Deep learning approach to predict photometric redshifts of galaxies in the Sloan Digital Sky Survey DR12

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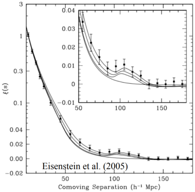
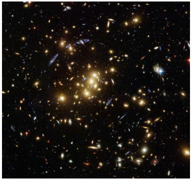
Need accurate redshits for cosmology

Reliable redshifts are necessary to constrain the dark energy equation-of-state and to study the large scale structure of the universe

- Baryonic Acoustic Oscillations



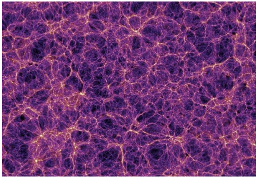
- Weak lensing



Strong gravitational lensing around galaxy cluster CL0024+17

Credit : NASA/ESA/M.J. Jee (John Hopkins University)

- Cosmic web

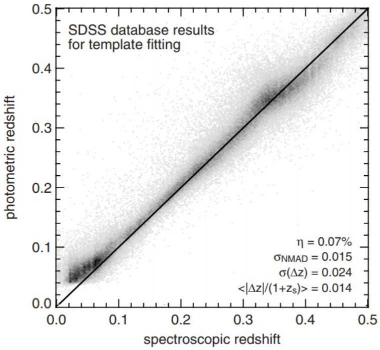


Results of a digital simulation showing the large-scale distribution of matter, with filaments and knots.

Credit: V.Springel, Max-Planck Institut für Astrophysik, Garching bei München

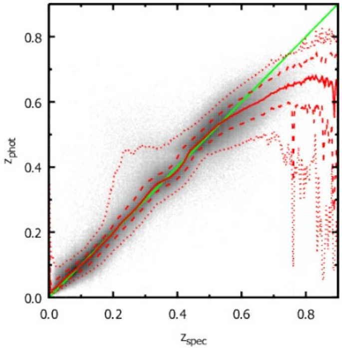
Existing methods

A template fitting method



From Greisel et al. 2013

A machine learning method (KNN)



From Beck et al. 2016

Preliminary results with Deep Learning methods (Hoyle 2016, D'Isanto 2018)

Photometric redshifts with Deep Learning

Photometric redshifts from SDSS images using a Convolutional Neural Network (**J. Pasquet**, E. Bertin, M. Treyer, S. Arnouts and D. Fouchez, just submitted)

Key elements :

1. A representative and a complete training database with r-band magnitude $\leq$ 17.8 and redshift, $z \leq 0.4$ (516,525 galaxies)
2. Photoz values + associated Probability Distribution Functions
3. Photoz immune to IQ variations and neighbours contamination
4. A dedicated Neural Network architecture

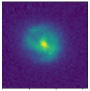
Results obtained :

Clear improvements compared to other methods!

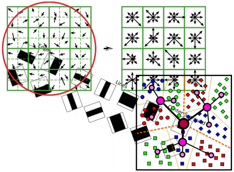
Why use Deep Learning methods?

