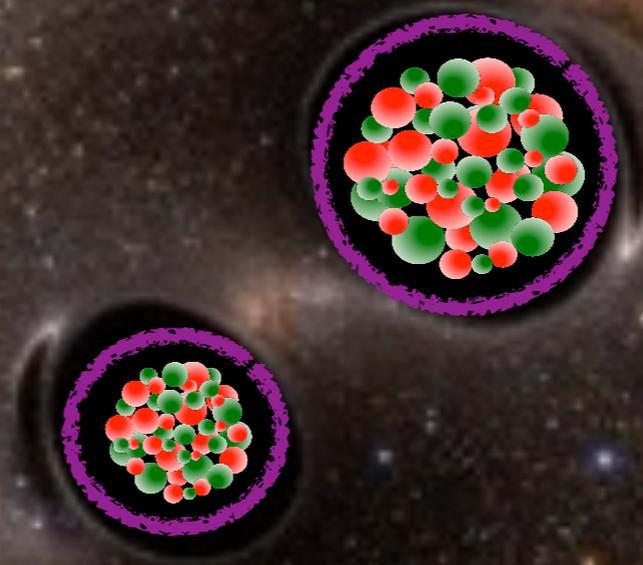


Microstate Geometries and the Microstructure of Black Holes



IPhT colloquium

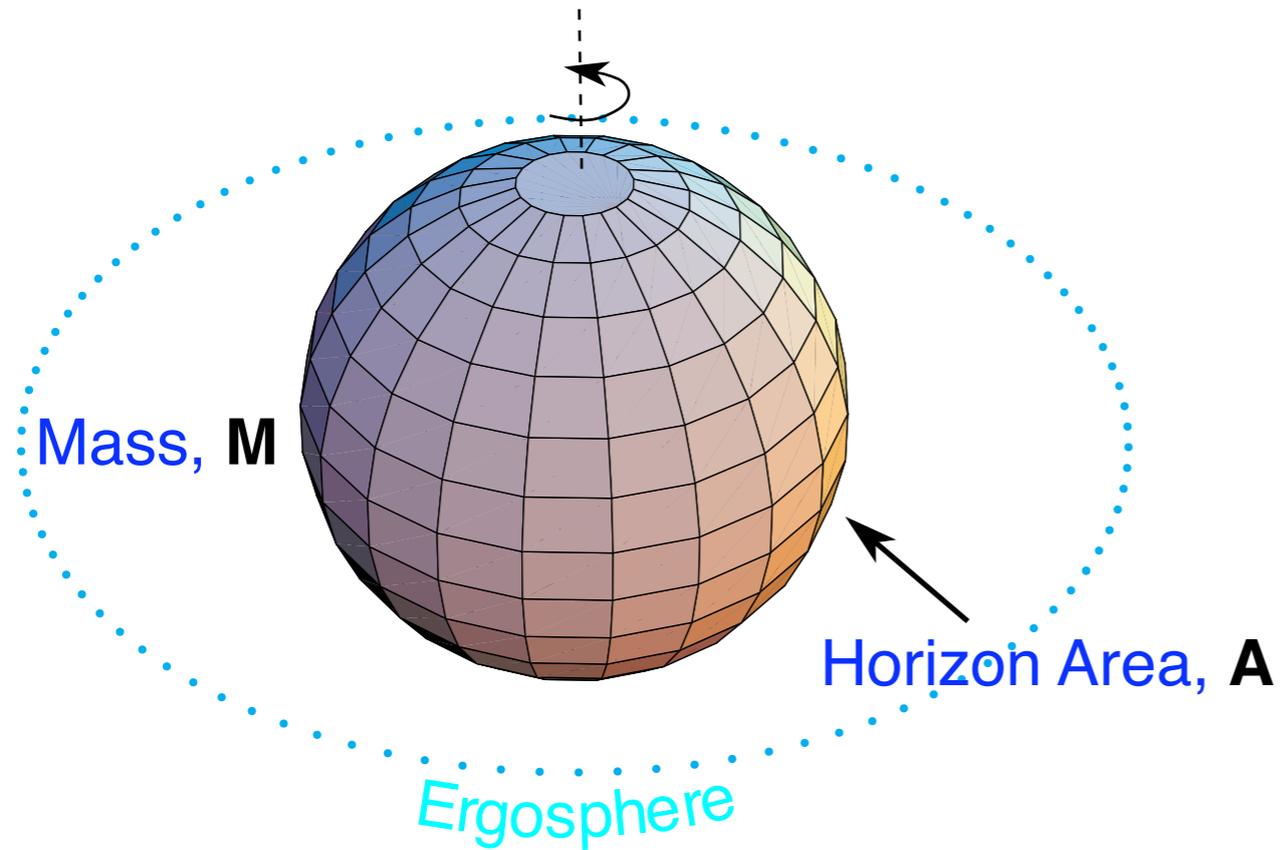
Nick Warner, October 15, 2018

Research supported supported in part by DOE grant DE- SC0011687

Original photo credit:
LIGO/Caltech

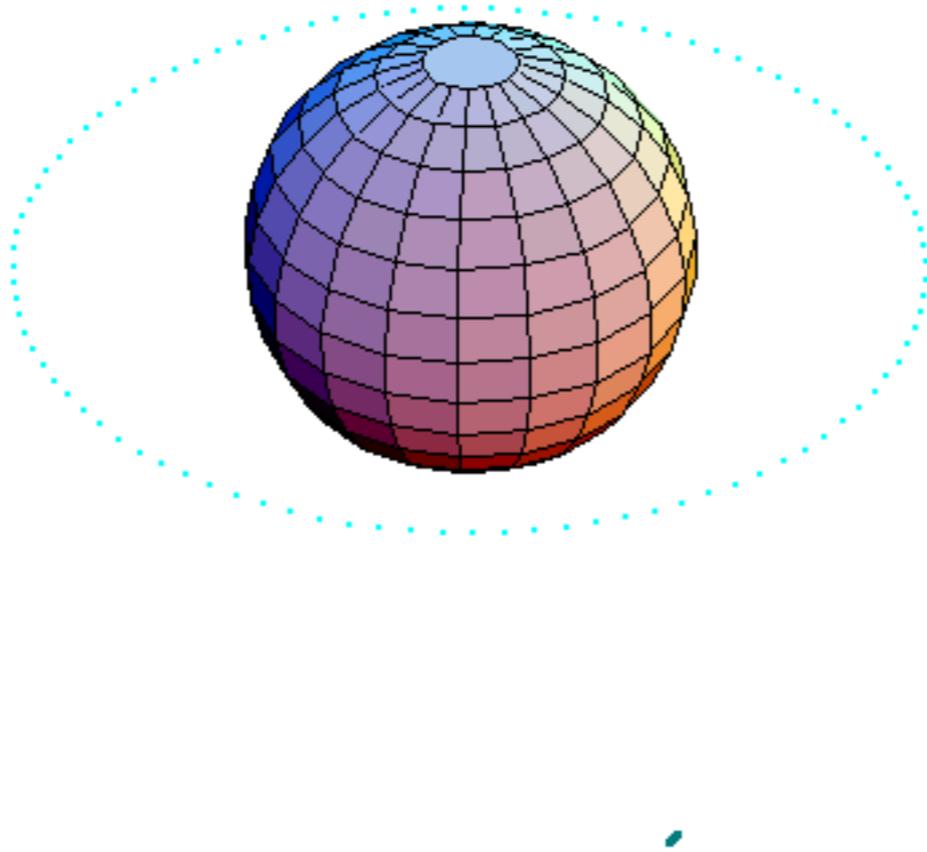
Black-hole thermodynamics and the Penrose Process

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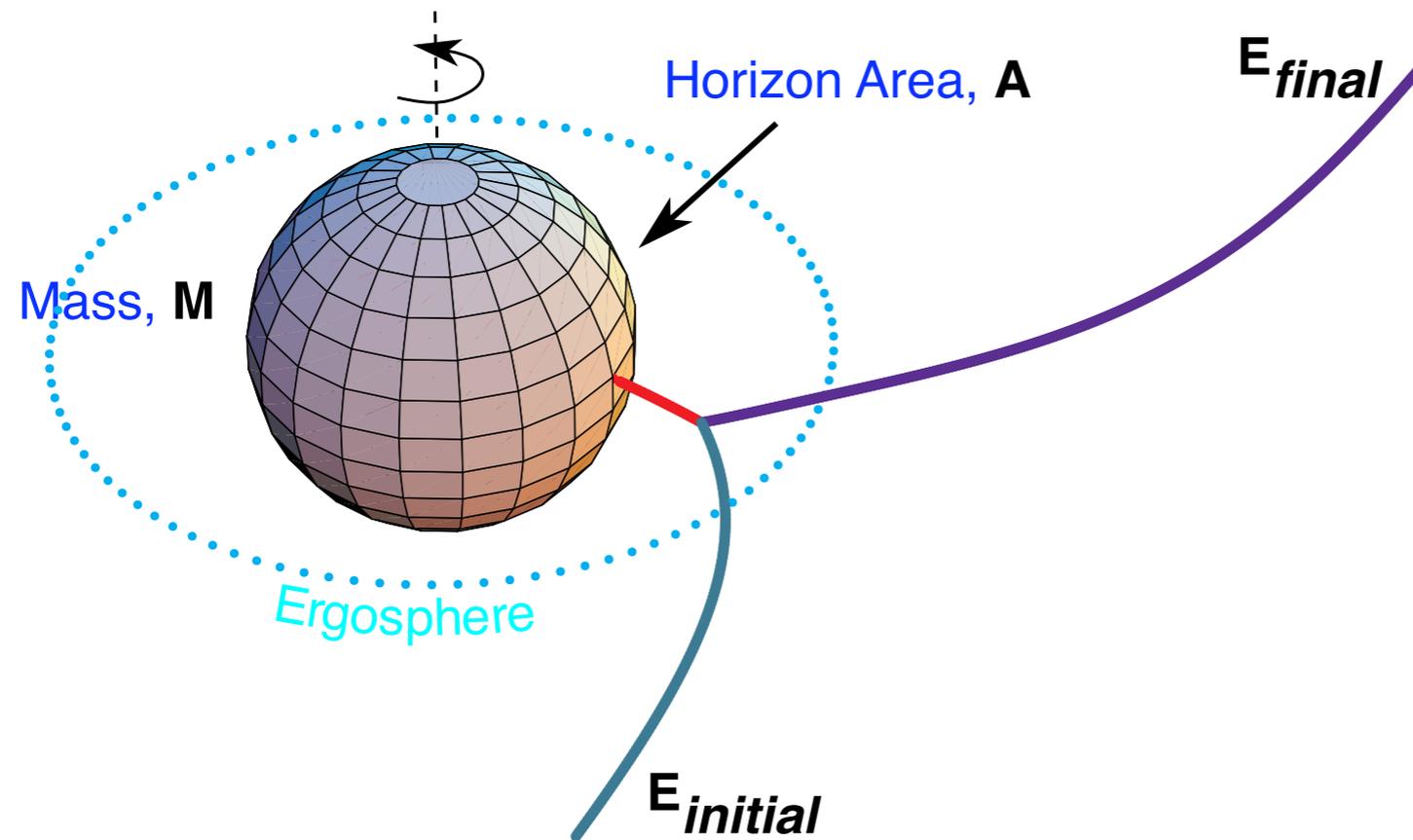
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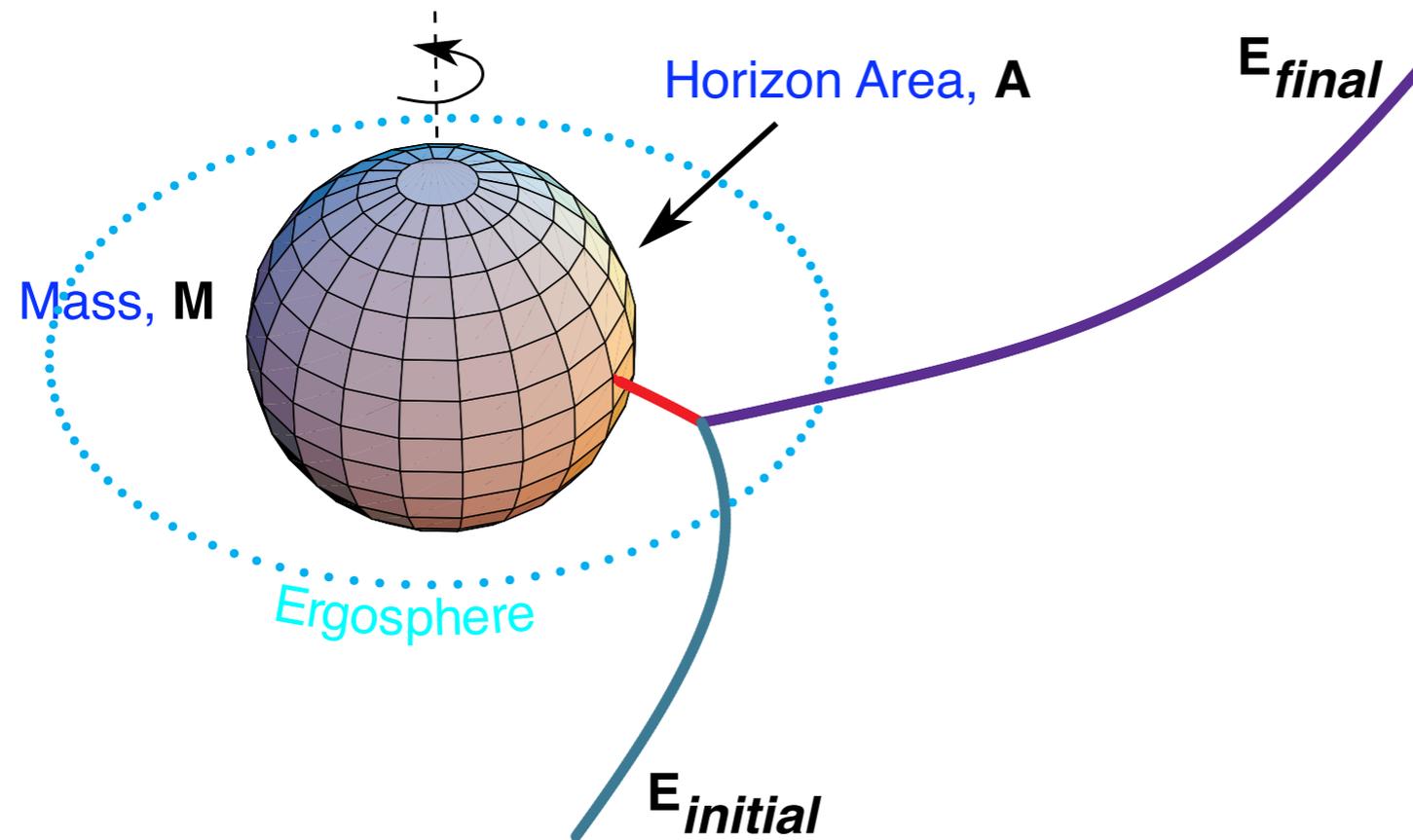


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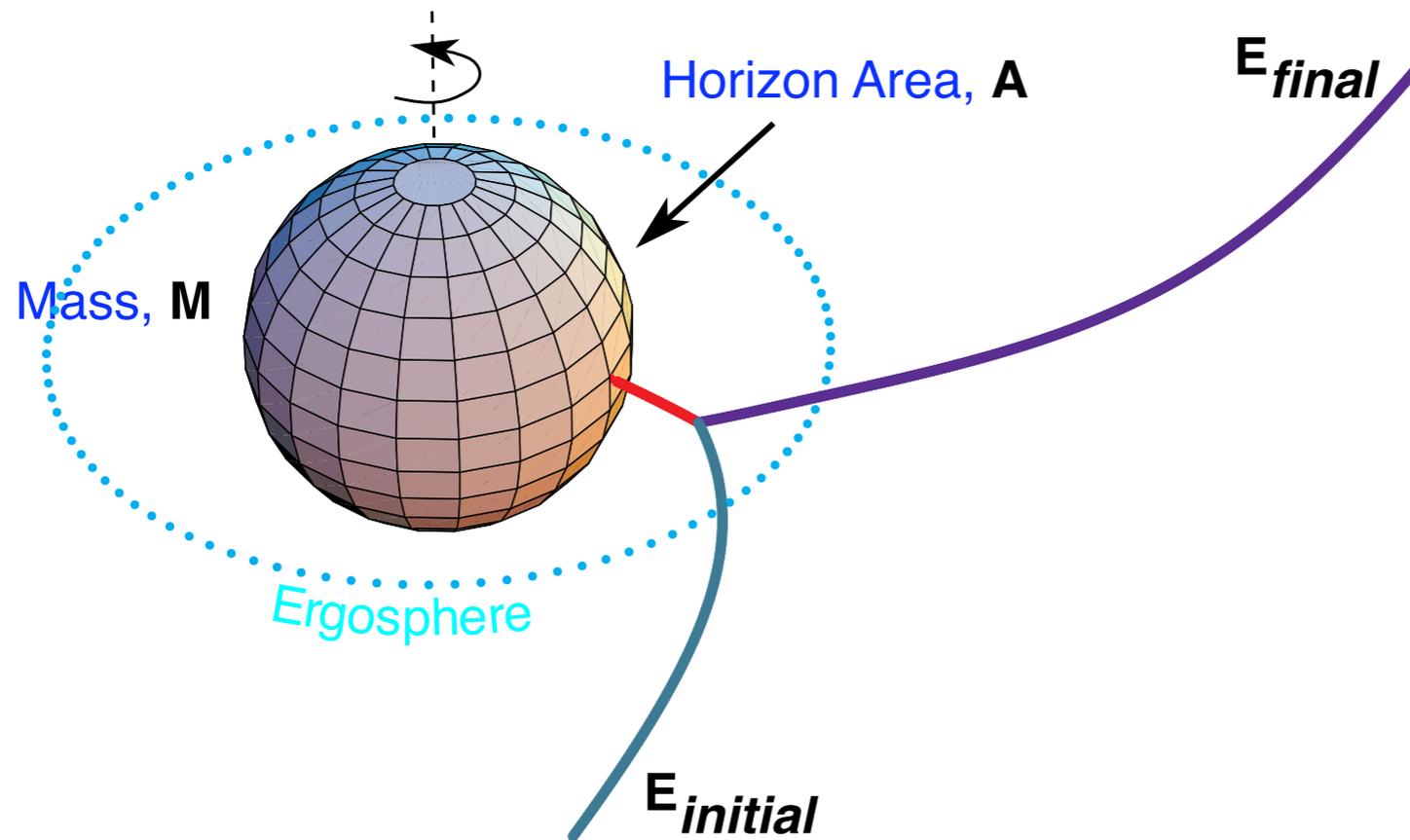
- Work done by black hole:

