

Cosmology: the Cosmic Dawn and Epoch of Reionization

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Why is this topic important?

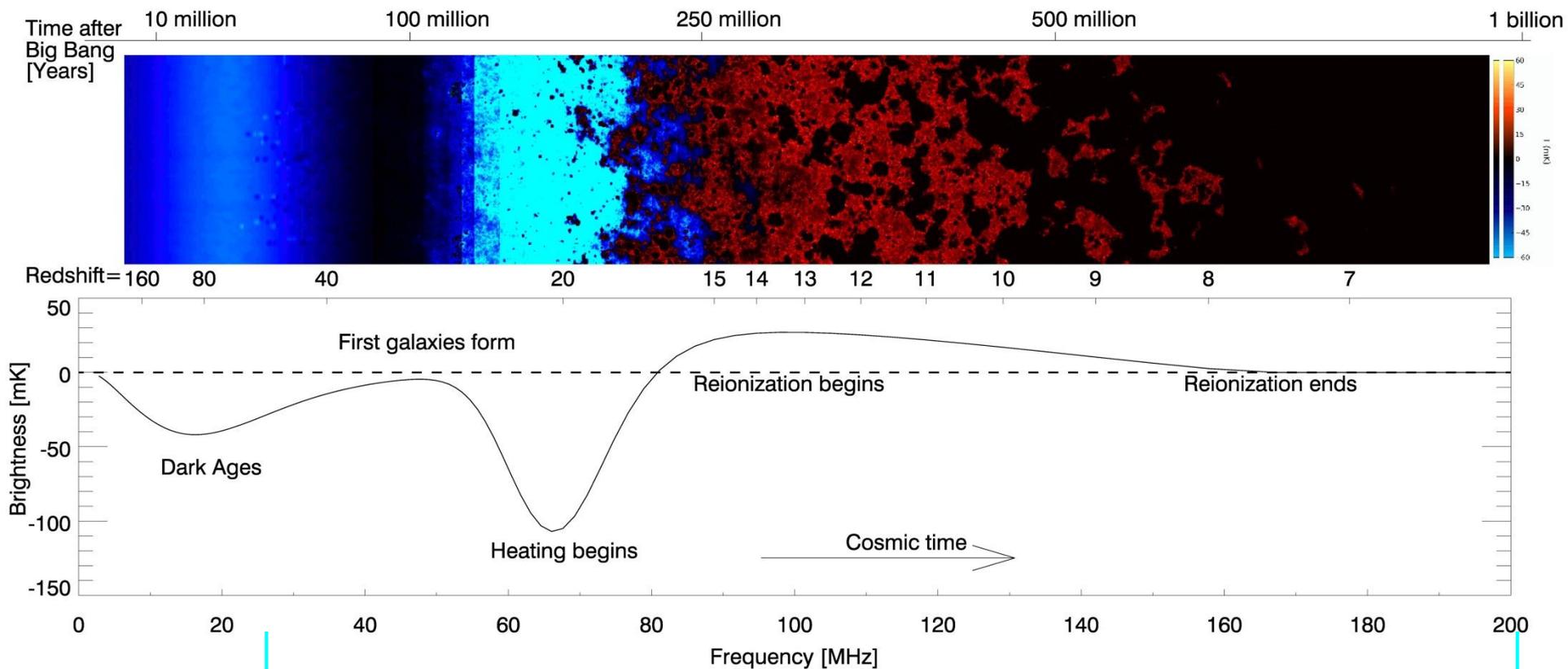
What is currently the biggest challenge in this field? Why?

What do we need to solve it?

