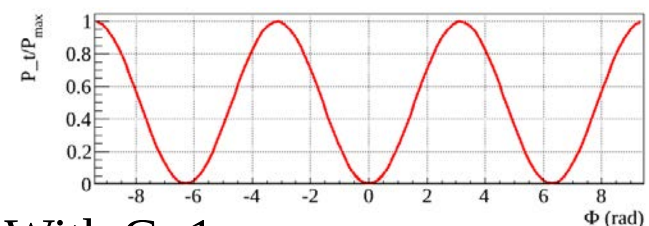
- Virgo interference pattern much larger than the beam size:

  ~1 m between two consecutive fringes

  → we do not study the fringes in nice images !



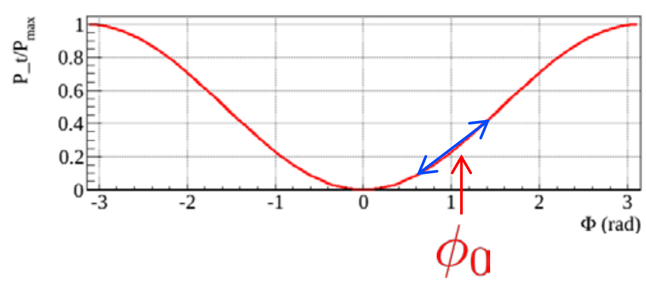
Equivalent size of Virgo beam



With C=1

Freely swinging mirrors

Setting a working point →



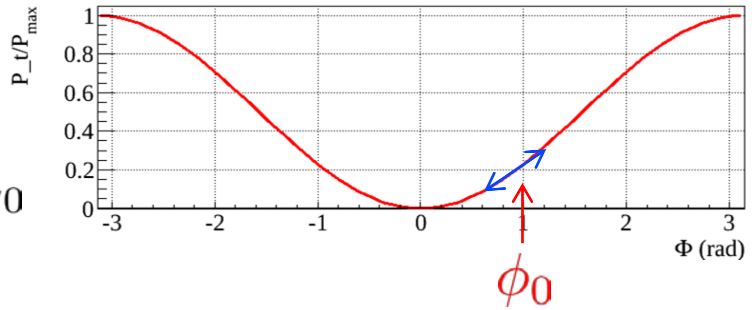
Controlled mirror positions

From the power to the gravitational wave

$$P_t = \frac{P_i}{2}\left(1 - C\cos(\phi)\right) \quad \text{where } \phi = 2\frac{2\pi}{\lambda}(l_y - l_x)$$

- Around the working point:

$$\left.\frac{\mathrm{d}P_t}{d\phi}\right|_{\phi_0} = \frac{P_i}{2} C\sin(\phi_0) \quad \text{where } \phi_0 = \frac{4\pi}{\lambda}\Delta L_0$$



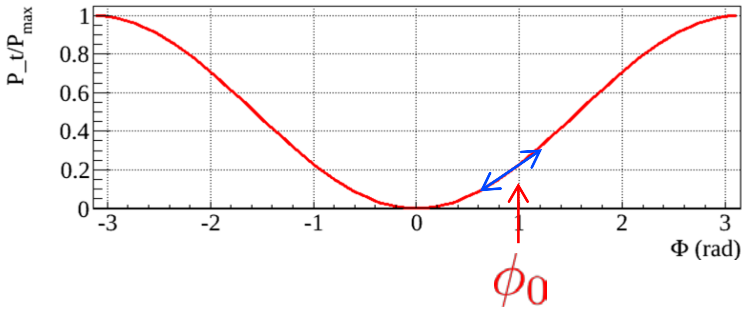
- Power variations as function of small differential length variations:

$$\delta P_t = \frac{P_i}{2} C\sin(\phi_0)\delta\phi$$

$$\delta P_t = P_i\, C\, \frac{2\pi}{\lambda}\sin\left(\frac{4\pi}{\lambda}\Delta L_0\right)\delta\Delta L$$

$\delta P_t \propto \delta\Delta L = hL_0$ around the working point !

From the power to the gravitational wave



- Around the working point:

$$\delta P_t = P_i\, C\, \frac{2\pi}{\lambda}\, \sin\left(\frac{4\pi}{\lambda}\Delta L_0\right) \delta\Delta L$$

$$\delta P_t = \text{(Interferometer response)} \times \delta\Delta L$$

(W/m)

Measurable physical quantity

Physical effect to be detected

Improving the interferometer sensitivity

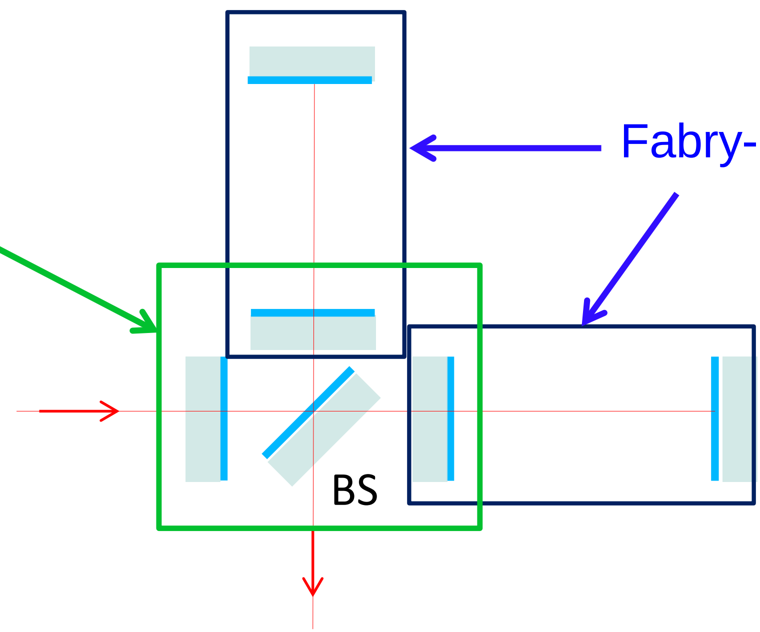
$$\delta P_t = P_i C \sin\left(\frac{4\pi}{\lambda}\Delta L_0\right)\underbrace{(k\,\delta\Delta L)}_{\propto\, \delta\phi}$$

Increase the input power on BS

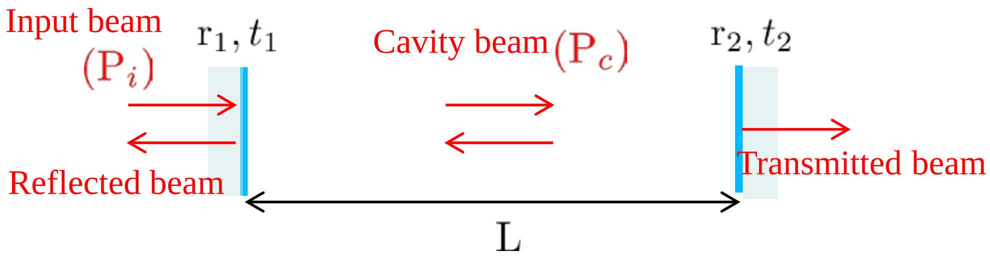
Increase the phase difference between the arms for a given differential arm length variation

Recycling **cavity**

Fabry-Perot **cavities** in the arms

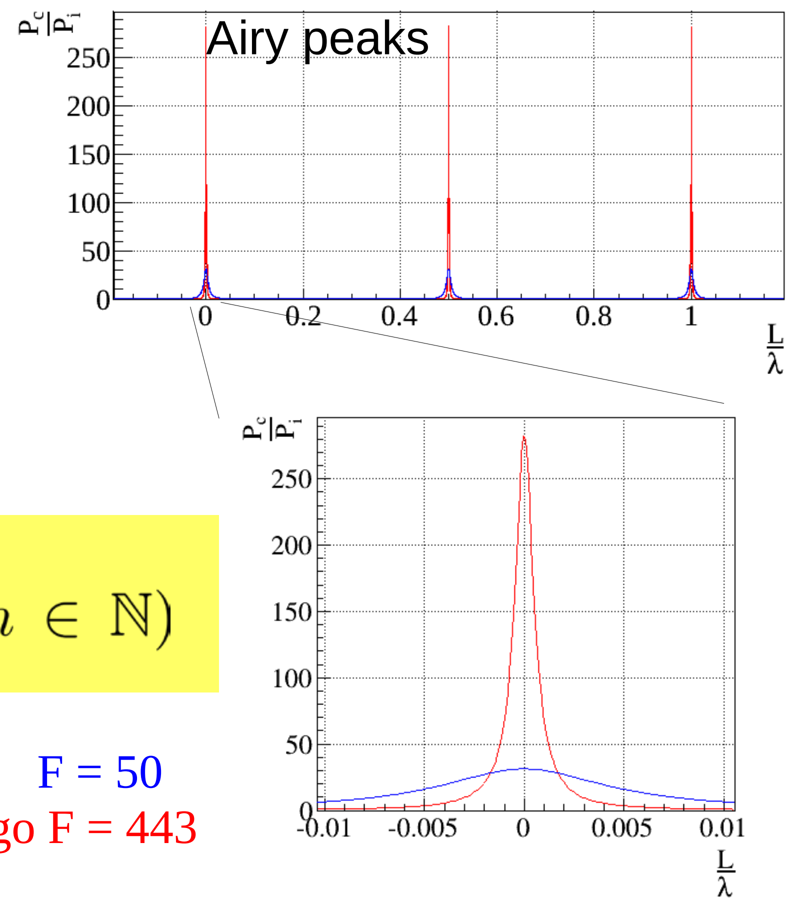


Beam resonant inside the cavities



$$P_c = P_i \frac{t_1^2}{(1 - r_1 r_2)^2} \frac{1}{1 + \left(\frac{2\mathcal{F}}{\pi}\right)^2 \sin^2(kL)}$$

$$\text{Finesse } \mathcal{F} = \frac{\pi\sqrt{r_1 r_2}}{1 - r_1 r_2}$$



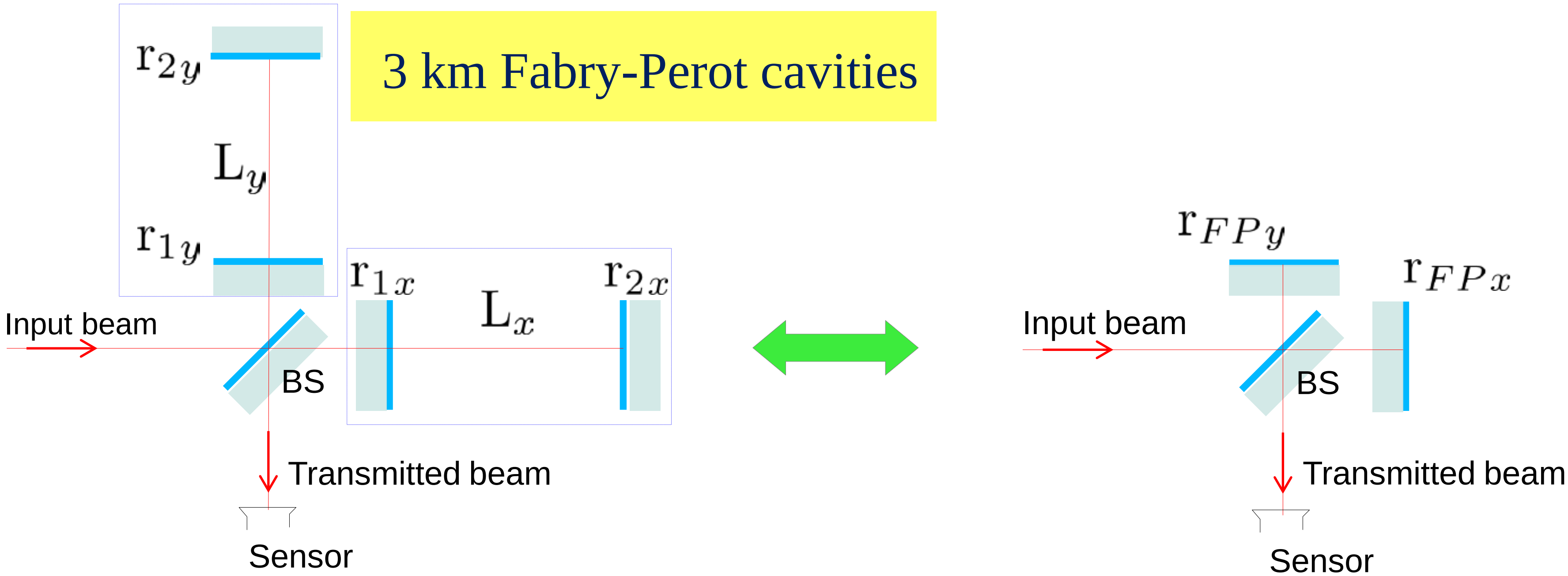
Virgo cavity at resonance: $L = n\frac{\lambda}{2} \quad (n \in \mathbb{N})$

Virgo F = 50
AdVirgo F = 443

Average number of light round-trips in the cavity: $N = \frac{2\mathcal{F}}{\pi}$

How do we amplify the phase offset?

3 km Fabry-Perot cavities



$$r_{FPx} = -1 \times e^{J \frac{2\mathcal{F}}{\pi} 2\,k\,\delta L_x}$$

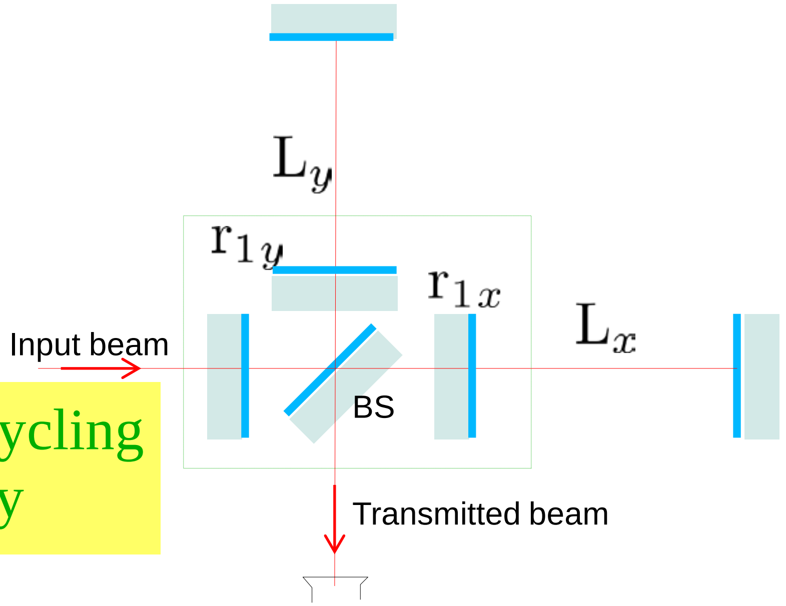
~number of round-trips in the arm
~300 for AdVirgo

(instead of $r_{armx} = -1 \times e^{J 2k(L_x + \delta L_x)}$ in the arm of a simple Michelson)

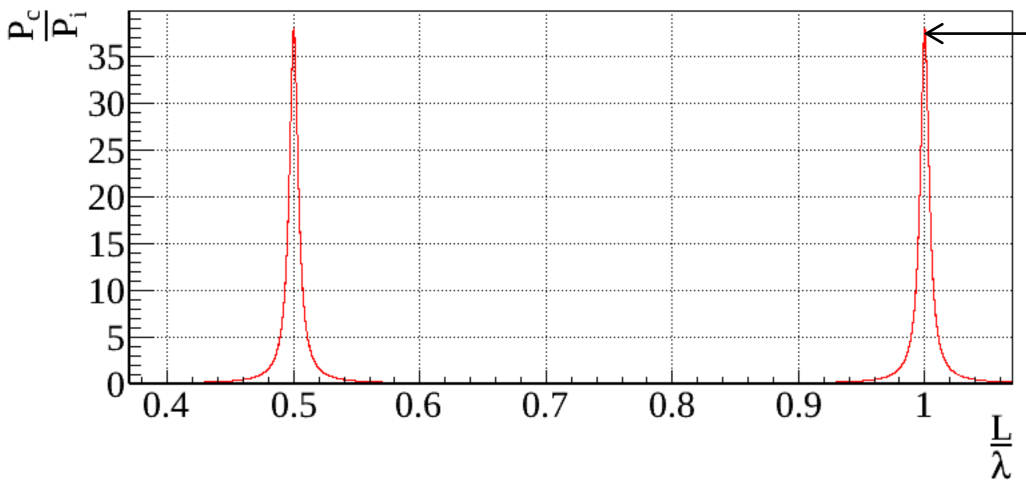
How do we increase the power on BS?

Detector working point close to a dark fringe
$\longrightarrow$ most of power go back towards the laser

Power recycling cavity



Resonant power recycling cavity



$$G_{PR} = 38 \quad (r_{PR}^2 = 0.95)$$

$\longrightarrow$ input power on BS increased by a factor 38!