$$dW = E_{final} - E_{initial}$$

Accompanied by a change in

- Mass, **M**, of the black hole
- Horizon Area, **A**

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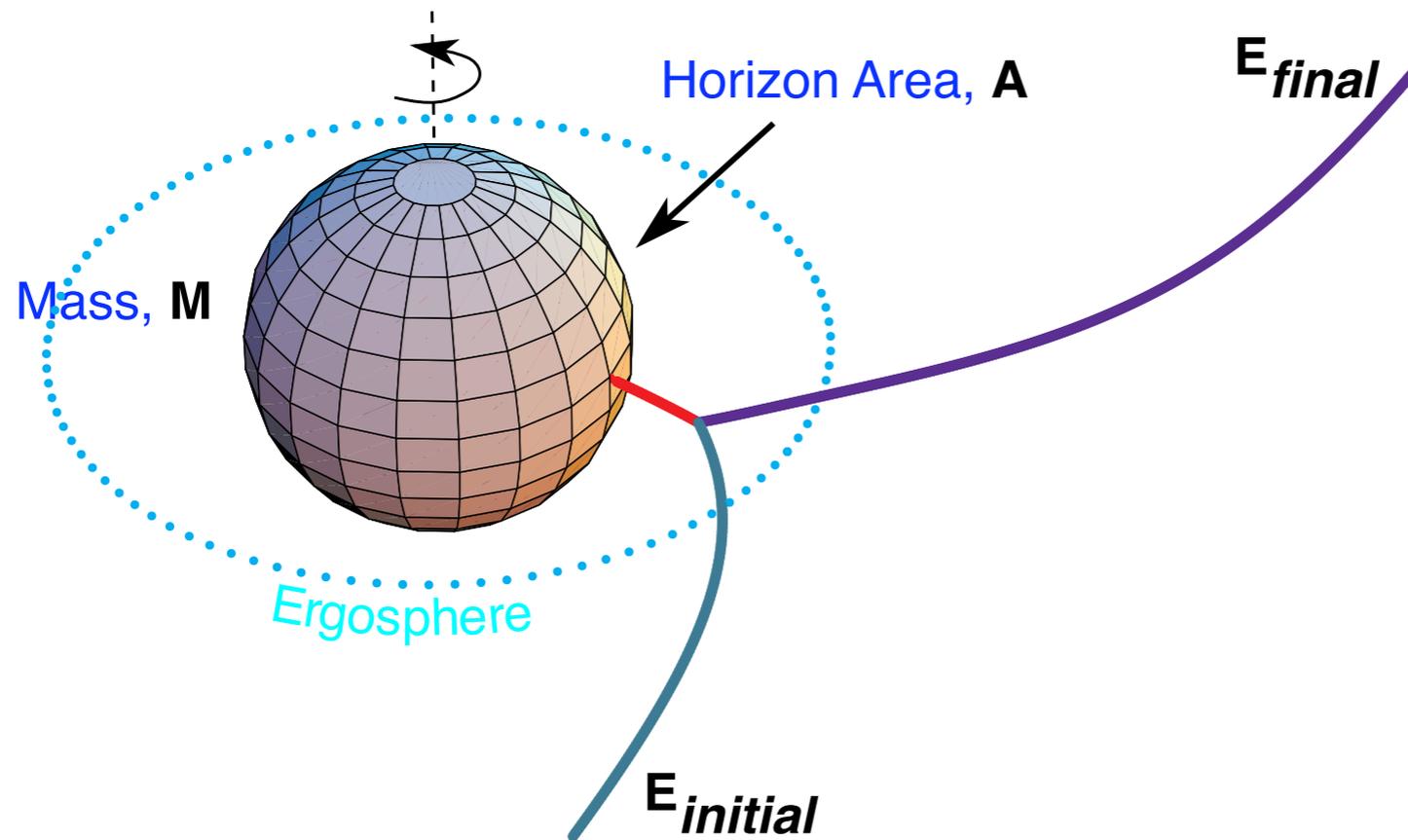
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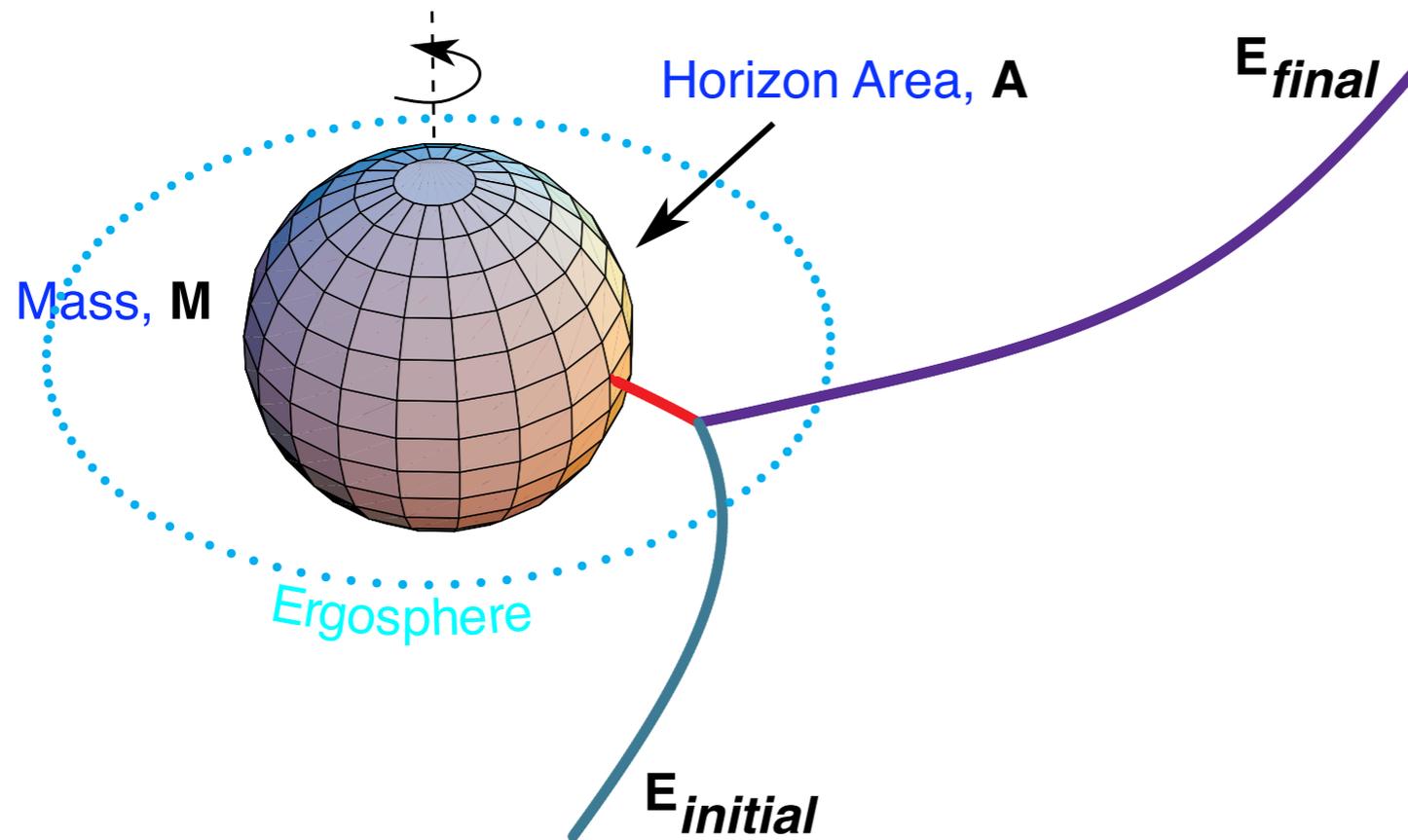
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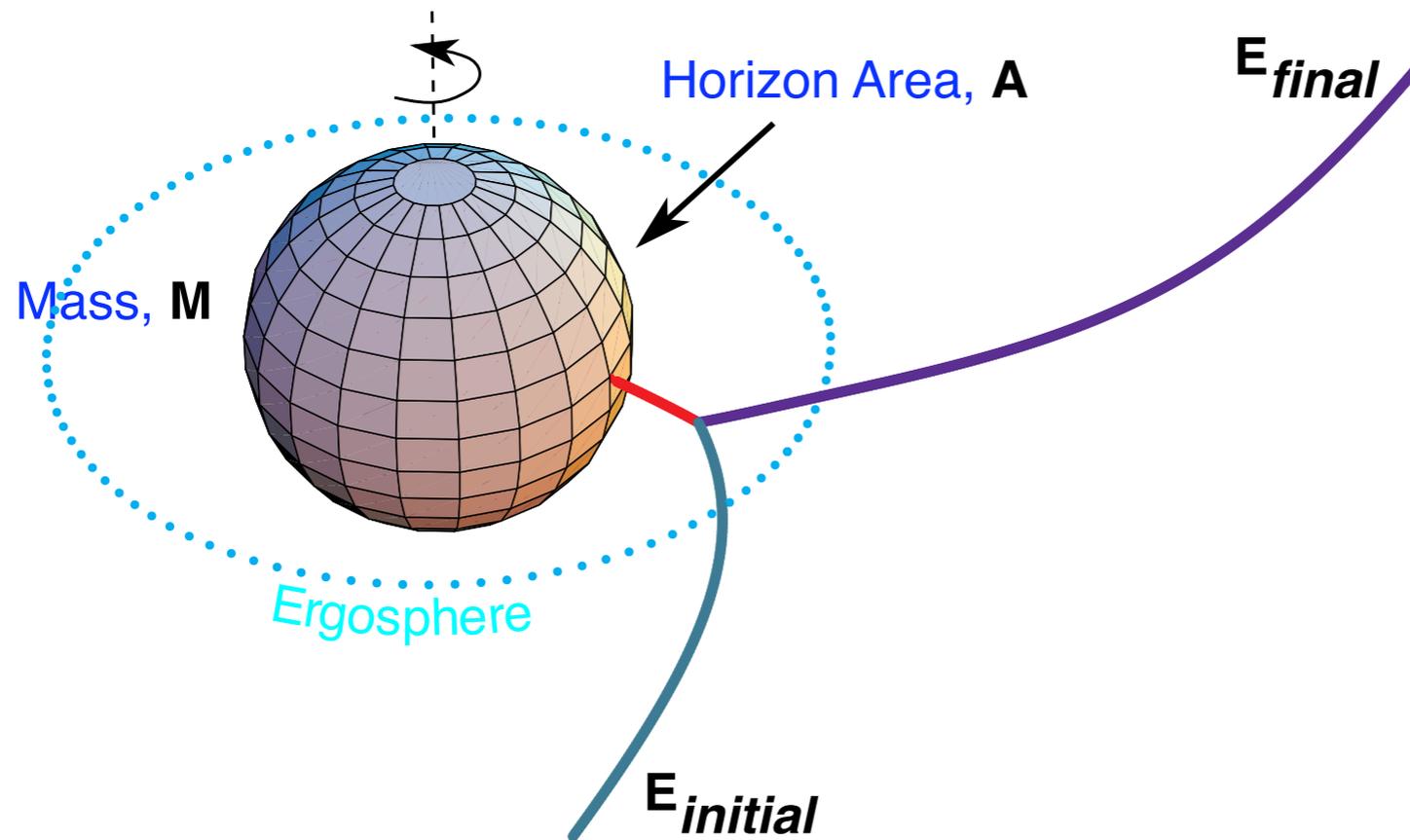
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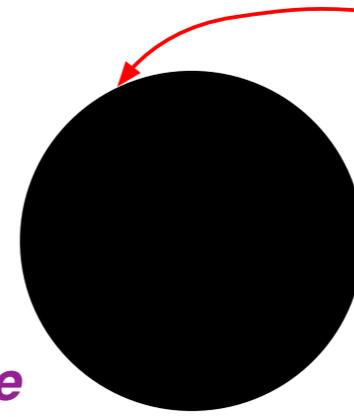
Reversible, Carnot ideal: $dA = 0 \Leftrightarrow$ Infalling particle grazes horizon

Microstates?

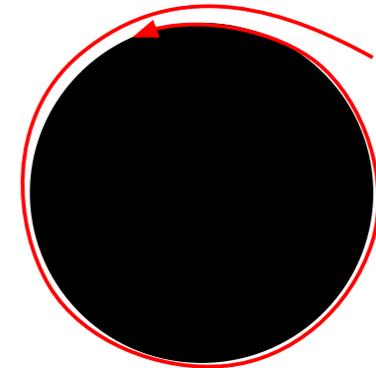
Microstates?

Reversible Infall: $dA = 0$ and
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Irreversible



$dA > 0$



$dA \rightarrow 0$

*Nearly
reversible*

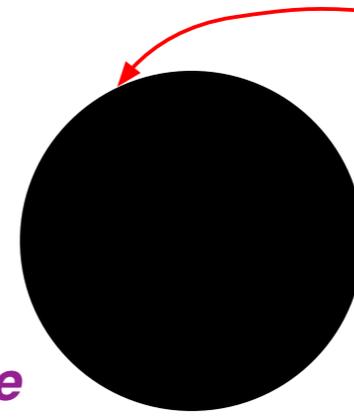
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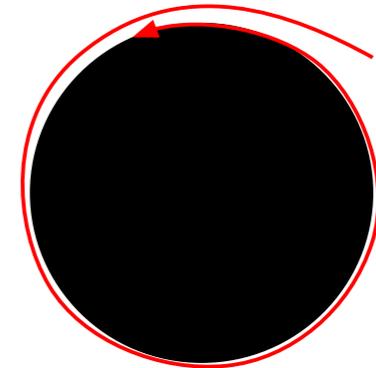
Bekenstein:

- Particle must have finite size \sim Compton Wavelength \Rightarrow non-grazing $\Rightarrow dA > 0$
- Particle falls in \Rightarrow Minimal loss of information, $\Delta S = \ln 2$
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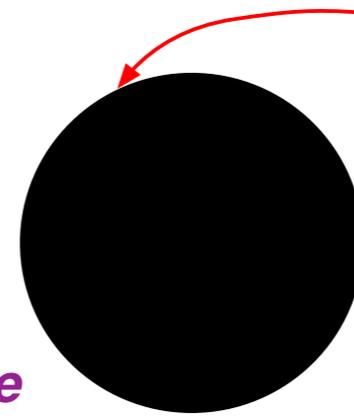
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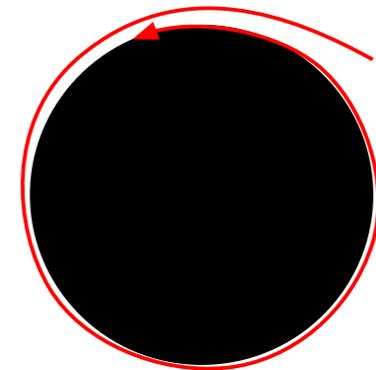
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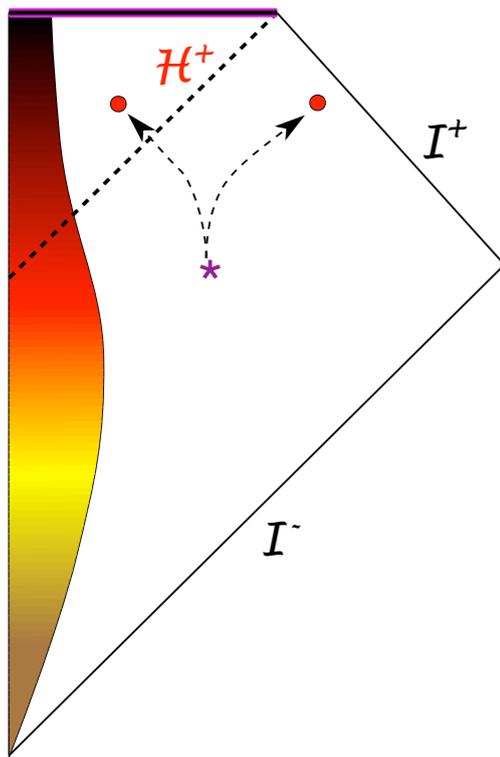
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Hawking radiation

Quantum analog of Penrose process:

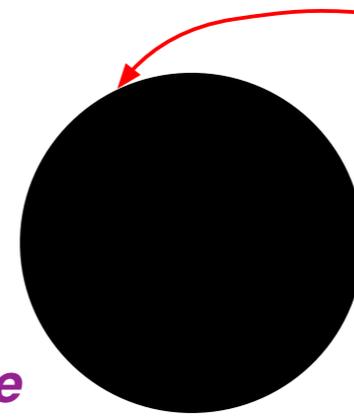
Negative energy falls in; Positive energy escapes to infinity



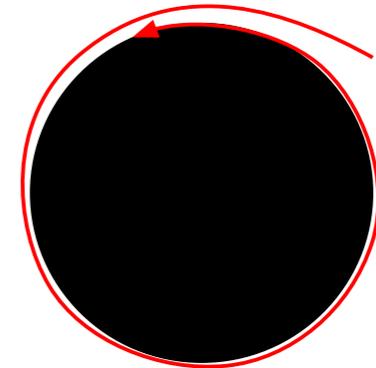
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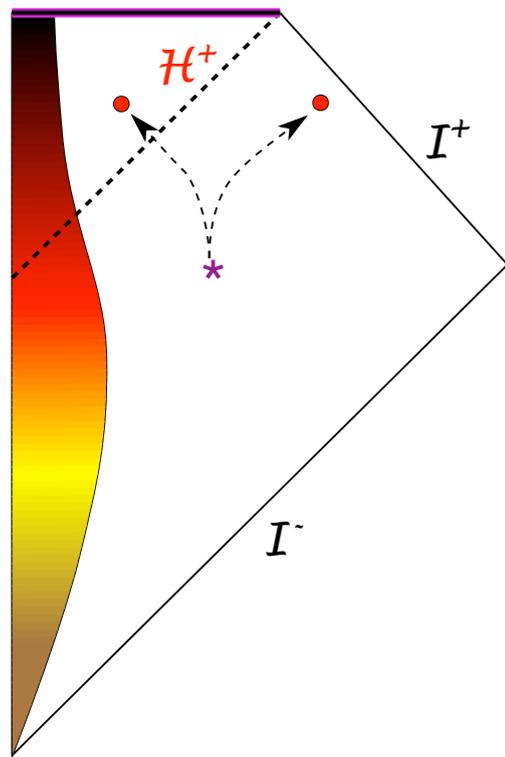
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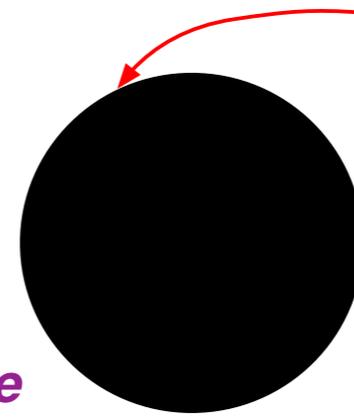
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$$T = \frac{\kappa}{2\pi} = \frac{\hbar c^3}{8\pi G k_B M}$$

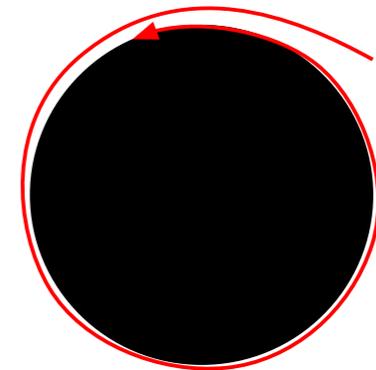
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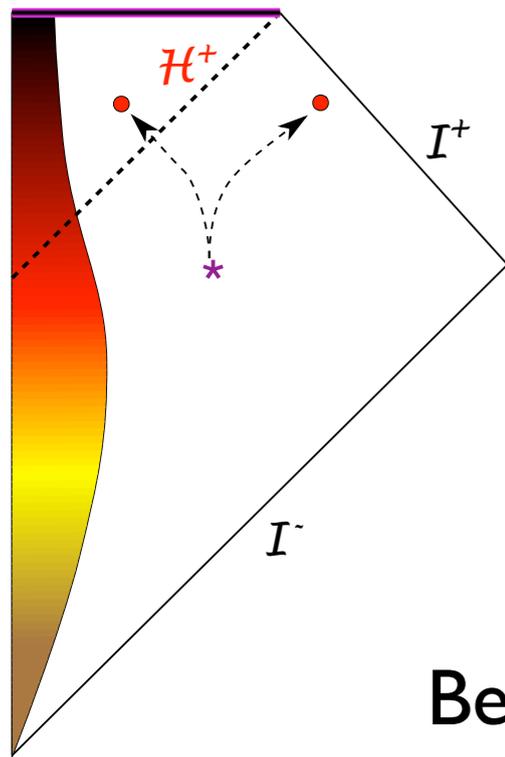
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Bekenstein-Hawking entropy:
$$S = \frac{k_B c^3}{4 G \hbar} A = \frac{1}{4} \frac{A}{\ell_P^2}$$

