Comprendre l'infiniment grand

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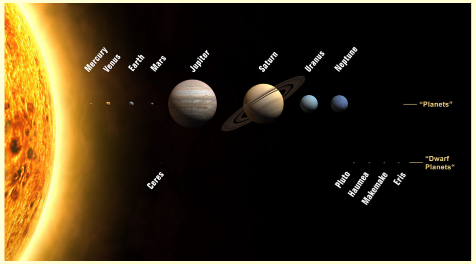
1) **Du modèle du Big Bang au contenu de l'Univers**
2) Mesures cosmologiques
3) Les grandes structures de l'Univers

The rise of the Big Bang model

Orders of magnitude. The three pillars of the Big Bang model

- Distances, cosmological scales
- Redshift
- Expansion of the Universe, the metric
- The rise of the Big Bang model

1. Distances



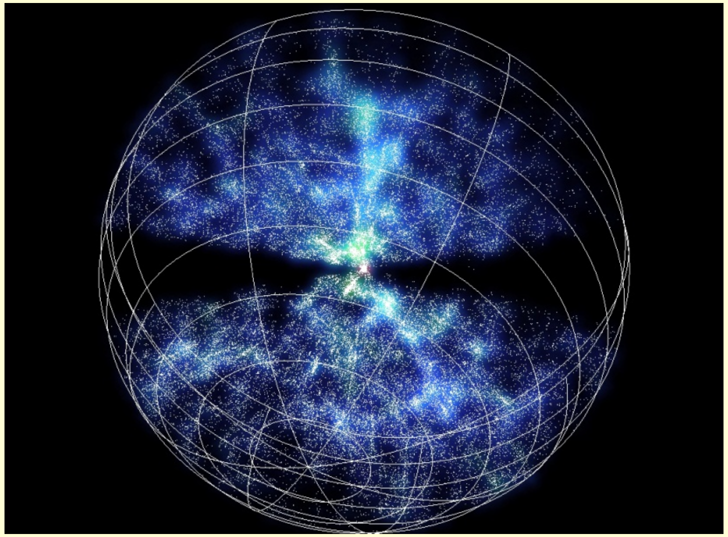
- solar system: Earth-Sun ~ 150. 10^6 km ~ 1 AU

- galaxies, eg the Milky Way:

  ∅ ~100,000 lyr ~ 30kpc Sun-Gal. center ~ 10kpc

1lyr=63,240 AU   1pc=3.26 lyr 1Mpc=3.3 10^6 lyr

- galaxy clusters: largest and most massive gravitationally bound structures, 50 to 1,000's of galaxies, ∅ ~ 2 to 10 Mpc



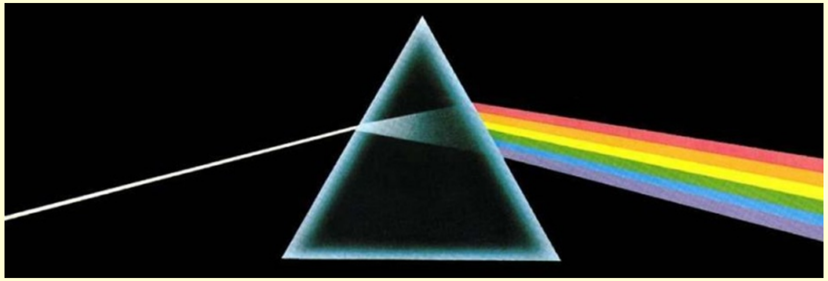
- large scale structures: galaxies → clusters → superclusters (15 - 100Mpc) making a network of voids (25 - 125Mpc) and filaments (90 - 300 Mpc)

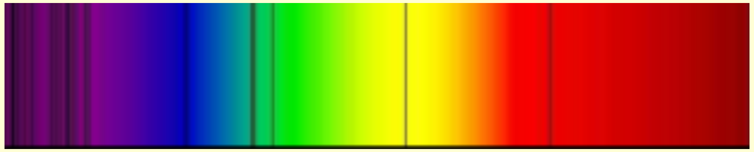
  ⇒ beyond 100Mpc, homogeneous and isotropic Universe

6dF Galaxy Redshift Survey, (2009)

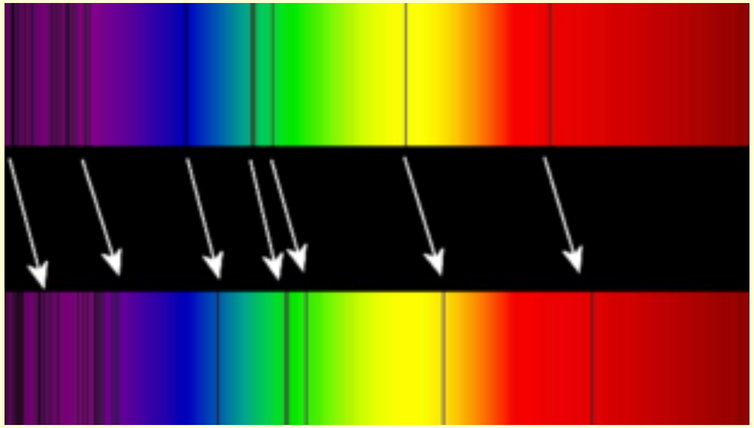
2. Redshift

- Emitted light spectrum ⇒ spectral lines ⇒ astro. object composition, environment … and motion relative to Earth.





- Redshift : the object moves away from us ⇒ λ_γ increases

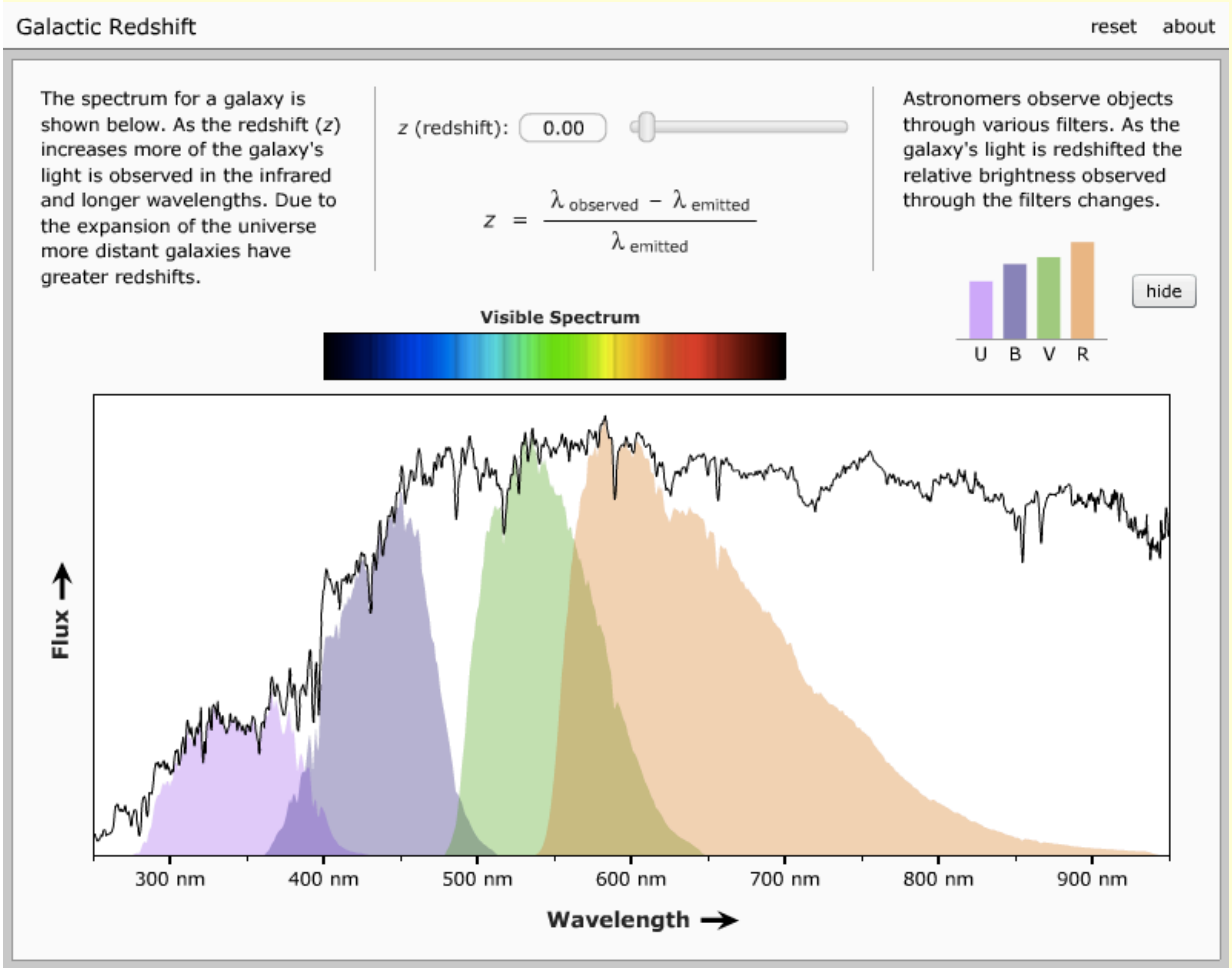


$$z \equiv \frac{\lambda_{observed} - \lambda_{emitted}}{\lambda_{emitted}}$$

or

$$1 + z = \frac{\lambda_{observed}}{\lambda_{emitted}}$$

See http://astro.unl.edu/classaction/animations/cosmology/galacticredshift.html

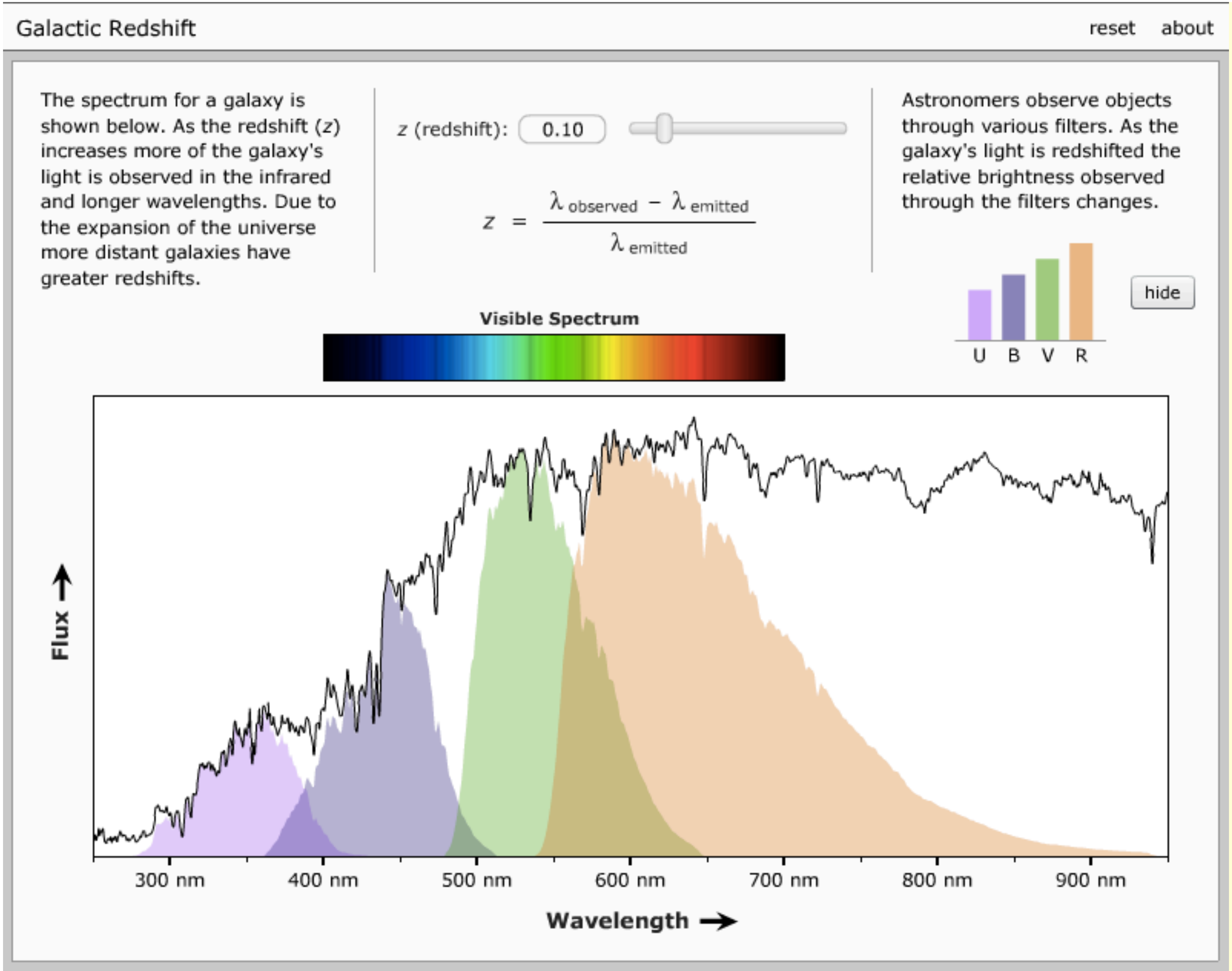


The spectrum for a galaxy is shown below. As the redshift (z) increases more of the galaxy's light is observed in the infrared and longer wavelengths. Due to the expansion of the universe more distant galaxies have greater redshifts.

z (redshift): 0.00

$$z = \frac{\lambda_{\text{observed}} - \lambda_{\text{emitted}}}{\lambda_{\text{emitted}}}$$

Astronomers observe objects through various filters. As the galaxy's light is redshifted the relative brightness observed through the filters changes.



