Constraining the equation of state for dense matter through thermal emission of neutron stars

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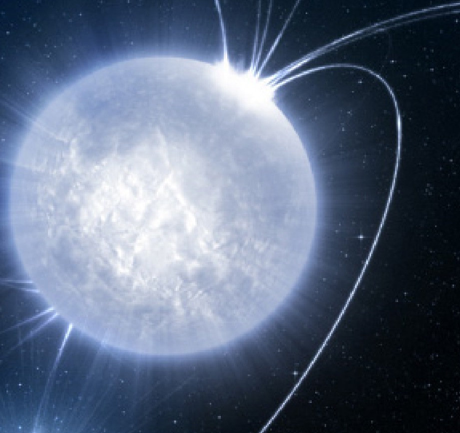
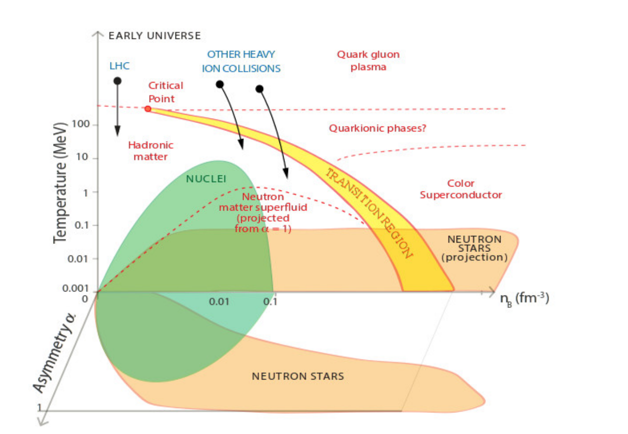


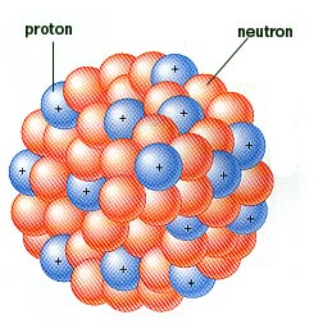
Neutron stars are best laboratory for very high densities :

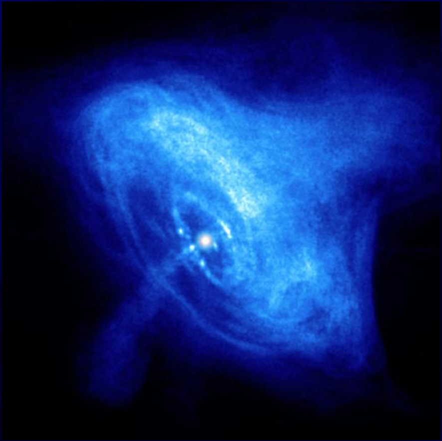
- Hadronic matter at very high density / isospin asymmetry
- **Phase transition ?**



Hadronic matter, anyway !

Comparision between two "isotopes" :





$\simeq 10^{-15}$ m $\longrightarrow$ $\simeq 10^{4}$ m

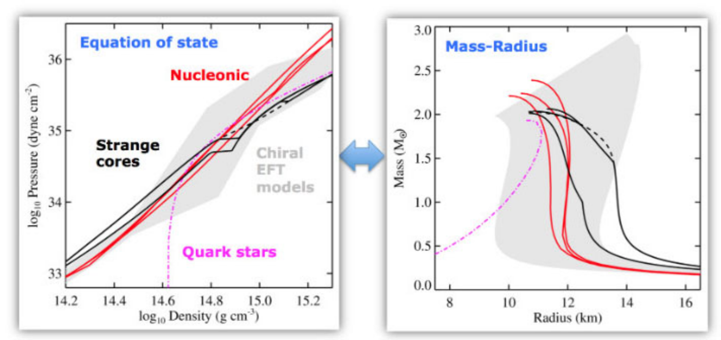
Link between EoS and M-R relation :