Black-hole uniqueness

In General Relativity coupled to electromagnetism in 3+1 dimensions:

Bulk State Functions (**M** + **angular momentum** + **conserved charges**)

uniquely specify the solution/metric outside the horizon

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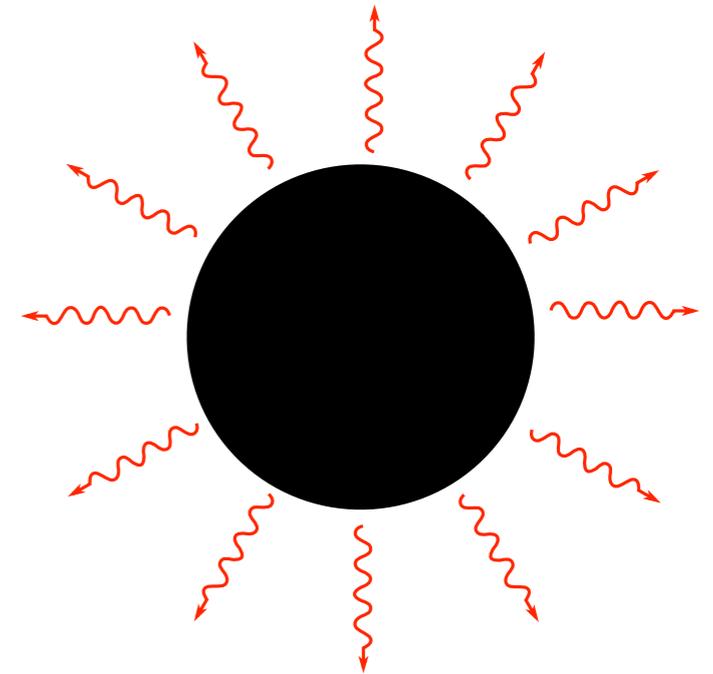
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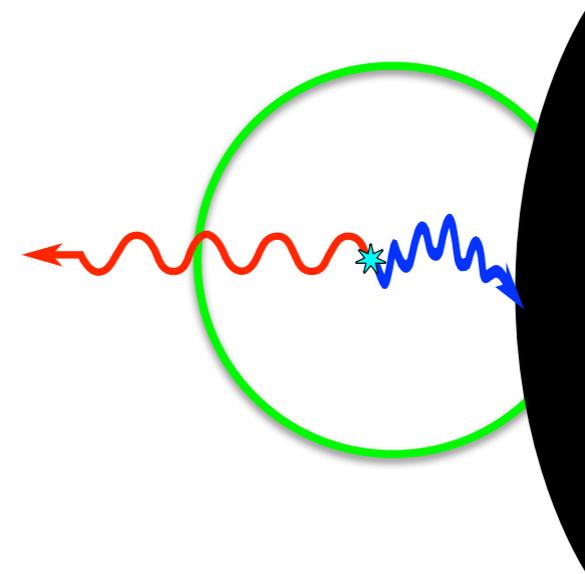
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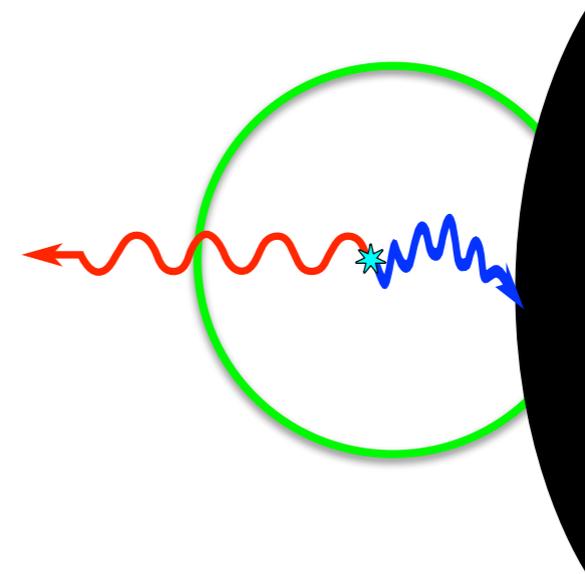
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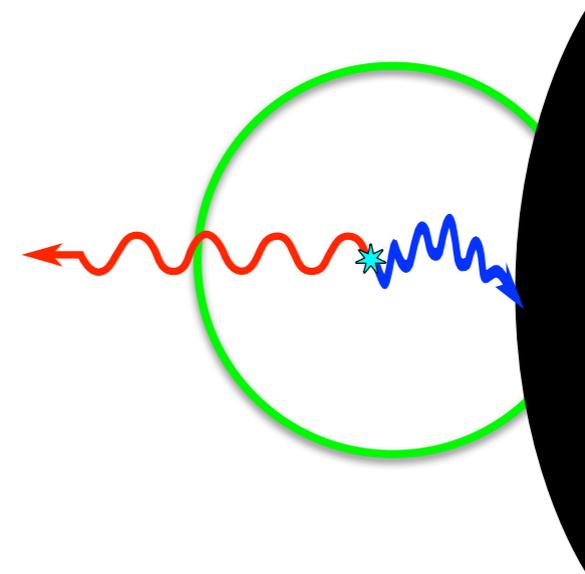
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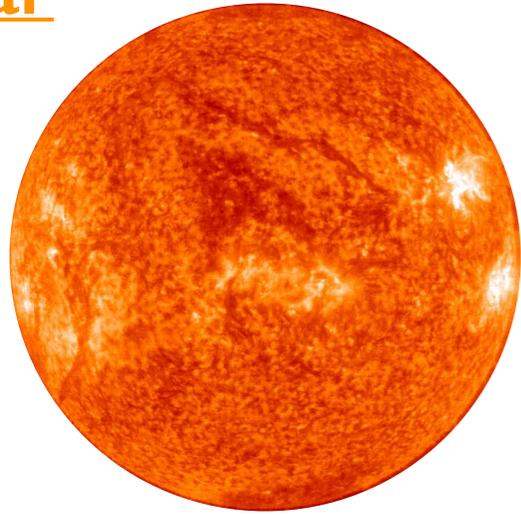


General Relativity + Quantum Mechanics

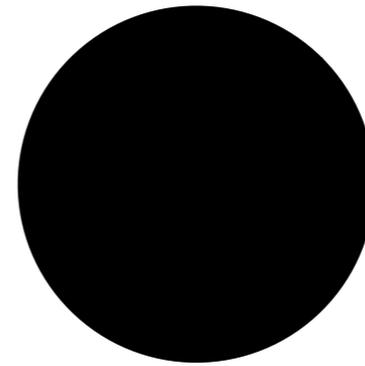
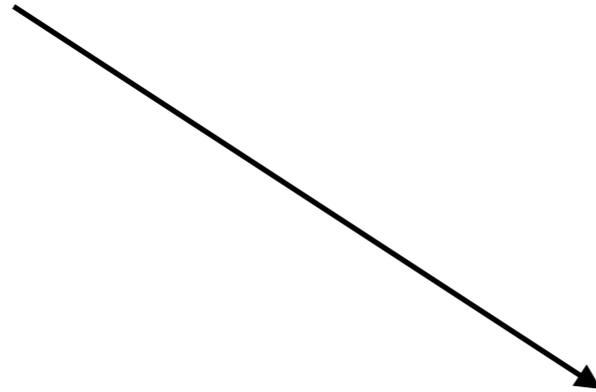
⇒ *Black holes, no matter how they form, evaporate into the same (largely featureless) cloud of Hawking Radiation*

The exterior structure of a black hole, and
the *Hawking radiation* is unique

A star

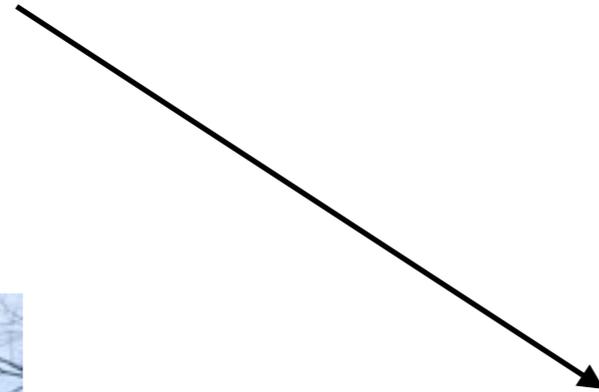
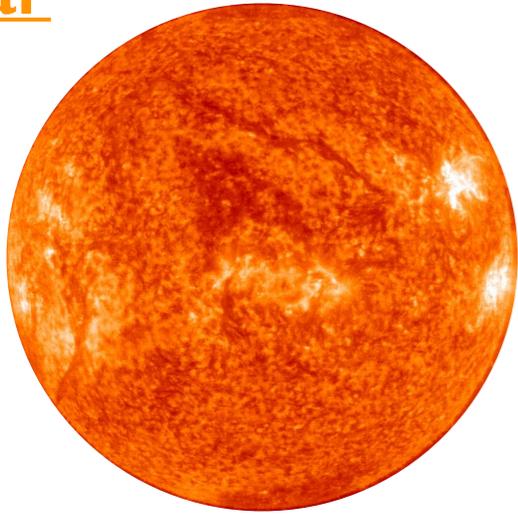


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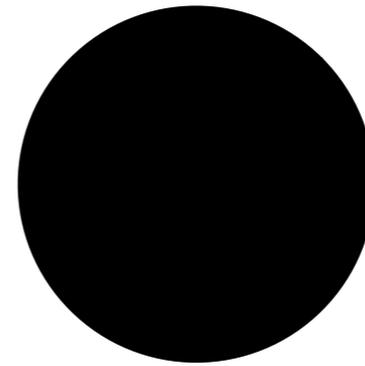


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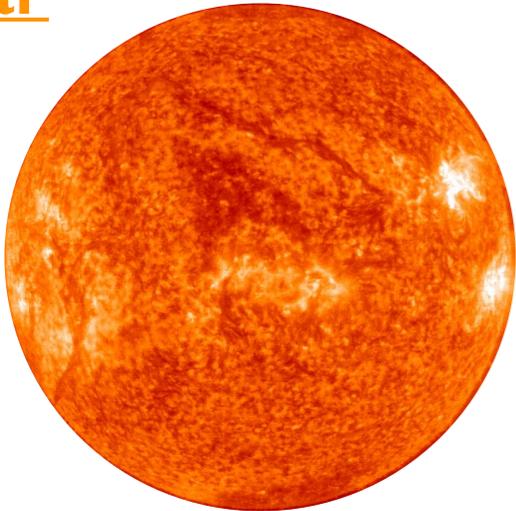


Trash



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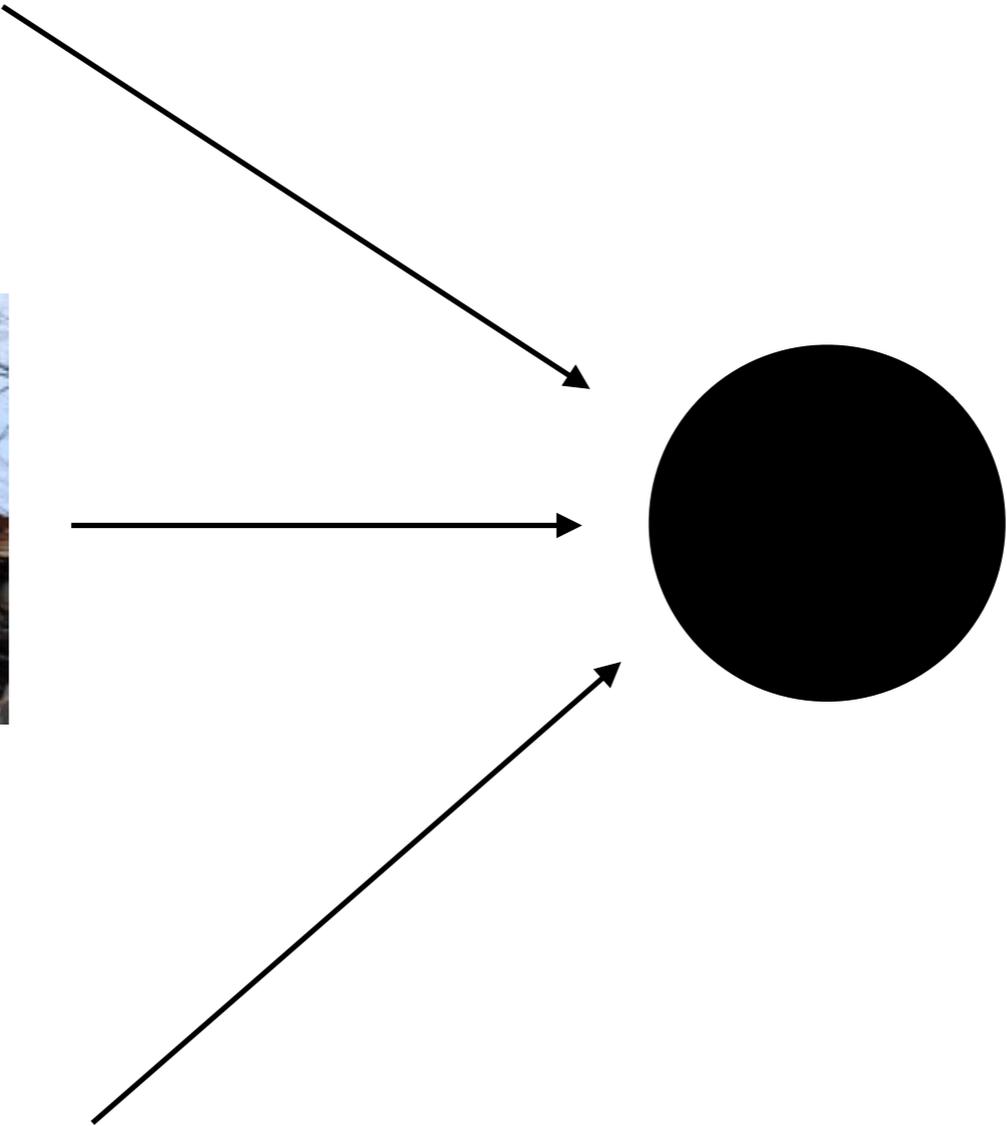
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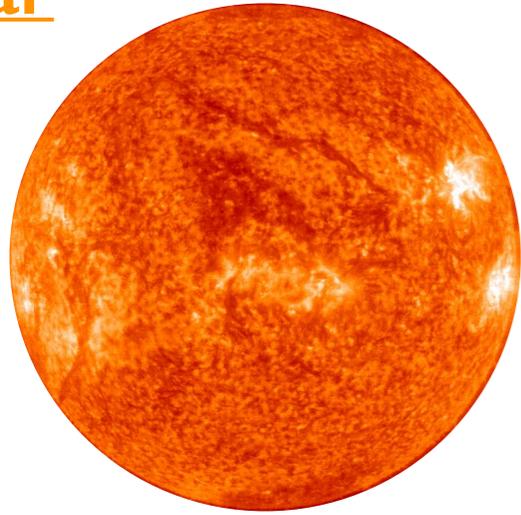
Trash



Politicians

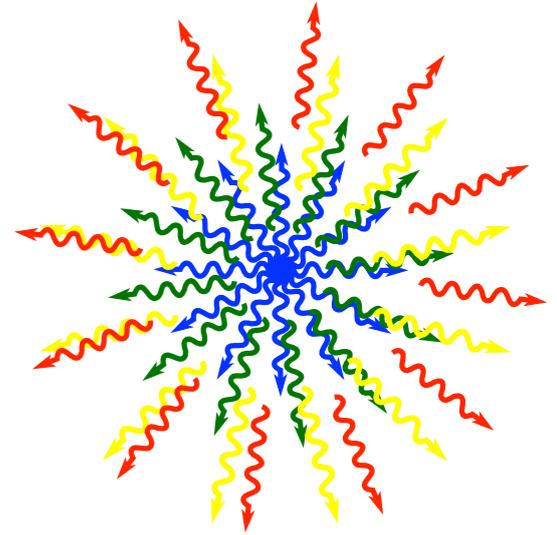


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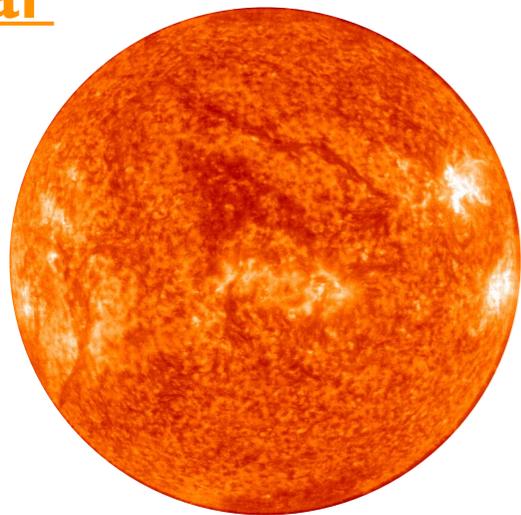
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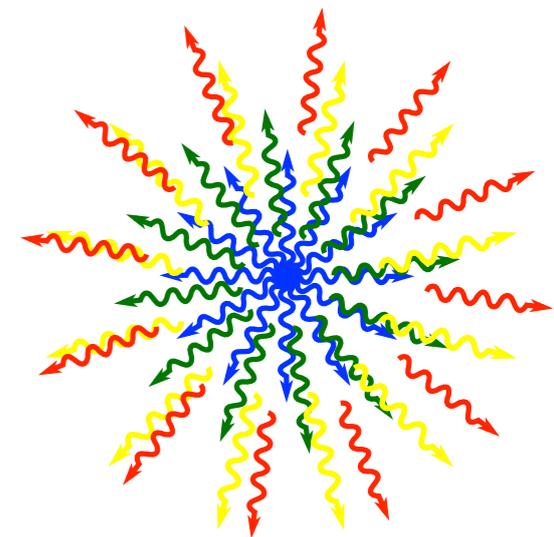
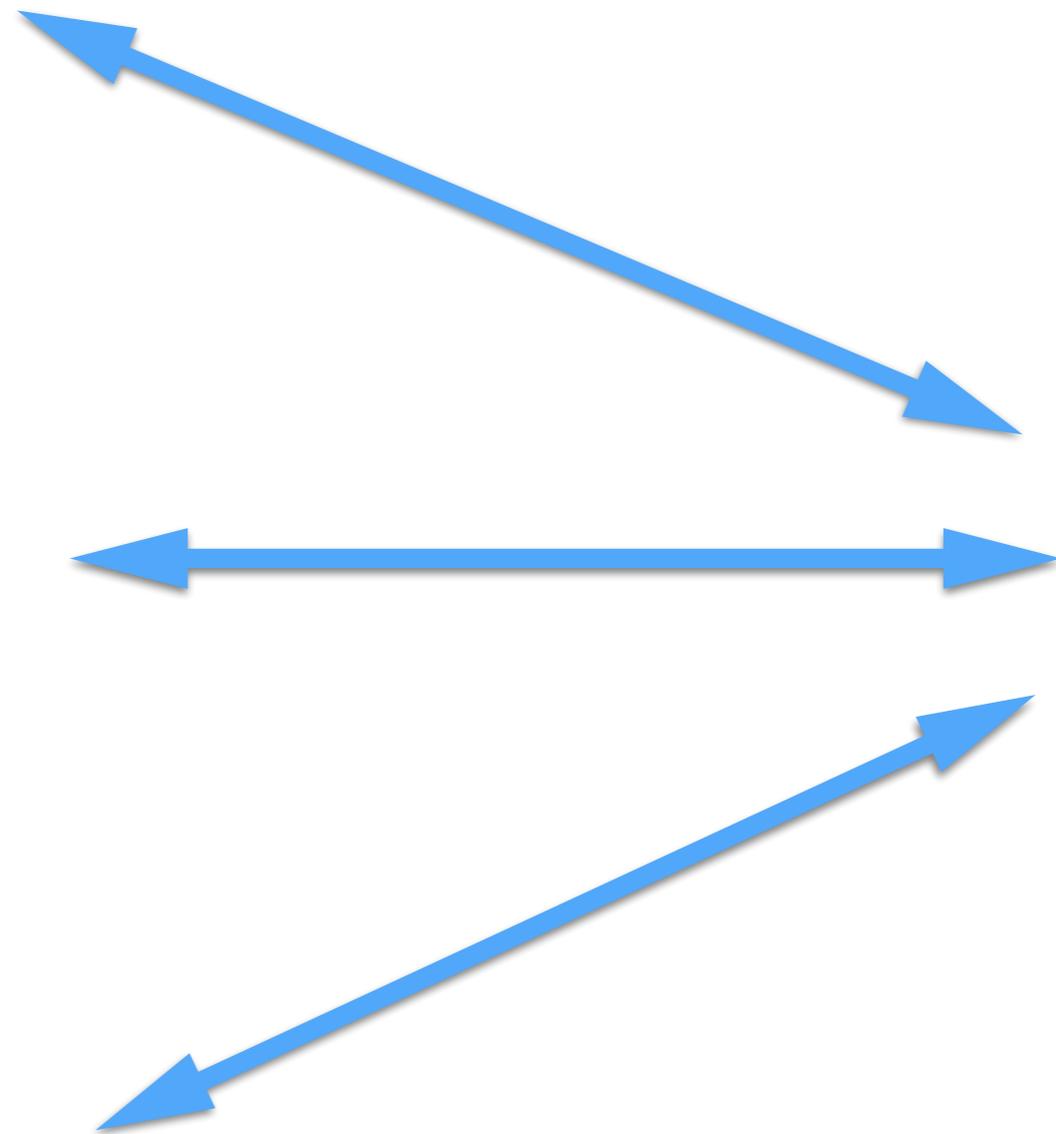