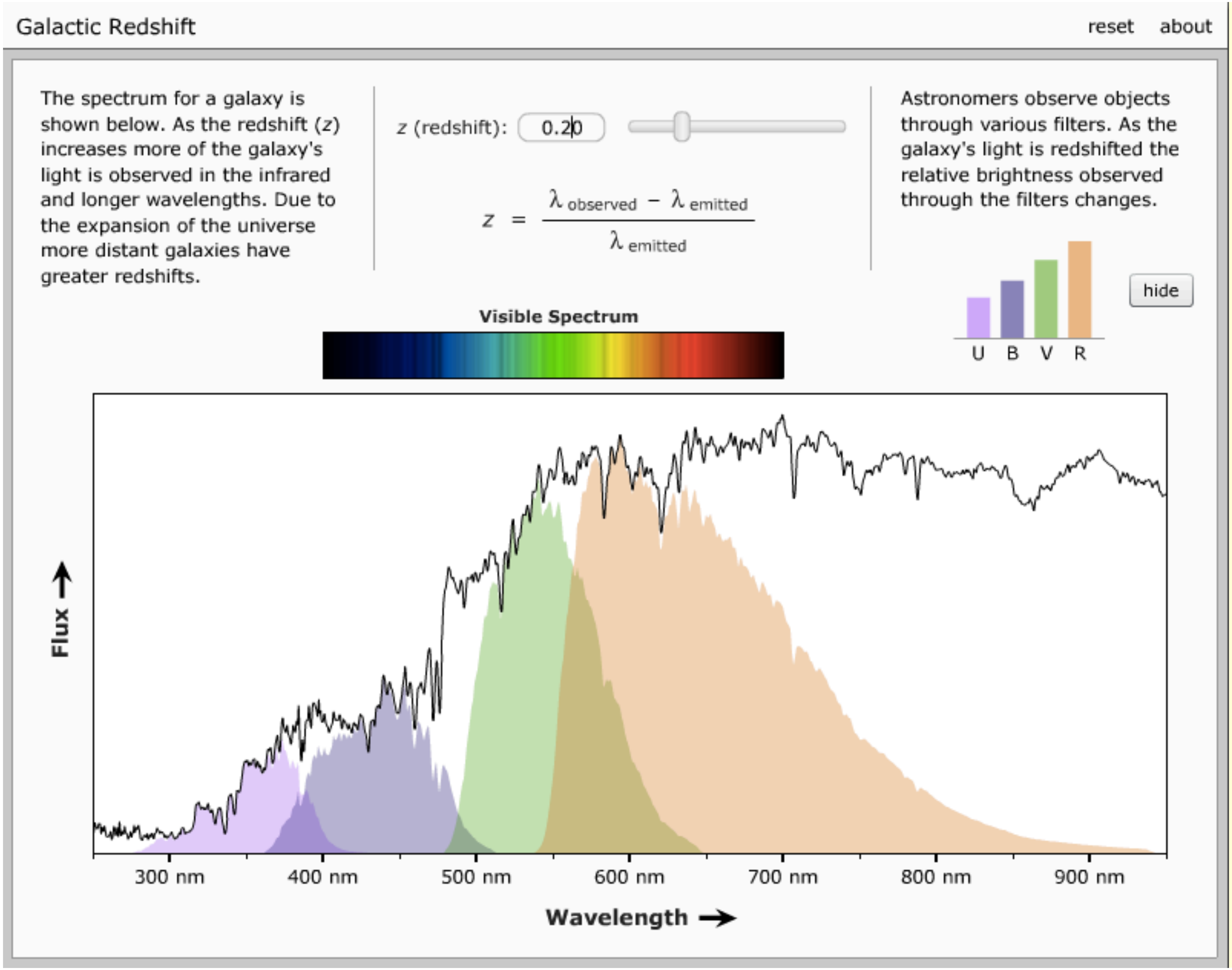
Galactic Redshift

The spectrum for a galaxy is shown below. As the redshift (z) increases more of the galaxy's light is observed in the infrared and longer wavelengths. Due to the expansion of the universe more distant galaxies have greater redshifts.

z (redshift): 0.10

$$z = \frac{\lambda_{\text{observed}} - \lambda_{\text{emitted}}}{\lambda_{\text{emitted}}}$$

Astronomers observe objects through various filters. As the galaxy's light is redshifted the relative brightness observed through the filters changes.

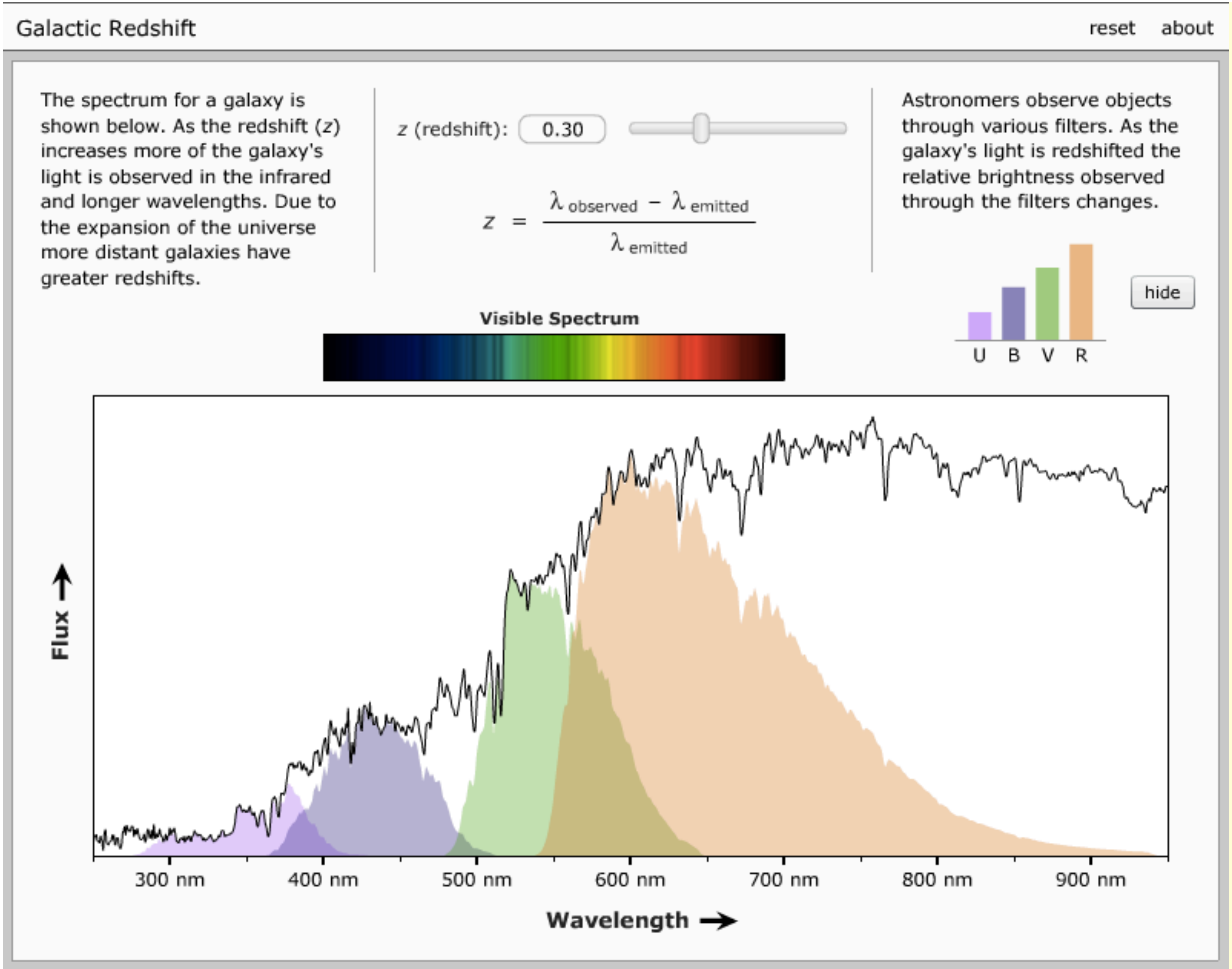


The spectrum for a galaxy is shown below. As the redshift (z) increases more of the galaxy's light is observed in the infrared and longer wavelengths. Due to the expansion of the universe more distant galaxies have greater redshifts.

z (redshift): 0.20

$$z = \frac{\lambda_{\text{observed}} - \lambda_{\text{emitted}}}{\lambda_{\text{emitted}}}$$

Astronomers observe objects through various filters. As the galaxy's light is redshifted the relative brightness observed through the filters changes.

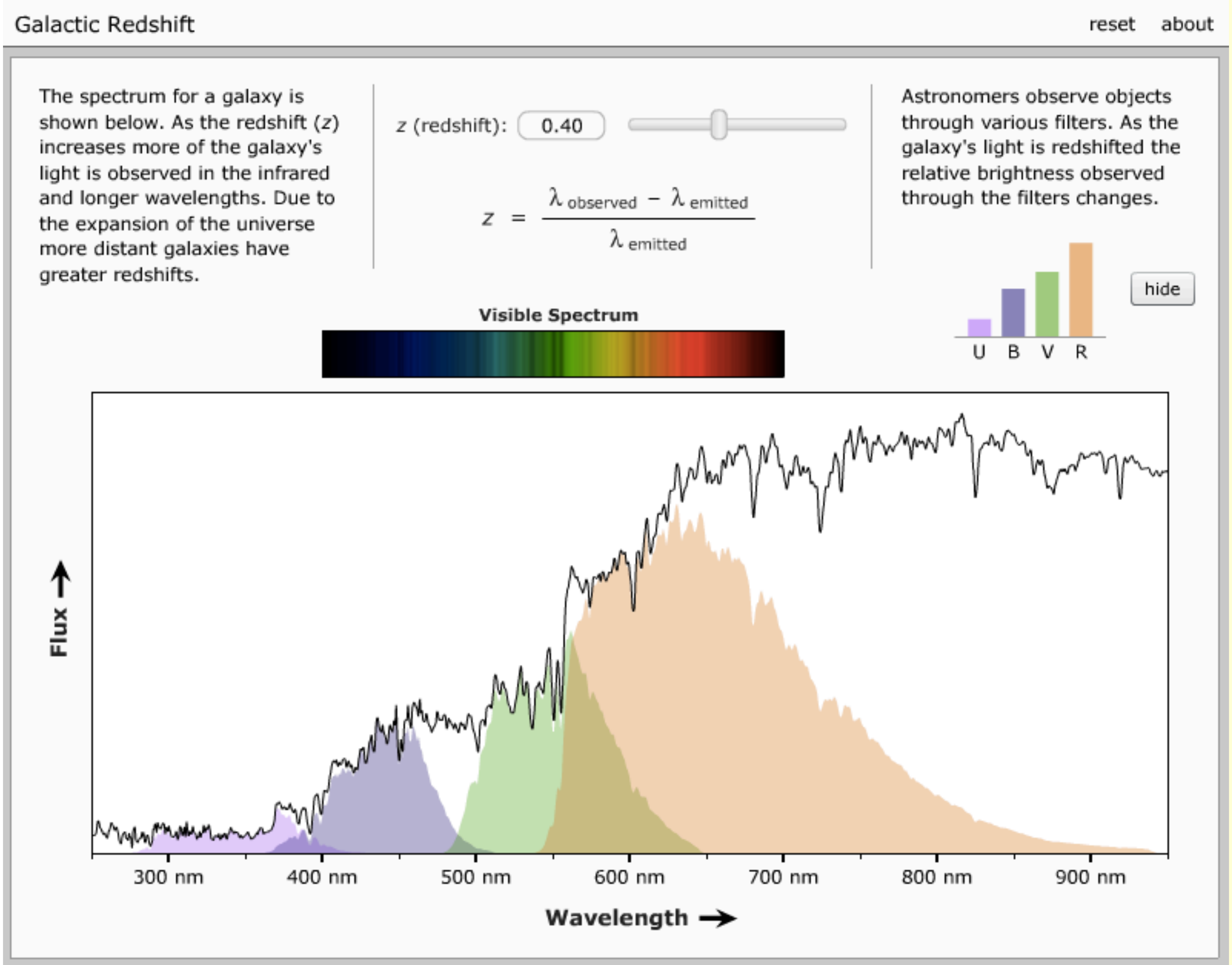


The spectrum for a galaxy is shown below. As the redshift (z) increases more of the galaxy's light is observed in the infrared and longer wavelengths. Due to the expansion of the universe more distant galaxies have greater redshifts.

z (redshift): 0.30

$$z = \frac{\lambda_{\text{observed}} - \lambda_{\text{emitted}}}{\lambda_{\text{emitted}}}$$

Astronomers observe objects through various filters. As the galaxy's light is redshifted the relative brightness observed through the filters changes.



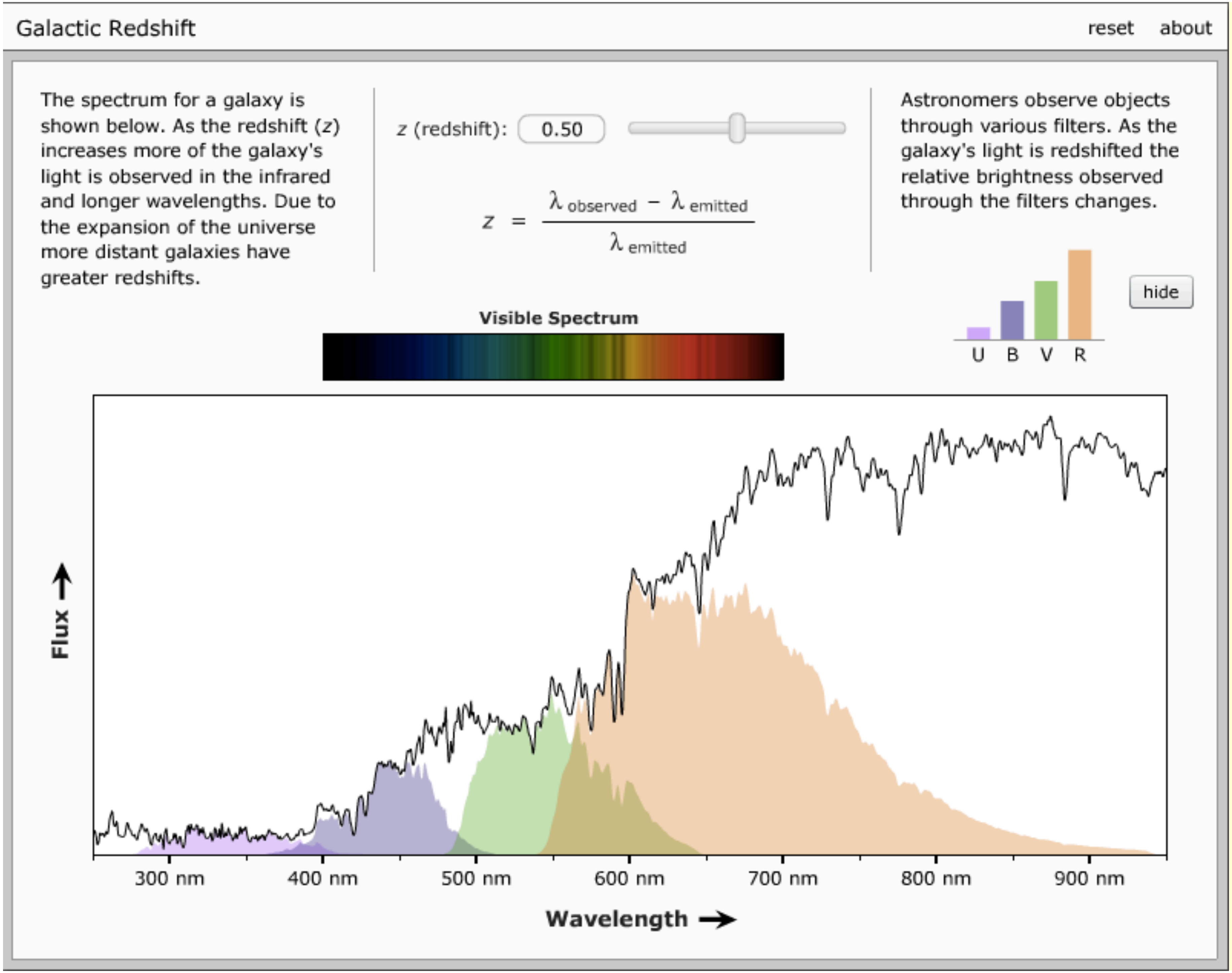
Galactic Redshift

The spectrum for a galaxy is shown below. As the redshift (z) increases more of the galaxy's light is observed in the infrared and longer wavelengths. Due to the expansion of the universe more distant galaxies have greater redshifts.

z (redshift): 0.40

$$z = \frac{\lambda_{\text{observed}} - \lambda_{\text{emitted}}}{\lambda_{\text{emitted}}}$$

Astronomers observe objects through various filters. As the galaxy's light is redshifted the relative brightness observed through the filters changes.



