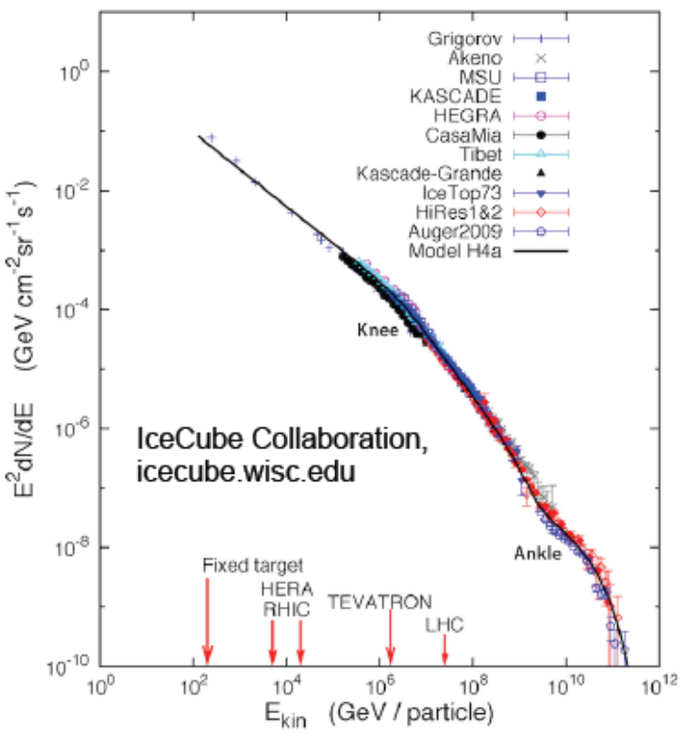
Toward Neutrino Astronomy with IceCube+ANTARES+KM3NeT

DAMIEN DORNIC (CPPM)

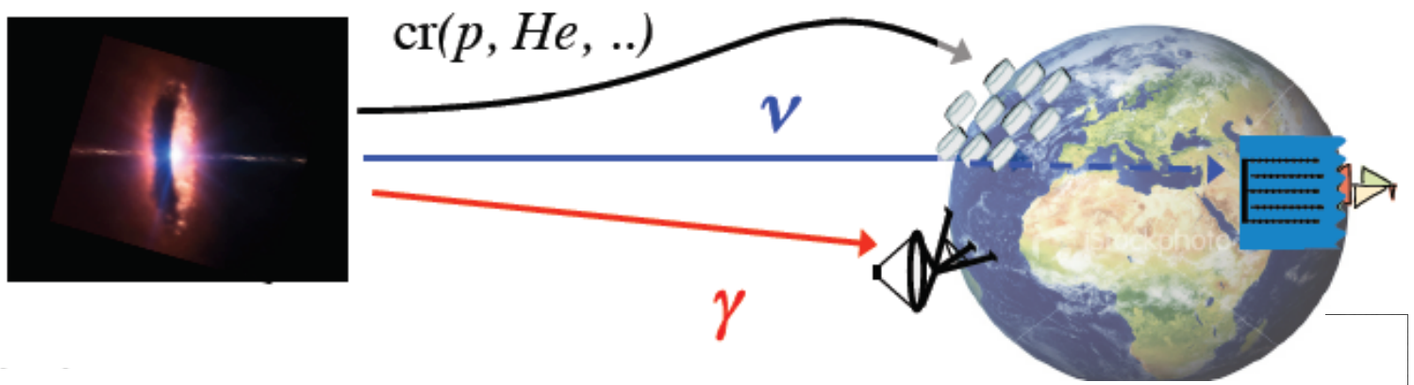
3rd SVOM science workshop

Les Houches - May 13-18 2018

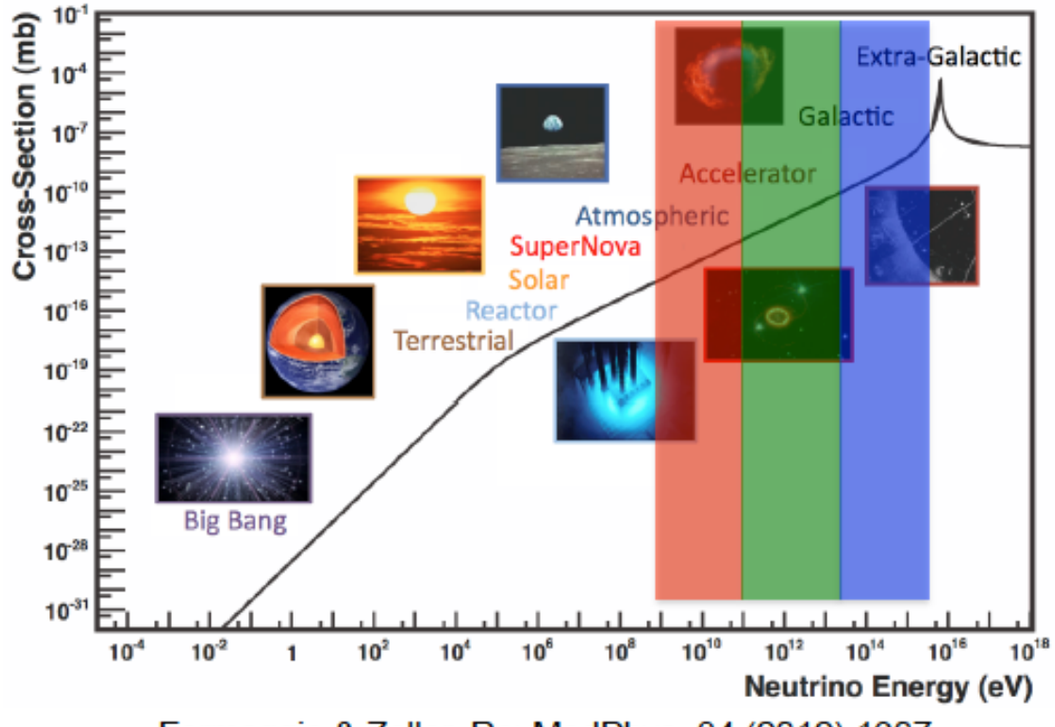
NEUTRINO AS COSMIC MESSENGER



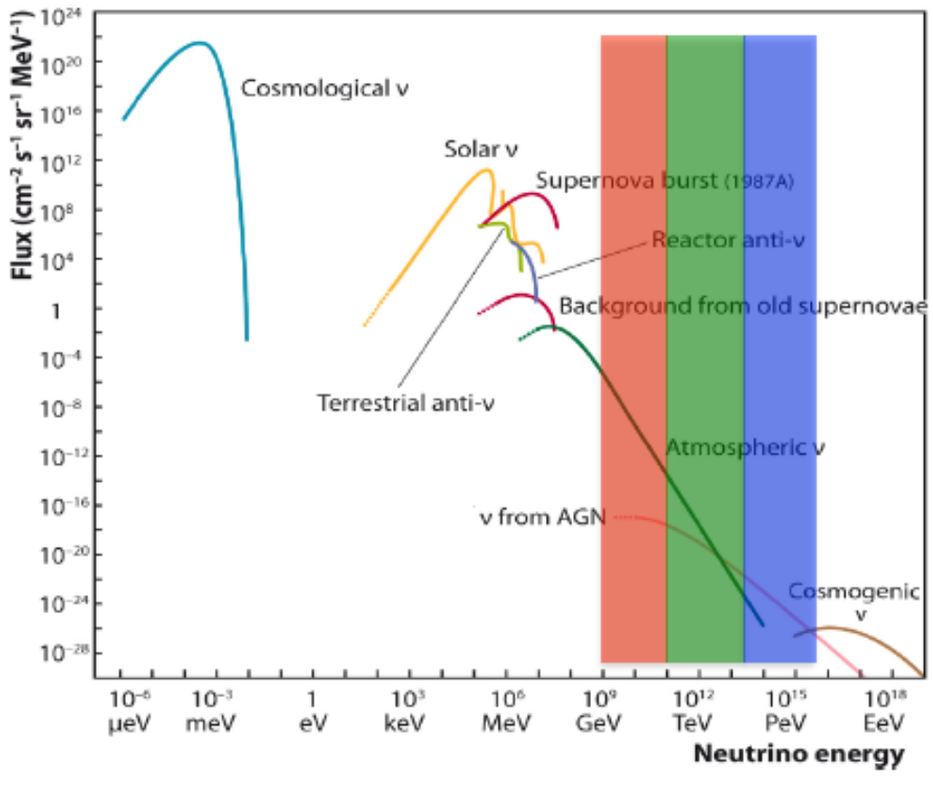
Neutrinos: smoking gun for cosmic-ray interactions



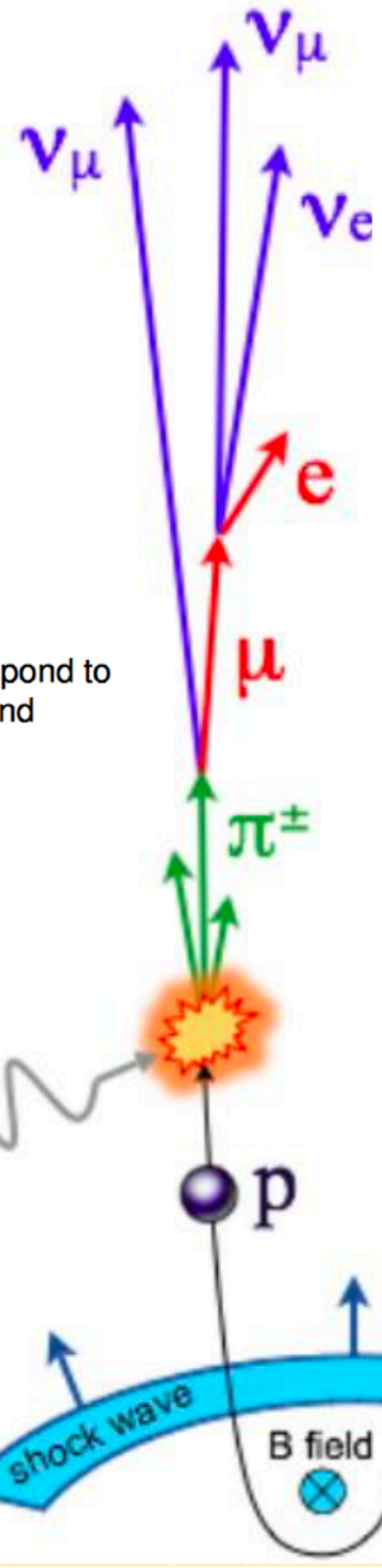
- 3 GeV – 1 TeV: atmospheric neutrinos, dark matter… ORCA
- 100 GeV - 30 TeV: various galactic (TeV gamma) sources ANTARES
- 30 TeV – 3 PeV: IceCube signal (astrophysical flux) ARCA



Formaggio & Zeller, RevModPhys 84 (2012) 1307



HE NEUTRINO PRODUCTION

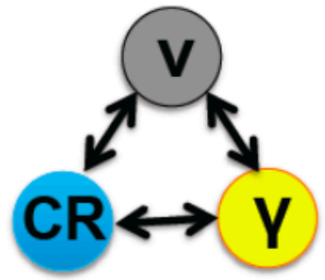


1 PeV neutrinos correspond to
20 PeV CR nucleons and
2 PeV γ-rays

The sources of HE ν are not necessary the sources of UHECR

Hadronuclear (e.g. star burst galaxies and galaxy clusters)

$$pp \rightarrow \begin{cases} \pi^0 \rightarrow \gamma\,\gamma \\ \pi^+ \rightarrow \mu^+ \nu_\mu \rightarrow e^+ \nu_e \nu_\mu \bar{\nu}_\mu \\ \pi^- \rightarrow \mu^- \bar{\nu}_\mu \rightarrow e^- \bar{\nu}_e \bar{\nu}_\mu \nu_\mu \end{cases}$$



Photohadronic (e.g. gamma-ray bursts, active galactic nuclei)

$$p\gamma \rightarrow \Delta^+ \rightarrow \begin{cases} p\,\pi^0 \rightarrow p\,\gamma\,\gamma \\ n\,\pi^+ \rightarrow n\,\mu^+ \nu_\mu \rightarrow n\,e^+ \nu_e \bar{\nu}_\mu \nu_\mu \end{cases}$$

cosmic ray + gamma

cosmic ray + neutrinos

Neutrino flavour ratio at source:

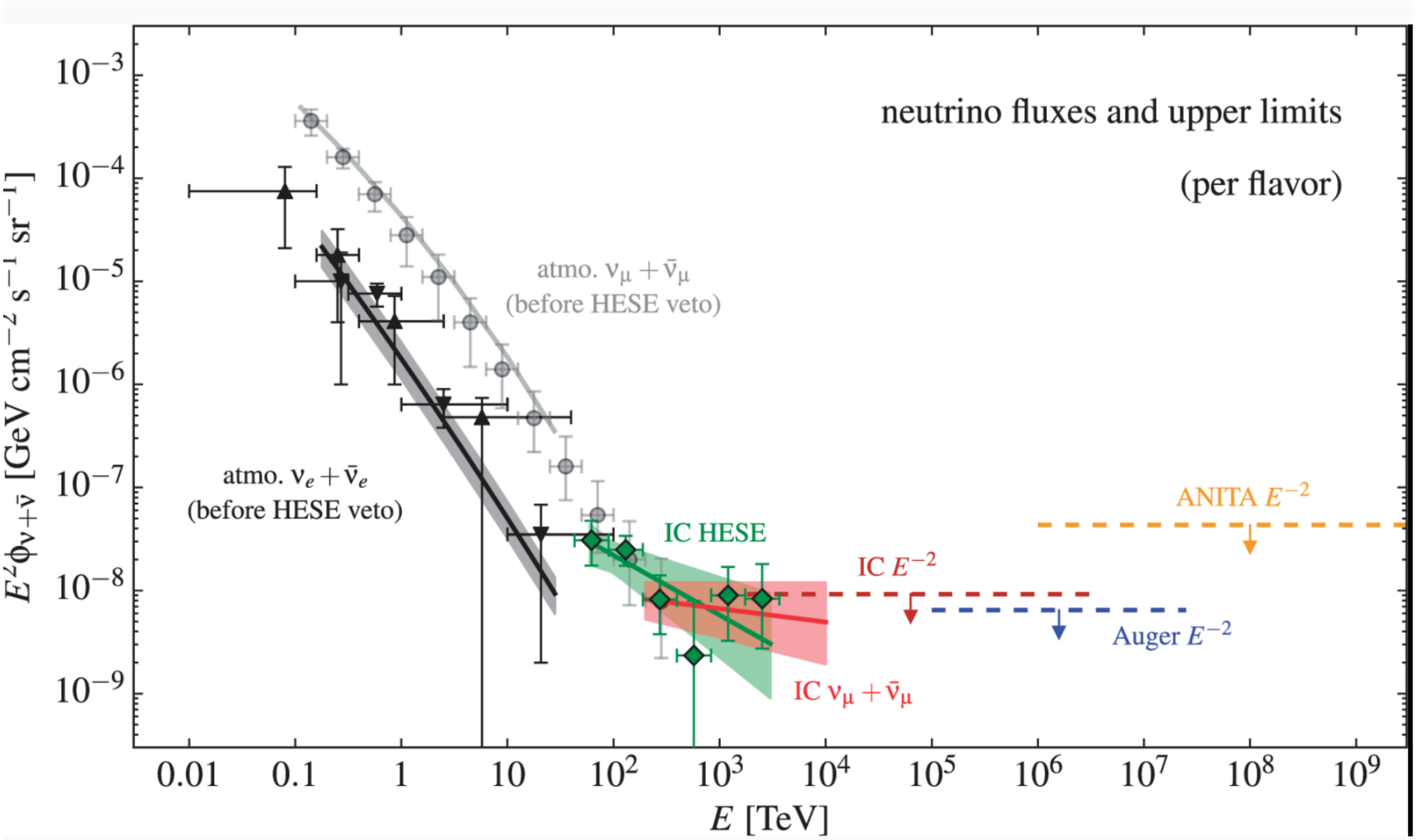
pion-muon decay

$$\nu_e : \nu_\mu : \nu_\tau \sim 1 : 2 : 0$$

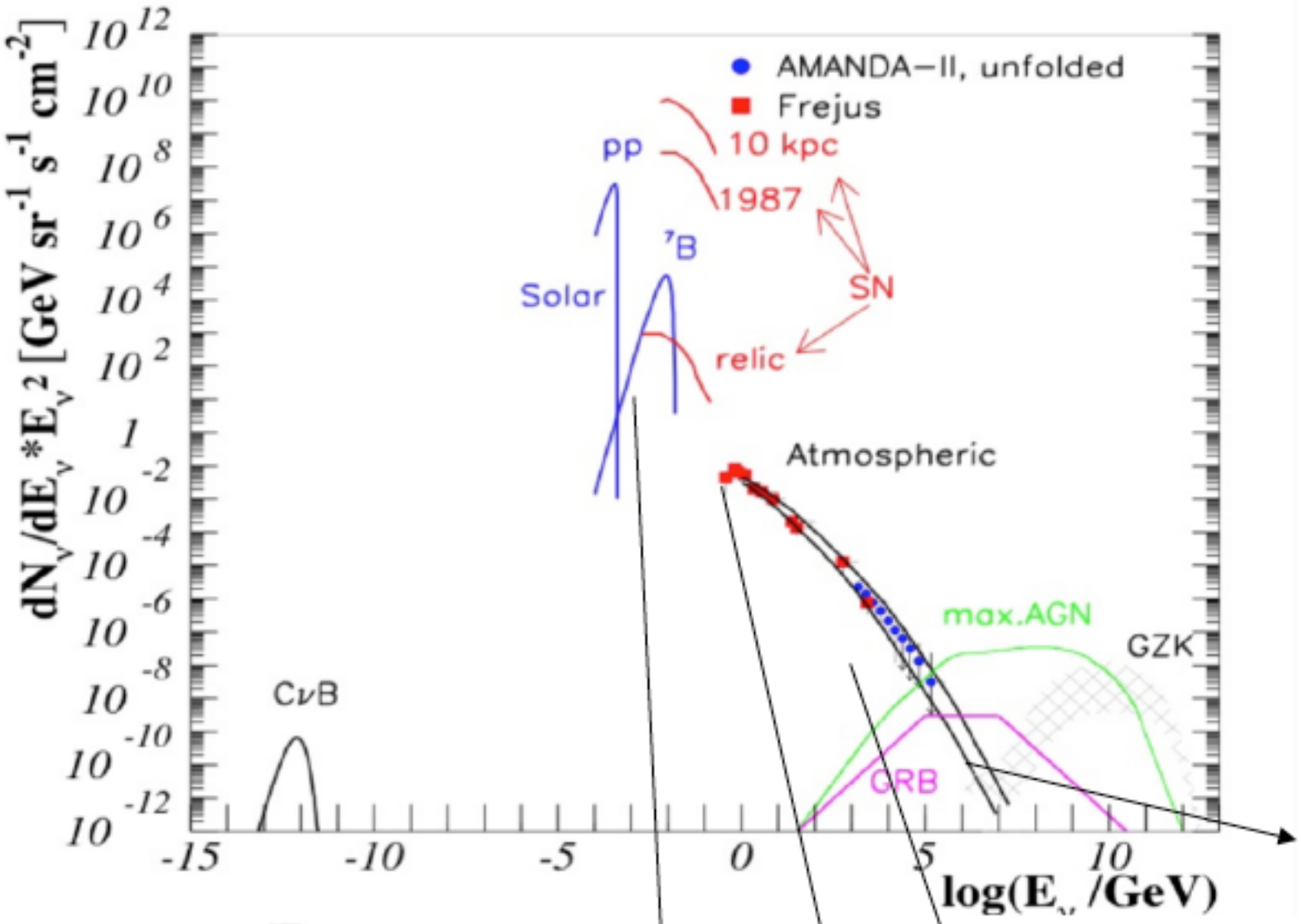
Oscillations average out over cosmic baselines