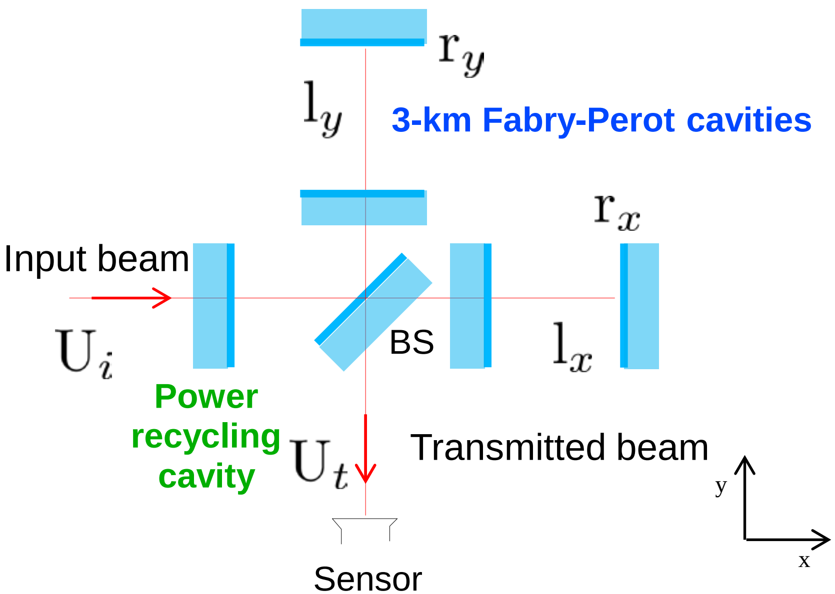
Improved interferometer response

- **Response of simple Michelson:**

$$\delta P_t = P_i \, C \, \frac{2\pi}{\lambda} \sin\left(\frac{4\pi}{\lambda}\Delta L_0\right) \delta\Delta L$$

$$\delta P_t = \underbrace{(\text{Michelson response})}_{(\text{W/m})} \times \delta\Delta L$$

- **Response of recycled Michelson with Fabry-Perot cavities:**



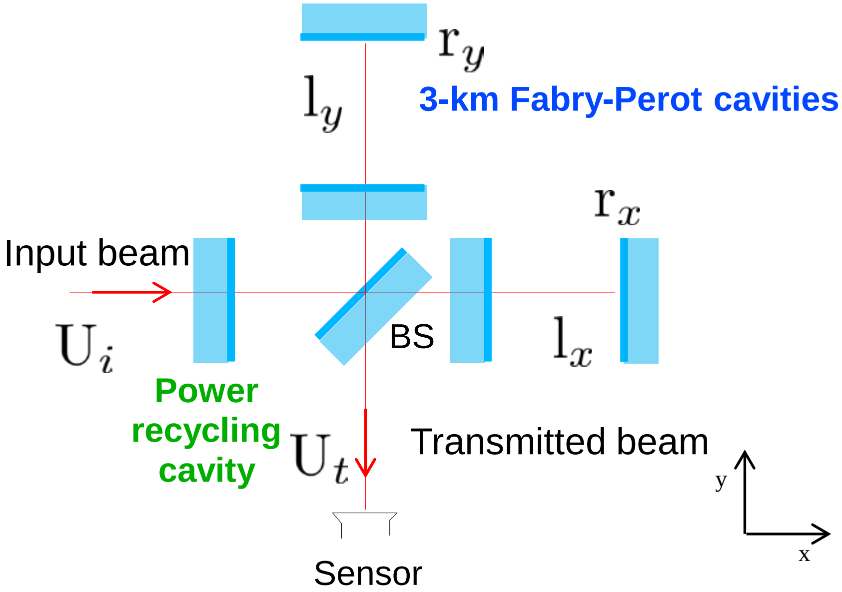
$$\delta P_t = \underbrace{G_{PR}}_{\sim 38} \, P_i \, C \, \frac{2\pi}{\lambda} \sin\left(\frac{4\pi}{\lambda}\Delta L_0\right) \underbrace{\frac{2\mathcal{F}}{\pi}}_{\sim 300} \delta\Delta L$$

For the same δΔ*L*,
δ*P*$_t$ has been increased by a factor ~12000

Order of magnitude of the « sensitivity »

$$\delta P_t = G_{PR} P_i \, C \, \frac{2\pi}{\lambda} \sin\left(\frac{4\pi}{\lambda}\Delta L_0\right) \frac{2\mathcal{F}}{\pi} \, \delta\Delta L$$

Laser wavelength	$\lambda = 1064$ nm
Input power	$P_i \sim 100$ W
Interferometer contrast	$C \sim 1$
Cavity finesse	$\mathcal{F} \sim 450$
Power recycling gain	$G_{PR} \sim 38$
Working point	$\Delta L_0 \sim 10^{-11}$ m



Shot noise due to output power of $\sim 50\,\mathrm{mW}$

$\rightarrow \delta P_{t,min} \sim 0.1\,\mathrm{nW}$ $\longrightarrow$ $\delta\Delta L_{min} \sim 5 \times 10^{-20}$ m

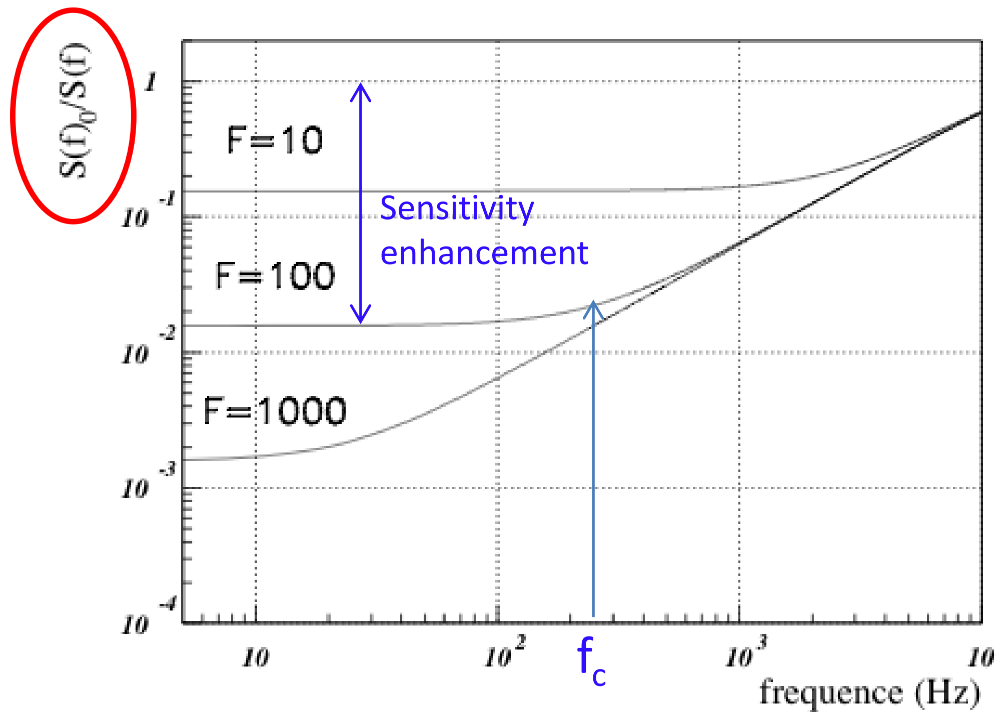
$$\rightarrow h_{min} = \frac{\delta\Delta L_{min}}{L} \sim 10^{-23}$$

In reality, the detector response depends on frequency...

Example of frequency dependency of the ITF response

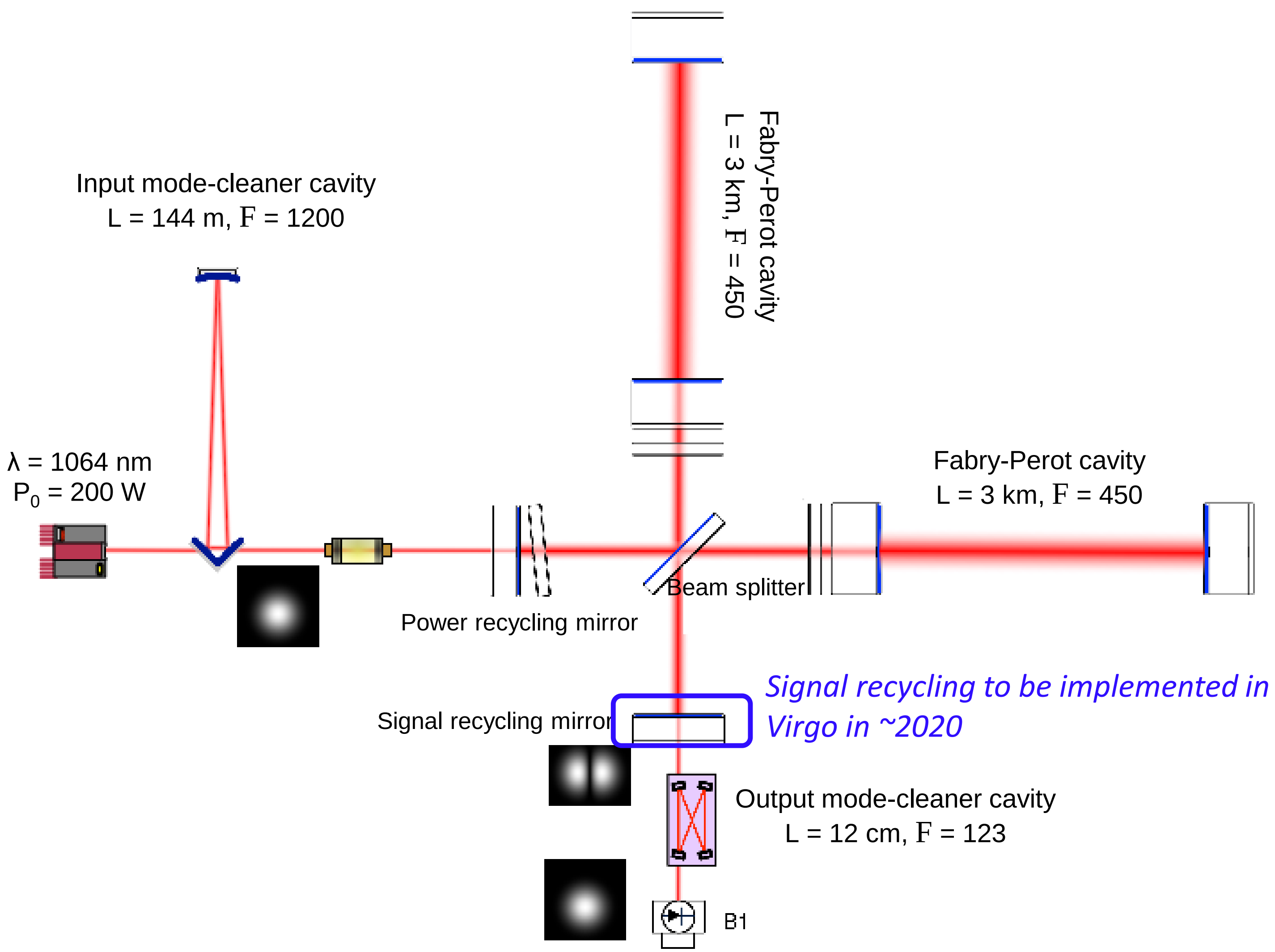
- Light travel time in the cavities must be taken into account
- Fabry-Perot cavities behave as a low pass filter
- Frequency cut-off: $$f_c = \frac{c}{4\mathcal{F}L}$$

Ratio between the sensitivity of an interferometer with Fabry-Perot cavities versus the sensitivity of an interferometer without cavities



- Finesse of Virgo Fabry Perot cavities: F = 450, L= 3 km → $f_c = 55$ Hz

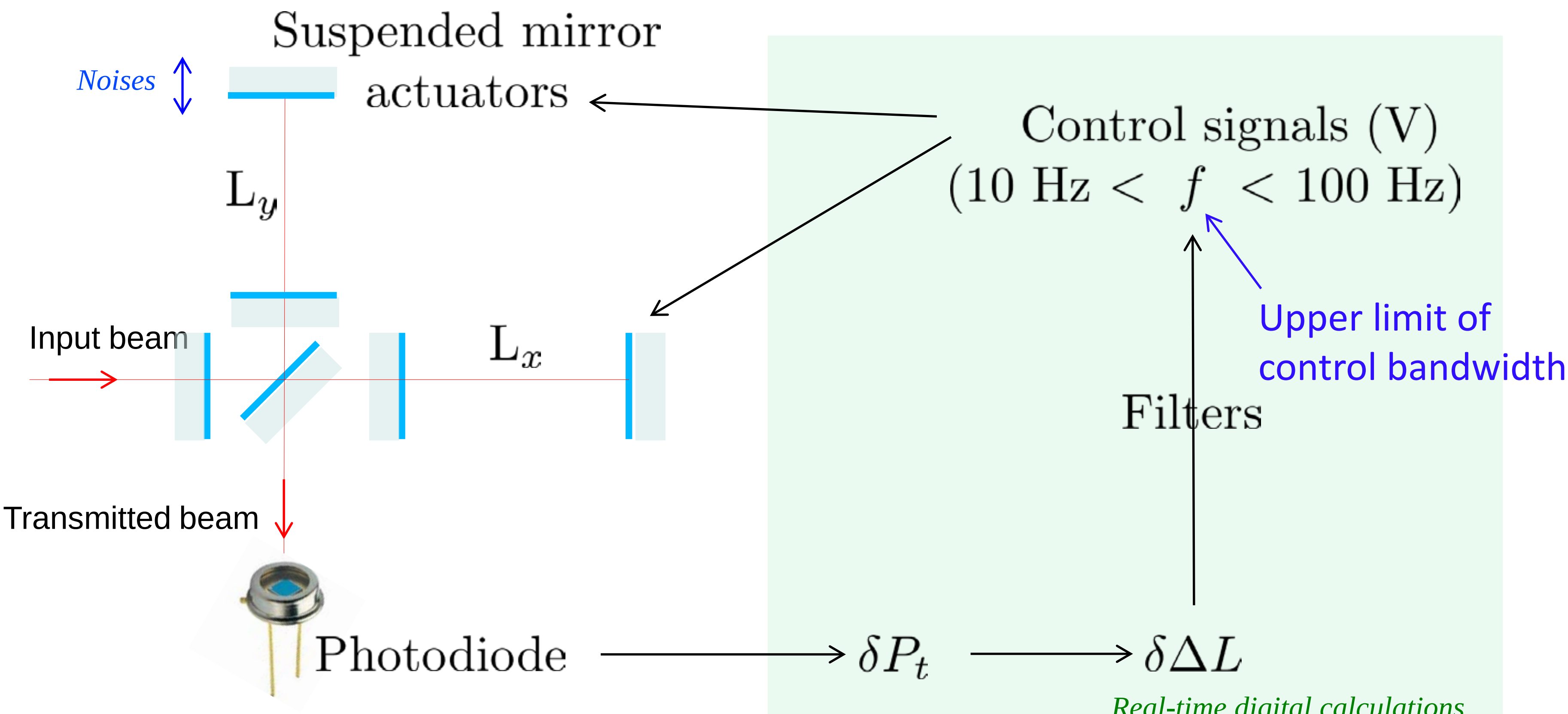
Optical layout of Virgo



How do we control the working point?

Small offset from a dark fringe: $\Delta L_0 = n\frac{\lambda}{2} + 10^{-11}\,\mathrm{m}$

- Controls to reduce the motion up to ~100 Hz
- Precision of the control $\delta\Delta L_{true}$ ~ 10^{-15} m

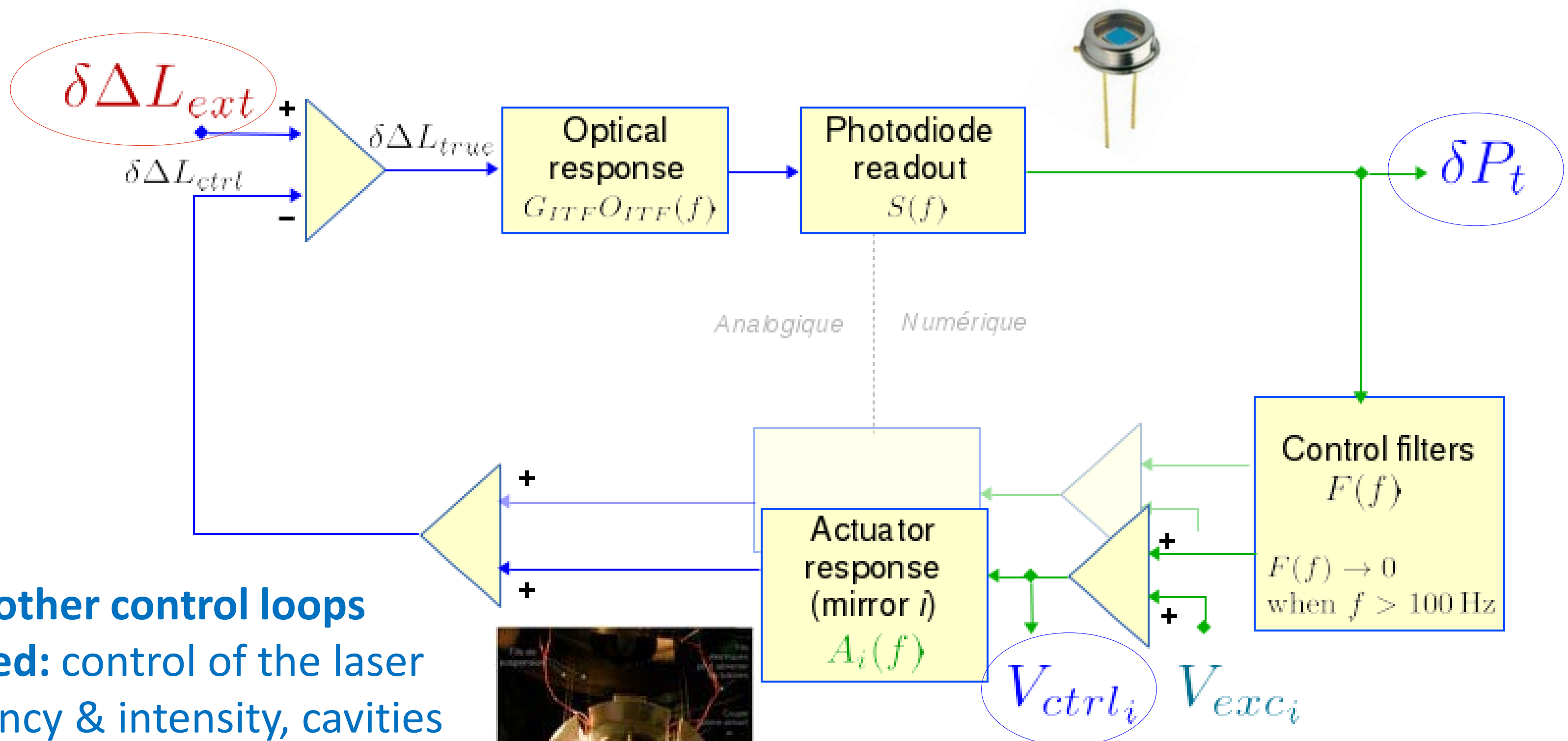


How do we control the working point?

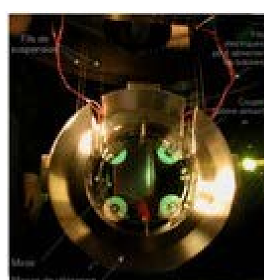
Small offset from a dark fringe: $\Delta L_0 = n\frac{\lambda}{2} + 10^{-11}\,\mathrm{m}$

- Controls to reduce the motion up to ~100 Hz
- Precision of the control $\delta\Delta L_{true}$ ~ 10^{-15} m

$$\delta\Delta L_{ext} = \delta\Delta L_{noise} + \delta\Delta L_{GW}$$



Many other control loops required: control of the laser frequency & intensity, cavities maintained at resonance, mirrors angular controls, …



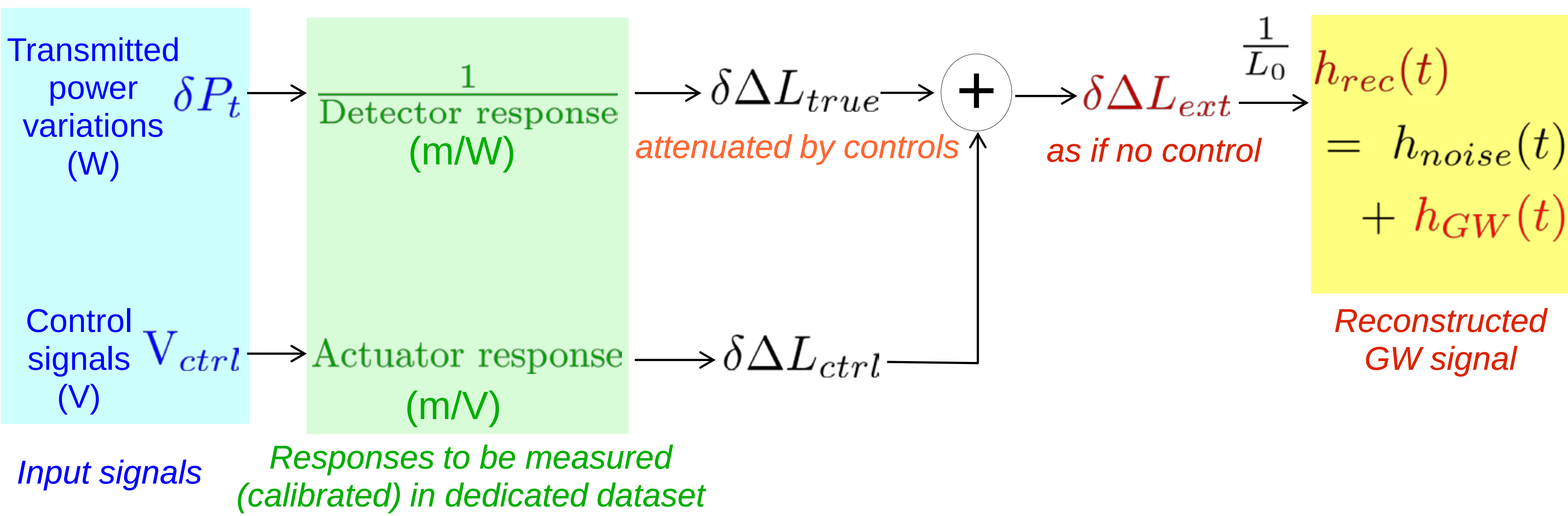
From the detector data to the GW strain h(t)