T.O.V equations (Tolmann-Oppenheimer-Volkov) :

$$\frac{dm(r)}{dr} = 4\pi r^2 \epsilon(r), \tag{1}$$

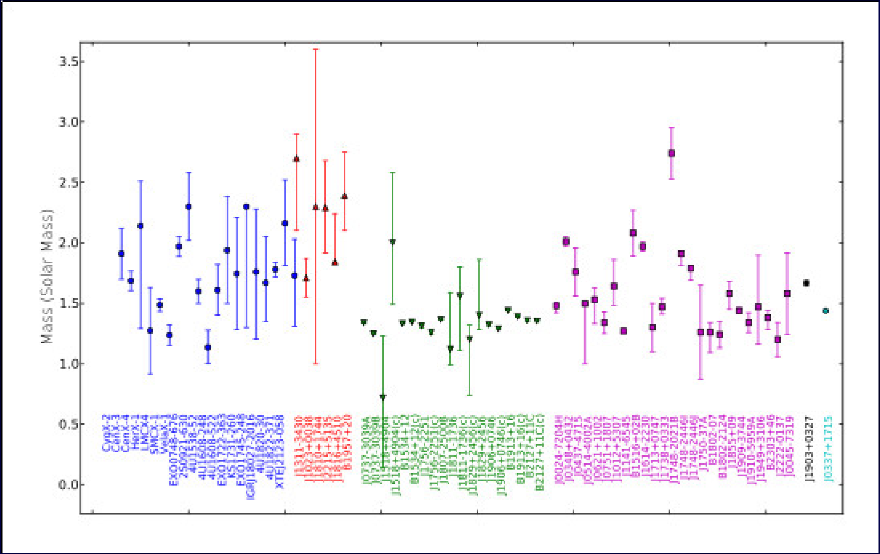
$$\frac{dP(r)}{dr} = -\frac{G\epsilon(r)m(r)}{r^2}\left(1+\frac{P(r)}{\epsilon(r)c^2}\right)\left(1+\frac{4\pi P(r)r^3}{m(r)c^2}\right)\times\left(1-\frac{r_{sh}}{r}\frac{m(r)}{M}\right)^{-1}. \tag{2}$$

Equation of state : $P(\epsilon) \leftrightarrow$ mass-radius relation.



Anna Watts et al, 2014

Masses distribution of NSs :



Radii distribution of NSs :

Recent works :

$R \in [9; 15]$ km for a 1.4 $M_\odot$ NS

Tools :

- Thermal emission of NSs (Guillot et al. 2013, Steiner 2018)
- X-ray bursts (Ozel et al. 2016, Poutanen and Suleimanov 2013)
- Present and future : Gravitational waves from NS-NS mergers (Tews et al. 2018)
- Future : NICER mission (Ozel et al. 2016, Gendrau et al. 2017)
- Future : Athena mission (Barcons et al. 2017)

Constraining nuclear equation of state from thermal radiation of neutron stars :

- Developpement of pulsar X-ray astronomy $\rightarrow$ strong limits on the equation of state at high density (Heinke et al. 2014, Özel et al 2016 , Steiner et al 2018).
- Tool : Thermal radiation from the surface of neutron stars.
- Developpement of atmosphere models (Heinke et al 2006).
- Applying Bayesian analysis for several qLMXBs (Guillot et al 2013, Lattimer & Steiner 2014, Heinke et al 2014, Özel et al 2015, Bogdanov 2016).

The equation of state (EoS) for neutron stars interior :

Various kinds of EoS :

- Purely nucleonic → no phase transition, **smooth EoS** with n, y_e, T.
- Phase transition :
  - hadronic matter : onset of hyperons, pion condensate ...
  - pure quark stars (absolutely stable strange matter) ?
  - hybrid stars (QGP/color superconductivity in the core) ?