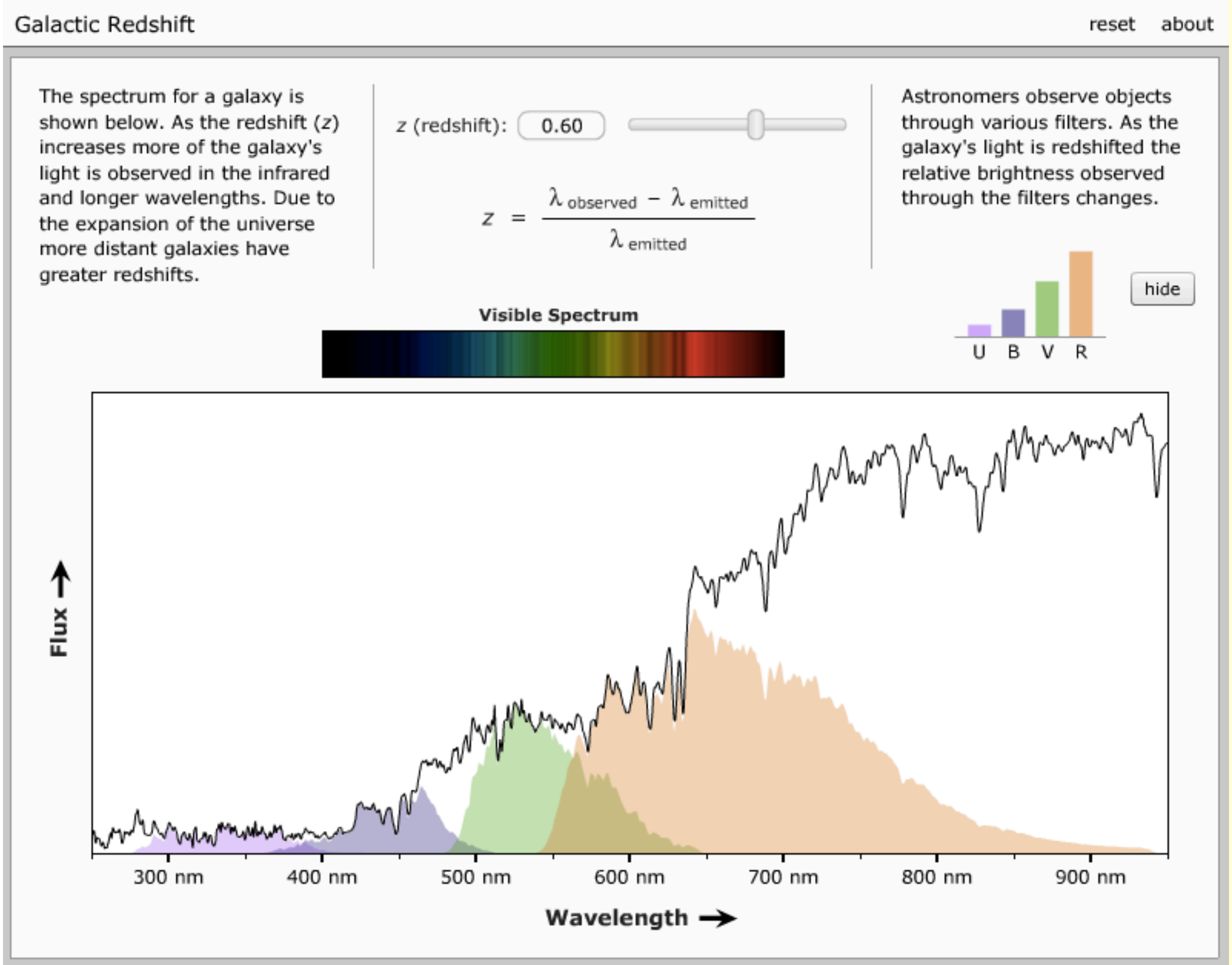
Galactic Redshift

The spectrum for a galaxy is shown below. As the redshift (z) increases more of the galaxy's light is observed in the infrared and longer wavelengths. Due to the expansion of the universe more distant galaxies have greater redshifts.

z (redshift): 0.50

$$z = \frac{\lambda_{\text{observed}} - \lambda_{\text{emitted}}}{\lambda_{\text{emitted}}}$$

Astronomers observe objects through various filters. As the galaxy's light is redshifted the relative brightness observed through the filters changes.

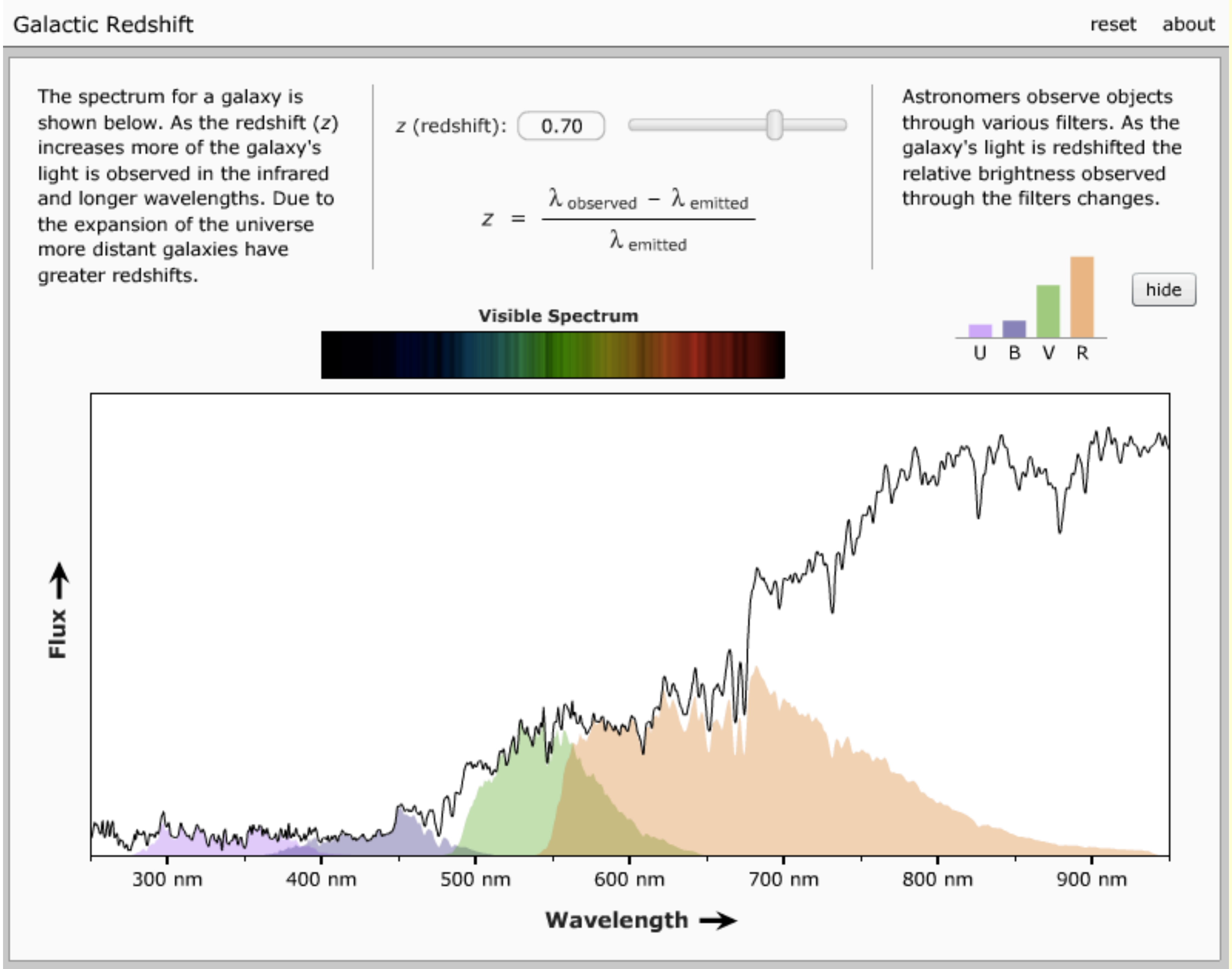


The spectrum for a galaxy is shown below. As the redshift (z) increases more of the galaxy's light is observed in the infrared and longer wavelengths. Due to the expansion of the universe more distant galaxies have greater redshifts.

z (redshift): 0.60

$$z = \frac{\lambda_{\text{observed}} - \lambda_{\text{emitted}}}{\lambda_{\text{emitted}}}$$

Astronomers observe objects through various filters. As the galaxy's light is redshifted the relative brightness observed through the filters changes.



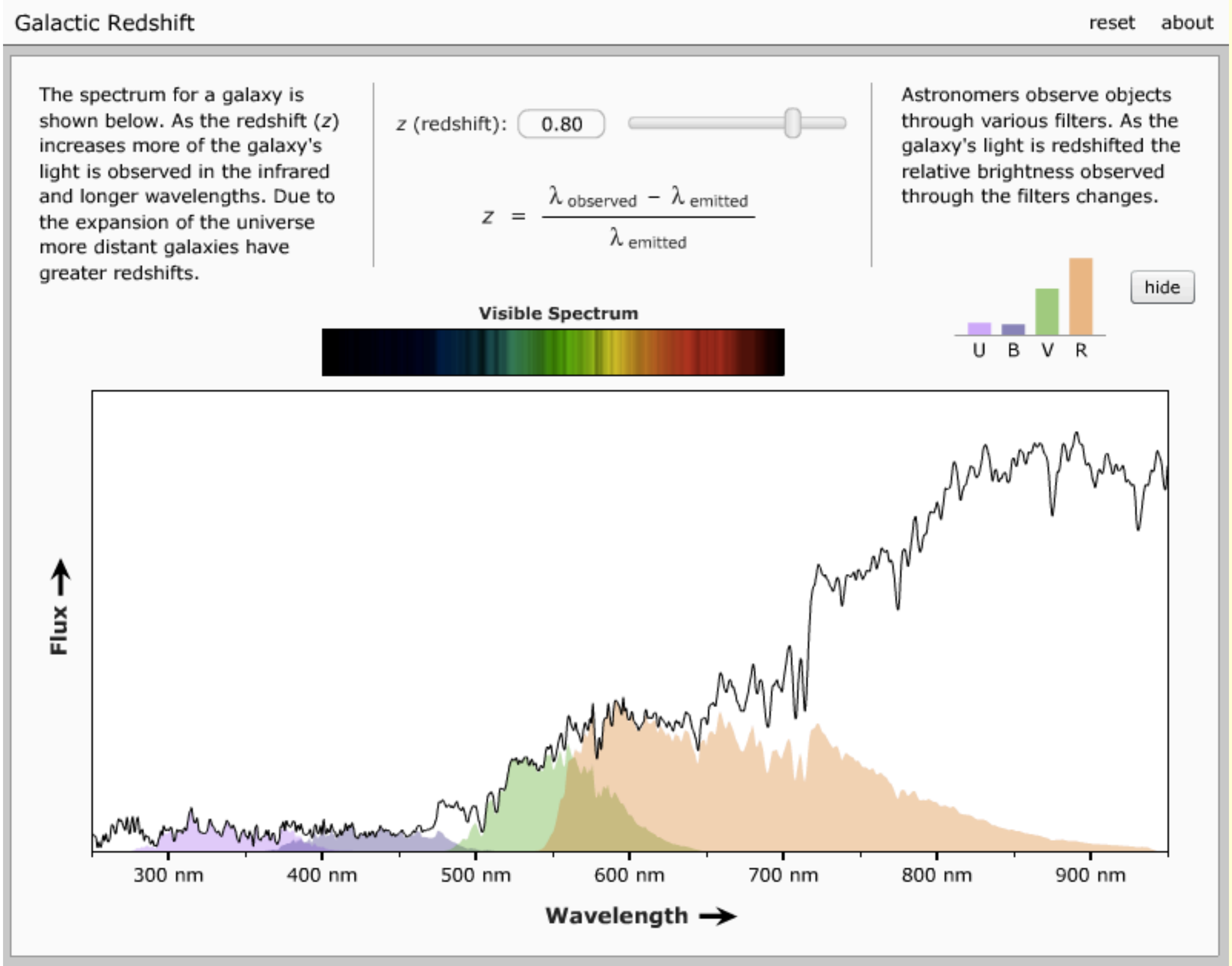
Galactic Redshift

The spectrum for a galaxy is shown below. As the redshift (z) increases more of the galaxy's light is observed in the infrared and longer wavelengths. Due to the expansion of the universe more distant galaxies have greater redshifts.

z (redshift): 0.70

$$z = \frac{\lambda_{\text{observed}} - \lambda_{\text{emitted}}}{\lambda_{\text{emitted}}}$$

Astronomers observe objects through various filters. As the galaxy's light is redshifted the relative brightness observed through the filters changes.