$$\nu_e : \nu_\mu : \nu_\tau \sim 1 : 1 : 1$$

HE NEUTRINO FLUXES

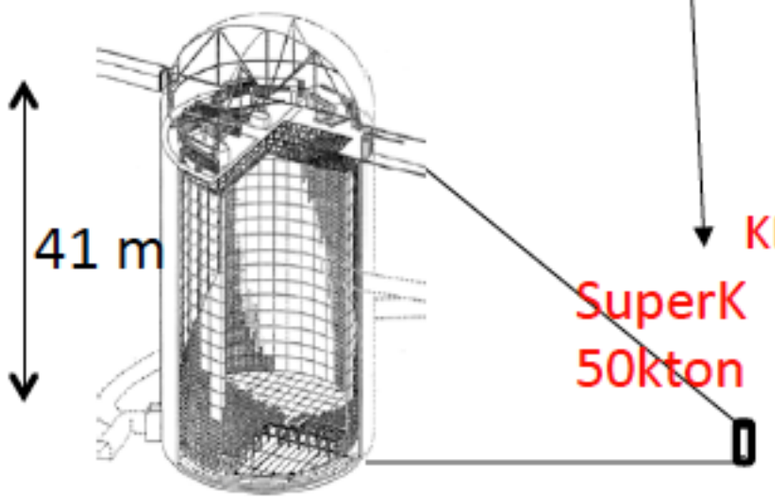


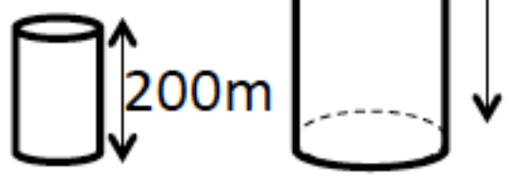
HE NEUTRINO DETECTORS

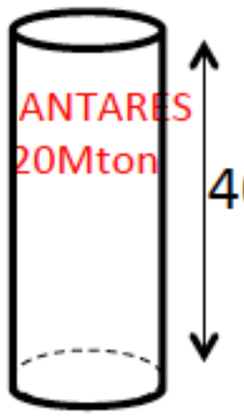


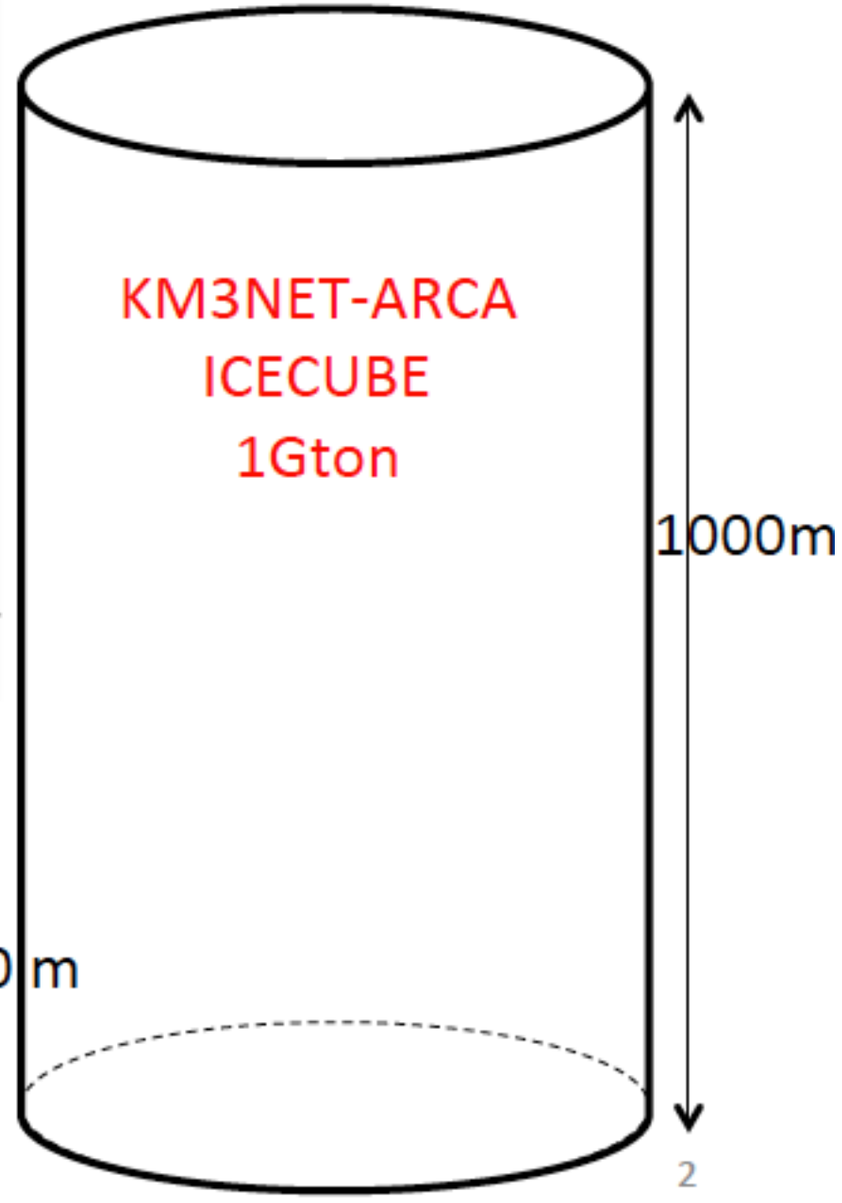
$\sigma(\nu p)/\sigma(\gamma p) = 10^{-7}$ at 1 TeV

Need very large detectors

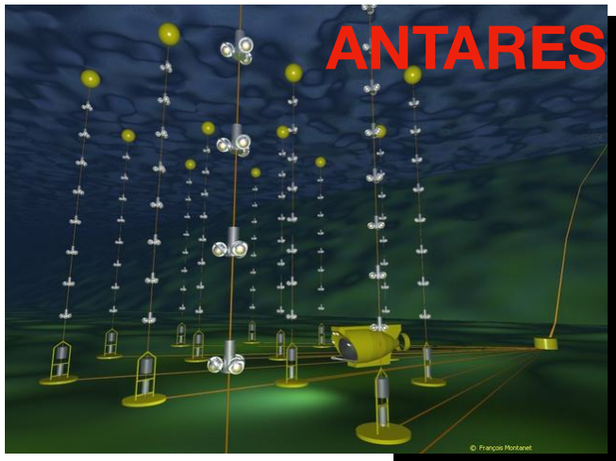


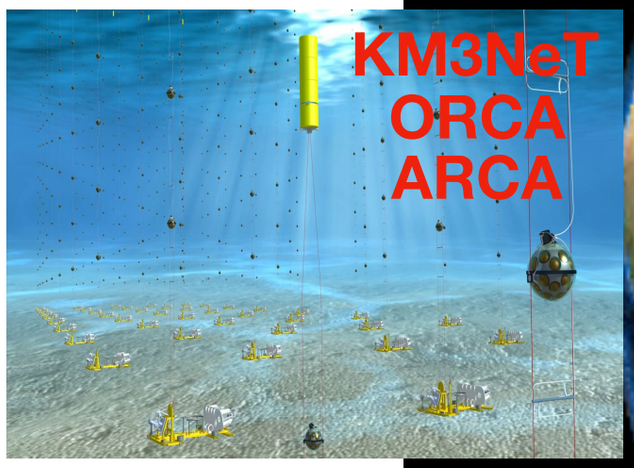


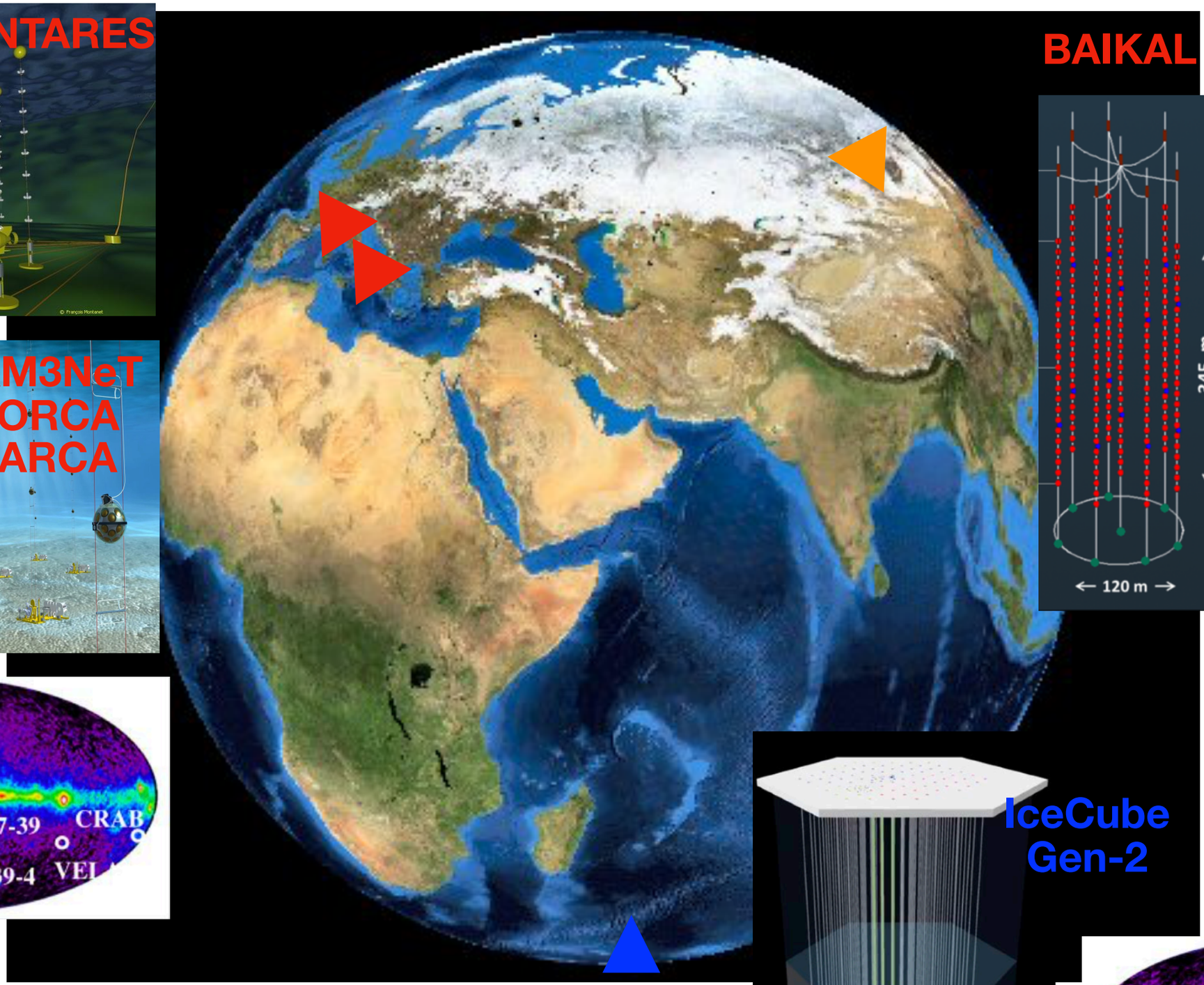


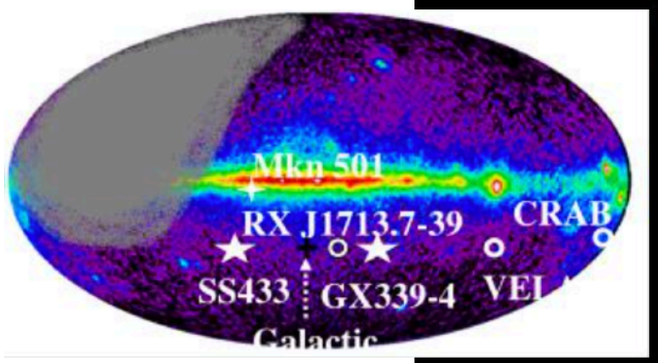


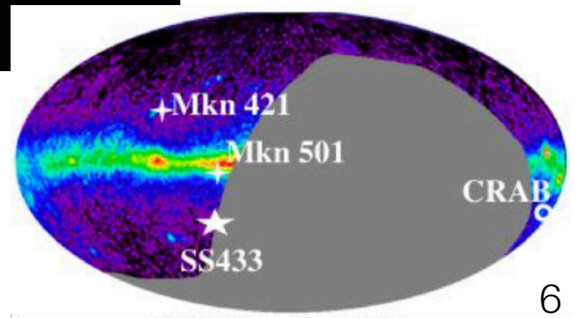
HE NEUTRINO DETECTORS





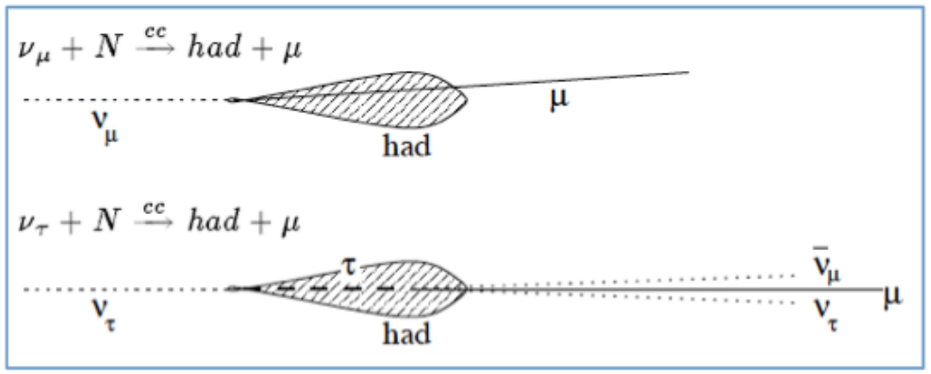




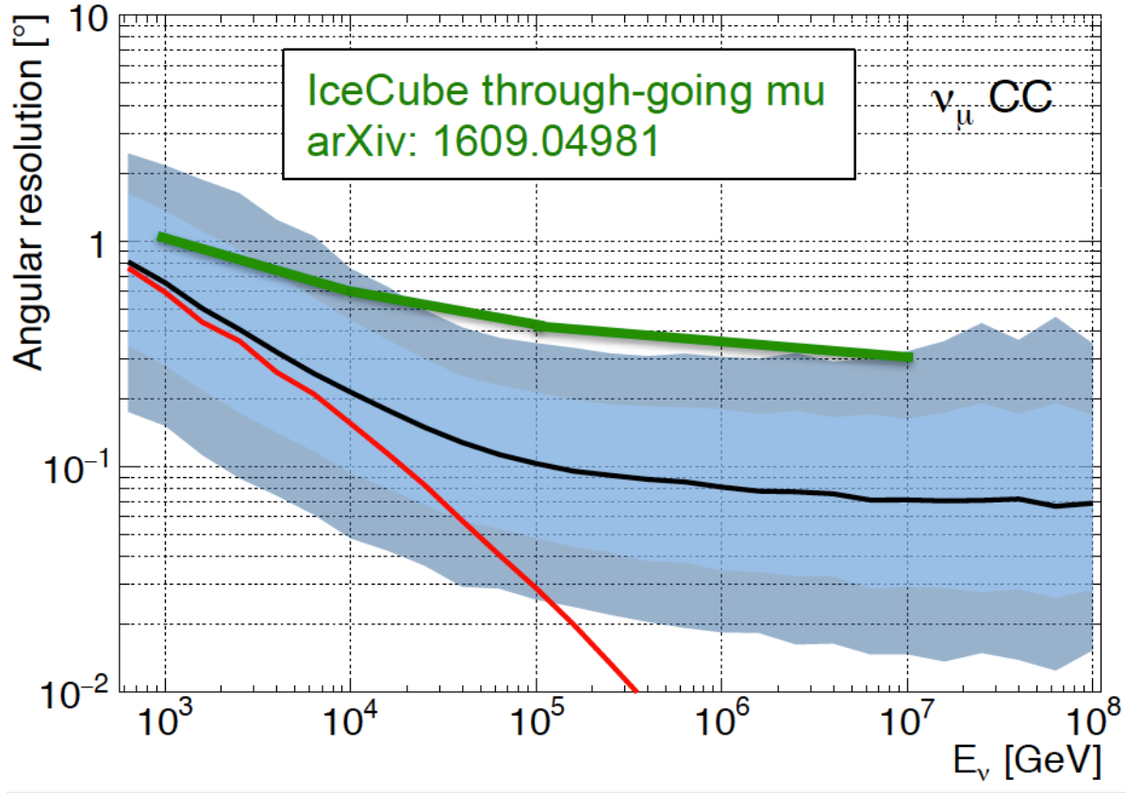


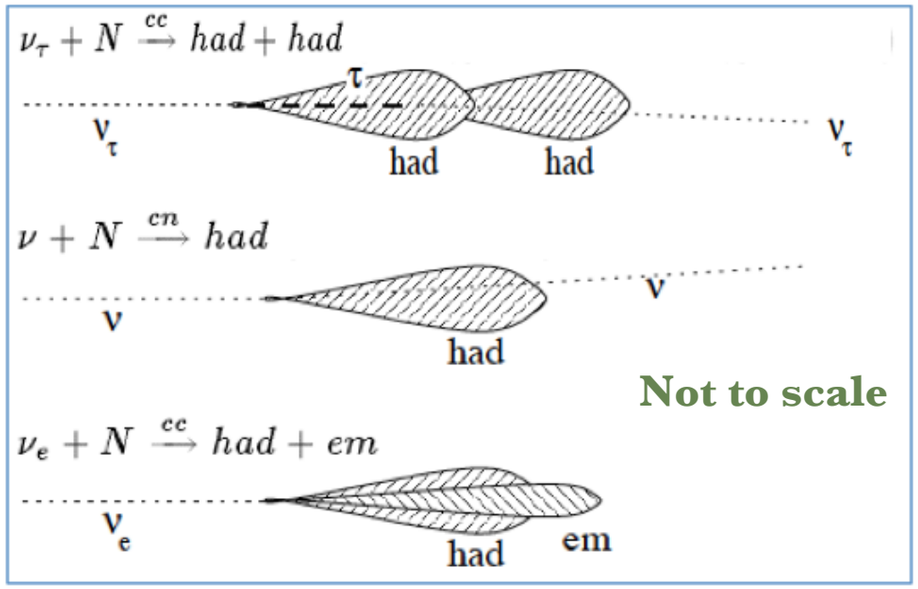
GNN: Global Neutrino Network linked all HE neutrino telescopes + provide framework for regular combined meetings and combined analysis

ALL-FLAVOR NEUTRINO TOPOLOGIES

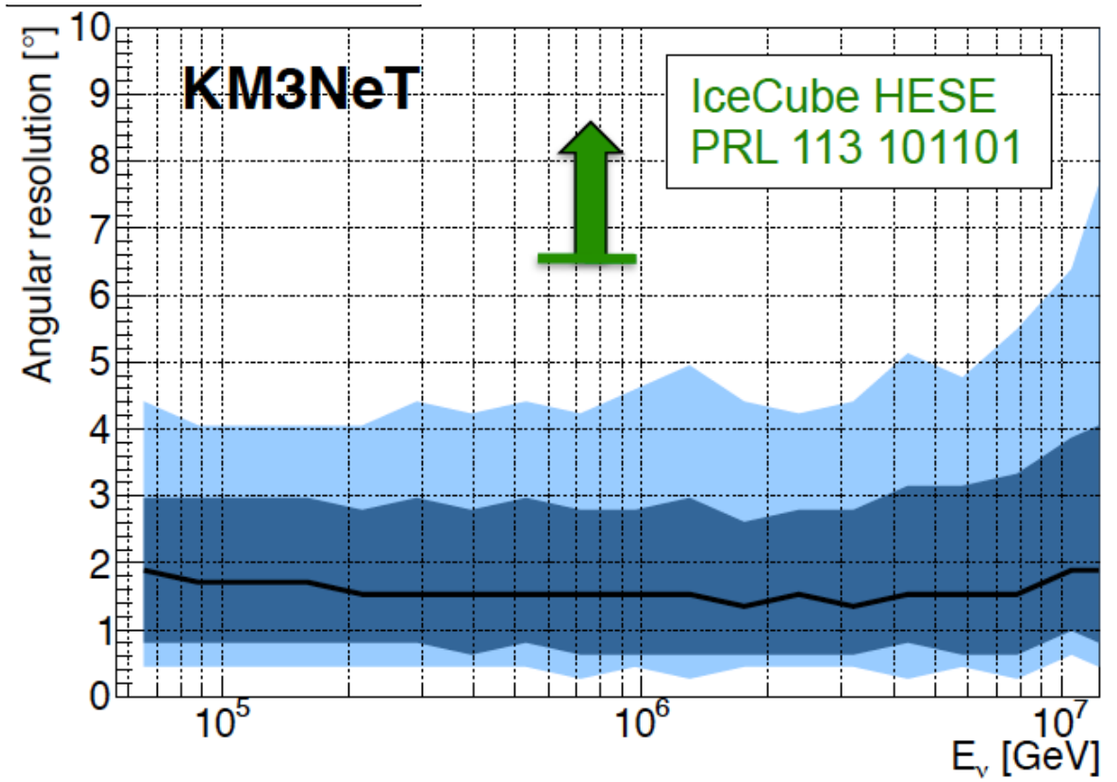


- Direction:
- ➡ Gal. srcs: 0.2° at 10TeV [0.4° for ANTARES]
- ➡ Extra-gal. srcs: 0.1° at 100TeV [0.3° for ANTARES]
- Energy: 0.27 in Log10(E)



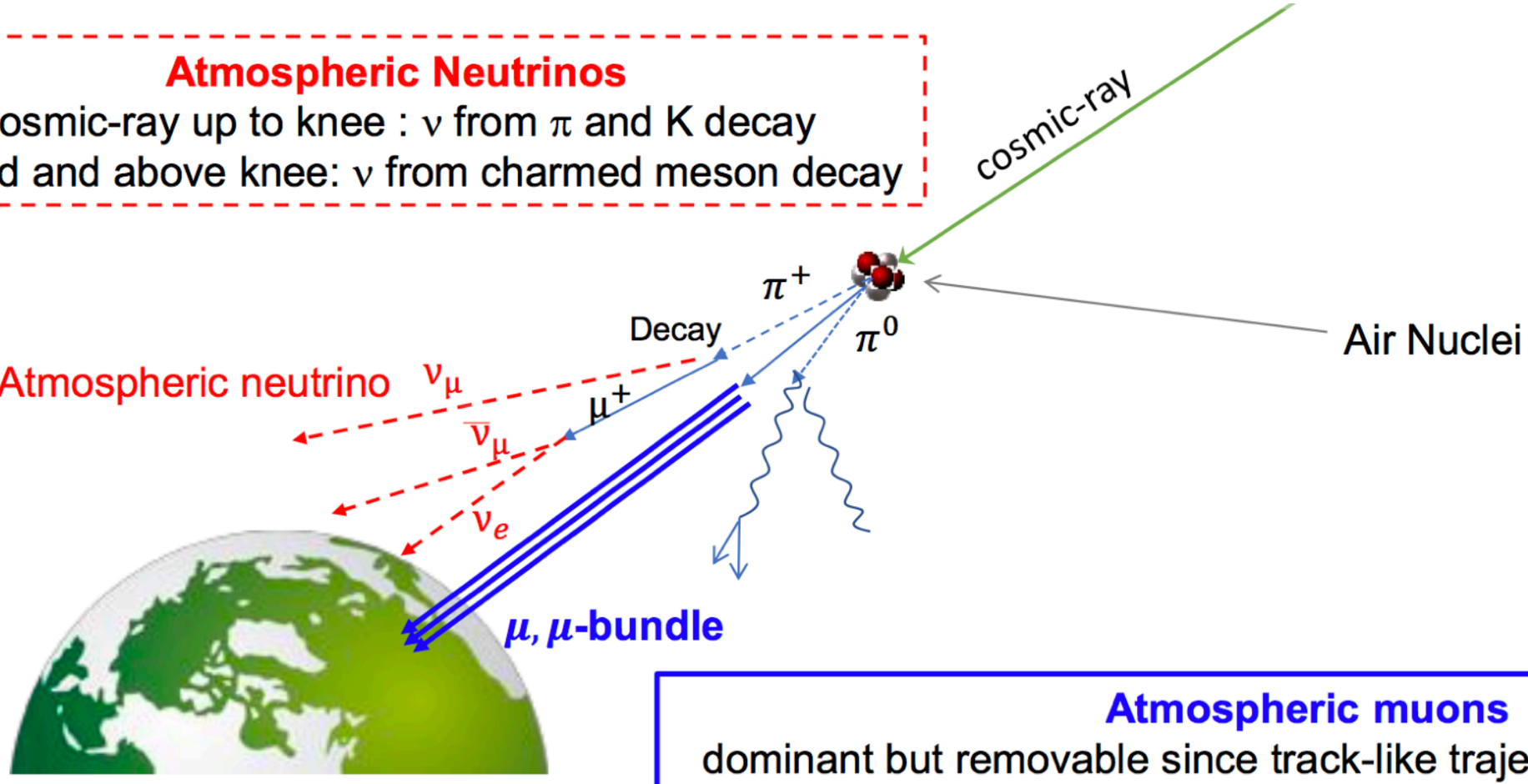


- Vertex: 6-8m (long), 0.5m (perp)
- Direction: ~1.5° [3° for ANTARES]
- Energy: 5%