Motivation for choosing **pure nucleonic matter** :

- **extrapolation** of **nuclear physics** knowledge (and uncertainties) towards high densities and low Y_p.
- Define M-R boundaries for smooth EoS to **Confront** with **observational data.**

The EoS model based on empirical parameters :

Definition of the empirical parameters :

$$\epsilon(n,\delta) = \epsilon_{IS} + \delta^2 \epsilon_{IV} \tag{3}$$

$$\epsilon_{IS} = E_{sat} + \frac{1}{2} K_{sat} x^2 + \frac{1}{3!} Q_{sat} x^3 + \frac{1}{4!} Z_{sat} x^4 + o(x^5) \tag{4}$$

$$\epsilon_{IV} = E_{sym} + L_{sym} x + \frac{1}{2} K_{sym} x^2 + \frac{1}{3!} Q_{sym} x^3 + \frac{1}{4!} Z_{sym} x^4 + o(x^5) \tag{5}$$

Taylor expansion around n_{sat} for the energy density :

$$\epsilon(n,\delta) = t(n,\delta) + \sum_{\alpha \geq 0}^{N} v_\alpha(\delta) \frac{x(n)^\alpha}{\alpha!} u_\alpha^N(x), \quad \delta = \frac{(n_n - n_p)}{n}, \quad x = \frac{(n - n_{sat})}{3 n_{sat}}$$

$v_\alpha(\delta) = v_{\alpha,IS} + v_{\alpha,IV}\delta^2$, and $u_\alpha^N(x)$ ensure the convergence.

(Margueron et al. PRC 2018)

Our present knowledge of the empirical parameters :

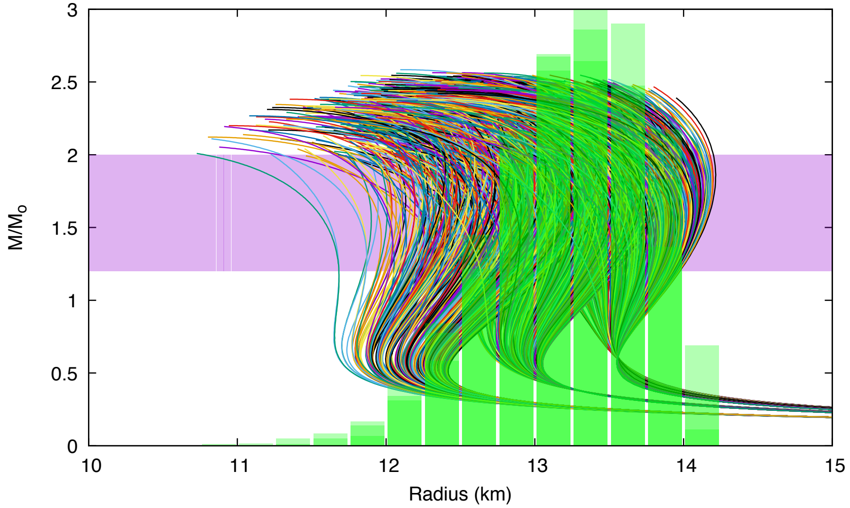
Empirical parameters for various effective approaches :

		fixed			Explore in small interval		Explore in large interval				
Model (N_α)	der. order	E_{sat} MeV 0	E_{sym} MeV 0	n_{sat} fm^{-3} 1	L_{sym} MeV 1	K_{sat} MeV 2	K_{sym} MeV 2	Q_{sat} MeV 3	Q_{sym} MeV 3	Z_{sat} MeV 4	Z_{sym} MeV 4
Skyrme	Average	-15.88	30.25	0.1595	47.8	234.2	-129.8	-357.0	377.9	1500.0	-2218.9
(16)	σ	0.15	1.70	0.0011	16.8	10.2	66.0	22.4	110.3	169.3	617.6
Skyrme	Average	-15.87	30.82	0.1596	49.6	237.3	-131.7	-349.0	370.0	1447.6	-2175.1
(35)	σ	0.18	1.54	0.0039	21.6	26.6	89.1	88.5	187.9	510.4	1069.4
RMF	Average	-16.24	35.11	0.1494	90.2	268.0	-4.6	-1.9	271.1	5058.3	-3671.8
(11)	σ	0.06	2.63	0.0025	29.6	33.5	87.7	392.5	357.1	2294.1	1582.3
RHF	Average	-15.97	33.97	0.1540	90.0	248.1	128.2	389.2	523.3	5269.1	-9955.5
(4)	σ	0.08	1.37	0.0035	11.1	11.6	51.1	350.4	236.8	838.4	4155.7
Total	Average	-16.03	33.30	0.1543	76.6	251.1	-2.7	12.7	388.1	3925.0	-5267.5
(50)	σ_{tot}	0.20	2.65	0.0054	29.2	28.6	131.7	431.1	289.4	2269.7	4281.9
	Min	-16.35	26.83	0.1450	9.9	201.0	-393.9	-748.0	-86.2	-903.0	-16916.2
	Max	-15.31	38.71	0.1746	122.7	355.4	213.2	949.8	846.3	9997.3	-4.9