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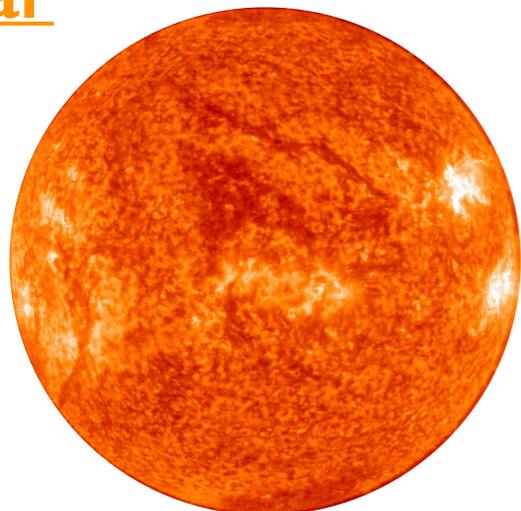
Politicians



One (1) state

Vast number of initial states

A star



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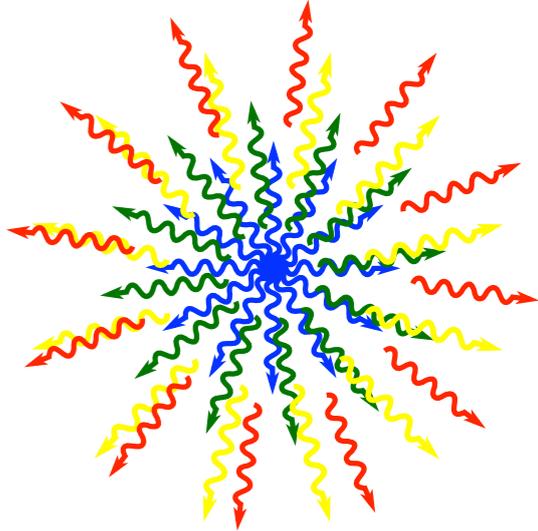
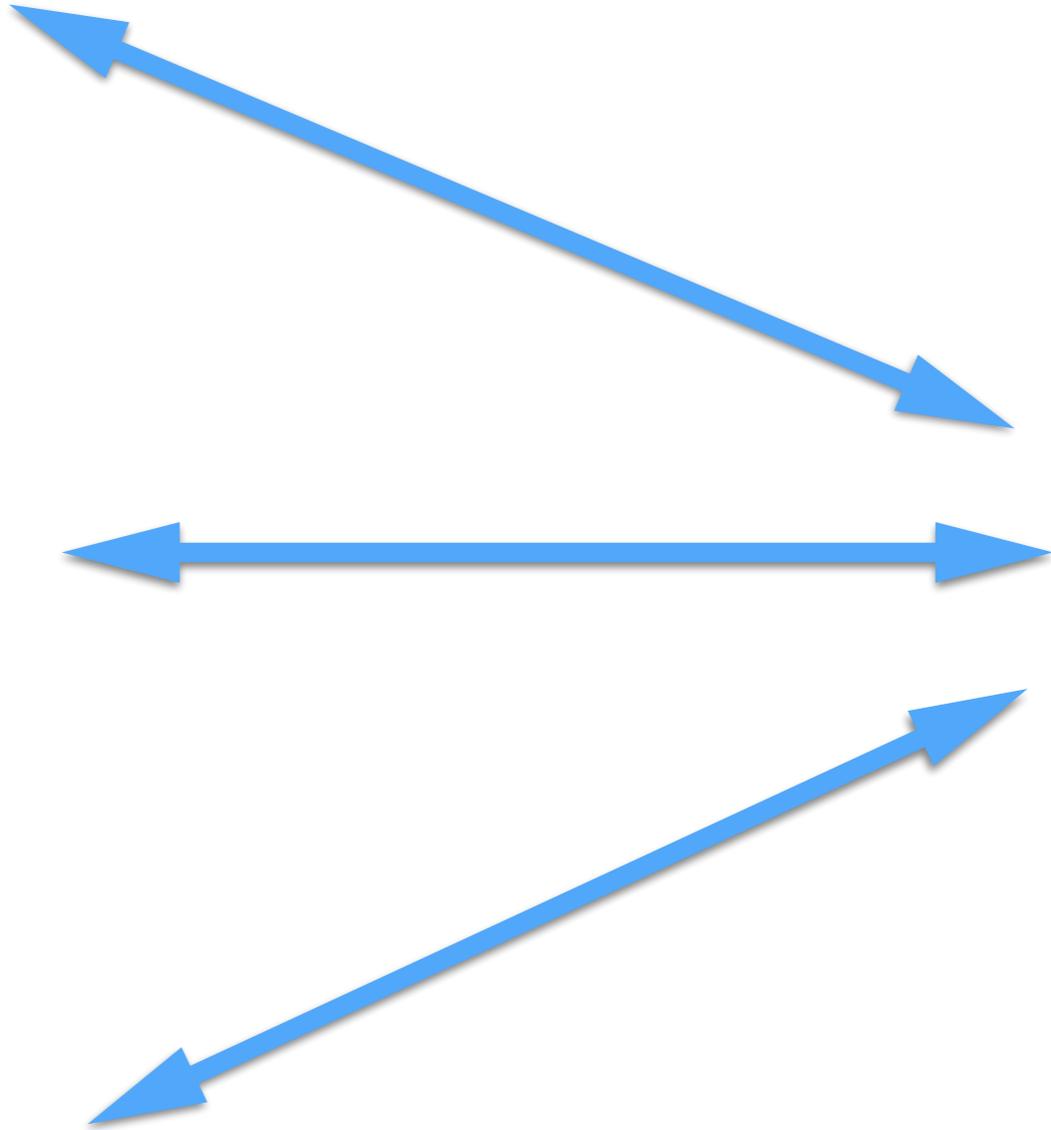
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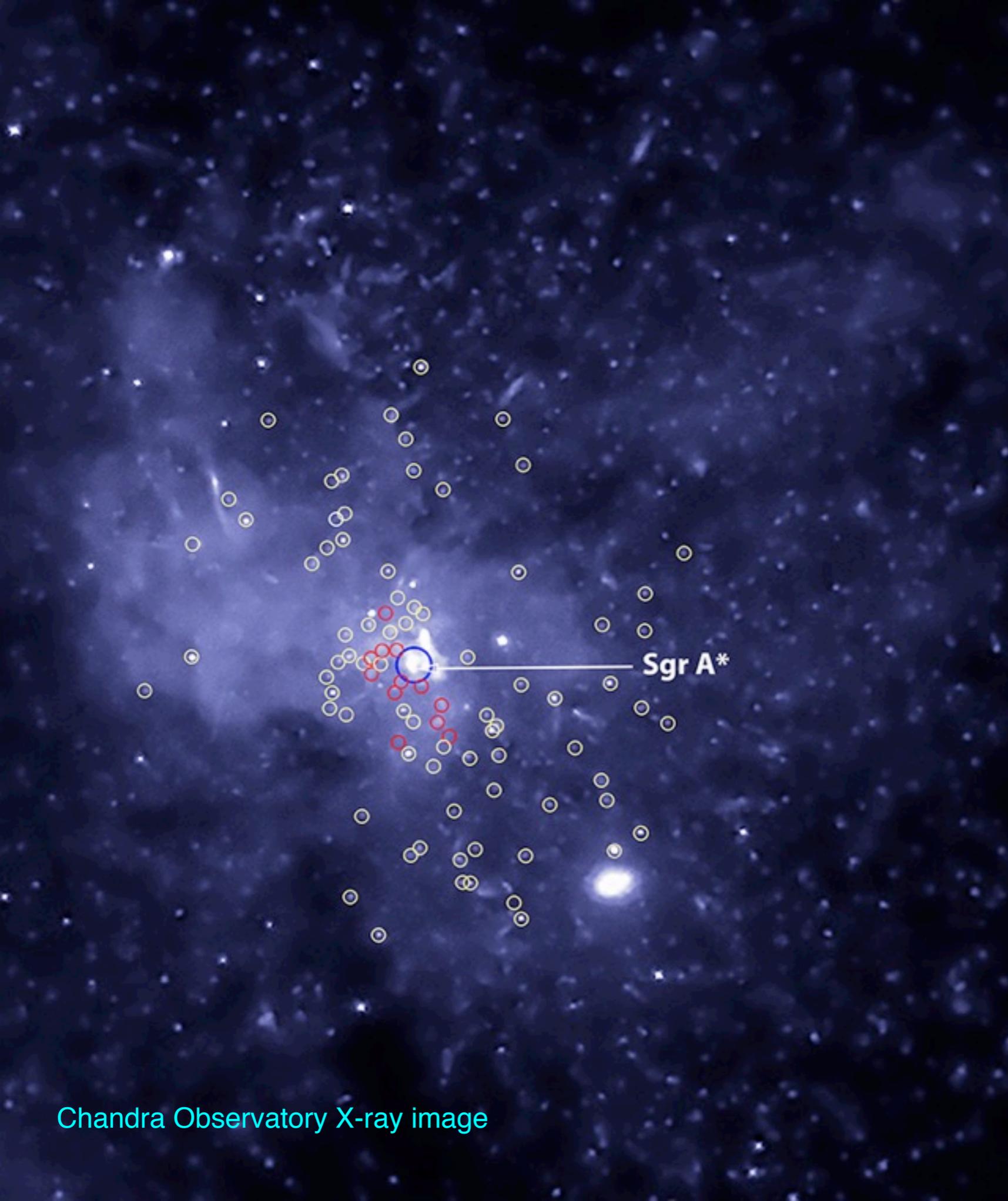


One (1) state

The Information Paradox:
This process cannot be described by Unitary Evolution

Sgr A: Black hole at
the core of Milky Way

Mass $\approx 4 \times 10^6 M_{\odot}$



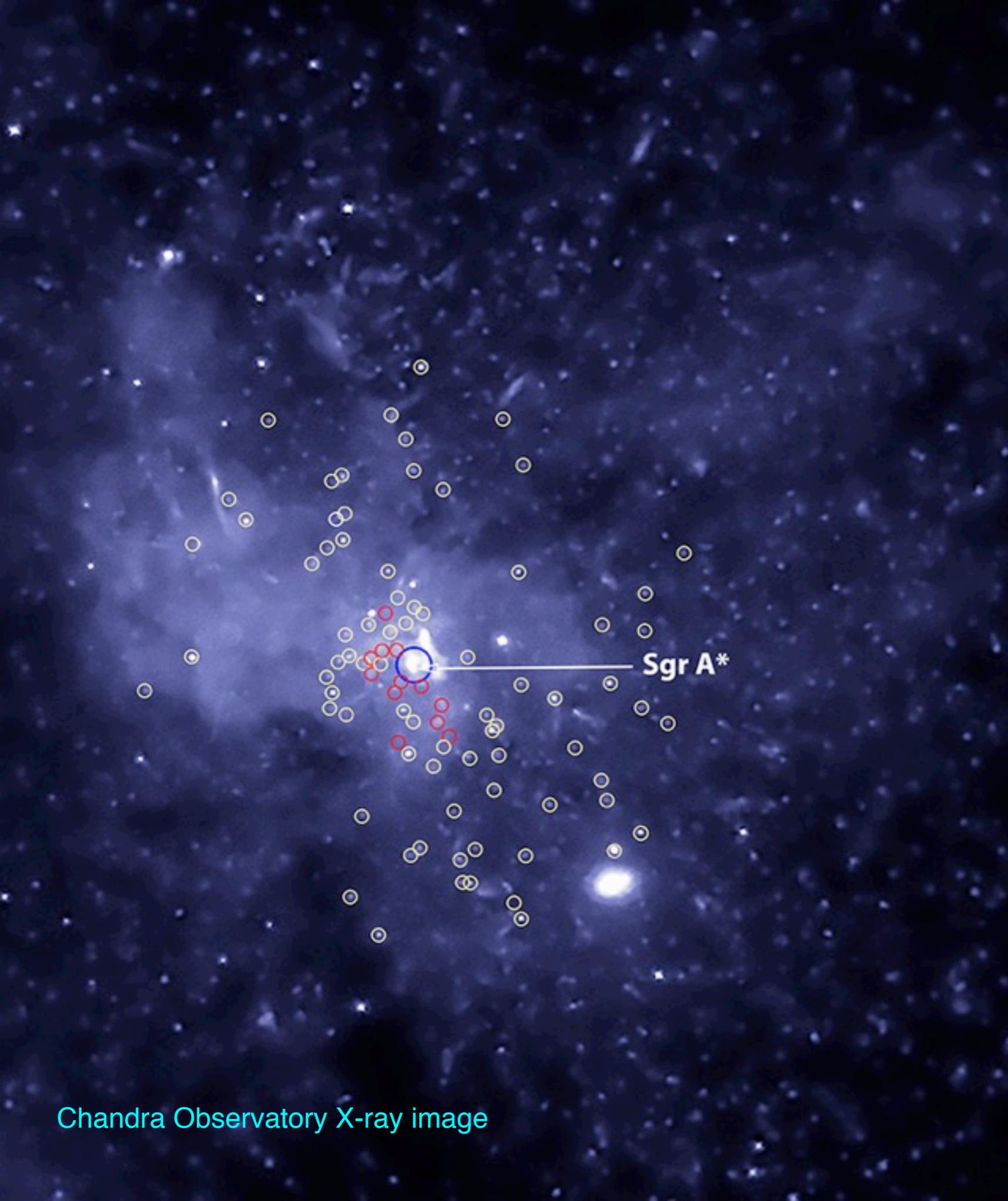
Chandra Observatory X-ray image

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$A \approx 7 \times 10^{90} L_{\text{Planck}}^2$