Why would GRAND help solve it?

$$\nu = 1420 \text{ MHz} / (1 + z)$$

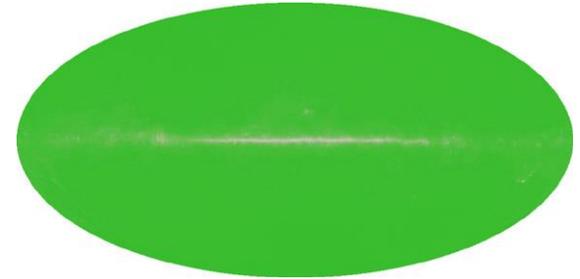
$$\lambda = 21 (1 + z) \text{ cm}$$



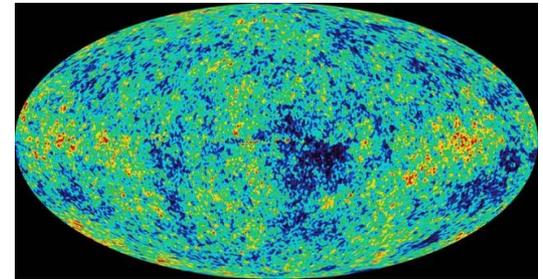
The best window

Measurements of the CD/EOR

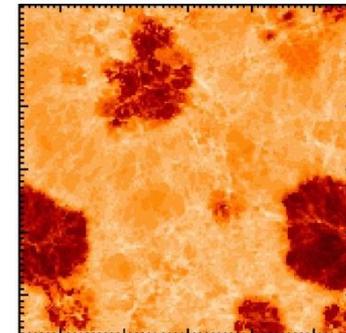
1. Global Signatures (total power)



2. Fluctuations (power spectrum)



3. Imaging (structures)



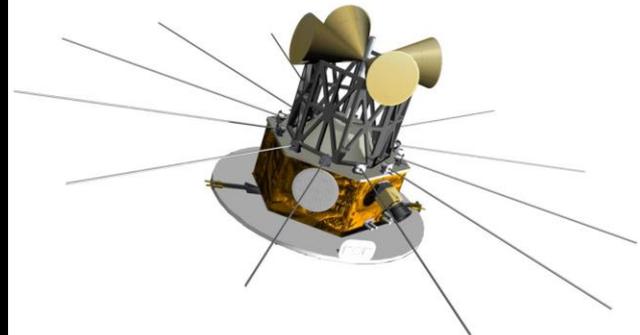


EDGES



LEDA

Global Signature



DARE



BIGHORNS



SCI-HI



SARAS



CORE



LOFAR



21CMA

Power Spectrum



LWA



PAPER



MITEoR

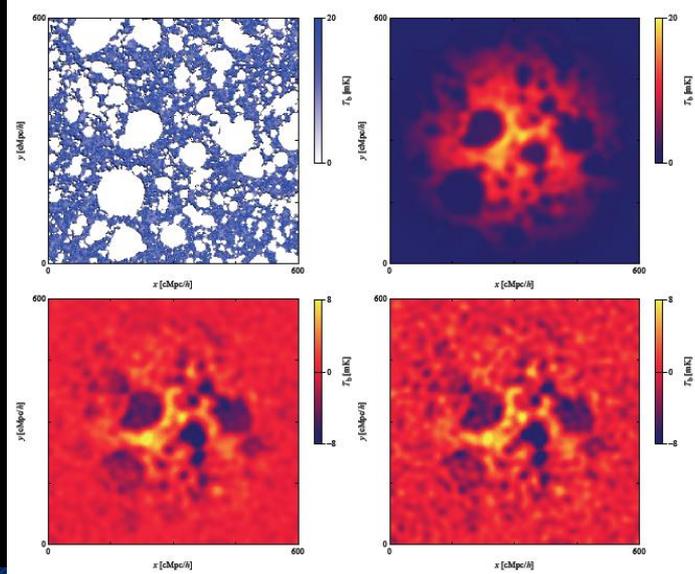


HERA



MWA

Deep imaging of EoR



Square Kilometre Array (SKA)