(Margueron et al. PRC 2018)

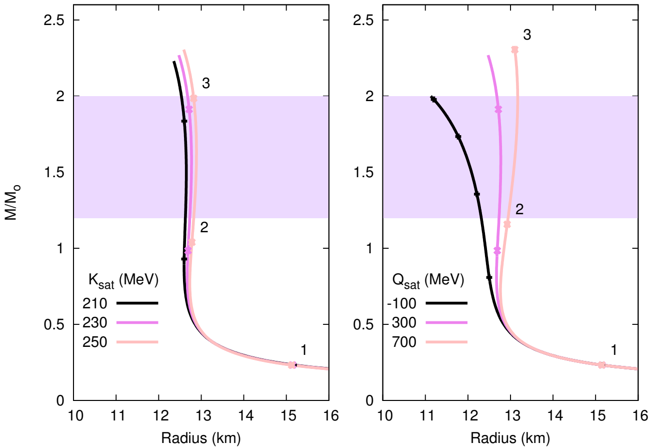
Dependence of M-R on the empirical parameters :

+ constrains : $0 < v_s^2 < c^2$ and $\epsilon_{IV}(n) > 0$ for $n < 4n_{sat}$



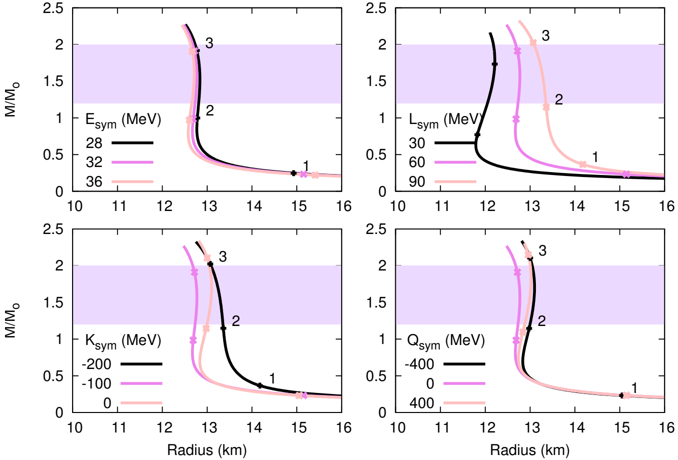
(Baillot d'Étivaux, Thesis, 2018)

Dependence of M-R on the empirical parameters : isoscalar chanel



(Margueron et al. PRC 2018)

Dependence of M-R on the empirical parameters : isovector chanel



(Margueron et al. PRC 2018)

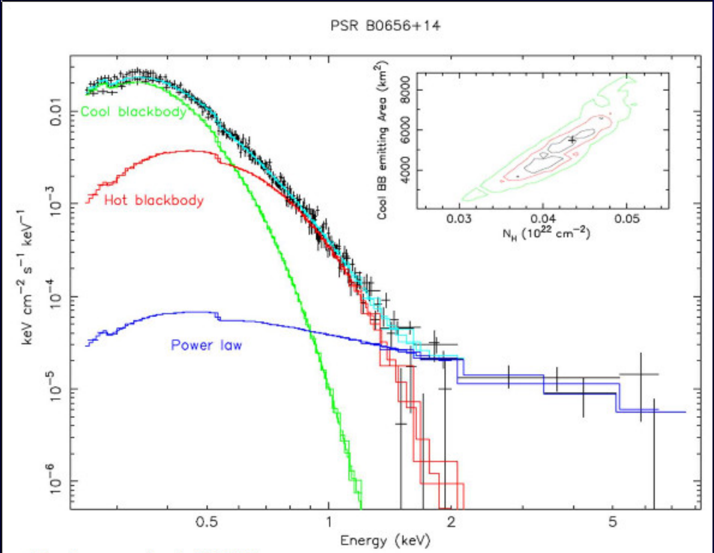
X-ray observations of neutron stars :