Classical methods

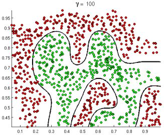
Input data



Feature crafting

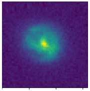


Separation with a classifier

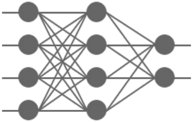


Deep learning

Input data

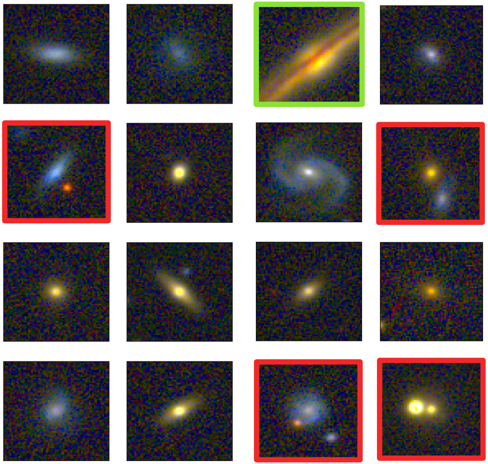


Feature learning



→ The best feature space representation is found by the network

Input SDSS galaxy images transmitted to the CNN

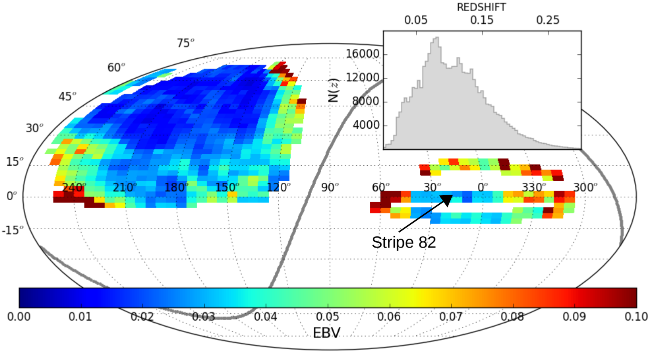


– large galaxies

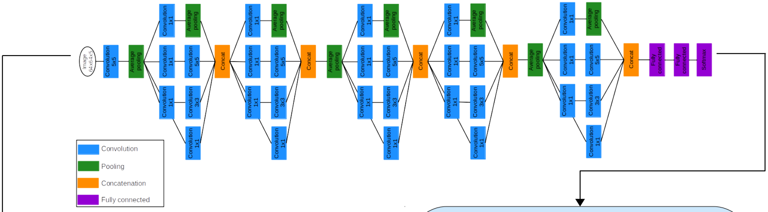
— crowded images

Main Galaxy Sample SDSS

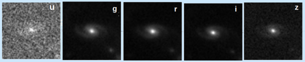
A multi-band imaging and spectroscopic redshift survey



Our architecture

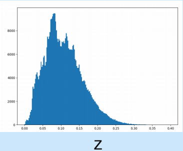


Input images of size 64x64



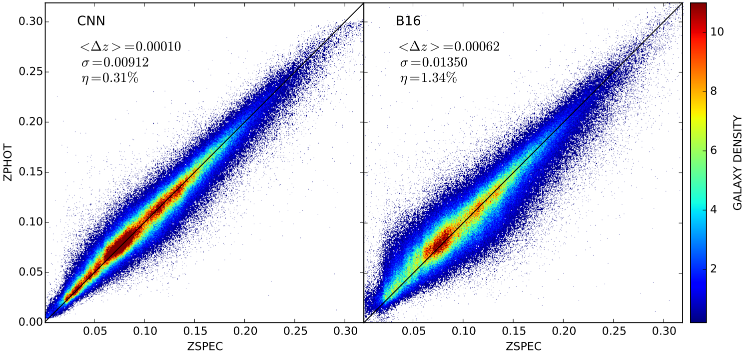
Output probabilities

Vector of posterior probabilities for the galaxy redshift to be in a z-bin



The estimator for z-phot is the centroid of the zPDF

Results of the method



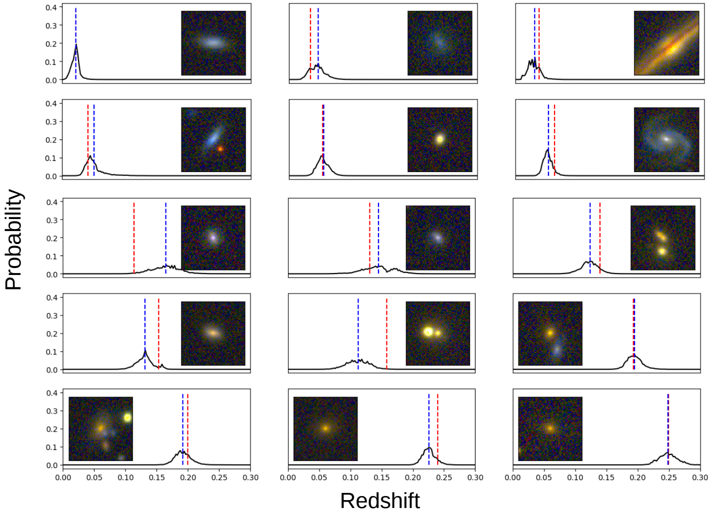
$\langle\Delta z\rangle=1.0\times10^{-4}$ ← Factor of 6 improvement $\langle\Delta z\rangle=6\times10^{-4}$

$\sigma=9.1\times10^{-3}$ ← 30 % improvement $\sigma=1.3\times10^{-2}$

$\eta=0.31\,\%$ ← Factor of 4 improvement $\eta=1.35\,\%$

$$\Delta z=(z_{phot}-z_{spec})/(1+z_{spec})$$
$$\sigma=1.4826\times MAD$$
$$MAD=Median(|\Delta z-Median(\Delta z)|)$$
$$\eta=|\Delta z|>0.05$$