LARGELY DOMINATED BY ATM BKG

Atmospheric Neutrinos
cosmic-ray up to knee : ν from π and K decay
around and above knee: ν from charmed meson decay



Atmospheric muons
dominant but removable since track-like trajectories of Cherenkov photons and its directions is able to be reliably reconstructed

<u>To have better discovery potential:</u>

- **Have the lowest angular precision (tracks)**
- **Have the lowest background contamination (cascades)**
- **Search for time+space-correlations**

THE ICECUBE SIGNAL

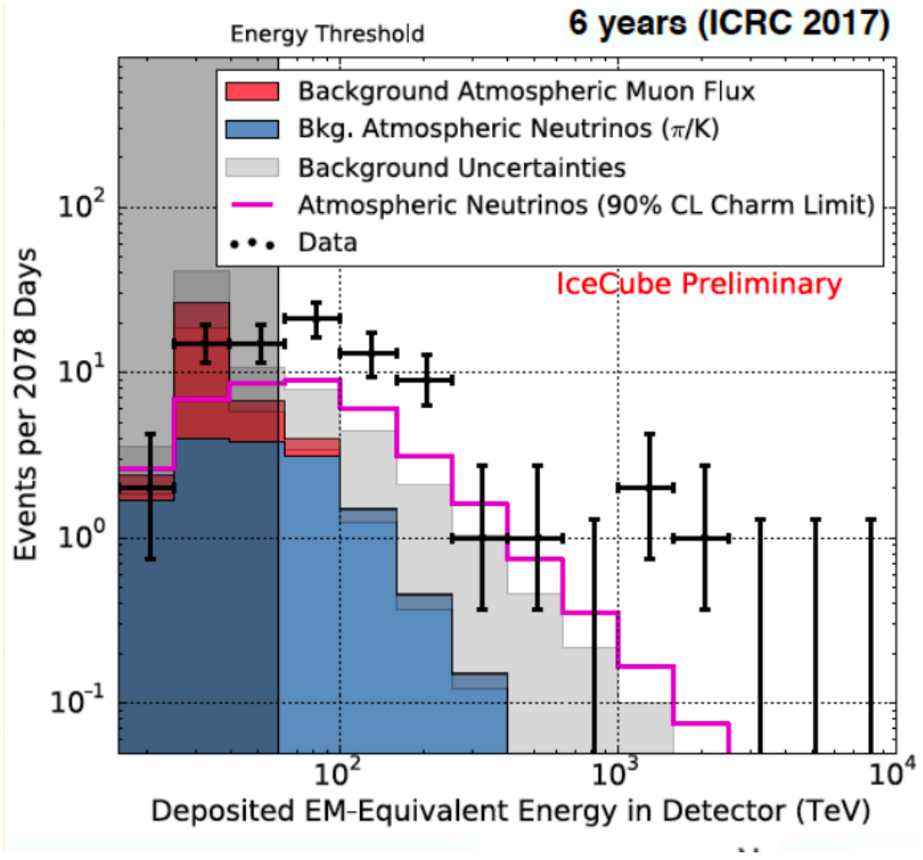
6 year HESE analysis (ICRC 2017)

80(+2) events

Bkg: 15.6+11.4-3.9 atm nu + 25.2+/-7.3 atm mu

Hemipshere North and South

E_{th}: 60 TeV



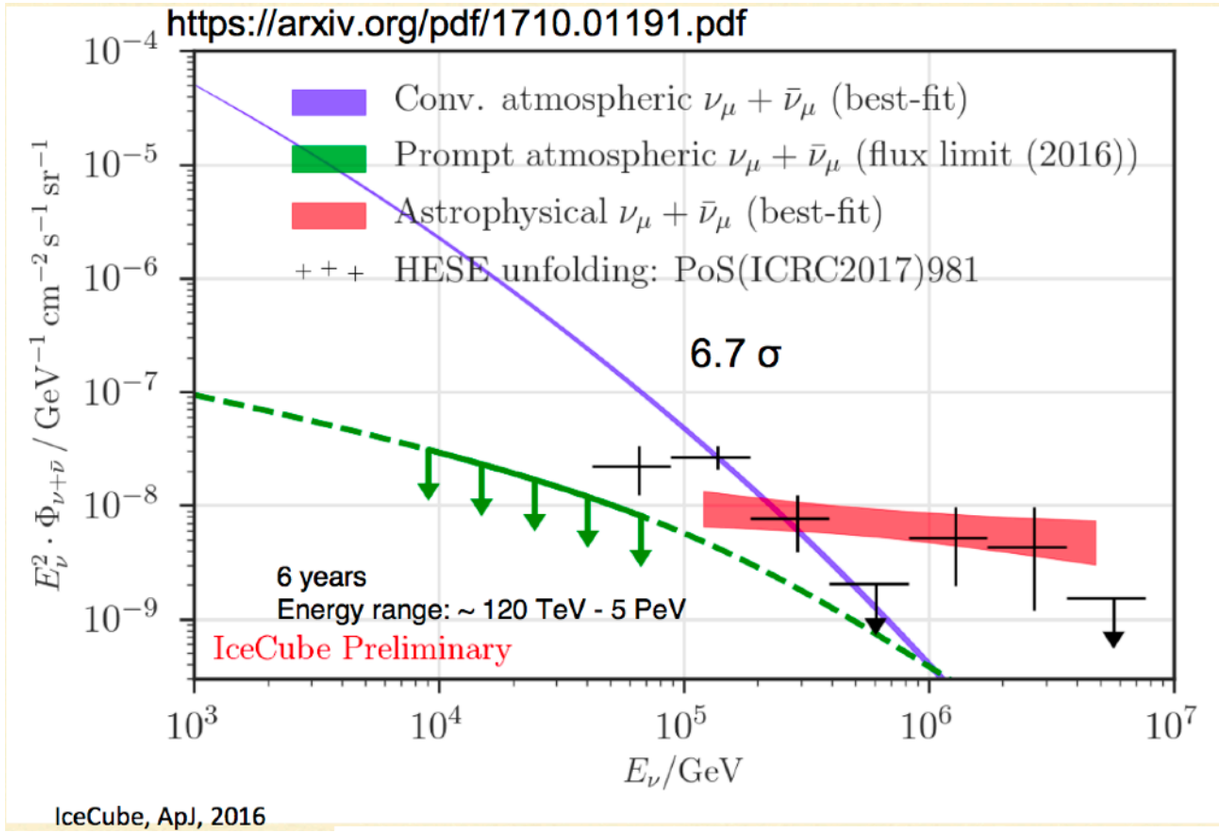
Significance: 6.5 sigma

Spectra: $E^{-2.92(+0.33\ -0.29)}$

8 year upgoing muon

E_{th}: 200 TeV

E_{event} >5 PeV !



Significance: 6.7 sigma

Spectra: $E^{-2.19(+/-0.10)}$

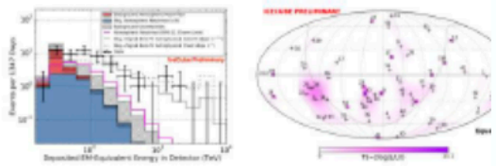
➡ Indication of a break in spectrum? (energy threshold different)

➡ Indication of galactic and extra-galactic components? (different hemispheres)

THE ICECUBE SIGNAL

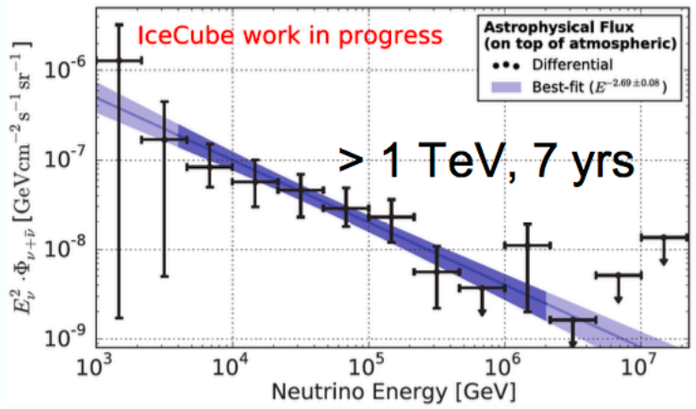
Last update for the starting track analysis

$E > 50\ TeV$



IceCube, ICRC 2017

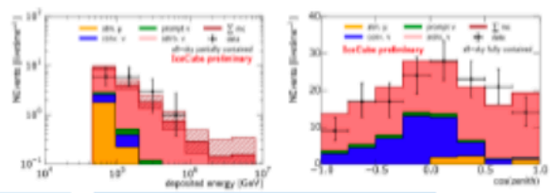
$\gamma = -2.92 \pm 0.30$



IceCube, PRD, 2015And TeVPA2017

$\gamma = 2.69 \pm 0.08$

showers (+partially cont.)



IceCube, ICRC, 2015

$\gamma = -2.7 \pm 0.12$

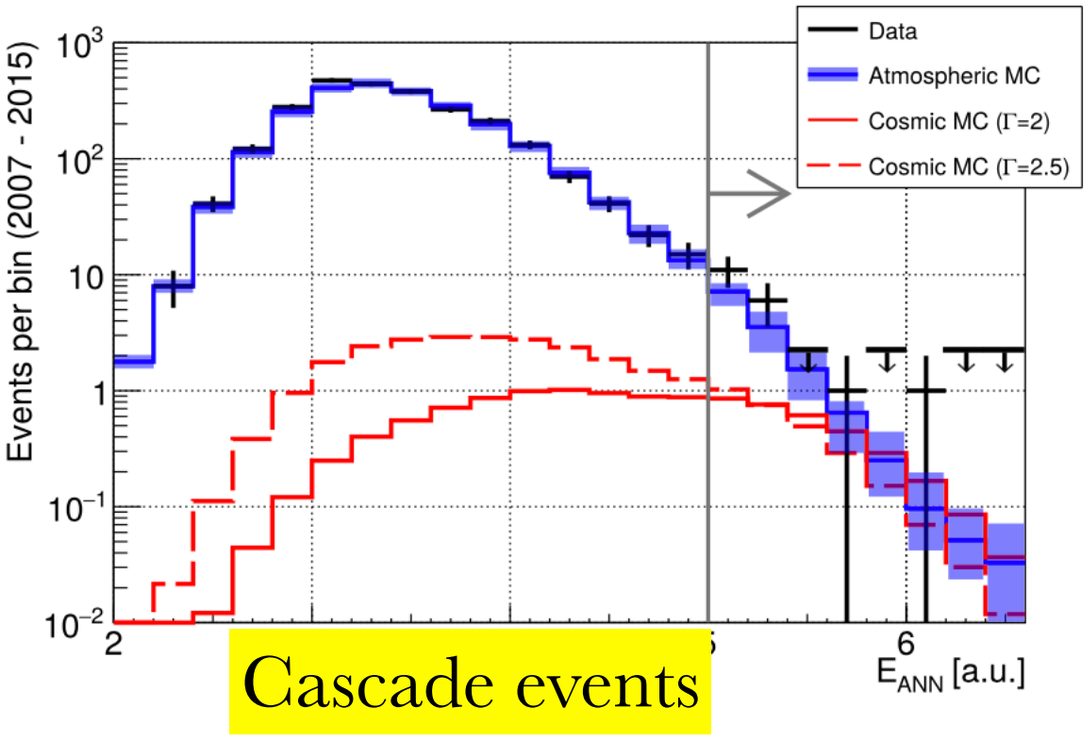
Soft
Spectral index?

5

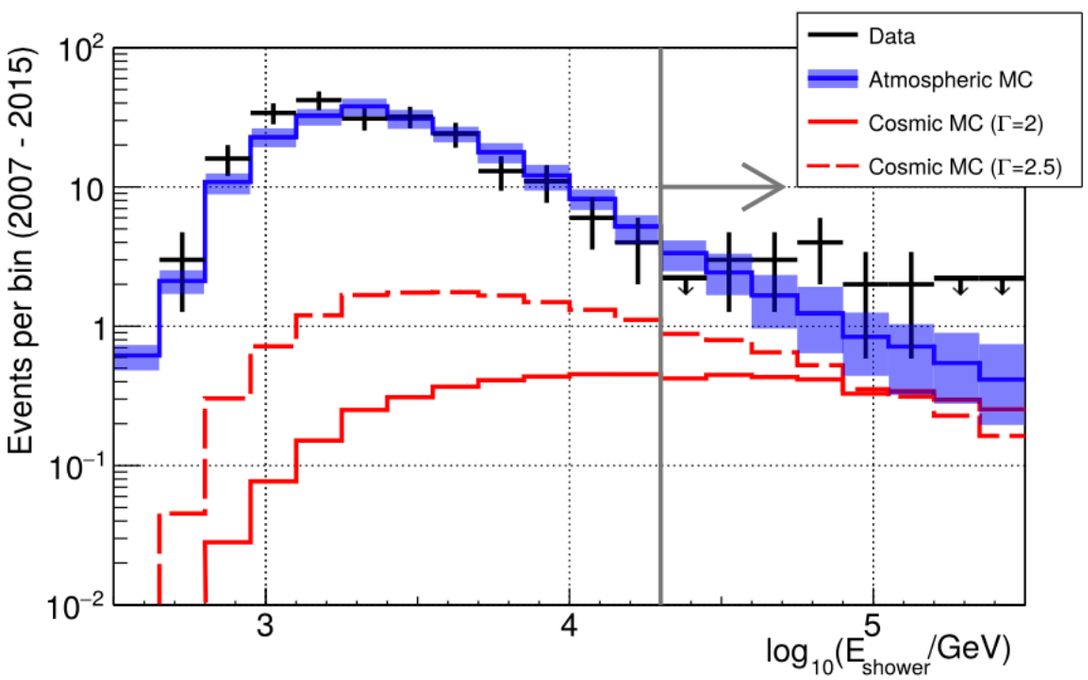
Analysis	Index	Normalization @ 100 TeV	Significance (σ)	Energy range
HESE 6 yr	2.92 ± 0.3	2.46 ± 0.8	8	60 TeV to 3 PeV
Northern tracks 6 yr	2.19 ± 0.10	1.01 +0.26 -0.23	6.7	119 TeV to 4.8 PeV
Cascades 4 yr	2.48 ± 0.08	1.57 +0.23 -0.22	4.7 (2 year)	10 TeV to 1 PeV
Global fit	2.50 ± 0.09	2.2 ± 0.4		25 TeV to 2.8 PeV

ANTARES DIFFUSE RESULTS

Track events



Cascade events

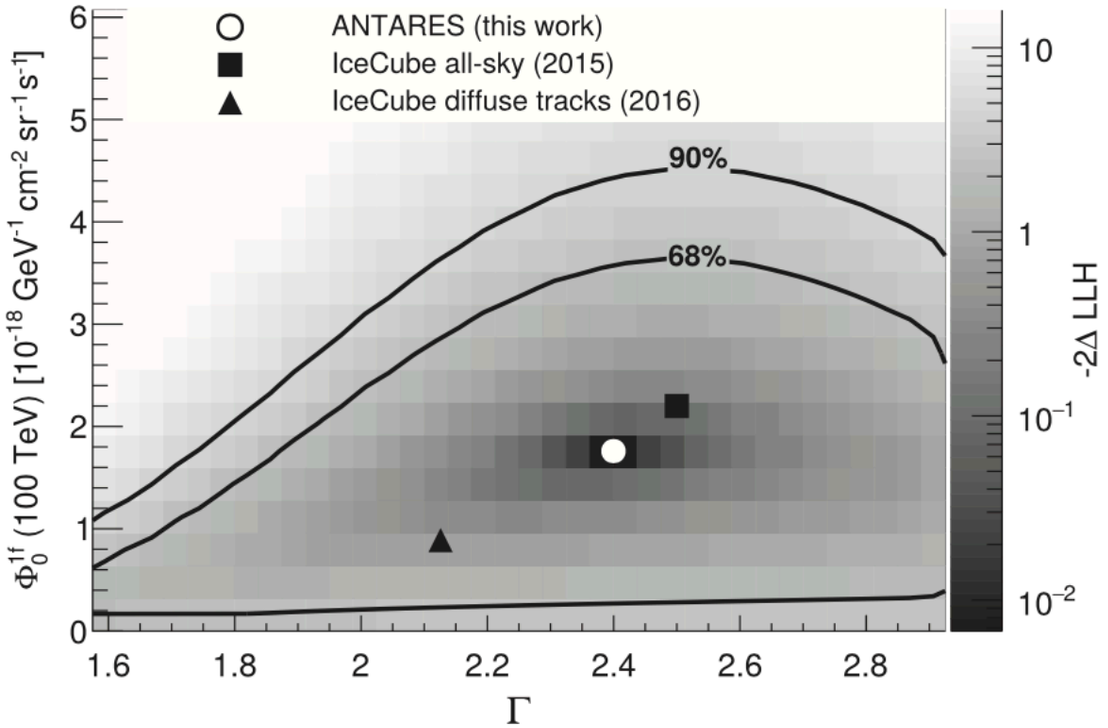


All-sky / All-flavor neutrino search

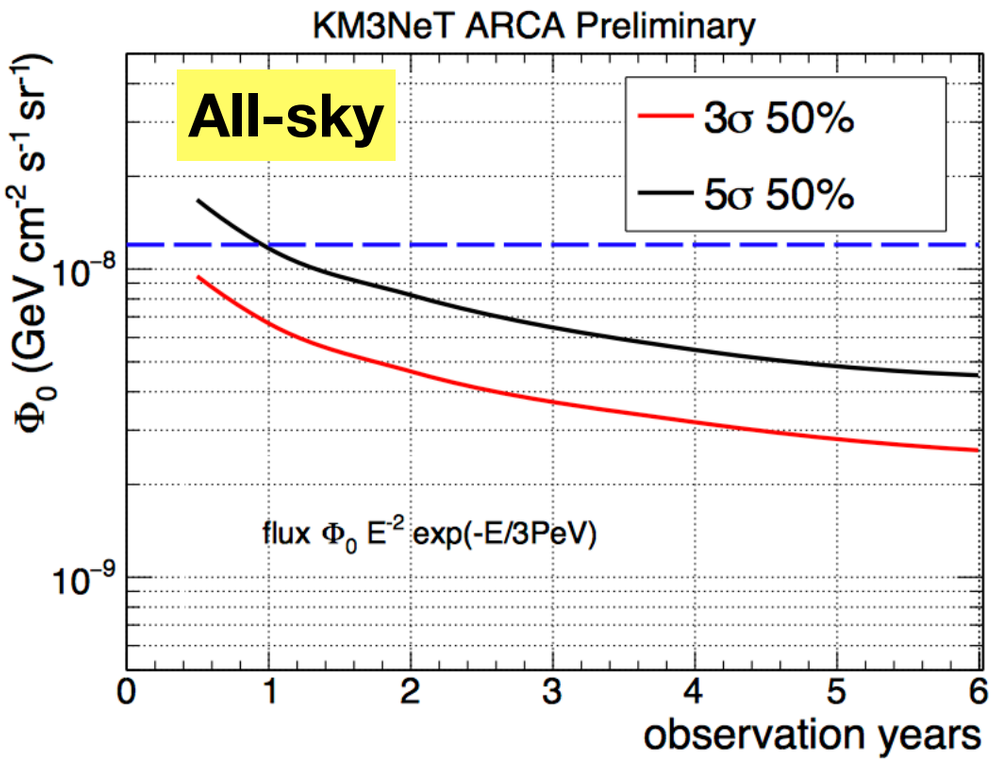
- Look for excess above a given E_{th}
- 9 yrs of data for tracks (cascades)