- High frequency (>100 Hz): mirrors behave as free falling masses

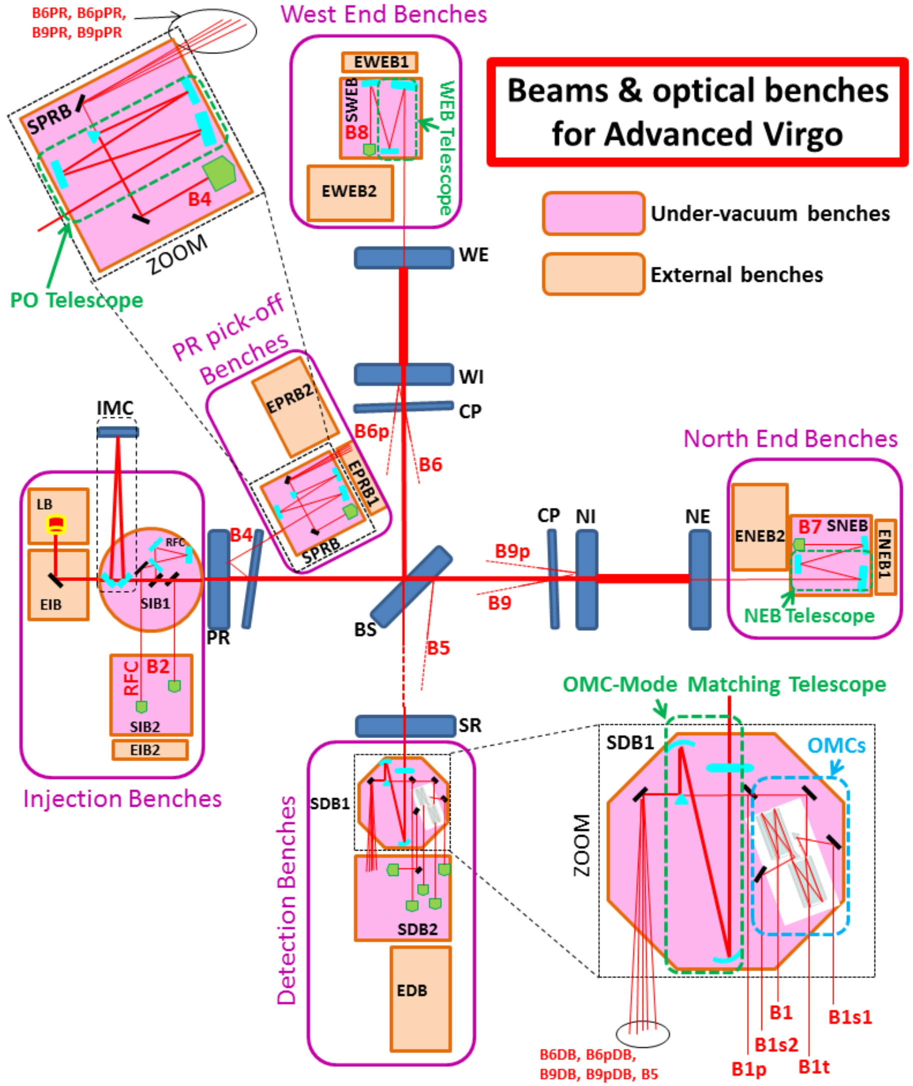
  $\rightarrow \quad h(t) = \frac{\delta \Delta L_{true}(t)}{L_0}$

- Lower frequency: the controls attenuate the noise... but also the GW signal!

  → the control signals contain information on *h(t)*



How to extract all error signals? Interferometer optical ports

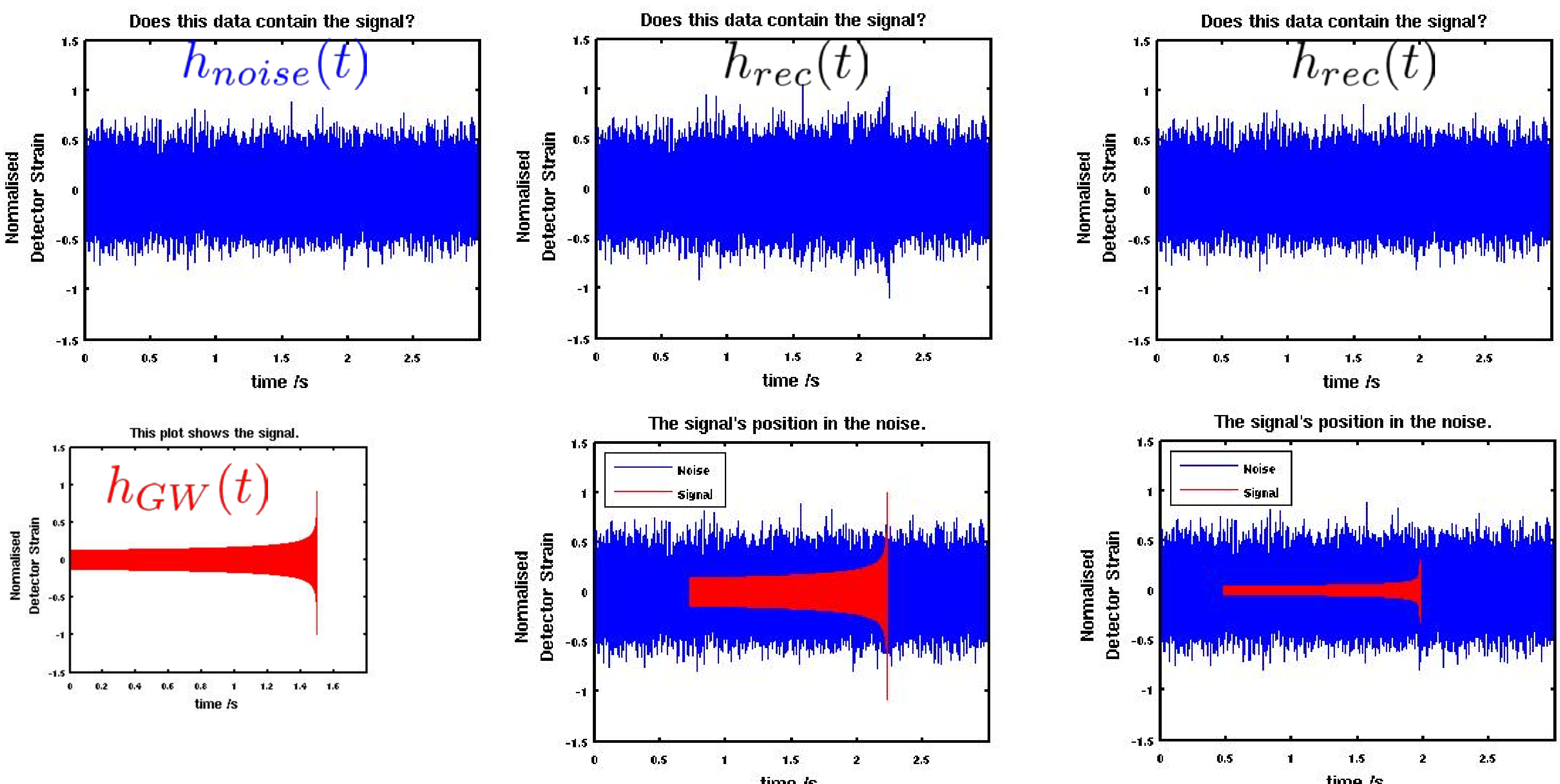


Noises limiting interferometer sensitivity: How to tackle them ?

Reminder: what is noise in Virgo?

- Stochastic (random) signal that contributes to the signal $h_{rec}(t)$ but does not contain information on the gravitational wave strain $h_{GW}(t)$

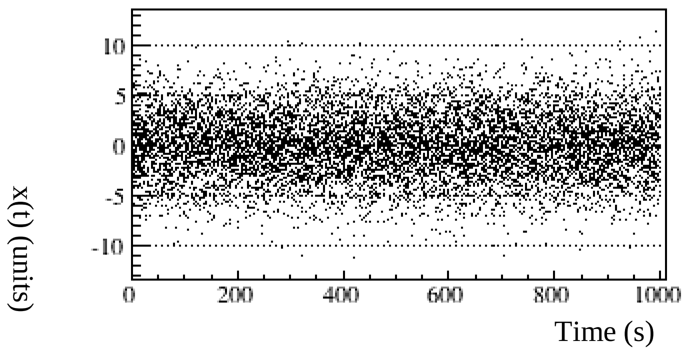
$$h_{rec}(t) = h_{noise}(t) + h_{GW}(t)$$

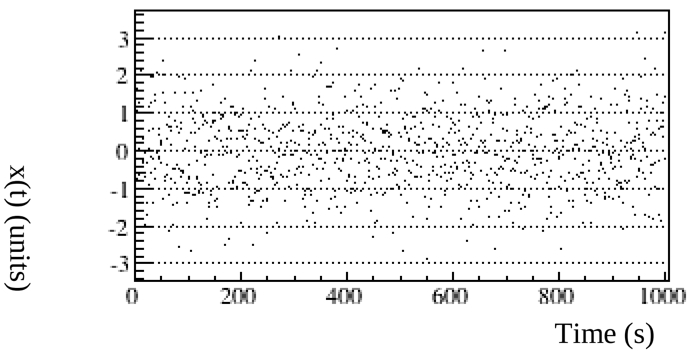


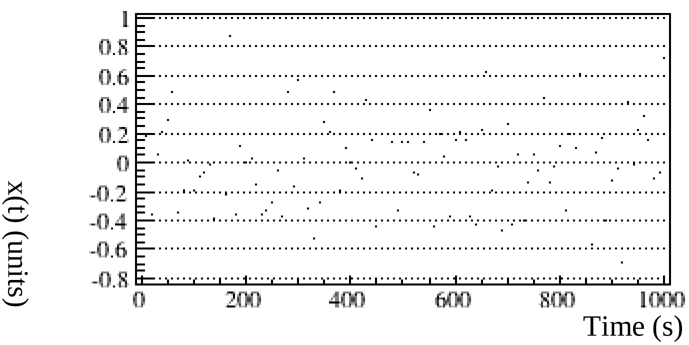
Extracted from Black Hole Hunter: http://www.blackholehunter.org/

How do we characterize noise?

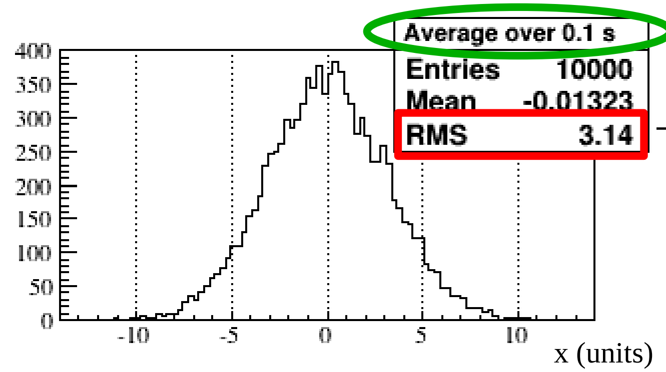
Data points (noise)





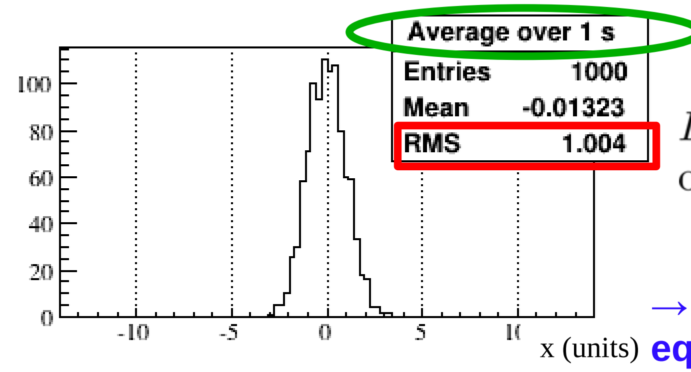


Distribution of the data



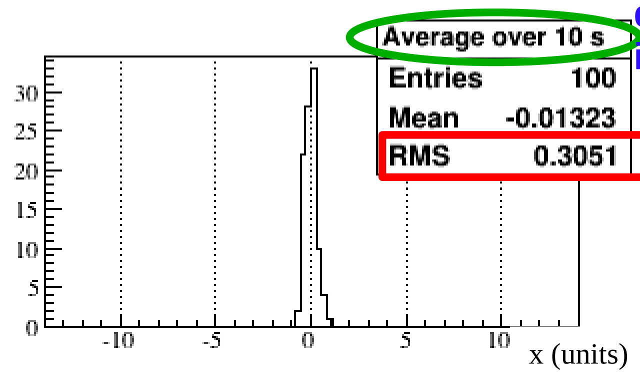
→ Noise characterised by its standard deviation σ_x

$$\sigma_x = \frac{D}{\sqrt{\text{average duration}}}$$



D is in (Data units × $\sqrt{\mathrm{s}}$) or $\frac{\text{Data units}}{\sqrt{\mathrm{Hz}}}$

→ its absolute value is equal to the standard deviation of the noise when it is averaged over 1 s

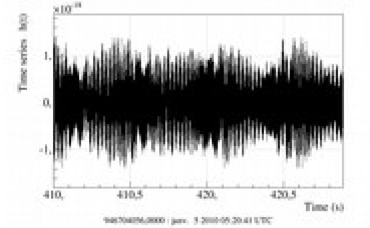


→ D = 1 units/√Hz

From hrec(t) to Virgo sensitivity curve

1/ Reconstruction of h(t)

$$h_{rec}(t) = h_{noise}(t) + \cancel{h_{GW}(t)}$$



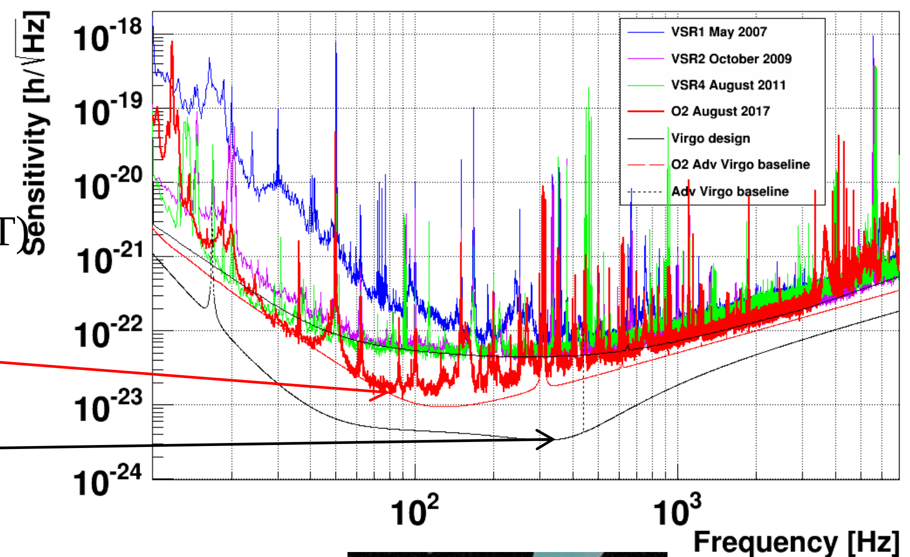
2/ Amplitude spectral density of *h(t)*
(noise standard deviation over 1 s)

$$ASD = \sqrt{PSD} = \sqrt{\frac{|DFT|^2}{T}}$$

Discrete Fourier Transform (DFT)

~5 x10^{-20} m/√Hz (Advanced Virgo O2, 2017)

~10^{-20} m/√Hz (Advanced Virgo nominal)



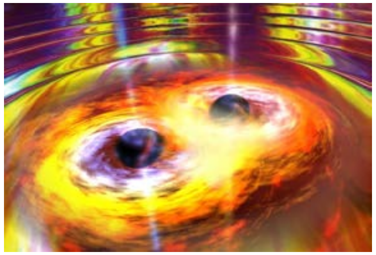


Image: Danna Berry/SkyWorks/NASA

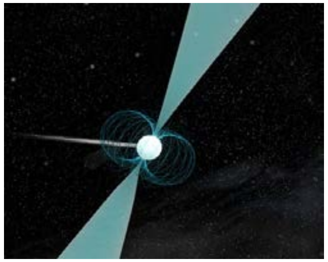


Image: B. Saxton (NRAO/AUI/NSF)

Compact Binary Coalescences

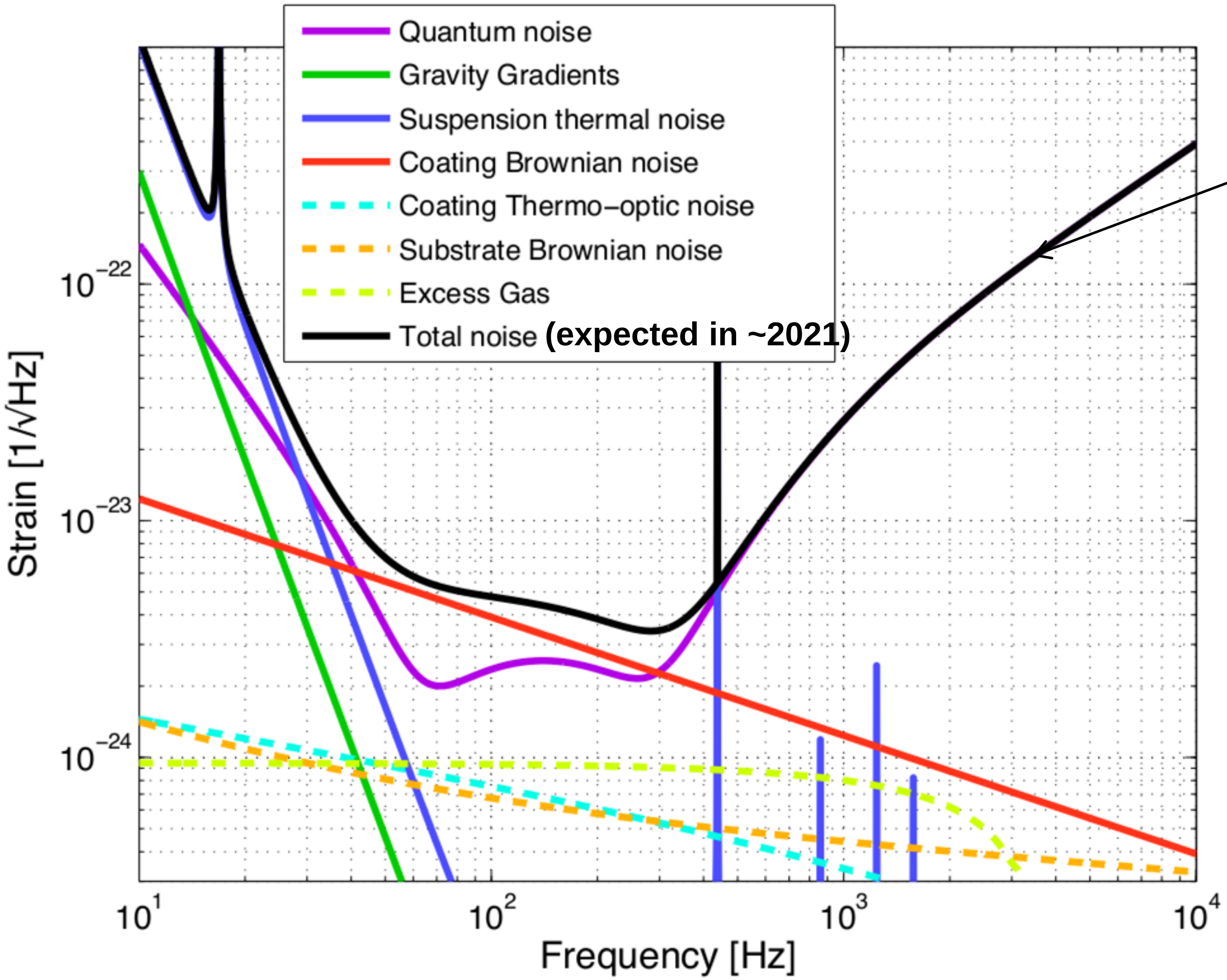
Signal lasts for a few seconds
→ can detect h ~ 10^{-23}