In collaboration with Natalie Webb and Sebastien Guillot

XMM-Newton

Chandra Observatory

Observation of hermal emission of neutron stars :



(De Luca et al. 2005)

Observations of thermal emission from NSs :

Black body like emission :

$$F \propto T^4 (R_\infty / D)^2 \ .$$

(Rutledge et al. 1999)

- **7 low mass X-ray transcients :**
  constant flux, purely H atmosphere,
  low magnetic fields $\longrightarrow$ almost pure thermal components,
  in globular clusters $\longrightarrow$ well constrained distances.
- **a single EoS $\rightarrow$ simultaneous analysis.**
- **Theory directly implemented in the data analysis**

Globular clusters with qLMXBs

Source	distance (recent pub) (kpc)	distance (GAIA DRII-2018) (kpc)
47-Tuc	4.53 ± 0.08	4.50 ± 0.06
ω-Cen	4.59 ± 0.08	5.20 ± 0.09
M13	7.1 ± 0.62	7.9 ± 0.62
M28	5.5 ± 0.3	5.50 ± 0.13
M30	8.2 ± 0.62	8.10 ± 0.12
NGC6397	2.51 ± 0.07	2.30 ± 0.05
NGC6304	6.22 ± 0.26	5.90 ± 0.14

Modeling the spectra with Xspec

Spectrum model used :

- "pile-up" (Davis 2001, Bogdanov 2016), "TBgas" absorbtion and "nsatmos" for the atmosphere (Heinke et al. 2006).
- $E < 2$ keV : only thermal component and no power-law

Parameters which are allowed to vary :

- pile-up α parameter
- hydrogen collumn density on the line of site $n_{H,22}$ (10^{22} cm^{-2})
- distance to the stars D (kpc), surface effective temperature T_{eff} (K), the mass of the stars M ($M_{\odot}$)
- the nuclear empirical parameters $\{P_\alpha\}$: EoS($\{P_\alpha\}, M$) $\rightarrow R$

$\approx$ 40 free parameters for $\approx$ 1000 degrees of freedom.

$\longrightarrow$ MCMC methods

Prior distribution for the distances

Multiplying the likelihood by gaussian priors :

$$\frac{1}{N}\exp\left(-\frac{\chi^2}{2}\right) \longrightarrow \frac{1}{N}\exp\left(-\frac{\chi^2}{2}\right) \times \prod_i P^i(d^i),$$

with reference distances d^i_{ref}, and reference uncertainties $\sigma^i_{d,ref}$:

$$P^i(d^i) = \exp\left(-\frac{1}{2} \times \frac{d^i - d^i_{ref}}{\sigma^i_{d,ref}}\right)^2 .$$

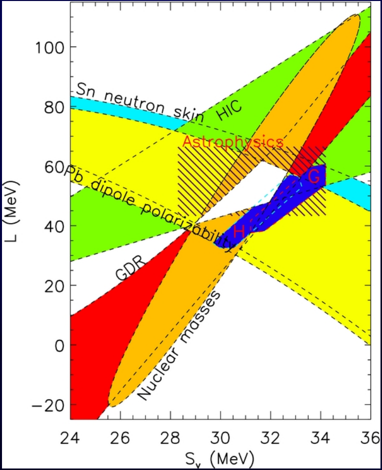
Constraining L_{sym} from nuclear knowledge