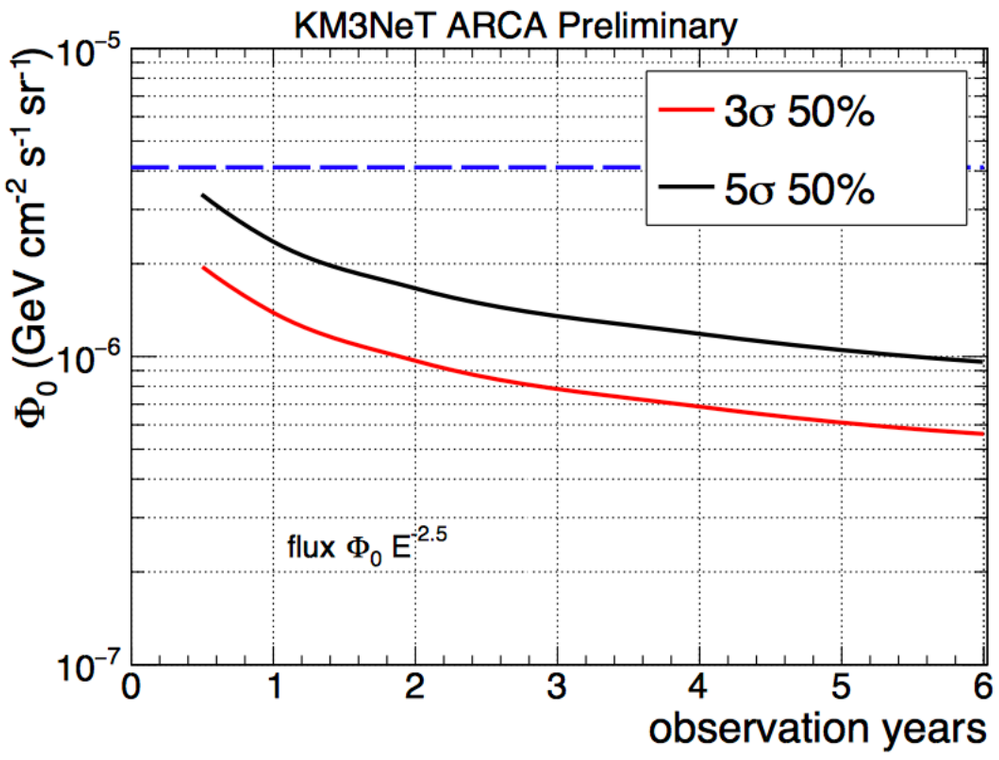
	Bkg expectation	Signal exceptation	Nb events measured
Track	**13.5+/-4**	**3-3.5**	**19**
Shower	**10.5+/-4**	**3-3.5**	**24**

=> Small excess (not significant)



KM3NeT: DIFFUSE FLUX

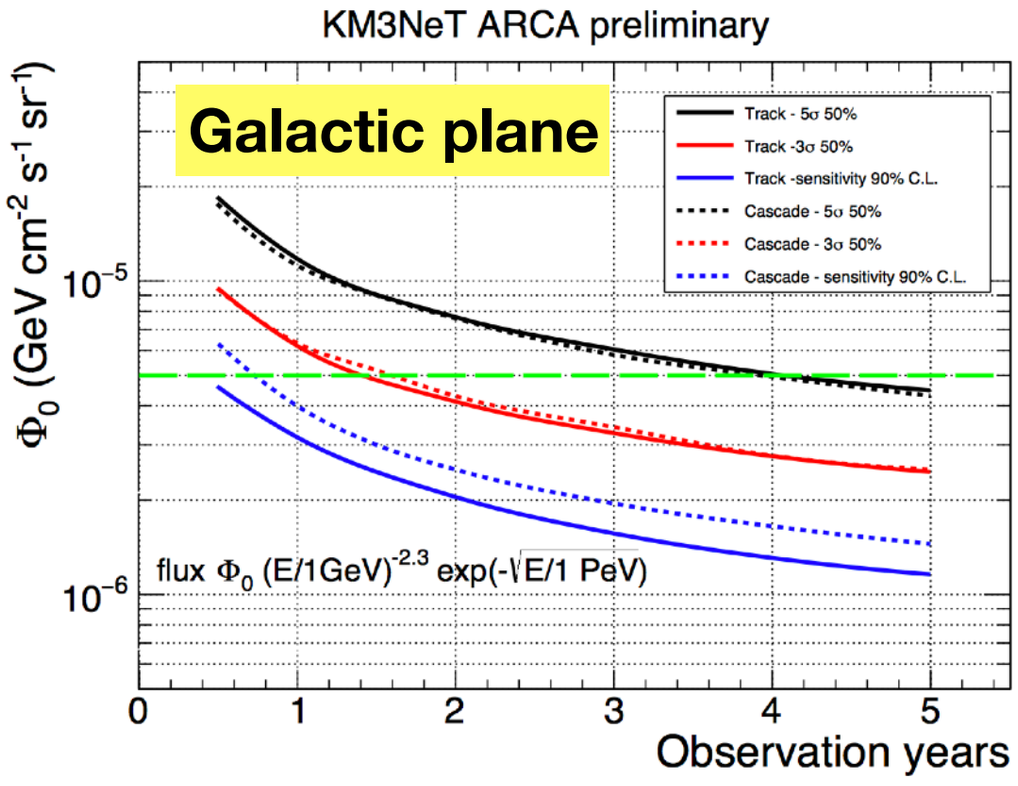




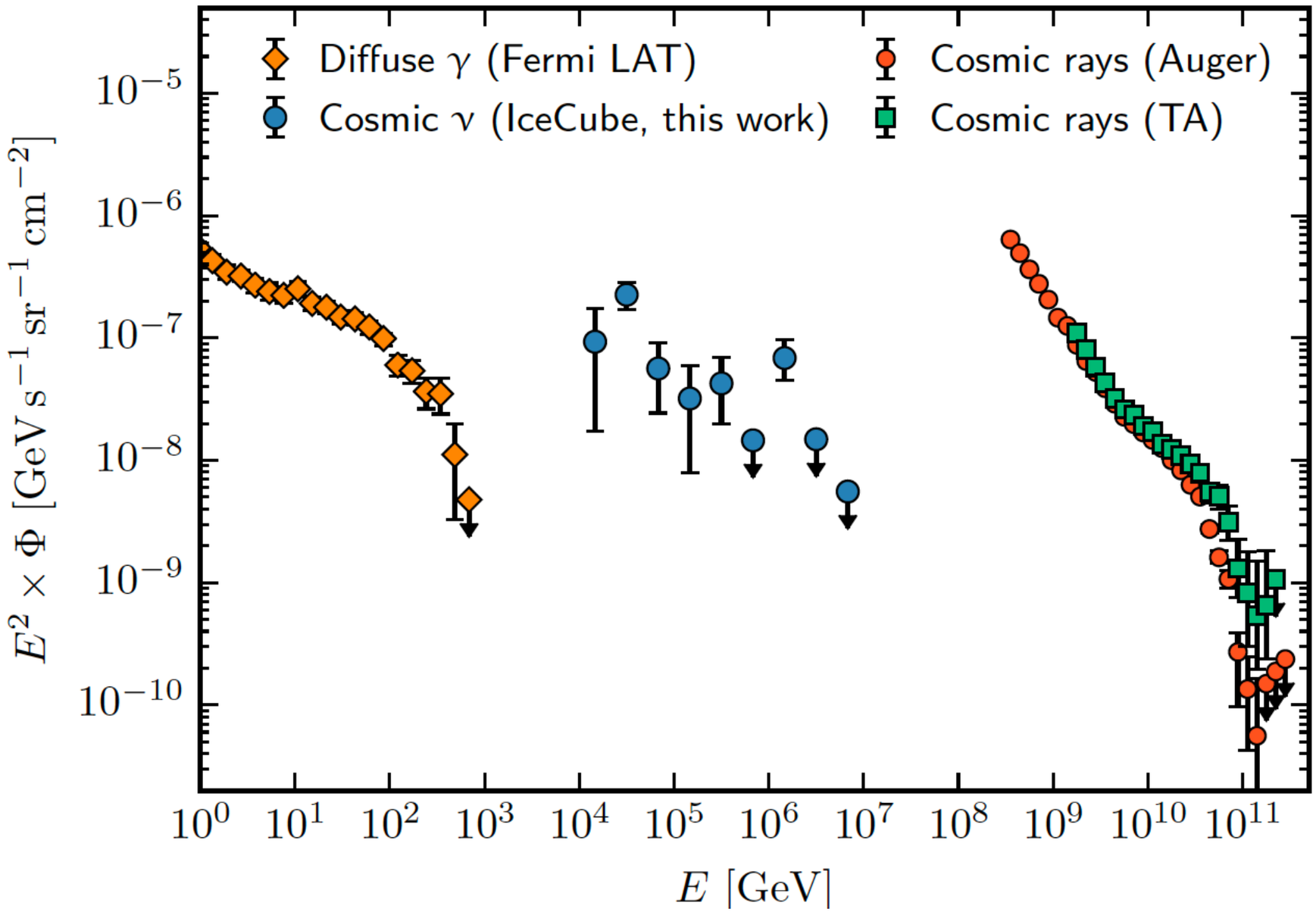
KM3NeT/ARCA is expected to observe the IC signal in less than 1 yr.

➡ Precise characterization (spectral shape, flavor composition, anisotropy)

➡ Excellent sensitivity in the galactic plane: identify gal/extra-gal components ?



γ-ν-RC DIFFUSE FLUXES



⇒ Energy density of neutrinos in the non-thermal Universe is the same or higher as that in Fermi gamma-rays.

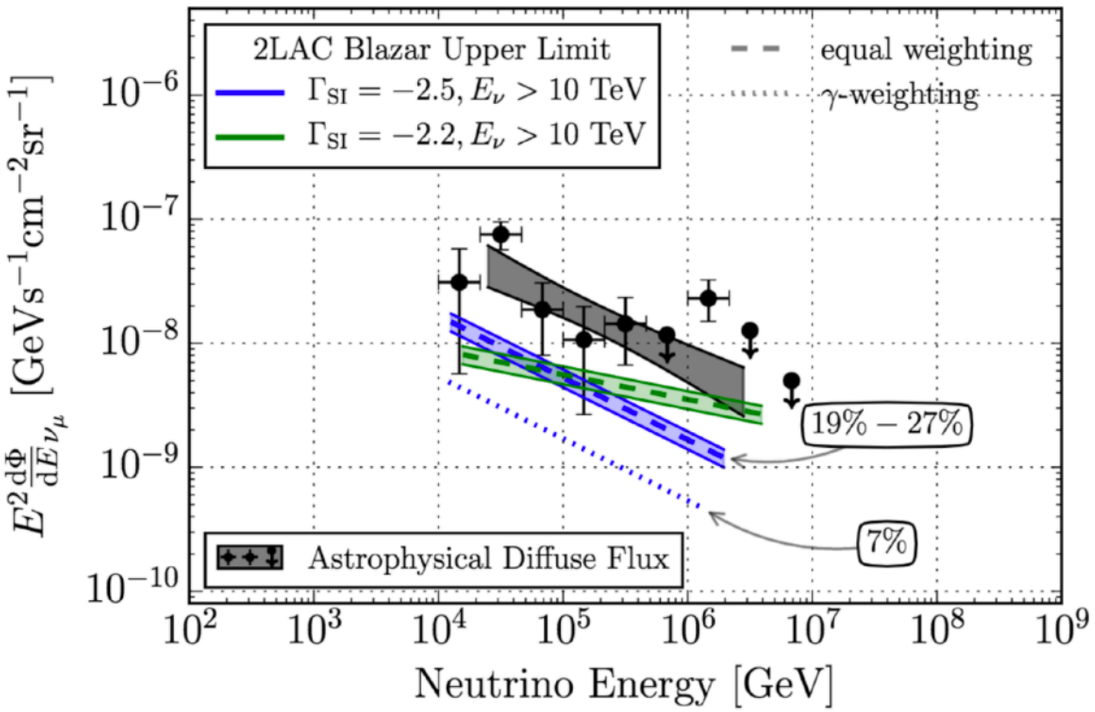
⇒ Common sources ? Fermi/LAT γ-ray flux dominated by AGN/blazars (~ 85%)

POPULATION STUDIES

Blazar space correlation

(862 '2LAC' blazars)

Contribution max of the 2LAC blazars < 27% (10 TeV - 2 PeV), assuming equal weighting among blazars and single power-law with γ = –2.5.



7% of neutrino signal assuming ν flux ⇔ γ-ray flux

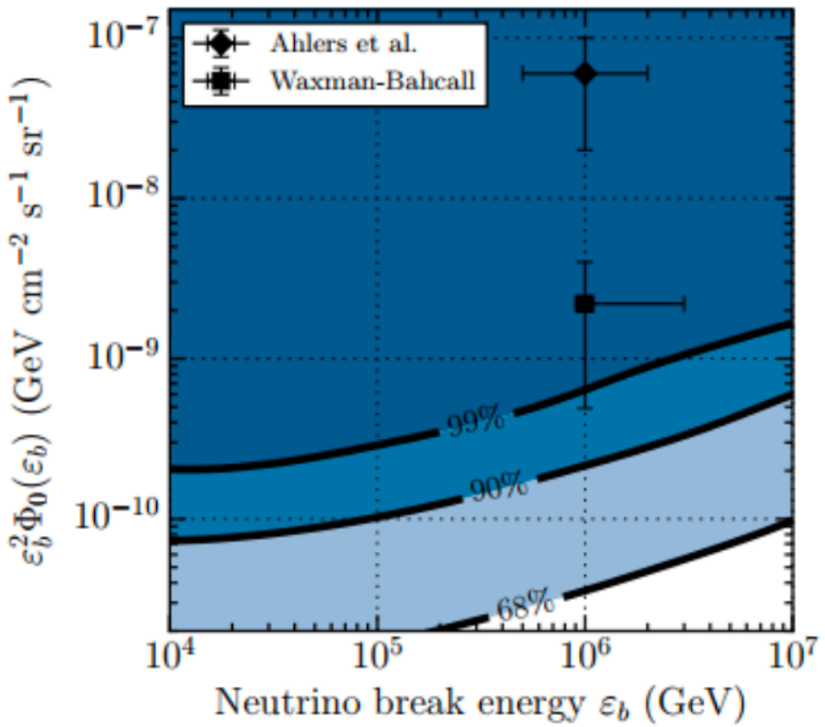
(correlation with 2FHL: < few % of the IC flux)

Astrophysical Journal 835 (2017) 1

GRB time/space correlation

contribute **no more than 1%** of the observed diffuse flux

$$\Phi_\nu(E_\nu) = \Phi_0 \times \begin{cases} \varepsilon_b^{-1} E_\nu^{-1}, & E_\nu \leq \varepsilon_b \\ E_\nu^{-2}, & \varepsilon_b < E_\nu \leq 10\varepsilon_b \\ E_\nu^{-4} (10\varepsilon_b)^2, & 10\varepsilon_b < E_\nu, \end{cases}$$



(1172 GRBs - benchmark parameters)

arXiv:1702.06868

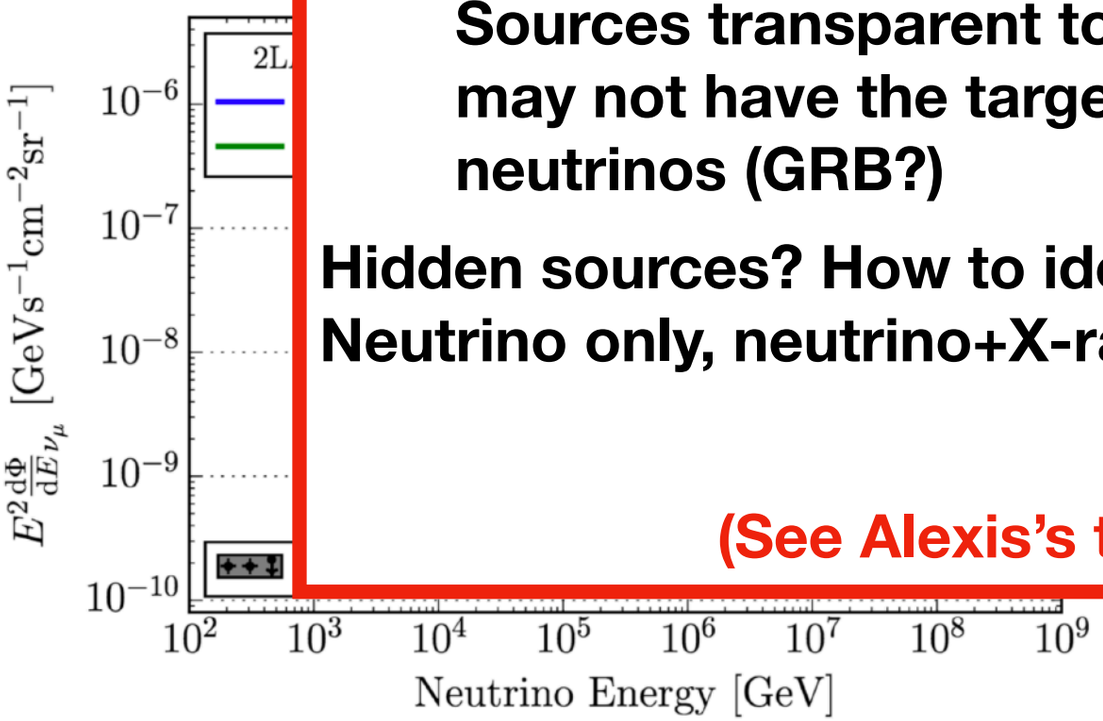
POPULATION STUDIES

Blazar space correlation

(862 '2LAC' blazars)

Contribution max of the 2LAC blazars < 27% (10 TeV - 2 PeV), assu

single power



BUT, neutrinos originate from a larger volume

50% of blazars not identified

Sources transparent to high energy gamma rays may not have the target density to produce neutrinos (GRB?)

Hidden sources? How to identify these sources ?
Neutrino only, neutrino+X-ray ?

(See Alexis's talk tomorrow)

7% of neutrino signal assuming ν flux ⇔ γ-ray flux

(correlation with 2FHL: < few % of the IC flux)

Astrophysical Journal 835 (2017) 1

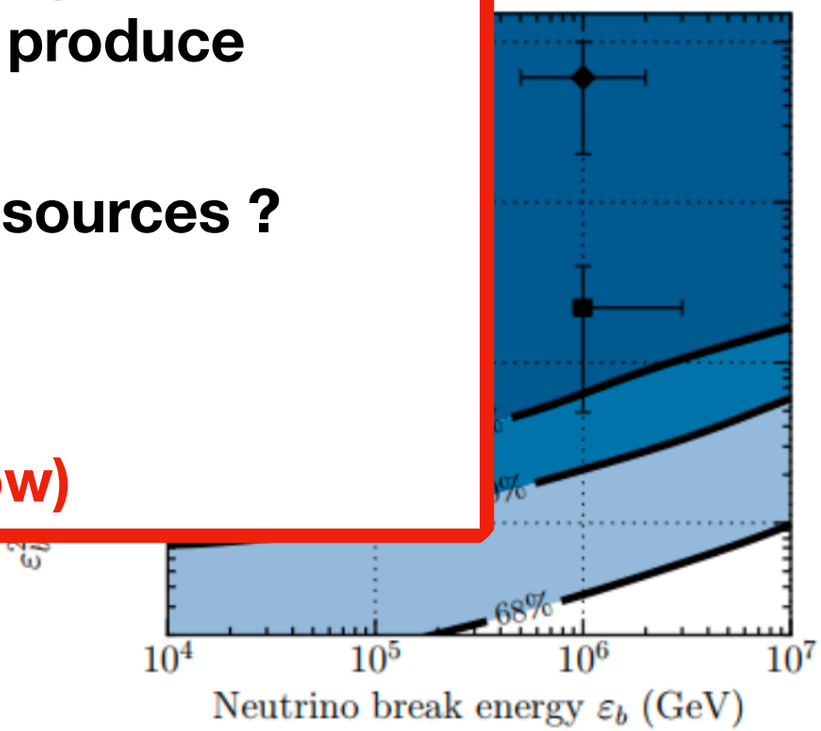
GRB time/space correlation

contribute **no more than 1%** of the observed diffuse flux

^{1}, $E_\nu \leq \varepsilon_b$

$\varepsilon_b < E_\nu \leq 10\varepsilon_b$

$0\varepsilon_b)^2$, $10\varepsilon_b < E_\nu$,



(1172 GRBs - benchmark parameters)