Rotating neutron stars

Signal averaged over days (~10^6 s)
→ can detect h ~ 10^{-26}

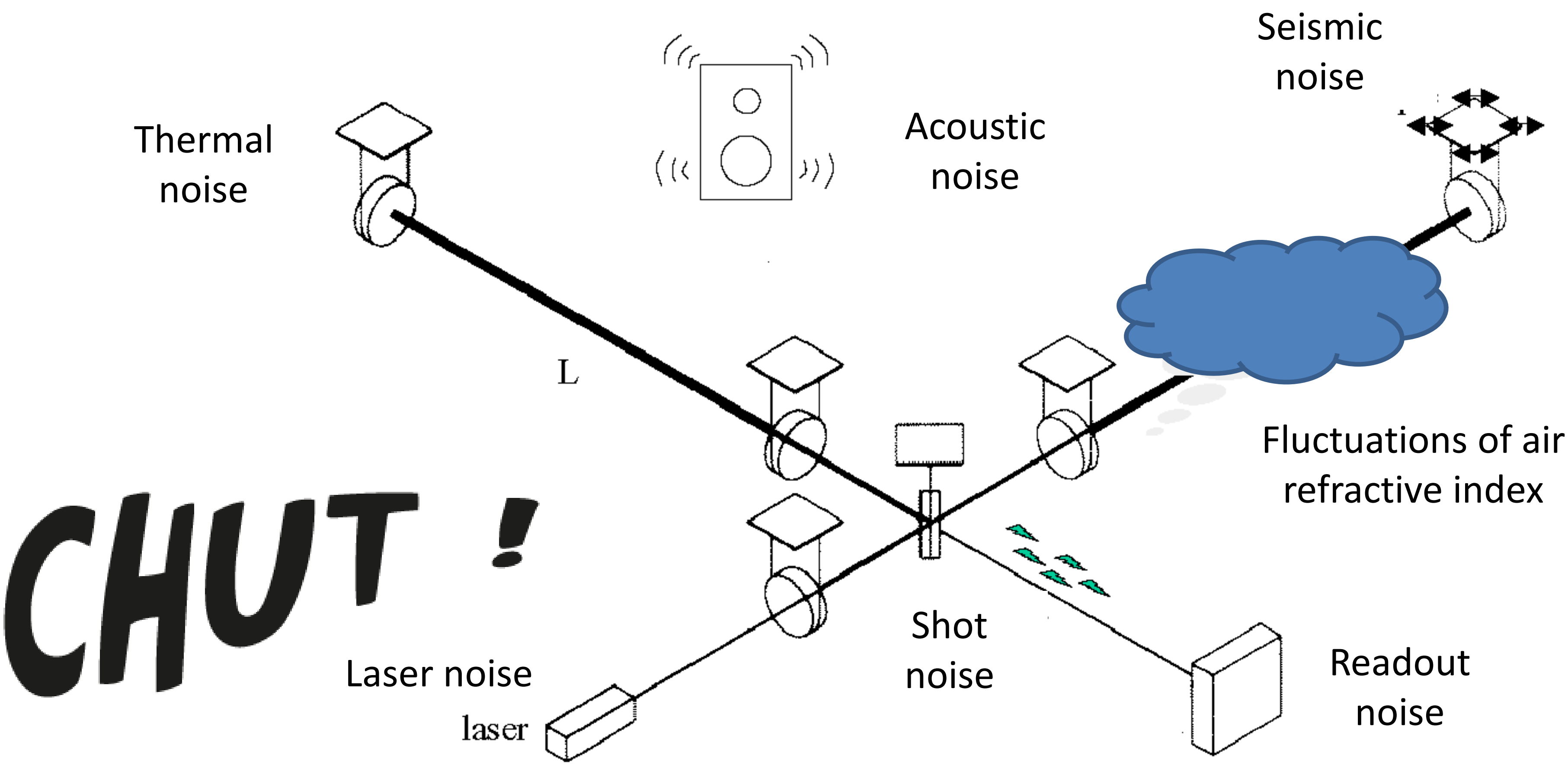
Nominal sensitivity of Advanced Virgo

Fundamental noise only
Possible technical noise not shown



Sum of all the noises

Fundamental noise sources



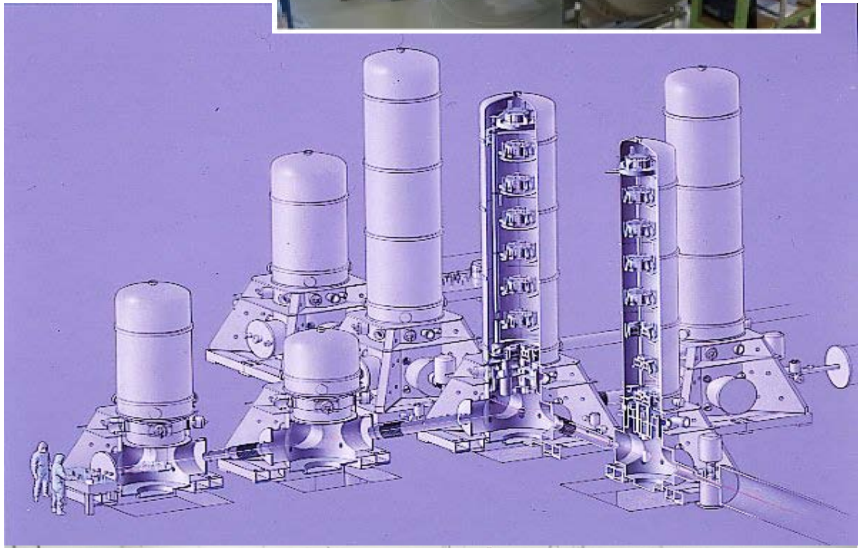
Under vacuum

Goals

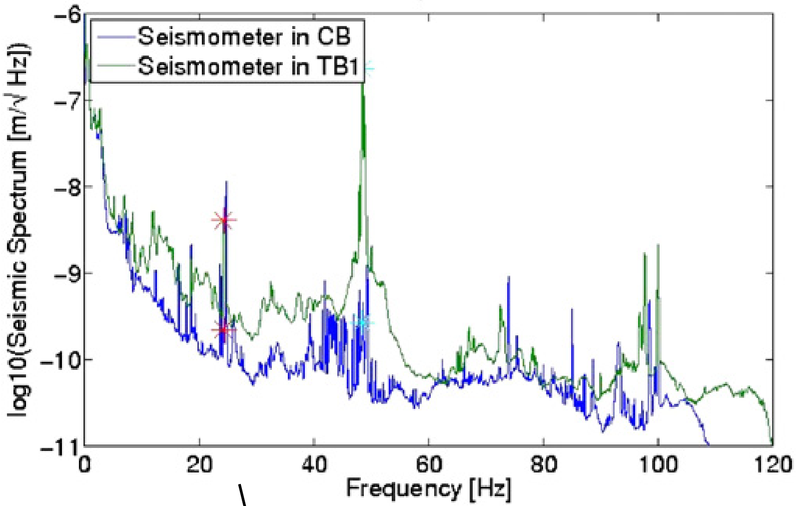
- Isolation against acoustic noise
- Avoid measurement noise due to fluctuations of air refractive index
- Keep mirrors clean

Advanced Virgo vacuum in a few numbers:

- Volume of vacuum system: 7000 m^3
- Different levels of vacuum:
  - 3 km arms designed for up to 10^{-9} mbar (Ultra High Vacuum)
  - $\sim 10^{-6}$ - 10^{-7} mbar in mirror vacuum chambers (« towers »)
- Separation between arms and towers with cryotrap links

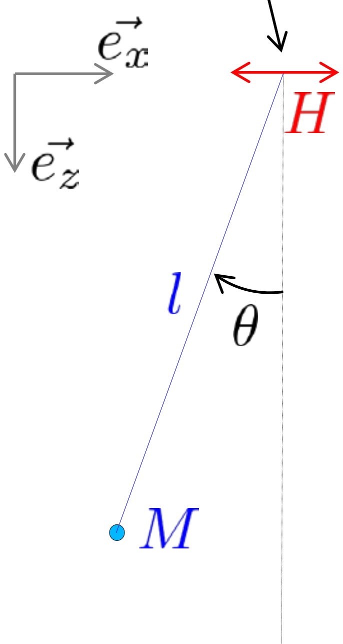


Seismic noise and suspended mirrors



Ground vibrations up to ~1 μm/√Hz at low frequency decreasing down to ~10 pm/√Hz at 100 Hz

$\gg 10^{-19}\,\mathrm{m}/\sqrt{\mathrm{Hz}}$ needed to detect GW !!

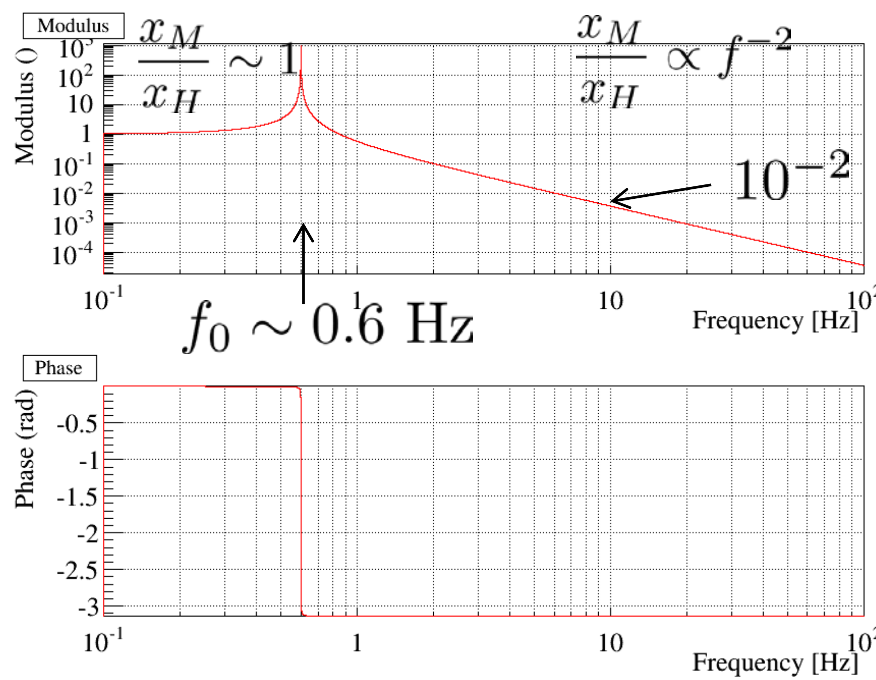


Assuming

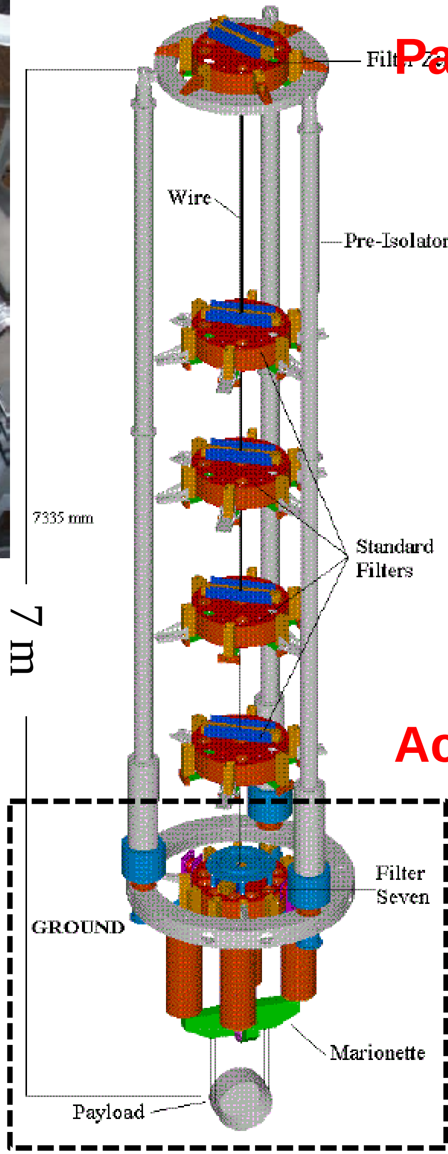
δx_H small and sinusoidal and θ small:

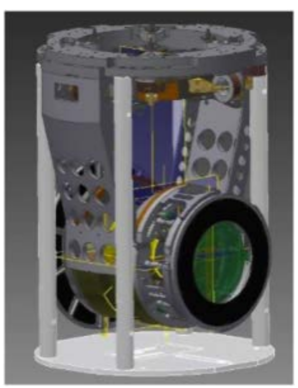
$$\underline{x_M} = \underline{\mathcal{H}} \times \underline{x_H}$$

Transfer function



Seismic noise: Virgo super-attenuators





Passive attenuation: 7 pendulum in cascade

$$\text{At 10 Hz: } \frac{x_{mirror}}{x_{ground}} \sim \left(10^{-2}\right)^7 = 10^{-14}$$

$$x_{ground} \sim 10^{-9}\,\mathrm{m}/\sqrt{\mathrm{Hz}}$$

$$\rightarrow x_{mirror} \sim 10^{-23}\,\mathrm{m}/\sqrt{\mathrm{Hz}}$$

This noise directly modifies the positions of the mirror surfaces, and thus $\delta\Delta L$ and $h_{rec}(t)$!

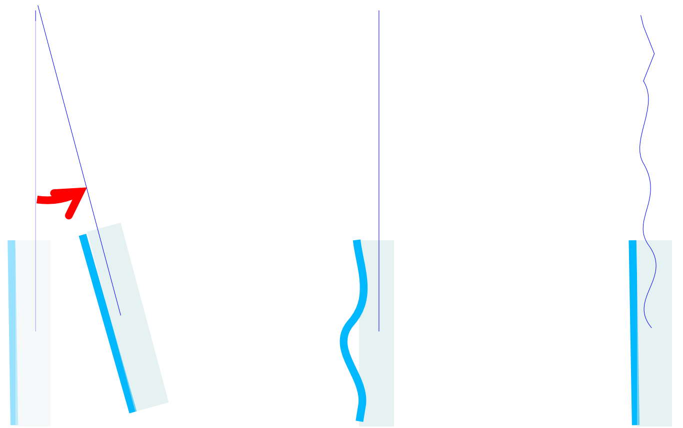
Active controls at low frequency

- Accelerometers or interferometer data
- Electromagnetic actuators
- Control loops

Thermal noise (pendulum and coating)

Microscopic thermal fluctuations

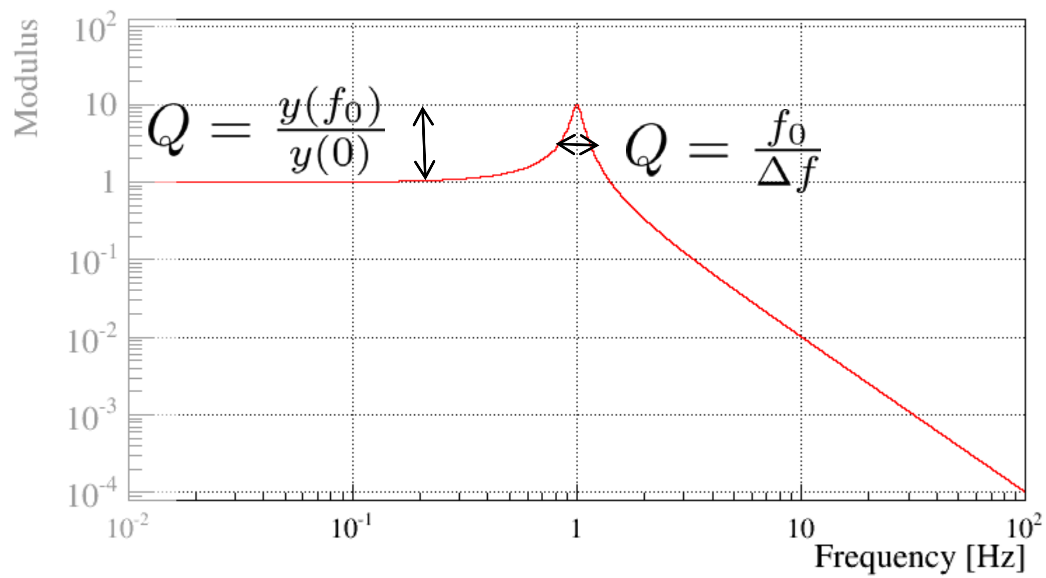
→ dissipation of energy through excitation of the macroscopic modes of the mirror

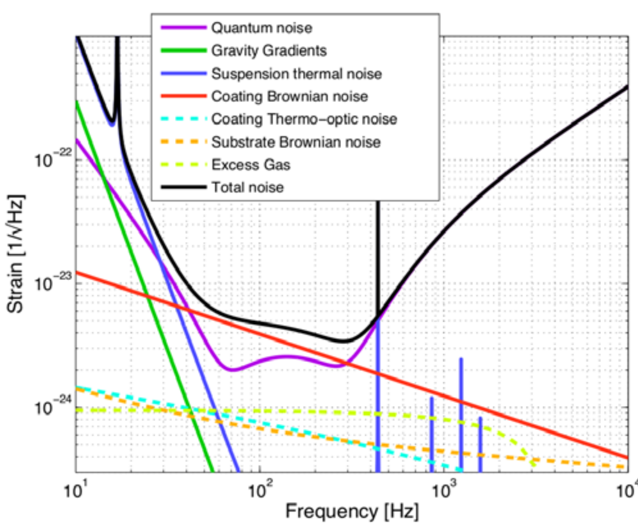


Pendulum mode
f < 40 Hz

"Mirror" mode
f> few kHz

"Violin" modes
f > 40 Hz

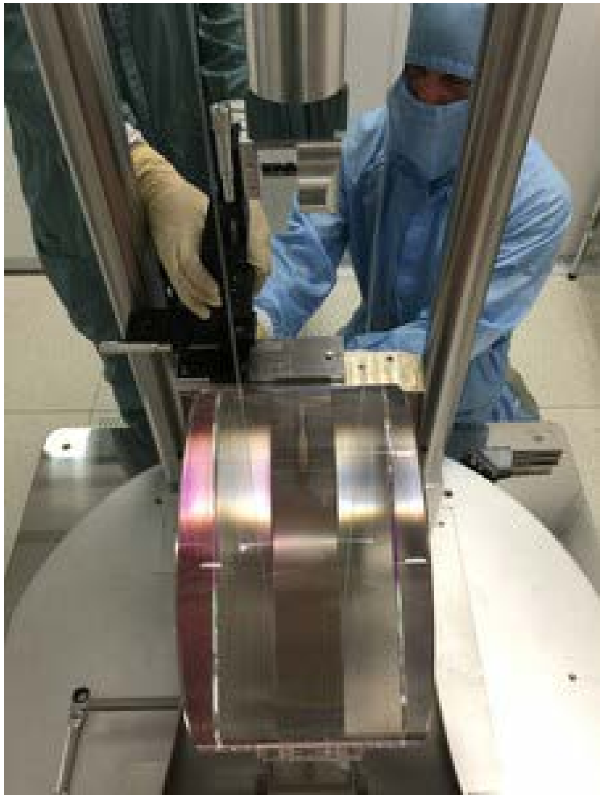




This noise directly modifies the positions of the mirror surfaces, and thus $\delta\Delta L$ and $h_{rec}(t)$!