Examples of PDFs



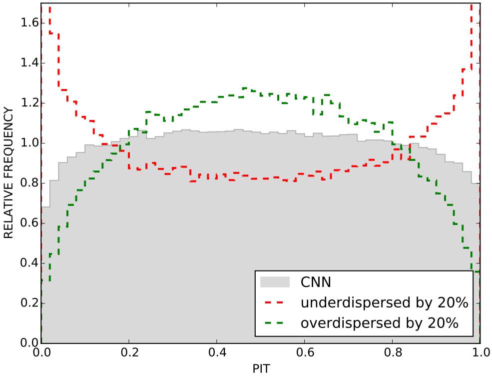
-- Spectroscopic redshift

-- Photometric redshift

Assess the prediction quality of our PDFs

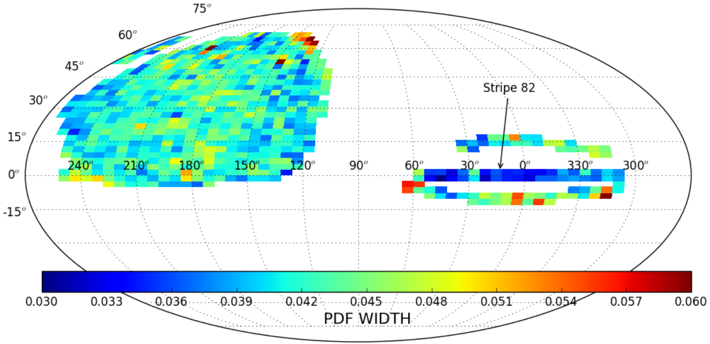
The PIT statistic (Dawid 1984) is based on the histogram of the cumulative probabilities at the true value. For galaxy i with spectroscopic redshift z_i in the test sample :

$$\mathrm{PIT}_i = \int_{-\infty}^{z_i} PDF_i(z)dz$$



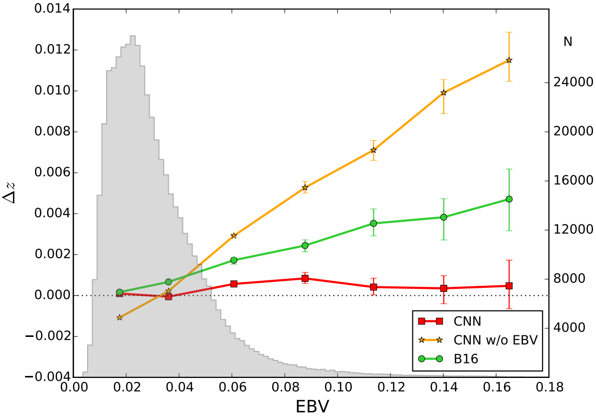
Impact of Signal-to-Noise Ratio (SNR) on widths of PDFs

The Stripe 82 region, which combines repeated observations of the same part of the sky, gives us the opportunity to look into the impact of SNR



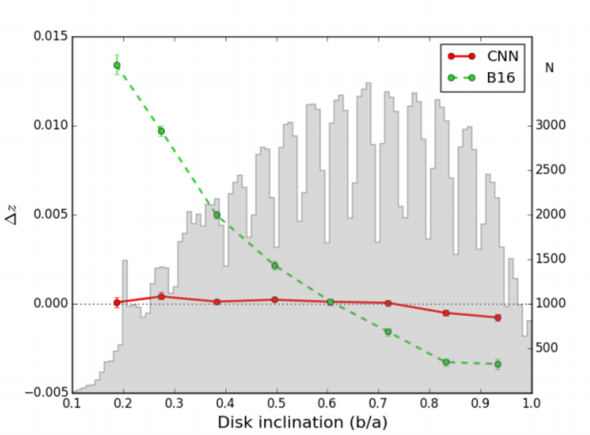
Impact of the extinction of our Galaxy on photometric redshifts

Our method tends to overestimate redshifts in obscured regions (confusing galactic dust attenuation with redshift dimming), unless $E_{(B-V)}$ is used for training



Impact of the disk inclination of galaxies on photometric redshifts

Our method automatically corrects for galactic dust reddening which increases with disk inclination