$L_{sym} = 50 \pm 10$ MeV :

$$P(L_{sym}) = \exp\left(\frac{L_{sym} - 50}{2 \times 10}\right)^2$$

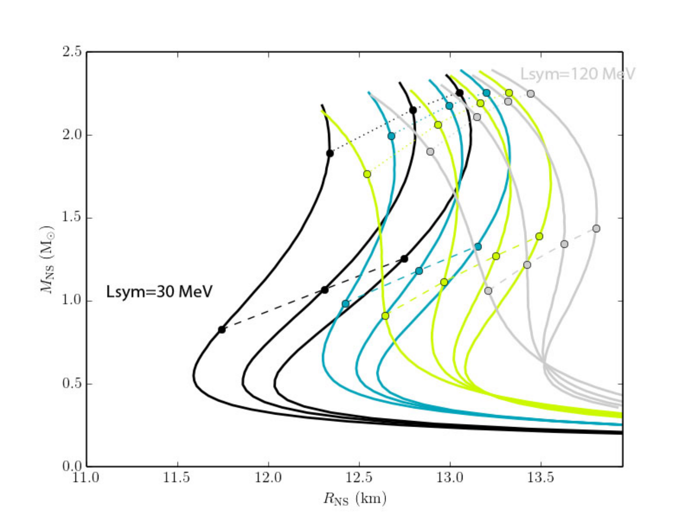
$K_{sym} \in [-400 : 200]$ MeV

$Q_{sat} \in [-1300 : 1900]$ MeV



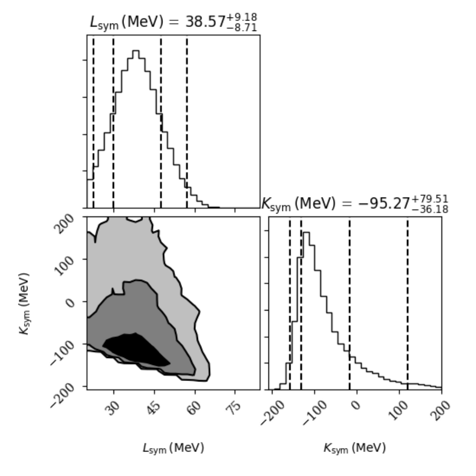
(Lattimer and Lim 2013)

M-R exploration varying L_{sym} and K_{sym}



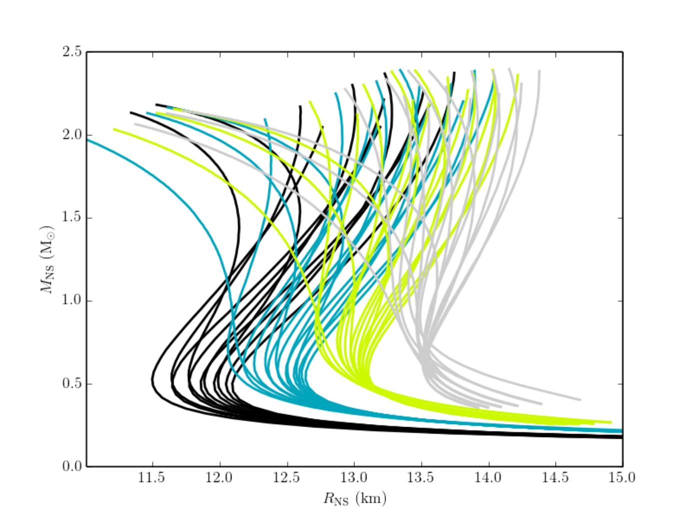
(Baillot d'Étivaux et al. (in preparation))

Constraining L_{sym} and K_{sym}



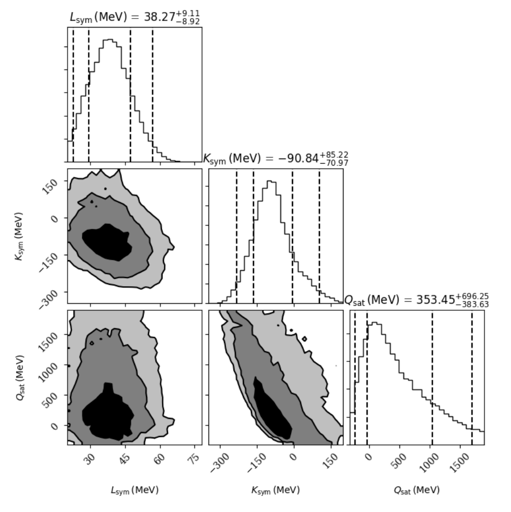
(Baillot d'Étivaux et al. (in preparation))