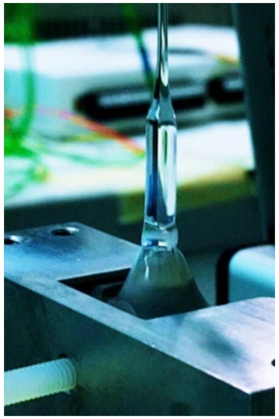
We want high quality factors Q to concentrate all the noise in a small frequency band

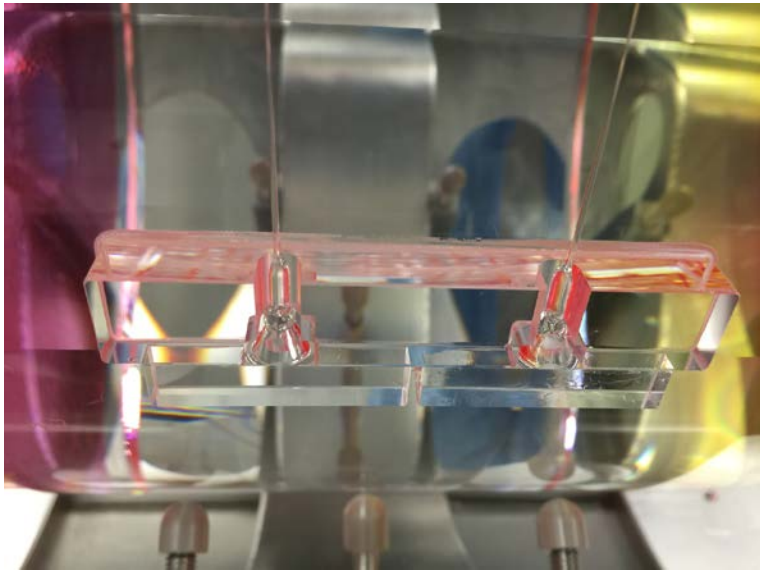
Reduction of thermal noise: monolithic suspensions



Increase the quality factor of the mirrors (wrt to steel wires):

Fused silica
- 400 µm diameter, increasing to ~ 1 mm at both ends
- 0.7 m length
- Load stress: 800 Mpa





Installed in Virgo in 2010
But failures in 2015/2016… (vacuum cleanliness issues)
… now fixed, re-installed beginning 2018

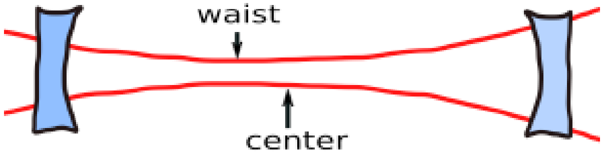
Reduction of thermal noise: mirror coating

40 kg mirrors of Advanced Virgo
35 cm diameter, 40 cm width
Suprasil fused silica

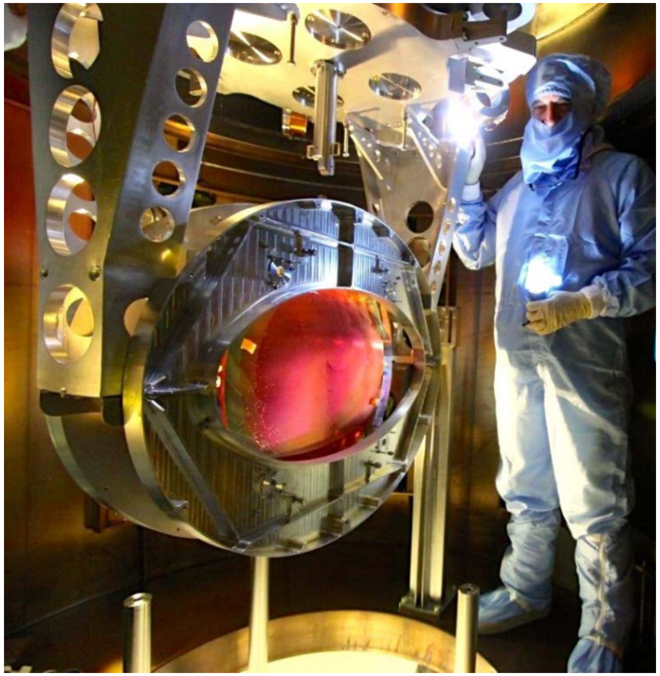
- Currently the main source of thermal noise
- Very high quality mirror coating developed in a lab close to Lyon (Laboratoire des Matériaux Avancés)
- R&D to improve mechanical properties of coating
- Cryogenics mirrors (at Kagra, future detectors)
  - other substrate
  - other coating
  - other wavelength

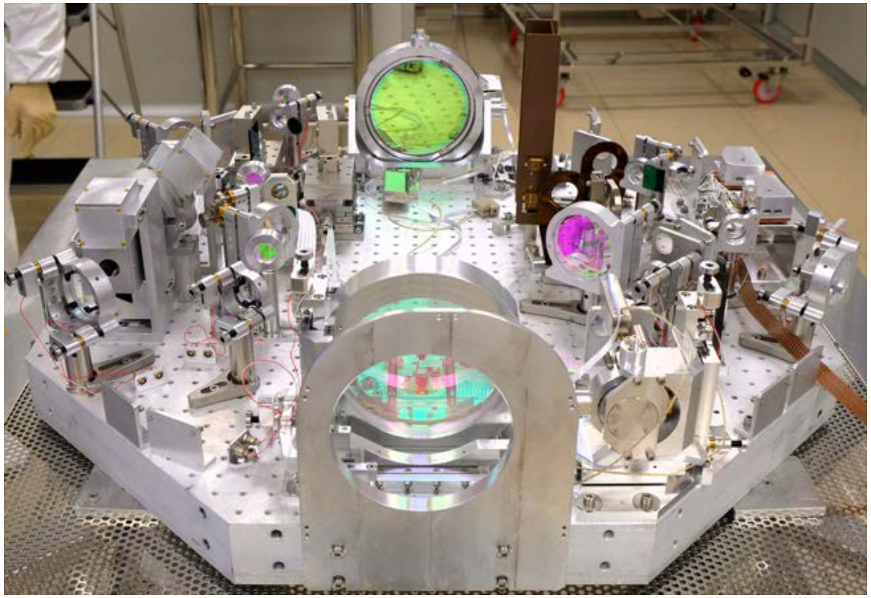
Thermal noise: coupling reduction

- Reduce the coupling between the laser beam and the thermal fluctuations
  → **use large beams**: fluctuations averaged over larger surface
  → Thermal Noise ~1/D, with D = beam diameter



- Impact of large beams:
  - Require large mirrors (and heavier):
    > Advanced Virgo beam splitter diameter = 55 cm
  - High magnification telescopes to adapt beam size to photodetectors (from w=50 mm on mirrors to w=0.3 mm on sensors) > require optical benches





Shot noise

Fluctuations of arrival times of photons (**quantum noise**)

Power received by the photodiode: P_t

$\rightarrow \quad N = \frac{P_t}{\mathrm{h}\nu}$ photons/s on average.

Arrival time of single photons

Standard deviation on this number: $\sigma_N = \sqrt{N}$

$$\rightarrow \quad \sigma_{P_t} = \sigma_N \times \mathrm{h}\nu = \sqrt{\frac{P}{\mathrm{h}\nu}}\mathrm{h}\nu = \sqrt{P_t \mathrm{h}\nu}$$

Virgo laser: $\lambda = 1.064\,\mu\mathrm{m} \rightarrow \nu = \frac{\mathrm{c}}{\lambda} \sim 2.8 \times 10^{14}\,\mathrm{Hz}$

Working point: $P_t \sim 80\,\mathrm{mW} \quad \rightarrow \quad \sigma_{P_t} = 0.1\,\mathrm{nW}/\sqrt{\mathrm{Hz}}$

$\rightarrow$ a variation of power is interpreted as a variation of distance $\delta \Delta L$

$$\delta P_t = (\text{Virgo response}) \times L_0 \times h$$

(in W/m)

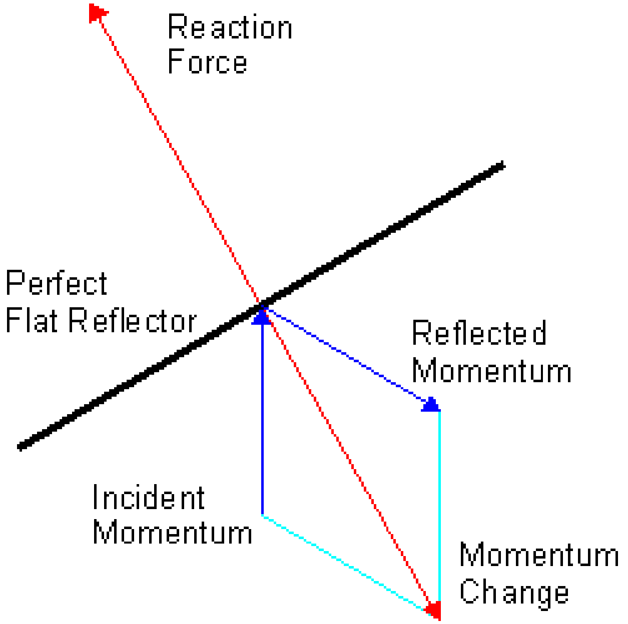
$$h_{equivalent} = \frac{1}{L_0} \frac{\sigma_{P_t}}{(\text{Virgo response})}$$

$\rightarrow h_{equivalent}\ \alpha\ 1/\sqrt{P_{in}}$

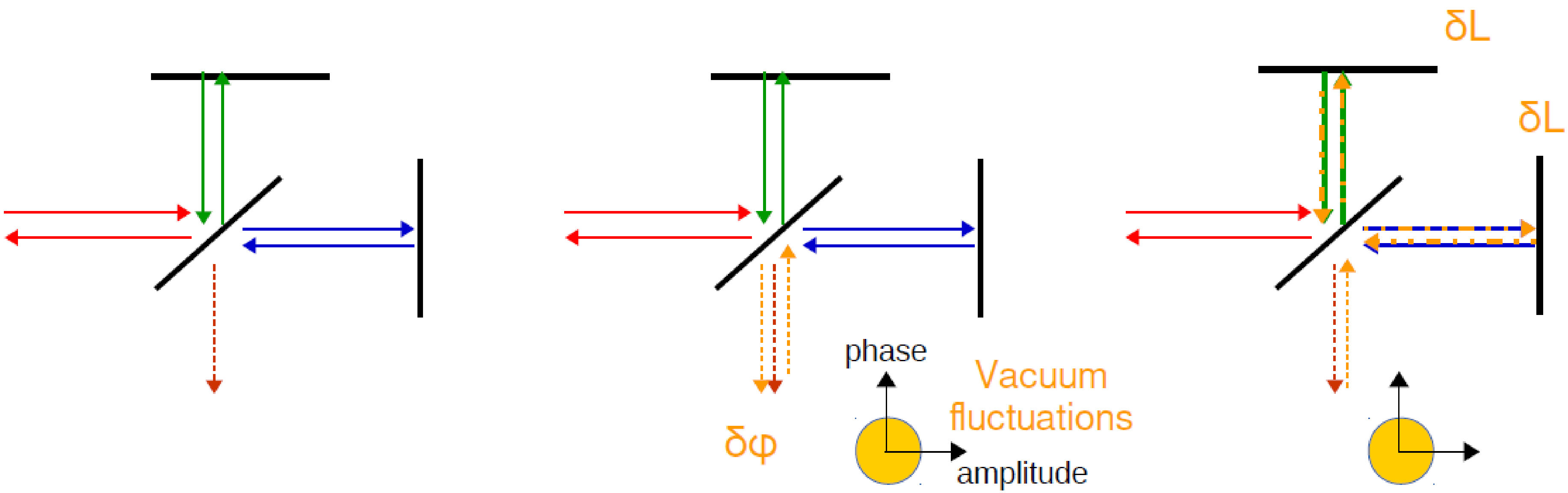
Radiation pressure noise

- Radiation pressure: transfer of photon's momentum to the reflective surface (recoil force)
- Radiation pressure noise: due to fluctuations of number of photons hitting the mirror surfaces > mirror motion noise
- Radiation pressure noise impact at low frequency:

  > Mirror motion filtered by pendulum mechanical response



Quantum noise: shot noise and radiation pressure noise



Working point close to a dark fringe
Signal related to phase offset

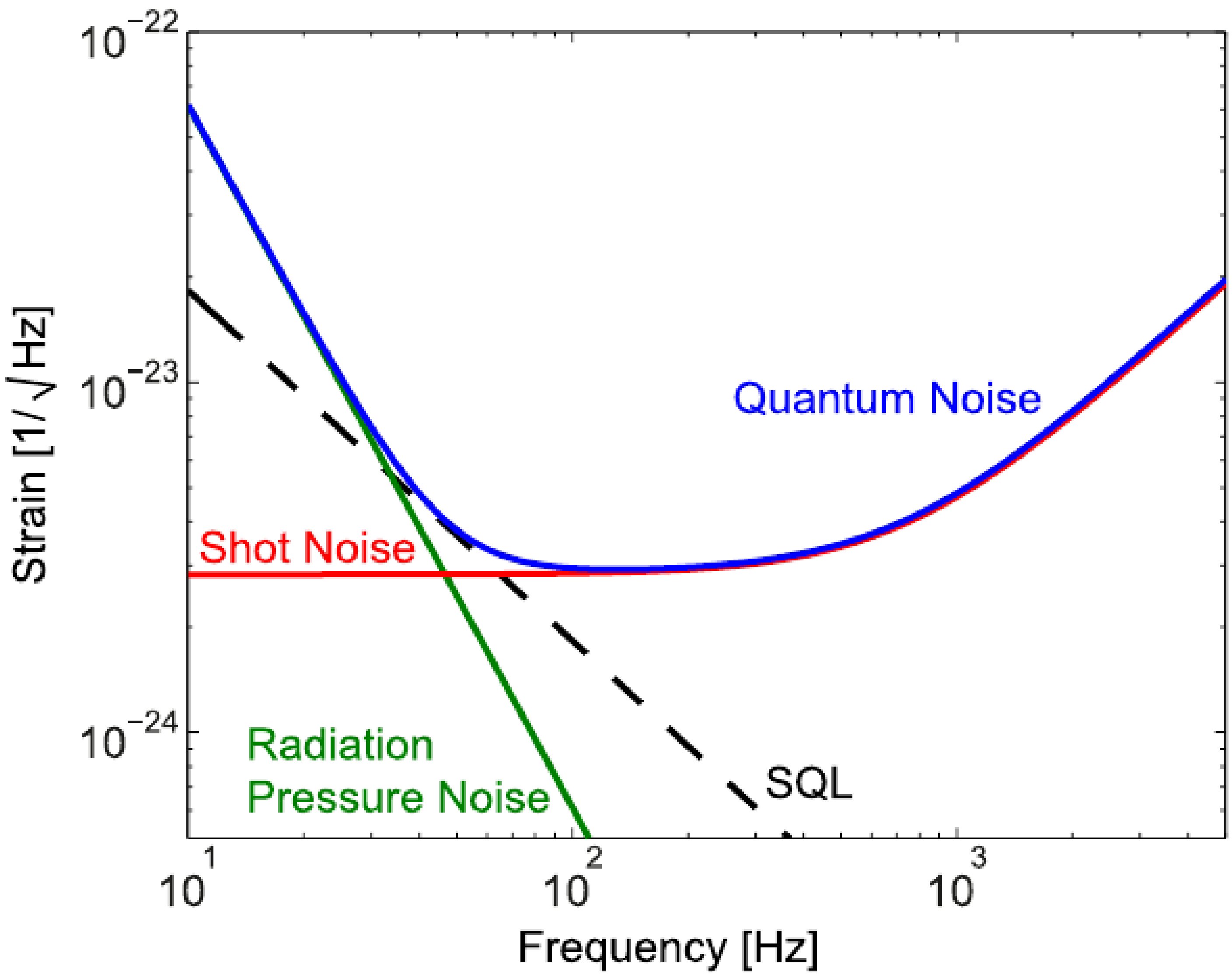
→ fluctuations of the phase
Photon shot noise

$$\propto \frac{\text{Phase fluctuation}}{\text{Laser power}}$$

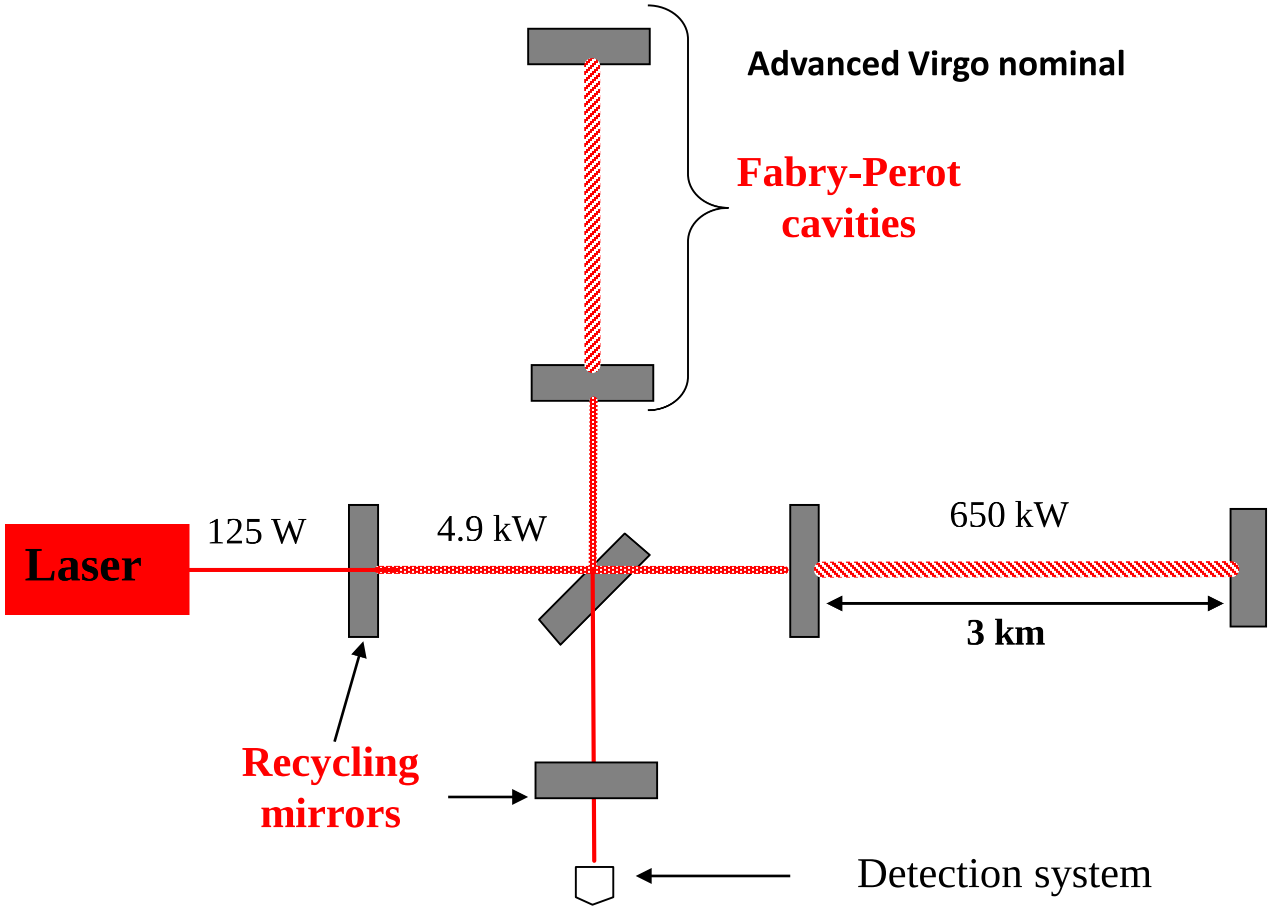
→ fluctuations of the mirror positions
Radiation pressure noise