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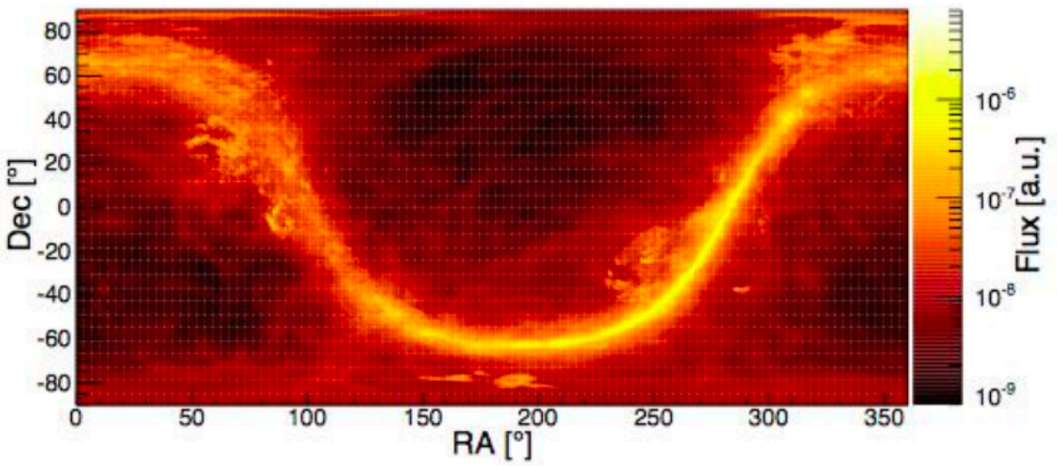
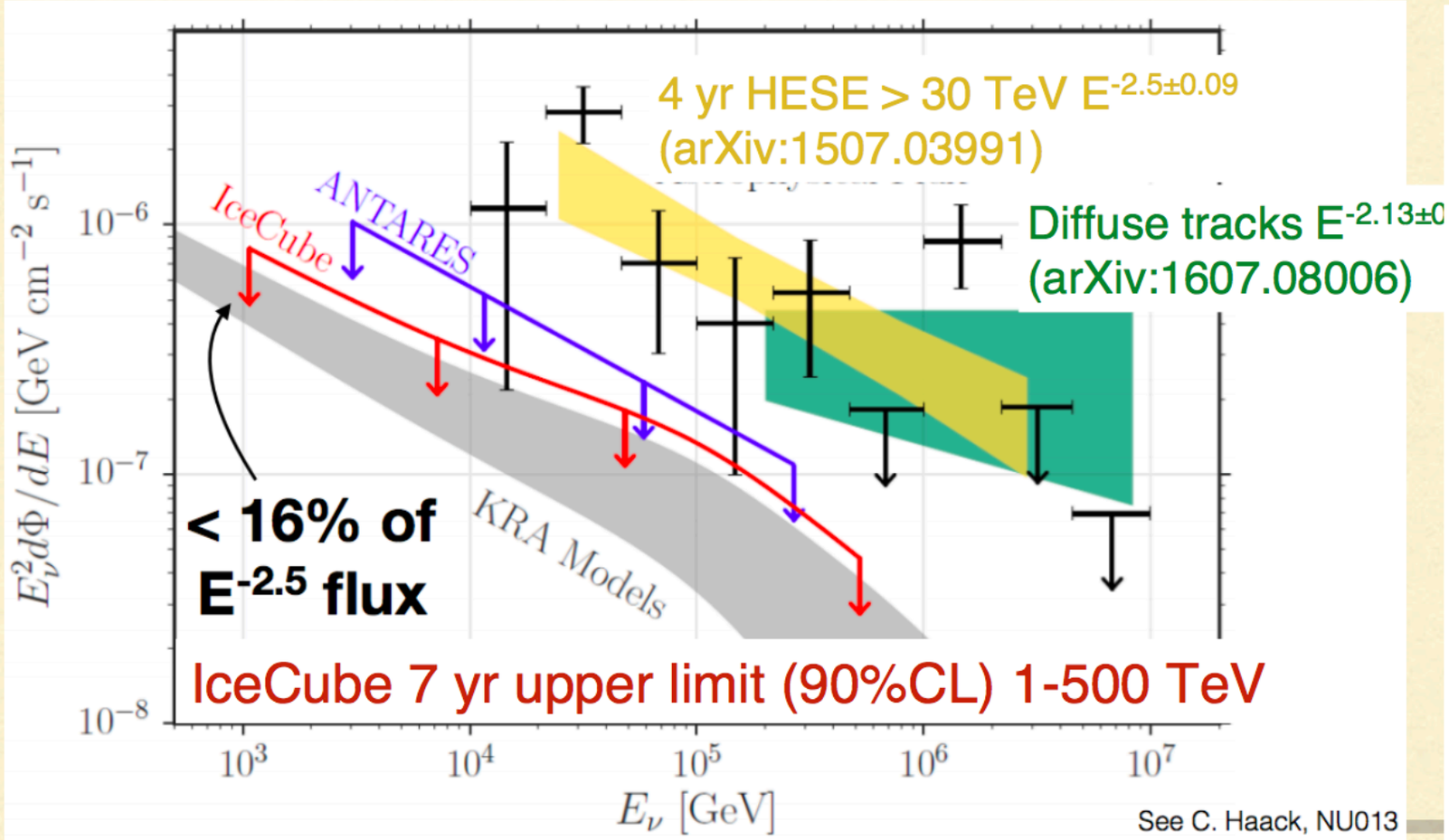
GALACTIC DIFFUSE FLUX



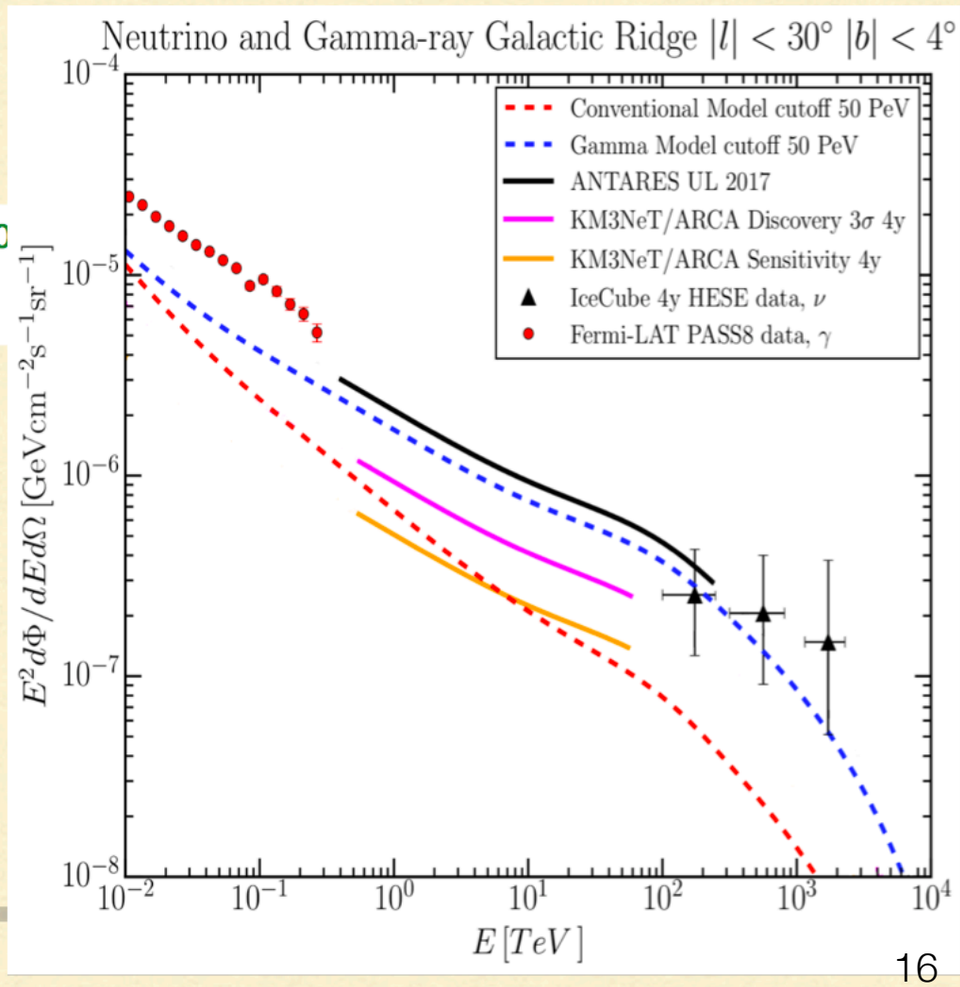
(a) KRA-γ (50 PeV cutoff) template

- Analysis of correlation with template map derived from interstellar gas distribution reproducing Fermi-LAT data Models in Gaggero et al, arXiv:1504.00227
- Only small fraction of signal can originate from CR interactions in the Galaxy. UL for IC and ANTARES 1.2 x KRA-γ (50 PeV)

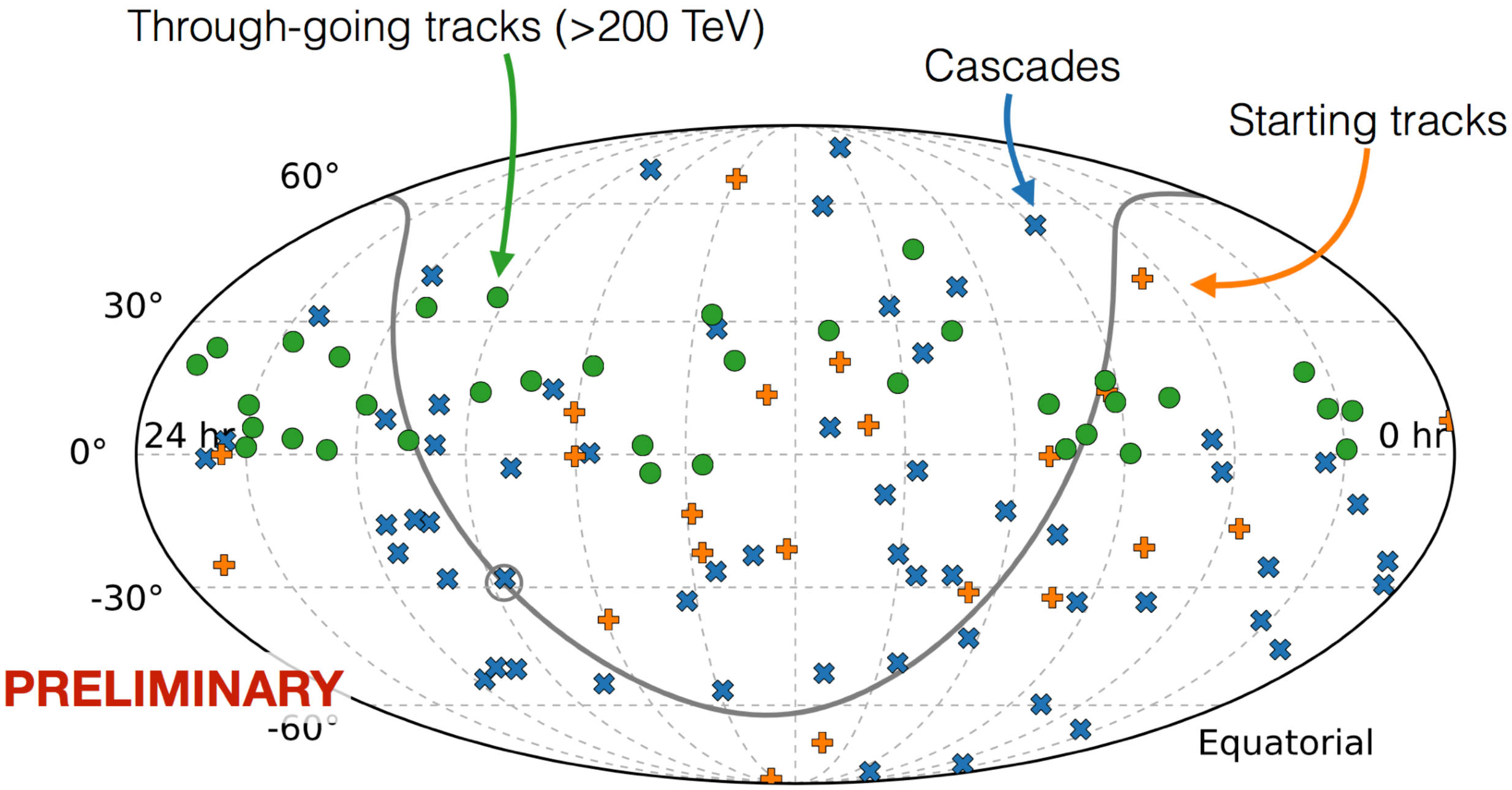
ANTARES arXiv:1602.03036 updated at this conference



See C. Haack, NU013
arXiv:1707.03416



IC NEUTRINO SKYMAP



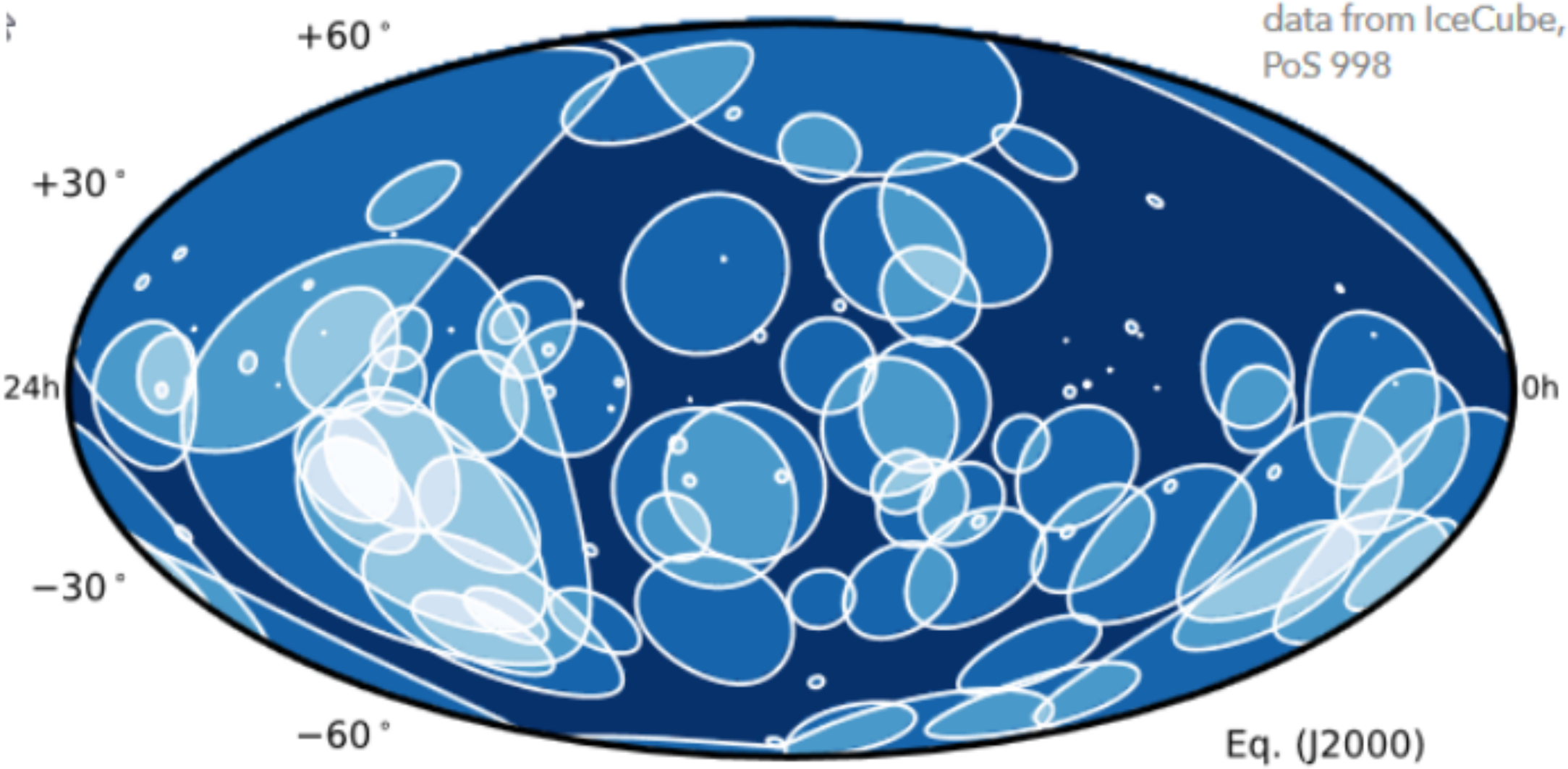
IceCube, 2017

No evidence of clustering in high-energy neutrino directions

mostly isotropic ⟹ neutrinos of extragalactic origin

Where are the PeV γ-rays together with PeV neutrinos ?

IC NEUTRINO SKYMAP



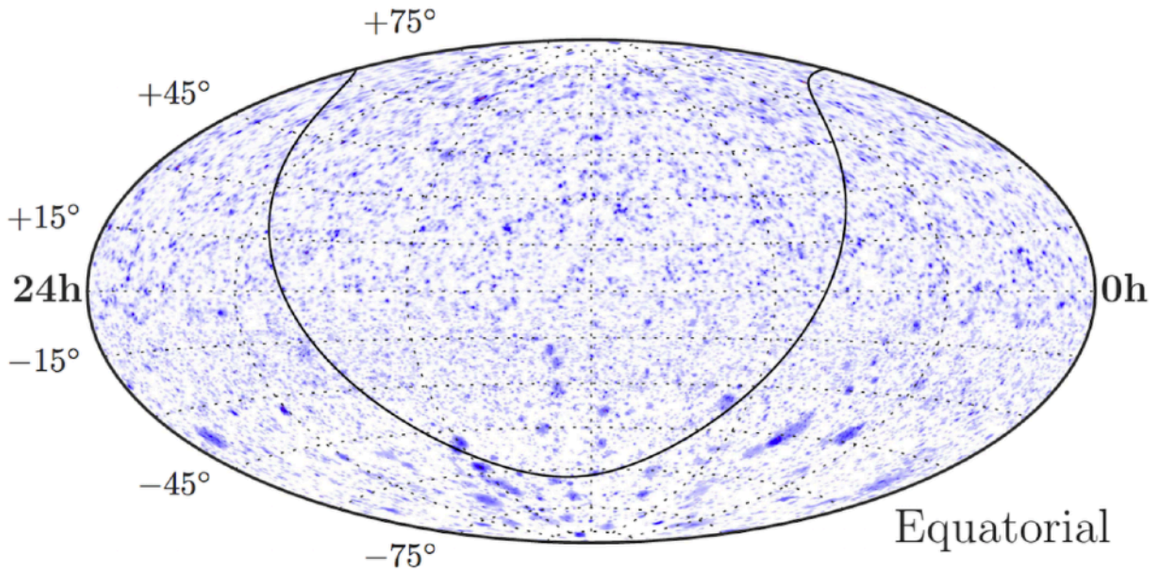
Resolution for ν_e

ANTARES ○

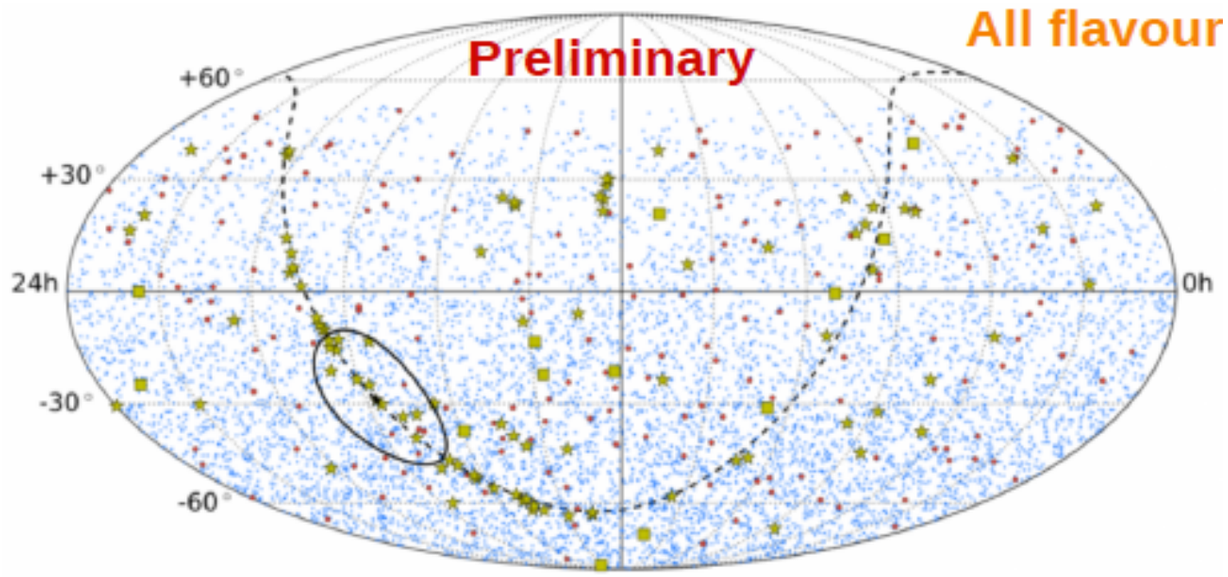
KM3NeT ○

LOOKING FOR POINT-SOURCES

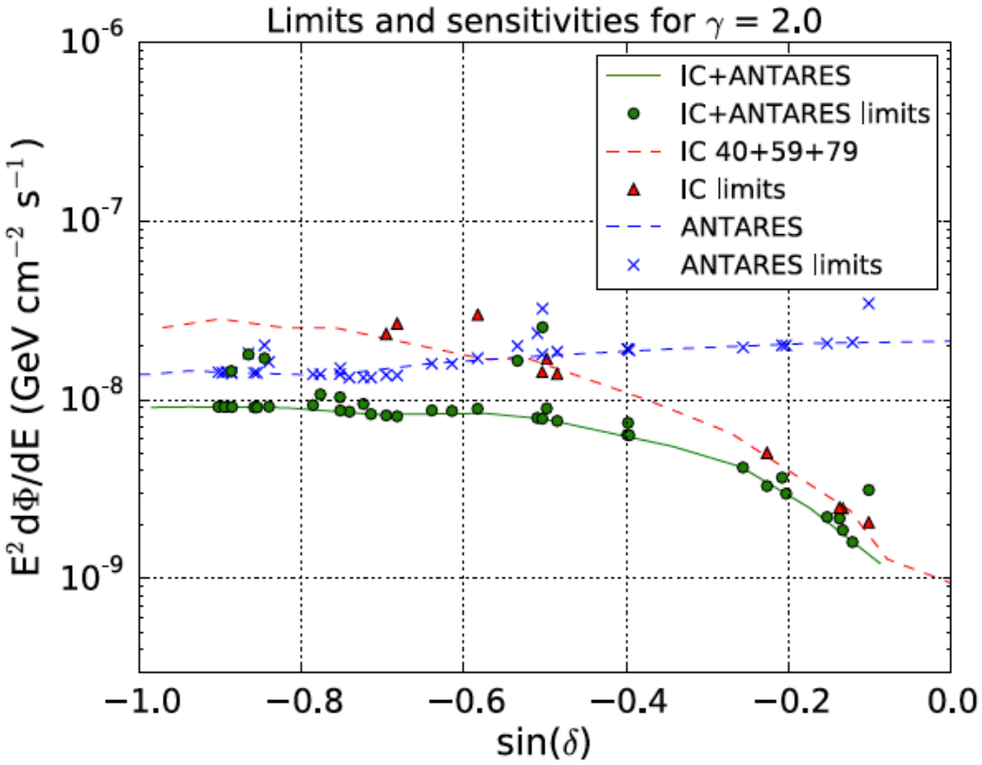
IceCube (7 yrs - tracks)