Summary results

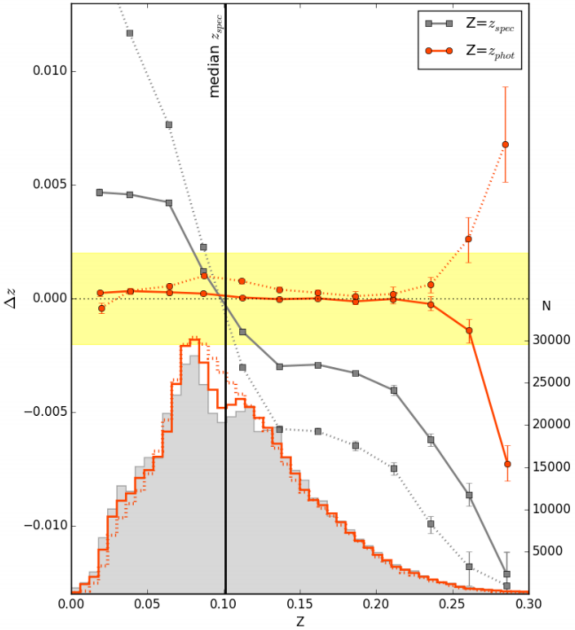
Trial	training sample size	bias	σ	η
Training with 80% of the dataset	393,219			
Full test sample		**0.00010**	**0.00912**	**0.31**
(B16)		(0.00062)	(0.01350)	(1.34)
Widest 20% of PDFs		**0.00005**	**0.00789**	**0.06**
Stripe 82 only		-0.00009	0.00727	0.34
Stripe 82 with widest 20% of PDFs removed		**0.00004**	**0.00635**	**0.09**
Training with 50% of the dataset⋆	250,000	0.00007	0.00910	0.29
Training with 20% of the dataset	**99,001**	**-0.00001**	**0.00914**	0.30
Training with 2% of the dataset	10,100	-0.00017	0.01433	1.26
Training and testing on Stripe 82	**15,771**	-0.00002	**0.00795**	0.38

Summary

- We developed a **Deep Convolutional Neural Network** (CNN) used as a classifier to estimate photometric redshifts and their associated PDFs.
- Our work shows **significant improvements** for:
  - the **dispersion** of photometric redshifts, σ_{MAD}
  - the **PDFs** that are **well calibrated**
  - **no measurable bias** with the reddening and the inclination of galaxies
- A high SNR tends to improve the results
- This work opens very **promising perspectives** for the exploitation of large and deep photometric surveys which encompass a larger redshift range and where spectroscopic follow-up is necessarily limited

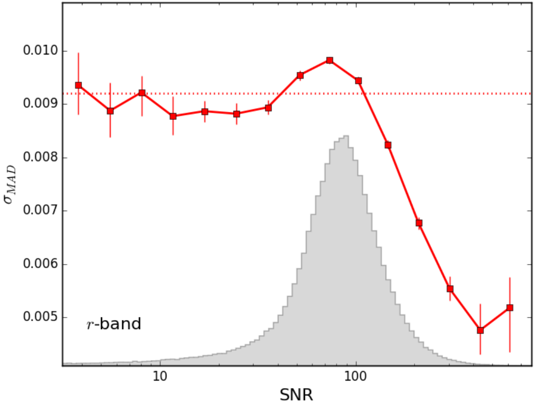
Thank you!

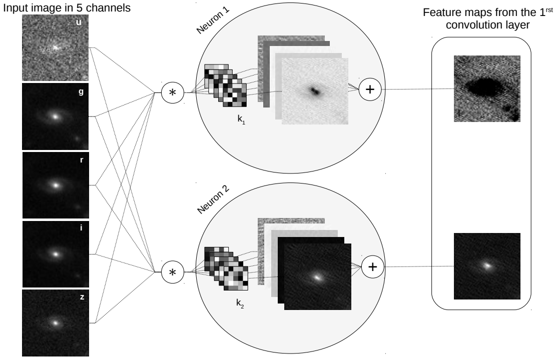
Supported by the OCEVU Labex (ANR-11-LABX-0060) funded by the "Investissements d'Avenir" French government program


median z_{spec}
$Z=z_{spec}$
$Z=z_{phot}$
0.010
0.005
0.000
−0.005
−0.010
Δz
N
30000
25000
20000
15000
10000
5000
0.00
0.05
0.10
0.15
0.20
0.25
0.30
Z