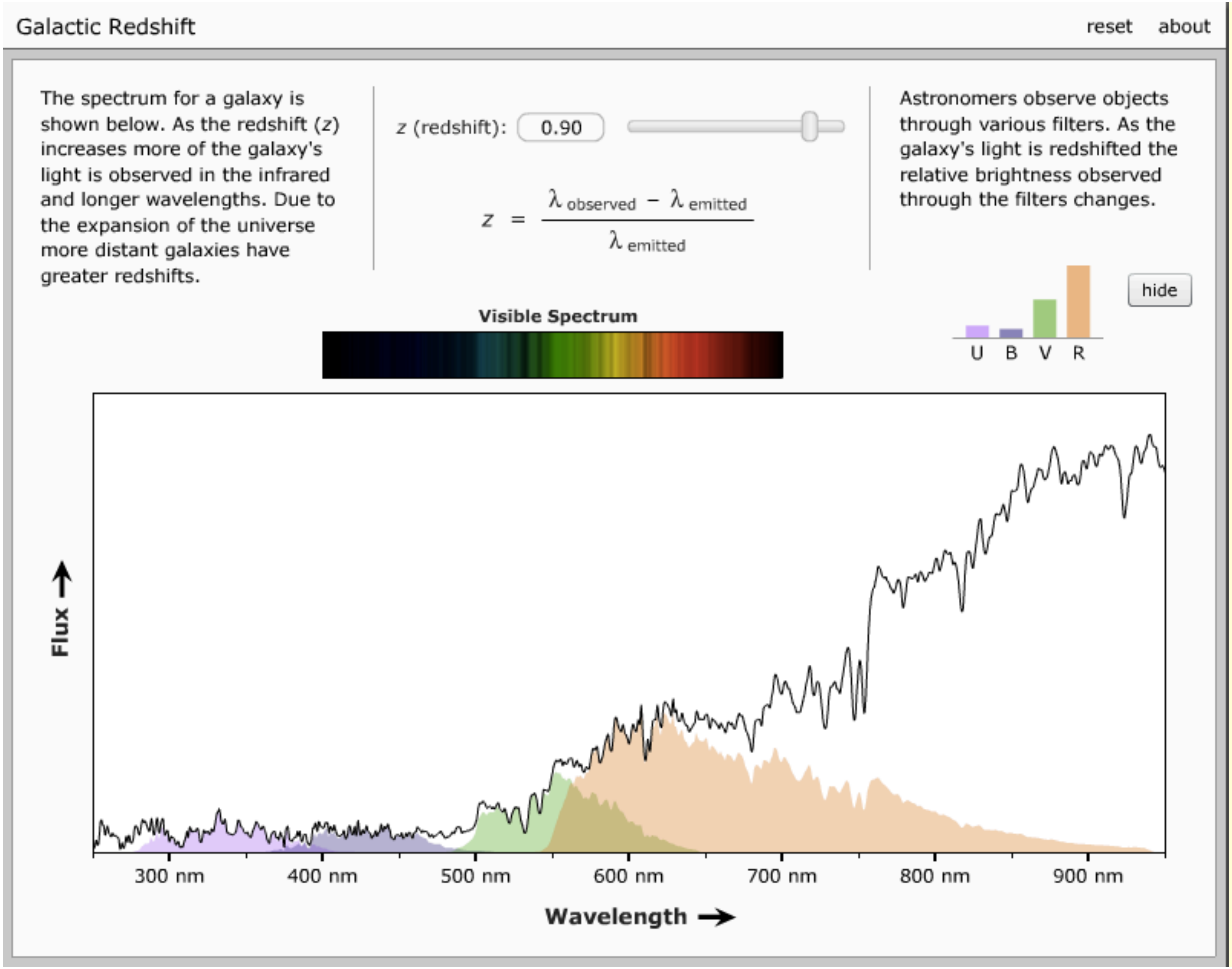


The spectrum for a galaxy is shown below. As the redshift (*z*) increases more of the galaxy's light is observed in the infrared and longer wavelengths. Due to the expansion of the universe more distant galaxies have greater redshifts.

z (redshift): 0.80

$$z = \frac{\lambda_{\text{observed}} - \lambda_{\text{emitted}}}{\lambda_{\text{emitted}}}$$

Astronomers observe objects through various filters. As the galaxy's light is redshifted the relative brightness observed through the filters changes.



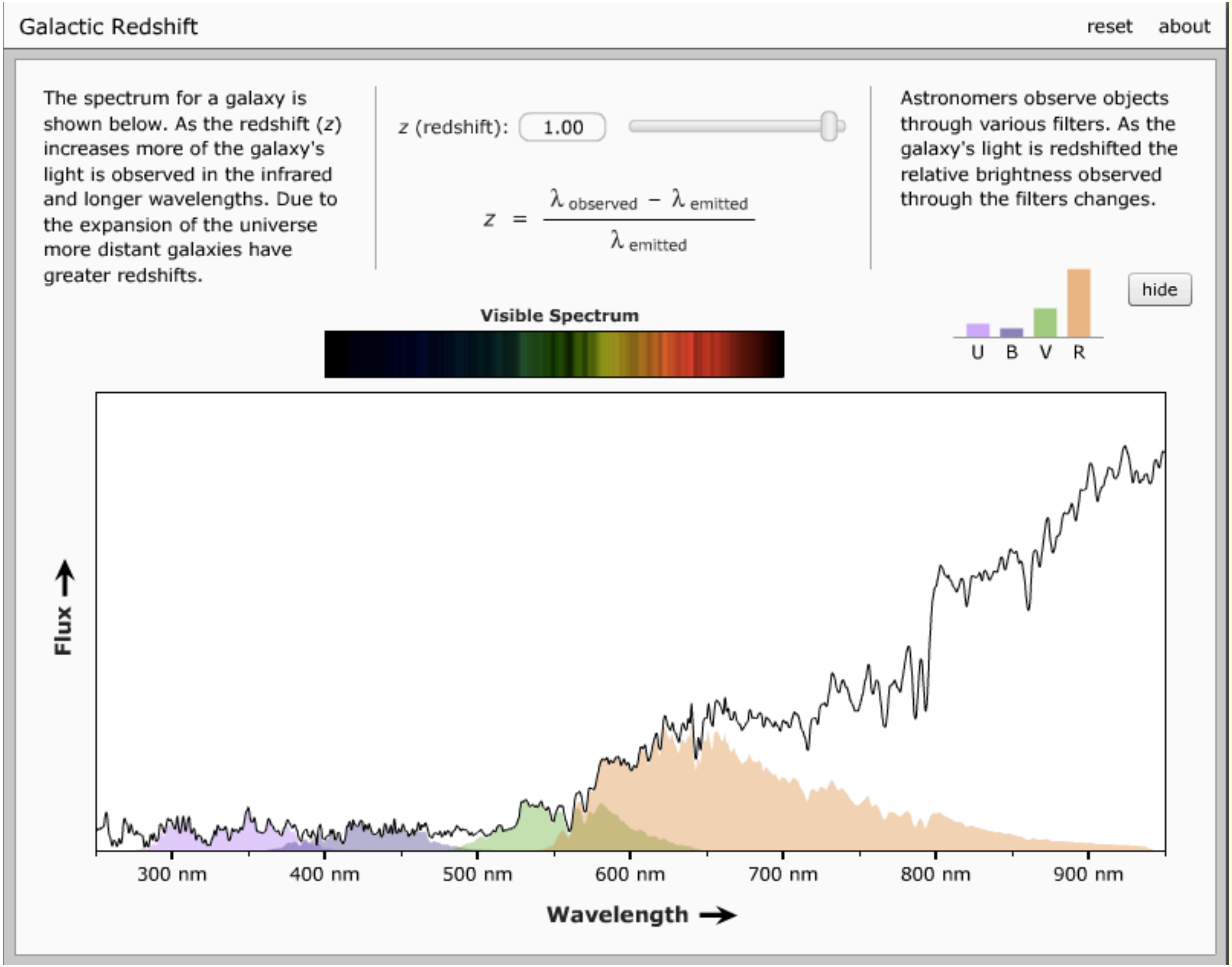
Galactic Redshift

The spectrum for a galaxy is shown below. As the redshift (*z*) increases more of the galaxy's light is observed in the infrared and longer wavelengths. Due to the expansion of the universe more distant galaxies have greater redshifts.

z (redshift): 0.90

$$z = \frac{\lambda_{\text{observed}} - \lambda_{\text{emitted}}}{\lambda_{\text{emitted}}}$$

Astronomers observe objects through various filters. As the galaxy's light is redshifted the relative brightness observed through the filters changes.



The spectrum for a galaxy is shown below. As the redshift (z) increases more of the galaxy's light is observed in the infrared and longer wavelengths. Due to the expansion of the universe more distant galaxies have greater redshifts.

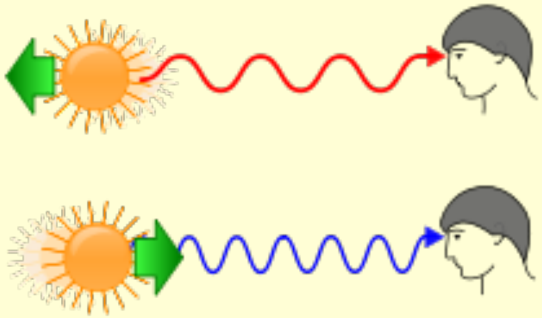
z (redshift): 1.00

$$z = \frac{\lambda_{\text{observed}} - \lambda_{\text{emitted}}}{\lambda_{\text{emitted}}}$$

Astronomers observe objects through various filters. As the galaxy's light is redshifted the relative brightness observed through the filters changes.

Different origins of redshifts/blueshifts

- Doppler effect : redshift/blueshift due to relative motion



$$1+z=\gamma\left(1+\frac{v_{//}}{c}\right) \qquad z\approx\frac{v_{//}}{c} \quad \textit{for small } v_{//}$$

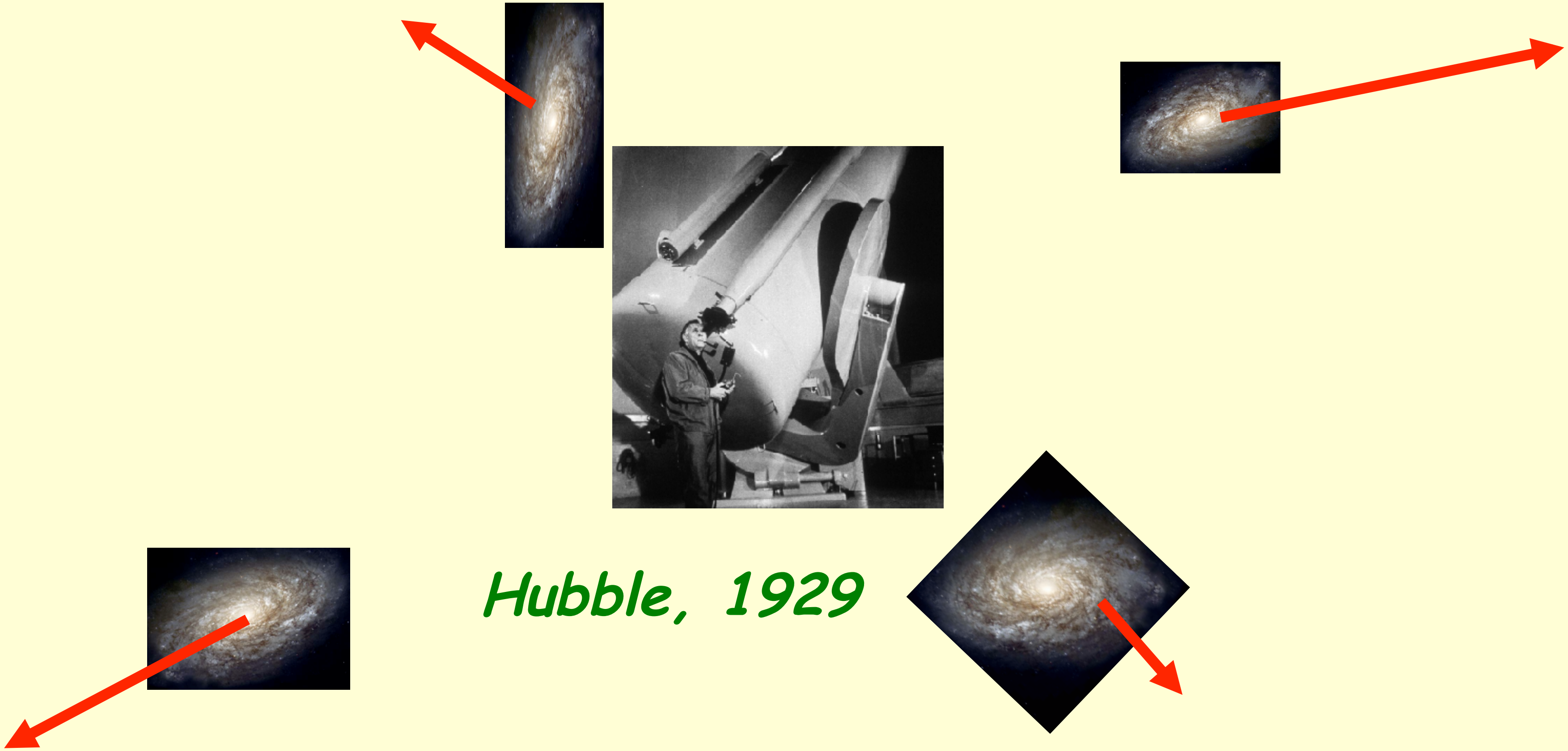
(in Minkowski space i.e. flat spacetime)

- Cosmological redshift : dominant for distant sources (above tens of Mpc or $z>0.01$). Due to Universe expansion.