M-R exploration varying L_{sym}, K_{sym} and Q_{sat}



(Baillot d'Étivaux et al. (in preparation))

Constraining L_{sym}, K_{sym} and Q_{sat}



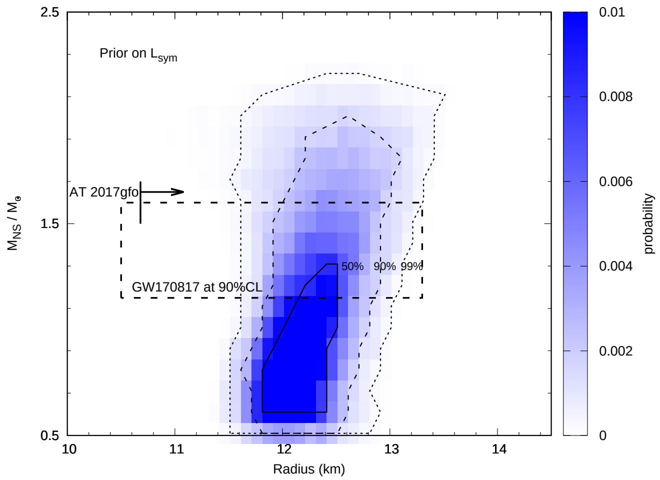
(Baillot d'Étivaux et al. (in preparation))

Tests of sensibility

Sources	Distances	prior L_{sym}	L_{sym} MeV	K_{sym} MeV	Q_{sat} MeV
all	Gaia DR2	yes	$37.2^{+9.2}_{-8.9}$	-85^{+82}_{-70}	318^{+673}_{-366}
all	qLMXB publ.	yes	$38.3^{+9.1}_{-8.9}$	-91^{+85}_{-71}	353^{+696}_{-484}
all	qLMXB publ.	yes	$38.6^{+9.2}_{-8.7}$	-95^{+80}_{-36}	-
all	qLMXB publ.	no	$27.9^{+11.4}_{-5.8}$	-60^{+107}_{-79}	498^{+738}_{-479}
all	qLMXB publ. 3σ	yes	$40.2^{+9.5}_{-9.2}$	-86^{+111}_{-85}	555^{+740}_{-494}
High S2N	Gaia DR2	yes	$37.5^{+9.0}_{-8.9}$	-88^{+76}_{-70}	263^{+764}_{-361}
Low S2N	Gaia DR2	yes	$50.3^{+9.8}_{-9.6}$	-1^{+134}_{-143}	881^{+671}_{-705}
all/47-Tuc	Gaia DR2	yes	$43.4^{+9.7}_{-9.3}$	-66^{+137}_{-102}	622^{+763}_{-560}
all/NGC6397	Gaia DR2	yes	$42.6^{+9.9}_{-9.5}$	-77^{+129}_{-96}	623^{+757}_{-544}
all/M28	Gaia DR2	yes	$42.5^{+9.5}_{-9.5}$	-80^{+124}_{-91}	597^{+717}_{-510}

(Baillot d'Étivaux et al. (in preparation))

Constraining L_{sym}, K_{sym} and Q_{sat} : M-R distribution

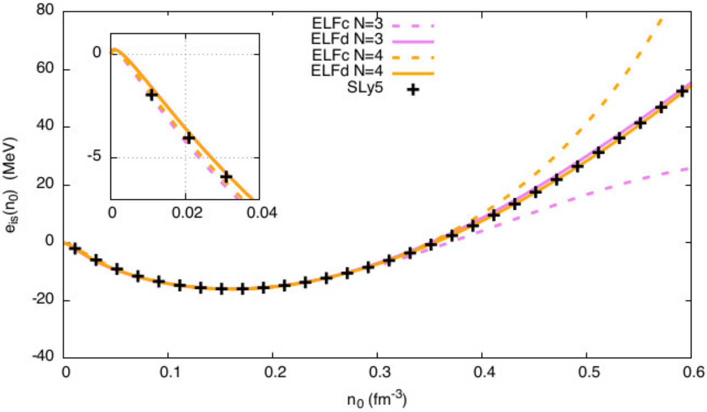


Conclusions

- Link between the EoS and the M-R relation.
- Nuclear EoSs can hardly produce radii below 11.5 km.
- Nucleonic models are able to reproduce the data.
- The parameters which have the most impact on the EoS are : L_{sym}, K_{sym}, and Q_{sat}.
- NSs observation can bring new constraints on these parameters.