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Following Mauricio's Order

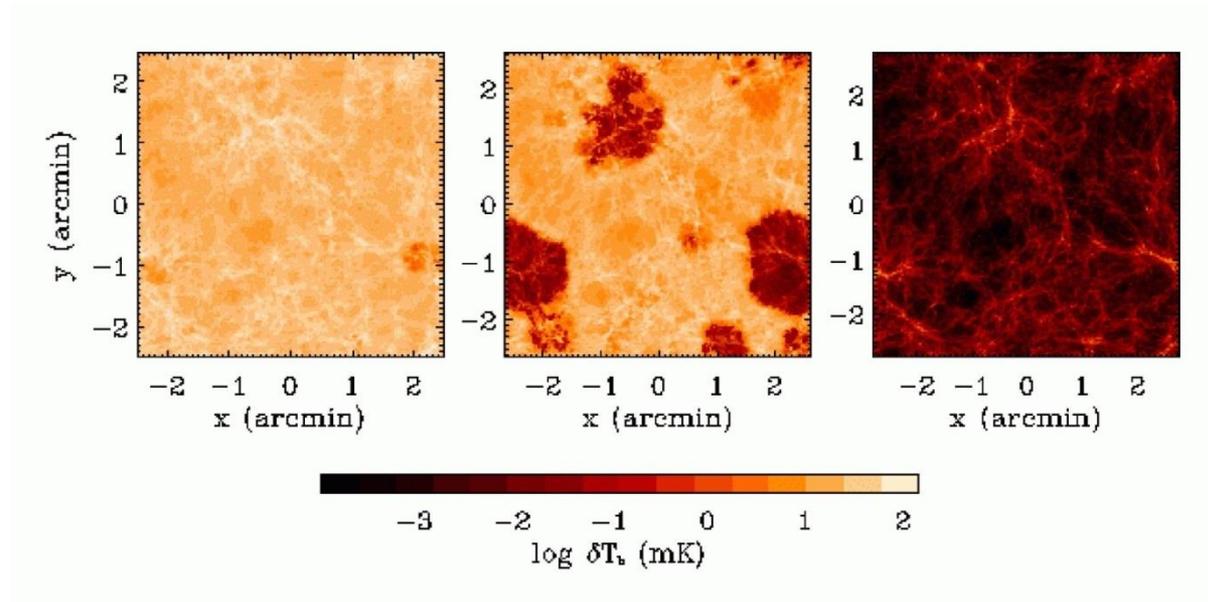
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What is currently the biggest challenge in this field? Why?

What do we need to solve it?

Why would GRAND help solve it?