- Gravitational red/blue shift : radiation moving out of/into a gravitational field.

e.g. grav. redshift $1+z=\dfrac{1}{\sqrt{1-2GM/rc^2}}$

The rise of modern cosmology



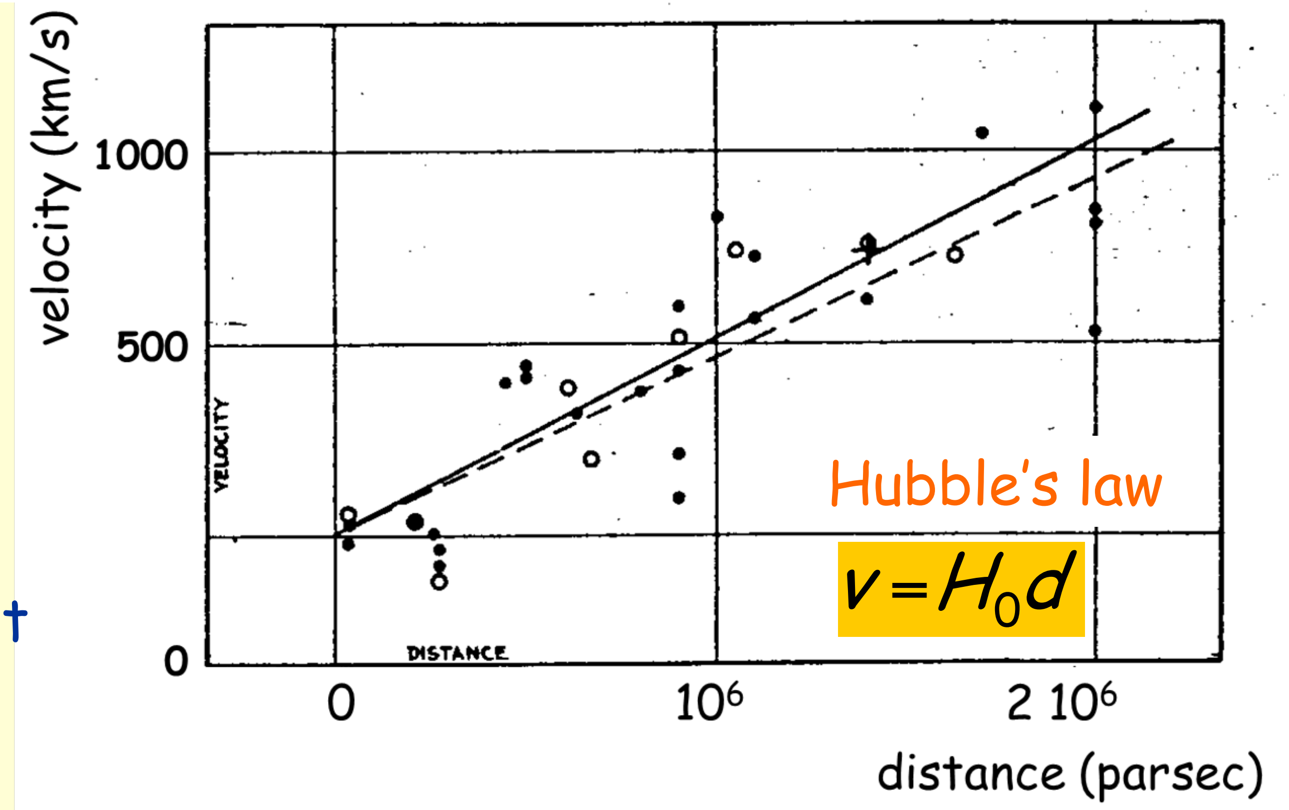
Hubble, 1929

⇨ distant galaxies are receding : the Universe is in expansion
(as predicted in General Relativity by A.Friedmann 1922 and G.Lemaître 1927)

Hubble data

Redshift
converted
into
velocity

$z \sim v/c$



Hubble's law

$v = H_0 d$

- H_0: Hubble constant
- recent, precise measurements (2011): $H_0 = 73.8 \pm 2.4 km/s/Mpc$

  (2014): $H_0 = 70.6 \pm 3.3 km/s/Mpc$

  (2019): $H_0 = 74.0 \pm 1.4 km/s/Mpc$
- critical density today

$$\rho_c^0 = 3H_0^2 \Big/ 8\pi G = 1.04 h^2 10^{10} eV \cdot m^{-3} \qquad h = H_0 \Big/ 100 kms^{-1}\ Mpc^{-1}$$

$\Updownarrow$

same energy density as $1 gal/Mpc^3 \sim 5 p/m^3$

Describing an expanding universe

- General Relativity : the simplest relativistic theory of gravitaty consistent with data. Gravity described as a geometric property of spacetime.

- Metric: allows to compute distances between two points

$$ds^2 \equiv g_{\mu\nu} dx^\mu dx^\nu$$

$g_{\mu\nu}$: metric tensor ds^2 : line element, invariant

  - Reminder: in special relativity (no gravity):

$$ds^2 \equiv \eta_{\mu\nu} dx^\mu dx^\nu = c^2 dt^2 - dx^2 - dy^2 - dz^2$$

$$= c^2 dt^2 - \left(dr^2 + r^2 \left(d\theta^2 + \sin\theta^2 d\varphi^2 \right) \right)$$

$\eta_{\mu\nu}$: Minkowski metric

  - in a particle's rest frame:

$$ds = c \; d\tau$$

$d\tau$: particle proper time

- Friedmann-Lemaître-Robertson-Walker (FLRW) metric : metric for a spatially homogeneous and isotropic expanding universe, with scale/expansion factor a(t) and curvature k

$$ds^2 = c^2 dt^2 - a(t)^2 \left(\frac{dr^2}{\boxed{1-kr^2}} + r^2 \left(d\theta^2 + \sin^2\theta \, d\varphi^2 \right) \right)$$

expansion factored out

$$H(t) \equiv \frac{\dot{a}(t)}{a(t)}$$ expansion rate or Hubble parameter

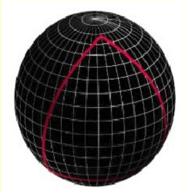
$$H_0 \equiv H(t_0) \quad t_0 = today$$

r,θ,φ spherical comoving coordinates:

r : dimensionless & stationary wrt expansion

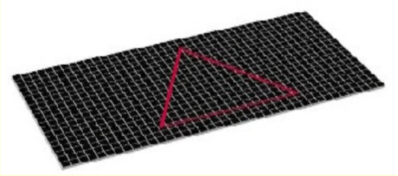
k, gaussian curvature:

closed universe



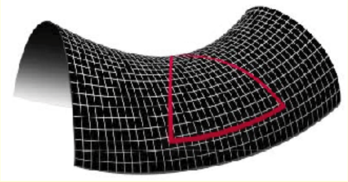
k=+1

flat universe



k=0

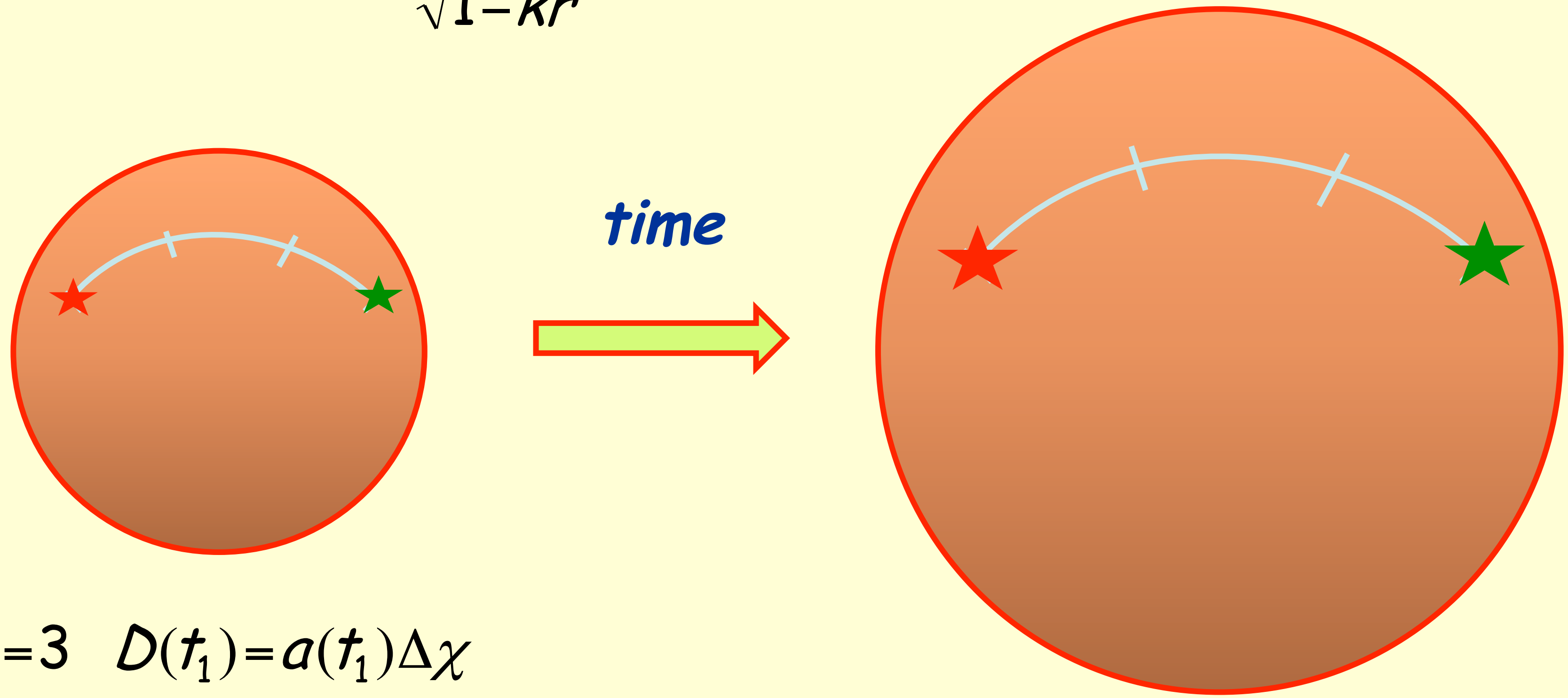
open universe



k=-1

More on FLRW metric

- Introducing $d\chi \equiv \frac{dr}{\sqrt{1-kr^2}}$ to account for curvature :



$\Delta\chi = 3 \quad D(t_1) = a(t_1)\Delta\chi$

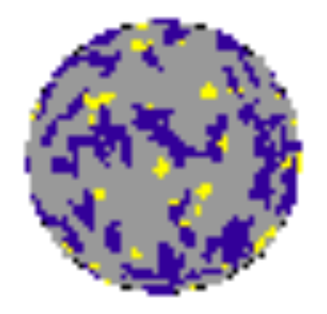
$\Delta\chi = 3 \quad D(t_2) = a(t_2)\Delta\chi \geq D(t_1)$

a(t): scale factor

χ: comoving coordinate, stationary wrt expansion

D(t) : proper/physical distance $\Delta\chi$: comoving distance

The cosmological redshift



Universe in expansion

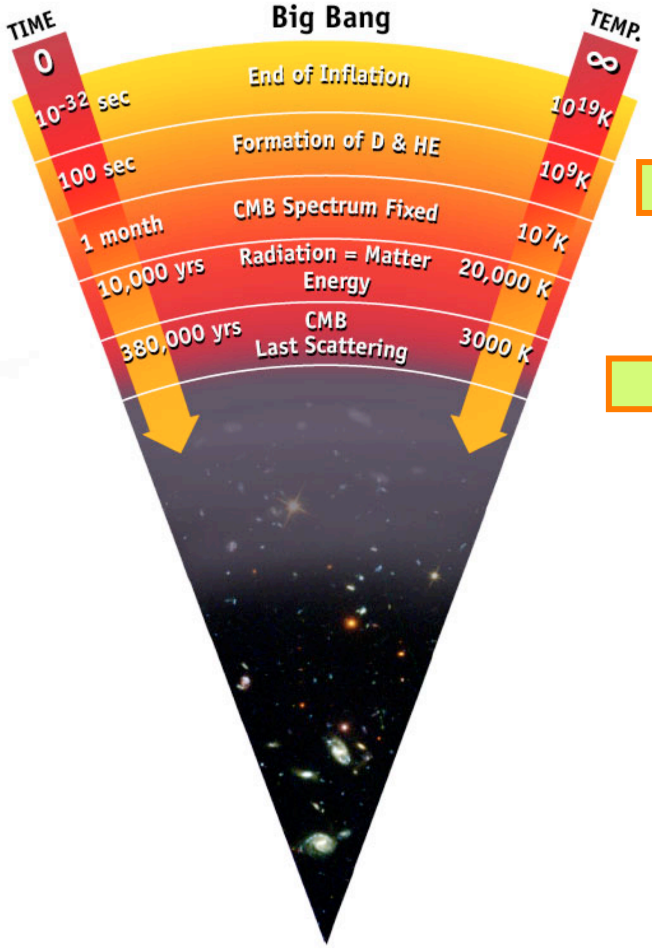
↓

light from distant sources is redshifted

$$\frac{\lambda_{observed}}{\lambda_{emitted}} \equiv 1+z = \frac{a(t_{observation})}{a(t_{emission})}$$