Thank you !

Ability of the EoS to mimic known EoS : example of SLy5



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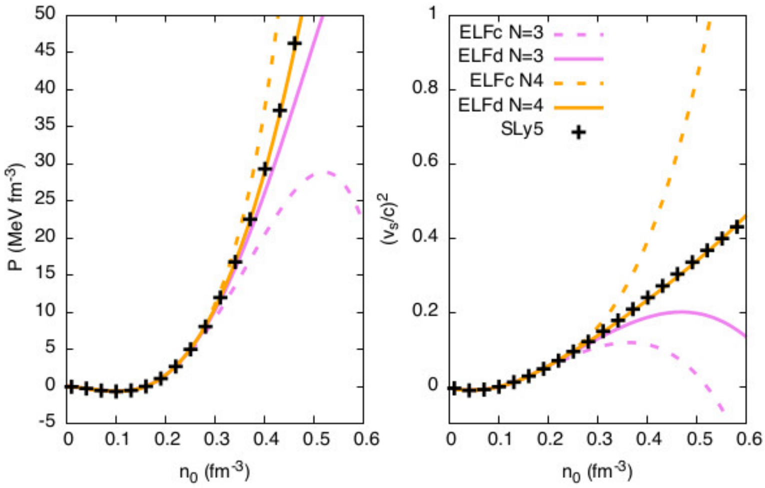


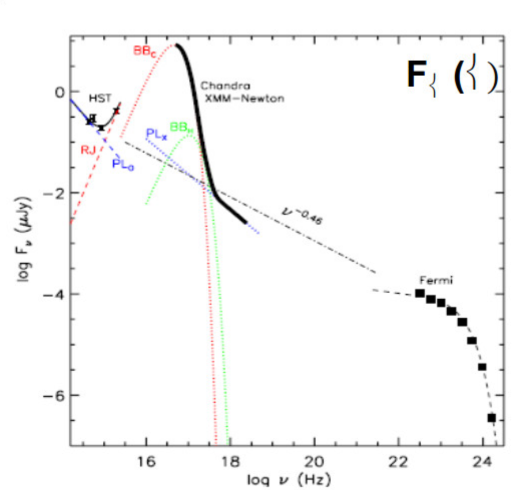
Parameters for isoscalar channel :

$$
\begin{aligned}
E_0 &= e_b(x=0,\delta=0), && (6)\\
P_0 &= n_{sat}^2(1+3x)^2 \frac{\partial}{\partial x} e_b(x,\delta)\Big|_{x=0,\delta=0} = 0, && (7)\\
K_0 &= \frac{\partial^2}{\partial x^2} e_b(x,\delta)\Big|_{x=0,\delta=0}, && (8)\\
Q_0 &= \frac{\partial^3}{\partial x^3} e_b(x,\delta)\Big|_{x=0,\delta=0}, && (9)\\
Z_0 &= \frac{\partial^4}{\partial x^4} e_b(x,\delta)\Big|_{x=0,\delta=0}. && (10)
\end{aligned}
$$

Parameters for isovector channel :

$$\begin{aligned}
S_{sym} &= e_{s,iv}(x=0), && (11)\\
L_{sym} &= \frac{\partial}{\partial x} e_{s,iv}(x)\Big|_{x=0,\delta=0}, && (12)\\
K_{sym} &= \frac{\partial^2}{\partial x^2} e_{s,iv}(x)\Big|_{x=0,\delta=0}, && (13)\\
Q_{sym} &= \frac{\partial^3}{\partial x^3} e_{s,iv}(x)\Big|_{x=0,\delta=0}, && (14)\\
Z_{sym} &= \frac{\partial^4}{\partial x^4} e_{s,iv}(x)\Big|_{x=0,\delta=0}. && (15)
\end{aligned}$$

Spectrum 1 :



Mignani, Pavlov, Kargaltsev 2011)