∝ Amplitude fluctuation x Laser power

Quantum noise in the sensitivity



Minimizing shot noise with optical configuration



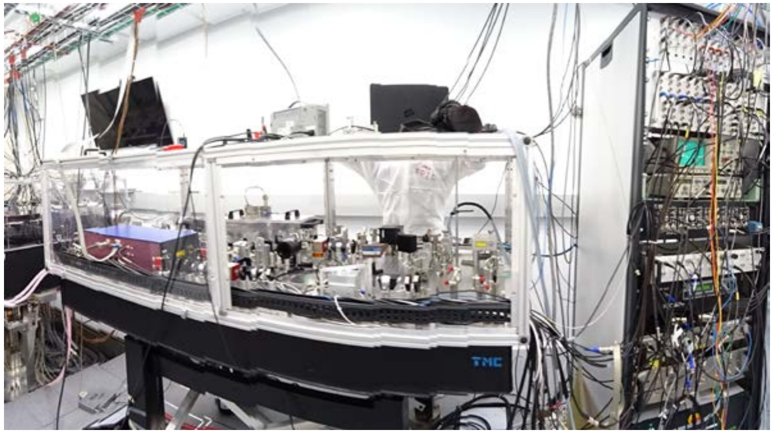
Reduction of shot noise: high power laser

Goal for AdV (nominal):

- continuous 200 W laser, stable monomode beam (TEM00), 1064 nm

Only 18W currently injected in Advanced Virgo

→ decrease shot noise contribution



But limited by side-effects:

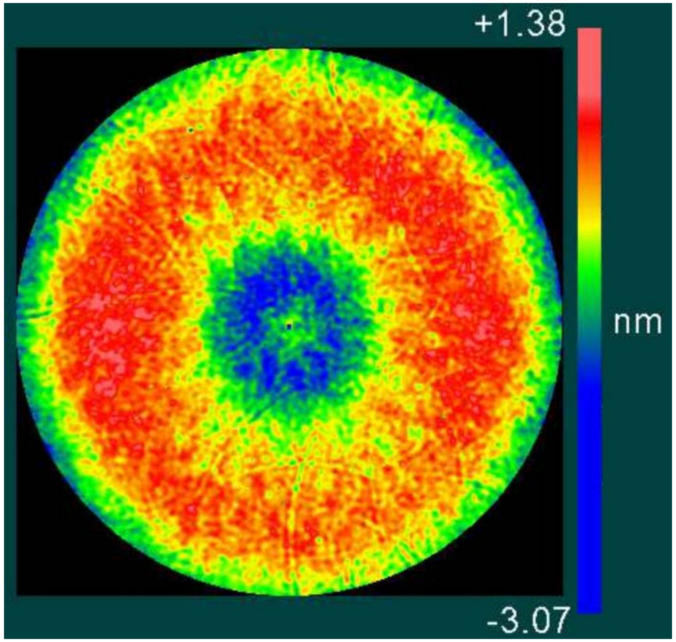
- **Radiation pressure**
  - Increase of radiation pressure noise
  - Cavities more difficult to control
  - Parametric instabilities: coupling of laser high order modes with mirrors mechanical modes

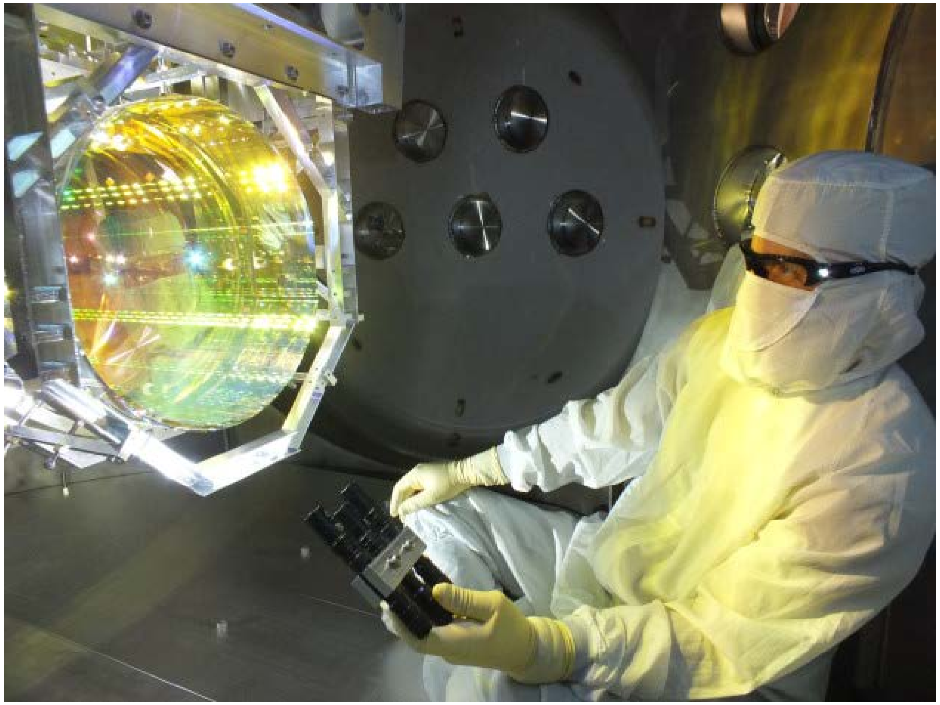
- **Thermal absorption in the mirrors (optical lensing)**
  → Need of thermal compensation system

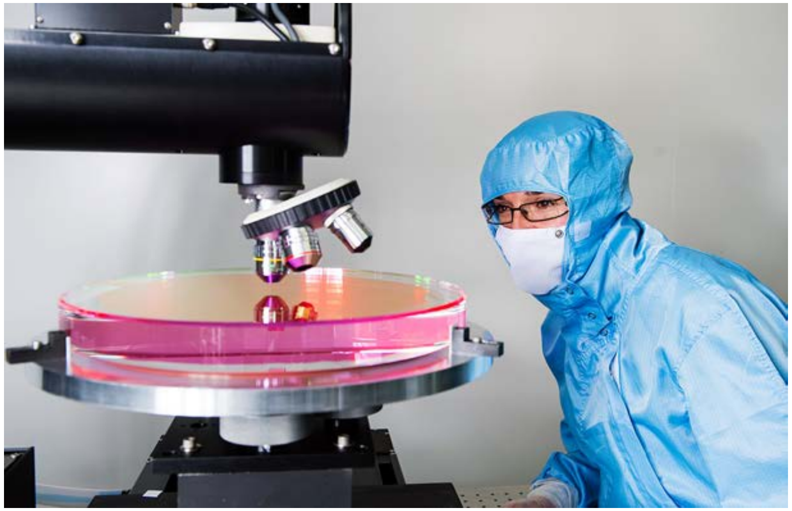
Avoid optical losses to not spoil high power → high quality mirrors

« Perfect » mirrors

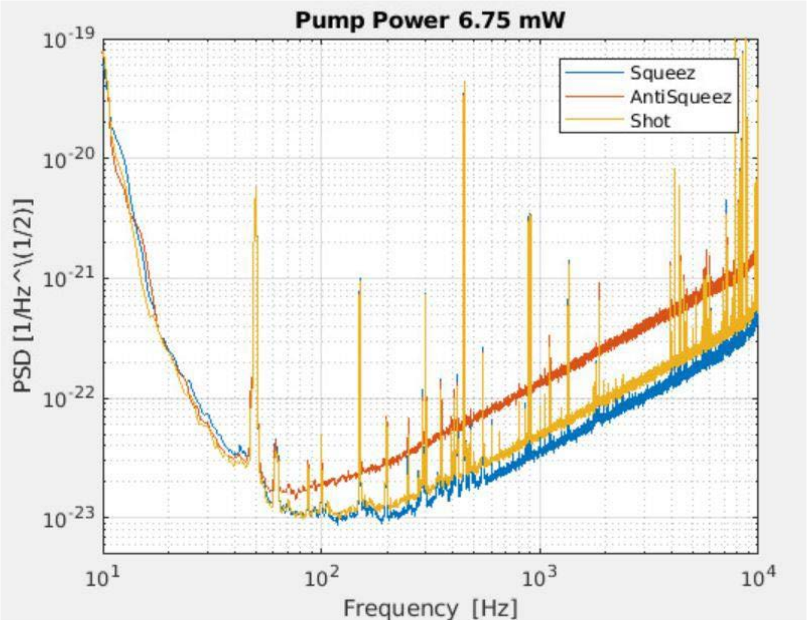
- 40 kg, 35 cm diameter, 20 cm thickness in ultra pure silica
- Uniformity of mirrors is unique in the world:
  - a few nanometers peak-to-valley
  - flatness < 0.5 nm RMS (over 150mm diameter)

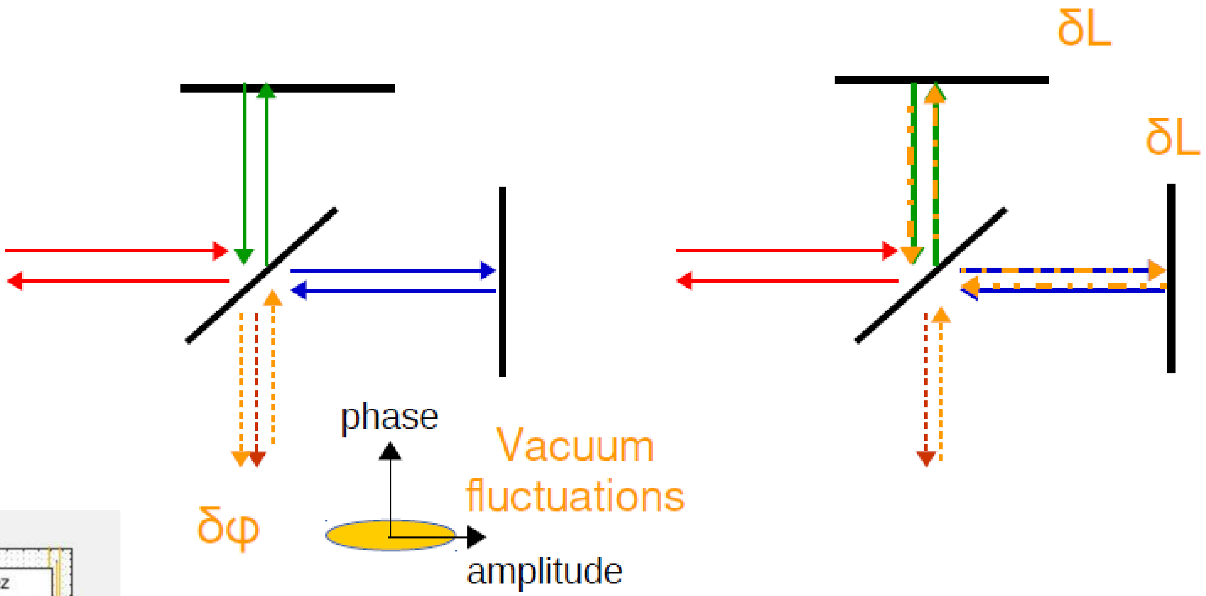






Reduction of shot noise: squeezing



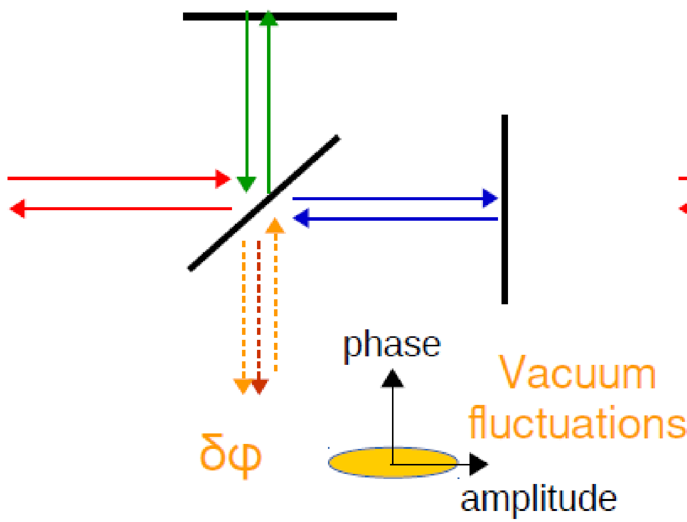


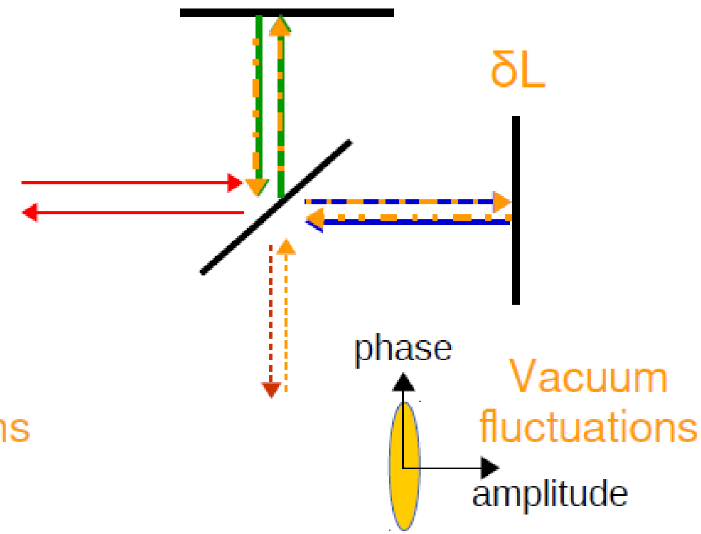
→ **Decrease shot noise**

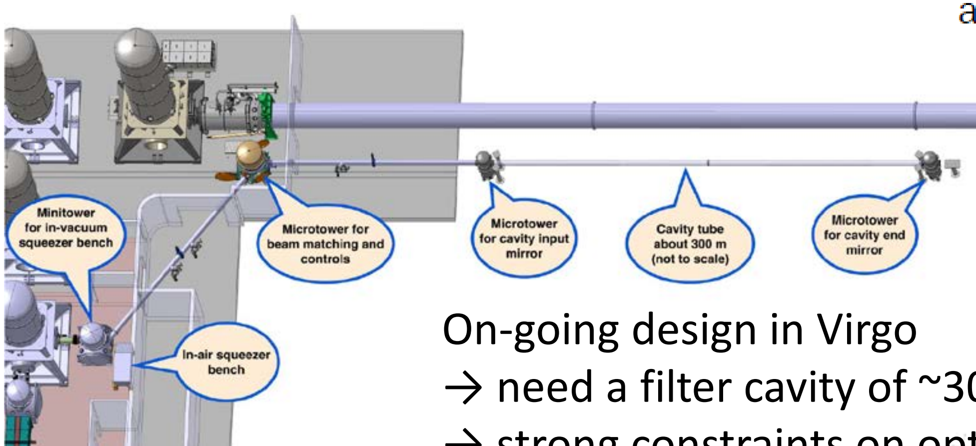
But increase radiation pressure noise

Installed in Virgo and LIGO

Reduction of quantum noise: frequency dependent squeezing





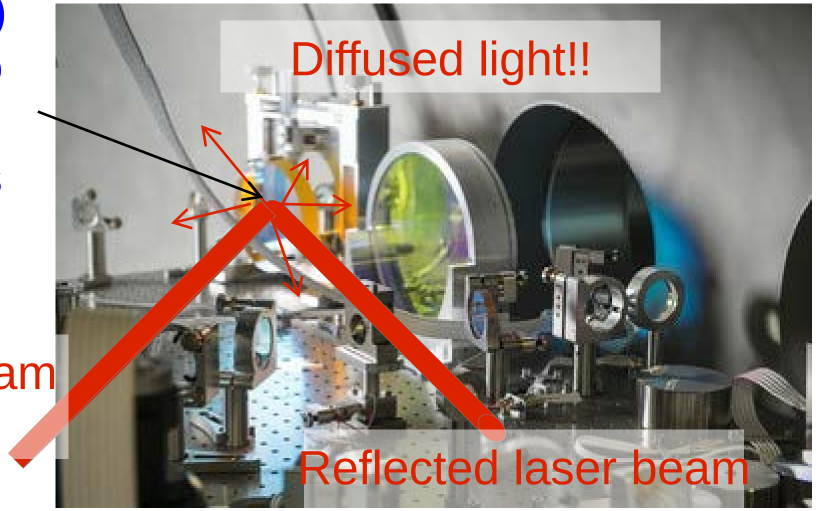


→ **Decrease shot noise AND radiation pressure noise**

On-going design in Virgo
→ need a filter cavity of ~300 m, with finesse ~10000
→ strong constraints on optical losses, beam matching and alignment, …
(suspended in vacuum optical benches)

Example of technical noise: Diffused light

Optical element (mirror, lens, ...) vibrating due to seismic or acoustic noises



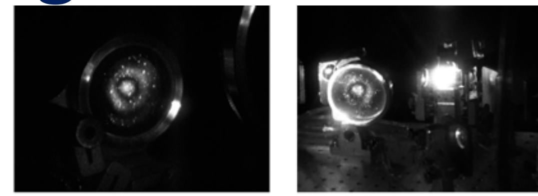
some photons of the diffused light gets recombined with the interferometer beam