Signal of EoR

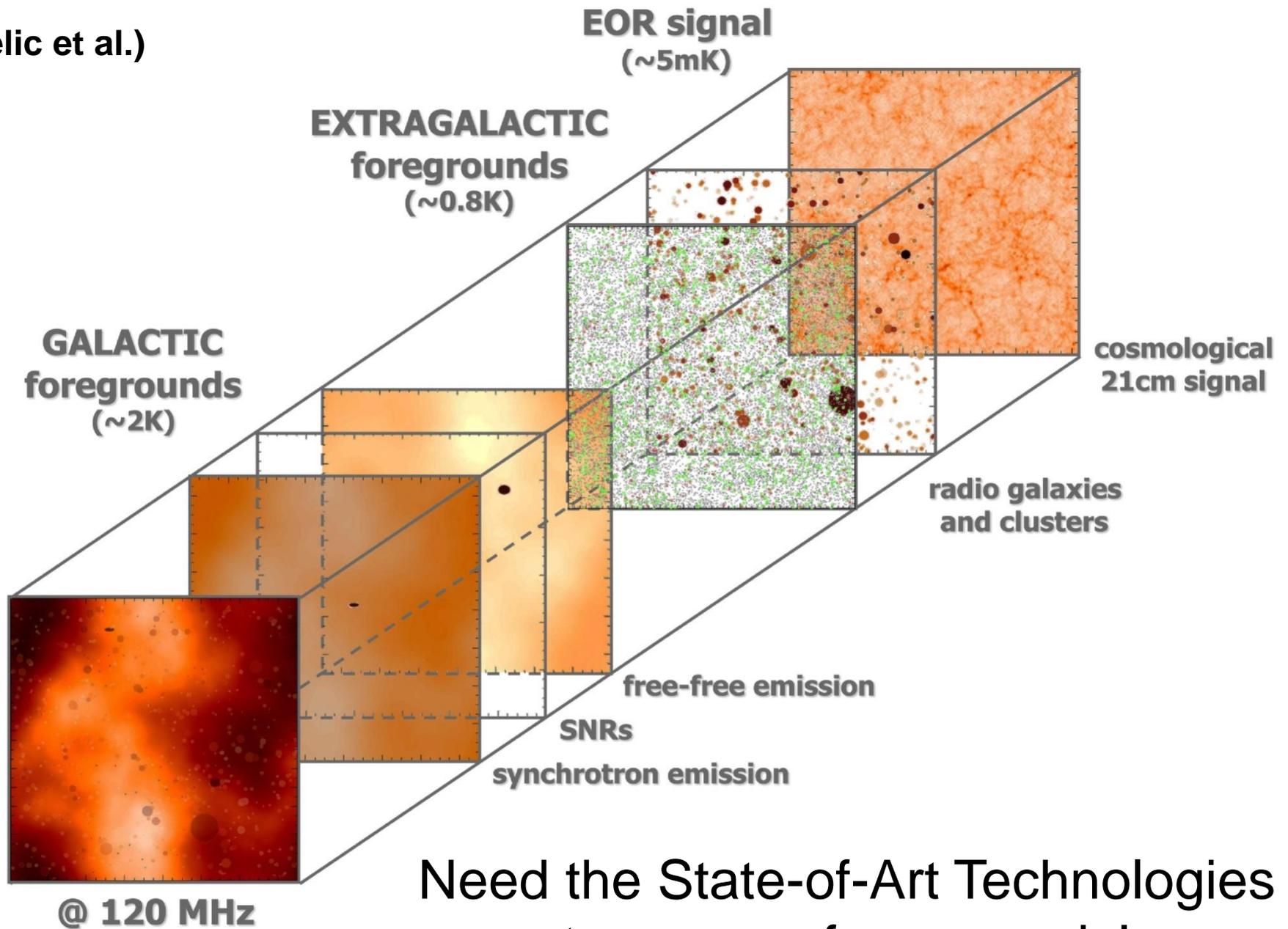


Furlanetto et al. (2003)

$$\delta T \approx 23.5(1 + \delta) x_H \left(\frac{T_S - T_{\text{CMB}}}{T_S} \right) \left(\frac{\Omega_b h^2}{0.02} \right) \left(\frac{0.15}{\Omega_M h^2} \right)^{1/2} \left(\frac{1+z}{10} \right)^{1/2} \text{ mK}$$

Need High Sensitivity
to see images and structures!

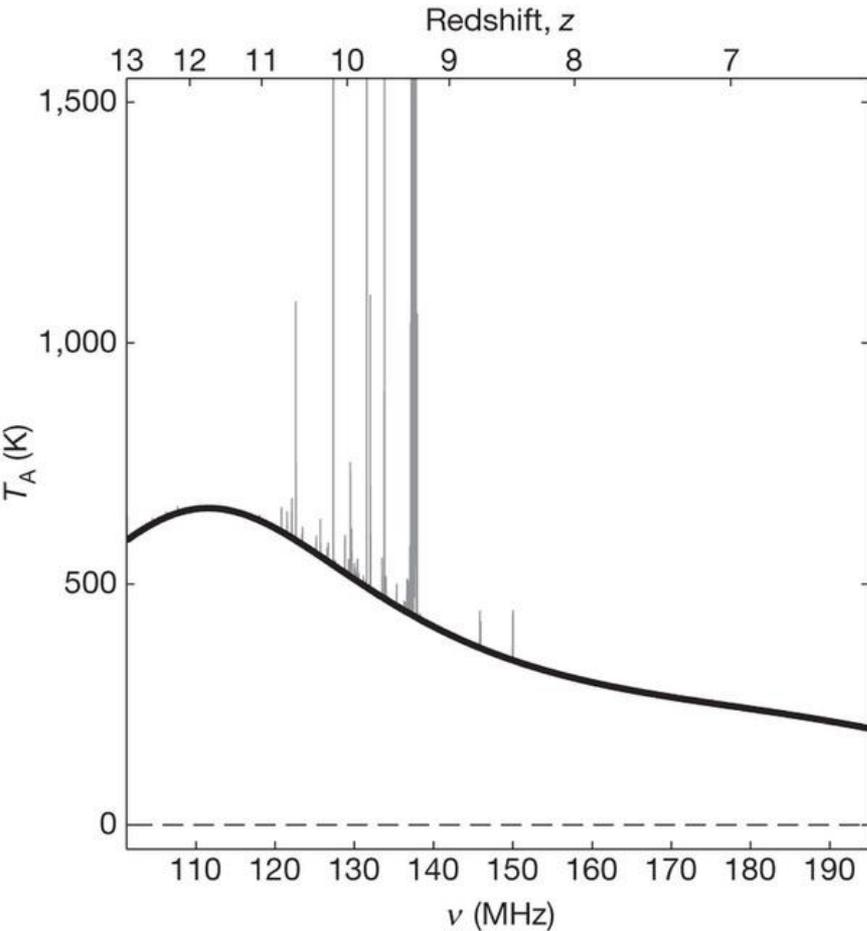
(Jelic et al.)