$S = \frac{1}{4} A \approx 10^{90}$



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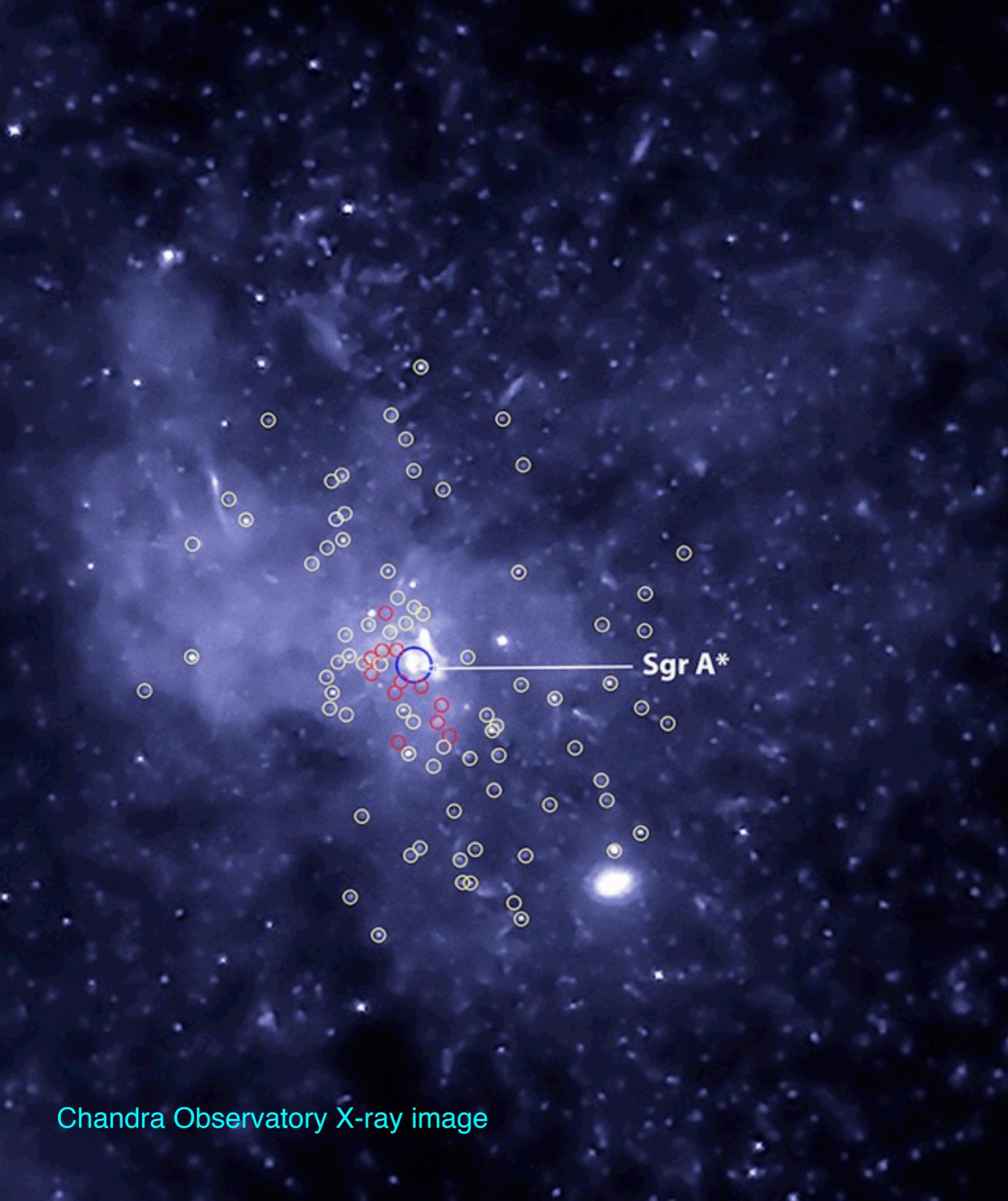
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Number of microstates
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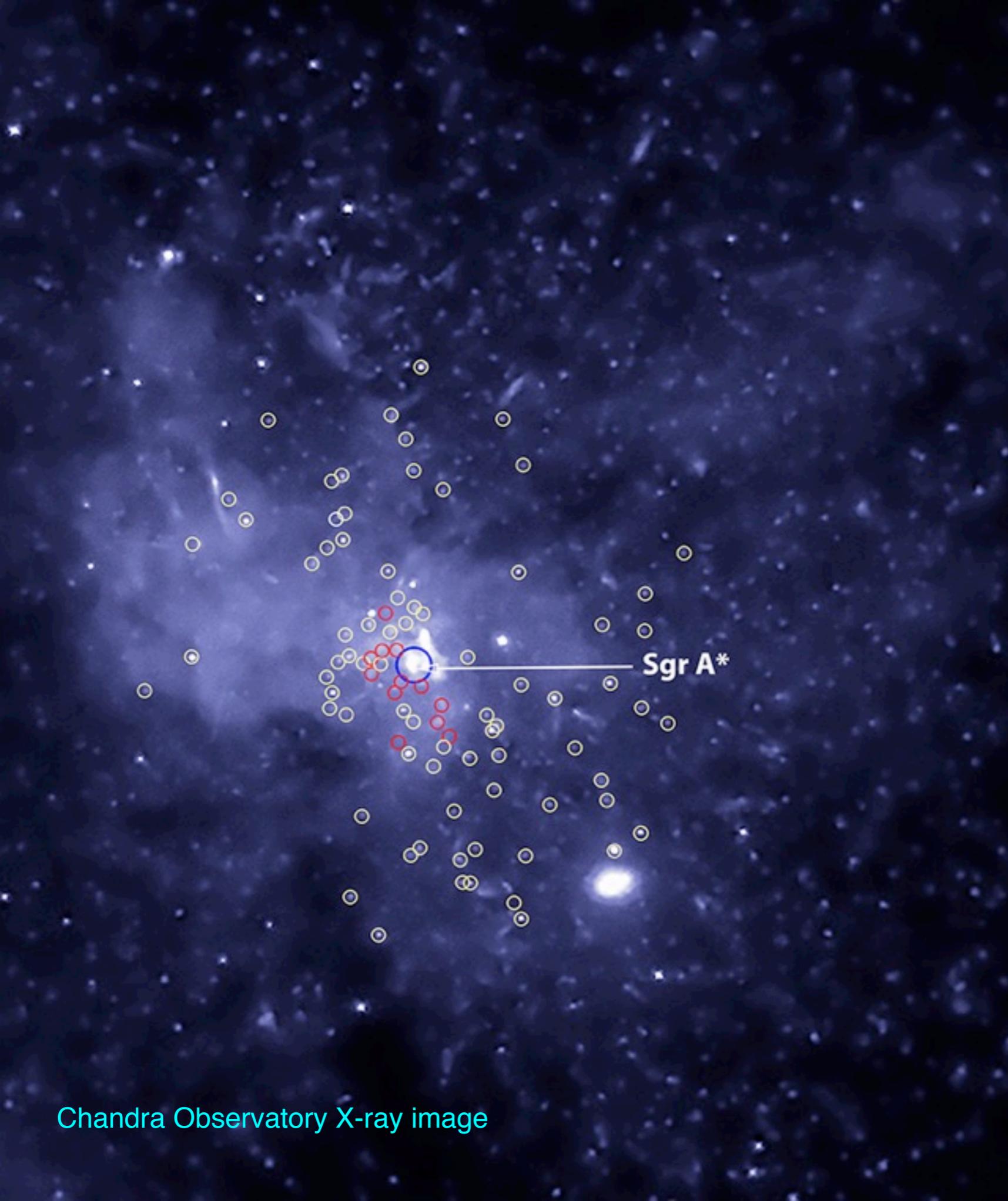
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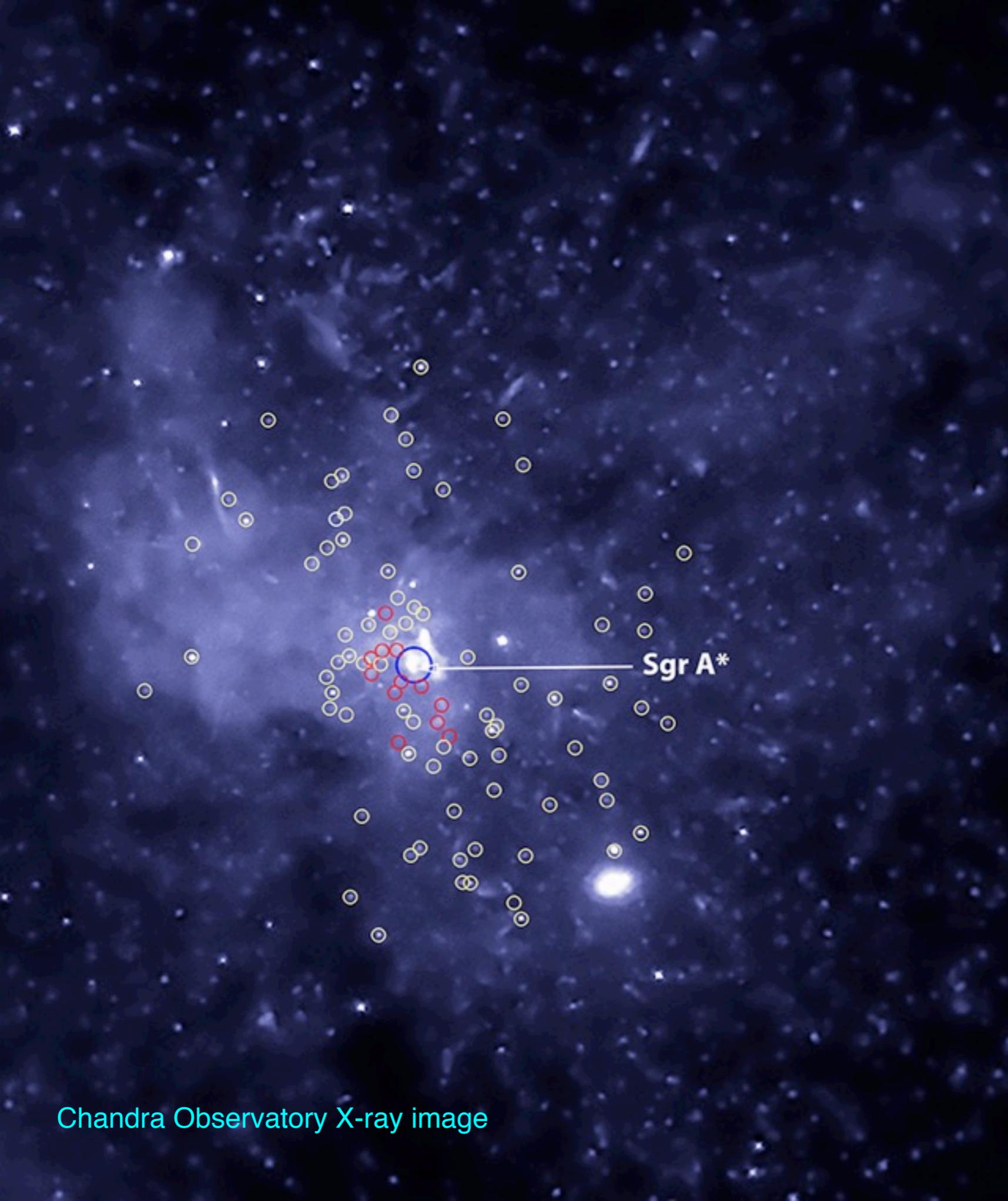
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The information problem:

$$e^{10^{90}} \neq 1$$

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The First Real Progress: BH State Counting in String Theory

Strominger and Vafa: [hep-th/9601029](https://arxiv.org/abs/hep-th/9601029)

Simplify the problem: (i) Impose supersymmetry (ii) $G_{\text{Newton}}, g_{\text{String}} = 0$

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 - ★ Stable and time independent: Hawking Temperature = 0
The *information problem* simplifies to the information storage problem.
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 - ★ Microstates “protected by supersymmetry;” preserved under variation of the string coupling

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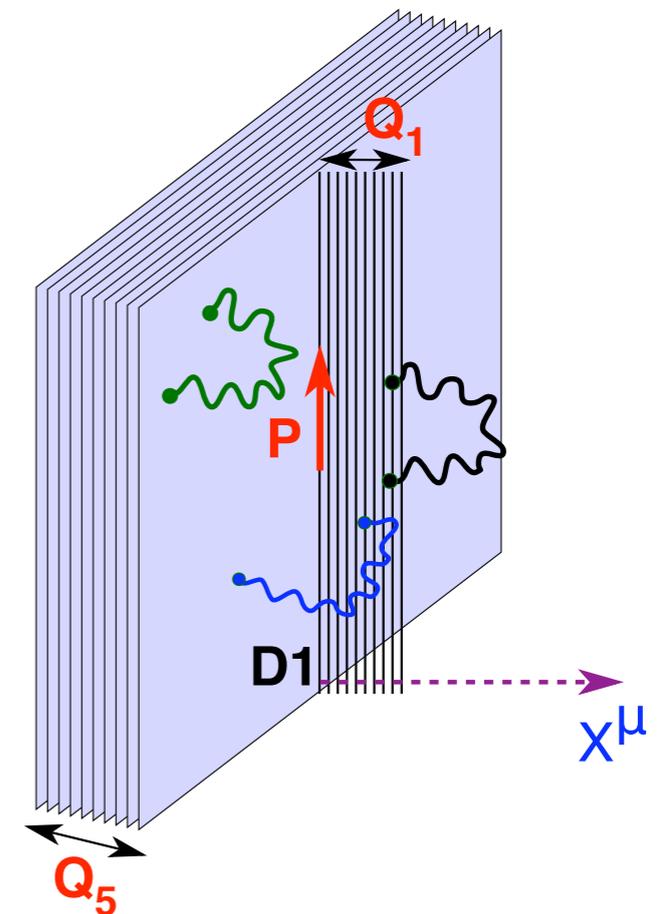
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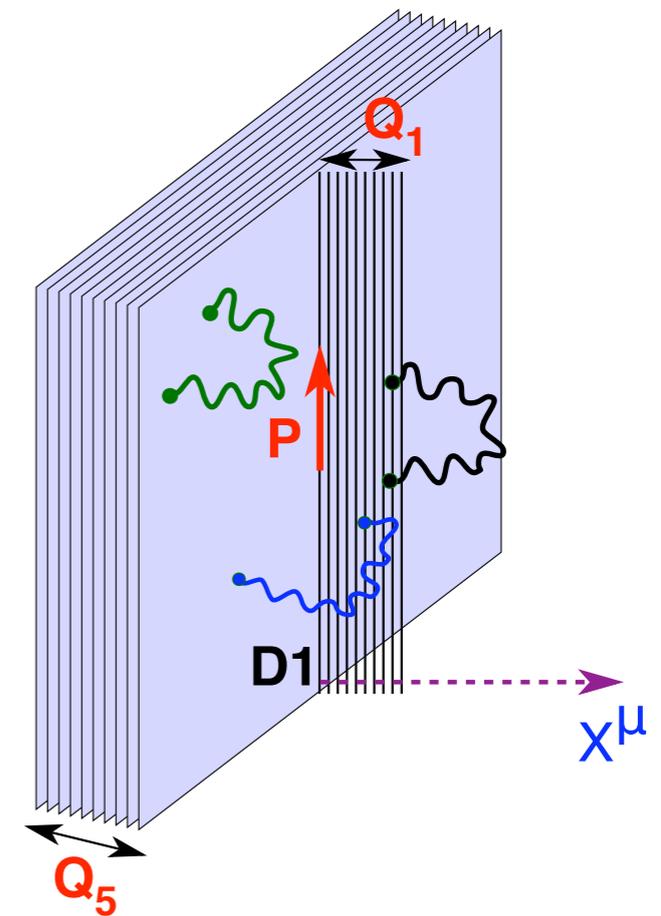
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Vast number of perturbative BPS string states carry momentum, P , along branes.



The First Real Progress: BH State Counting in String Theory

Strominger and Vafa: hep-th/9601029

Simplify the problem: (i) Impose supersymmetry (ii) $G_{\text{Newton}}, g_{\text{String}} = 0$

- Supersymmetric/BPS black holes “ $M = Q$ ”

- ★ Stable and time independent: Hawking Temperature = 0

The information problem simplifies to the information storage problem.

- ★ Computationally far simpler. Microstates are all BPS states

- ★ Microstates “protected by supersymmetry;” preserved under variation of the string coupling

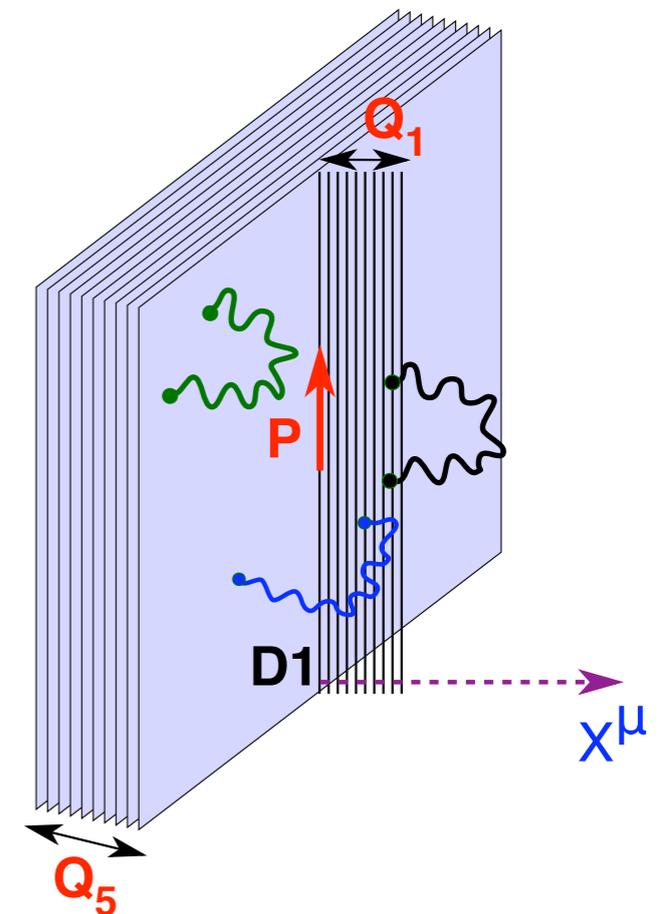
- Vanishing $G_{\text{Newton}}, g_{\text{String}}$

At $g_{\text{String}} = 0$, look for D-brane configurations that become BPS black holes with macroscopic horizon areas at finite $G_{\text{Newton}}, g_{\text{String}}$

Vast number of perturbative BPS string states carry momentum, P , along branes.

Count these microstates:

$$S \equiv \log[\Omega(Q_1, Q_5, Q_P)] = 2\pi \sqrt{Q_1 Q_5 Q_P}$$



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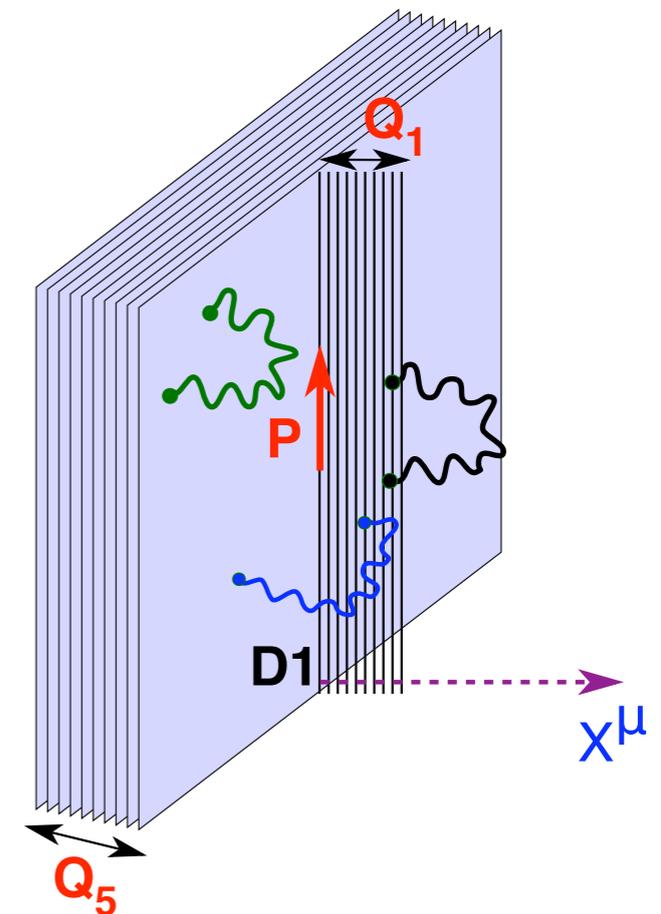
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$$\begin{aligned} S &\equiv \log[\Omega(Q_1, Q_5, Q_P)] = 2\pi \sqrt{Q_1 Q_5 Q_P} \\ &= \frac{A_{\text{Horizon}}}{4 \ell_{\text{Planck}}^2} \quad \text{at finite } G_{\text{Newton}}, g_{\text{String}} \end{aligned}$$



Goals since late 1990's:

(i) G_{Newton} , $g_{\text{String}} \neq 0$

(ii) No supersymmetry; broken supersymmetry

State counting using a supersymmetric index

The perturbative microstates are supersymmetric index states and so should be preserved under deformations, including changes in G_{Newton} and g_{String} .

*So what do these perturbative string states become at finite $G_{\text{Newton}}, g_{\text{String}}$?
What do black-hole microstates actually look like?*

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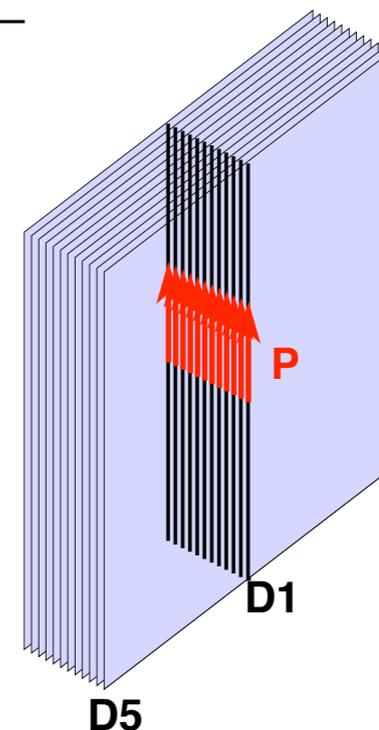
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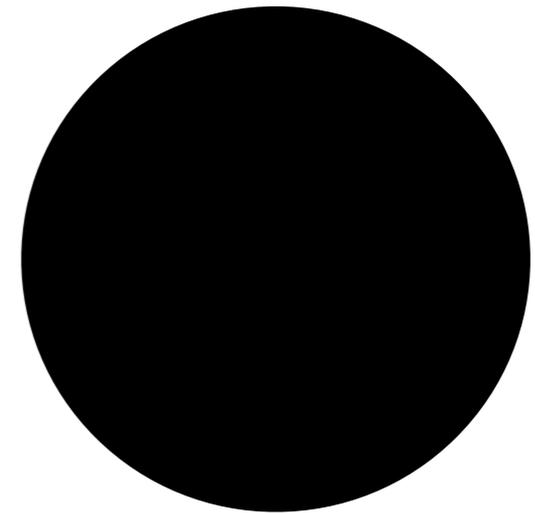
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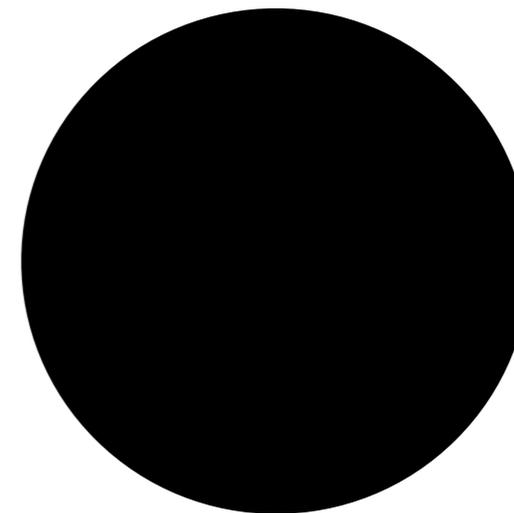
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We seem to be back where we started ...