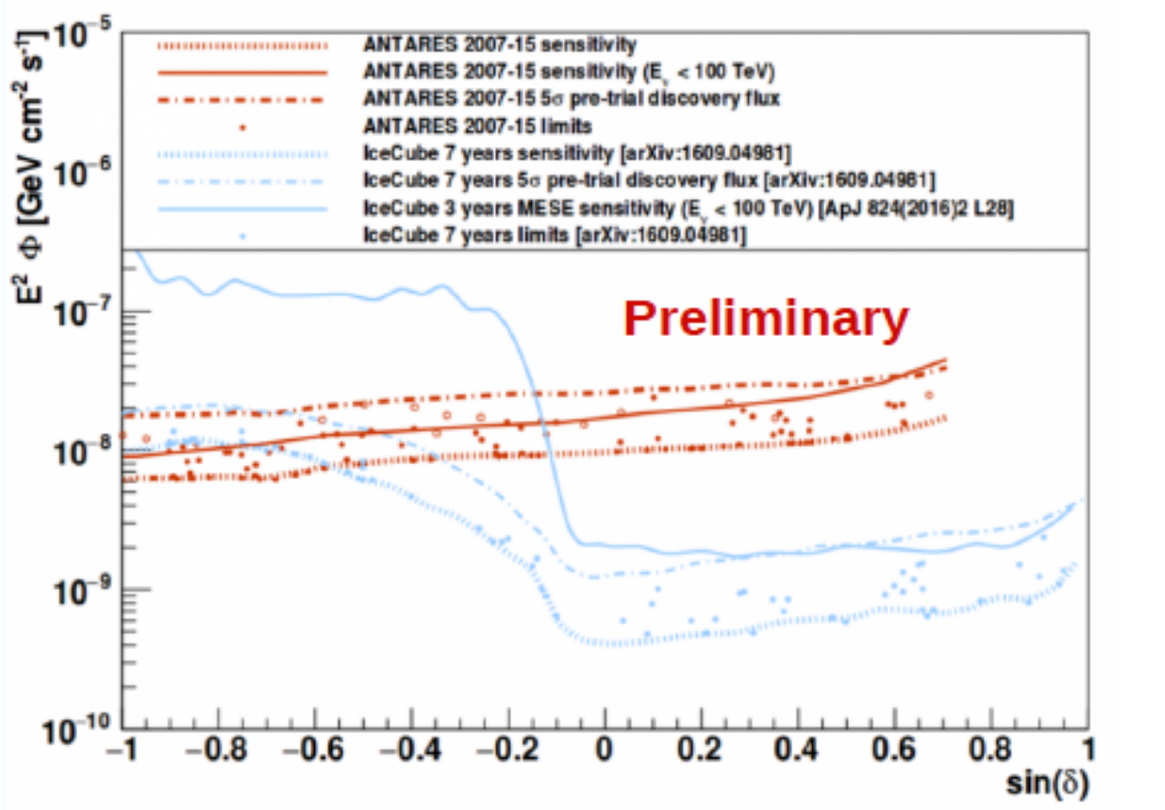


ANTARES



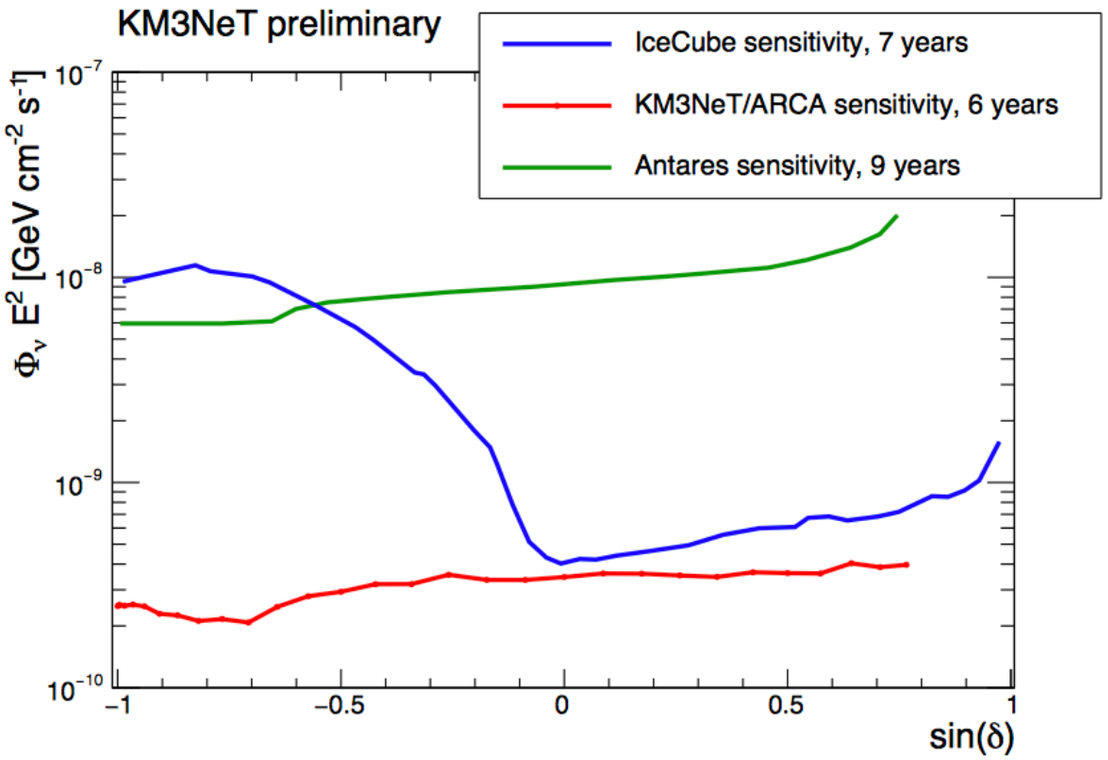
ANTARES+IceCube

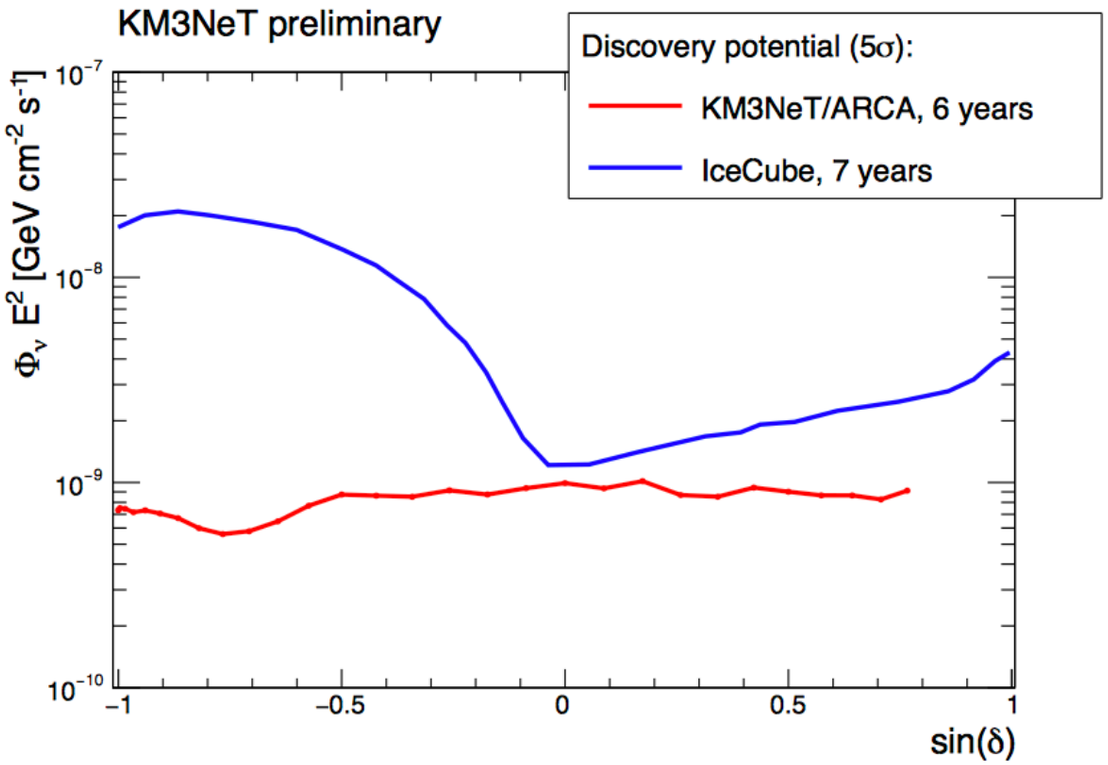




KM3NeT: POINT-SOURCE

Generic source

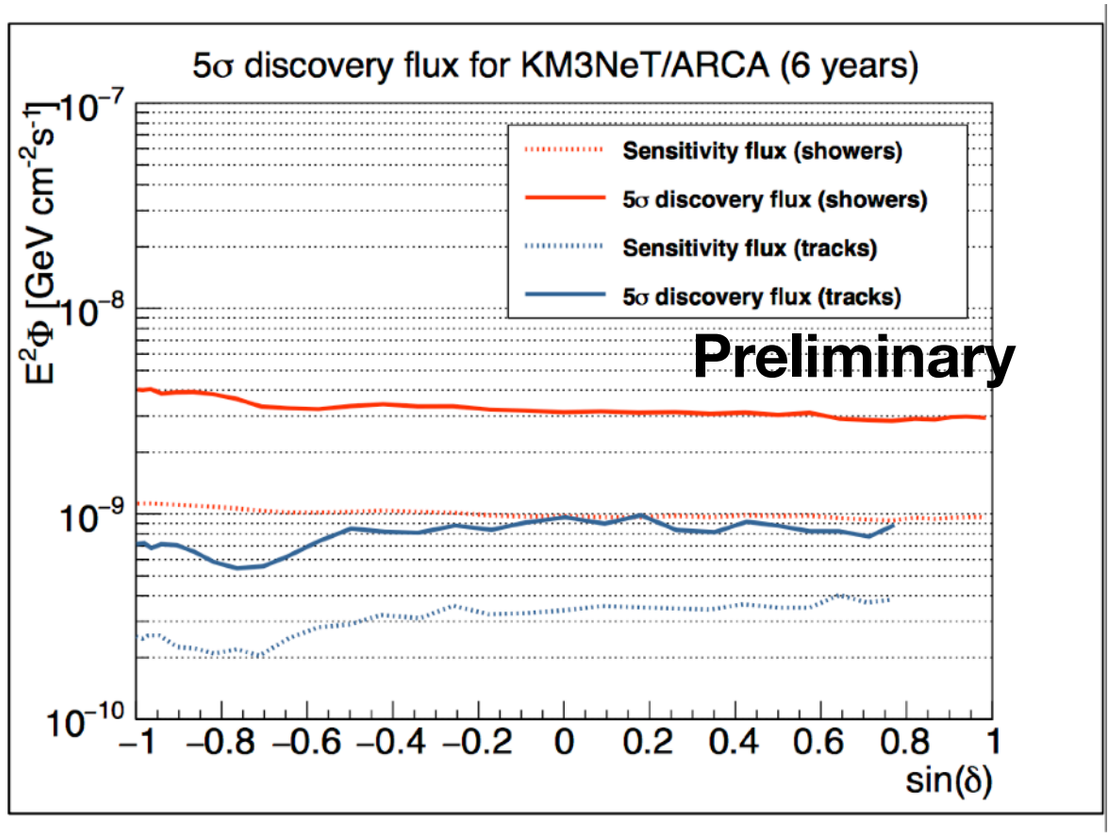




KM3NeT/ARCA is expected to have more than one order of magnitude better sensitivity than IC in the Southern sky.

➡ Due to the quite good angular resolution for cascade events, the point-source search is also very efficient.

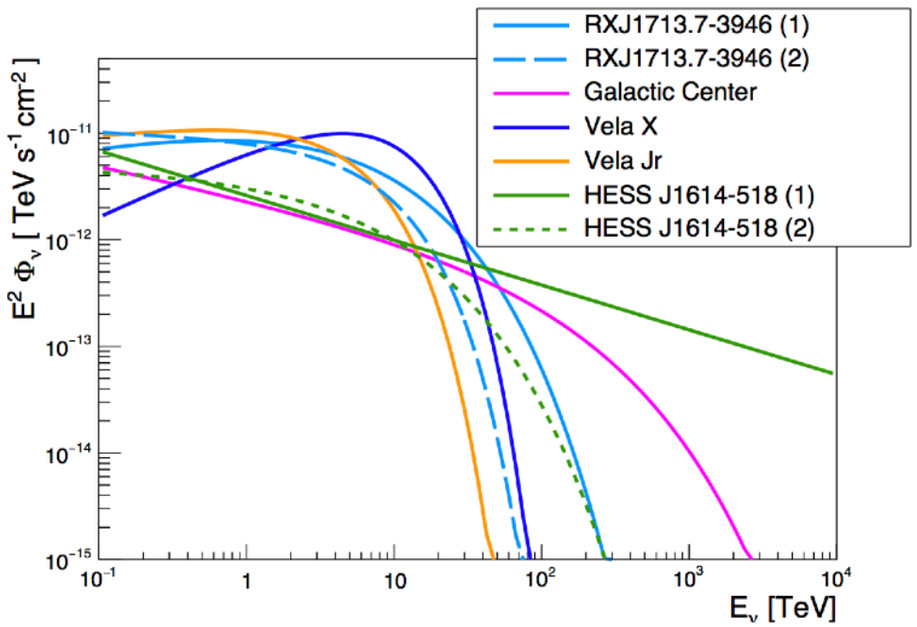
➡ Expected better performances for the transient neutrino sources (GRB, AGN…)



KM3NeT: POINT-SOURCE

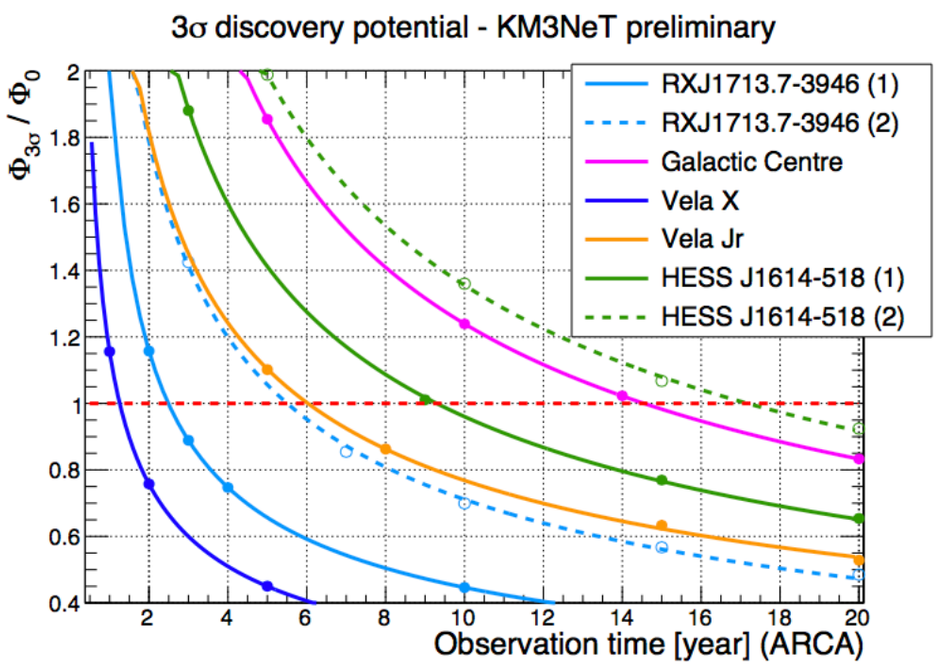
Specific galactic sources

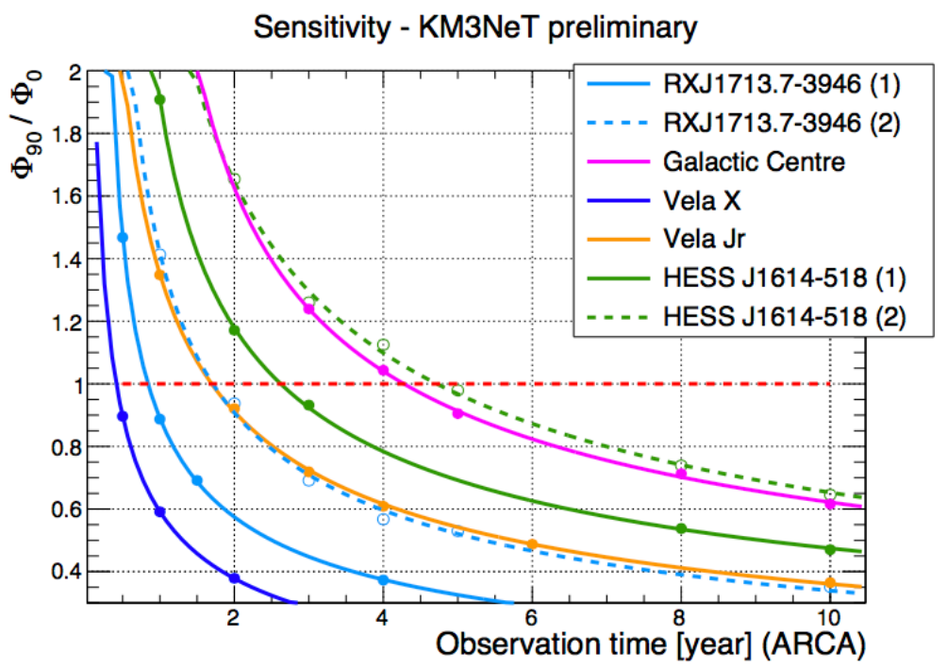
Source	δ	extension	Φ_0	Γ	E_{cut}	β	γ-ray data
RX J1713.7-3946 (1)	-39.77°	0.6°	1.68	1.72	2.1	0.5	[13]
RX J1713.7-3946 (2)	-39.77°	0.6°	0.89	2.06	8.04	1	[14]
Vela X	-45.6°	0.8°	0.72	1.36	7	1	[15]
Vela Jr	-46.36°	1°	1.30	1.87	4.5	1	[16]
HESSJ1614-518 (1)	-51.82°	0.42°	0.26	2.42	-	-	[17]
HESSJ1614-518 (2)	-51.82°	0.42°	0.51	2	3.71	0.5	[17]
Galactic Centre	-28.87°	0.45°	0.25	2.3	85.53	0.5	[18]
MGRO J1908+06 (1)	6.27°	0.34°	0.18	2	17.7	0.5	see text
MGRO J1908+06 (2)	6.27°	0.34°	0.16	2	177	0.5	see text
MGRO J1908+06 (3)	6.27°	0.34°	0.16	2	472	0.5	see text



γ→ν flux conversion:

F. Vissani, *Astropart. Phys. 26* (2006), 310.
F. L. Villante and F. Vissani, *Phys. Rev. D 78* (2008), 103007.
F. Vissani and F. Villante, *Nucl. Instrum. Methods A 588* (2008), 123.