Need the State-of-Art Technologies to remove foregrounds!

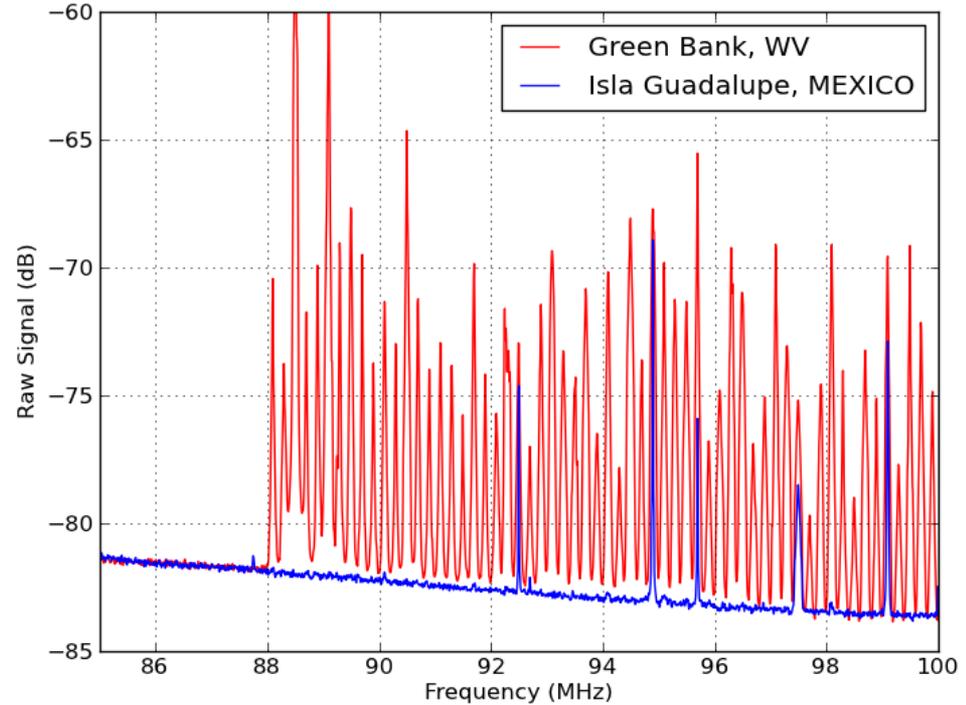
Examples: Measurements of Global Signature of EoR

