

cherenkov telescope array





#### **Real-Time Analysis Discussion: On-Site Data Deduction**

Lenka Tomankova (RUB)

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#### **Target Data Levels: R0 → DL0**

- Required reduction factor ~100
- Data reduction process must be **reversible**
- DL0 data families
  - Event (EVT) } 80–90%

Calibration (CAL)
Technical (TECH)
10–20% (acquired at lower frequencies)

# **Types of Data Reduction**

- 1. Waveform reduction/integration
- 2. Pixel selection (zero-suppression, RoI/VoI selection)
- 3. Suppression of background events
- 4. Data compression (lossless)

#### 1. Waveform Reduction (1)

- Most cameras record signal as function of time in each pixel
- Reduce waveform to a few parameters (amplitude, time, width, ...) used in high-level analysis
- Requires real-time knowledge of:
  - Pedestal values & timing calibration
  - Reconstructed image orientation & time gradient estimate
- Implications for calib.

LST:  $125 \rightarrow 15$  Bytes

- Understanding of use of waveform shape in reconstruction and gamma/hadron separation
- Reduction factor per pixel ~10 8.3 for LST

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### 1. Waveform Reduction (2)

- ACTL criterion for planning purposes:
  - Complete waveform for 3% (on average) of pixels
  - Time-integrated waveform for all others

	Relative data rate due to	
	3% complete waveforms (125 Bytes/pixel)	97% integrated waveforms (15 Bytes per pixel)
LST	21%	<b>79%</b>

- **Reduction factor** for whole camera read-out: **6.8**
- Majority of data volume comes from integrated waveforms

#### 2. Pixel Selection (1)

- Large FOV → only a small fraction of pixels contain useful information
- Discard pixels without useful information early → save processing & storage
- Requires pre-calibration & preliminary image reconstruction

### 2. Pixel Selection (2)

#### <u>Approach</u>

- 1. Apply image cleaning
  - Test different algorithms (tail-cuts, iterative tail-cuts (MAGIC), wavelets, SVD, ...)
- 2. Store values from
  - all selected pixels,
  - additional **boundary region** of thickness **N**, and
  - any **pixels above X pe**
  - · Determine optimal N and X from simulations (LST optimization)
  - Additional pixels necessary if cleaning thresholds change later on due to re-calibration

#### 2. Pixel Selection (3)

Performance study

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- Low-level parameters: width, length, multiplicity, ...
- High-level parameters: energy bias and resolution, angular resolution, sensitivities, ...
- Using simulations with different levels of NSB

## 3. Background Suppression

- Vast majority of events are initiated by protons or helium nuclei
   → can be discarded at low energies without loss of science
- Required pre-calibration & preliminary event reconstruction
   confidence level on the background hypothesis
- **Very-high-amplitude** always of interest (spectra of heavy cosmicray nuclei) → simple threshold criterion (~400 pe) for storage
- Weak  $\gamma$ /hadron separation power for threshold events (most numerous)
- Factor ~2 reduction

#### 4. Data Compression

- Factor ~2 without loss of sensitivity
- Final compression factor depends on the previous level of data reduction

#### **Strategy for LST-1 Commissioning**

- **Performance** of reduction will be studied on **commissioning data**, in which **all raw data will (be attempted to) be stored**.
- After performance verification, data reduction will be applied on-site (on-the-fly).
- For additional checks, every N<sup>th</sup> (N~100) event will be kept in full form.