↓

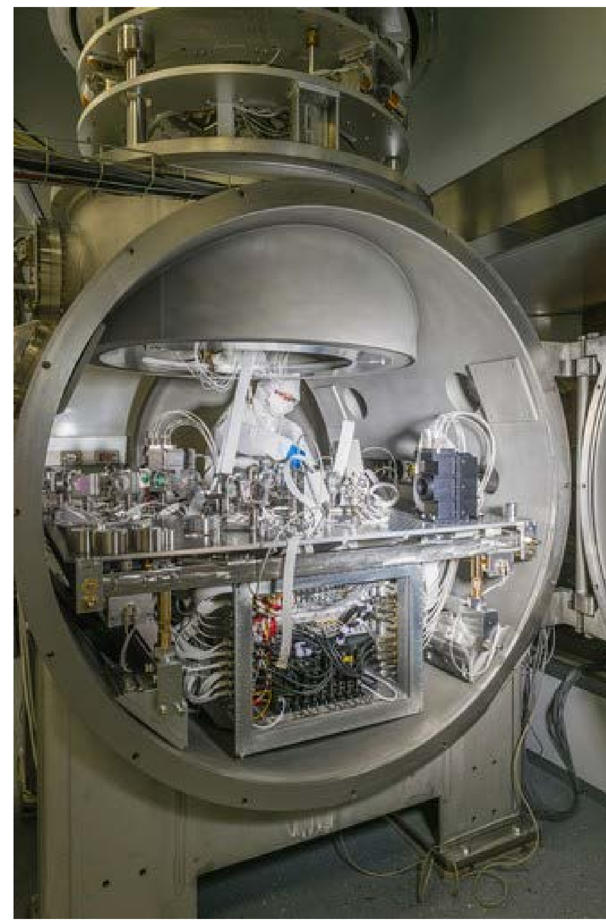
phase noise

↓

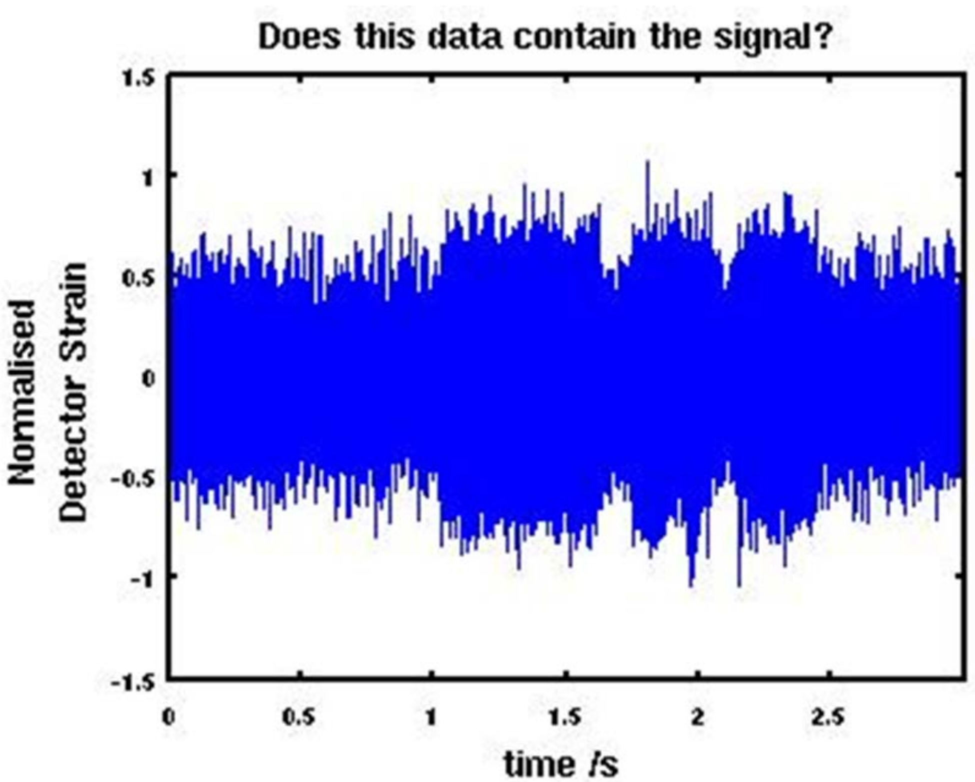
extra power fluctuations
(imprint of the optical element vibrations)



Evolution for AdVirgo: suspend the optical benches and place them under vacuum



Noises are not always stationary



“Glitches” are impulses of noise.
They might look like a transient GW signal

- environmental disturbances monitored with an array of sensors: seismic activities, magnetic perturbations, acoustic noises, temperature, humidity → used to veto false alarm triggers due to instrumental artifacts

- requires coincidence between 2 detectors to reduce false alarm rate

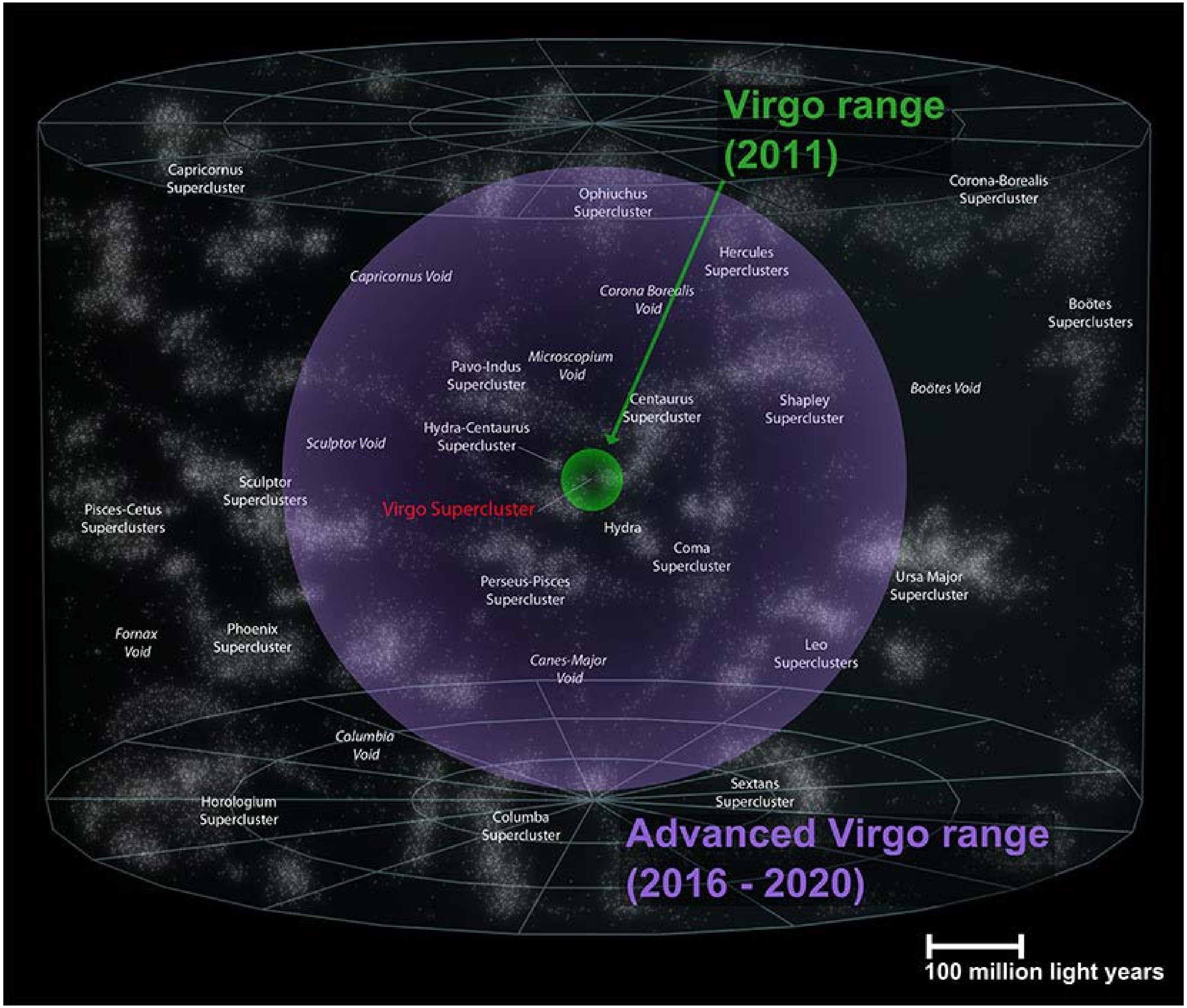
Table of Contents

- **How can we detect gravitational waves with laser interferometers?**

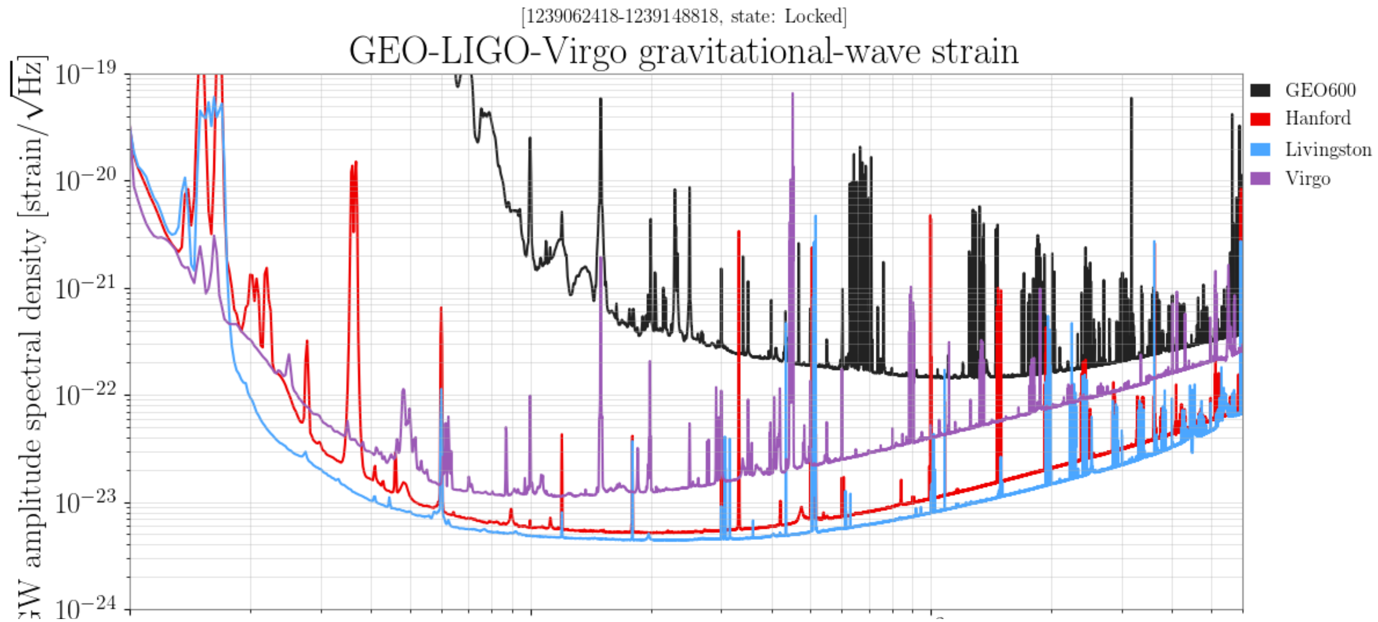
- **How do ground-based interferometers work?**
  - ➢ The Virgo optical configuration or how to measure 10^{-20} m
  - ➢ How to maintain the ITF at its working point?
  - ➢ How to measure the GW strain h(t) from this detector?
  - ➢ Noises limiting the ITF sensitivity: how to tackle them?

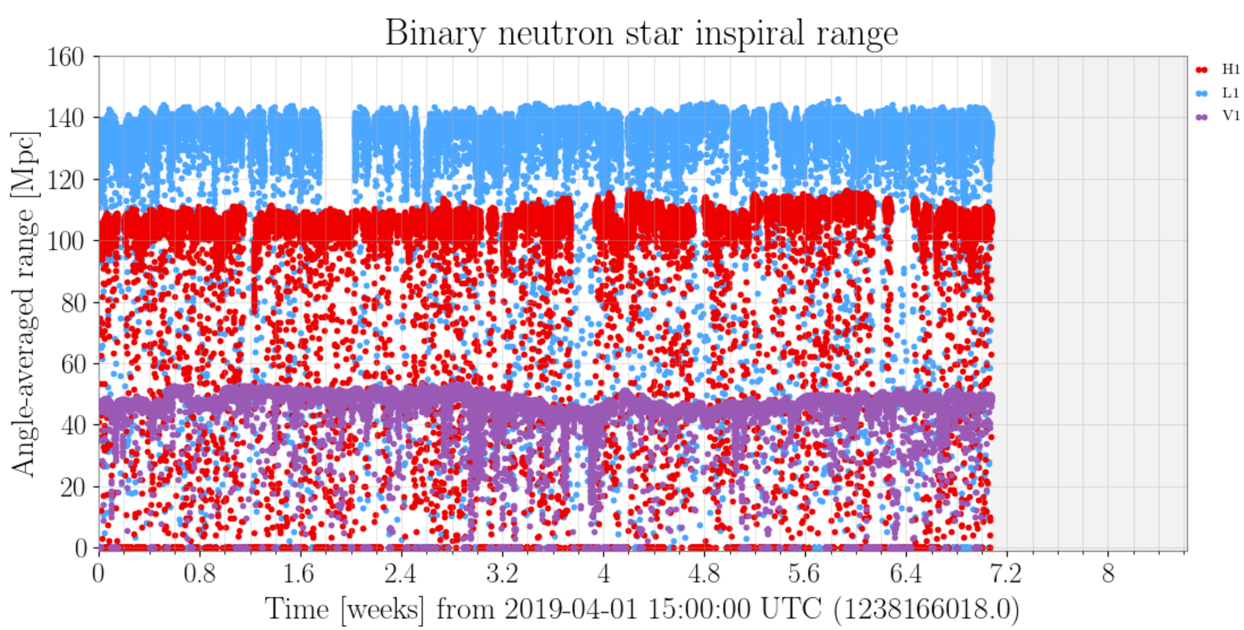
- **Prospectives for interferometers and other detectors**

From initial to advanced detectors



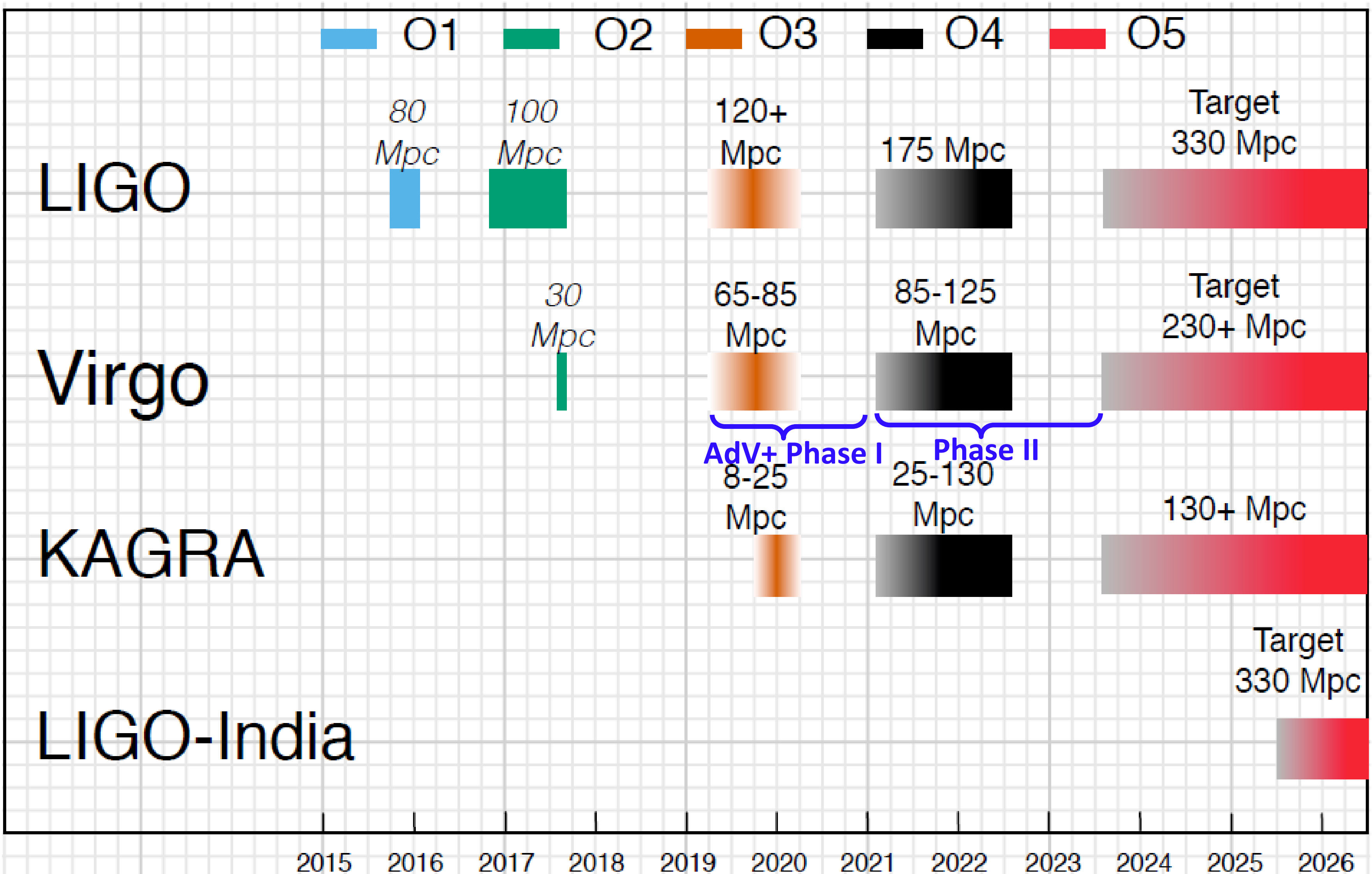
Current interferometers sensitivity





BNS Range:
Distance at which a neutron star binary coalescence with averaged orientation over the sky can be seen with signal-to-noise ratio of 8