Impact of the SNR on the performance

σ_{MAD} decreases with the signal-to-noise ratio (SNR), achieving values below 0.007 for SNR > 100, as in the deep stacked region of Stripe 82




Input image in 5 channels
u
g
r
i
z
Neuron 1
Neuron 2
k1
k2
Feature maps from the 1rst convolution layer

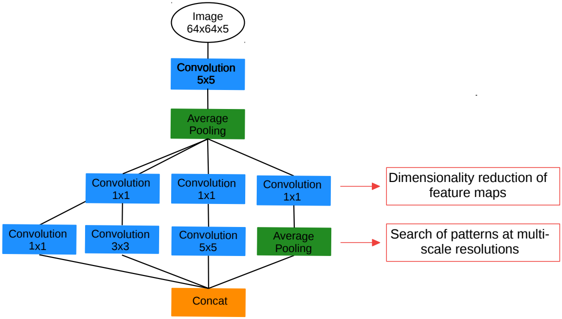
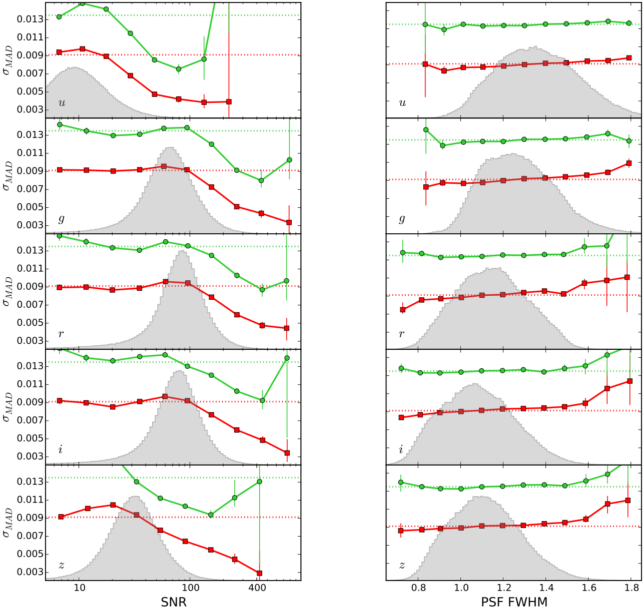


Image
64x64x5
Convolution
5x5
Average
Pooling
Convolution
1x1
Convolution
1x1
Convolution
1x1
Convolution
1x1
Convolution
3x3
Convolution
5x5
Average
Pooling
Concat
Dimensionality reduction of feature maps
Search of patterns at multi-scale resolutions

Trial	training sample size	size of 1 test sample	bias	σ_{MAD}	η	$< CRPS$
Training with 80% of the dataset	393,219					
Full test sample		103,306	0.00010	0.00912	0.31	0.00674
(B16)		(103,306)	(0.00062)	(0.01350)	(1.34)	
Suspect zone (SZ) removed		101,499	0.00004	0.00908	0.31	0.00672
Widest 10% of PDFs		91,543	0.00006	0.00848	0.09	0.00606
Widest 20% of PDFs		79,897	0.00005	0.00789	0.06	0.00556
Stripe 82 only		3,943	-0.00009	0.00727	0.34	0.00574
Stripe 82 with widest 20% of PDFs removed		3,131	0.00004	0.00635	0.09	0.00467
Training with 50% of the dataset*	250,000	252,500	0.00007	0.00910	0.29	0.00672
Training with 20% of the dataset	99,001	385,970	-0.00001	0.00914	0.30	0.00677
Training with 2% of the dataset	10,100	434,228	-0.00017	0.01433	1.26	0.01009
Training on Stripe 82	15,771					
Stripe 82 removed*		478,274	0.00194	0.01341	1.15	0.00988
Stripe 82 only		3,942	-0.00002	0.00795	0.38	0.00622
Training w/o Stripe 82	486,560					
Stripe 82 removed*		97,607	0.00000	0.00914	0.33	0.00680
Stripe 82 only*		19,714	-0.00077	0.00760	0.41	0.00606


0.013
0.011
0.009
0.007
0.005
0.003
σ_{MAD}
u
g
r
i
z
10
100
400
SNR
0.8
1.0
1.2
1.4
1.6
1.8
PSF FWHM