EDGES



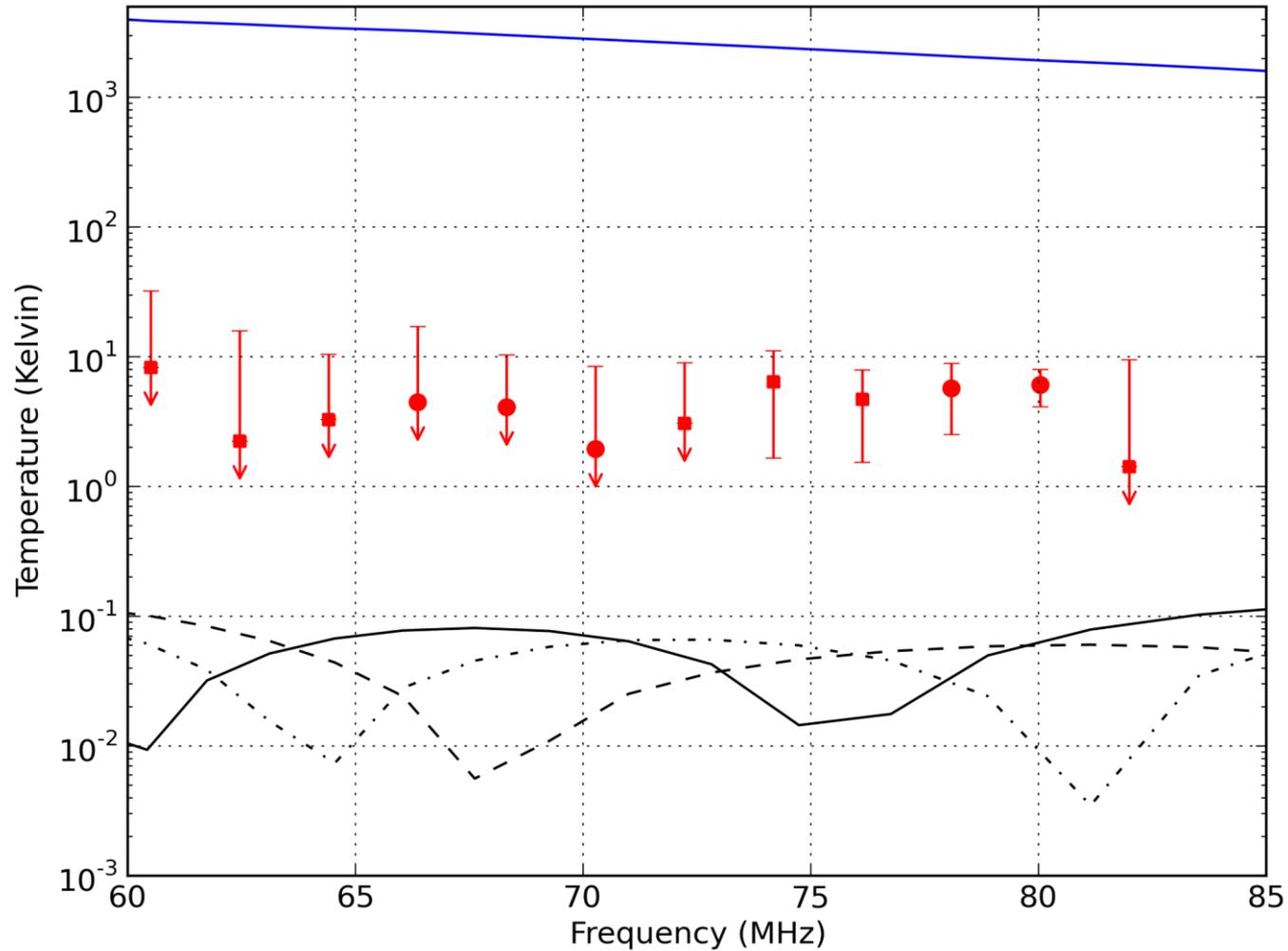
Bowman & Rogers
(2010, Nat 468)

SCI-HI



arXiv.1311.0014

SCI-HI



Magnitude comparison of foregrounds (blue), residuals from 4.4h of integration (red) and predictions (black)

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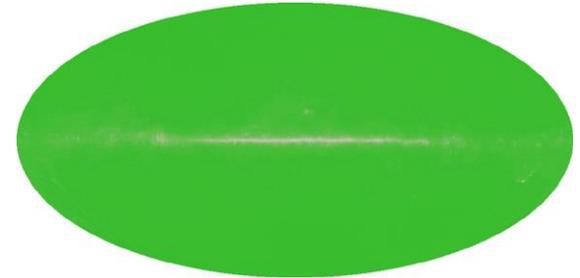
What do we need to solve it?

Why would GRAND help solve it?