©Ned Wright, UCLA

Initial dense and hot phase followed by expansion and cooling



13,8 billions of years after Big Bang

More confirmations of the Big Bang model

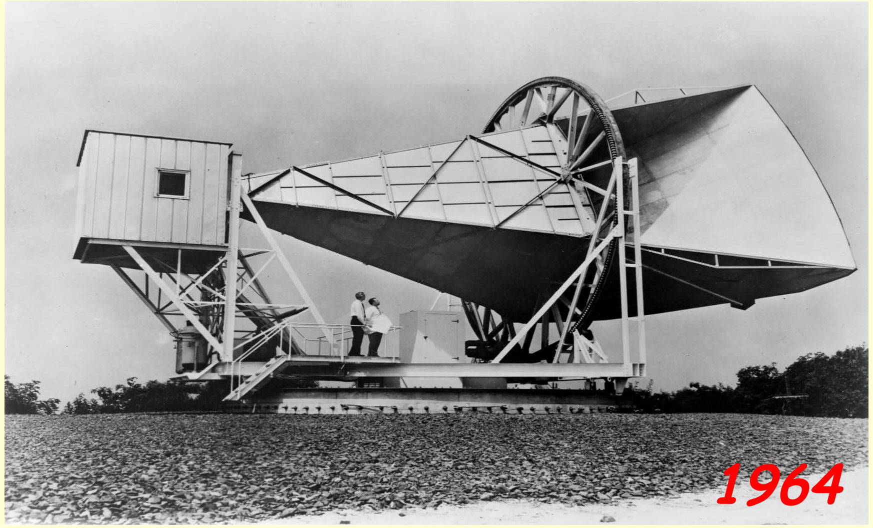
Primordial nucleosynthesis (He, D): T~1MeV predicted light element abundances = data

$$\frac{n_B - n_{\bar{B}}}{n_\gamma} \approx 10^{-10}$$

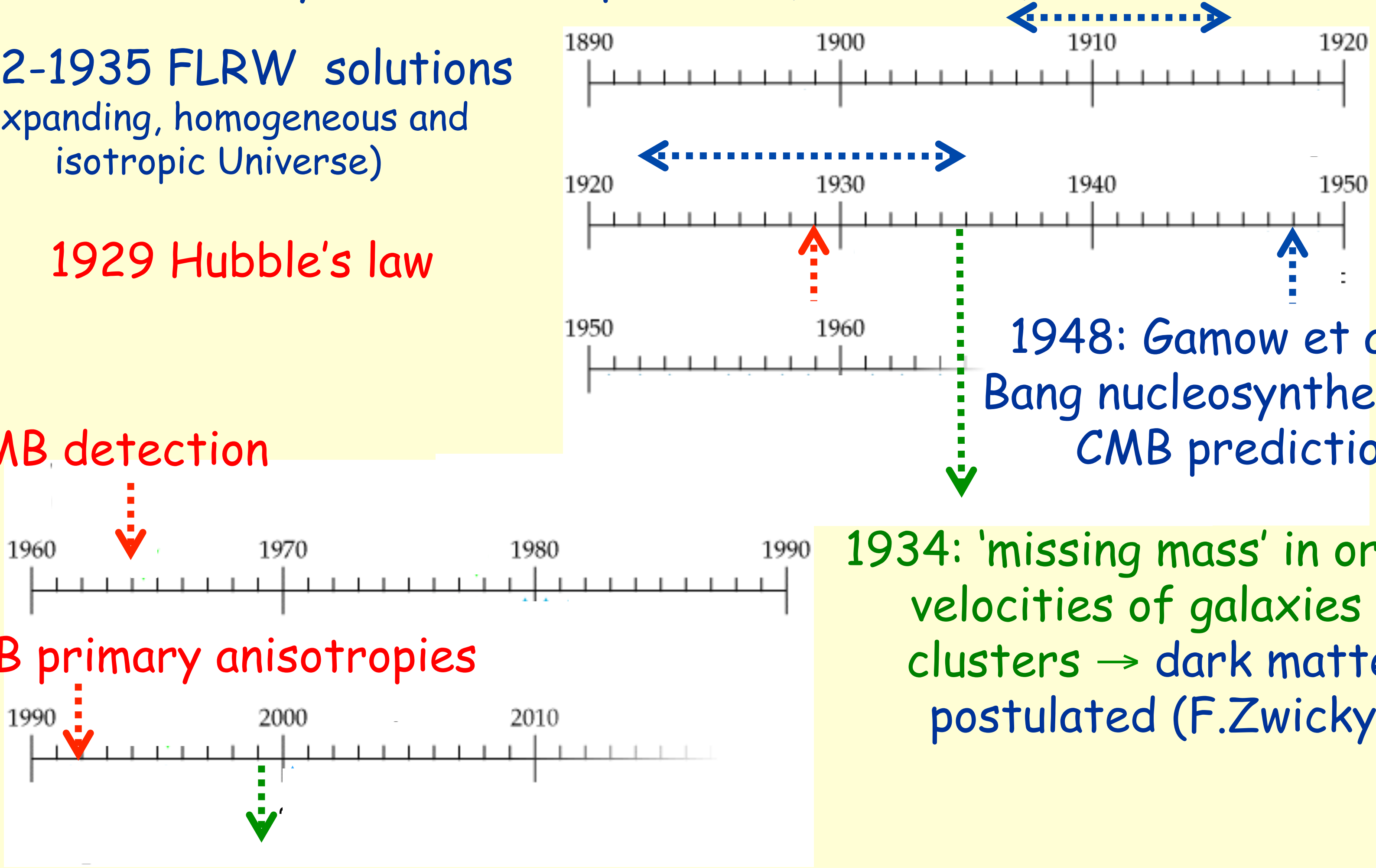
1948: G.Gamow, R.Alpher

Matter-radiation decoupling: z~1100 T~3000K **C**osmic **M**icrowave **B**ackground, relic radiation predicted and observed (T~2.725K 1992)

1948: G.Gamow, R.Alpher, R.Herman
1964: A.Penzias, R.Wilson



The rise of the Big Bang model



General Relativity, Einstein's equations (1907-1915)

1922-1935 FLRW solutions
(expanding, homogeneous and
isotropic Universe)

1929 Hubble's law

1948: Gamow et al: Big
Bang nucleosynthesis and
CMB predictions

1934: 'missing mass' in orbital
velocities of galaxies in
clusters → dark matter
postulated (F.Zwicky)

CMB detection

CMB primary anisotropies

Accelerated expansion: dark energy (or modified Einstein's gravity,
e.g. cosmological constant)

CONCLUSIONS (1)

- The Universe is spatially homogeneous and isotropic on cosmological scales (above 100Mpc)

- The Universe, initially in a hot, dense phase, is expanding

$$H_0 \equiv H(t_{today}) \approx 70 km/s/Mpc$$

  - cosmological redshift: due to the Universe expansion

$$\frac{\lambda_{observed}}{\lambda_{emitted}} \equiv 1+z = \frac{a(t_{observation})}{a(t_{emission})}$$

- General relativity and FLRW metric are the basics of the Big Bang model

- Observational confirmations: primordial abundances, CMB

Content of the Universe

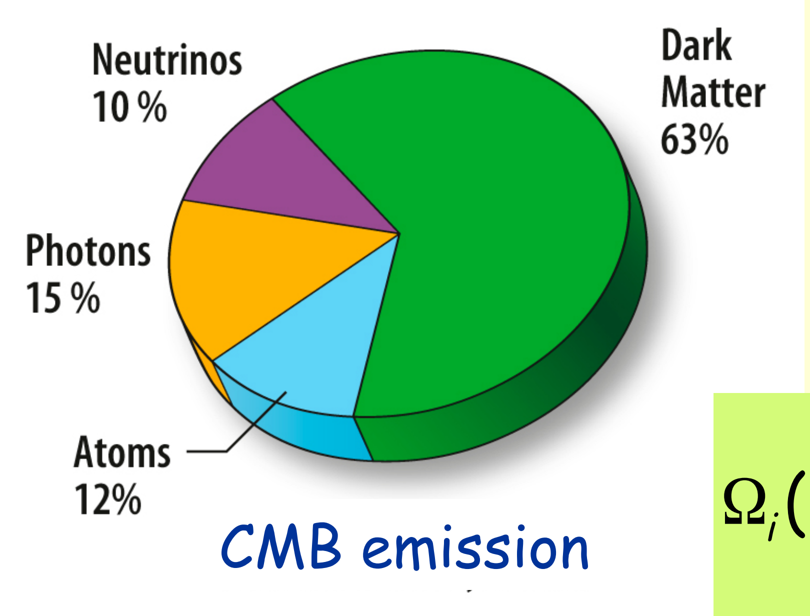
From General Relativiy to the matter-energy content of the Universe

- The Ω_i parameters and the expansion history
- Thermal history of the Universe

1. Densities ρ_i and Ω_i parameters

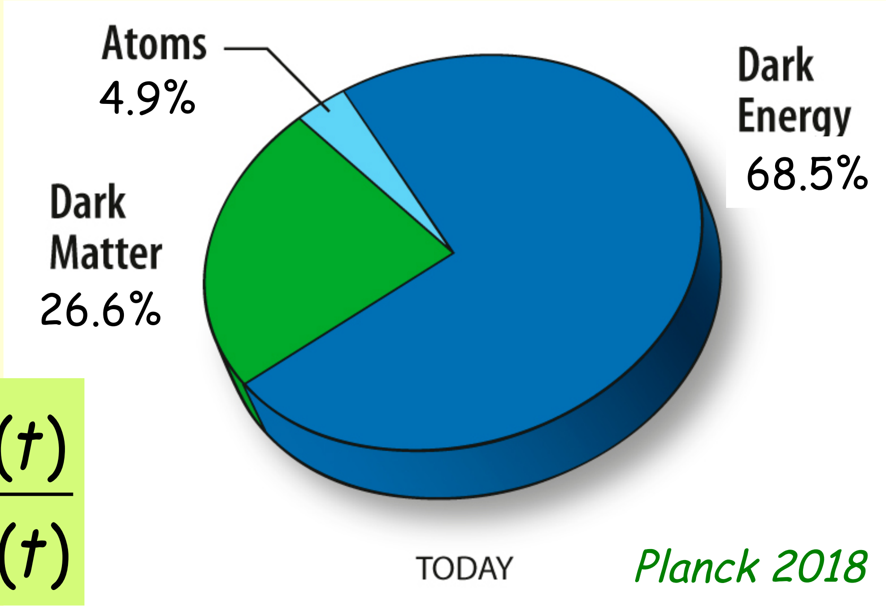
Big Bang cosmology model+ **many** cosmological data

energy balance of the Universe



CMB emission

$$\Omega_i(t) \equiv \frac{\rho_i(t)}{\rho_c(t)}$$



Planck 2018

$$1 = \Omega_m(t) + \underbrace{\Omega_r(t) + \Omega_k(t)}_{} + \Omega_\Lambda(t)$$

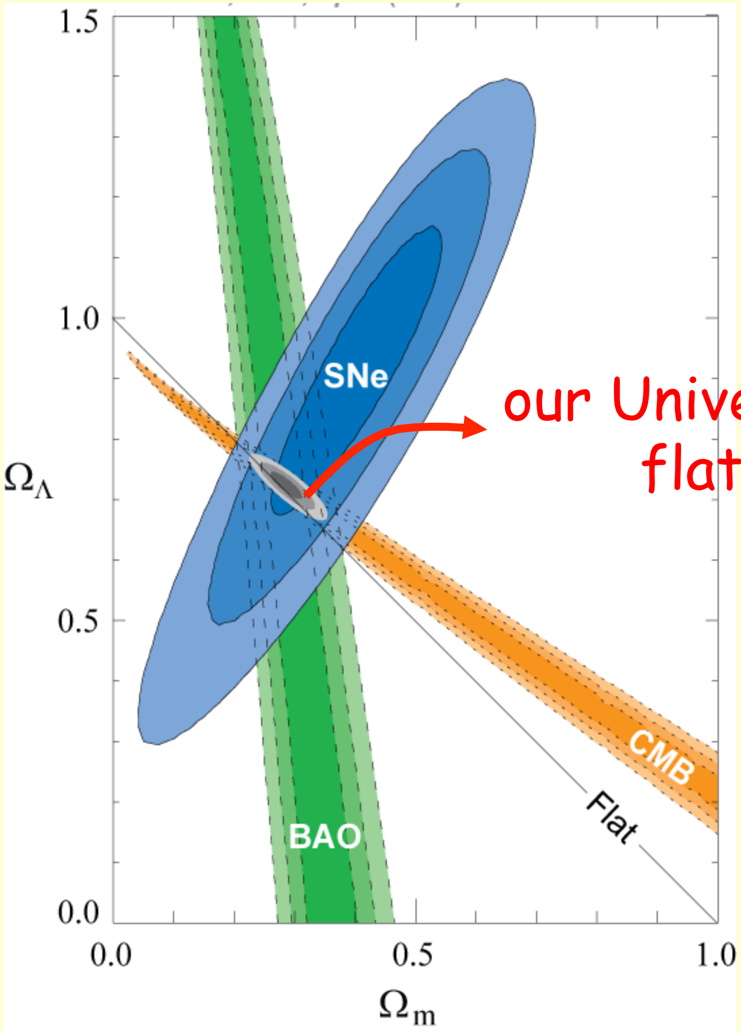
NR matter; radiation; curvature; "dark energy"

negligible today

e.g. cosmological constant Λ

Present values of the Ω_i's : the Ω_i^0 values

- CMB + BAO + SNe Ia data :



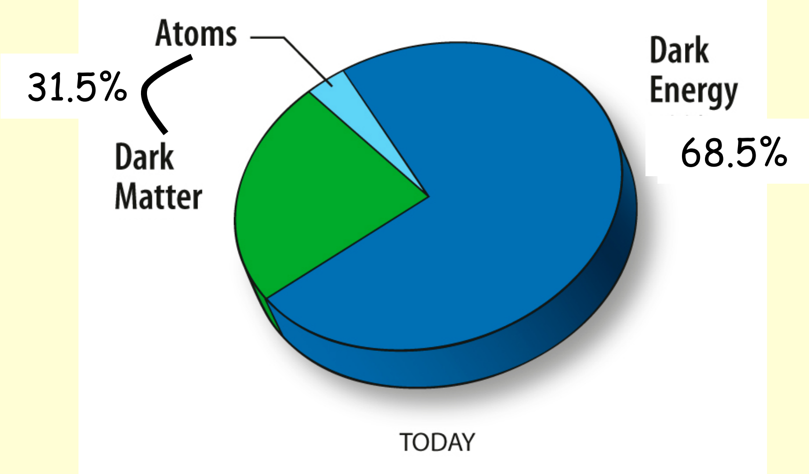
our Universe is flat !