Dodging the bullets

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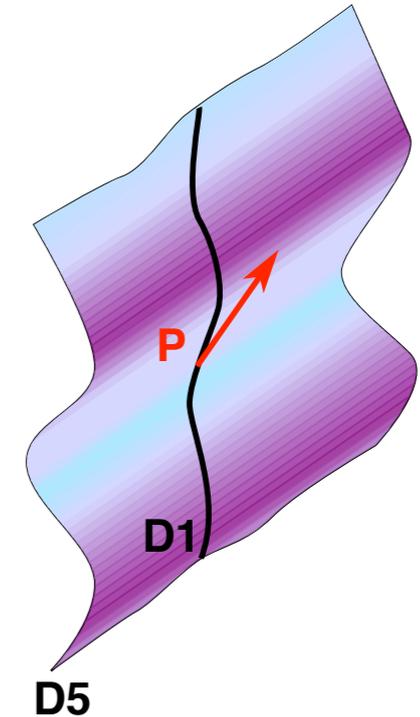
Microstate scaling

D-brane tension $\sim g_s^{-1}$ + momentum, P

\Rightarrow Branes spread out with increasing g_s

Can arrange brane configurations that grow in size with g_s at exactly the same rate as the size of a black hole.

Bena and Kraus, hep-th/0402144



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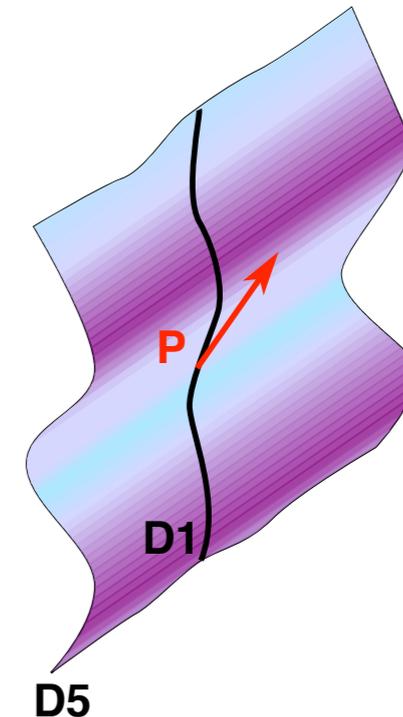
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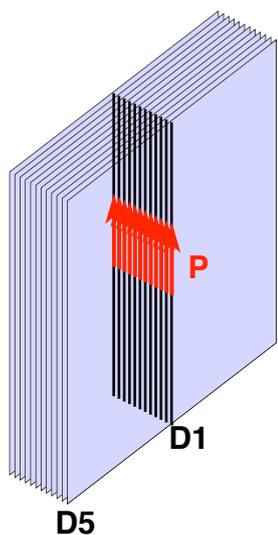
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Geometric transitions in more than (3+1) dimensions

Dp-branes are electric charges of p-form Maxwell fields whose equations of motion typically involve Chern-Simons terms:

$$d * F^{(p)} = \delta^{(D-p+1)} + \sum_k G^{(k)} \wedge G^{(D-p-k+1)}$$



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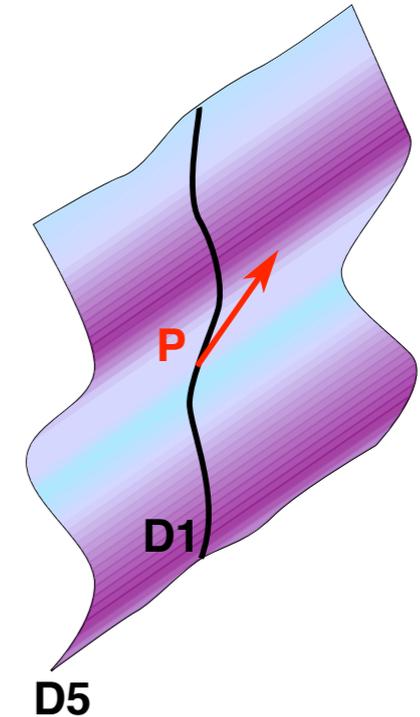
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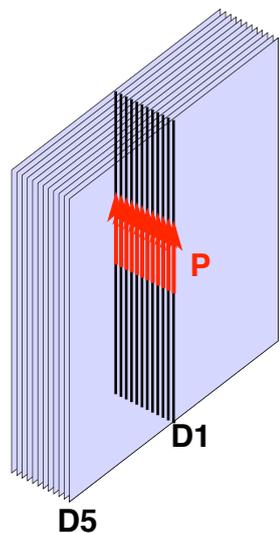
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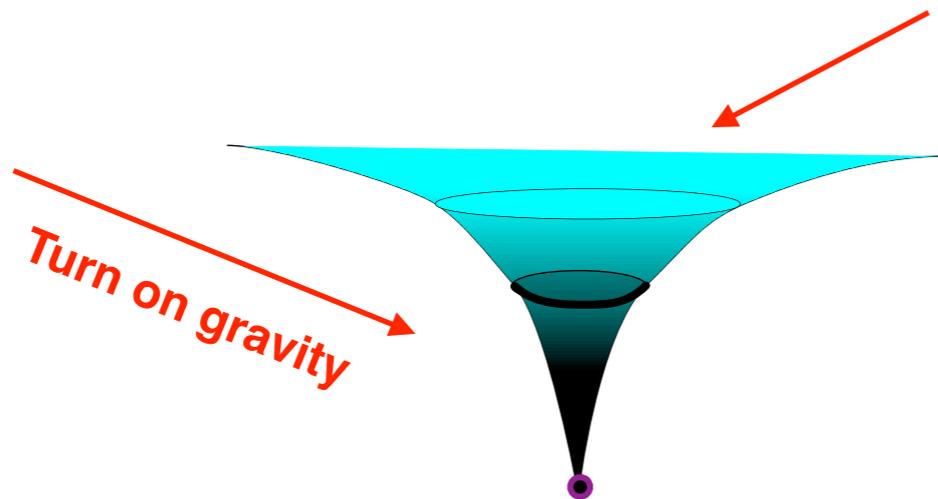


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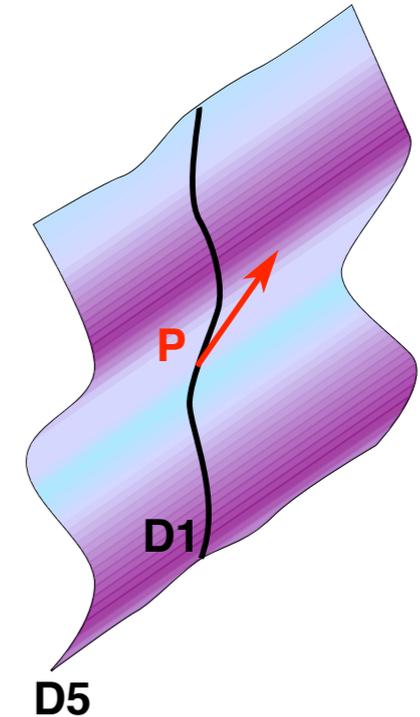
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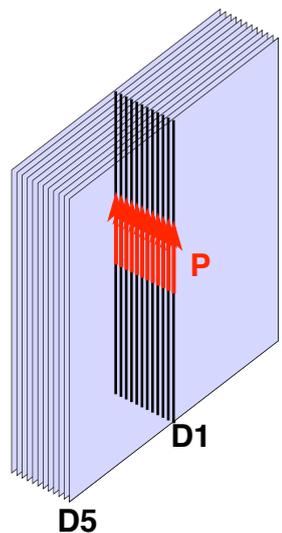
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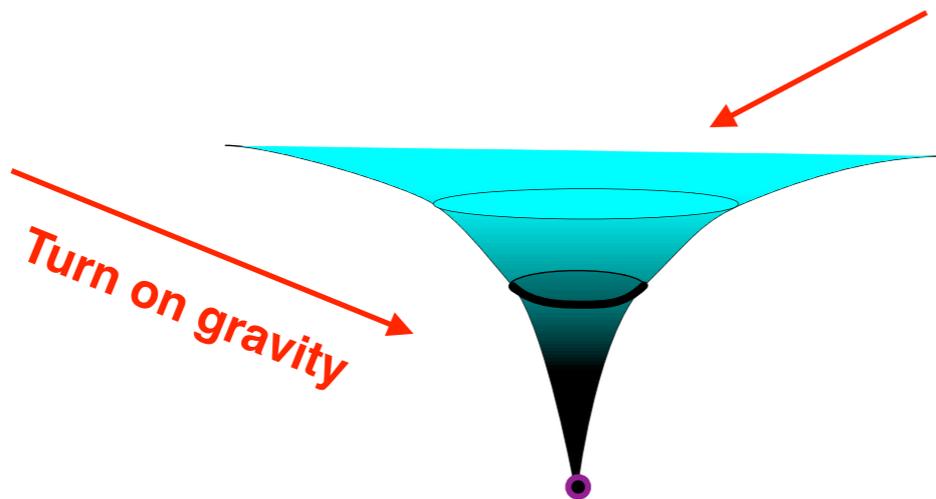


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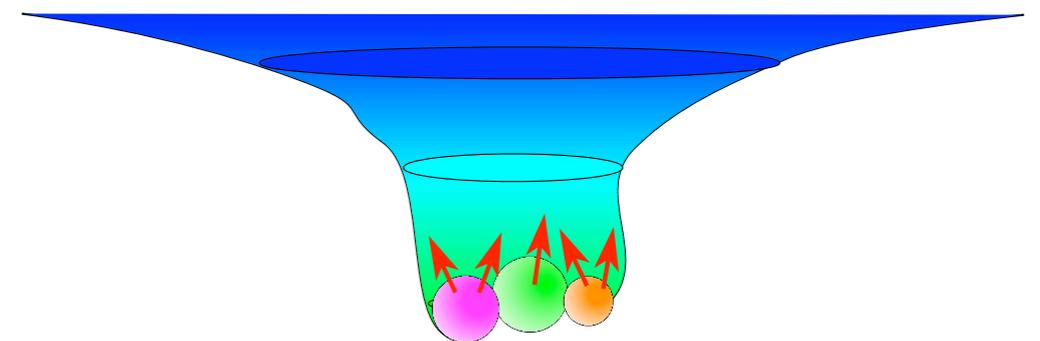
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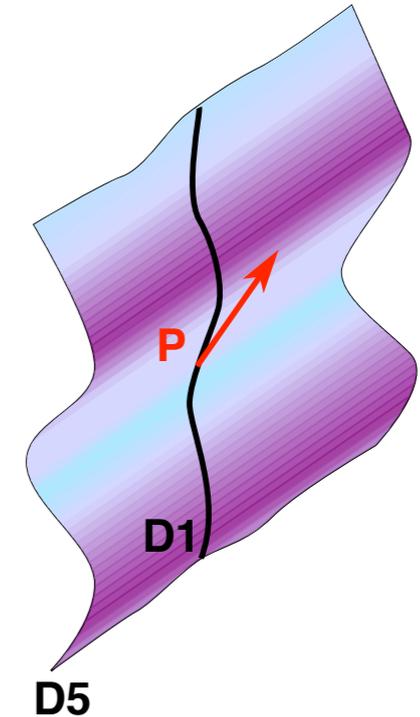
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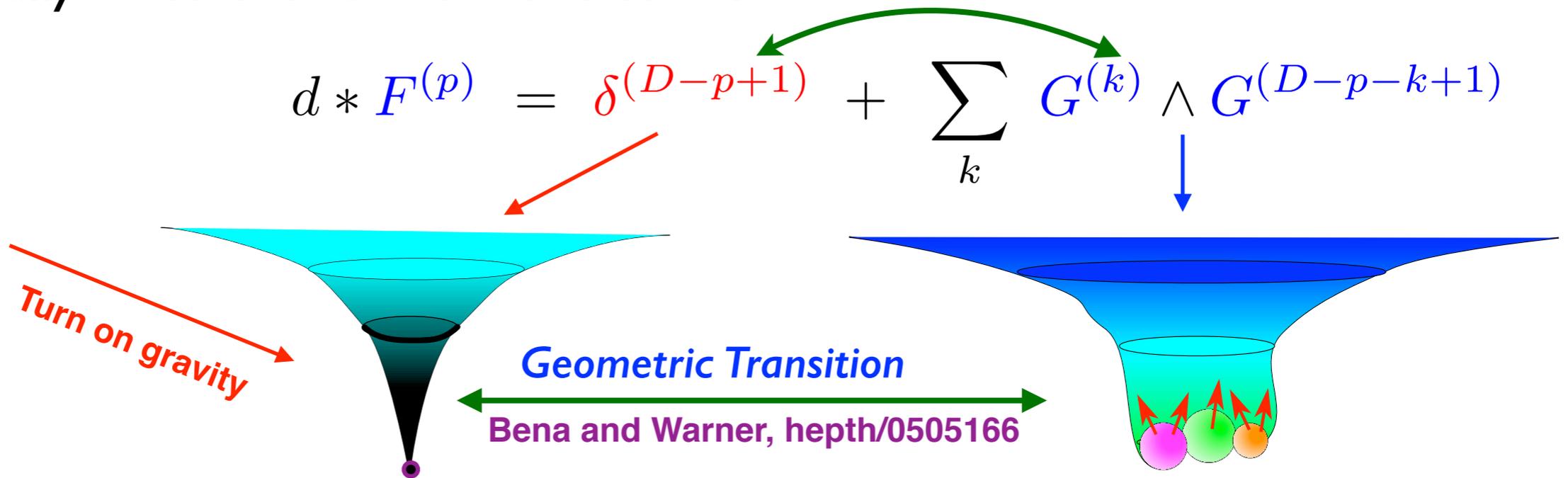
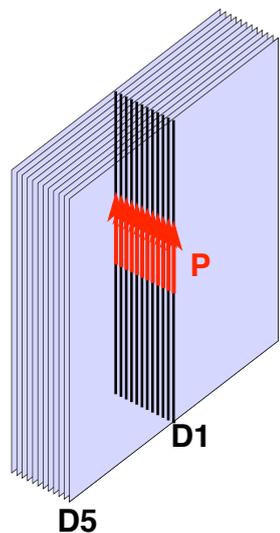
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String theory can resolve singularities:



... in more than (3+1) dimensions

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But it must do more

The information problem is caused by the presence of a horizon ...

\Rightarrow Event horizons must go!

Microstate Geometry Program

- Goals:
- ★ Resolve singularities
 - ★ Remove horizons
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- Stringy resolutions on horizon scale \Rightarrow *Very long-range effects* \Rightarrow Massless fields (*only other scale in String theory is the Planck scale*)



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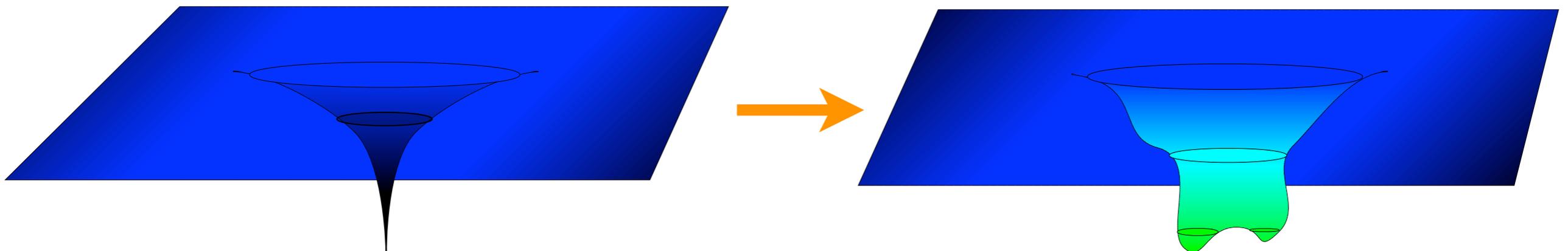


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Microstate Geometries

- ▶ *Smooth, horizonless solutions* to the *bosonic* sector of supergravity (the low-energy limit of string theory) *with the same asymptotic structure as a given black hole or black ring*

Singularity resolved; Horizon removed



Microstate Geometries: Some Major Developments

- ★ 2005: First large families of explicit examples **Bena and Warner, hep-th/0505166**
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An old conceit: Hawking evaporation is extremely slow:

$$t_{evap} = \frac{5120 \pi G^2 M_{\odot}^3}{\hbar c^4} \approx 6.6 \times 10^{74} s \approx 2.1 \times 10^{67} years \quad \text{for a one solar mass black hole}$$

leak the information slowly via tiny quantum gravity/string (*Riemann*)ⁿ corrections to radiation.

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⇒ There must be *O(1)* to the Hawking states at the horizon.

(from strong subadditivity of quantum information)

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- ★ 2015-8: Precision holography of Microstate Geometries

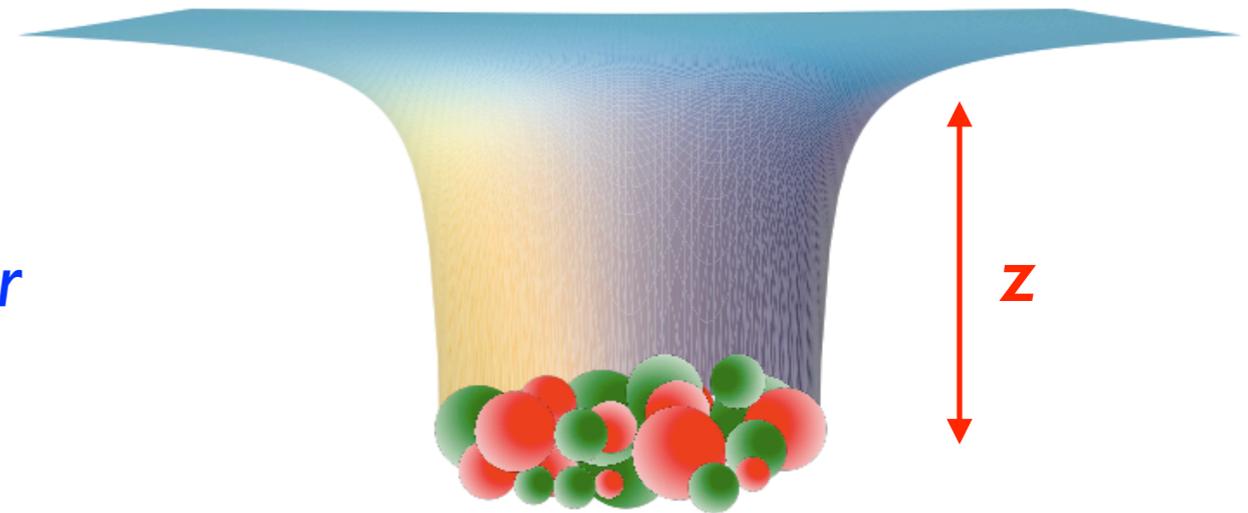
Holographic duals of families of some Strominger-Vafa microstates

Bena, Giusto, Martinec, Russo, Shigemori, Turton and Warner
arXiv:1503.01463, arXiv:1607.03908, arXiv:1711.10474

Scaling Microstate Geometries and Holography

Scaling microstate geometries:

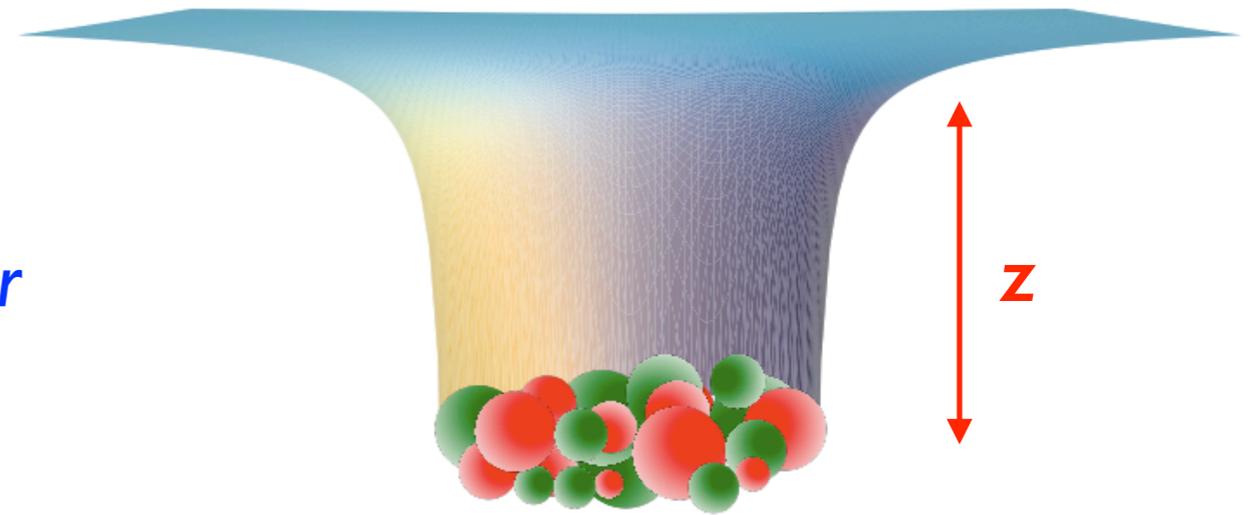
- ✦ Look like black holes until arbitrarily close to the horizon *but cap off smoothly*
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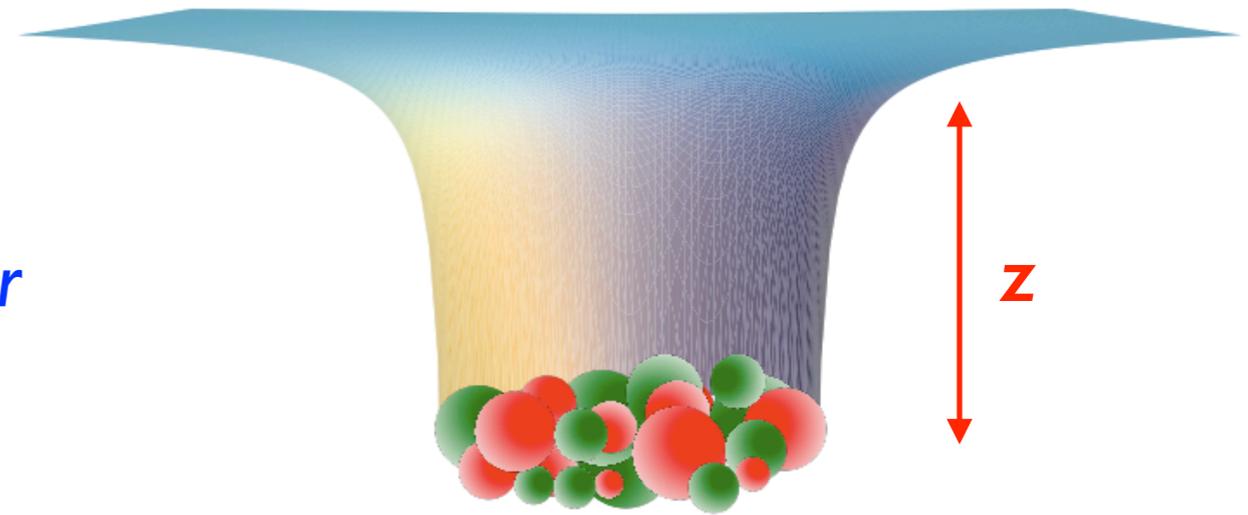
New physics *at the horizon scale:*

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The Dual CFT

In the system built with *momentum waves* on a system of D1 and D5 branes

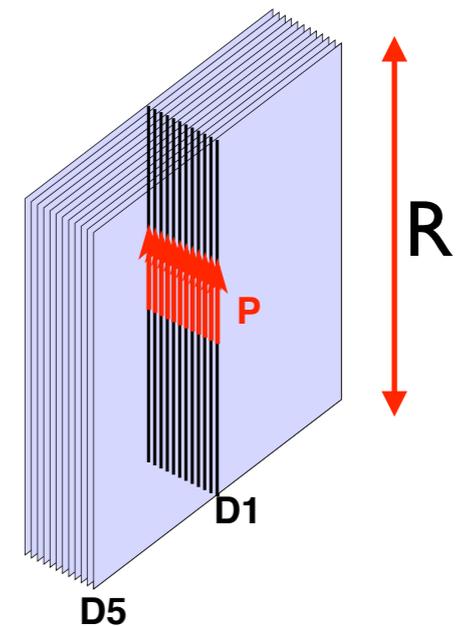
Black-hole CFT: a marginal deformation of an orbifold (1+1)-dimensional CFT

Central charge

$$c = 6 N_1 N_5$$

Energy gap

$$E_{gap} = \frac{1}{N_1 N_5 R}$$

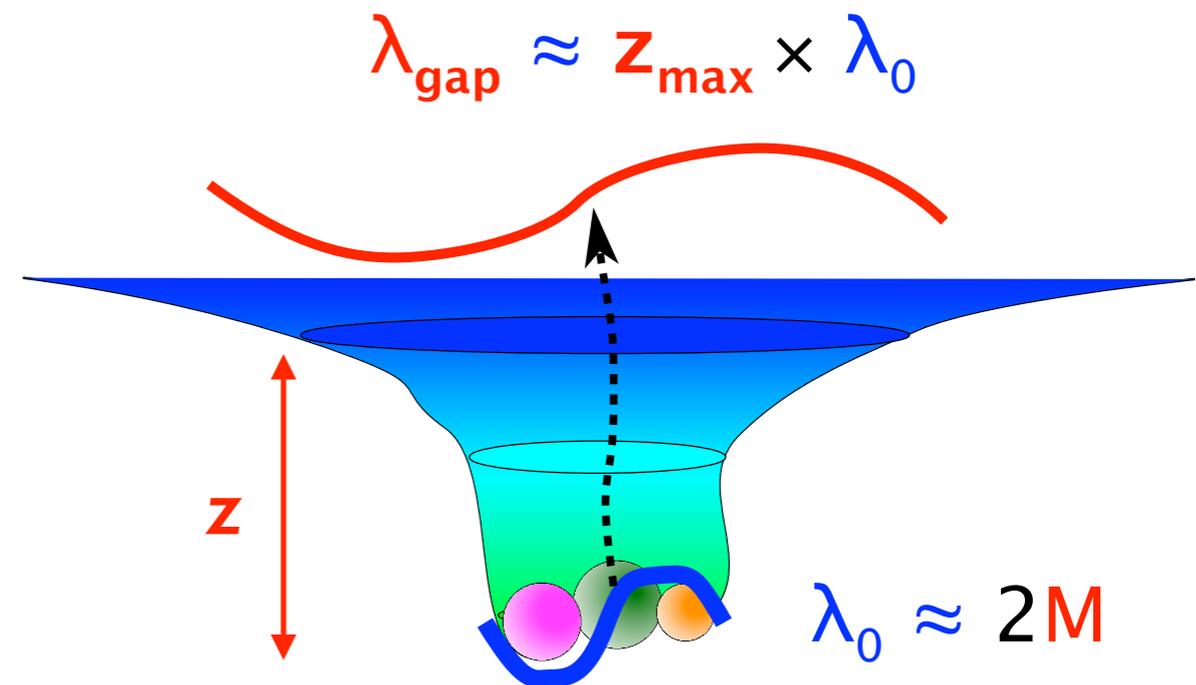


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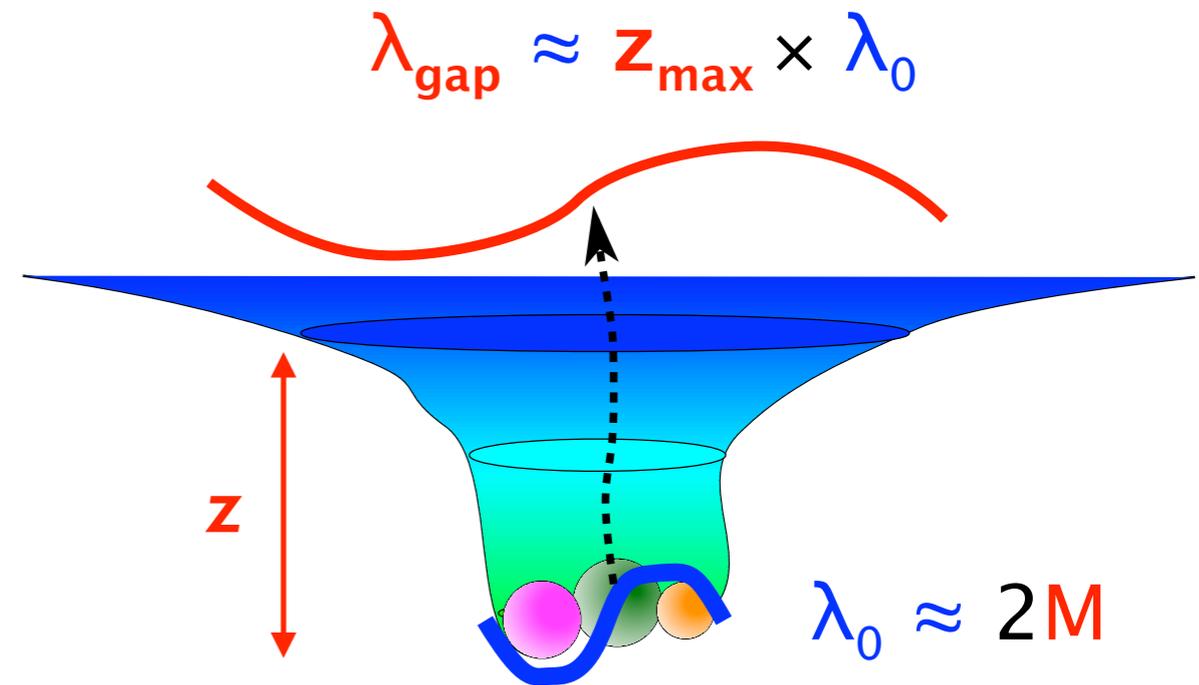
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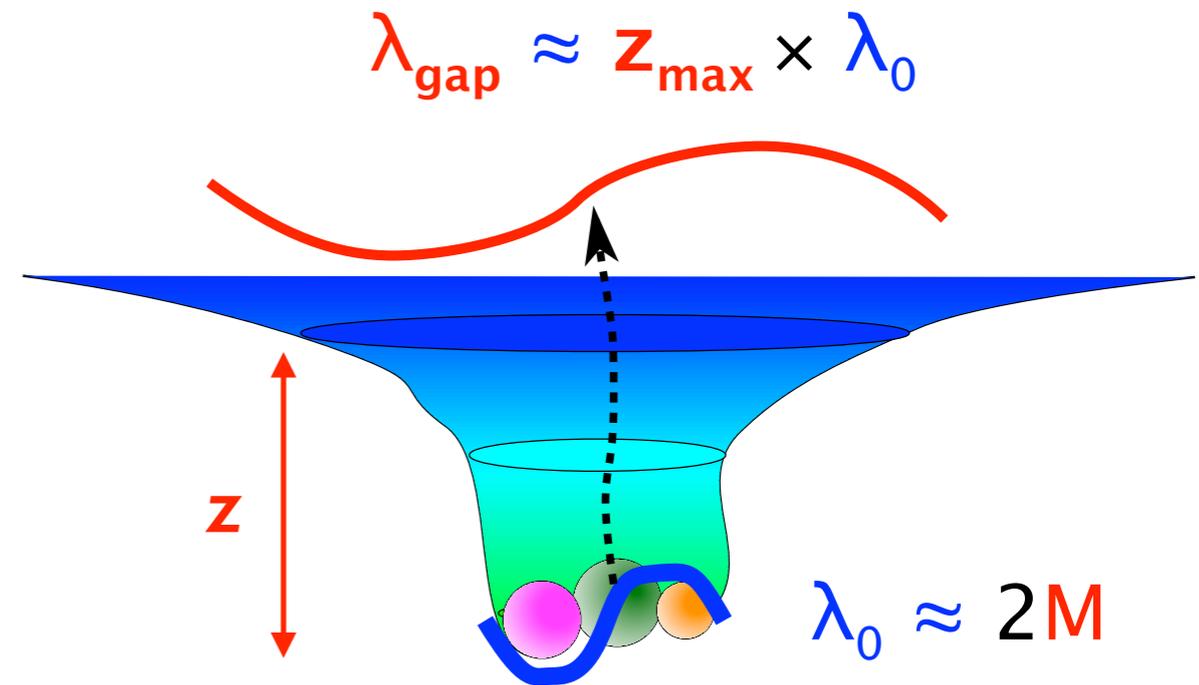
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Bena, Wang and Warner, arXiv:hep-th/0608217

de Boer, El-Showk, Messamah, Van den Bleeken, arXiv:0807.4556

The throat depth, or z , is *not* a free parameter: It is quantized and *there is a maximal depth* \Rightarrow There is a *minimum energy* for excitations of the geometry

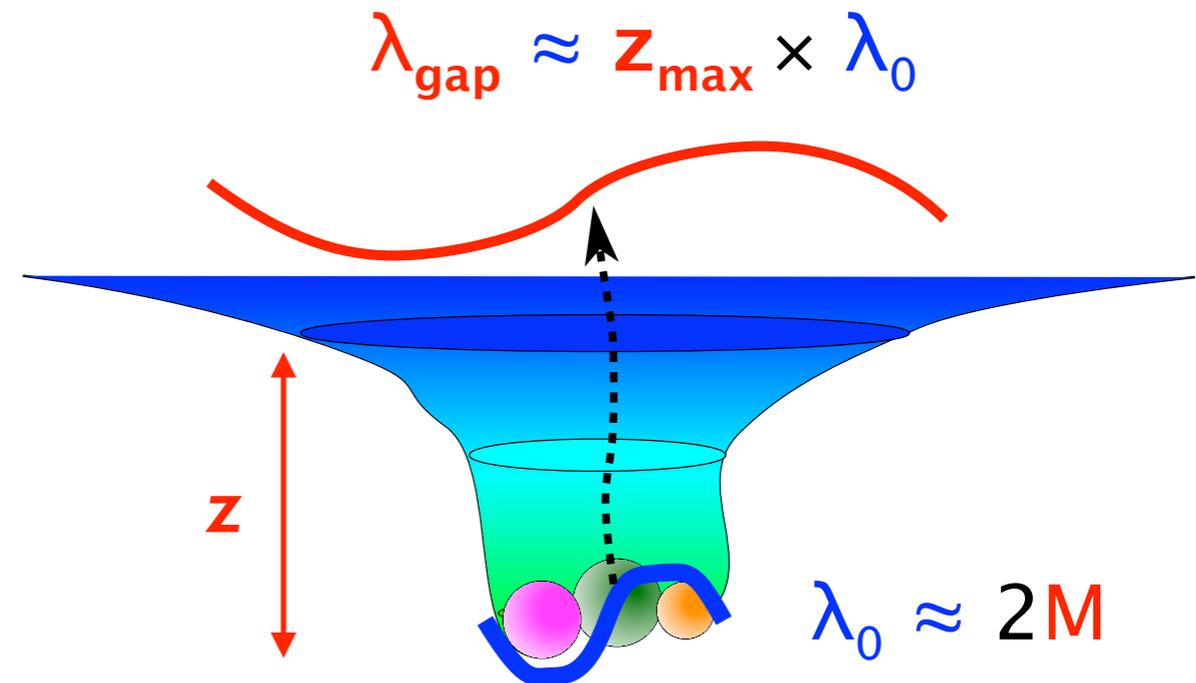
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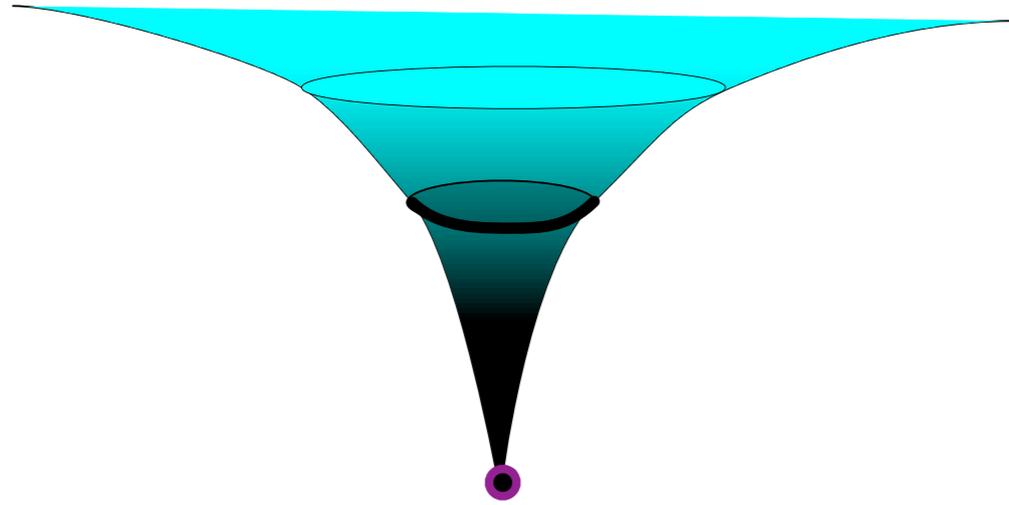
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We found that this *exactly matches* $E_{\text{gap}} \sim (C_{\text{cft}})^{-1}$ for the *softest stringy excitations* underlying the original state counting of Strominger and Vafa

\Rightarrow Excitations of scaling microstate geometries access the softest modes of the “*typical sector*” that provides the dominant contribution to the entropy ...