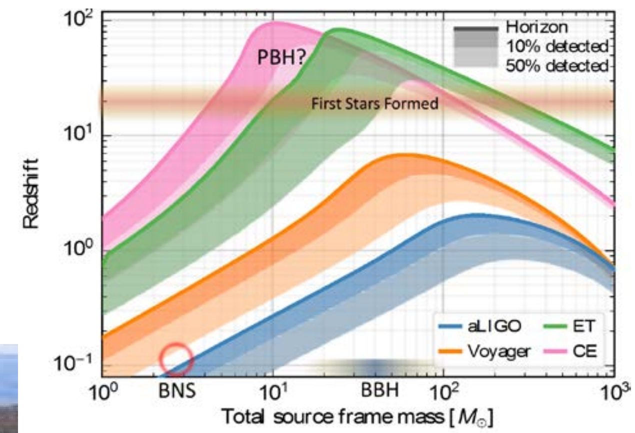
Future observing runs

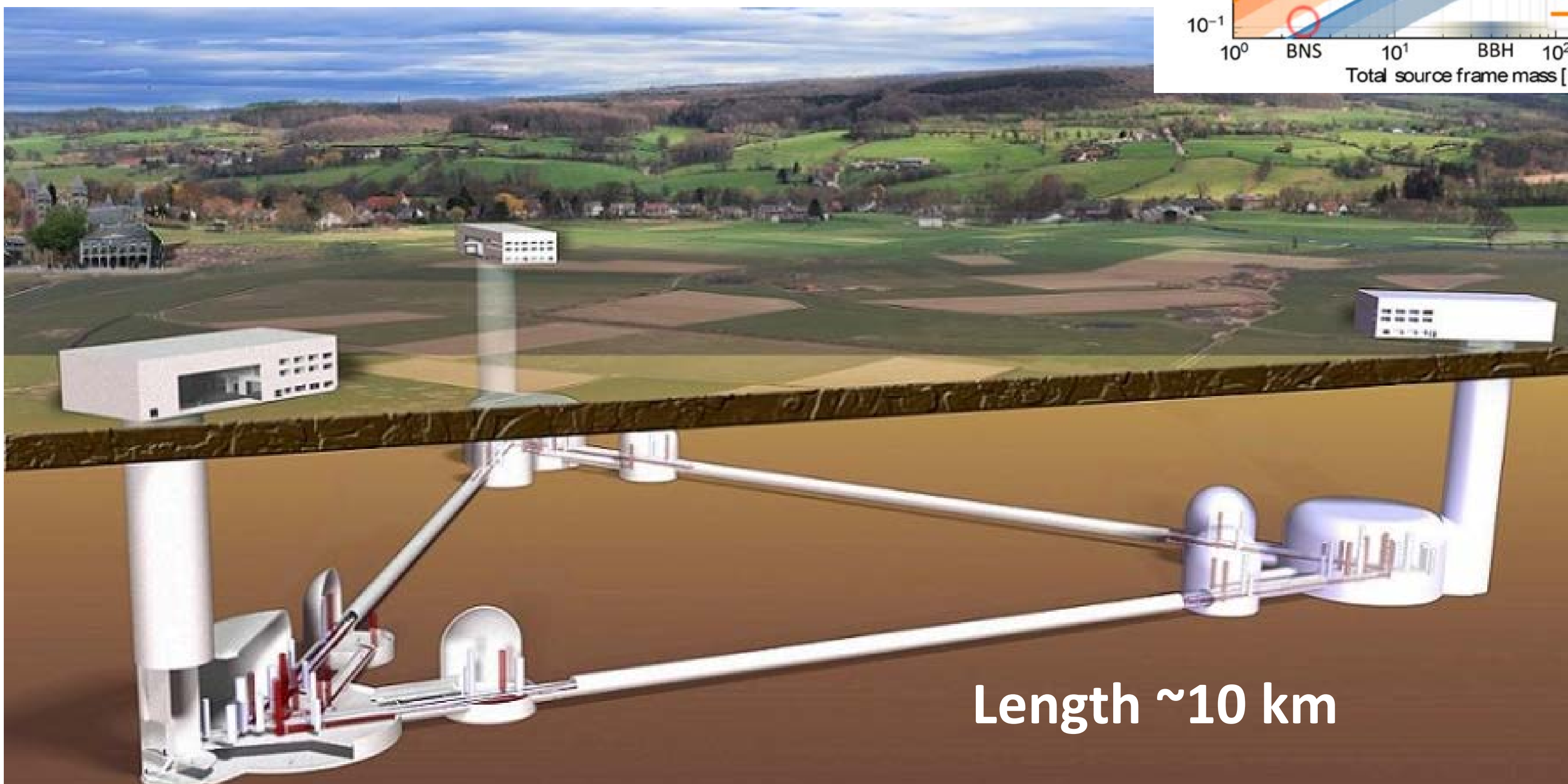


Einstein Telescope

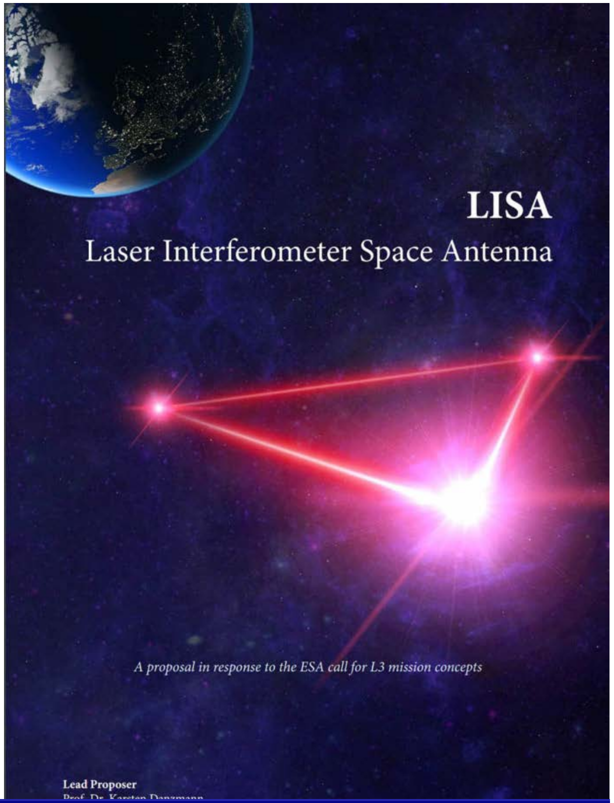
- Third generation interferometer: gain another factor 10 in sensitivity and enlarge bandwidth
- Located underground, ~10 km arms
- Thermal noise reduction with cryogenics
- Xylophone detector?
- In operation after 2030?

Could probe CBC signals from a large fraction of the Universe





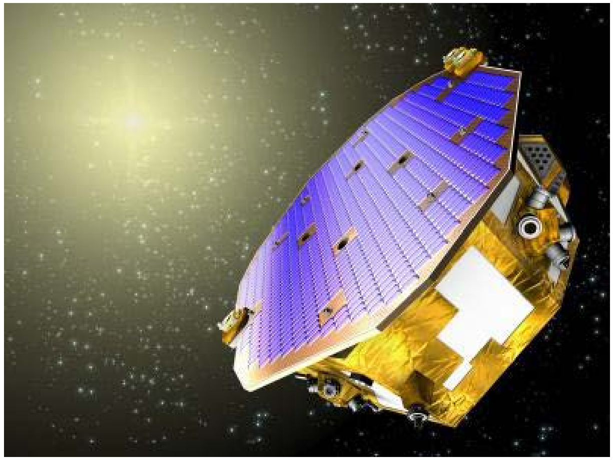
Spatial interferometer: LISA

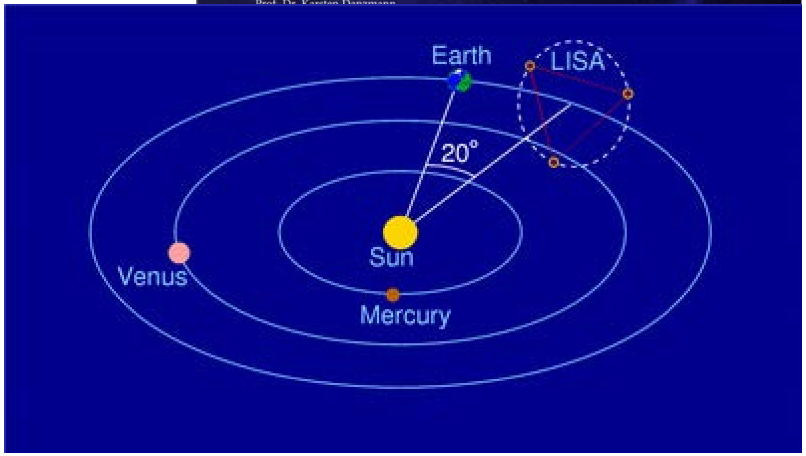


- **Bandwidth: 0.1 mHz to 1 Hz (2.5 million km arm length)**
- Launch of LISA in the years 2030?

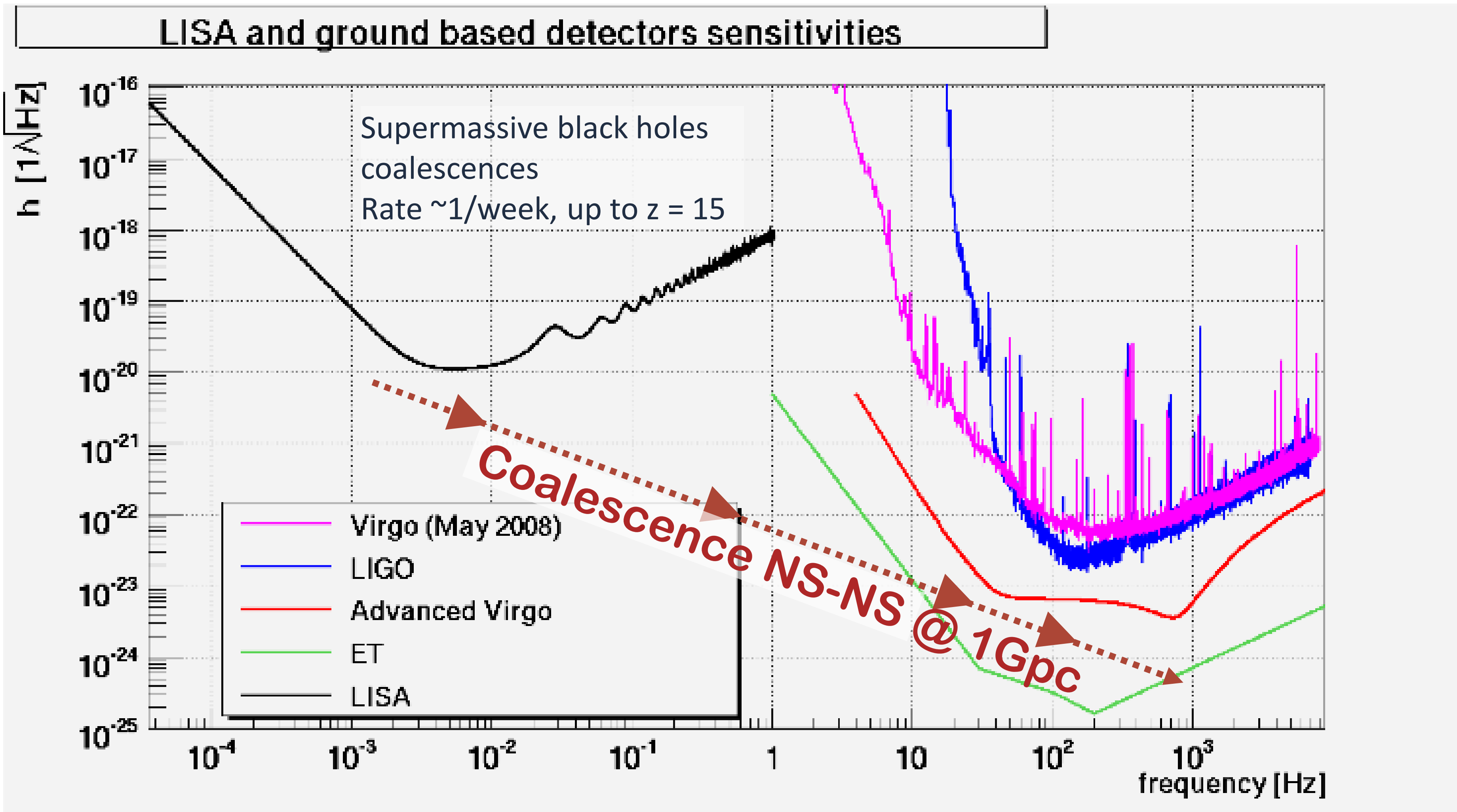
  → operation for 5 to 10 years

- Successful intermediate step: LISA Pathfinder
  - ➢ launched end 2015
  - ➢ test of free-fall masses
  - ➢ validation of differential motion measurements



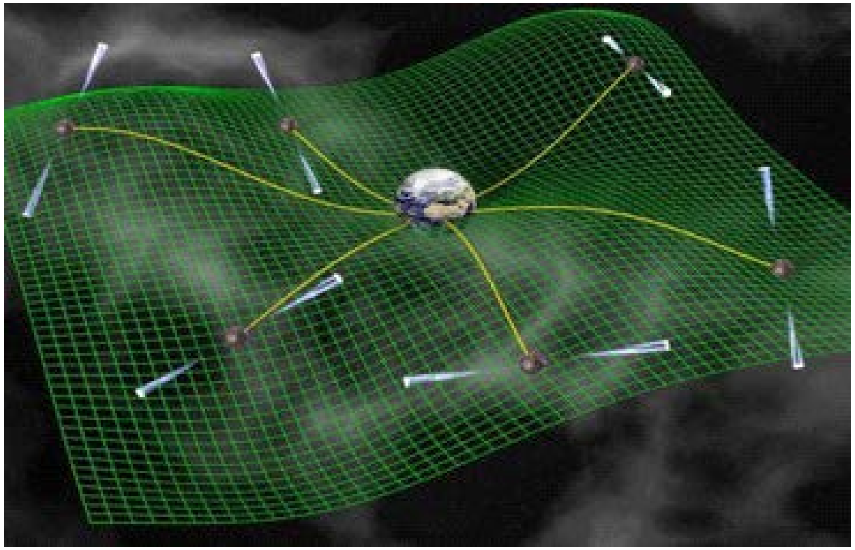


ET and LISA performances



Pulsars timing arrays

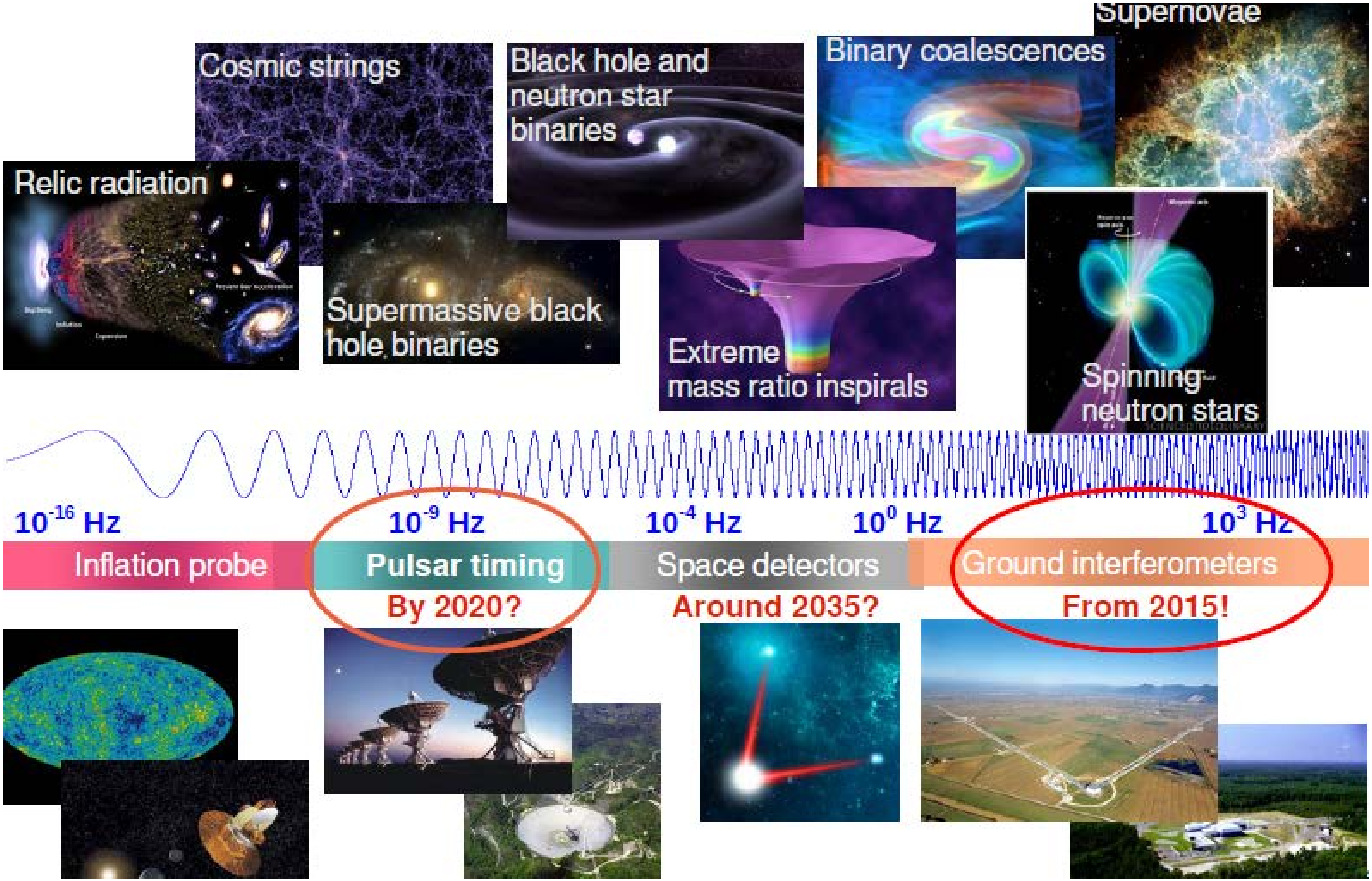
- **Bandwidth: 1 nHz to 10 nHz**
- Observation of 20 ms pulsars in radio
  - Residuals of modellisation < 100 ns
  - Weekly sampling over 5 years



- International network
  - Parkes PTA
  - North American NanoHertz Gravitationnal Wave Observatory
  - European PTA

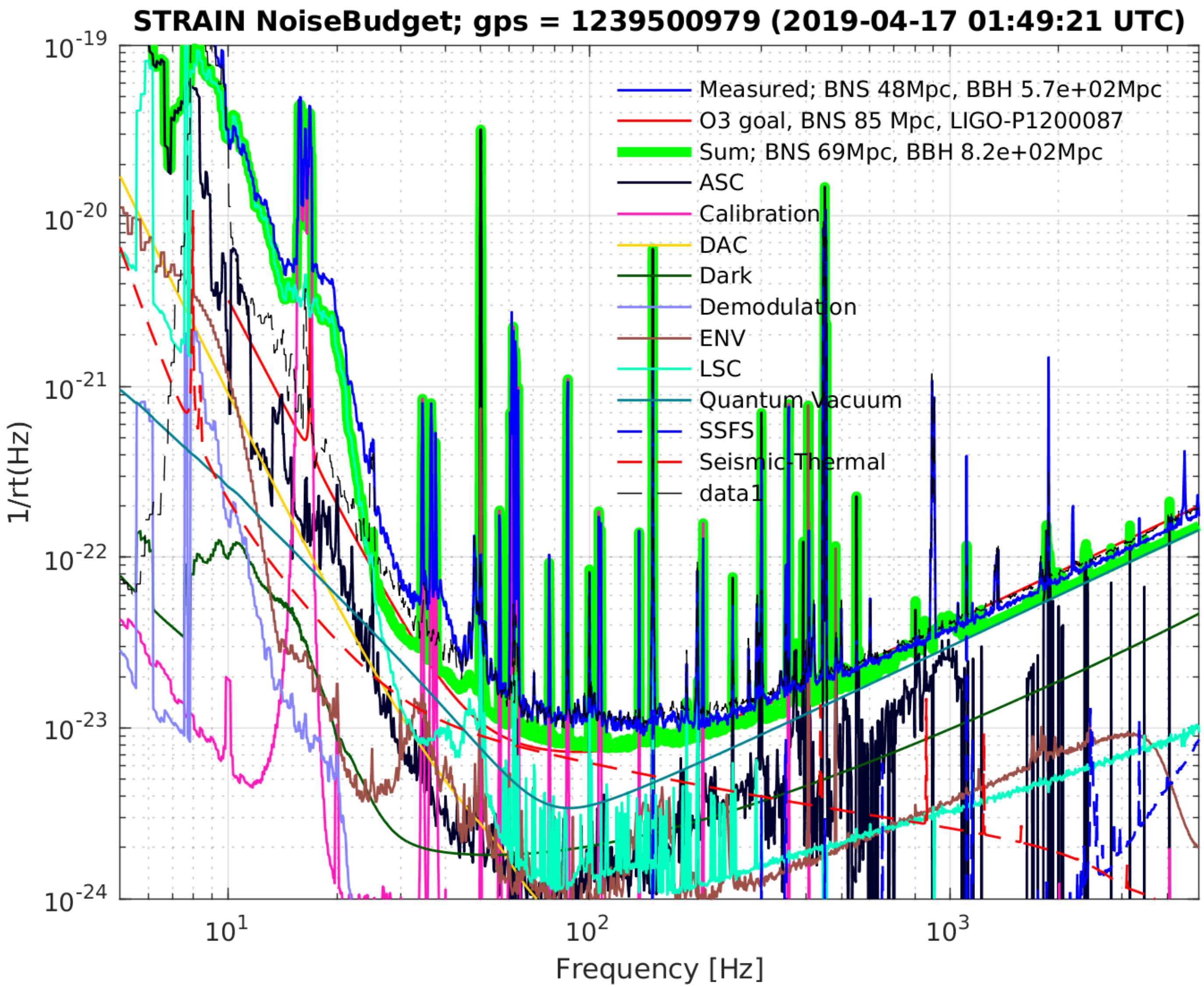
- **First detections expected in the coming years!**

A large GW spectrum to be studied…



SPARES

Example of Advanced Virgo noise budget (O3 run)



Virgo data acquisition summary

- Photodiodes
- Currents in the mirror actuators
- Seismometers
  Magnetometers
  Thermometers
- Cameras
- **Sensors....**

↓

GPS-synchronised digitisation
Sampling frequencies from 1 Hz to ~50 kHz
~2500 sensor signals to be digitised

↓

Data acquisition processes
(data formatting, timestamp, write to disks..)
~ 10000 channels

→

Continuous flow of ~3 TBytes/day (40 to 60 MBytes/s)
Disk space on Virgo site: ~400 TB for 4 months of data

Longer storage: data sent via Ethernet to computing centers (Lyon, Bologna)

Output Mode Cleaner

- 2 bow-tie Fabry Perot cavities:
  - Get rid of high order modes and controls signals.