M. Kowalski et al., 2008, ApJ, 686, 749

2.2% precision

$$\Omega_m^0 = 0.315 \pm 0.007$$
$$\Omega_\Lambda^0 = 0.685 \pm 0.007$$

Planck collaboration. 2018, arXiv:1807.06209



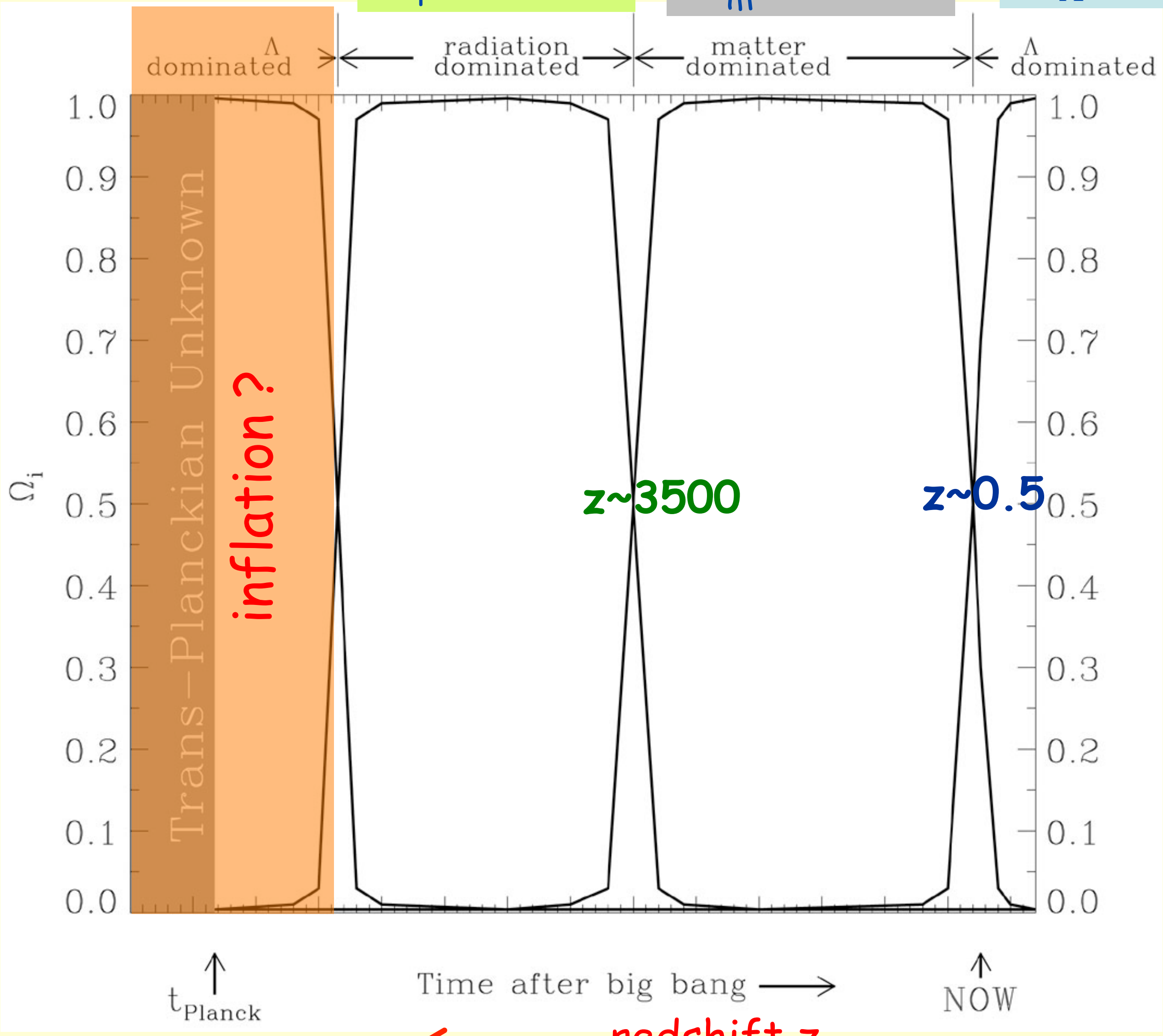
atoms: gas (4%), stars (0.4%), ν (0.3%), heavy elements (0.03%)

The rise and fall of the Ω_i's

$\Omega_r \propto a^{-4} H^{-2}$

$\Omega_m \propto a^{-3} H^{-2}$

$\Omega_\Lambda \propto H^{-2}$



redshift z

Expansion and temperature

- Photons emitted at t (e.g. CMB), received today:

$$1+z \equiv \frac{\lambda_{observed}}{\lambda_{emitted}} = \frac{a_0}{a(t)} = \frac{E_\gamma(t)}{E_\gamma(t_0)} = \frac{T(t)}{T_0}$$

at emission

today

- Temperature was hotter in the past, expansion implies cooling down
- CMB now: $T_{CMB}^{now} \approx 2.725K \Rightarrow T_0 \approx 2.35\ 10^{-4}\,eV \Rightarrow \Omega_\gamma^0 \approx 5.10^{-5}$

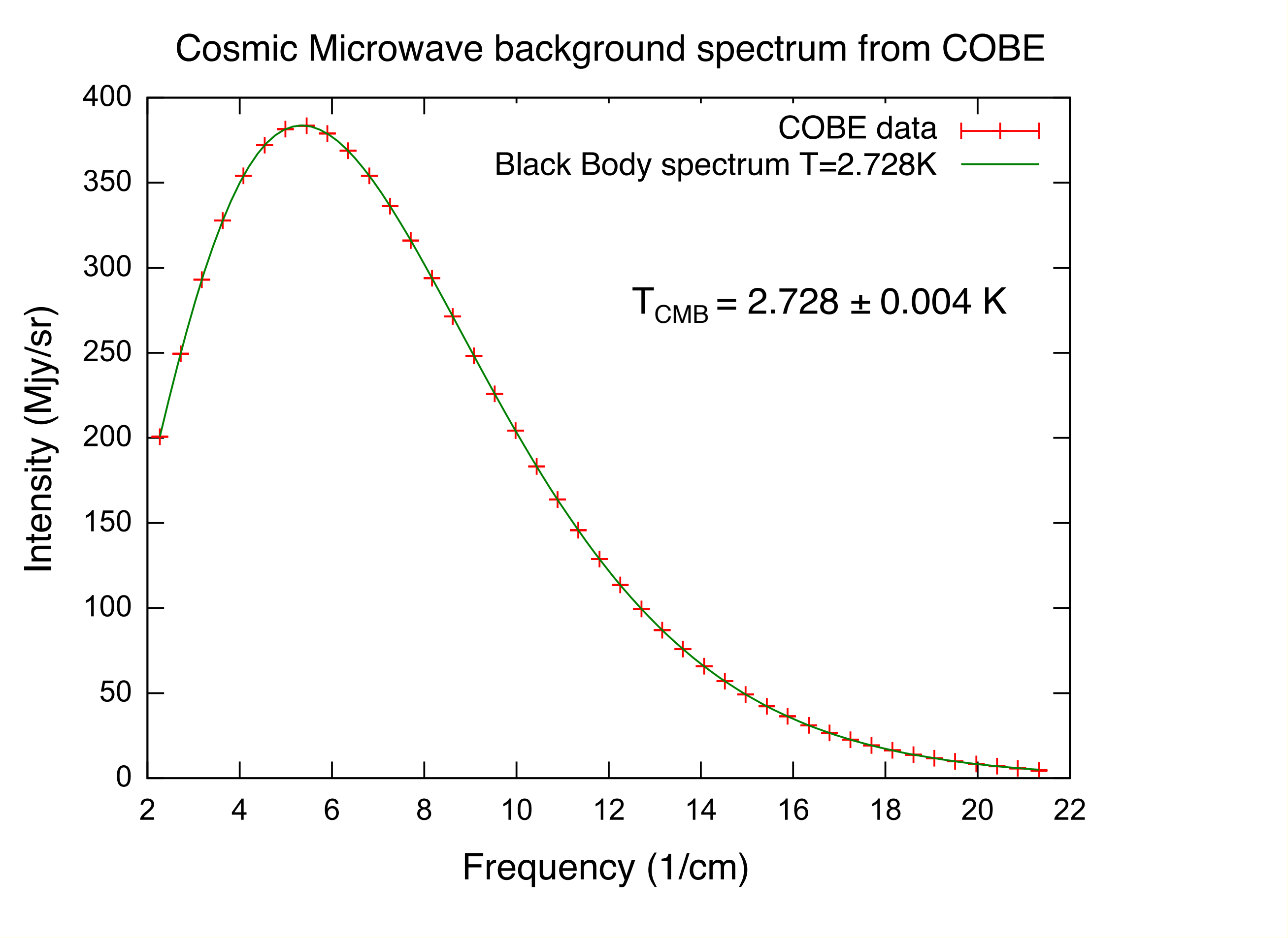
  (COBE) $(k = 8.617\ 10^{-5}\,eV.K^{-1})$ negligible (today)
- CMB at emission (from anisotropy measurements):

$$z_{CMB}^{emission} \approx 1100 \Rightarrow T_{CMB}(t_{emission}) \approx 0.26eV \approx 3{,}000K$$

- matter-radiation equality:

$$\rho_r(t_{eq}) = \rho_m(t_{eq}) \Rightarrow a_0/a(t_{eq}) \Rightarrow T_{eq} \approx 1eV \quad z \approx 3500$$

The CMB spectrum as seen by COBE (1989-1993)



D.J. Fixsen et al. 1996, ApJ, 473, 576F

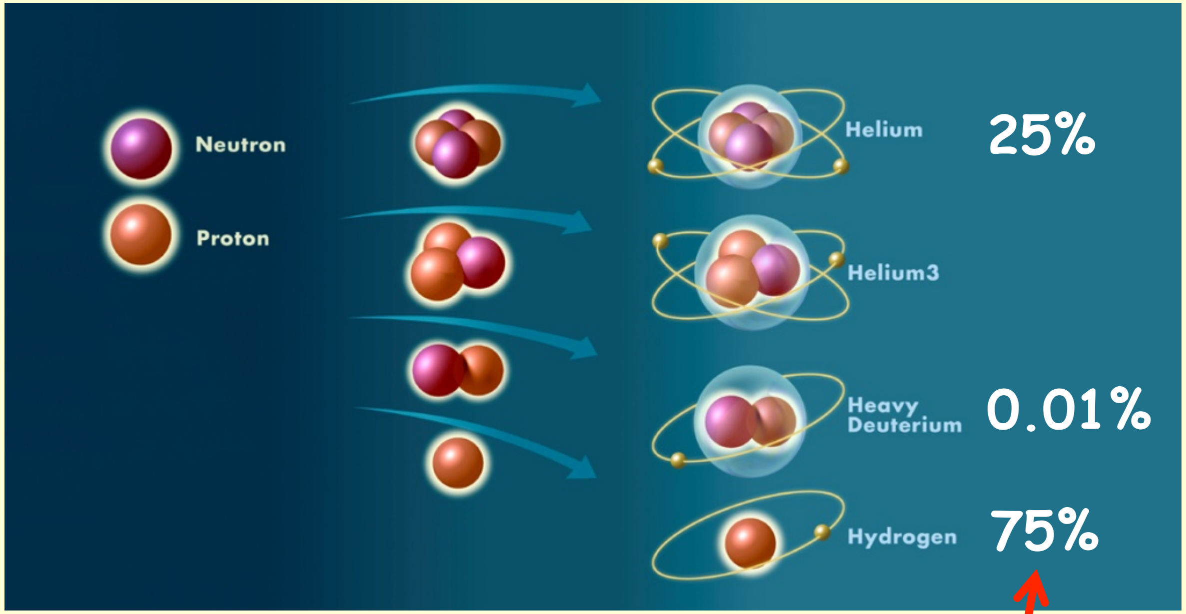
- Final : $T_{CMB} = 2.7255 \pm 0.0006$ K

D.J. Fixsen, 2009, AJ,707,916

Primordial nucleosynthesis

- 1948: G.Gamow, R.Alpher "The origin of Chemical Elements"

Expanding universe : deuterium and helium nuclei are formed by nuclear reactions inside the primordial plasma of p,n,e,γ when temperature and densities are adequate, leading to light element abundances as measured.