GRAND's Targets

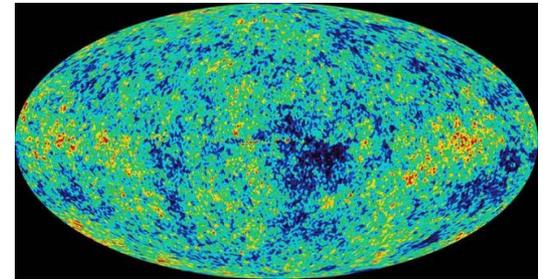
1. Global Signatures (total power)

(next a few years)



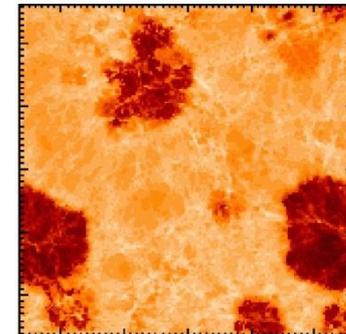
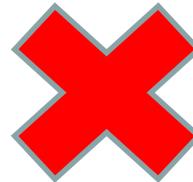
2. Fluctuations (power spectrum)

(next a few years)



3. Imaging (structures)

(next 10 years @ SKA only)



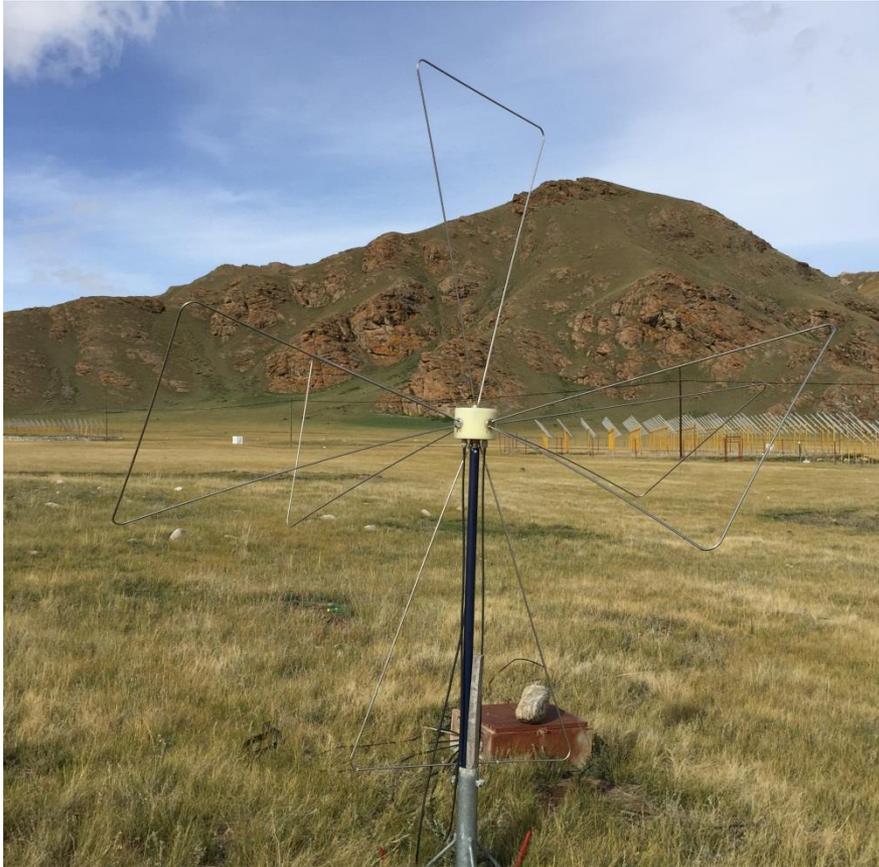
Detection of the Global Signature of CD/EoR @ GRAND



Technical Requirements

Frequencies:	10 - 200MHz
Sensitivity:	1mK
Integration time:	24h for 1 antenna
Stability:	1mK over 24 h
Calibration :	1mK
Foregrounds:	1mK

Detection of the Global Signature of CD/EoR @ GRAND



Disadvantages

Cosmic Rays & Neutrinos:

time domain: pulses@triggers

CD/EoR Detection :

frequency domain: noise

Different Working Modes !

Detection of the Global Signature of CD/EoR @ GRAND



Advantages

Many independent Antennas:

- 1. Good for Stability Control**
- 2. Good for Statistics**
- 3. Polarization Information**