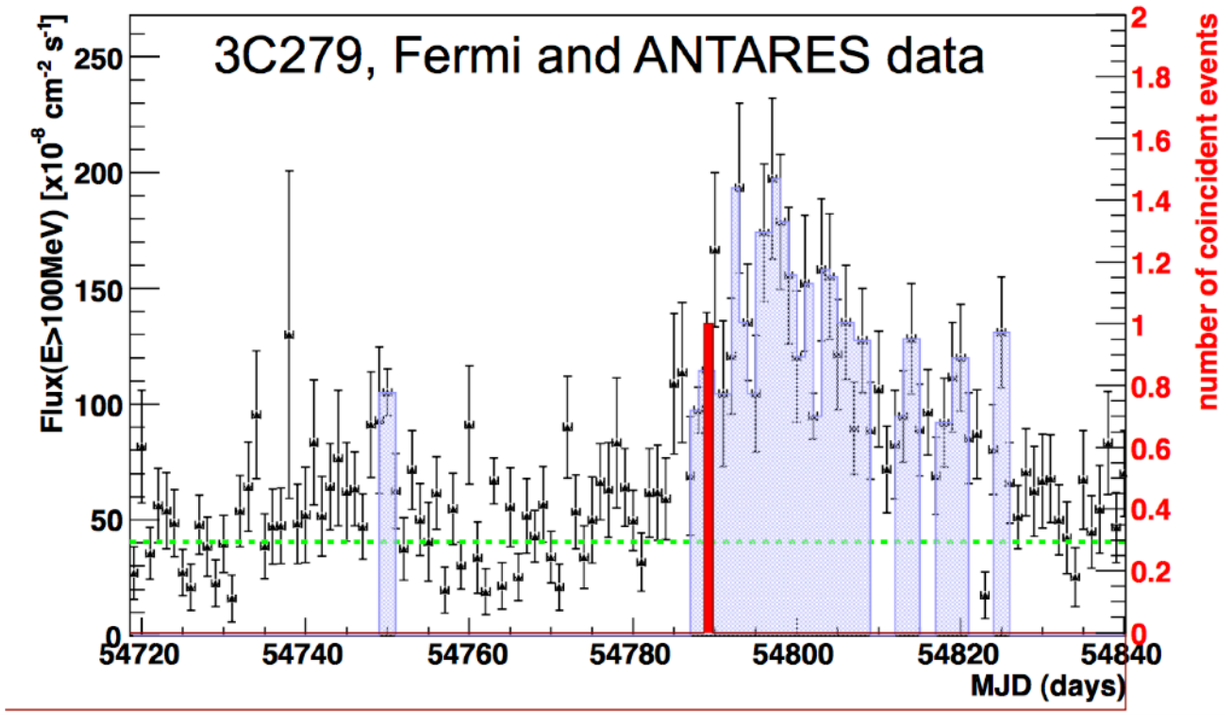


With reasonable 100% hadronic models, large probabilities to observe individual neutrino sources in the Galactic Plane

LOOKING FOR VARIABLE SOURCES

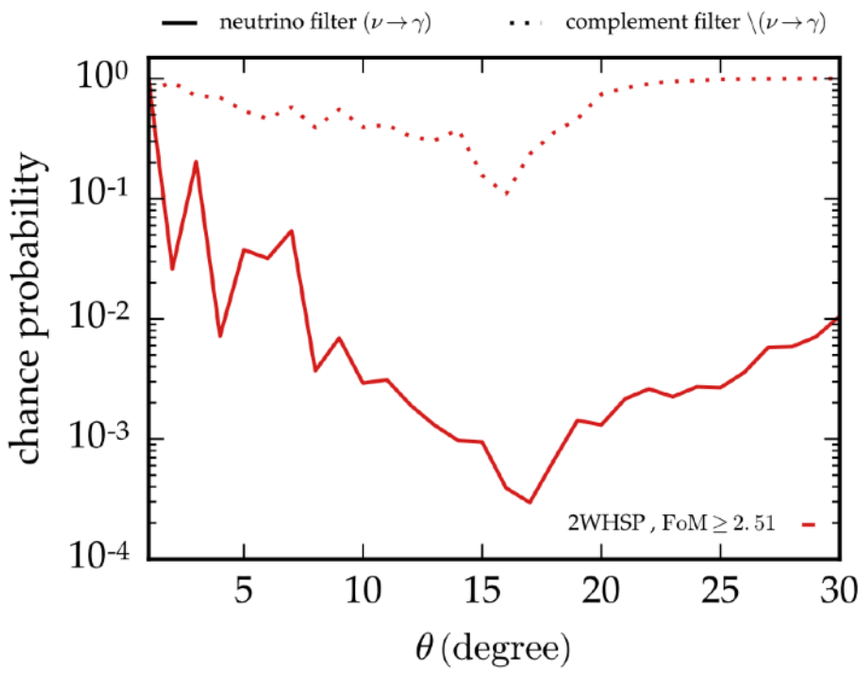
- ➡ No correlation with GRB, FRB
- ➡ Few hints with blazars (nothing significant)
- ➡ One hint with SN Ic (IC160427)

(See Alexis's talk tomorrow)

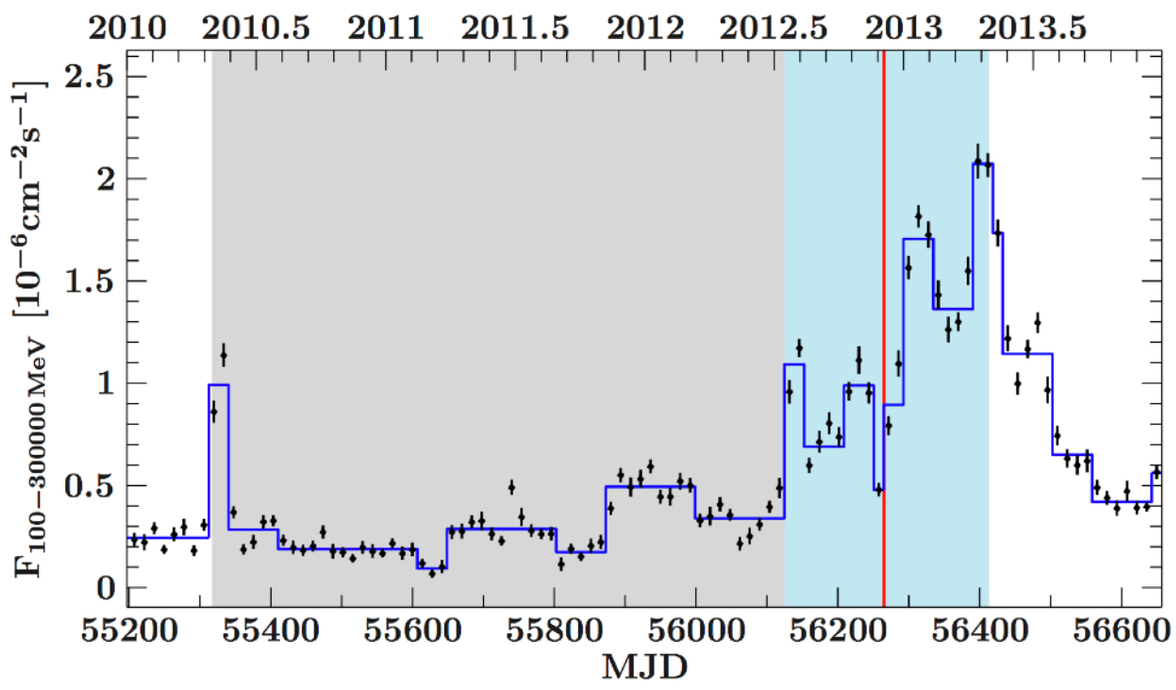


Connection ν-γ-UHECR

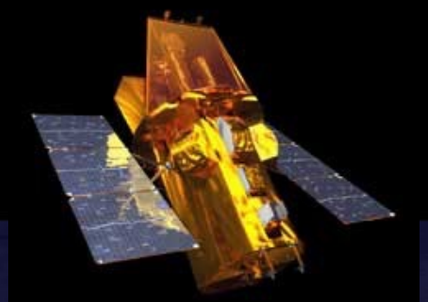
Resconi et al 2017, 2.9 sigma correlation with sub-sample of HBLs, IC nu and Auger UHECR



IceCube- Big Bird PKS B1424-418



Multi-λ observatories linked to ANTARES for the real-time analysis



SWIFT

INTEGRAL

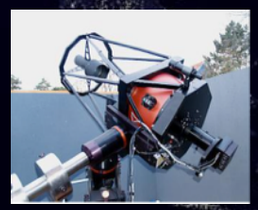
MASTER

HAWC

ANTARES

SVOM

GWAC



TAROT

HESS

MWA

ZADKO

~~ROTSE~~

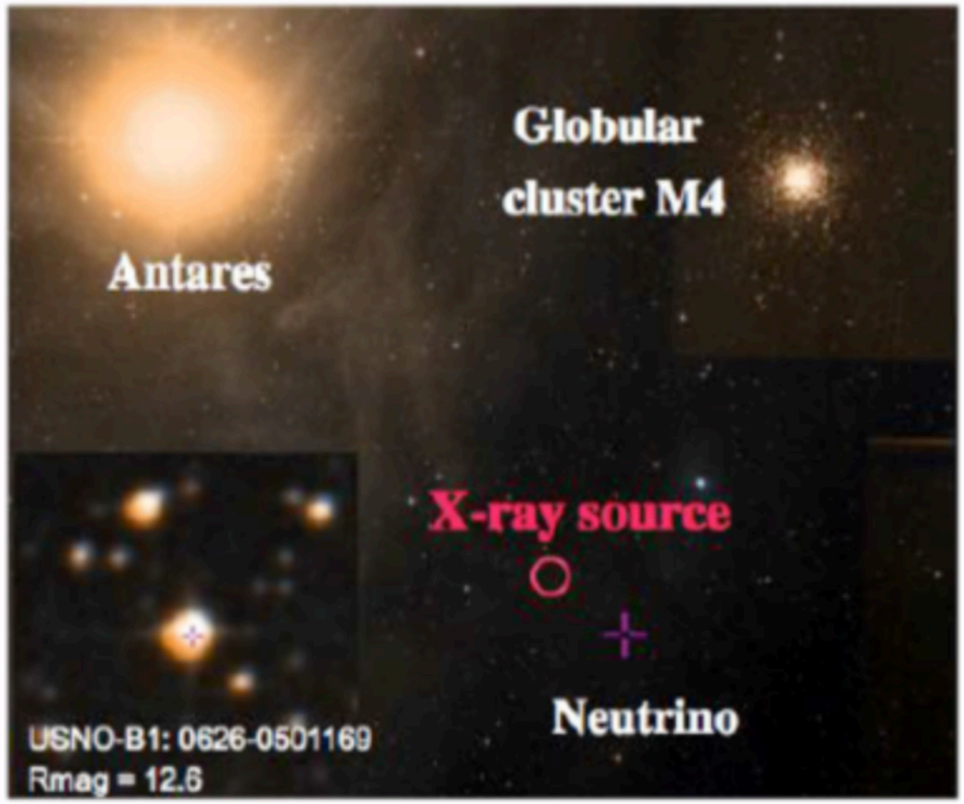
+ GW alerts from LIGO/VIRGO
+ UHECR/UHEν with Auger/TA
+ Neutrino alerts from IceCube
+ AMON

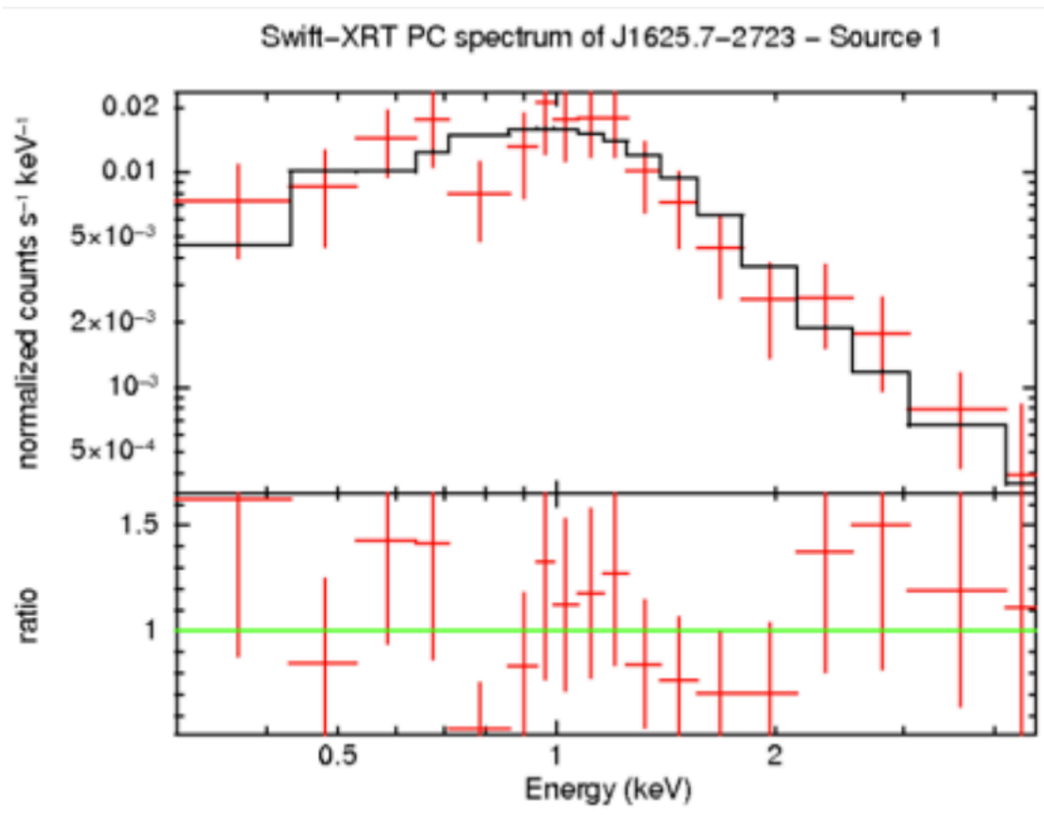
(Update 01/2018)

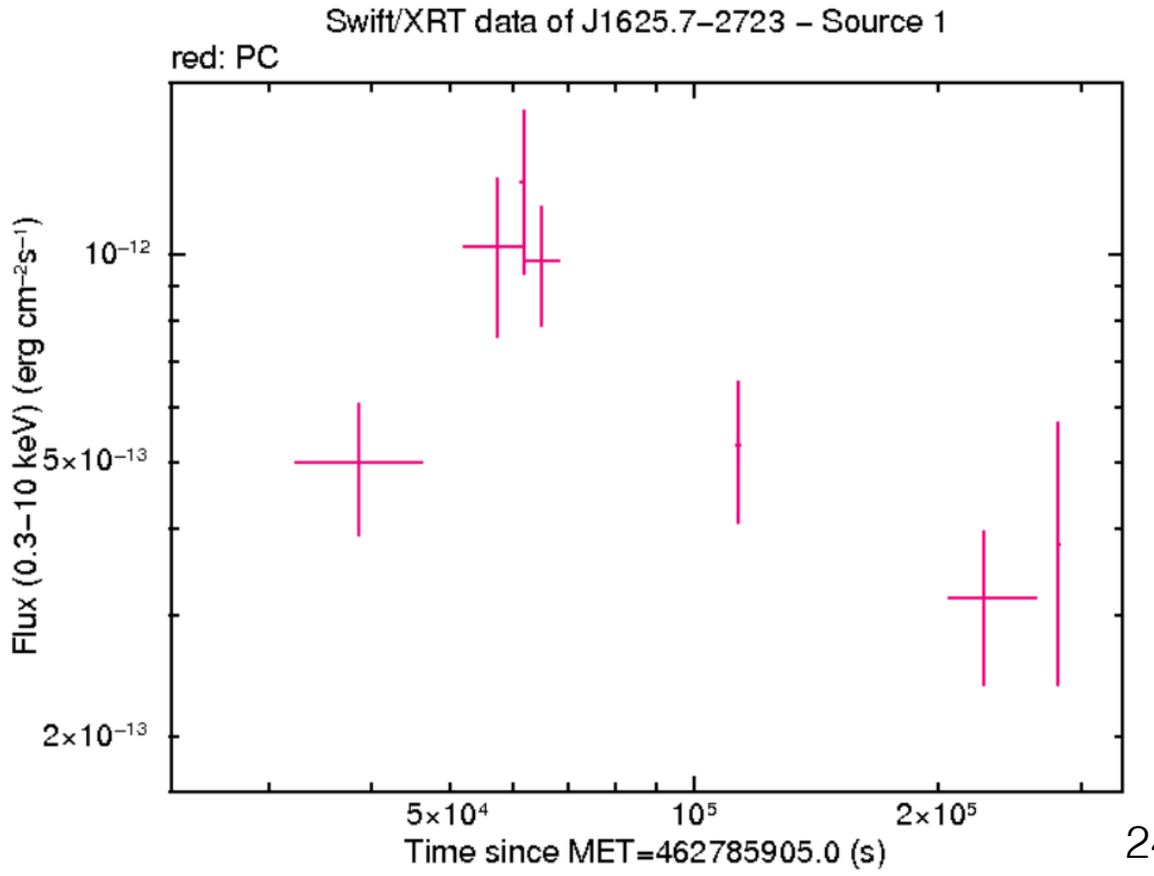
TAToO, one exemple: ANT150901A

TATOO: ANT150109A ALERT

- E ~50-100 TeV
- Error box=18 arcmin
- Sent in 10s to Swift and Master
- Swift obs: +9h
- Master obs: +10h







TAToO, one exemple: ANT150901A

- Neutrinos
  - IceCube: ATel 8097
- Optical
  - Pan-STARRS: ATel 7992, 8027
  - SALT: ATel 7993
  - NOT: ATel 7994 GCN18236
  - WiFeS: ATel 7996
  - CAHA: ATel 7998, GCN18241
  - MASTER: ATel 8000 GCN18240
  - LSGT: ATel 8002
  - NIC: ATel 8006
  - ANU: GCN18242
  - GCM: GCN18239
  - VLT/X-shooter
- X-rays
  - Integral: ATel 7995
  - MAXI: ATel 8003
  - Swift: ATel 8124, GCN18231
- Radio
  - Jansky VLA: ATel 7999, 8034
- Gamma-rays
  - MAGIC: ATel 8203
  - Fermi-GBM: GCN18352
  - HAWC
  - HESS

Great interest by astro-community

25

TAToO ⟹ GWAC@SVOM

- MoU between ANTARES-SVOM signed summer 2017

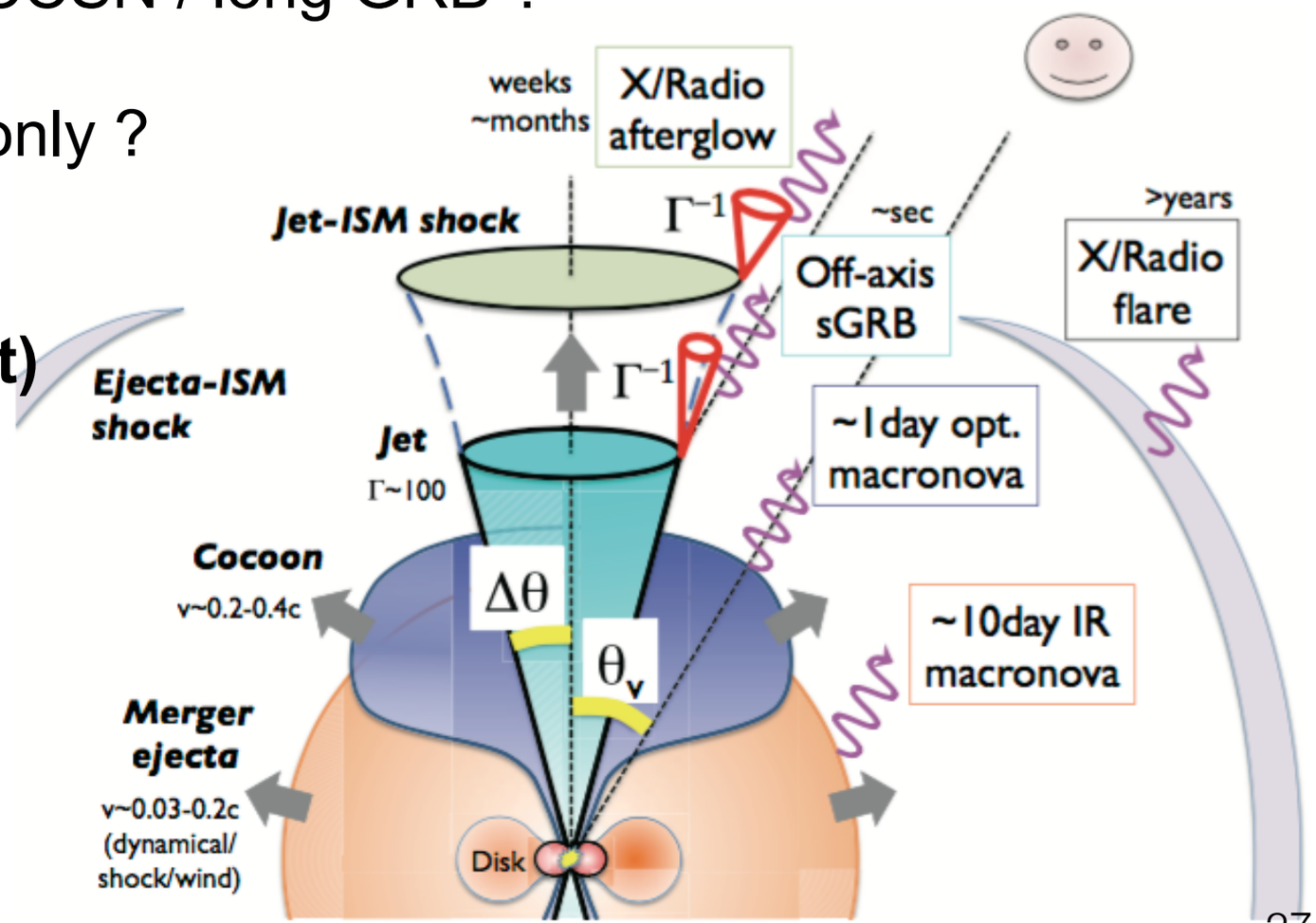
- 29% of the ANTARES neutrino triggers visible at Xinglong, 20% have been followed with mini-GWAC + 30 cm telescope [in agreement with the location of the telescopes in Xinglong and the weather constrains]
  ⇒ Delay: [0, 50min] for mini-GWAC (auto), [40, 1h40] for 30cm (manual)

  ⇒ **No counterpart identified [mag<12 for mini-GWAC, mag<≈18 for 30cm]**

- **Use large FoV of (mini)GWAC and the fact that the optical transients detected in the images are stored to look for fortuitous/offline follow-up of ANTARES cascade events [need to have 2017 cascade reconstruction]**