25%

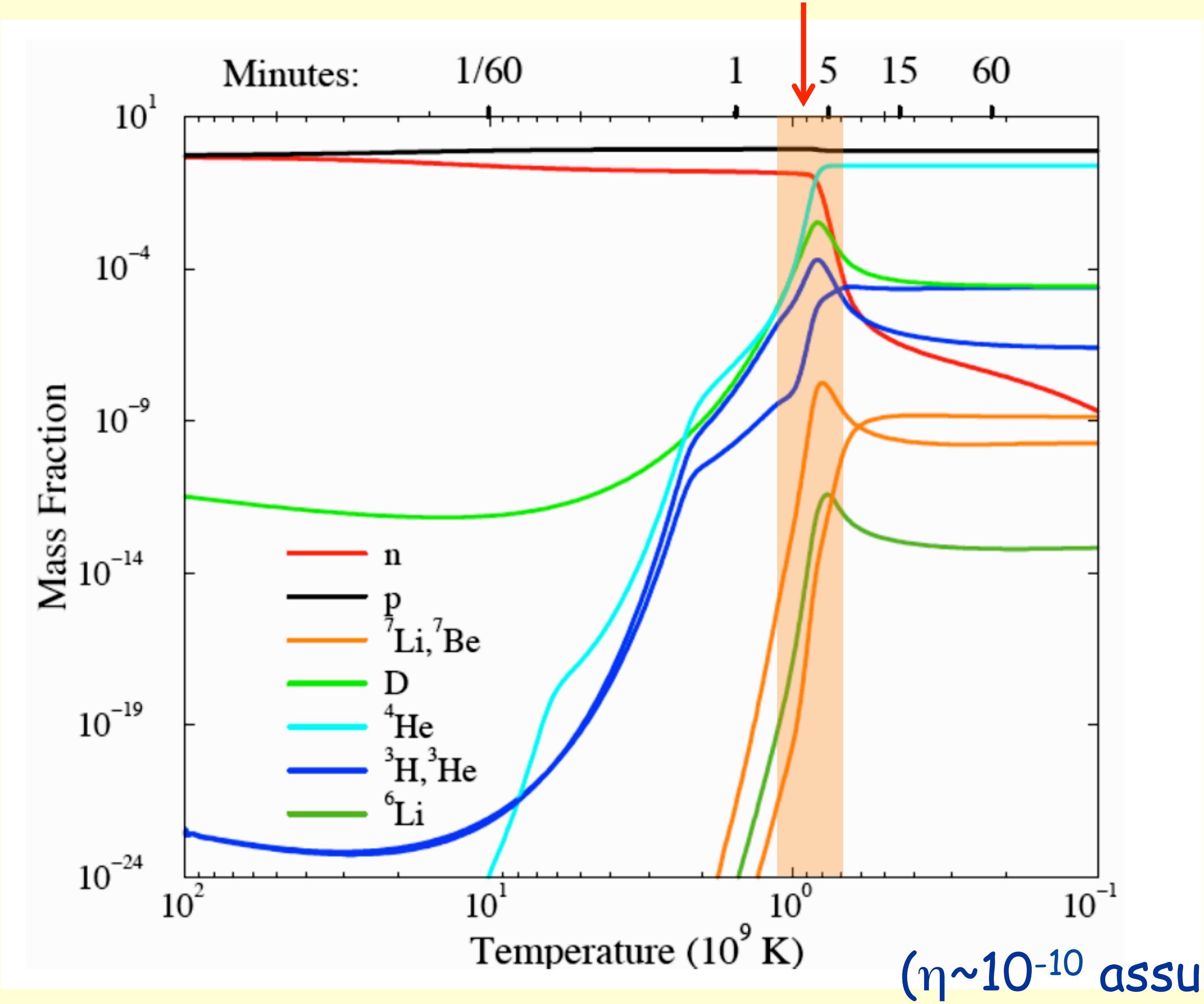
0.01%

75%

measured primordial abundances

The evolution of abundances

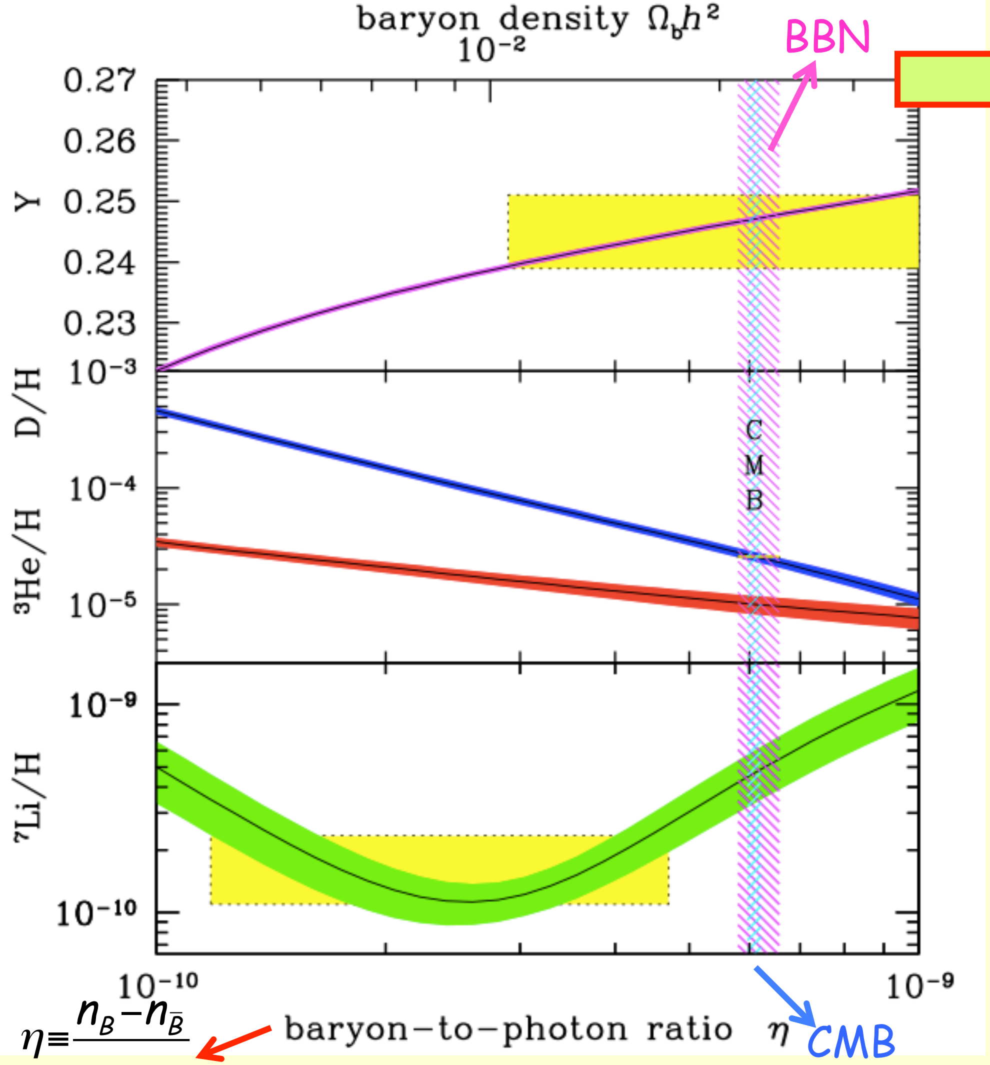
t_{BBN}~3mn : abundances frozen



(η~10^{-10} assumed)

Present values of the Ω_i's

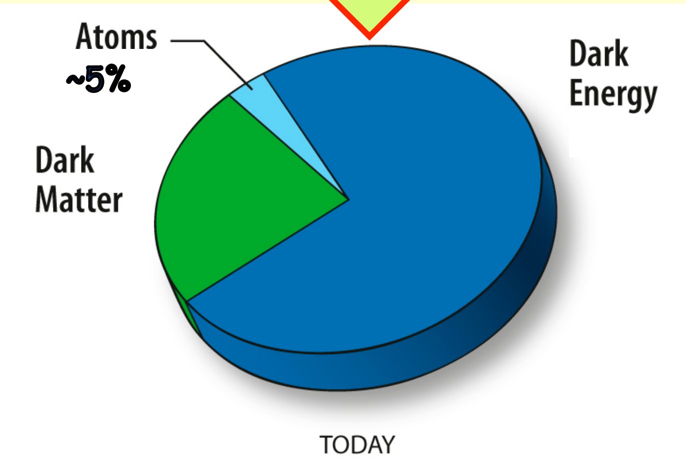
- Measurements of light element abundances vs predictions:



$$\eta \equiv \frac{n_B - n_{\bar{B}}}{n_\gamma}$$

$$0.046 \leq \Omega_b^0 \leq 0.053$$

(confirmed independently by CMB - Planck 2018)

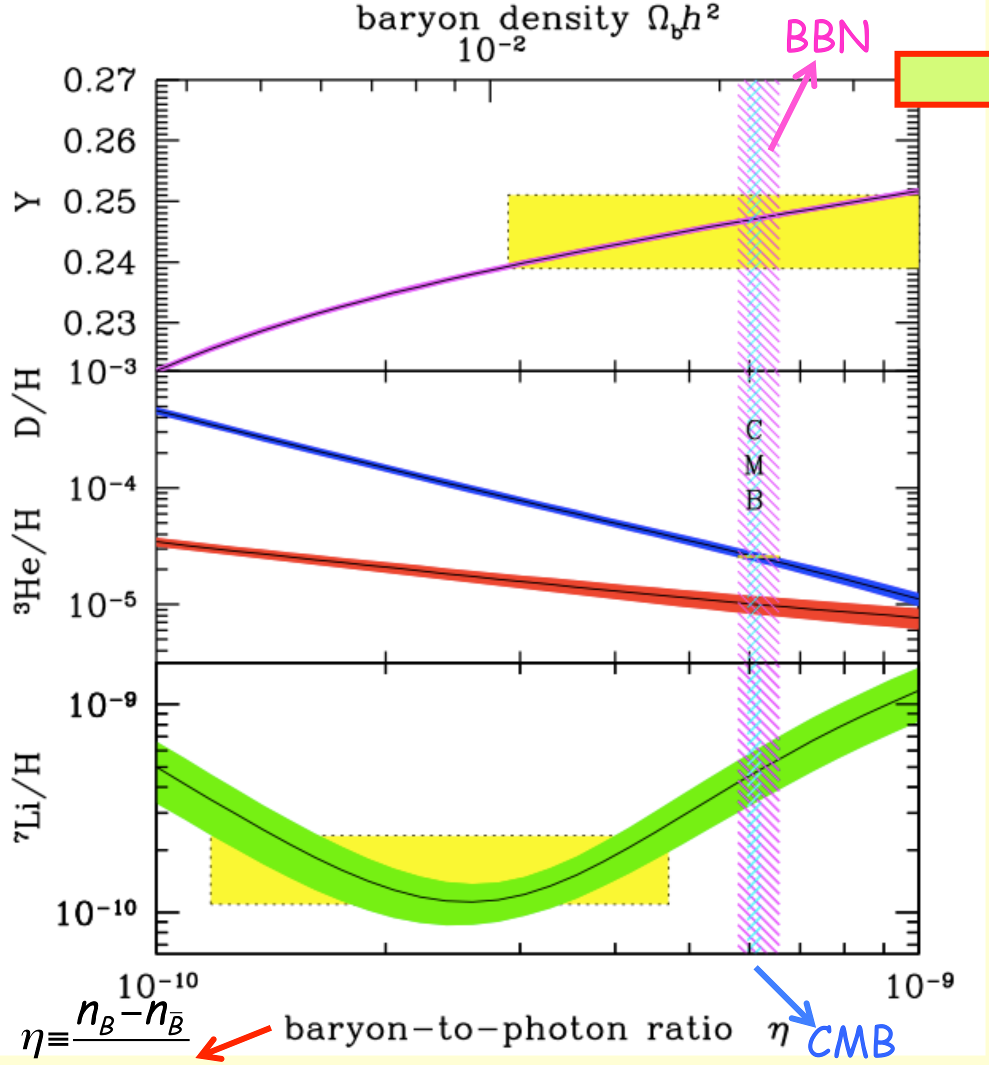


Most matter in the Universe is not ordinary atoms : dark matter !

M. Tanabashi et al. (PDG), Phys. Rev. D 98, 030001 (2018)

Present values of the Ω_i's

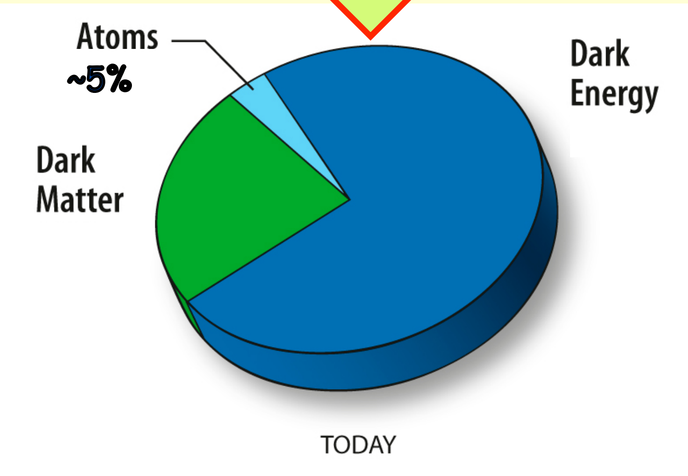
- Measurements of light element abundances vs predictions:



$$\eta \equiv \frac{n_B - n_{\bar{B}}}{n_\gamma}$$

$0.046 \le \Omega_b^0 \le 0.053$

(confirmed independently by CMB - Planck 2018)



atoms: gas (4%), stars (0.4%), ν (0.3%), heavy elements (0.03%)

M. Tanabashi et al. (PDG), Phys. Rev. D 98, 030001 (2018)

CONCLUSIONS (2)

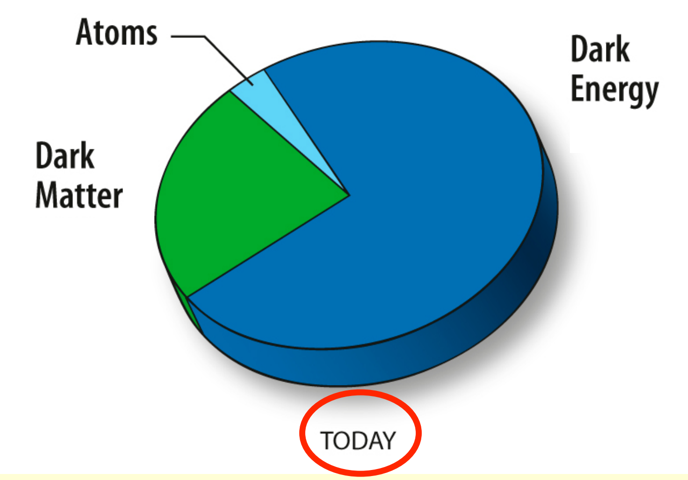
- Einstein GR equations + Λ + FLRW metric :

$$1 = \Omega_m(t) + \Omega_r(t) + \Omega_k(t) + \Omega_\Lambda(t)$$

- COBE measurement: radiation today is negligible
- Cosmological measurements vs predictions:
  - CMB+SNeIa+BAO:

$$\Omega_m^0 = 0.315 \pm 0.007$$
$$\Omega_\Lambda^0 = 0.685 \pm 0.007$$

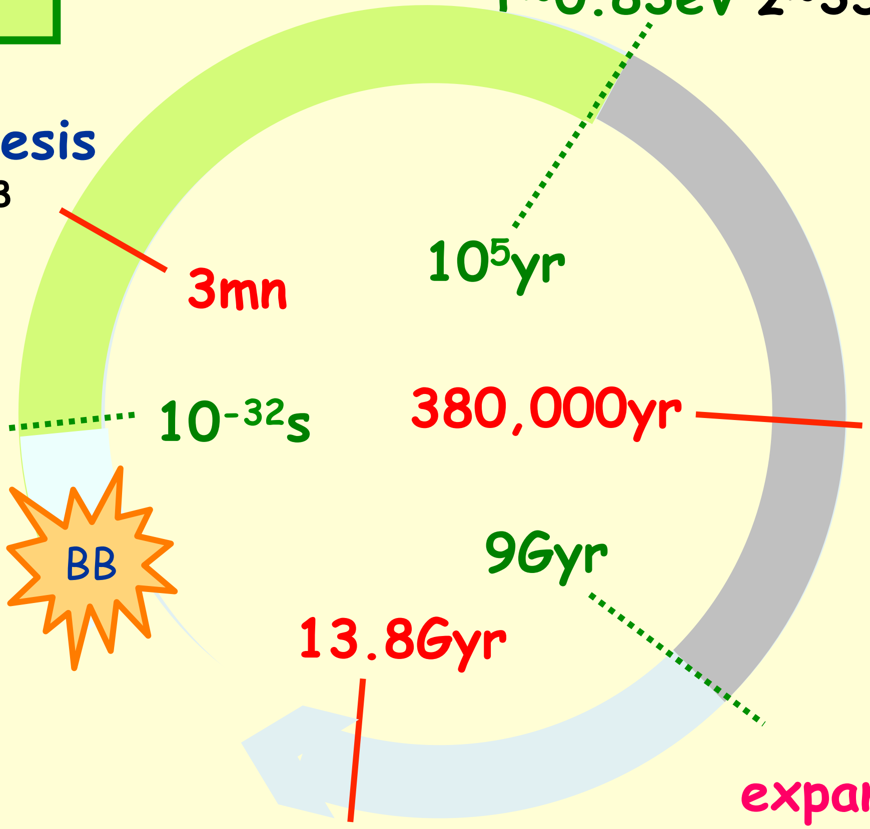
➔flat universe



  - Primordial abundances:

$$0.043 \leq \Omega_b^0 \leq 0.049$$

Concordance model : Λ_{CDM}



radiation
domination

matter-radiation equality
T~0.83eV z~3500

matter
domination

D, ^{4}He nucleosynthesis
T~60keV z~10^8

3mn

10^5yr

CMB
T~0.26eV z=1100

10^{-32}s

380,000yr

inflation ?

9Gyr

13.8Gyr

expansion acceleration
T~3.5 10^{-4}eV z~0.5

today
T~2.35 10^{-4}eV z=0

Λ
domination