The Holography of Topological Bubbles



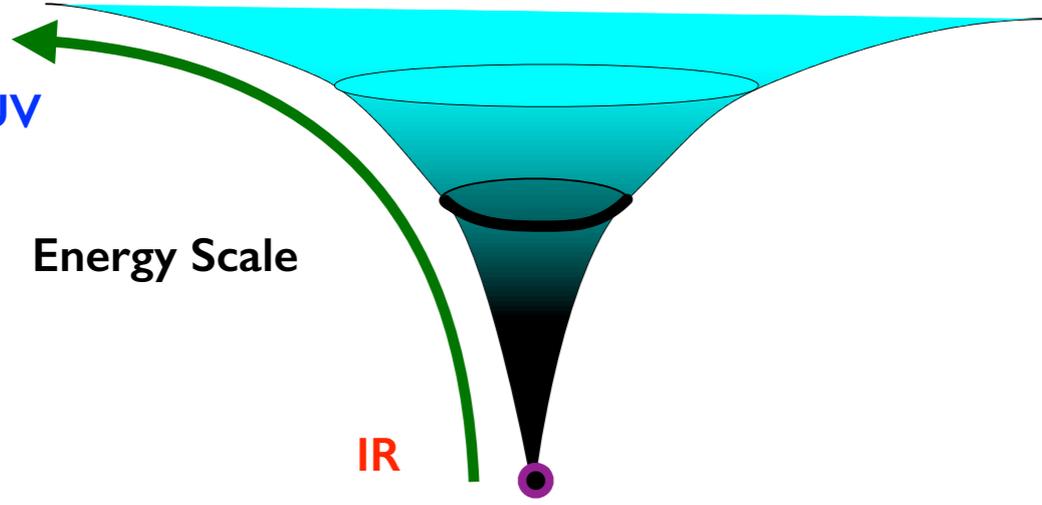
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UV CFT

UV

Energy Scale

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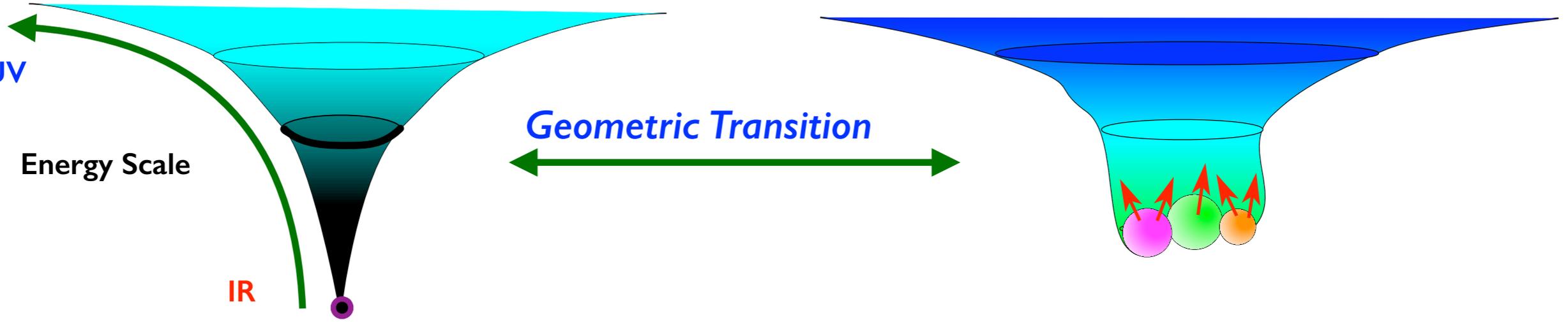
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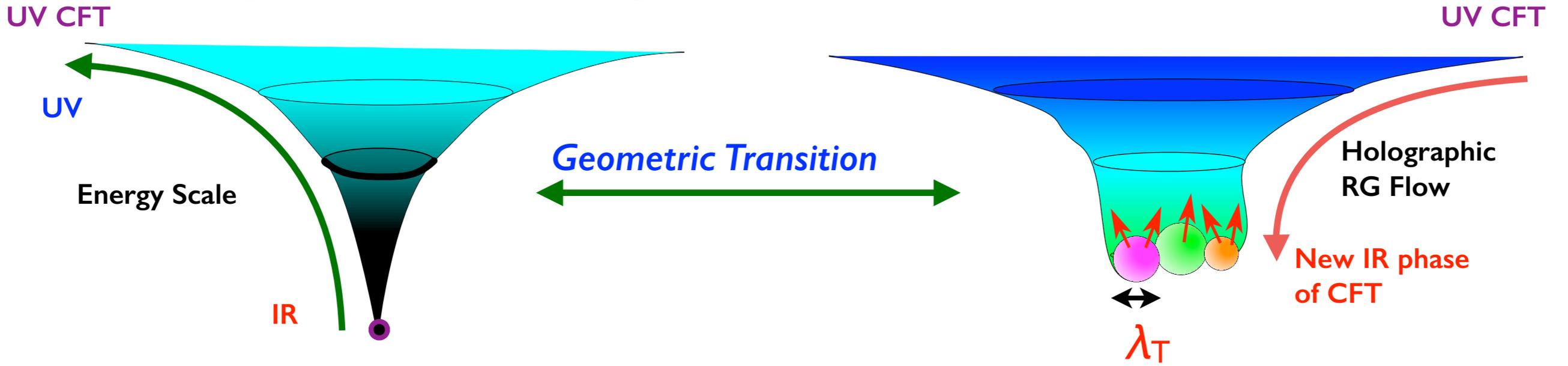
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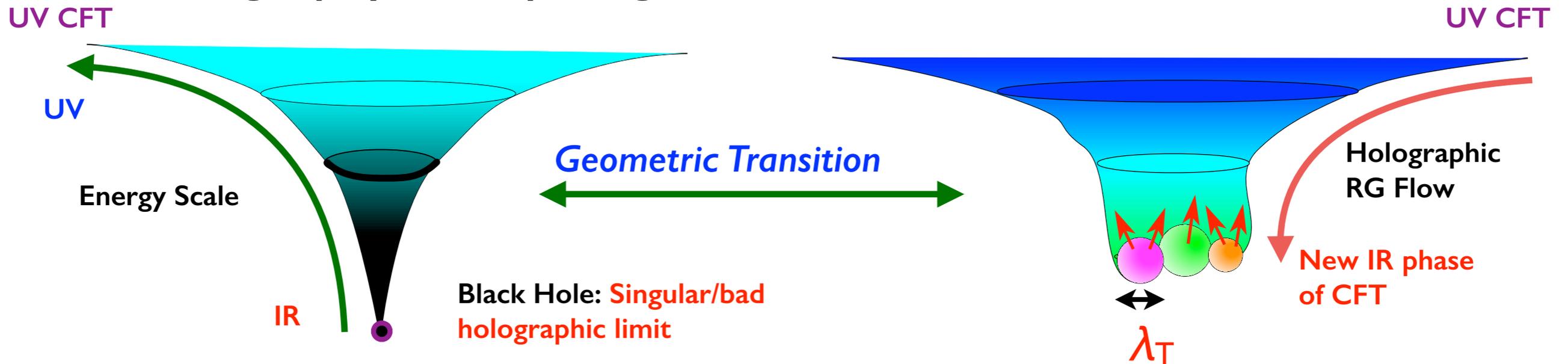
Geometric Transition



The Holography of Topological Bubbles



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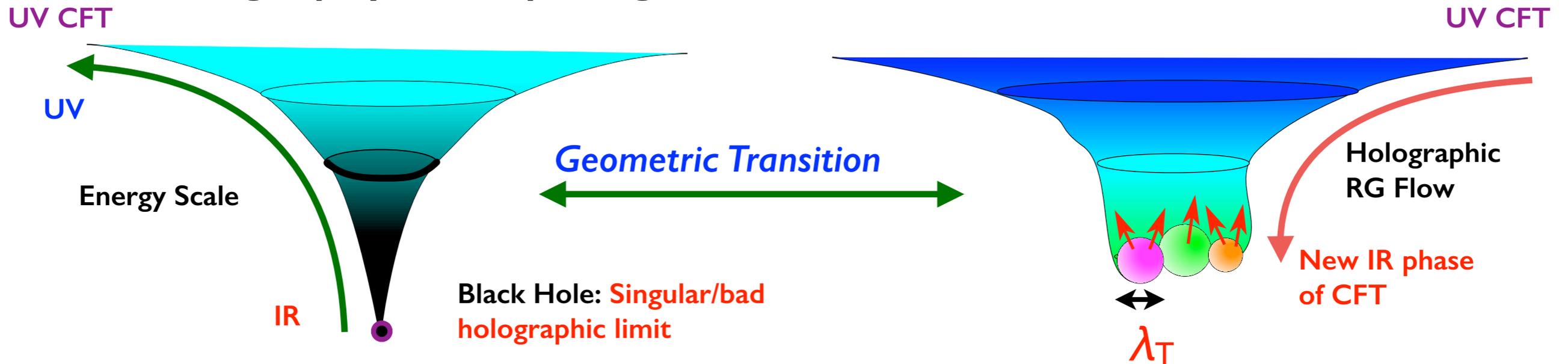


Holographic field theory:

- *Magnetic fluxes* \longleftrightarrow *vevs of order parameters of new phases*
- *Scale, λ_T , of topological bubbles* \longleftrightarrow *Magnitude of vevs*
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Essential ideas:

- *Microstate geometries* \longleftrightarrow *IR phases of the CFT*
- *Fluxes* \longleftrightarrow *Order parameters*
- *New scales in the physics of black holes:*
 - *Order parameters: λ_T*
 - *Energy gap: E_{gap}*
- *Fluctuations of microstate geometries* \longleftrightarrow *Excited states of the CFT*

Current Foci of the Microstate Geometry Program

- ◆ How much of the microstate structure can be encoded semi-classically in the fluctuations spectrum of the fields and the geometry?
 - *Classify these states holographically*
 - *Classify the phase structure of the microstate geometries*

Can these states sample phase space with enough fidelity to give a semi-classical description of black-hole thermodynamics?

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- ◆ Use microstate geometries to study intrinsically stringy effects within given coherent phases of the black-hole system

- *Stringy excitations within these geometries* Martinec, Massai, Turton
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Find intrinsically stringy descriptions of microstates in horizon-scale environment

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- ◆ Study the holographic field theory of microstate geometries

- *States and partition functions*
- *Correlation functions in the CFT and the geometry*

Giusto, Russo, Shigemori, Turton

Raju, arXiv:1804.10616

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The ERC Project: Quantum Black Hole Structure

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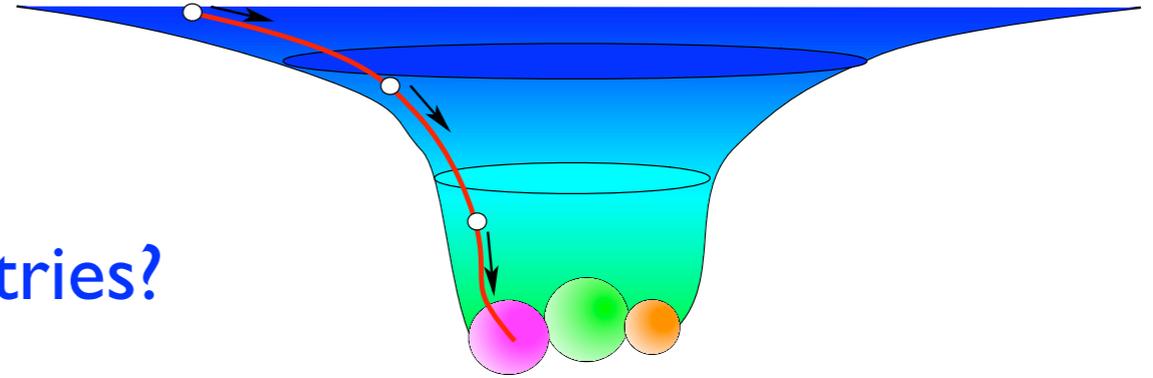
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- ◆ Can we use microstate geometries to model ideas of emergent space-time?

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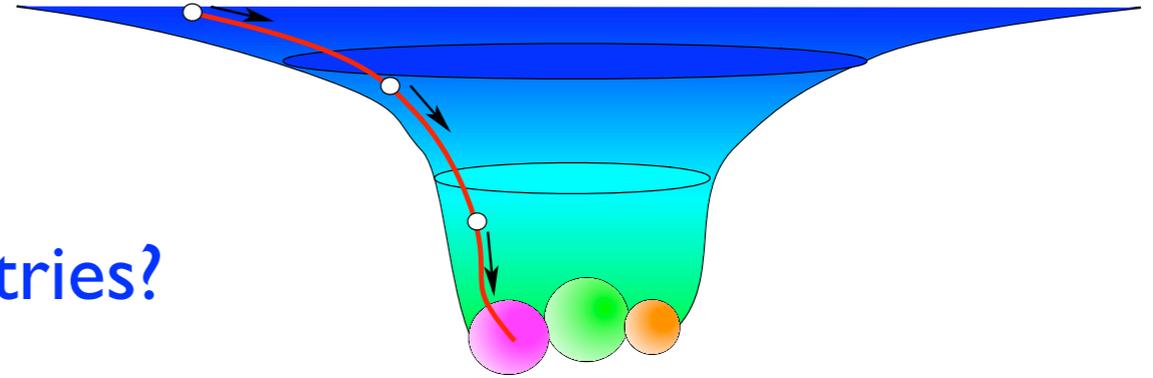
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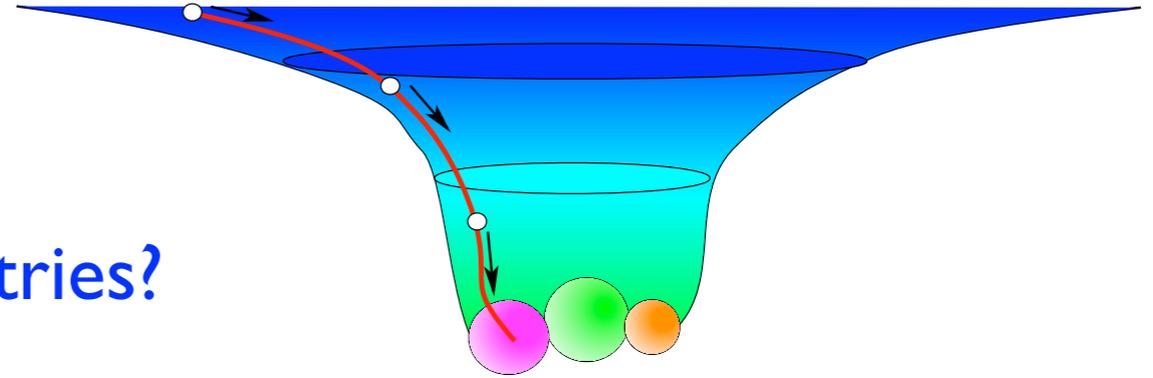


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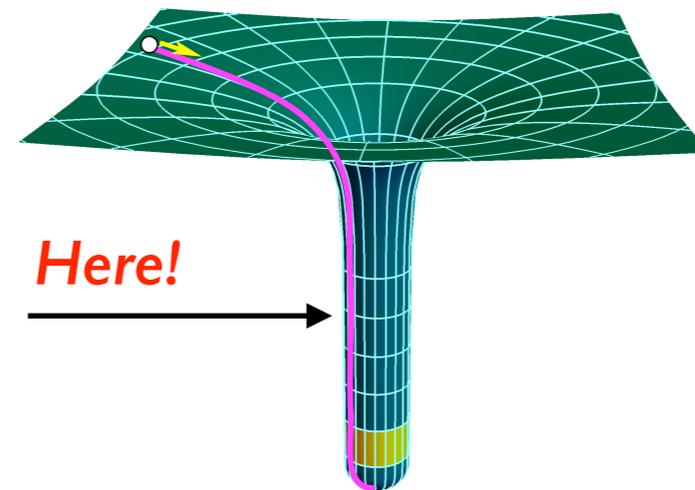
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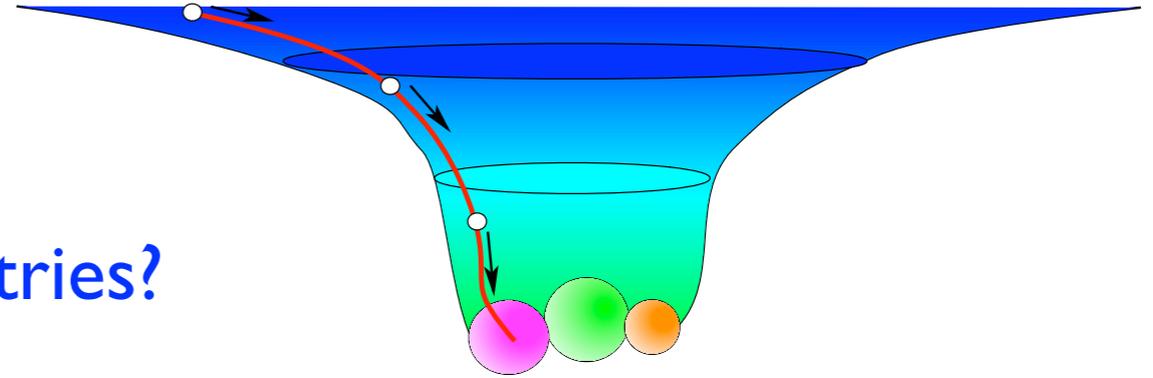


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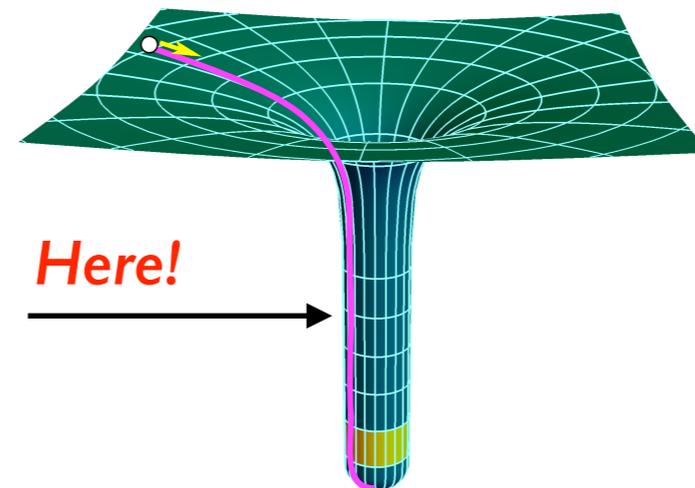
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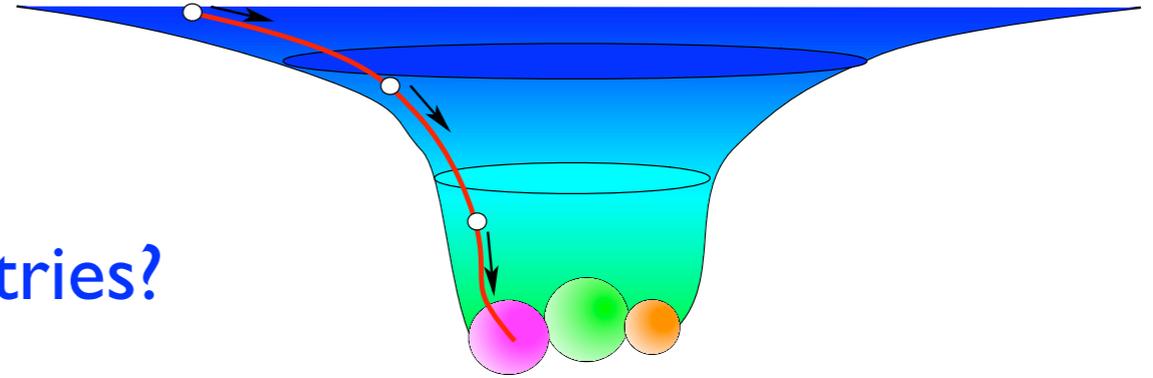
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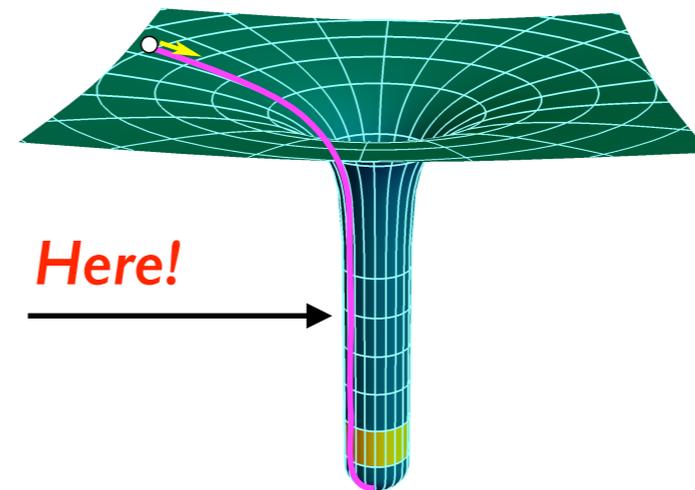
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What is the evolution of this scrambling? What is the back-reaction?
How is infalling matter incorporated into the microstate structure?

Return of information and Hawking Radiation

Green function calculations

Raju, arXiv:1804.10616

Bena, Guica, Heidemann, Monten, NPW arXiv:18XX.XXXX

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- ◆ Perturbative methods
 - *Non-BPS fluctuations and motions on moduli spaces*
 - *Topology of BPS microstate geometries make them robust under perturbation*
 - *BPS holographic dictionary: track the evolution in dual CFT*

New horizon-scale physics?

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Image credit: NASA/JPL-Caltech

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From RHIC and the study of quark-gluon plasmas:

Supergravity calculations in *Holographic Field Theories* are extremely good at giving **universal, effective hydrodynamics, like viscosity, coming from complex, strongly coupled quantum systems.**