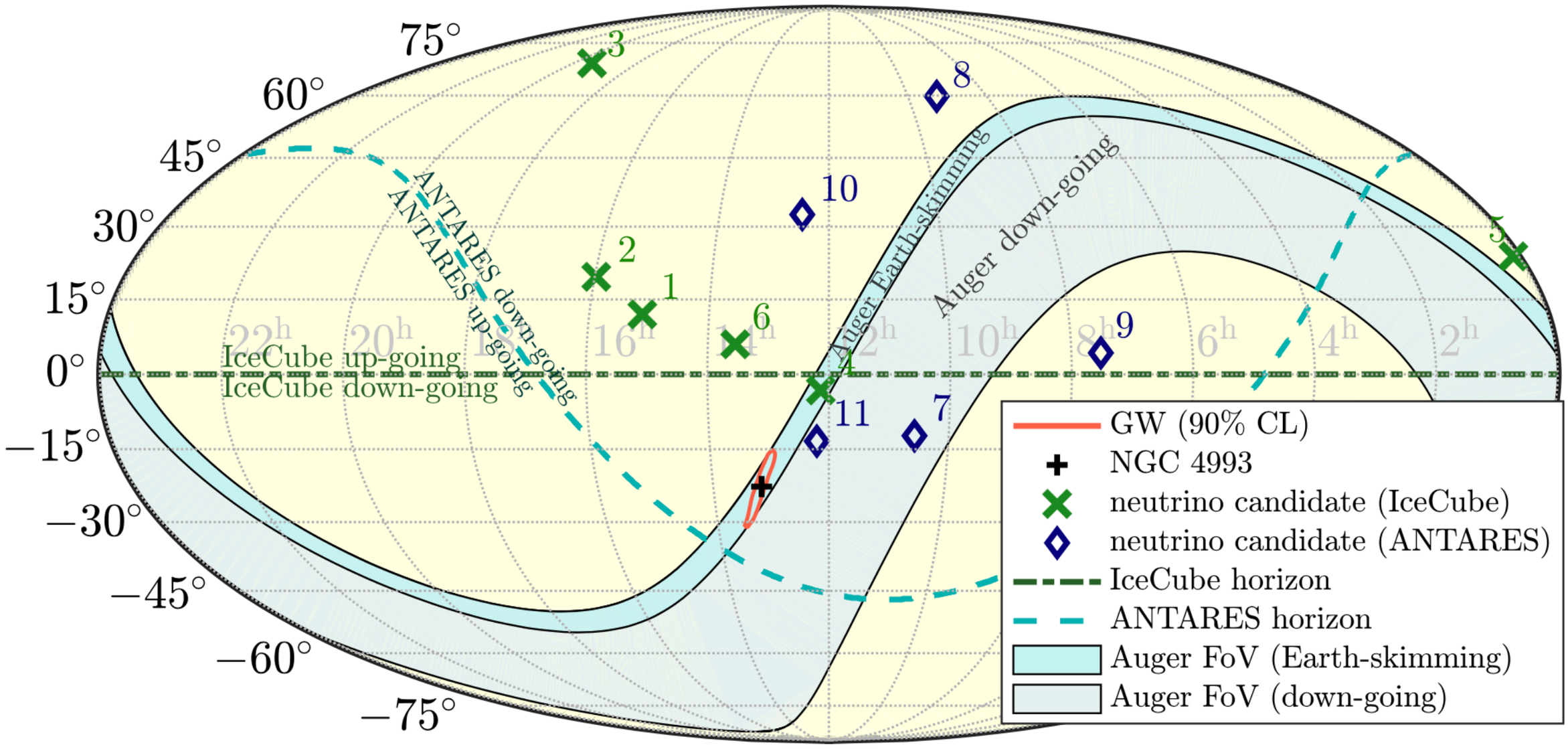
GW + neutrinos

- **Coincident neutrino observation could constrain the source position** (LIGO/Virgo 90% contour ~30°² vs ANTARES/KM3NeT: ~0.5°)
- Low level of background for transients → a few neutrino enough to have significant implications for GRB physics.
- Would confirm hadronic content of relativistic jets + dissipation processes.
- Chocked jets and unified picture CCSN / long GRB ?
- Dark bursts observed in nu+GW only ?
- **Relativistic jet formation/ dynamics (hadronic component)**
  - ✳ on-axis / off-axis
  - ✳ jet aperture
  - ✳ shape of the outflow (cocoon / chocked jet ?)
  - ✳ lower energy neutrinos with ORCA (~10 - 100 GeV)



NEUTRINO FROM GRB170817/GW170817 ?

Joint analysis of all HE-UHE neutrino experiments: ANTARES + IceCube + Auger + LVC

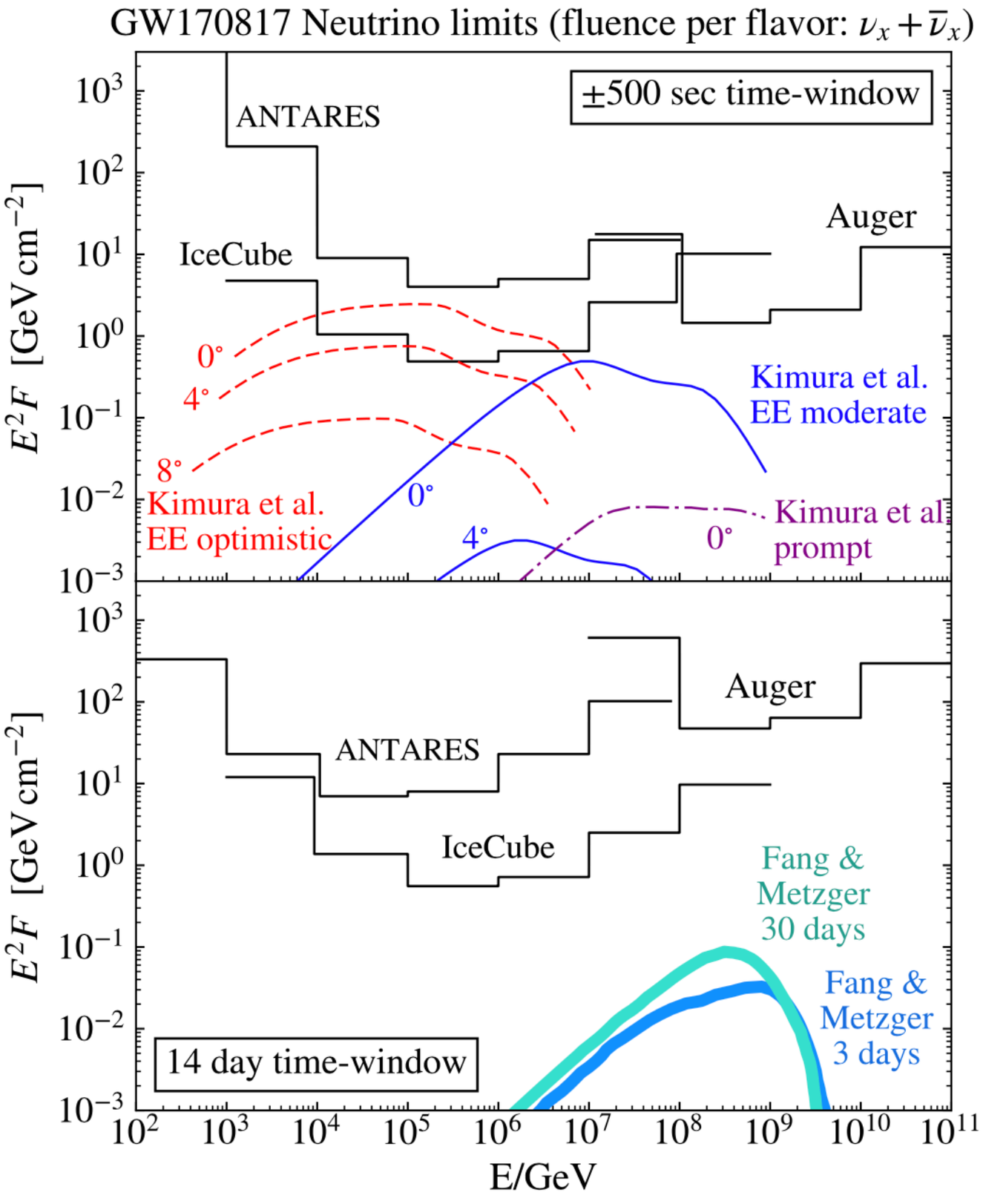


$\Longrightarrow$ **Bad luck localisation for ANTARES/IC, very very lucky for Auger**

$\Longrightarrow$ **No neutrino in space & time coincidence (+/-500s)**

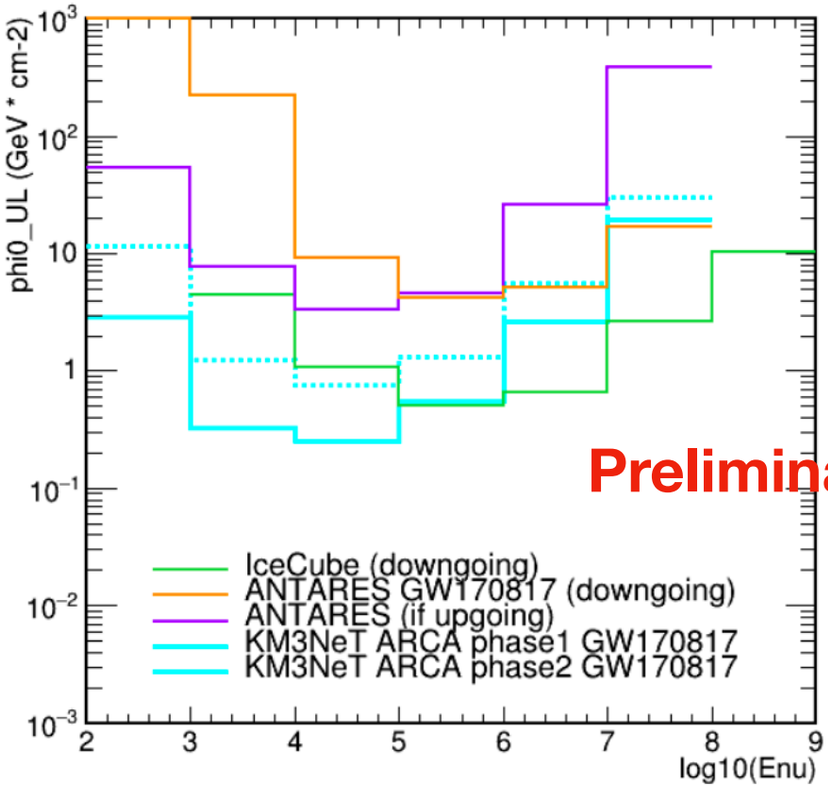
ApJL 850 L35 (2017)

NEUTRINO FROM GRB170817/GW170817 ?



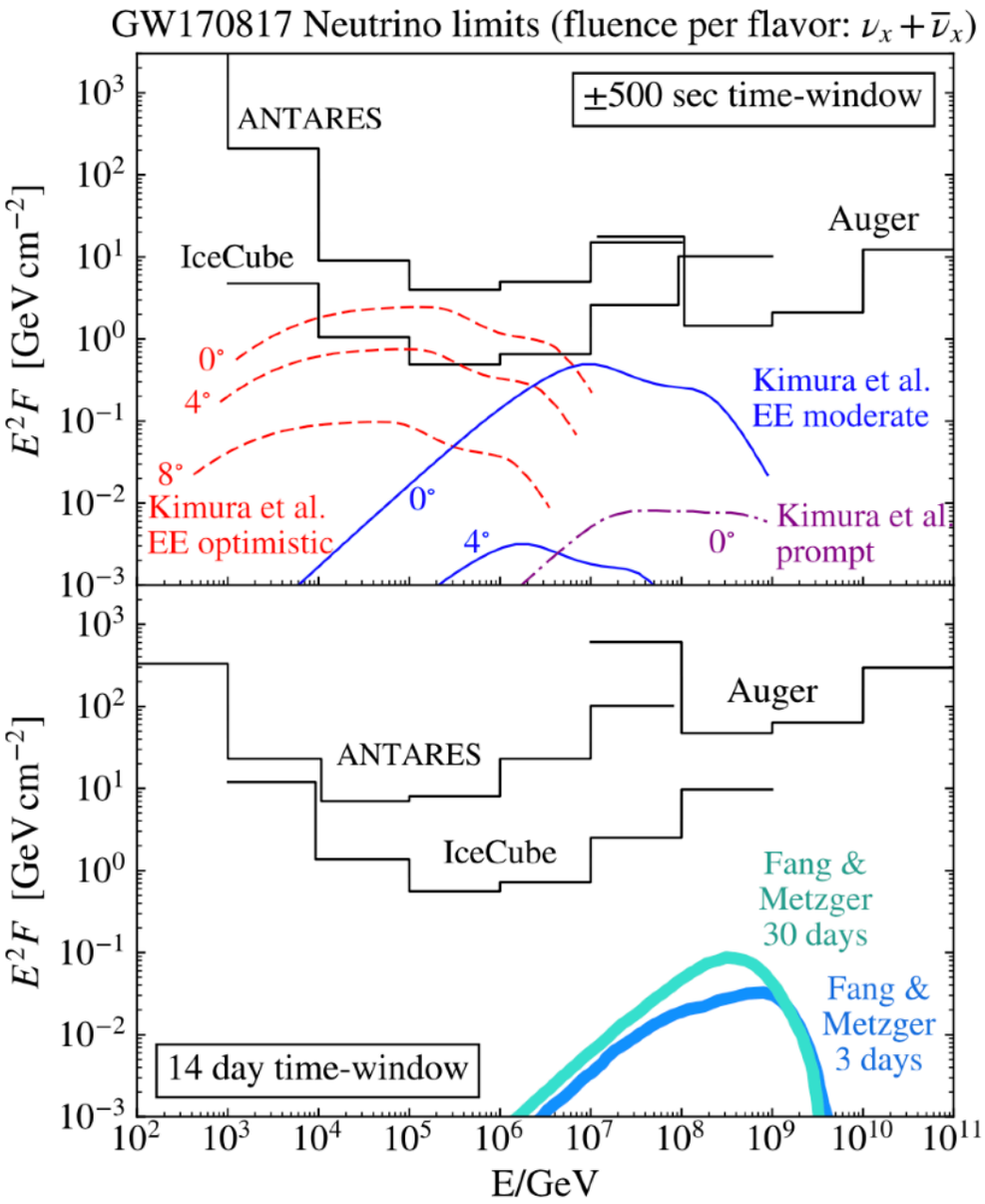
⟹ **Limited upper-limits from ANTARES/IC** (very high background contaminations)

⟹ **Comparison with hadronic model predictions (prompt emission from off-axis GRB or extended emission)**



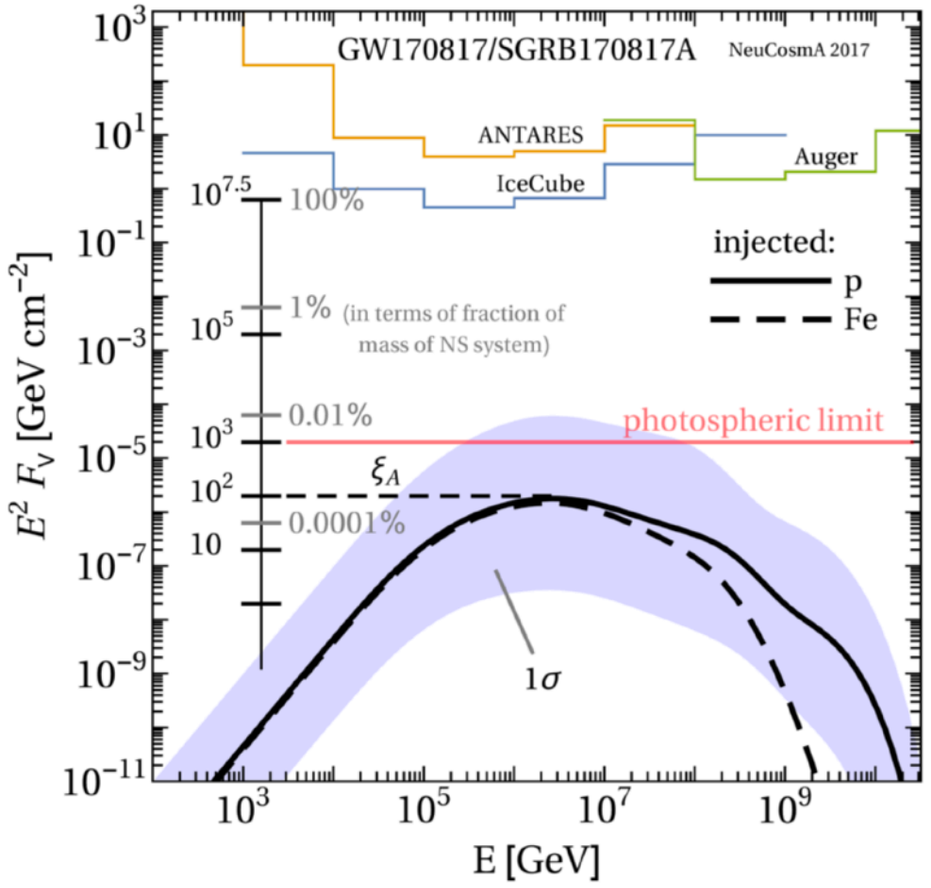
ApJL 850 L35 (2017)

NEUTRINO FROM GRB170817/GW170817 ?



Kimura, Murase, Mészaros & Kiuchi, ApJL, 848 (2017) L4

≠



Biehl, Heinze, Winter, arXiv:1712.00449

➡ **For this special event, very different computations. Need to be ready for the next interesting events**

Summary

Multi-messenger astronomy era ! (GW + neutrino)

- **Diffuse flux of cosmic neutrinos observed by IceCube**
- **Higher level of hadronic activity in the non-thermal universe than previously thought**
- **Sources remain to be identified. Hints are pointing in MM analyses. We are quite closed !**

Exciting times ahead !

⇒ KM3NeT: phased approach to next-generation neutrino telescope

ARCA (KM3NeT-It) **for HE neutrino astronomy (tracks & showers)**
ORCA (KM3NeT-Fr) **for measurement of neutrino mass hierarchy**
→ First strings performing well !!!

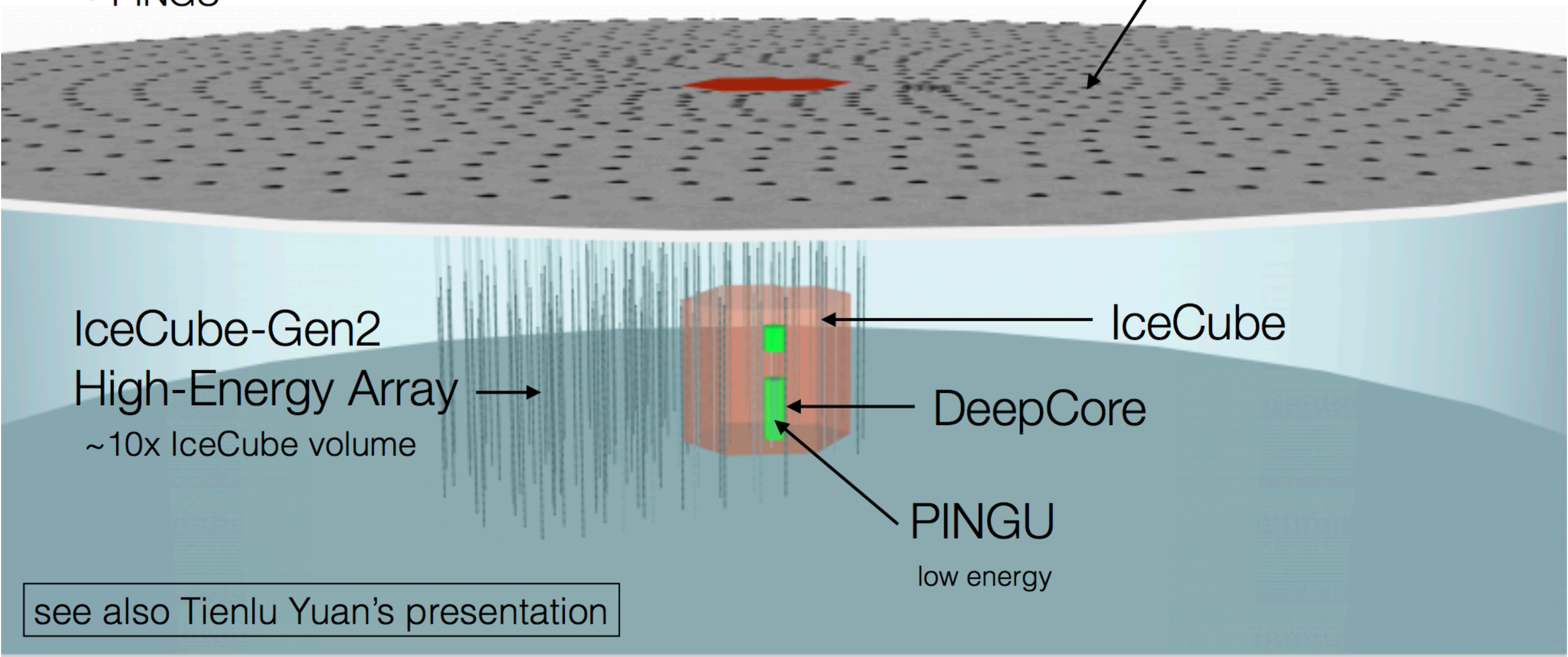
- **Start to implement the multi-messenger programs in KM3NeT for both ORCA and ARCA based on the successful experience of ANTARES.**
- *The follow-up of gravitational waves have worked very well and the community is organizing itself to get an even better follow-up of GW events. Neutrinos are a bit left in this structurant process. Need to think more in a multi-messenger manner rather than separated the messenger.*

The IceCube-Gen2 facility

A wide band neutrino observatory (MeV – EeV) using several detection technologies – optical, radio, and surface veto – to maximize the science

Multi-component observatory:

- IceCube-Gen2 High-Energy Array
- Surface air shower detector
- Sub-surface radio detector
- PINGU



see also Tienlu Yuan's presentation

KM3NeT data policy:

- ➡ KM3NeT neutrino data are proprietary but become public after a latency of 2 years after the data taking.
- ➡ However, significant events might trigger alerts that will be distributed publicly to the astro community using standard VO event format within ~10s after the neutrino detection [Open Public Alerts]
- ➡ Sub-threshold alerts and multiplets will be distributed though private channel to observing teams upon MoU agreements.